

Searching far and long I: Pilot ALMA 2 mm follow-up of bright dusty galaxies as a redshift filter

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ABSTRACT

A complete census of dusty star-forming galaxies (DSFGs) at early epochs is necessary to constrain the obscured contribution to the cosmic star formation rate density, however DSFGs beyond $z \sim 4$ are both rare and hard to identify from photometric data alone due to degeneracies in submillimeter photometry with redshift. Here, we present a pilot study obtaining follow-up Atacama Large Millimeter Array (ALMA) 2 mm observations of a complete sample of 39 850 μm -bright dusty galaxies in the SSA22 field. Empirical modeling suggests 2 mm imaging of existing, bright samples of DSFGs selected at 850 μm -1 mm can quickly and easily isolate the “needle in a haystack” DSFGs that sit at $z > 4$ or beyond. With available redshift constraints, we find a positive correlation between redshift and 2 mm flux density. Specifically, our models suggest the addition of 2 mm to a moderately constrained IR SED will improve the accuracy of a millimeter-derived redshift from $\Delta z/(1+z) = 0.3$ to $\Delta z/(1+z) = 0.2$. Combining archival submillimeter imaging with our measured ALMA 2 mm photometry ($1\sigma \sim 0.08 \text{ mJy beam}^{-1} \text{ rms}$), we characterize the galaxies’ IR SEDs and use them to constrain redshifts. Our IR SED characterizations provide evidence for relatively high emissivity spectral indices ($\langle \beta \rangle = 2.4^{+0.4}_{-0.3}$) in the sample. From millimeter photometric redshifts we identify 7/39 high- z candidates each with $> 50\%$ likelihood for $z > 4$, and demonstrate the utility of 2 mm imaging as a redshift filter for DSFGs. We measure that especially bright ($S_{850\mu\text{m}} > 5.5 \text{ mJy}$) galaxies contribute $\sim 10\%$ to the cosmic-averaged CSFRD from $2 < z < 5$, confirming findings from previous work with similar samples.

1. INTRODUCTION

A primary goal in extragalactic astronomy is to understand the cosmic star formation rate density (CSFRD) across cosmic time (Madau & Dickinson 2014). The history of cosmic star formation measured at rest frame UV/optical wavelengths currently extends to $z \gtrsim 10$ due to the relative ease of identifying Lyman Break Galaxies (LBGs) in deep *Hubble Space Telescope* and *Spitzer Space Telescope* imaging (e.g. Ellis et al. 2013; Oesch et al. 2013; Bouwens et al. 2015; Finkelstein et al. 2015; Finkelstein 2016). However, measuring solely unobscured star formation by surveying optical and UV

emission alone presents an incomplete picture of cosmic star formation, as at least half of the star formation activity in the Universe is obscured by dust.

Dusty Star-Forming Galaxies (DSFGs) dominate cosmic star formation at its peak epoch at $z \sim 2 - 2.5$ (Casey, Narayanan, & Cooray 2014), with extreme star formation rates (SFRs) $\gtrsim 100 \text{ M}_\odot \text{ yr}^{-1}$, stellar masses $\gtrsim 10^{10} \text{ M}_\odot$, and high gas mass fractions $\sim 40 - 80\%$ (Carilli & Walter 2013). While high SFRs imply the presence of young, massive UV-emitting stars, the majority of the UV and optical stellar emission is absorbed and then re-radiated at longer wavelengths in the (sub)millimeter by dust. Therefore, in order to take a full census of cosmic star formation, multi-wavelength

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observations are necessary to capture both unobscured and obscured star formation.

Most DSFGs are discovered in wide-area observations using single dish far-infrared/millimeter instruments such as the Submillimeter Common-User Bolometer Array (SCUBA/SCUBA-2) on the James Clerk Maxwell Telescope (JCMT; e.g. Smail et al. 1997; Hughes et al. 1998; Chapman et al. 2005; Koprowski et al. 2017; Simpson et al. 2019), the *Herschel Space Observatory*¹ (Pilbratt et al. 2010; Eales et al. 2010; Oliver et al. 2012), and the AzTEC instrument (Scott et al. 2008; Aretxaga et al. 2011). Multi-wavelength follow up of hundreds of survey-identified DSFGs reveals that most sit between $1 < z < 3$ (Casey et al. 2012a; Magnelli et al. 2011, 2013; Gruppioni et al. 2013; Le Floc'h et al. 2005). While there exists handful of individually studied DSFGs at redshifts as high as $z \sim 5 - 7$ (e.g. Cooray et al. 2014; Strandet et al. 2017; Marrone et al. 2018; Zavala et al. 2018b; Casey et al. 2019; Reuter et al. 2020), such high- z systems have proven difficult to both identify and spectroscopically confirm. This is because (a) DSFGs at $z > 4$ are outnumbered by the dominant DSFG population at $z \approx 1 - 3$ and (b) there are serious physical and evolutionary degeneracies that make DSFG photometric redshifts highly uncertain (with precision often $\sigma_{\Delta z/1+z} \gtrsim 1$; Casey 2020). This latter point is often seen as a benefit: their strongly negative k -correction means that the flux density of DSFGs at $z > 1$ remains constant with increasing z for $\lambda_{\text{obs}} \gtrsim 850 \mu\text{m}$, meaning a DSFG at $z \sim 10$ can be observed as readily as a DSFG at $z \sim 1$ (Blain et al. 2002). However, when searching for high- z DSFGs, this negative k -correction is also a hindrance as it becomes difficult to identify redshifts for galaxies with only long-wavelength emission. This effect is further exacerbated by the (sub)millimeter color degeneracy between dust temperature and redshift. Thus, barring clear identification at other wavelengths, it is easy to confuse $z \sim 2$ and $z \sim 6$ DSFGs with solely submillimeter observations.

Previous studies lacked the completeness necessary to resolve whether high- z DSFGs play a significant role in cosmic star formation at $z > 3$ (e.g. Zavala et al. 2018a; Casey et al. 2018a,b). For example, Koprowski et al. (2017), Rowan-Robinson et al. (2016), Dudzevičiūtė et al. (2020), and Gruppioni et al. (2020) all present very different conclusions on the number density of IR luminous galaxies at $z > 3$, due to either limited sample sizes at high redshift or fundamental differences in

survey strategy. Additionally, using a blank-field 2 mm ALMA survey (Mapping Obscuration to Reionization with ALMA, or MORA), Zavala et al. (2021) found that DSFGs contribute $\sim 35\%$ at $z = 5$ and only $20 - 25\%$ at $z = 6 - 7$, suggesting dust-obscured star formation plays a minor role in total star formation at early epochs. On the other hand, Gruppioni et al. (2020) find a flatter IRLF slope from $4.5 < z < 6$, with a significant contribution to the CSFRD from obscured sources at high redshift, though again, using a very different survey strategy. While studies like Zavala et al. (2021) and Gruppioni et al. (2020) provide a good benchmark for the contributions of obscured and unobscured sources to the CSFRD, large uncertainties remain, necessitating a complete census of DSFGs at early epochs.

Empirical modeling by Casey et al. (2018a,b) derived from simulated DSFGs emphasizes the value of longer wavelength observations in easily and cheaply selecting high- z candidates. DSFGs selected at $850/870 \mu\text{m}$ with single wavelength measurements can be sorted roughly by redshift using interferometric 2 mm follow up imaging, which simultaneously provides higher precision astrometry due to a $\sim 10\times$ smaller beam size (e.g. da Cunha et al. 2021). Staguhn et al. (2014) report the first deep IRAM/GISMO 2 mm observations, and detect 15 sources in the Hubble Deep Field with a median redshift of $z = 2.9 \pm 0.9$, higher than the average redshift of $850 \mu\text{m}$ -selected DSFGs (e.g. Chapman et al. 2005, median $z = 2.2$ and an interquartile range $1.7 < z < 2.8$). Magnelli et al. (2019) present deep GISMO 2 mm observations in COSMOS and detect four out of five $z > 3$ DSFGs with known (sub)millimeter counterparts in COSMOS catalogs; they suggest 2 mm surveys favor detection of massive, extremely star-forming, high- z galaxies. Predictions from the SHARK model (Lagos et al. 2020) are also in line with this observation. The MORA survey — the largest ALMA 2 mm blank-field contiguous survey (184 arcmin^2) — demonstrates this, finding an average redshift $\langle z \rangle = 3.5^{+0.3}_{-0.2}$, with 77% of sources at $z > 3$ and 30% at $z > 4$, effectively filtering out lower redshift DSFGs (Casey et al. 2021). Other millimeter wavelengths have been leveraged to select for higher redshift sources; for example, Zavala et al. (2018a) conduct a blind search at 3 mm and detect 16 sources at $> 5\sigma$, and Williams et al. (2019) serendipitously discover a $z \sim 5 - 6$ source at similar wavelengths. Reuter et al. (2020) find $\langle z \rangle = 3.9$ for a sample of gravitationally-lensed DSFGs selected at 1.4 mm, where the lensing and millimeter selections combine to filter out low-redshift sources.

In this paper, we present Atacama Large Millimeter/submillimeter Array (ALMA) 2 mm observations of

¹ *Herschel* is an ESA space observatory with science instruments provided by European-led Principal Investigator consortia and with important participation from NASA.

$> 5\sigma$ bright sources ($S_{850\mu\text{m}} > 5 \text{ mJy}$) identified with SCUBA-2. Our goals are to identify the highest- z sources for further follow-up, provide an independent measurement of the volume density of $z > 3$ DSFGs, and test the practical utility of 2 mm follow up observations as a redshift filter for larger DSFG surveys. We describe the sample and observations in §2, and in §3 we present analysis of source redshifts and physical characteristics. §4 discusses the implications of our measurements, and §5 summarizes. We assume a Chabrier IMF (Kroupa 2001) and *Planck* cosmology throughout this paper, adopting $H_0 = 67.7 \text{ km s}^{-1} \text{ Mpc}^{-1}$ and $\Omega_\lambda = 0.6911$ (Planck Collaboration et al. 2016).

2. SAMPLE & OBSERVATIONS

2.1. Sample of DSFGs

In this project, we set out to measure 2 mm flux densities for a sample of bright 850 μm -selected DSFGs. To select our sample, we drew from the 850 μm -selected submillimeter galaxies (SMGs) observed with SCUBA-2 as part of the SCUBA-2 Cosmology Legacy Survey (S2CLS; Geach et al. 2017). S2CLS, initially conducted starting in 2011 on the James Clark Maxwell Telescope (JCMT), surveyed ~ 5 square degrees of sky at 850 μm . The largest and deepest survey of SMGs, the wide survey component of S2CLS includes seven extragalactic fields (UKIDSS-UDS, COSMOS, *Akari*-NEP, Extended Groth Strip, Lockman Hole North, SSA22, and GOODS-North) and reports 2851 submillimeter sources at 850 μm with a $\geq 3.5\sigma$ detection, with 1313 of those sources at $S_{\text{obs}} > 5 \text{ mJy}$. Over 90% of the total survey area reaches a sensitivity of below 2 mJy beam $^{-1}$, with a median depth per field of $\sim 1.2 \text{ mJy beam}^{-1}$ (Geach et al. 2017). The S2CLS survey is ongoing; see additional deep imaging in COSMOS presented in Simpson et al. (2019).

Our targets were first selected to lie in ALMA-accessible fields (COSMOS, UDS, and SSA22), and second, detected at $S_{850\mu\text{m}} > 5 \text{ mJy}$ with $> 5\sigma$ significance. We selected the brightest 850 μm -selected DSFGs for 2 mm follow up due to the observation that brighter DSFGs have a higher average redshift than galaxies at fainter submillimeter flux densities (Béthermin et al. 2015; Casey et al. 2018a,b). This paper presents a complete sample of 39 such bright DSFGs in SSA22 only; given the depth of SCUBA-2 observations, the $> 5\sigma$ threshold translates to an effective flux density cut of $S_{850\mu\text{m}} > 5.55 \text{ mJy}$ in SSA22 given the $\sim 1.2 \text{ mJy rms}$. The remaining targets in COSMOS and UDS are approved for observations in ALMA Cycle 8 and will be presented in a future work. Though the SSA22 subsample constitutes $\sim 9\%$ of the full sample to be observed,

the subsample is sufficiently large to generalize conclusions on the nature of the brightest $\gtrsim 5 \text{ mJy}$ DSFGs.

