

QUOKKA: A code for two-moment AMR radiation hydrodynamics on GPUs

Benjamin D. Wibking^{1*} and Mark R. Krumholz¹

¹*Research School of Astronomy & Astrophysics, Mount Stromlo Observatory, Cotter Road, Weston Creek, ACT 2611 Australia*

Accepted XXX. Received YYY; in original form ZZZ

ABSTRACT

We present a new subcycling-in-time, block-structured adaptive mesh refinement (AMR) radiation hydrodynamics code for graphics processing units (GPUs). The equations of hydrodynamics are solved with the piecewise parabolic method (PPM) in a method-of-lines formulation and the radiative transfer is solved via the variable Eddington tensor (VET) radiation moment equations with a local closure. We combine explicit-in-time evolution of the radiation moment equations with the reduced speed-of-light approximation. We show results for a wide range of test problems for hydrodynamics, radiation, and coupled radiation hydrodynamics. On uniform grids in 3D, we achieve a peak of 69 million hydrodynamic updates per second per GPU. For radiation hydrodynamics problems on uniform grids in 3D, our code scales from 4 GPUs to 256 GPUs with an efficiency of > 93 per cent. The code is publicly released under an open-source license on [GitHub](#).

Key words: radiation hydrodynamics – numerical methods

1 INTRODUCTION

1.1 Methods for radiation-hydrodynamics

In many astrophysical systems, the radiation field carries a substantial portion of the total momentum and energy budget, and therefore must be included in any numerical simulation. However, simulating such systems, particularly at high resolution, presents a fundamental challenge in both physics and numerics. Part of this challenge is dimensionality: in full generality the radiation field is governed by the time-dependent equation of radiative transfer,

$$\frac{1}{c} \frac{\partial}{\partial t} I_{\nu} + \hat{\mathbf{n}} \cdot \nabla I_{\nu} = \eta_{\nu} - \kappa_{\nu}^{(\text{tot})} I_{\nu} + \frac{1}{4\pi} \int \kappa_{\nu}^{(\text{sca})} I_{\nu} d\Omega, \quad (1)$$

where I_{ν} is the radiation intensity travelling in the direction specified by the unit vector $\hat{\mathbf{n}}$, η_{ν} is the matter emissivity, and $\kappa_{\nu}^{(\text{tot})}$ and $\kappa_{\nu}^{(\text{sca})}$ are the total and scattering opacities, respectively. This is a time-dependent integro-differential equation with six dimensions: three positions, two angles (parameterised by $\hat{\mathbf{n}}$), and the frequency. Full numerical solution of a problem of this dimensionality, at least if it must be done millions of times to run in tandem with a hydrodynamic (HD) or magnetohydrodynamic (MHD) simulation, remains out of reach for most applications.

Within the astrophysics community, there are two general approaches to this problem. One is characteristic methods, which solve [equation 1](#) (or its time-independent form), but only along rays pointing back to particular sources (e.g., [Abel & Wandelt 2002](#); [Rijkhorst et al. 2006](#); [Krumholz et al. 2007b](#)) or rays randomly assigned by Monte Carlo (e.g., [Fleck & Cummings 1971](#); [Tsang & Milosavljević 2015](#)). A second approach, which we will pursue here, is moment methods (e.g., [Mihalas & Mihalas 1984](#); [Castor 2004](#)), whereby one

takes moments of the transfer equation, thereby eliminating the angular dimensions of the problem.¹ This results in a series of moment equations that one can solve in lieu of solving the equation of radiative transfer directly, but at the price of introducing the need for a closure relation to approximate the higher moments that appear in the equations being solved. Some authors also combine moment and characteristic methods (e.g., [Rosen et al. 2017](#)).²

One of the simplest closures is the flux-limited diffusion (FLD) approximation ([Alme & Wilson 1973](#)), which retains only the first moment equation, for the radiation energy density, and closes the system by assuming both that the time derivative of the radiation flux is zero and that the radiation pressure tensor and has a particular, specified relationship with the radiation energy density. A more accurate approximation is to evolve both the radiation energy

¹ Fully eliminating the angular dependence requires some care, since in general for a moving medium the emissivity and opacity depend on direction, even if the medium itself interacts with light isotropically in its own rest frame. While one might guess velocity-dependent beaming effects might be unimportant in non-relativistic problems, it turns out that one cannot formulate a consistent, energy-conserving theory of non-relativistic radiation-hydrodynamics without including them, at least to leading order ([Mihalas & Klein 1982](#); [Lowrie et al. 1999](#); [Krumholz et al. 2007a](#)). Systems of moment equations where the radiation moments are written in the lab frame but the emissivity and absorption are written in the comoving frame, where they can be assumed to be isotropic, are known as mixed-frame formulations. Most modern radiation-hydrodynamics codes use the mixed-frame formulation.

² While characteristic and moment methods are the only ones widely used in astrophysics, in terrestrial applications (for example neutron transport calculations for nuclear reactor design) there are two other widely-used approaches to handle the angular dependence of the transfer equation. One is to discretise the unit sphere using weighted quadratures (the discrete ordinates, or S_n , method; e.g., [Lathrop & Carlson 1964](#); [Adams 1997](#)). The other is to expand the angular integration in spherical harmonics (the P_n method; e.g., [Modest 1989](#)).

* E-mail: ben.wibking@anu.edu.au (BDW)

density and the radiation flux, while still invoking a closure relation for the radiation pressure tensor; this is called a two-moment approach, since one solves for the first two moments of the radiation field. When the radiation pressure tensor (or rather, the ratio of the radiation pressure tensor to the radiation energy density, known as the Eddington tensor) is computed via a formal solution of the angle-dependent radiative transfer equation, we obtain the quasidiffusion or variable Eddington tensor (VET) method (Gol'din 1964). When retaining a local closure for the radiation pressure tensor in terms of the radiation energy density E and the flux F , we obtain a local VET method, commonly referred to as the M1 ('moment-one') method (Minerbo 1978; Levermore 1984; Dubroca & Feugeas 1999; González et al. 2007). There are a number of moment-based astrophysical radiation-hydrodynamic codes, implementing a wide variety of closures, in wide use in astrophysics, including ZEUS (Turner & Stone 2001), FLASH (Fryxell et al. 2000), ORION (Krumholz et al. 2007a; Shestakov & Offner 2008; Li et al. 2021), RAMSES (Commerçon et al. 2011; Rosdahl et al. 2013), ATHENA (Davis et al. 2012; Jiang et al. 2012), ENZO (Reynolds et al. 2009; Bryan et al. 2014), CASTRO (Zhang et al. 2011, 2013; Almgren et al. 2020), and FORNAX (Skinner et al. 2019), to give a partial list.

While the use of moment methods removes the dimensionality problem, it leaves a second problem, which is the strong mismatch in signal speeds between radiation and sound (or MHD) waves in non-relativistic problems, which travel at far less than the speed of light. This mismatch renders simple explicit methods, as are commonly used for HD and MHD, impractically slow for radiative transfer, due to the tiny time steps that would be imposed by the Courant-Friedrichs-Lewy (CFL) condition. For this reason, numerical methods for coupled radiation-hydrodynamics (RHD) or radiation-magnetohydrodynamics (RMHD) either use an implicit method for the radiation part of the problem (e.g., ZEUS, ORION, some versions of RAMSES and ATHENA) or adopt the reduced speed of light approximation (Gnedin & Abel 2001; Skinner & Ostriker 2013; FORNAX, other versions of RAMSES and ATHENA). The reduced speed of light approximation (RSLA) consists of replacing the speed of light c that defines the signal speed in the radiation moment equations, with a lower speed \hat{c} , while keeping the terms that describe the rate of momentum and energy exchange between gas and radiation unchanged. The lower speed \hat{c} , while still substantially larger than the HD or MHD signal speeds, is close enough to those speeds to allow radiation time steps large enough to render explicit methods computationally feasible.