2.2. ALMA Observations

ALMA observations of our sample in SSA22 were taken on 8 January 2020 as one target field of Project #2019.1.00313.S (PI: Casey). In this field, 39 individual targets were observed with ALMA’s Band 4 centered at 2 mm. ALMA pointings were centered on the reported S2CLS 850 μm positions (Geach et al. 2017). Each target was observed for $\sim 1 \text{ min}$ using the 12-m array, with angular resolution 1. $''$ 2 and local oscillator tuning 145 GHz. The observations had a PWV of 2.05 mm, with 43 antennae and total on-source time of 34.12 min ($< 1 \text{ min}$ per source). The phase calibrator was J2217+0220 and the bandpass calibration was J2253+1608. The baseline limit with good phase (80%) was 279 m (roughly a C43-2 configuration).

The ALMA images were reduced using the Common Astronomy Software Application (**CASA**) version 5.6.0.² We adopted robust = 2 natural weighting to optimize imaging depth given the sources were unresolved at this resolution. The images were primary beam corrected, accounting for the primary beam response decreasing radially outwards from the center of the field. We derive flux density and noise measurements from primary-beam-corrected images using the **CASA** task **imstat**. The median rms achieved was 0.08 mJy beam $^{-1}$ (range [0.07–0.11] mJy beam $^{-1}$), better than the requested rms of 0.1 mJy. The synthesized beam for these observations is 1. $''$ 7 \times 1. $''$ 4. On these scales ($\sim 12 \times 14.5 \text{ kpc}$ at $z = 2$), it is expected that all detected sources are unresolved, therefore we adopt the peak flux density from **CASA** **imstat** within the SCUBA-2 beam (7.5 arcsec radius) as the measured 2 mm flux density of the source. We confirmed the robustness of this method by comparing to measurements of the source flux using **CASA viewer** across an extended aperture, and found good agreement between the two measurement methods.

Given that all of our targets had prior SCUBA-2 detections, we invoked a 3σ detection threshold for our ALMA observations. While a 3σ detection threshold is below the nominal 5σ threshold for blank field detection, other works have demonstrated that prior-based measurements are far less burdened by contaminants (e.g. Hodge et al. 2013; Dunlop et al. 2017). Out of the 39 targets, 35 targets were detected in the 2 mm observations above this threshold. Only one detected target (SSA.0007) was resolved into two sources, and the re-

² <https://casa.nrao.edu>

mainder were singletons. The $850\mu\text{m}$ and 2mm flux densities for all targets are listed in Table 1.

Comparing S2CLS reported coordinates to our measured coordinates of sources with flux densities $> 3\sigma$, we find an average RA offset of $1''.64 \pm 0''.92$ and an average Dec offset of $-1''.0 \pm 0''.7$, for a total net average offset of $1''.9 \pm 1''.2$. Some offset is expected, as SCUBA-2 is a single dish facility and achieves lower resolution and therefore lower astrometric precision than ALMA. We do find some systematic offset as illustrated by Figure 1, likely originating from the astrometric imprecision of SCUBA-2. Other studies find similar few arcsec systematic offsets between single-dish and ALMA detections (e.g. Hodge et al. 2013; Simpson et al. 2015).

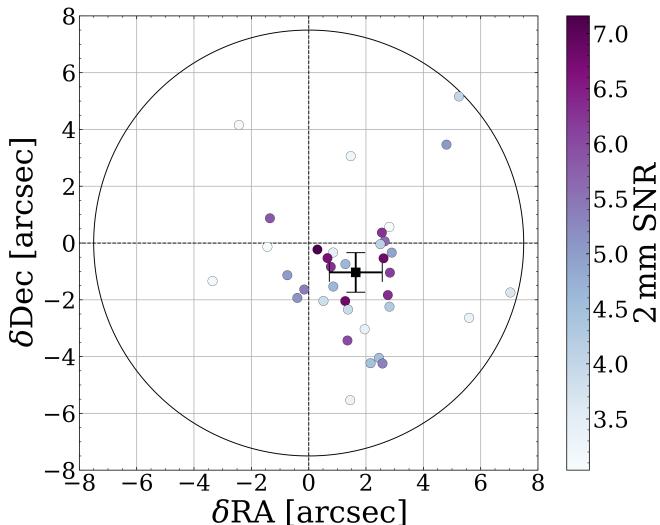


Figure 1: Positional offset between ALMA and SCUBA-2 centroids for the $S_{2\text{mm}} > 3\sigma$ sources in our sample, colored by single-to-noise ratio (SNR) in ALMA data centered on SCUBA-2 positions (pointing centers). The black square shows the average offset and uncertainties. While there is systematic offset between the SCUBA-2 and ALMA positions, nearly all detected sources lie well within the SCUBA-2 beam FWHM, represented by the black circle. Those far offset tend to have lower SNR suggesting a higher probability of chance alignment.

We adopt deboosted flux densities from S2CLS (Geach et al. 2017) to account for confusion and Eddington boosting, i.e. flux-limited sample bias wherein more sources have fluxes boosted rather than decreased by noise. The median offset between deboosted and observed $S_{850\mu\text{m}}$ flux densities in our catalog is 0.8 mJy ($\lesssim 20\%$ of the total flux density), less than the re-

ported uncertainty for each individual source's flux density (which averages to 1.17 mJy).

2.3. The SSA22 Field

The SSA22 field (RA 1950: 22:15:01.0, Dec 1950: 00:00:05.8; Cowie et al. 1994) was first observed as one of four small selected areas (SSA) chosen for deep multiband imaging followed by complete follow-up spectroscopy. The initial deep imaging was part of the Hawaii K-Band Galaxy Survey, and was first presented in Cowie et al. (1994), with follow-up spectroscopy detailed in Songaila et al. (1994) and Cowie et al. (1996).

The now well-known protocluster structure in SSA22 at $z = 3.09$ was first revealed by Steidel et al. (1998). Protoclusters are rich, overdense regions thought to be the progenitors of massive clusters in the local Universe (Overzier 2016). The SSA22 protocluster hosts intense star formation activity, and has an abundance of Lyman break galaxies (LBGs, e.g. Steidel et al. 1998) and other photometrically-selected galaxies (e.g. Uchimoto et al. 2012), Lyman- α emitters (LAEs, e.g. Hayashino et al. 2004; Yamada et al. 2012), Lyman- α blobs (LABs, e.g. Steidel et al. 2000; Matsuda et al. 2004, 2011), and a rare overdensity of DSFGs (e.g. Stevens et al. 2003; Tamura et al. 2009; Geach et al. 2005; Umehata et al. 2015, 2017).

Several works have noted that DSFGs may act as signposts for overdense regions or protoclusters at high redshift (Casey 2016, Lewis et al. 2018, c.f., Miller et al. 2015). A number of DSFGs have been detected in the SSA22 field, originally discovered with SCUBA-2 observations (e.g. Chapman et al. 2001, 2004a), and later detected in AzTEC or ALMA 1.1 mm maps (e.g. Tamura et al. 2009; Umehata et al. 2014, 2015, 2018), some of which have optical to near-infrared and/or millimeter to radio photometric redshifts consistent with the protocluster. Umehata et al. (2017, 2018) map a small central region (ADF22) of the protocluster with ALMA, revealing an unusually high number of DSFGs at the protocluster core, including intrinsically fainter DSFGs ($S_{1.1\text{ mm}} < 1\text{ mJy}$). Of these DSFG protocluster members, very few of those are $S_{850\mu\text{m}} > 5\text{ mJy}$. Thus, while SSA22 is home to this well known protocluster structure and overabundance of DSFGs, we wish to highlight that the field's population of $S_{850\mu\text{m}} > 5.55\text{ mJy}$ DSFGs are not overabundant or more common than the cosmic average based on measurements of bright $850\mu\text{m}$ number counts. In other words, there is no direct evidence for a statistical excess in the field due to the well-studied overdensity. We explore this further in §4.4.

2.4. Redshifts from the Literature

Name	R.A. hms	Dec. (J2000) dms	$S_{850\mu\text{m}}^{\text{deboost}}$ mJy beam $^{-1}$	$S_{2\text{mm}}$ mJy beam $^{-1}$
SSA.0001	22:17:32.41	+00:17:43.8	14.5±1.1	0.55±0.09
SSA.0002	22:16:55.62	+00:28:46.1	10.7±1.4	0.26±0.07
SSA.0003	22:16:59.83	+00:10:39.8	10.2±1.5	[0.2±0.07]
SSA.0004	22:16:51.24	+00:18:20.7	10.0±1.4	0.36±0.07
SSA.0005	22:17:18.79	+00:18:09.5	7.9±1.3	[0.23±0.08]
SSA.0006	22:18:06.61	+00:05:20.6	8.8±1.8	0.57±0.08
SSA.0007	22:17:37.02	+00:18:22.6	7.2±1.3	0.67±0.11
SSA.0007.1	22:17:37.02	+00:18:22.6	7.2±1.3†	0.34±0.08
SSA.0007.2	22:17:36.98	+00:18:20.61	7.2±1.3†	0.33±0.08
SSA.0008	22:18:06.46	+00:11:34.5	7.7±1.5	0.21±0.07
SSA.0009	22:17:33.93	+00:13:52.0	7.34±1.09	0.47±0.07
SSA.0010	22:17:01.11	+00:33:31.4	8.9±2.0	0.49±0.08
SSA.0011	22:18:27.88	+00:25:36.4	7.1±1.5	0.53±0.08
SSA.0012	22:17:43.24	+00:12:31.9	7.0±1.4	0.33±0.07
SSA.0013	22:16:57.31	+00:19:24.0	6.9±1.5	0.4±0.07
SSA.0014	22:18:15.26	+00:19:56.9	7.3±1.4	0.59±0.09
SSA.0015	22:16:52.19	+00:13:42.3	6.8±1.5	0.44±0.08
SSA.0016	22:18:06.19	+00:04:01.5	8.7±2.0	0.44±0.07
SSA.0017	22:17:44.04	+00:08:21.8	6.5±1.5	0.21±0.07
SSA.0018	22:17:28.31	+00:20:26.2	6.3±1.3	0.52±0.08
SSA.0019	22:17:42.26	+00:17:02.3	6.0±1.4	0.36±0.07
SSA.0020	22:18:27.20	+00:19:31.3	5.6±1.5	0.25±0.08
SSA.0021	22:17:41.35	+00:26:41.5	5.8±1.4	0.36±0.07
SSA.0022	22:18:23.61	+00:26:33.1	6.0±1.4	0.38±0.08
SSA.0023	22:18:17.21	+00:29:32.7	5.8±1.3	0.31±0.08
SSA.0024	22:18:10.12	+00:15:55.2	5.7±1.4	[0.18±0.07]
SSA.0025	22:17:09.51	+00:14:08.9	5.5±1.5	0.38±0.08
SSA.0026	22:18:13.51	+00:20:31.2	5.5±1.3	0.59±0.09
SSA.0027	22:16:32.20	+00:17:46.4	5.6±1.4	0.39±0.09
SSA.0028	22:16:50.06	+00:22:48.4	5.3±1.3	0.35±0.08
SSA.0029	22:18:27.10	+00:21:36.4	5.13±1.25	0.28±0.07
SSA.0030	22:18:33.06	+00:18:42.2	5.41±1.23	0.24±0.08
SSA.0031	22:17:17.43	+00:31:37.4	5.1±1.4	0.26±0.08
SSA.0032	22:17:32.31	+00:29:30.7	5.1±1.3	0.27±0.08
SSA.0033	22:17:03.52	+00:26:03.7	4.9±1.3	0.25±0.08
SSA.0034	22:17:02.27	+00:15:53.9	4.9±1.3	0.23±0.07
SSA.0035	22:18:06.72	+00:06:30.9	5.6±1.4	0.65±0.10
SSA.0036	22:17:31.78	+00:14:54.5	4.9±1.3	[0.20±0.07]
SSA.0037	22:18:29.06	+00:08:35.0	6.3±1.9	0.25±0.07
SSA.0038	22:17:02.95	+00:24:39.2	4.8±1.3	0.29±0.07
SSA.0041	22:18:34.99	+00:21:42.6	6.0±1.3	0.45±0.08

Table 1: All SSA22 targets observed in Band 4 with ALMA. Positions are in J2000, with 2 mm detected source positions from CASA `imstat` measurements. $S_{850\mu\text{m}}^{\text{deboost}}$ is the 850 μm flux density in mJy beam $^{-1}$ reported from S2CLS (Geach et al. 2017) and $S_{2\text{mm}}$ is our measured 2 mm flux density in mJy beam $^{-1}$. All sources with $\text{S/N} > 3$ from ALMA are considered detections given the SCUBA-2 detection as a prior; flux densities of non-detections at 2 mm are listed in brackets. †Reported flux density here is the sum of the two components as it is unresolved with this instrument.

We crossmatch spectroscopic and photometric redshifts from literature on the SSA22 field to our sample. The target SSA.0001 is a confirmed protocluster member reported in Umehata et al. (2014) with a spectroscopic redshift of $z = 3.092$. SSA.0007 is resolved into two components in our 2 mm data and is also spectroscopically confirmed as part of a dense group within the protocluster at $z = 3.09$, possibly in a multiple merger phase (Umehata et al. 2015; Kubo et al. 2016; Umehata et al. 2017, 2018, 2019). In addition to the sources confirmed in the protocluster, three galaxies are confirmed spectroscopically to be in the foreground of the protocluster: SSA.0009 at $z = 2.555$ (Chapman et al. 2005), SSA.0019 at $z = 2.278$ (Alaghband-Zadeh et al. 2012), and SSA.0031 at $z = 2.6814$ (confirmed via Lyman- α emission; Cooper et al. in prep). About a third of our targets (14/39) also have optical-infrared (OIR) photometric redshifts and about one quarter (9/39) have mm/radio photometric redshifts (beyond those we calculate herein) reported in Umehata et al. (2014). In §3.1, we analyze protocluster membership for our sample relative to likely foreground and background sources.

2.5. Ancillary Data

We crossmatch our sample to the X-ray catalog for SSA22 presented in Lehmer et al. (2009) to check for the presence of luminous active galactic nuclei (AGN). AGN could be a concern in using millimeter photometry to constrain redshifts, as AGN can heat ISM dust to temperatures warmer than galaxies without AGN. Fifteen sources have X-ray coverage but just two sources (SSA.0001 and SSA.0007) have significant X-ray luminosities (0.5–8 keV luminosities within $10^{43} - 10^{45}$ erg s $^{-1}$) indicating potentially powerful AGN activity for those targets. This suggests the sample as a whole is not dominated by X-ray bright AGN (extremely obscured AGN could produce hotter dust temperatures and cannot be ruled out without follow-up data from e.g. *JWST*), and that the presence of AGN is unlikely to impact the millimeter photometric analysis presented later in this paper as both SSA.0001 and SSA.0007 have secure spectroscopic redshifts.

We crossmatch our SSA22 sources with reported AzTEC 1.1 mm flux densities from Umehata et al. (2014) and find $\sim 70\%$ (27/39) of our sample is detected down to SNR of 3.8. This threshold corresponds to a flux density detection limit of $S_{1.1\text{ mm}} \gtrsim 2.45 - 4.55$ mJy beam $^{-1}$. Flux densities at 1.1 mm for the other 12 sources were extracted from the AzTEC map at the best available positions (precise ALMA 2 mm positions for 2 mm detections, else SCUBA-2 850 μ m po-

sitions), with one source (SSA.0010) lying outside the AzTEC field.

Herschel/SPIRE measurements of SSA22 at 250 μ m, 350 μ m, and 500 μ m were extracted using cross-matched positional priors from MIPS 24 μ m, data (Swinbank et al. 2014; Kato et al. 2016). Note that we do not use the MIPS 24 μ m data directly because the depth varies significantly across the field and the coverage is incomplete. Furthermore, at the expected redshifts of our sources ($z \approx 2 - 6$), observed 24 μ m does not probe rest-frame FIR emission, but instead measures the mid-IR, rich with complex Polycyclic Aromatic Hydrocarbon (PAH) emission and Silicon absorption that we are not attempting to constrain in this work. SPIRE maps were deblended by oversampling to 1"/pixel from the original *Herschel*/SPIRE 8" pixel maps, and the extended cirrus emission from the Milky Way's interstellar medium was subtracted. About half of our sample (20/39) overlaps with the MIPS 24 μ m coverage, and from this subset we identified 18/39 counterparts within the MIPS 24 μ m FWHM of 7". For another 19 sources, we directly extract flux densities from deblended *Herschel*/SPIRE maps using our best positional constraints. The flux densities of the final two sources were extracted from the point source *Herschel*/SPIRE maps (optimized for point source extraction, ideal for our unresolved sources), as they lay outside the Swinbank et al. (2014) catalog and maps. We compared flux densities of overlapping sources from the point source *Herschel*/SPIRE maps and the deblended maps/catalog, and find no statistical differences within uncertainty. All of our measurements from the *Herschel* data are confusion limited (Nguyen et al. 2010).

For sources without positional priors in the MIPS 24 μ m catalogs (i.e. those with flux densities extracted from the deblended maps directly) we conduct a statistical analysis of cataloged sources from Swinbank et al. (2014) to infer an appropriate rms for our *Herschel* flux density measurements. For each flux density measurement, we calculate the 68% confidence interval of the catalog rms values within a 0.01 dex bin centered on the measured flux density. As the data are confusion limited, we adopt the confusion error from Nguyen et al. (2010) for any sources with derived errors less than these confusion error limits. This is a conservative estimate of the noise given the uncertainty in positional priors of deblended catalogs; in other words, at or below the confusion limit, confidence in the accuracy of positional priors is greatly reduced.

For our photometric redshift and SED fitting, we add absolute flux scale calibration uncertainty in quadrature with the statistical photometric uncertainties, tak-

ing that calibration error to be 4% for *Herschel* (Bendo et al. 2013), 5% for SCUBA-2 (Simpson et al. 2020), 10% for AzTEC (Wilson et al. 2008), and 5% for ALMA (Braatz 2021).

3. REDSHIFT AND FIR SED ANALYSIS

In this section, we combine existing FIR/mm photometry with new 2 mm data to identify the highest redshift candidates and constrain dust SEDs for our targets.

3.1. Photometric Redshifts

We derive a millimeter-based photometric redshift using the **MMpz** tool for each galaxy in our sample, as detailed in Casey (2020). **MMpz** uses rest-frame FIR/mm reprocessed dust emission to derive a photo- z probability density distribution. The probability distributions are based on the measured distribution of galaxy SEDs in the empirical relation between rest-frame peak wavelength and total IR luminosity, i.e. the $L_{\text{IR}}-\lambda_{\text{peak}}$ plane. The technique accounts for intrinsic SED breadth as it probes a wide range of dust temperatures at fixed IR luminosity. Estimating redshifts in the long wavelength regime suffers from a strong degeneracy between (sub)millimeter flux density and colors with redshift; this algorithm is suited for data like these that lack other redshift constraints from spectroscopy or OIR photometry. Further, here we target a bright subset of DS-FGs, a regime where OIR photo- z estimates can often be less accurate given differential attenuation with wavelength due to complex geometries (Casey et al. 2012b; da Cunha et al. 2015). Note that **MMpz** assumes the dust emissivity spectral index $\beta = 1.8$, which is a parameter fit during the dust SED characterization described in §3.2. Fixing β in this way has a minimal impact for use in photometric redshift fitting given the large uncertainties on redshift constraints from temperature degeneracies; see Casey (2020) for further discussion.

Comparing **MMpz** photo- z results to 12 OIR photo- z 's from Umehata et al. (2014), we calculate $\Delta z/(1+z) = 0.13$ with no systematic offset. Note that our comparison to OIR photo- z 's excludes SSA.0028, which has a reported OIR photo- z of $z = 0$, inconsistent with both our FIR results and publicly available shallow SDSS optical imaging, which shows no obvious local universe source within $0.2'$ (Ahumada et al. 2020). Additionally, we exclude source SSA.0006 which has an OIR photo- z of $z = 6$; its detection at $24 \mu\text{m}$ likely precludes such a high redshift solution. While the sample of galaxies with existing spectroscopic redshifts is much smaller (limited to SSA.0001, SSA.0007, SSA.0009, SSA.0019, and SSA.0031), we find good agreement of $\Delta z/(1+z) = 0.04$ between our **MMpz** results and spectroscopic redshifts.

In order to characterize bulk properties of these galaxies in the absence of precise redshift information, we use **MMpz** photo- z probability density distributions to sort the sample into three redshift bins, with the middle bin at $2.6 < z < 3.6$, centered on the SSA22 protocluster redshift of $z = 3.1$. Note that while we use a protocluster category to sort the sample into redshift bins, we set a bin size much larger than the spectroscopically confirmed protocluster redshift range as we do not expect the presence of the protocluster to bias our sample; we quantify this further in Section §4.4. To bin the sample, we integrate the PDF in intervals of width $\Delta z = 0.1$ and then consider the 3 highest probability redshift intervals. If any of the redshift intervals with highest probability are within $\Delta z = 0.5$ of $z_{\text{SSA22}} = 3.1$, the galaxy is categorized in the protocluster bin. If the highest probability redshift intervals are lower or higher than the protocluster, they are binned as low- z or high- z , respectively. In a couple of cases, the 3 highest probability redshift intervals span multiple bins, with only one in the protocluster bin (e.g. protocluster and low- z , or protocluster and high- z). In these cases we categorize the galaxy in the protocluster bin. This binning technique accounts for the width of the photo- z PDF, rather than only considering the PDF peak. This results in low- z -binned sources ranging $1.2 < z < 2.4$, protocluster-binned sources ranging $2.5 < z < 3.6$, and high- z -binned sources ranging $3.4 < z < 4.7$. We verify this binning method by visual inspection to account for variation in PDF shape.

From our millimeter photo- z measurements, we find 15 low- z galaxies with median $\langle z_{\text{low-}z} \rangle = 2.03 \pm 0.13$, 17 candidate protocluster members with median $\langle z_{\text{protocluster}} \rangle = 2.98 \pm 0.17$, and 7 high- z galaxies with median $\langle z_{\text{high-}z} \rangle = 3.83 \pm 0.16$, as shown in Figure 2. For the full sample, we find a median $\langle z \rangle = 2.6 \pm 0.2$. Errors on the median redshifts are derived from bootstrapping.

Note that **MMpz** takes both color and luminosity into account, and presumes no redshift evolution of λ_{peak} (and therefore, T_{dust}). Therefore, we expect z_{MMpz} to have some correlation with 2 mm flux density. This correlation, as well as the utility of 2 mm flux densities as a redshift filter, is discussed further in §4.2.

3.2. IR SED Characterization

Most galaxies in the sample lack precise redshift constraints; for sources lacking spectroscopic redshifts, we use photometry to derive a redshift from **MMpz**. We then fix the best available redshift solution to infer basic characteristics about the galaxies' SEDs. Note this degeneracy limits the scope of our conclusions as it may lead to underestimated uncertainties. Since the new 2 mm data

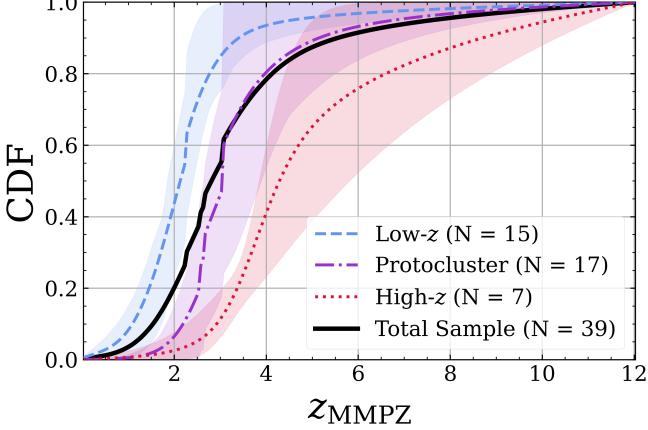


Figure 2: Redshift cumulative probability distribution functions (CDFs) from MMpz, with any spectroscopic redshift represented as delta functions at the measured redshift. The full sample is shown in black, and in blue, purple, and red are the CDFs for the low redshift, protocluster, and high redshift bins, respectively. The shaded regions show the 68% confidence interval for each redshift bin based on the individual MMpz fits. For the high- z sample, $\sim 60\%$ of the aggregate probability distribution is at $z > 4$.

is on the Rayleigh-Jeans tail, we focus our analysis on that portion of the spectrum.