1.2 Why a new radiation-hydrodynamics code?

In this paper we describe QUOKKA³, a new code for RHD. In terms of the taxonomy introduced above, QUOKKA is a two moment code that uses the RSLA to allow an explicit treatment of radiation transport. QUOKKA is also an adaptive mesh refinement (AMR) code, so it begins with a base grid at uniform spatial resolution, but then dynamically adds higher-resolution grids as needed to achieve user-specified accuracy goals (Berger & Oliger 1984; Berger & Colella 1989). However, these features do not make QUOKKA unique: ORION and RAMSES (among others) offer moment-based AMR RMHD, while FORNAX uses RSLA on a dendritic (though not adaptive) grid.

The unique feature of QUOKKA is that it has been designed from the ground up to run efficiently on graphics processing units (GPUs).

This design goal motivated our choice of both algorithms and low-level implementation details. While QUOKKA is not the first GPU hydrodynamics code in astrophysics (others include GAMER, Schive et al. 2010, 2018, CHOLLA, Schneider & Robertson 2015, CASTRO, Almgren et al. 2020, and ARK-RT, Bloch et al. 2021), nor even the first AMR GPU code, it is the first to feature two-moment AMR RHD on GPUs.

Bringing RHD to GPUs creates some unique challenges. Contemporary compute nodes are often limited by data bandwidth, both in terms of moving data between main memory and the CPU or GPU, and in terms of moving data between CPUs or GPUs. For this reason, implicit methods generally have poor scalability, due to the need for global communications during an implicit solve. This imbalance between computation and communication is magnified on GPUs. Likewise, robust implicit methods require iterative sparse matrix solvers, which achieve lower peak efficiency on GPUs compared to CPUs due to their heavy use of indirect addressing and highly branching control flow. These considerations motivate our choice of an explicit RSLA method. It also motivates our choice of time integration strategy, which as we detail below has been designed to maximize computation (and therefore minimize the relative amount of communication) on each hydrodynamic timestep. We show that, with this strategy, we are able to achieve update computation rates of approximately 70 million zone updates per second per GPU for pure HD, and almost 10 million for RHD. We also achieve > 93% parallel efficiency out to 256 GPUs. This combination of performance and scaling makes QUOKKA orders of magnitude faster than any other public RHD code.

The remainder of this paper is organized as follows. In Section 2 we introduce the set of equations that QUOKKA solves, and detail our numerical methods for solving them. In Section 3, we present a wide range of tests that demonstrate the accuracy and capabilities of the code. Section 4 covers our tests of code performance and scalability. In Section 5, we discuss the range of applicability of our methods, and our plans for application and future code expansions. Finally, the code itself, including all test problems, is freely available on [GitHub](#) under an open-source license.

2 METHODS

2.1 Equations

We solve the equations of radiation hydrodynamics (Pomraning 1973; Mihalas & Mihalas 1984; Castor 2004) for an inviscid, nonrelativistic fluid in local thermodynamic equilibrium in the mixed-frame formulation, where the radiation variables are defined in an inertial frame (i.e., Eulerian simulation coordinates) and the radiation-matter interaction terms are written in the frame comoving with the fluid, with the transformations between the frames accounted for via the addition of radiation-matter exchange terms that depend explicitly on the ratio of fluid velocity to the speed of light, $\beta = v/c$. We write

³ Quadrilateral, Umbra-producing, Orthogonal, Kangaroo-conserving Kode for Astrophysics!

the equations as follows:

$$\frac{\partial \rho}{\partial t} + \nabla \cdot (\rho \mathbf{v}) = 0, \quad (2)$$

$$\frac{\partial(\rho \mathbf{v})}{\partial t} + \nabla \cdot (\rho \mathbf{v} \mathbf{v} + \mathbf{P}) = \mathbf{G}, \quad (3)$$

$$\frac{\partial E}{\partial t} + \nabla \cdot [(E + \mathbf{P})\mathbf{v}] = cG^0, \quad (4)$$

$$\frac{\partial E_r}{\partial t} + \nabla \cdot \mathbf{F}_r = -cG^0, \quad (5)$$

$$\frac{1}{c^2} \frac{\partial \mathbf{F}_r}{\partial t} + \nabla \cdot \mathbf{P}_r = -\mathbf{G}, \quad (6)$$

where ρ is the gas density, \mathbf{v} is the gas velocity, E is the total energy density of the gas, $\mathbf{P} = \delta_{ij}P$ is the gas pressure tensor, E_r is the radiation energy density, \mathbf{F}_r is the radiation flux, \mathbf{P}_r is the radiation pressure tensor, $\nabla \cdot \rho \mathbf{v} \mathbf{v}$ denotes the sum $(\rho v_i v^j)_{,j}$, and G^i is the radiation four-force, with G^0 the time-like component and \mathbf{G} consisting of the space-like components. In the mixed-frame formulation, the radiation four-force to order β is

$$-cG^0 = \rho(\kappa_P 4\pi B - \kappa_E cE_r) + \rho\kappa_F \left(\frac{\mathbf{v}}{c} \cdot \mathbf{F}_r \right), \quad (7)$$

$$-\mathbf{G} = -\rho\kappa_F \frac{\mathbf{F}_r}{c} + \rho\kappa_P \left(\frac{4\pi B}{c} \right) \frac{\mathbf{v}}{c} + \rho\kappa_F \frac{\mathbf{v} \mathbf{P}_r}{c}, \quad (8)$$

where κ_F , κ_E , and κ_P are the flux-mean, energy-mean, and Planck-mean specific opacities evaluated in the comoving frame, B is the Planck function evaluated at the gas temperature, and $\mathbf{v} \mathbf{P}_r$ is the tensor contraction $v_j P_r^{ij}$ (Mihalas & Mihalas 1984). The latter two terms in the expression for \mathbf{G} correspond to the relativistic work term of Krumholz et al. (2007a) and are only important in the regime $\beta\tau \gtrsim 1$ (where τ is a characteristic optical depth), to which we cannot apply the RSLA (as discussed below), so we neglect them. However, the term of order β in the expression for cG^0 corresponds to the work done by the radiation force on the gas and can be the dominant term for problems of interest.

To apply the RSLA to these equations, we first rewrite the radiation moment equations so that they have a factor of exactly $1/c$ next to each of the time derivatives:

$$\frac{1}{c} \frac{\partial E_r}{\partial t} + \nabla \cdot \left(\frac{\mathbf{F}_r}{c} \right) = -G^0, \quad (9)$$

$$\frac{1}{c} \frac{\partial}{\partial t} \left(\frac{\mathbf{F}_r}{c} \right) + \nabla \cdot \mathbf{P}_r = -\mathbf{G}, \quad (10)$$

then we replace this $1/c$ factor with a factor of $1/\hat{c}$, where \hat{c} is the reduced speed of light, and multiply through by factors of \hat{c} to obtain the conservation law form of the reduced speed of light radiation moment equations (e.g., Skinner & Ostriker 2013):

$$\frac{\partial E_r}{\partial t} + \nabla \cdot \left(\frac{\hat{c}}{c} \mathbf{F}_r \right) = -\hat{c}G^0, \quad (11)$$

$$\frac{\partial \mathbf{F}_r}{\partial t} + \nabla \cdot (c\hat{c} \mathbf{P}_r) = -c\hat{c} \mathbf{G}. \quad (12)$$

The maximum wave speed of this system of equations is bounded by \hat{c} (as long as the flux satisfies causality, i.e. $F_r \leq cE_r$). As emphasised by Skinner & Ostriker (2013), all other factors of c remain unchanged, and, since the factors of c are unchanged on the right-hand side of the hydrodynamic equations, the reduced speed of light radiation hydrodynamic system does not conserve total energy or momentum for $\hat{c} \neq c$. When the left-hand side flux divergence terms are negligible, this nonconservation implies that the equilibrium temperature of the reduced speed of light system is slightly modified with

respect to the correct equilibrium temperature, implying that we cannot apply the RSLA to problems in the equilibrium diffusion limit in general (see section ??).

Writing out the right-hand side terms explicitly, we obtain

$$\frac{\partial \rho}{\partial t} + \nabla \cdot (\rho \mathbf{v}) = 0, \quad (13)$$

$$\frac{\partial(\rho \mathbf{v})}{\partial t} + \nabla \cdot (\rho \mathbf{v} \mathbf{v} + \mathbf{P}) = \rho\kappa_F \mathbf{F}_r / c, \quad (14)$$

$$\frac{\partial E}{\partial t} + \nabla \cdot [(E + \mathbf{P})\mathbf{v}] = -c\rho(\kappa_P a_r T^4 - \kappa_E E_r) - \rho\kappa_F \left(\frac{\mathbf{v}}{c} \cdot \mathbf{F}_r \right), \quad (15)$$

$$\frac{\partial E_r}{\partial t} + \nabla \cdot \left(\frac{\hat{c}}{c} \mathbf{F}_r \right) = \hat{c}\rho(\kappa_P a_r T^4 - \kappa_E E_r) + \rho\kappa_F \left(\frac{\hat{c}}{c} \frac{\mathbf{v}}{c} \cdot \mathbf{F}_r \right), \quad (16)$$

$$\frac{\partial \mathbf{F}_r}{\partial t} + \nabla \cdot (c\hat{c} \mathbf{P}_r) = -\hat{c}\rho\kappa_F \mathbf{F}_r. \quad (17)$$

These equations make no approximations about the frequency dependence of the radiation field. However, for computational tractability, in what follows we will approximate κ_F with the Rosseland mean opacity κ_R , which yields the correct radiation force in the diffusion limit, and approximate κ_E with the Planck mean opacity κ_P , which yields the correct energy absorption and emission in the optically-thin limit for fluids at rest (Mihalas & Mihalas 1984). However, we emphasise that the choice to set $\kappa_F \approx \kappa_R$ and $\kappa_E \approx \kappa_P$ is an additional approximation, and that others might be preferable depending on the physical system being simulated. In future work, we plan to address the limitations of these approximate grey opacities via an extension of our method to the multigroup solution of the radiation moment equations. Our present set of equations is sufficient for grey nonrelativistic radiation hydrodynamics in the semi-transparent regime, where we can neglect the ‘relativistic work term’ that is important only in the dynamic diffusion ($\beta\tau \gtrsim 1$) regime, as described earlier.

2.2 Solution method

We solve the system formed by equation 13–equation 17 using an operator split approach, whereby we first advance the hydrodynamic transport subsystem (Section 2.2.1), then the radiation transport subsystem (Section 2.2.2), and finally update the local coupling terms (Section 2.2.3). The first subsystem uses a single explicit update step, the second a set of subcycled explicit updates, and the third an implicit update. We describe each of these steps below.

This update cycle operates within a Berger & Olinger (1984) / Berger & Colella (1989) adaptive mesh refinement (AMR) framework, whereby each spatial variable is represented by a volume average in each cell, on a rectangular, Cartesian grid. We cover the entire computational domain with a coarse grid with cell spacings Δx_0 , Δy_0 , Δz_0 in the x , y , and z directions, the grid spacings need not be the same, but for most applications we choose them to be the same. We denote this coarse grid level 0. We then dynamically add (or remove) additional, finer grids over parts of the domain in response to user-specified refinement criteria. We denote these additional levels 1, 2, ..., with each grid on level l having cells a factor of 2 smaller than those on level $l - 1$, so that the cell spacing on level l is $\Delta x_0/2^l$ in the x direction, and similarly for y and z . We use only factor of 2 refinements in order to minimize numerical glitches arising from

the discontinuous change in resolution, which can arise especially in problems where shocks cross the coarse-fine mesh interface at an oblique angle (e.g., [Fryxell et al. 2000](#)). When adding finer grids we conservatively interpolate the underlying coarse data, and when removing finer grids we conservatively average down the fine data. Time steps on different AMR levels are sub-cycled, such that the time step on level l is $\Delta t_l = \Delta t_0/2^l$. At the end of every two time steps on level $l > 0$, we perform a synchronization step to ensure that we maintain machine-precision conservation for all conserved quantities ([Section 2.2.4](#)).

2.2.1 Hydrodynamics

For the solution of the hydrodynamic subsystem ([equation 13–equation 15](#), omitting the matter-radiation coupling terms on the right hand sides), we adopt a method-of-lines (or semi-discrete) approach, discretizing the spatial variables while initially keeping the time variable continuous, thereby transforming the partial differential equations into a large set of ordinary differential equations that can be integrated in time using a standard ordinary differential equation (ODE) integrator ([Jameson et al. 1981](#)). For the latter, we use the second-order strong stability preserving Runge-Kutta method (RK2-SSP; [Shu & Osher 1988](#)). Such an approach has been successfully employed in several recent astrophysical hydrodynamics codes ([Skinner et al. 2019; Stone et al. 2020](#)). Somewhat surprisingly, we find that such a method-of-lines scheme is not stable when combining higher-order spatial reconstruction with forward Euler time integration, but is stable for timesteps $O(\Delta x/(c_s + |v|))$ when used with higher-order Runge-Kutta methods. This is consistent with the findings of [Stone et al. \(2020\)](#), who also find that the maximum stable timestep corresponding to the linear stability region of a given (second order or higher) Runge-Kutta integrator divided by the number of spatial dimensions. As [Skinner et al. \(2019\)](#) notes, in contrast to fully-discrete unsplit hydrodynamic methods such as the corner transport upwind (CTU) method ([Colella 1990](#)), the coupling across corners of each cell is achieved via the use of a multi-stage time integrator, rather than via direct computation of fluxes from diagonal neighbors of each cell. While we are formally limited to a smaller timestep compared to the CTU method, our method may be more robust in practice, as the CTU integrator has been found to be unstable in supersonic turbulence with strong radiative cooling unless very small (≤ 0.1) CFL numbers are employed ([Schneider & Robertson 2017](#)).

We reconstruct the hydrodynamic variables on each face of each cell from the cell-average variables of the neighbouring cells. We perform this reconstruction using the piecewise parabolic method (PPM) ([Colella & Woodward 1984](#), hereafter [CW84](#)) using the primitive hydrodynamic variables (density, velocity, and pressure). As is standard, the conversion from conserved (density, momentum, and energy) to primitive variables is carried out assuming that the volume average and cell centered states are equivalent, which is an approximation accurate to $O(\Delta x^2)$. As noted by several authors, the PPM algorithm is therefore formally second-order accurate in spatial resolution. After the primitive variables have been defined, for the reconstruction step proper, we use the standard interface-centered PPM stencil:

$$q_{j+1/2} = \frac{7}{12}(q_j + q_{j+1}) - \frac{1}{12}(q_{j+2} + q_{j-1}). \quad (18)$$

We follow the implementation of [Stone et al. \(2020\)](#) in re-grouping the above terms symmetrically with respect to the interface $i + 1/2$ so as to preserve exact symmetry in floating point arithmetic. On a mesh

with uniform spacing between cells along the direction perpendicular to the reconstructed face, this stencil is fourth-order accurate.