Each galaxy's FIR/mm SED is fit to a modified blackbody added piecewise with a mid-infrared power law. We utilize a technique similar to that described in Casey (2012), but replace least-squares fitting with Bayesian analysis; the full tool will be presented in a forthcoming publication (Drew & Casey, submitted). The mid-infrared power law is joined to the modified blackbody at the point where the blackbody slope is equal to the power law index $\alpha_{\text{MIR}} = 2$ (consistent with other works, e.g. Kovács et al. 2010; Casey 2012; U et al. 2012). The general opacity model is assumed, where the optical depth (τ) equals unity at $\lambda_{\text{rest}} = 200 \mu\text{m}$ (e.g. Conley et al. 2011; Greve et al. 2012). Best fit SEDs are found based on a Markov Chain Monte Carlo convergence.

Input into the IR SED fitting routine is the best available redshift prior (fixed to the z_{MMPz} solution or spec- z if available) and the FIR/mm photometry (at $250 \mu\text{m}$, $350 \mu\text{m}$, $500 \mu\text{m}$, $850 \mu\text{m}$, $1100 \mu\text{m}$, and $2000 \mu\text{m}$) with associated uncertainties. Our fixed parameters are $\alpha_{\text{MIR}} = 2$ and $\lambda_0 = 200 \mu\text{m}$ (the wavelength where optical depth equals unity) near the intrinsic peak of the dust SED. As neither can be directly constrained in this dataset, these broad population averaged values are assumed. We do include a CMB correction term in our fitting procedure to account for ISM dust heating from

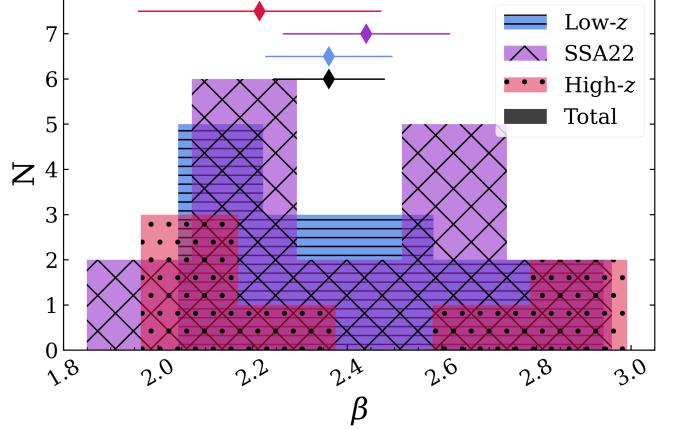


Figure 3: A histogram of measured emissivity spectral index (β) grouped by redshift bin. Low redshift sources are shown in blue, sources within the protocluster redshift in purple, and high redshift sources in red. The median β with uncertainty is plotted above the histogram for each redshift bin. All three distributions are consistent within error with median $\beta = 2.4 \pm 0.3$.

the CMB at high redshift (da Cunha et al. 2013), however, since the galaxies in our sample are predominantly at $z < 5$, this CMB correction is quite small for our sample in practice.

We used the SED fitting algorithm to find the best fit SED with measurements for each of the following free parameters: emissivity spectral index (β), total infrared luminosity (L_{IR} , taken from 8-1000 μm), dust temperature (T_{dust}), and rest-frame peak wavelength (λ_{peak}). The last two variables have a fixed relationship given our assumed opacity model with $\lambda_0 = 200 \mu\text{m}$ (see Figure 20 of Casey, Narayanan, & Cooray 2014).

Overall, our median sample properties are comparable to other samples of DSFGs. As our targets are selected for being among the brightest 850 μm -selected DSFGs in the SCUBA-2 survey, we find generally high IR luminosities such that all may be categorized as ultraluminous IR galaxies (ULIRGs, $L_{\text{IR}} \geq 10^{12} L_{\odot}$), with a median of $\log L_{\text{IR}} = 12.66 \pm 0.16 L_{\odot}$. Our sample is comparable to other samples with IR luminosities a few times $10^{12} L_{\odot}$ (Dudzevičiūtė et al. 2020; da Cunha et al. 2015; Miettinen et al. 2017).

We convert our L_{IR} measurements to a SFR using the TIR calibrator from Hao et al. (2011) and Murphy et al. (2011), consistent with the review Kennicutt & Evans (2012). We find characteristically high SFRs $\sim 500 - 1000 M_{\odot} \text{ yr}^{-1}$, with median SFR = $690 \pm 270 M_{\odot} \text{ yr}^{-1}$. While this is somewhat higher than other studies (e.g. $290 M_{\odot} \text{ yr}^{-1}$ in Dudzevičiūtė et al. (2020), and $280 M_{\odot} \text{ yr}^{-1}$ in da Cunha et al. (2015)),

we expect higher SFRs corresponding to our bright $S_{850\mu\text{m}} > 5.55 \text{ mJy}$ sample selection (the effective flux density cut given our criteria of $S_{850\mu\text{m}} > 5.5 \text{ mJy}$ detection at $> 5\sigma$ in SSA22 given the $\sim 1.2 \text{ mJy rms}$).

Peak wavelength measurements span $92 \mu\text{m} < \lambda_{\text{peak}} < 128 \mu\text{m}$, and consequently dust temperatures span $36 \text{ K} < T_{\text{dust}} < 53 \text{ K}$, with a median $T_{\text{dust}} = 44 \pm 4 \text{ K}$. [Dudzevičiūtė et al. \(2020\)](#) find a slightly lower median dust temperature of 30.4 K for their sample, with no evolution in T_{dust} at constant IR luminosity from $1.5 < z < 4$. [da Cunha et al. \(2021\)](#) also find a median $T_{\text{dust}} = 30^{+14}_{-8} \text{ K}$. We note that T_{dust} is not consistently defined and relies on assumptions of β and $\lambda - 0$, which are degenerate with T_{dust} (e.g. Figure 6 in [Spilker et al. 2016](#)). Further, [da Cunha et al. \(2021\)](#) find that optically-thin models tend to bias the temperature low compared with optically thick models.

Though our sample is $850 \mu\text{m}$ selected, and thus prone to bias towards intrinsically colder systems at $z \sim 2$ ([Eales et al. 2000](#); [Chapman et al. 2004b](#); [Casey et al. 2009](#)), our selection of brighter targets at higher L_{IR} is expected to correlate to hotter dust temperatures. [Lee et al. \(2013\)](#) use *Herschel*-selected DSFGs, which is unbiased at $z \sim 2$ with respect to dust temperature, and show that there are not many hot dust sources to be found unless they have luminous AGN. While obscured AGN could be the culprit, we cannot discern the presence of AGN without additional data. Therefore, our dataset represents a good sampling of galaxies above characteristic luminosity $L_{\text{IR}} > 3 \times 10^{12} L_{\odot}$ at the $S_{850\mu\text{m}} > 5.55 \text{ mJy}$ flux density limit; though there exists a bias with respect to dust temperature for this study, we don't expect many intrinsically hotter sources.

Overall, the measured emissivity spectral index, β , is consistent (within uncertainties) for the three subsamples at $\beta \sim 2.4$, with a median of $\langle \beta \rangle = 2.4 \pm 0.3$ for the full sample. The median value for the low- z bin is $\langle \beta_{\text{low-}z} \rangle = 2.36 \pm 0.13$, for the protocluster bin is $\langle \beta_{\text{protocluster}} \rangle = 2.44 \pm 0.17$, and for the high- z bin is $\langle \beta_{\text{high-}z} \rangle = 2.2 \pm 0.3$ (see Figure 3). Errors on the medians are derived from bootstrapping. The use of ALMA data with single-dish SCUBA-2 data, possibly suffering from confusion boosting, could impact individual derivations of β in this sample. Though we have accounted for deboosting as best as possible, the precision to which any individual β can be measured can only be improved with matched-beam ALMA data at both frequencies.

Using our 2 mm dust continuum emission, we derive a gas mass for each galaxy by Equation 16 in the Appendix of [Scoville et al. \(2016\)](#). We adopt a single mass-weighted dust temperature of 25 K (consistent with [Scoville et al. 2016](#); [Casey et al. 2019](#)) for

this calculation rather than our SED measurements of T_{dust} , which are luminosity-weighted. Our gas mass estimates show most of our targets are gas-rich with $M_{\text{gas}} \sim 10^{11} M_{\odot}$. With gas mass and SFR estimates in hand, gas depletion timescales can be derived in a statistical sense for the sample overall. We find gas depletion times around $\tau \sim 200 \text{ Myr}$ with median and standard deviation $\tau = 220 \pm 70 \text{ Myr}$, consistent with the majority of $z \gtrsim 1$ DSFGs (e.g. [Swinbank et al. 2014](#); [Dudzevičiūtė et al. 2020](#); [Sun et al. 2021](#)).

The best fit SEDs are presented with photometry overplotted in Figure 4, with best fit parameters listed in Table 2.

4. DISCUSSION

The goals of this work are to identify the highest- z sources from a sample of luminous DSFGs using 2 mm observations, characterize 2 mm emission and FIR SEDs for the sample, provide an independent measurement of the volume density of $z > 3$ DSFGs, and test the practical utility of 2 mm follow up observations as a useful and efficient redshift filter for large DSFG surveys. Based on FIR SED analysis for our galaxies, we place our sample into context with other DSFG populations, characterize the density of $850 \mu\text{m}$ -bright sources in the field, and evaluate the sample's contribution to the CSFRD.

4.1. Redshift Distribution of $S_{850\mu\text{m}} > 5.55 \text{ mJy}$ Sources

Based on the cumulative probability distribution function (Figure 2), $\sim 80\%$ of the aggregate probability density distribution is at $z > 2$. We categorize 15/39 galaxies as low- z , with a minimum redshift solution for an individual source of $z = 1.2^{+0.4}_{-0.4}$ for SSA.0023. We categorize 17/39 sources as potential protocluster members, or those most likely to be in the redshift range of the protocluster. Based on crossmatching as detailed in Section §2.3, we find most of our protocluster-binned sample has not been previously cataloged as protocluster members, nor are they expected to be members; rather they have a higher likelihood of being members than the DSFGs in low- z and high- z bins. Indeed, only $\sim 5\%$ of the volume contained in the protocluster redshift bin corresponds to the protocluster volume itself, leaving substantial room for our DSFGs to lie in the foreground or background population. Taking an overdensity of $\delta_{\text{rare}} = 10$ ([Casey 2016](#)) and the protocluster comoving volume within the S2CLS SSA22 coverage (assuming a redshift range informed by [Topping et al. 2018](#)) we may only expect $\sim 4 - 7$ of these sources to be protocluster members. Comparing our results to the typical density of $S_{850\mu\text{m}} > 5 \text{ mJy}$ sources in a general field ([Swinbank](#)