We do not perform the slope-limiting and contact steepening steps of [CW84](#). We instead enforce monotonicity of the reconstructed state by re-setting the interface state to lie between the values of the cells adjacent to the interface, following [Mignone et al. \(2005\)](#). This is followed by the extrema detection and overshoot correction step within each cell as described by [CW84](#). In this step, the parabola assumed to exist across each cell is examined. If an ‘overshoot’ (as defined by [CW84](#)) of the parabola is detected, we follow the original [CW84](#) prescription of performing linear reconstruction on the side of the cell affected by the overshoot. If an extremum is instead detected, rather than forcing the reconstruction to a constant value across the cell as done by [CW84](#), we revert to performing a linear reconstruction within the affected cell, following [Balsara \(2017\)](#). Either of these steps may make the interface states discontinuous, with distinct states associated with each of the two cells adjacent to an interface, which reduces the order of accuracy of the stencil. Since the overall algorithm is a combination of second order and (usually) fourth order steps, the overall order of accuracy of PPM is often referred to as ‘third order’ (e.g., [Stone et al. 2020](#)). However, we note that there exist fully fourth-order extensions of PPM ([Felker & Stone 2018](#)), but these methods do not allow local source terms to be evaluated independently for each cell while maintaining fourth-order accuracy.

We also implement reconstruction based on a piecewise-linear method (PLM) using the monotonized-central (MC) slope limiter ([Van Leer 1977](#)). We use PPM reconstruction by default, but allow PLM reconstruction via a compile-time option.

In some cases, especially in underresolved strong shocks, the previous steps do not provide sufficient dissipation to avoid oscillations. This problem was recognized by [CW84](#), who proposed a shock flattening procedure in combination with a small amount of artificial viscosity. We find that this shock flattening procedure is not sufficient in multidimensional problems. Instead, we follow [Miller & Colella \(2002\)](#), who generalize the [CW84](#) procedure for multidimensional hydrodynamics. Using this latter method, we find that no artificial viscosity is needed and we do not include any in our implementation.

Finally, in order to compute the flux of mass, momentum and energy between cells, we use the HLLC Riemann solver with the ‘primitive variable Riemann solver’ wavespeeds and intermediate states ([Toro 2013](#)). We make the standard approximation that the face-average flux is the same as the face-centered flux, and therefore this step is also second-order accurate in spatial resolution. For each cell, the fluxes across each face are then added together to produce an unsplit spatial divergence term used by each stage of the Runge-Kutta integrator to advance the cell in time.

In multidimensional simulations, it has been long recognized that in strong grid-aligned shocks, the HLLC Riemann solver can unphysically amplify the so-called ‘carbuncle’ instability ([Quirk 1994](#)). In astrophysical problems, this is most often encountered in strong explosions. Implementing additional dissipation in the form of artificial viscosity (e.g., [Gittings et al. 2008](#)), the ‘H-correction’ ([Sanders et al. 1998](#)), or by adaptively switching to an HLL Riemann solver ([Harten et al. 1983](#)) for computing fluxes perpendicular to strong shocks (e.g., [Quirk 1994; Skinner et al. 2019](#)) are possible solutions to this issue. In future work we plan to implement an adaptive procedure to fix the carbuncle instability via the latter method.

Future work may also include implementing an adaptive method to reduce the order of reconstruction in order to preserve density and pressure positivity in near-vacuum regions, such as the multidimensional optimal order detection (MOOD) method of [Clain et al.](#)

(2011). An alternative solution may be to adaptively switch to an exact (iterative) Riemann solver depending on the flow conditions (Toro 2013).

2.2.2 Radiation

We solve the radiation transport subsystem (equation 16–equation 17, again omitting the terms on the right-hand side) in a similar method-of-lines fashion. Our approach is most similar to that of Skinner et al. (2019), who also evolve the radiation moment equations with a time-explicit method-of-lines approach; however, they do not use either PPM reconstruction or a reduced speed of light. Because even with the RSLA the signal speed for the radiation subsystem is substantially larger than for the hydrodynamic subsystem, we evolve the former explicitly in time with several radiation timesteps per hydrodynamic timestep. In the regime of applicability of the RSLA, this approach allows a much more computationally efficient solution to the radiation moment equations, due to the fact that explicit methods have a greater arithmetic intensity per byte of data and do not require global communication across the computational domain in order to advance the solution in time. This high arithmetic intensity is greatly beneficial on GPUs, where the ratio of floating-point arithmetic performance to memory bandwidth is typically greater than on CPUs.

We carry out each radiation subcycle using the same RK2-SSP integrator (Shu & Osher 1988) that we use for hydrodynamics. We likewise use a finite volume representation of the radiation variables, with a PPM spatial reconstruction (or optionally, PLM) of the radiation energy density E_r and reduced flux $\mathbf{f} = \mathbf{F}_r/cE_r$; the only difference in our procedures for hydrodynamics and radiation is that for radiation we do not employ a shock flattening procedure. There can exist unphysical radiation shocks when using local closures, since in general such closures make the radiation subsystem nonlinear, but there is no applicable shock flattening procedure to suppress this effect. We carry out reconstruction in terms of the reduced flux \mathbf{f} rather than the absolute flux \mathbf{F}_r in order to suppress unphysical fluxes $|\mathbf{F}_r| > cE_r$. This is effective in 1D problems, but in multidimensional problems, the magnitude of the radiation flux may still exceed cE_r , which is an unphysical state in which local closures cannot compute the Eddington factor at all. Reducing the order of reconstruction to first order (piecewise constant) when the interface states violate this constraint helps but does not eliminate the issue in all cases. For the purpose of computing the local closure only, we use rescale the flux such that $|\mathbf{F}_r| = cE_r$ whenever $|\mathbf{F}_r| > cE_r$. For particularly difficult problems, especially in order to avoid unphysical instabilities in the propagation of non-grid-aligned optically-thin radiation fronts, we find that it is necessary to reconstruct the radiation variables using PLM reconstruction.

One drawback to upwind finite volume methods for radiation transport is that in naive form, they do not give the correct behavior for diffusive regions where the optical depth per cell is much greater than unity. This failure occurs because numerical diffusion dominates over physical diffusion when using upwind methods when the mean free path of photons is not resolved (Lowrie & Morel 2001). One common approach to fix this incorrect behavior is to modify the Riemann solver in the optically thick regime to reduce the upwind bias of the spatial derivative (Audit et al. 2002; Skinner et al. 2019; Mezzacappa et al. 2020). However, this can lead to violations of causality (i.e., $F_r > cE_r$) when the radiation flux is in the streaming regime (Audit et al. 2002), which occurs especially at discontinuities in the opacity between optically-thin cells and optically-thick cells. The only apparent fix for this problem, which we adopt, is to

disable the optical-depth correction in the Riemann solver for those cells where it produces a causality-violating state. We find that this condition is only activated when $f \rightarrow 1$, so it may not qualitatively affect the solution. However, we also advocate refining on the gradient in the optical depth per cell in order to resolve the boundary layers in such situations whenever it is computationally feasible. In Appendix A, we show that such an optical-depth correction substantially improves the behavior of the method in the diffusion regime, but is not sufficient to obtain the correct diffusion solution in thermal waves.