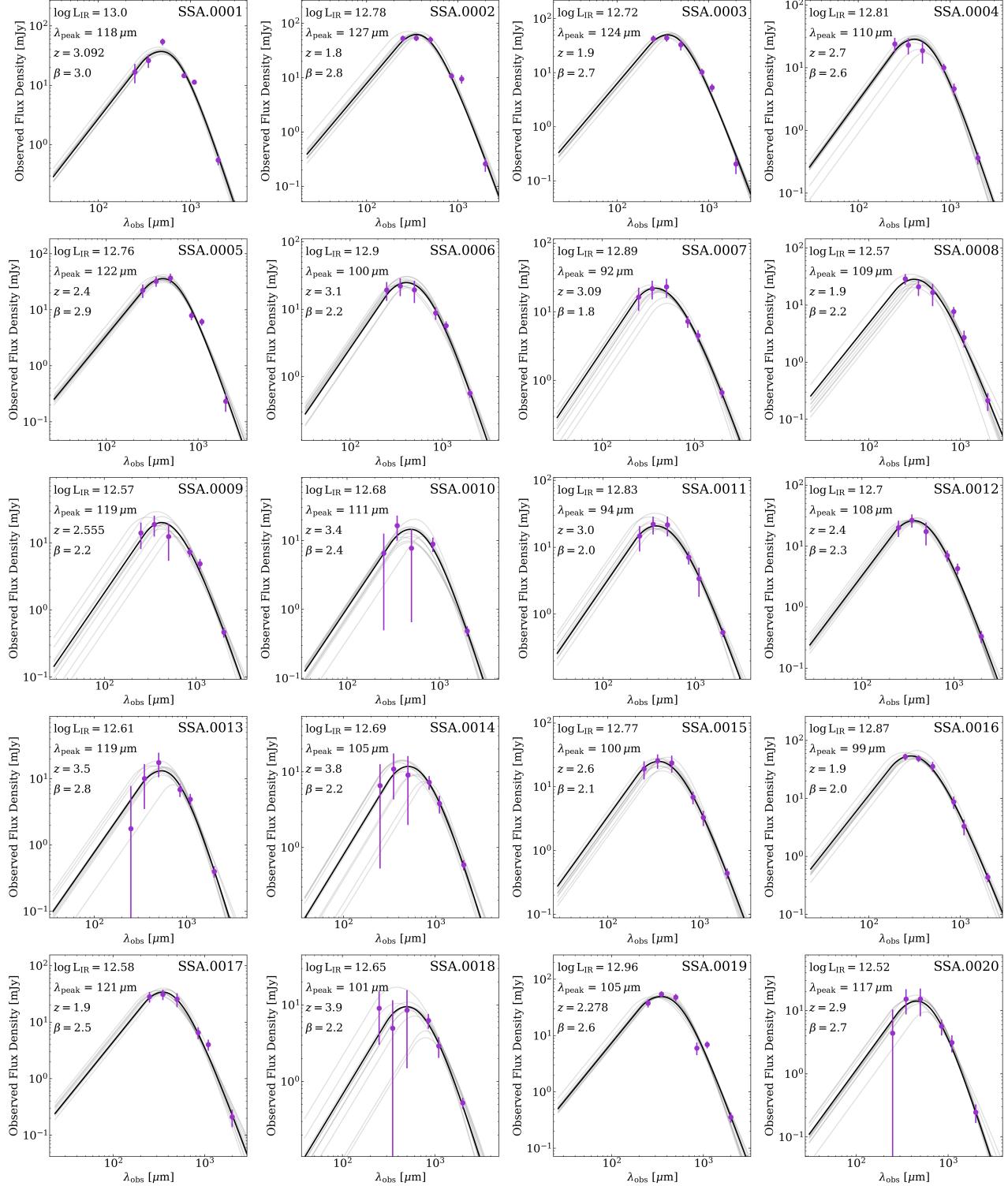
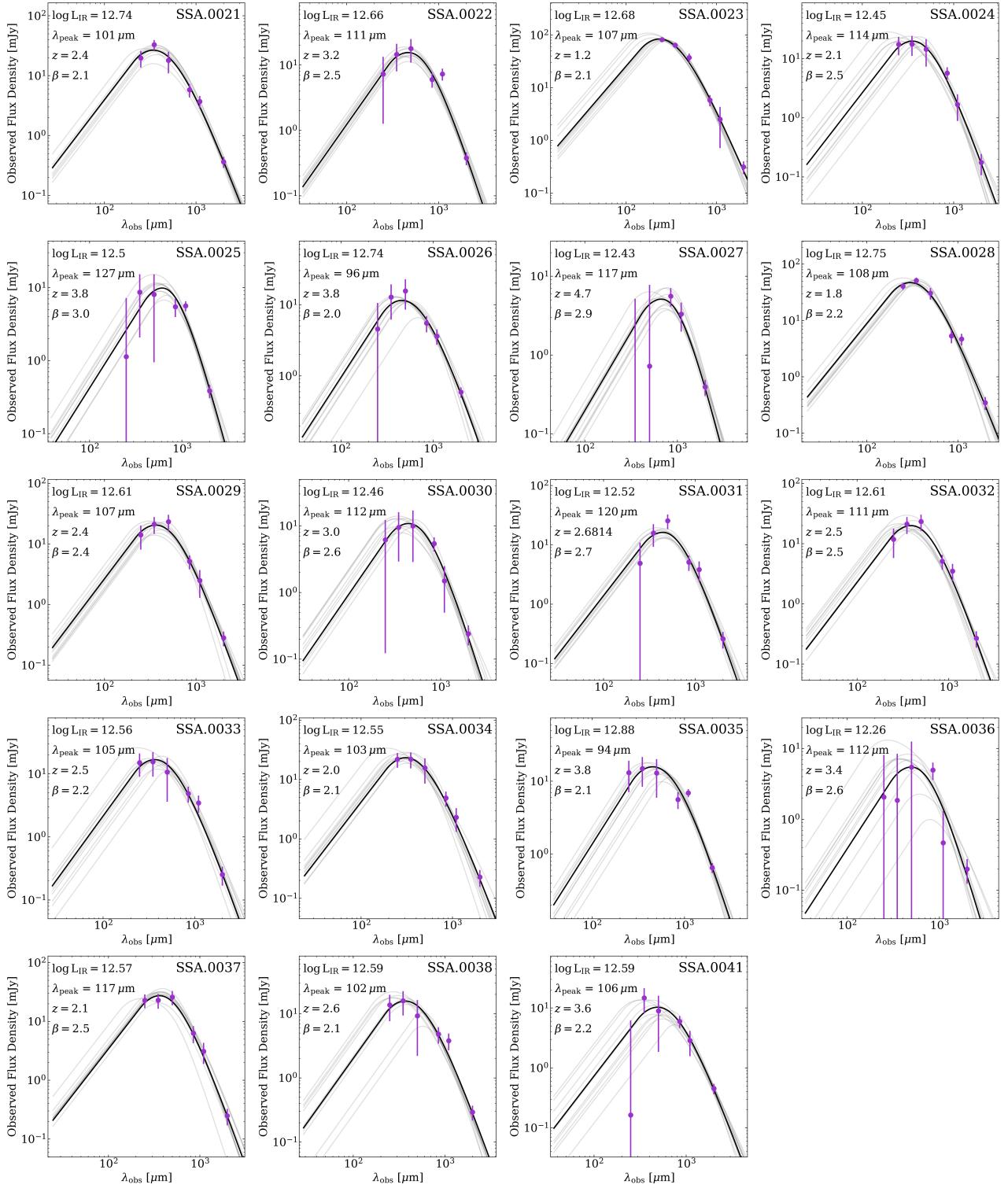


Figure 4: Millimeter photometry (purple points) with the best fit dust SED (black line) for each target in the sample. SEDs are constructed as a modified blackbody with a mid-infrared powerlaw component. For the modified blackbody we assume a general opacity such that $\tau = 1$ at $200 \mu\text{m}$, near the intrinsic peak of the dust SED. The uncertainty of the fit is shown via the gray SED fits, drawn randomly from the successful MCMC trials. Measured SED characteristics are noted for each target in their respective panels and given in Table 2.



— Continuation of Figure 4 —

et al. 2014; Umehata et al. 2015), we find no statistical excess of sources in the protocluster bin; therefore, in the absence of spectroscopic confirmation, we cannot definitively classify these galaxies as protoclusters members. Lastly, we categorize 7/39 sources as high- z , with a maximum redshift solution for an individual source of $z = 4.7^{+3.6}_{-1.7}$ for SSA.0027. For the high- z sample, $\sim 60\%$ of the aggregate probability density distribution is at $z > 4$ (see Figure 2).

While the redshift distribution of galaxies selected at 2 mm is expected to be relatively high with $\langle z \rangle \approx 3.6$ (Casey et al. 2021), the selection of these targets at $850 \mu\text{m}$ leads to a lower median redshift, consistent with what has been previously found for $850 \mu\text{m}$ -selected DSFGs (e.g. Dudzevičiūtė et al. 2020; da Cunha et al. 2021). Therefore, though we expect the follow up 2 mm observations to filter out lower redshift sources (Casey et al. 2021; Zavala et al. 2021, Manning et al. submitted), for this study the redshift distribution is reflective of the $850 \mu\text{m}$ -selection. Previous millimeter studies demonstrate that deeper surveys tend to select for lower redshift DSFGs; in other words brighter DSFGs tend to sit at higher redshifts (e.g. Béthermin et al. 2015). Though we focus on only the brightest subset of the population, we find a redshift distribution consistent with other $850 \mu\text{m}$ -selected SMG samples.

We find a positive correlation between 2 mm flux density and redshift, and sources with lower flux ratios $S_{850 \mu\text{m}}/S_{2 \text{ mm}}$ tend to have higher redshifts (see Figure 5). The loose positive correlation between 2 mm flux density and redshift implies that 2 mm flux density alone does not constrain redshift, but does suggest sources with higher 2 mm flux densities tend to sit at higher redshifts. The relation between FIR/mm flux density and redshift becomes more clear with the flux ratio $S_{850 \mu\text{m}}/S_{2 \text{ mm}}$; for a source with a given $850 \mu\text{m}$ flux density, a relatively higher 2 mm flux density tends to result in a higher redshift solution. This trend holds for sources in our sample with MMpz-derived redshifts as well as spectroscopic redshifts.

Our redshift distribution has good agreement with other $850/870 \mu\text{m}$ -selected surveys, including Chapman et al. (2005), who conduct an early assessment of the $850 \mu\text{m}$ -selected SMG redshift distribution (median $z = 2.2$), and Wardlow et al. (2011), with a similar redshift distribution peaking around $z \sim 2.5$ for their complete, unbiased $870 \mu\text{m}$ -selected sample. Similarly, Danielson et al. (2017) derive a median redshift of $z = 2.4 \pm 0.1$ for the 52 spectroscopically confirmed SMGs in their sample. Danielson et al. (2017) also find the distribution features a high redshift tail, with $\sim 23\%$ of the SMGs at $z \geq 3$; we find a comparable $\sim 35\%$ of the sample

at $z \geq 3$. The recent AS2UDS survey (Dudzevičiūtė et al. 2020) follow-up a comparable sample of $S_{850 \mu\text{m}} > 3.6 \text{ mJy}$ SCUBA-2-selected SMGs in the UKIDSS UDS field with ALMA at $870 \mu\text{m}$. They find a photometric redshift distribution with median $z = 2.61 \pm 0.08$ (1σ range of $z = 1.8\text{--}3.4$), in good agreement with our sample's median redshift of $z = 2.6 \pm 0.2$ (bootstrapped error on the median) and 1σ range of $z = 1.9 - 3.6$. In addition, they find $\sim 6\%$ of their sample at $z > 4$, and only 5/707 ($<< 1\%$) sources at $z < 1$. This is broadly consistent with our smaller, brighter sample, as we find zero sources at $z < 1$, and find one source ($\sim 3 \pm 3\%$ of the sample) formally at $z > 4$. Still, we do find a number of galaxies at $z > 3.8$ (6/39, most of the high- z sample).

We also find consistent results with the ALESS survey (da Cunha et al. 2021), a sample of 99 $870 \mu\text{m}$ -selected SMGs followed up with ALMA 2 mm observations, with a similar flux density limit as AS2UDS of $S_{870} \geq 3.5 \text{ mJy}$. To compare to our sample, we take the subset of 25/99 ALESS sources with $S_{870} \geq 5 \text{ mJy}$ from da Cunha et al. (2021), and find a median redshift and 68% confidence interval of $z_{\text{med}} = 3.2^{+0.4}_{-1.1}$. This is broadly in agreement with our median redshift. Note that the full ALESS sample has a median redshift $z_{\text{med}} = 2.8^{+0.7}_{-0.8}$, and includes fainter sources than our sample selection criteria.