For computing the flux of radiation quantities between cells, we use an HLL Riemann solver, with wavespeeds computed assuming the Eddington factors are fixed at the beginning of the timestep (Balsara 1999). This approach allows us to substitute different closure relations for the Eddington factors without requiring a modification of the Riemann solver, unlike previous implementations that are restricted to a single local closure (e.g., González et al. 2007; Skinner & Ostriker 2013; Skinner et al. 2019). In principle, we could even use Eddington tensors computed via a short characteristics formal solution of the radiative transfer equation (e.g., Davis et al. 2012), but we leave exploration of a non-local VET method to future work.

Our default closure relation for the Eddington tensor is the Levermore (1984) closure, which is derived by assuming that the radiation field is isotropic in some (unknown) reference frame and then computing a Lorentz transform from this reference frame to one in which the reduced flux \mathbf{f} matches the value in the cell under consideration. This procedure leads to a radiation pressure tensor (e.g., González et al. 2007; Rosdahl et al. 2013; Skinner & Ostriker 2013)

$$\mathbf{P}_r = \left(\frac{1-\chi}{2} \mathbf{I} + \frac{3\chi-1}{2} \hat{\mathbf{n}}\hat{\mathbf{n}} \right) E_r \quad (19)$$

where \mathbf{I} is the identity tensor, and the Eddington factor χ and the flux direction cosine $\hat{\mathbf{n}}$ are

$$\chi = \frac{3 + 4f^2}{5 + 2\sqrt{4 - 3f^2}}, \quad (20)$$

$$\hat{\mathbf{n}} = \frac{\mathbf{F}_r}{|\mathbf{F}_r|}. \quad (21)$$

When the radiation flux is exactly zero, we drop the direction-dependent term in equation 19. By considering a coordinate system where the radiation flux is aligned with a coordinate axis, we see that χ is the component of the Eddington tensor in the direction of the radiation flux.

We emphasise that this is only one possible choice of closure, and a variety of alternative local closures exist (e.g., Minerbo 1978; Levermore & Pomraning 1981). We refer readers to Janka (1992) and Koerner & Janka (1992) for systematic comparisons to angle-dependent transport solutions for neutrinos, and Olson et al. (2000) for comparisons to photon solutions. Because of its prominence in the neutrino transport literature, as well as marginally favorable performance on some test problems, we also provide an implementation of the Minerbo (1978) closure in addition to the default Levermore (1984) option. However, users can also implement any local closure of their choice simply by providing an implementation of a function that maps from the reduced flux \mathbf{f} to the Eddington factor χ for their preferred closure. Doing so does not come at any cost in computational performance.

2.2.3 Matter-radiation coupling

Following the computation of the hyperbolic part of the radiation subsystem, we use an implicit method to evaluate the source terms (those

appearing on the right-hand sides of [equation 13](#)–[equation 17](#)) for both the radiation and hydrodynamic subsystems; this update occurs once per radiation subcycle, and thus several times per hydrodynamic step. Since there are no spatial derivatives in these terms, each cell can be updated independently.

The radiation-matter coupling update occurs in three steps. The first is to handle the energy source terms $c\rho(\kappa_P a_r T^4 - \kappa_E E_r)$ that appear in [equation 15](#) and [equation 16](#). In the regime of problems to which we can apply the RSLA this term is often the stiffest, and we therefore update it using the backward-Euler implicit method of [Howell & Greenough \(2003\)](#), specialized to the case of a single material and extended to include a reduced speed of light. Let $E_g = E - \rho v^2/2$ be the gas internal energy, and let E_g^t and E_r^t be the gas internal energy and radiation energy at the end of the hyperbolic update. We compute the new gas internal energy E_g^{t+1} and radiation energy E_r^{t+1} after accounting for the exchange term by solving the implicit system

$$0 = F_G \equiv (E_g^{t+1} - E_g^t) + \left(\frac{c}{\hat{c}}\right) R \quad (22)$$

$$0 = F_R \equiv (E_r^{t+1} - E_r^t) - (R + S), \quad (23)$$

where

$$R \equiv \Delta t \rho \kappa_P (4\pi B - \hat{c} E_r), \quad (24)$$

Δt is the radiation substep timestep, and S is an optional source term that we include to allow, for example, addition of radiation by stellar sources. The quantities F_G and F_R are the residual errors in the gas energy and radiation energy, respectively.

To solve this system via Newton-Raphson iteration, we require the Jacobian matrix, the elements of which are

$$\frac{\partial F_G}{\partial E_g} = 1 + \left(\frac{c}{\hat{c}}\right) \frac{\partial R}{\partial E_g}, \quad (25)$$

$$\frac{\partial F_G}{\partial E_r} = -c \Delta t \rho \kappa_P, \quad (26)$$

$$\frac{\partial F_R}{\partial E_g} = -\frac{\partial R}{\partial E_g}, \quad (27)$$

$$\frac{\partial F_R}{\partial E_r} = 1 + \hat{c} \Delta t \rho \kappa_P, \quad (28)$$

where

$$\frac{\partial R}{\partial E_g} = \frac{\rho \Delta t}{C_v} \left[\kappa_P \frac{\partial B}{\partial T} + \frac{\partial \kappa_P}{\partial T} (4\pi B - \hat{c} E_r^{t+1}) \right], \quad (29)$$

C_v is the gas specific heat at constant volume. From the Jacobian, we can write the change in radiation and gas temperature for each iterative update as

$$\Delta E_r = -\frac{F_R + \eta F_G}{\frac{\partial F_R}{\partial E_r} + \eta \frac{\partial F_G}{\partial E_r}}, \quad (30)$$

$$\Delta E_g = -\frac{F_G + \Delta E_r \frac{\partial F_G}{\partial E_r}}{\frac{\partial F_G}{\partial E_g}}. \quad (31)$$

where $\eta \equiv -(\partial F_R / \partial E_g)(\partial F_G / \partial E_g)^{-1}$. We repeatedly apply [equation 30](#) and [equation 31](#) to the radiation and gas energies until the system converges. [Howell & Greenough \(2003\)](#) leave unspecified the convergence criteria they use for their solver. After experimenting with several possibilities, we decide to stop the Newton-Raphson

iterations when the residuals F_R and F_G satisfy

$$\left| \frac{F_G}{E_{\text{tot}}} \right| < \epsilon \text{ and} \quad (32)$$

$$\left| \frac{c}{\hat{c}} \frac{F_R}{E_{\text{tot}}} \right| < \epsilon, \quad (33)$$

where

$$E_{\text{tot}} \equiv E_g^t + \frac{c}{\hat{c}} (E_r^t + S). \quad (34)$$

When $\hat{c} = c$, E_{tot} is the total (internal gas plus radiation) energy at the end of the timestep. By default, the relative tolerance ϵ is set to 10^{-10} . We find that larger tolerances produce unacceptably inaccurate solutions for many problems. In especially stiff problems, it may be necessary to reduce the tolerance to the order of machine precision for double-precision floating point arithmetic ($\sim 10^{-15}$). If the solver exceeds a specified maximum number of iterations (400 by default) without converging, the code prints an error message and exits. Convergence failure usually occurs only when the initial timestep has not been sufficiently reduced compared to the CFL timestep at the start of a simulation.

Once the Newton-Raphson iterations have converged, we have obtained the updated gas internal energy and radiation energy, and we proceed to the next two steps of updating the coupling terms. We first update the radiation and gas momenta, accounting for the coupling term $\rho \kappa_F \mathbf{F}_r / c$. To do so compute the flux mean opacity κ_F using the updated gas temperature. Then, following [Skinner et al. \(2019\)](#), we use a backward-Euler discretization of the radiation flux source term (modified to include a reduced speed of light):

$$\mathbf{F}_r^{t+1} = \frac{\mathbf{F}_r^t}{1 + \rho \kappa_F \hat{c} \Delta t}. \quad (35)$$

In order to ensure momentum conservation when $\hat{c} = c$, we apply the difference in radiation flux in an equal and opposite manner to the gas momenta (as advocated by [Skinner et al. 2019](#)):

$$\Delta \mathbf{F}_r \equiv \mathbf{F}_r^{t+1} - \mathbf{F}_r^t, \quad (36)$$

$$(\rho \mathbf{v})^{t+1} = (\rho \mathbf{v})^t - \frac{\Delta \mathbf{F}_r}{\hat{c} c}. \quad (37)$$

The final step is to compute the work done by the radiation force on the gas. Since we are evolving the conserved variables, this term cannot be computed explicitly as written in [equation 16](#) without causing a significant error in the gas internal energy when the radiation force is stiff. We instead compute this term as the difference in gas kinetic energy over the timestep ΔE_{kin} , then add this quantity to the total gas energy and subtract this quantity from the radiation energy:

$$E^{t+1} \leftarrow E_g^{t+1} + (E_{\text{kin}}^t + \Delta E_{\text{kin}}), \quad (38)$$

$$E_r^{t+1} \leftarrow E_r^{t+1} - \left(\frac{\hat{c}}{c}\right) \Delta E_{\text{kin}}, \quad (39)$$

where E^{t+1} denotes the total gas energy at the end of the timestep. This completes the update for all radiation-matter coupling terms.

2.2.4 Level synchronization procedure

As explained by [Berger & Colella \(1989\)](#), in an AMR calculation it is necessary to adjust the solution on the coarse AMR level following the solution on any refined level in order to maintain conservation of the evolved quantities (e.g., mass, momentum, energy). For hyperbolic equations evolved explicitly in time, this is traditionally done by saving the flux at the coarse-fine grid boundary in a ‘flux register’

for both the flux computed on the fine level and the flux computed on the coarse level. In general, these fluxes are different due to the differing stencil used on the coarse and fine levels, and without correction, this would lead to a loss of conservation of energy (and any other conserved quantities). The flux register stores this mismatch, and in the synchronization step, adds the missing mass, momentum, or energy to the cells on the coarse level immediately adjacent to the coarse-fine boundary.

As noted by [Howell & Greenough \(2003\)](#), an implicit radiation update has additional difficulties in ensuring conservation, since radiation can propagate much further than a single grid cell on the coarse grid. Our radiation update is fully explicit, but we would like to advance each AMR level on the hydrodynamic timescale, rather than on the radiation timescale, so we have a similar long-range signal propagation difficulty. [Rosdahl et al. \(2013\)](#) outline three possible solutions to the problem: i) perform the radiation solve after each coarse hydrodynamic step, keeping subcycling-in-time on refined levels (which would be very inaccurate), ii) use a single global timestep for all AMR levels, which allows one to advance the radiation solution on all levels in each radiation substep (which would be very computationally expensive, since in our applications of interest, the global timestep is typically limited by the timestep of the highest-resolution level) or iii) restrict the timestep for each level to the minimum of the radiation and hydrodynamic timesteps. In our code, we set the coarse timestep such that the number of radiation substeps per level is limited to a maximum value $N_{\text{sub,max}}$ in order to minimize the signal propagation distance from the coarse-fine boundaries. The flux mismatch at the coarse-fine boundaries is added to the immediately adjacent cells on the coarse grid at the end of each level advance. When $N_{\text{sub,max}} = 1$, our solution is identical to the flux synchronization method used in the RAMSES AMR code ([Rosdahl et al. 2013](#)). However, as a default we set the parameter $N_{\text{sub,max}}$ to 10, which appears to be sufficient to avoid significant discontinuities in the radiation energy and flux at coarse-fine boundaries, but still allows for significant subcycling and thus a substantially lower computational cost. We use this value for all test problems shown in this work, but users are able to set this parameter as desired for either greater efficiency or greater consistency at refinement boundaries. When this parameter is too large, however, it is possible for the coarse level to fail to maintain positivity of the radiation energy or causality of the radiation flux.

3 TEST PROBLEMS

3.1 Hydrodynamics

3.1.1 Sound wave

We compute the propagation of a sound wave in one dimension in order to measure the convergence of our numerical method to the exact solution as a function of spatial resolution, following the test described by [Stone et al. \(2008\)](#). With \mathbf{U} denoting the vector of conserved variables $(\rho, \rho v_x, \rho v_y, \rho v_z, E)$, we initialize the simulations with the initial state $\mathbf{U}_0 + \delta\mathbf{U}$, where

$$\delta\mathbf{U} = A\mathbf{R}\sin(2\pi x) \quad (40)$$

where $\mathbf{R} = (1, -1, 1, 1, 1.5)$ is the right eigenvector of the linearized hydrodynamic system, and \mathbf{U}_0 is the background state with density $\rho = 1$, velocity $\mathbf{v} = 0$, and pressure $P = 1/\gamma$. We set the adiabatic index $\gamma = 5/3$ and the wave amplitude $A = 10^{-6}$. Periodic boundary conditions on the domain $[0, 1]$ are adopted and the simulations are evolved for one wave period, allowing us to compute the error of the

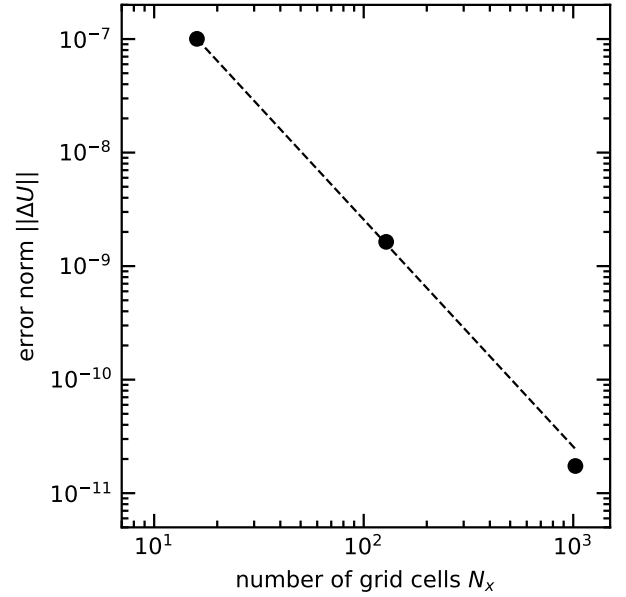


Figure 1. The error $\|\Delta\mathbf{U}\|_2$ in the solution (equation 41) for a sound wave as a function of spatial resolution per wavelength N_x is shown by the black circles. The dashed line is a power law that scales as N_x^{-2} normalised to the observed error at the smallest N_x .

solution by direct comparison of the initial conditions and the final state of the simulation. The error vector is computed as

$$(\Delta\mathbf{U})_k = \frac{1}{N_x} \sum_i |U_{i,k} - U_{i,k}^0| \quad (41)$$

where k denotes a component of each state vector, \mathbf{U}_i is the vector of conserved variables in cell i at the final timestep, and \mathbf{U}_i^0 is the vector of conserved variables in cell i in the initial conditions. The accuracy of the solution is assessed with the L_2 norm $\|\Delta\mathbf{U}\|_2$.