4.2. Utility of 2 mm as a Redshift Filter

A key motivation in this work is to analyze the utility of 2 mm in identifying the highest redshift galaxies among a population that lacks spectroscopy or high quality OIR photometric redshifts. This is a core problem in the study of obscured galaxy populations, where the search for redshifts has been the primary bottleneck for the past 20 years (Casey, Narayanan, & Cooray 2014). For example, following up these sources with spectral scans in the millimeter would take over 100 hours with ALMA (at least ~ 3 hours per source with overheads), versus the total time spent in this paper, with a total on-source time of just 34.12 min for 39 sources (< 1 min per source). Selection at 2 mm has been found to effectively select a higher redshift population of DSFGs, as shown by the blind 2 mm MORA survey in Casey et al. (2021), with $\langle z \rangle = 3.6^{+0.5}_{-0.9}$ for their sample. Our hypothesis is that 2 mm imaging can provide an efficient and targeted method for identifying high- z candidates for follow-up as a zeroth order redshift sorting of $850 \mu\text{m}$ -bright sources (limited to sources that are detected at $850 \mu\text{m}$ by construction, given our selection criteria).

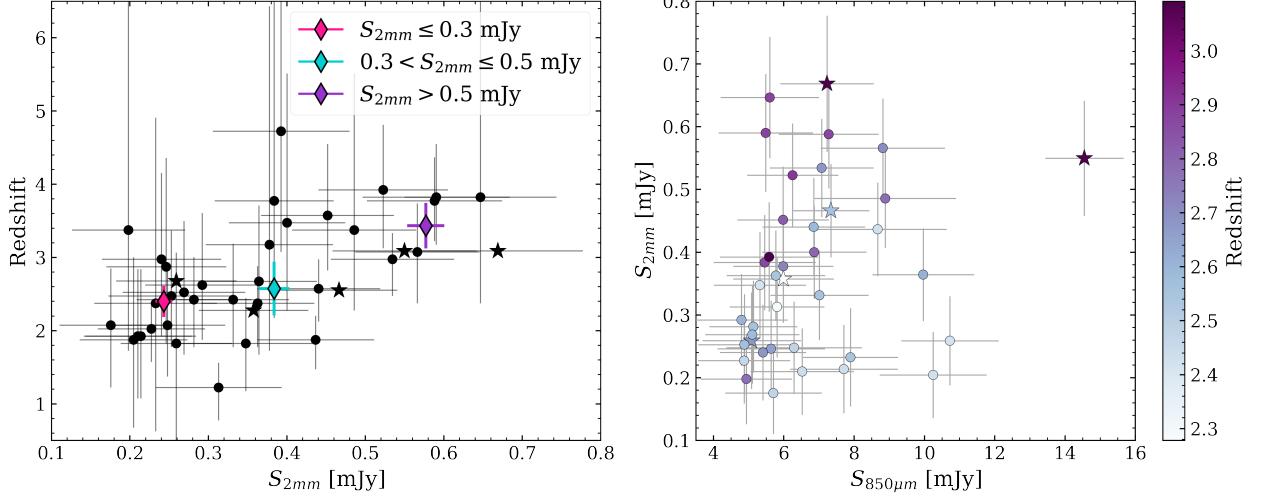


Figure 5: Redshift versus (sub)mm flux density for the $S_{850\mu\text{m}} > 5.55 \text{ mJy}$ sample. *Left:* Flux density $S_{2\text{mm}}$ versus redshift showing some positive correlation, with sources with higher $S_{2\text{mm}}$ tending to have higher redshift solutions. The stars indicate sources which have spectroscopic redshifts, otherwise **MMpz**-derived redshifts are shown. The average values within given flux density bins are shown as colored diamonds. *Right:* Flux densities $S_{2\text{mm}}$ versus $S_{850\mu\text{m}}$ with point color indicating redshift as shown by the colorbar on the right. As before, the stars denote spec- z 's while the circles are photo- z 's. Sources with lower flux ratios $S_{850\mu\text{m}}/S_{2\text{mm}}$ tend to sit at higher redshifts.

Definitively assessing the utility of 2 mm imaging for redshift filtering requires spectroscopic redshifts, but in the absence of spectroscopic data we can evaluate the impact of 2 mm data on our derivation of redshifts relative to the other bands for which we have data. To derive a photo- z , **MMpz** uses all flux densities for which galaxies have measurements; is our redshift solution primarily driven by 2 mm flux density, which would render the correlation of 2 mm flux density and redshift in Figure 5 unsurprising?

To evaluate the set of photometry that has the most influence on **MMpz** redshift solutions, we generate simulated galaxy FIR/mm photometry. We specifically test the sensitivity of **MMpz** on the addition or absence of different bands by generating fake SEDs using the same methodology as [Casey \(2020\)](#) where sources of different fixed redshifts are assigned realistic flux densities with associated uncertainties matched to those of this dataset. We then fit redshifts with **MMpz** for subsets of the data including the full data set: SPIRE 250 μm /350 μm /500 μm , SCUBA-2 850 μm , AzTEC 1.1 mm, ALMA 2 mm; just SCUBA-2 850 μm and ALMA 2 mm; and just SPIRE 250 μm /350 μm /500 μm and SCUBA-2 850 μm . [Casey \(2020\)](#) notes that using only 850 μm and 2 mm photometry – two points that should exclusively probe the Rayleigh-Jeans tail for most redshifts and temperature – is insufficient. SPIRE 250 μm /350 μm /500 μm data is needed to constrain the

dust peak for $z \lesssim 6$, where even non-detections provide very useful (though loose) constraints on the SED peak.

From these simulated SEDs, the addition of the 2 mm point – provided $\sim 500 \mu\text{m}$ constraints exist – improves accuracy from $\Delta z/(1+z) = 0.3$ (without the 2 mm data) to $\Delta z/(1+z) = 0.2$ (with the 2 mm data), in particular for galaxies at $z \gtrsim 3.5$. This demonstrates that given *Herschel* data, the 2 mm data is the next most impactful data for improved accuracy of the redshift solution, especially for high redshift sources. For example, the probability of a hypothetical source with $z_{\text{MMpz}} = 4.1$ lying at $z > 3.5$ increases from 81% without the 2 mm constraint, to 96% with the 2 mm constraint.

While 2 mm data does improve the accuracy of redshift constraints, they do not significantly change the precision of those redshift solutions. Quantitatively, the average breadth of redshift PDFs for our sample is $\delta z = 1.1 - 1.2$ with or without the 2 mm data folded into the photometry. This is not unexpected, as the breadth of the redshift PDF generated from millimeter data is dominated by the intrinsic spread in the $L_{\text{IR}}-\lambda_{\text{peak}}$ relation (discussed further in [Casey 2020](#)).

4.3. Emissivity Spectral Index β

With 2 mm flux densities in hand for a large sample of bright 850 μm -selected DSFGs, we have a unique dataset with which we analyze the galaxy-integrated slope of the Rayleigh-Jeans tail of blackbody emission. The slope is governed by the physical quantity β , the emissivity spec-

Name	z_{MMPz}	z bin	z_{OIR}	z_{spec}	β	$\log L_{\text{IR}}$ $\log_{10}(L_{\odot})$	SFR $M_{\odot} \text{ yr}^{-1}$	λ_{peak} μm	T_{dust}^* K
SSA.0001	$3.3^{+2.8}_{-1.6}$	Protocluster	$2.85^{+3.15}_{-0.98}$	3.092 ^a	$3.0^{+0.2}_{-0.2}$	$13.00^{+0.06}_{-0.07}$	1480^{+230}_{-210}	118^{+7}_{-7}	41^{+3}_{-3}
SSA.0002	$1.8^{+1.2}_{-1.4}$	Low- z			$2.8^{+0.3}_{-0.2}$	$12.78^{+0.05}_{-0.05}$	890^{+110}_{-90}	127^{+7}_{-8}	36^{+4}_{-3}
SSA.0003	$1.9^{+1.1}_{-1.2}$	Low- z			$2.6^{+0.3}_{-0.2}$	$12.71^{+0.06}_{-0.06}$	770^{+110}_{-100}	124^{+9}_{-10}	37^{+4}_{-4}
SSA.0004	$2.5^{+1.2}_{-0.9}$	Protocluster			$3.42^{+0.06}_{-0.11}$	$12.81^{+0.09}_{-0.10}$	960^{+230}_{-200}	110^{+13}_{-12}	44^{+6}_{-6}
SSA.0005	$2.4^{+2.5}_{-1.8}$	Low- z	$3^{+1.27}_{-1.67}$		$3.0^{+0.3}_{-0.3}$	$12.76^{+0.07}_{-0.07}$	860^{+150}_{-130}	122^{+9}_{-9}	39^{+4}_{-4}
SSA.0006	$3.1^{+0.7}_{-0.7}$	Protocluster			$2.2^{+0.3}_{-0.3}$	$12.90^{+0.09}_{-0.11}$	1200^{+300}_{-300}	100^{+12}_{-11}	49^{+7}_{-6}
SSA.0007	$3.2^{+0.5}_{-0.5}$	Protocluster	$3.7^{+0.72}_{-0.19}$	3.0854 ^b	$1.9^{+0.3}_{-0.2}$	$12.89^{+0.11}_{-0.12}$	1200^{+300}_{-300}	92^{+11}_{-10}	53^{+7}_{-7}
SSA.0008	$1.9^{+0.9}_{-0.9}$	Low- z	$2.2^{+0.06}_{-0.05}$		$2.2^{+0.4}_{-0.4}$	$12.57^{+0.11}_{-0.12}$	550^{+200}_{-130}	109^{+18}_{-18}	44^{+11}_{-8}
SSA.0009	$3.1^{+0.9}_{-0.7}$	Protocluster	$2.2^{+0.72}_{-0.32}$	2.555 ^c	$2.2^{+0.3}_{-0.3}$	$12.57^{+0.12}_{-0.15}$	550^{+180}_{-160}	119^{+16}_{-13}	38^{+6}_{-7}
SSA.0010	$3.5^{+2.0}_{-1.2}$	Protocluster			$2.4^{+0.6}_{-0.5}$	$12.68^{+0.16}_{-0.20}$	700^{+300}_{-300}	111^{+22}_{-19}	43^{+10}_{-10}
SSA.0011	$3.0^{+0.4}_{-0.5}$	Protocluster			$2.0^{+0.3}_{-0.3}$	$12.83^{+0.11}_{-0.13}$	1000^{+300}_{-300}	94^{+13}_{-11}	52^{+8}_{-7}
SSA.0012	$2.4^{+0.8}_{-0.7}$	Low- z	$6^{+0}_{-4.06}$		$2.3^{+0.4}_{-0.3}$	$12.70^{+0.10}_{-0.10}$	750^{+200}_{-160}	108^{+13}_{-13}	44^{+7}_{-6}
SSA.0013	$3.5^{+2.0}_{-1.2}$	Protocluster	$2.55^{+3.45}_{-0.69}$		$2.8^{+0.4}_{-0.4}$	$12.61^{+0.15}_{-0.19}$	610^{+260}_{-220}	119^{+19}_{-16}	40^{+8}_{-8}
SSA.0014	$3.8^{+0.6}_{-0.6}$	High- z			$2.2^{+0.5}_{-0.4}$	$12.69^{+0.19}_{-0.25}$	730^{+400}_{-320}	105^{+25}_{-18}	46^{+11}_{-12}
SSA.0015	$2.6^{+0.4}_{-0.5}$	Protocluster			$2.1^{+0.4}_{-0.3}$	$12.78^{+0.10}_{-0.11}$	880^{+230}_{-190}	100^{+12}_{-12}	49^{+8}_{-6}
SSA.0016	$1.9^{+0.3}_{-0.4}$	Low- z			$2.0^{+0.3}_{-0.2}$	$12.87^{+0.09}_{-0.08}$	1100^{+300}_{-200}	99^{+11}_{-13}	49^{+8}_{-6}
SSA.0017	$1.9^{+0.9}_{-0.9}$	Low- z	$2.25^{+0.2}_{-0.4}$		$2.5^{+0.3}_{-0.3}$	$12.58^{+0.09}_{-0.08}$	570^{+130}_{-100}	121^{+12}_{-13}	38^{+6}_{-5}
SSA.0018	$3.9^{+0.9}_{-0.8}$	High- z	$2.85^{+2.7}_{-0.69}$		$2.1^{+0.5}_{-0.4}$	$12.64^{+0.23}_{-0.28}$	700^{+400}_{-300}	101^{+27}_{-20}	48^{+14}_{-13}
SSA.0019	$2.1^{+1.5}_{-1.5}$	Low- z	$2.15^{+0.23}_{-0.21}$	2.278 ^d	$2.8^{+0.3}_{-0.2}$	$13.0^{+0.06}_{-0.06}$	1400^{+200}_{-200}	105^{+8}_{-8}	47^{+5}_{-4}
SSA.0020	$2.9^{+1.5}_{-1.1}$	Protocluster			$2.7^{+0.5}_{-0.5}$	$12.52^{+0.15}_{-0.18}$	490^{+200}_{-170}	117^{+17}_{-18}	41^{+9}_{-8}
SSA.0021	$2.3^{+0.6}_{-0.6}$	Low- z	$3.3^{+0.19}_{-0.21}$		$2.2^{+0.3}_{-0.3}$	$12.74^{+0.10}_{-0.10}$	820^{+220}_{-170}	101^{+12}_{-13}	48^{+8}_{-6}
SSA.0022	$3.2^{+3.3}_{-1.5}$	Protocluster			$2.5^{+0.4}_{-0.4}$	$12.66^{+0.15}_{-0.17}$	690^{+280}_{-230}	111^{+18}_{-16}	44^{+9}_{-8}
SSA.0023	$1.2^{+0.3}_{-0.5}$	Low- z			$2.1^{+0.3}_{-0.3}$	$12.68^{+0.14}_{-0.10}$	710^{+280}_{-150}	107^{+14}_{-16}	44^{+10}_{-6}
SSA.0024	$2.1^{+0.8}_{-0.9}$	Low- z			$2.5^{+0.5}_{-0.5}$	$12.45^{+0.16}_{-0.15}$	420^{+180}_{-120}	114^{+20}_{-23}	42^{+13}_{-9}
SSA.0025	$3.8^{+3.5}_{-1.6}$	High- z			$3.0^{+0.4}_{-0.5}$	$12.50^{+0.21}_{-0.21}$	470^{+280}_{-190}	127^{+20}_{-22}	36^{+10}_{-9}
SSA.0026	$3.8^{+0.7}_{-0.7}$	High- z			$2.0^{+0.4}_{-0.3}$	$12.74^{+0.18}_{-0.24}$	810^{+410}_{-340}	96^{+20}_{-15}	51^{+11}_{-11}
SSA.0027	$4.7^{+3.6}_{-1.7}$	High- z			$2.9^{+0.4}_{-0.6}$	$12.43^{+0.22}_{-0.24}$	410^{+270}_{-170}	117^{+20}_{-23}	41^{+12}_{-9}
SSA.0028	$1.8^{+0.6}_{-0.7}$	Low- z	$0^{+0.01}_{-0} \dagger$		$2.2^{+0.3}_{-0.3}$	$12.75^{+0.08}_{-0.07}$	830^{+170}_{-120}	108^{+10}_{-12}	45^{+6}_{-5}
SSA.0029	$2.4^{+0.6}_{-0.6}$	Low- z			$2.6^{+0.4}_{-0.4}$	$12.61^{+0.12}_{-0.13}$	610^{+200}_{-160}	107^{+15}_{-16}	45^{+9}_{-7}
SSA.0030	$3.0^{+1.2}_{-1.1}$	Protocluster			$2.6^{+0.5}_{-0.6}$	$12.46^{+0.20}_{-0.26}$	430^{+260}_{-190}	112^{+25}_{-23}	43^{+13}_{-12}
SSA.0031	$2.8^{+2.8}_{-1.5}$	Protocluster		2.6814 ^e	$2.7^{+0.5}_{-0.5}$	$12.56^{+0.13}_{-0.15}$	540^{+190}_{-160}	116^{+14}_{-16}	41^{+8}_{-7}
SSA.0032	$2.5^{+1.0}_{-0.8}$	Protocluster	$1.75^{+4.25}_{-0.82}$		$2.7^{+0.4}_{-0.4}$	$12.61^{+0.12}_{-0.13}$	610^{+200}_{-160}	111^{+15}_{-16}	44^{+9}_{-7}
SSA.0033	$2.5^{+0.8}_{-0.8}$	Protocluster			$2.1^{+0.5}_{-0.4}$	$12.56^{+0.16}_{-0.18}$	540^{+250}_{-180}	105^{+20}_{-19}	46^{+12}_{-10}
SSA.0034	$2.0^{+0.4}_{-0.5}$	Low- z	$2.05^{+0.32}_{-0.2}$		$2.1^{+0.4}_{-0.4}$	$12.55^{+0.15}_{-0.13}$	530^{+220}_{-140}	103^{+17}_{-19}	47^{+12}_{-8}
SSA.0035	$3.8^{+2.8}_{-1.5}$	High- z			$2.1^{+0.4}_{-0.4}$	$12.88^{+0.16}_{-0.23}$	1100^{+500}_{-500}	94^{+23}_{-14}	52^{+10}_{-12}
SSA.0036	3^{+5}_{-2}	High- z			$2.6^{+0.6}_{-0.7}$	$12.3^{+0.3}_{-0.4}$	270^{+270}_{-160}	110^{+30}_{-30}	43^{+18}_{-14}
SSA.0037	$2.1^{+0.7}_{-0.7}$	Low- z			$2.5^{+0.5}_{-0.4}$	$12.57^{+0.11}_{-0.10}$	550^{+150}_{-110}	117^{+15}_{-19}	40^{+9}_{-7}
SSA.0038	$2.6^{+1.0}_{-0.8}$	Protocluster			$2.1^{+0.5}_{-0.4}$	$12.59^{+0.16}_{-0.18}$	570^{+260}_{-200}	102^{+19}_{-18}	47^{+12}_{-10}
SSA.0041	$3.6^{+1.0}_{-0.8}$	Protocluster			$2.2^{+0.5}_{-0.4}$	$12.59^{+0.19}_{-0.26}$	600^{+300}_{-300}	106^{+25}_{-18}	45^{+11}_{-12}

Table 2: Measured and derived characteristics for each galaxy in our sample. Photometric redshift results from MMPz are listed in z_{MMPz} , with OIR photometric redshifts quoted as given in Umehata et al. (2014) under z_{OIR} , and any literature spec- z values listed in z_{spec} with the reference (a: Umehata et al. (2014), b: Kubo et al. (2016), c: Chapman et al. (2005), d: Alaghband-Zadeh et al. (2012), e: Cooper et al. in prep). Measured from the best fit SEDs (see Figure 4) are emissivity spectral index β , IR luminosity $\log L_{\text{IR}}$, dust temperature T_{dust} , and peak wavelength λ_{peak} . The star formation rates (SFR) are derived from L_{IR} using the Kennicutt & Evans (2012) scaling. *Note that we assume a general opacity such that $\tau = 1$ at 200 μm , therefore these estimates are only comparable to other works using a similar model. †Reported z_{OIR} from Umehata et al. (2014), which we note is inconsistent with both our FIR results and OIR archival data (see discussion of the source in §3.1).

tral index, or the frequency dependence of dust grain emissivity per unit mass. Still, drawing a physical interpretation for a measured β is not possible for spatially unresolved high redshift galaxies due to the complexity and heterogeneity of the ISM. This is especially true for sources lacking secure redshifts and with somewhat limited FIR/mm photometry.

Nevertheless, our data can be used to compare galaxy-integrated Rayleigh-Jeans slopes to other high- z samples and commonly adopted literature values of β for high- z datasets of similar quality. Here, our finding of $\langle\beta\rangle = 2.4^{+0.4}_{-0.3}$ suggests that the distribution of integrated β indices skews high, in line with other recent works, including a well-studied galaxy in SSA22 with $\beta = 2.3$ (Kato et al. 2018) as well as other samples with dust continuum data at $\lambda_{\text{obs}} \gtrsim 2$ mm (Jin et al. 2019; Casey et al. 2021). The aggregate best fit SED results for our sample (see Table 2) shows relatively steep β values ranging from $1.8 < \beta < 3.0$, all steeper than the standard value often adopted in the literature, typically $\beta = 1.8$ (e.g. Scoville et al. 2016), justified by measurements of the β from Milky Way's ISM (e.g. Paradis et al. 2009; Planck Collaboration et al. 2011). Nevertheless, while our median β skews high at $\langle\beta\rangle = 2.4^{+0.4}_{-0.3}$, it remains consistent with theoretical predictions for interstellar dust models (e.g. Draine & Lee 1984; Köhler et al. 2015), which predict $1 < \beta < 2.5$ depending on grain composition.

A steeper β could result from a fundamental difference in grain composition or size (Chihara et al. 2001; Mutschke et al. 2013; Inoue et al. 2020), although testing this would require probing the ISM with high spatial resolution and sensitivity. Many of the caveats influencing observed β measurements – including geometric effects, variations in optical depth, and temperature distributions – would flatten rather than steepen the SED at long wavelengths.

Figure 6 indicates the millimeter color distribution for the sample, represented by observed flux density ratio $S_{850\mu\text{m}}/S_{2\text{mm}}$. We compare the millimeter colors to SED tracks constructed from a modified blackbody spanning $1.8 < \beta < 3.0$ and $90\mu\text{m} < \lambda_{\text{peak}} < 120\mu\text{m}$. The sources follow a general trend of $S_{850\mu\text{m}}/S_{2\text{mm}}$ with redshift aligned with the shape of the model SEDs. Though this trend may be due to the use of millimeter photometry to derive redshifts, we note that MMpz fixes $\beta = 1.8$ and that modifications of β do not significantly impact the output redshift probability density distributions. The majority of sources have flux density ratios suggestive of high $\beta > 1.8$, with the sample largely consistent with SED tracks for $\beta \approx 2.0 - 2.5$ (cyan and purple tracks in Figure 6). This aggregate high β result suggests that steeper Rayleigh-Jeans slopes, i.e. higher

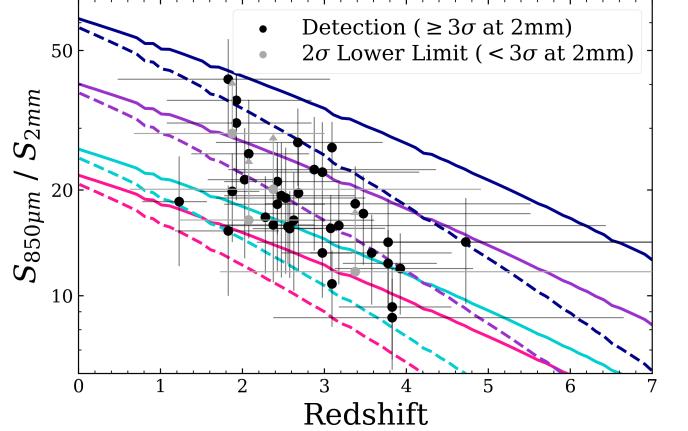


Figure 6: Flux density ratio $S_{850\mu\text{m}}/S_{2\text{mm}}$ versus redshift for the full sample; while $850\mu\text{m}$ flux densities are SCUBA-2 based, they have been deboosted for confusion noise and the errors here reflect that additional source of uncertainty. Redshifts are mm-derived from our analysis in all but five cases where a spec- z is available. Detected targets at 2 mm are denoted as secure detections ($\geq 3\sigma$ at 2 mm, black points) or plotted as 2σ lower limits for non-detections (detected at $< 3\sigma$ at 2 mm, in gray). Overlaid as colored lines are SED tracks, each a modified blackbody with mid-infrared power law for a given combination of emissivity spectral index ($\beta = 1.8, 2.0, 2.5$, and 3.0 in pink, cyan, purple, and navy, respectively) and dust temperature (given here as the observable, λ_{peak}), with solid lines for $\lambda_{\text{peak}} = 90\mu\text{m}$ and dashed for $\lambda_{\text{peak}} = 120\mu\text{m}$. Our sample is largely consistent with steeper slopes consistent with SED tracks for $\beta \approx 2.0 - 2.5$, with some variation (the median β measured is 2.4).