For a varying number of grid cells N_x , we run simulations using PPM reconstruction and a CFL number of 0.1. In [Figure 1](#), we show the error norm as a function of spatial resolution for these simulations. For $N_x = 16$, we obtain $\|\Delta\mathbf{U}\|_2 = 1.0 \times 10^{-7}$, for $N_x = 128$, we obtain $\|\Delta\mathbf{U}\|_2 = 1.6 \times 10^{-9}$, and for $N_x = 1024$, we obtain $\|\Delta\mathbf{U}\|_2 = 1.7 \times 10^{-11}$. The results for our code are in excellent agreement with those from the ATHENA hydrodynamic solver (Figure 7 of [Stone et al. 2008](#)). The N_x^{-2} scaling of the error norm indicates that our hydrodynamic solver converges at second order in spatial resolution, as expected from the formal order of accuracy of the method.

3.1.2 Contact discontinuity

The HLLC Riemann solver has the property that it can resolve an isolated stationary contact discontinuity with infinite resolution ([Toro 2013](#)). The HLL solver, on the other hand, introduces a large amount of numerical diffusion for this problem (see Figure 10.20 of [Toro 2013](#)). We show that a stationary contact discontinuity in [Figure 2](#) computed with the HLLC solver. The initial conditions have a left and right state separated with a discontinuity at $x = 0.5$. The left state is $\rho_L = 1.4$, $p_L = 1.0$, and the right state is $\rho_R = 1.0$ and $p_R = 1.0$. Since this is a pure contact discontinuity, the solution should not evolve from the initial state. We set the velocity to zero, and use an

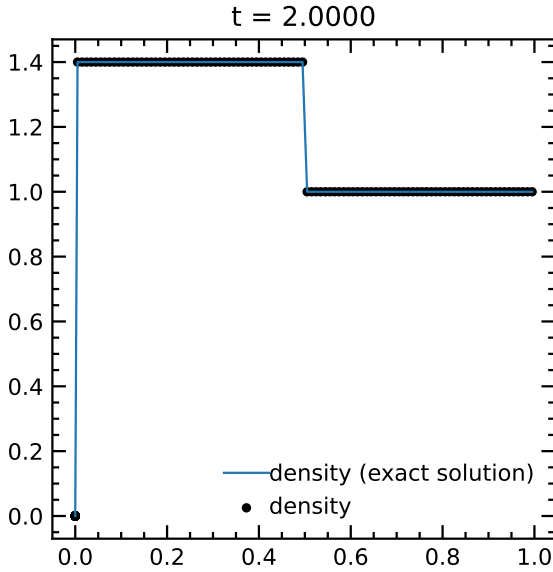


Figure 2. A contact discontinuity.

adiabatic index $\gamma = 1.4$. We evolve the solution numerically until $t = 2$. The error with respect to the correct solution is exactly zero.

3.1.3 Stationary shock tube

We show a stationary shock tube in Figure 3 using the parameters suggested on the website of F.X. Timmes⁴. This shock tube problem is substantially more difficult to solve than the standard Sod (1978) shock tube test due to the larger jump in pressure and density at the discontinuity. We initialize a left and right states with a discontinuity at $x_0 = 2$, with the left state $\rho_L = 10$, and $p_L = 100$, and the right state $\rho_R = 1$ and $p_R = 1$. The initial velocity is zero. We evolve the solution using a CFL number of 0.1 until $t = 0.4$ on a grid of 1000 cells on the domain $[0, 5]$. We use a small CFL number since the wave structure at the discontinuity creates waves that propagate faster than the linearized Roe eigenvalues would predict.

As for the sound wave test, we compute the L_1 error norm for each of the conserved variables, and then compute the root-mean-square of those error norms. The rms L_1 error norm divided by the rms norm of the exact solution is 1.12×10^{-3} . Inspecting the solution in Figure 3, we see that there agreement between the exact solution and the numerical solution is very good, with the only noticeable differences are small oscillations near discontinuities in the derivative of the solution at $x \approx 2.4$ and near the density discontinuity at $x \approx 3.6$. We find that the shock flattening method of Miller & Colella (2002) is essential to produce a reasonable solution to this problem. Without it, we find unacceptably large post-shock oscillations (not shown).

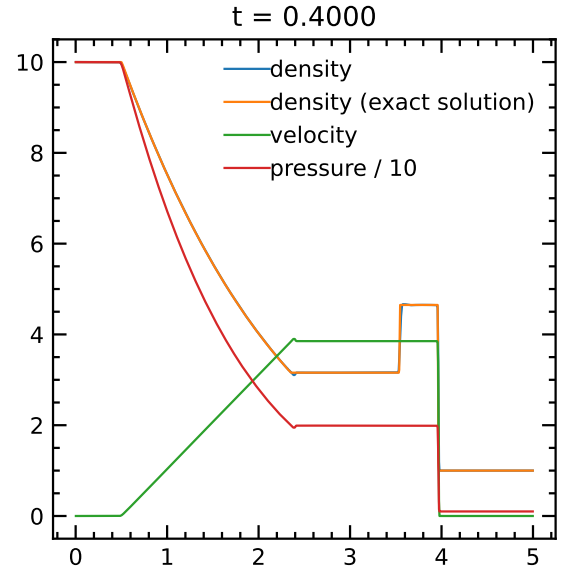


Figure 3. A shock tube.

3.1.4 ‘LeBlanc’ test

We show the ‘LeBlanc’ shock tube test in Figure 4. This was originally published by Benson (1992) and is further described by Pember & Anderson (2001). In this problem, we initialize a left state with $\rho_L = 1$ and $p_L = \frac{2}{3} \times 10^{-1}$, and a right state with $\rho_R = 10^{-3}$ and $p_R = \frac{2}{3} \times 10^{-10}$. We set the initial velocity to zero and use an adiabatic index $\gamma = 5/3$. This is an extreme shock tube that far exceeds any shock that may be encountered in any conceivable application, featuring a pressure jump of nine orders of magnitude, and is therefore an excellent test problem. We evolve this simulation until $t = 6$ using a grid of 2000 cells and a CFL number of 0.1. The resulting state is shown in Figure 4. The difficulty of obtaining the correct specific internal energy in the solution is highlighted by Pember & Anderson (2001). We show that we obtain the correct shock location and specific internal energy, with the exception of a small overshoot at the shock location. The use of shock flattening is essential for this problem. Overall, the performance of our code on this problem is excellent.

3.1.5 Wave-shock interaction (Shu-Osher) test

We show the Shu-Osher test in Figure 5 following the description of Shu & Osher (1989). The initial conditions are, on the left side,

$$\rho_L(x) = 3.857143, \quad (42)$$

$$v_L(x) = 2.629369, \quad (43)$$

$$P_L(x) = 10.33333, \quad (44)$$

and the right-hand state is

$$\rho_R(x) = 1 + 0.2 \sin(5x), \quad (45)$$

$$v_R(x) = 0, \quad (46)$$

$$P_R(x) = 1. \quad (47)$$

We compute a reference solution using ATHENA++ (Stone et al. 2020) with the VL2 integrator, PPM reconstruction in the characteristic

⁴ http://cococubed.asu.edu/code_pages/exact_riemann.shtml



Figure 4. The LeBlanc test problem.

variables, and the HLLC Riemann solver on a grid of 1600 cells. Our solution is computed using PPM reconstruction (in the primitive variables), the RK2-SSP integrator, and the HLLC Riemann solver on a grid of 400 cells. The agreement is very good, with comparable resolution of the high-frequency features to the third-order essentially-non-oscillatory (ENO) scheme of Shu & Osher (1989). When PLM reconstruction is used instead for the same number of grid cells, the high-frequency features are aliased (not shown; see also Figure 14 of Shu & Osher 1989), indicating a significantly higher effective resolution for PPM-based methods even in the presence of shocks.