β , should likely be applied for similar high- z DSFGs in the literature in future works.

To test the impact of SED fitting approach on our results, we fit our SEDs using isothermal modified blackbody models as in da Cunha et al. (2021). Given the known degeneracies and bias from opacity assumptions, we toggled λ_0 over multiple fits. For $\lambda_0 = 200\mu\text{m}$ (same assumption as in our modelling), we find a median of $\beta = 2.1 \pm 0.5$, statistically consistent within error with our previous result. Note that while in both SED fitting routines we assume isothermal dust, a primary difference between methods is the Drew & Casey (submitted) model accounts for the range in temperature on Wien's portion of the SED with a power-law, whereas da Cunha et al. (2021) do not attempt to fit any points below $70\mu\text{m}$ rest-frame with their simple single-component model. The discrepancy in results from these two methods demonstrates that care should be taken when di-

rectly comparing emissivity spectral indices obtained using different fitting methods. To this point, using identical fitting methods we do find a higher median emissivity spectral index for our sample ($\beta = 2.1 \pm 0.5$) than the ALESS sample ($\beta = 1.9 \pm 0.4$) [da Cunha et al. \(2021\)](#).

Though we have accounted for confusion noise as best as possible by using deboosted flux densities from SCUBA-2, measuring a higher β due to combining single dish $850\mu\text{m}$ data with interferometric ALMA 2 mm data is a concern. Still, [da Cunha et al. \(2021\)](#) use multiwavelength ALMA data and demonstrate $\beta \approx 2$ for high- z DSFGs, with a median $\beta = 1.9 \pm 0.4$ derived for their sample. It is possible the bias from single dish data could be responsible for the difference between a $\beta \approx 2$ result and our higher $\beta \approx 2.4$ measurement. Large samples with redshifts and all ALMA data on the Rayleigh-Jeans tail are needed to take better measurements with respect to confusion noise and multiplicity in detected sources.

4.4. CSFRD Contribution

We determine the sample's estimated contribution to the cosmic star formation rate density (CSFRD), as shown in Figure 7, to compare with other similar estimates of DSFG samples in the literature. While we do not expect the presence of the protocluster to impact our measurements, we verify this by analyzing the relative density of our sources compared to field samples. CSFRD estimates are found using a 100-trial Monte Carlo (MC) to sample the redshift and SFR probability density distributions for each of the 39 SSA22 sources. We calculate a comoving volume over the survey area (S2CLS coverage of SSA22 spans 0.28 deg^2) and redshift range based on discrete redshift bins (from $1 < z < 5$ with width $\Delta z = 0.5$), then divide the total SFR of all DSFGs in our sample by that volume when their draws fall in the given z -bin. The total CSFRD is measured for each MC trial, and from these measurements we take the mean and standard deviation for all trials to find the estimate for the CSFRD contribution from $1 < z < 5$ for this $S_{850\mu\text{m}} > 5.55\text{ mJy}$ sample.

Given the overall rarity of $S_{850\mu\text{m}} > 5.55\text{ mJy}$ sources among DSFGs, we estimate the impact of cosmic variance, and of incompleteness due to shallower S2CLS data of SSA22 relative to other fields like the UDS ([Geach et al. 2017](#)) and COSMOS ([Simpson et al. 2019](#)) SCUBA-2 maps. We conduct MC trials to count the number of sources in randomly placed SSA22-sized areas (0.28 deg^2) over UDS (0.74 deg^2) and COSMOS (1.94 deg^2); the 1σ $850\mu\text{m}$ depths are 0.9 mJy for UDS, 1.2 mJy for COSMOS, and 1.2 mJy for SSA22. We first enforce the same criteria used to select our SSA22

sources for the COSMOS/UDS sources: $S_{850\mu\text{m}} > 5\text{ mJy}$ and $> 5\sigma$. A first round of MC trials is run directly from the COSMOS/UDS catalogs for sources that meet these criteria. In this round, we find an average of 62^{+24}_{-30} sources/SSA22-sized field. This exceeds the 39 observed sources in SSA22 by a factor of $1.6^{+0.6}_{-0.8}$; though suggestive of an underdensity of bright DSFGs in SSA22, this does not account for the higher rms noise in SSA22 relative to UDS/COSMOS. Thus this factor is more representative of incompleteness. For the next round of trials, we artificially boost the rms of the UDS/COSMOS maps using the rms distribution of the SSA22 map to simulate the shallower SSA22 map. We do this by drawing rms values from the SSA22 map at random positions. After this elevation of noise for the COSMOS/UDS samples, we find an expectation value of 44^{+13}_{-14} sources/SSA22-sized field fulfilling our bright DSFG criteria. This is in line with our observed number of 39 DSFGs.

From this sampling of UDS/COSMOS maps, we determine that our SSA22 sample is not cosmologically over- or under-dense, but it is incomplete by a factor of $1.6^{+0.6}_{-0.8}$. While this is consistent with being complete, it does reveal a systematic offset that impacts our results based on the depth of SSA22 SCUBA-2 coverage, thus we adjust our CSFRD measurement accordingly by multiplying by the derived incompleteness factor.

We also run similar MC trials on simulated data to compare to our observational results. With the semiempirical model from [Popping et al. \(2020\)](#), we count the number of $S_{850\mu\text{m}} > 5\text{ mJy}$ sources in randomly placed SSA22-sized areas (0.28 deg^2), and find 61^{+8}_{-15} sources/SSA22-sized field. From similar MC trials on SHARK light cone results ([Lagos et al. 2020](#)) we find 48^{+12}_{-9} sources/SSA22-sized field. These results are consistent with the observational results from UDS & COSMOS S2CLS, supporting our estimates of incompleteness and cosmic variance for our sample from the real data.

The resulting CSFRD estimate for our sample is shown in Figure 7. This estimate is compared to the expected contribution by obscured galaxies to the CSFRD drawn from the IR luminosity function (IRLF) constraints in [Zavala et al. \(2021\)](#). We compare directly to the expected CSFRD for $S_{850\mu\text{m}} > 5.55\text{ mJy}$ from their work, corresponding to the $S_{850\mu\text{m}} > 5\text{ mJy}$ and $> 5\sigma$ selection for our sample at the 1σ depth ($\sim 1.2\text{ mJy}$) of the S2CLS data for SSA22. As shown in Figure 7, our incompleteness-corrected results are consistent with the [Zavala et al. \(2021\)](#) IRLF model within uncertainties.

Although we enforce a higher flux density cutoff than their sample, CSFRD estimates for our sample are com-

parable to those for the AS2UDS sample (Dudzevičiūtė et al. 2020). For our sample, we find the contribution to the CSFRD peaks around $z \sim 2.4$, slightly higher redshift than the total CSFRD peak at $z \sim 2$ from Madau & Dickinson (2014). Our $850\mu\text{m}$ -bright, obscured sources contribute $\sim 10\%$ (ranging 8-13%) to the cosmic-averaged CSFRD from $2 < z < 5$.

Using a similar MC trials technique, we also compute volume density. For a 100-trail MC, we find the volume density of DSFGs above $z \sim 4$ to be $6.3 \times 10^{-7} \text{ Mpc}^{-3}$ based on our sample. Note that this volume density is very specifically linked to our flux density cutoff, as different cuts will result in more or fewer sources (Long et al. in prep).

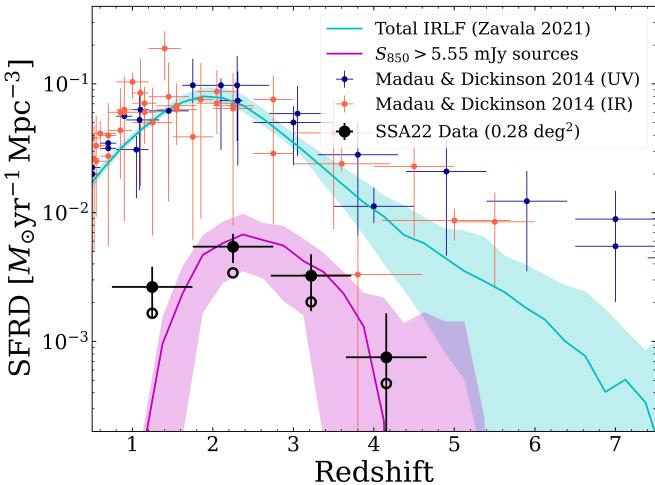


Figure 7: CSFRD as measured at rest-frame UV (blue) and IR-mm (orange) from Madau & Dickinson (2014). While rest-frame UV measurements now reach $z > 10$ thanks to deep HST NIR imaging campaigns, in contrast, surveys of obscured emission in galaxies remain uncertain at $z > 5$. The black points show the data, which lacks precision to constrain the CSFRD at high- z . The cyan region shows the IRLF from Zavala et al. (2021), and the magenta region shows the expected IRLF for sources with $S_{850\mu\text{m}} > 5.55 \text{ mJy}$, corresponding to the $S_{850\mu\text{m}} > 5 \text{ mJy}$ and $> 5\sigma$ selection for our sample at the SSA22 S2CLS depth. The filled black points show the CSFRD estimates from our SSA22 data corrected for incompleteness, while the open circles are the same data without correcting for incompleteness.

5. SUMMARY

In this paper we present ALMA 2 mm imaging for a complete sample of 39 SCUBA-2 detected DSFGs in SSA22 selected for $S_{850\mu\text{m}} > 5 \text{ mJy}$ at $> 5\sigma$ in S2CLS.

We detect 35/39 sources at $S_{2\text{mm}} > 3\sigma$, where our sensitivity is $0.08 \text{ mJy beam}^{-1}$ on average. With multi-wavelength (sub)millimeter data for the sample from *Herschel*/SPIRE, SCUBA-2, AzTEC, and ALMA, we characterize IR SEDs and measure and derive properties including IR luminosity, star formation rate, and emissivity spectral index. For each galaxy, we also estimate a millimeter photometric redshift with MMpz. Our main results are as follows:

- Based on our photometric redshifts and literature spectroscopic redshifts, we find a redshift distribution consistent with other $850\mu\text{m}$ -selected SMG samples, with $z_{\text{med}} = 2.6 \pm 0.7$ despite these sources being representative of only the brightest subset of the $850\mu\text{m}$ -selected population (often hypothesized to sit at higher redshifts). We also categorize 7/39 sources as high- z (all with redshift solutions $z \gtrsim 3.5$), to be spectroscopically confirmed in future follow-up observations.
- Provided the existence of $250 - 500\mu\text{m}$ photometry (i.e. from SPIRE) that brackets the dust SED peak for the majority of DSFGs, we find that the 2 mm photometric constraints are the next most impactful for refining redshift solutions for these sources. With available redshift constraints, we generally find a positive correlation between redshift and 2 mm flux density. This 2 mm data point is especially useful for high redshift sources at $z > 3.5$, where the addition of the 2 mm point improves accuracy from $\Delta z/(1+z) = 0.3$ (without the 2 mm data) to $\Delta z/(1+z) = 0.2$ (with the 2 mm data).
- Our sample has broadly steep emissivity spectral indices with median $\beta = 2.4 \pm 0.3$. Still, while the aggregate β skews high, measurements for any individual β would be improved with matched beam ALMA data at both frequencies, with respect to confusion noise and multiplicity in detected sources.
- For our sample of $850\mu\text{m}$ -bright sources, we estimate the contribution to the cosmic-averaged CSFRD is $\sim 10\%$ (ranging 8-13%).

Our study employs 2 mm imaging to filter out lower redshift DSFGs, advancing efforts to take a complete census of DSFGs at early epochs. While we do find that sources with higher 2 mm flux densities tend to sit at higher redshifts, spectroscopic follow up is needed to verify the efficacy of the 2 mm redshift filtering technique. Future 2 mm follow up of large samples of DSFGs will

be a useful tool to improve the efficiency of identifying the highest redshift obscured galaxies.

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This research made use of the following software: spyder (Raybaut 2009), astropy (Astropy Collaboration et al. 2018), Matplotlib (Hunter 2007), NumPy (van der Walt et al. 2011), SymPy (Meurer et al. 2017), SciPy (Virtanen et al. 2020), pandas (Wes McKinney 2010; pandas development team 2020).

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