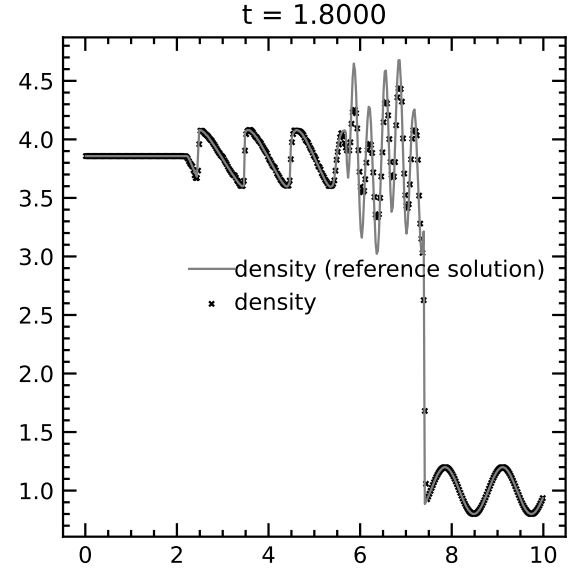


Figure 5. The Shu-Osher wave-shock interaction problem.

3.1.6 Slow-moving shock

We show a slow-moving shock in Figure 6 using the parameters from Jin & Liu (1996), where $\rho_L = 3.86$, $(\rho v)_L = -3.1266$, and $E_L = 27.0913$, and the right-side state $\rho_R = 1.0$, $(\rho v)_R = -3.44$, and $E_R = 8.4168$, with $\gamma = 1.4$. This corresponds to a shock moving to the right with a velocity $v_{\text{shock}} = 0.1096$. For a CFL number of 0.2, this corresponds to the shock taking ~ 250 timesteps to move across a single cell. This may not be a common scenario for our applications, but it may occur in a protostellar accretion shock, for instance. The quality of the solution is again significantly improved by shock flattening. The post-shock oscillations for slow-moving shocks may still be present with first-order reconstruction (Jin & Liu 1996; Lee 2011), so it is difficult to completely eliminate. We also find that adding a small amount of artificial viscosity does not significantly reduce the oscillations. A modification to PPM reconstruction based on a characteristic wave decomposition succeeds in significantly reducing this oscillation (Lee 2011), which we may consider implementing in a future version of the code.

3.1.7 Strong rarefaction

We show a strong rarefaction in Figure 7...

3.1.8 Kelvin-Helmholtz problem

We illustrate this problem in Figure 8...

3.1.9 2D Implosion problem

3.2 Radiation

3.2.1 Marshak wave

We show a Marshak wave in Figure 9...

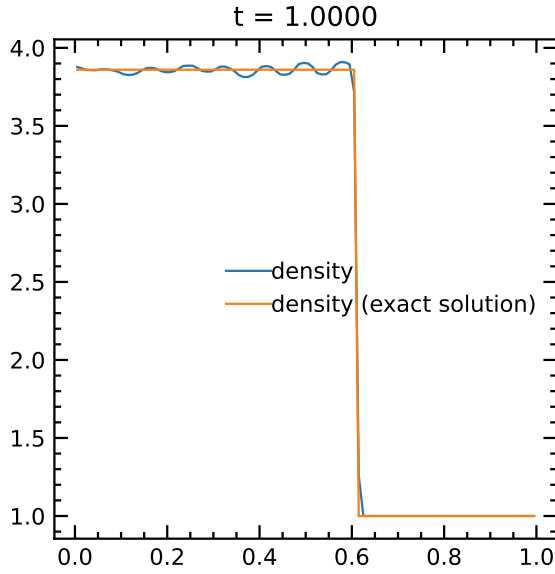


Figure 6. A slow-moving shock test problem.

3.2.2 Su-Olson problem

We show the Su-Olson problem in [Figure 10](#). . .

3.2.3 Radiation pressure tube

We show the radiation pressure tube ([Krumholz et al. 2007a](#)) in [Figure 11](#). . .

3.2.4 Radiation-matter energy exchange

We show the radiation-matter energy exchange test in [Figure 12](#). . .

3.2.5 Shadow test

3.2.6 Beam test

3.2.7 Optically-thin wind

3.3 Radiation hydrodynamics

3.3.1 Subcritical radiative shock

We show a radiative shock in [Figure 13](#). . .

3.3.2 Radiation-driven dust shell

4 PERFORMANCE AND SCALING

5 DISCUSSION AND CONCLUSIONS

5.1 Range of applicability

5.2 Planned applications

The future is bright for radiation hydrodynamics on GPUs. . .

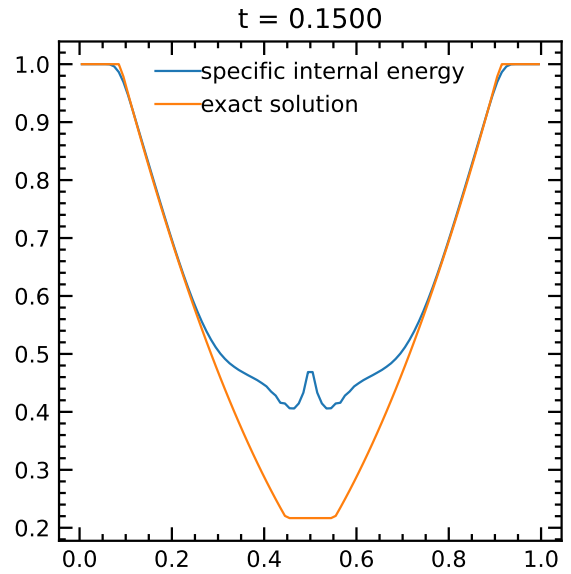
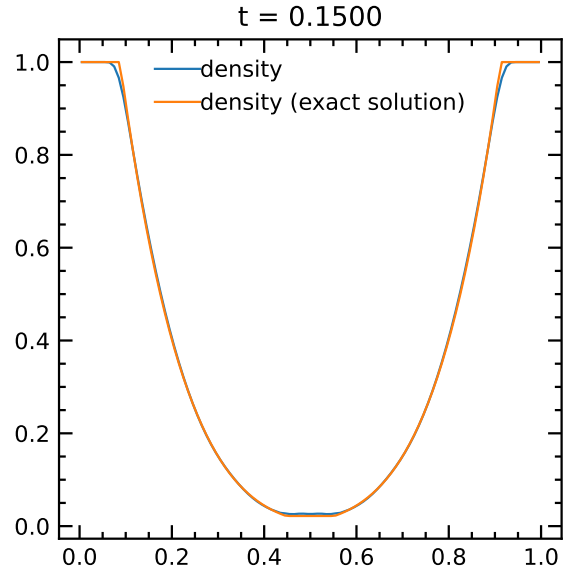


Figure 7. A strong rarefaction test problem.

ACKNOWLEDGEMENTS

This research was supported by the Australian Research Council through its Discovery Projects and Future Fellowship Funding Schemes, awards DP190101258 and FT180100375. This research was undertaken with the assistance of resources and services from the National Computational Infrastructure (NCI), which is supported by the Australian Government.

Software: AMReX ([AMReX Development Team et al. 2021](#)), matplotlib ([Hunter 2007](#)), numpy ([Harris et al. 2020](#)).

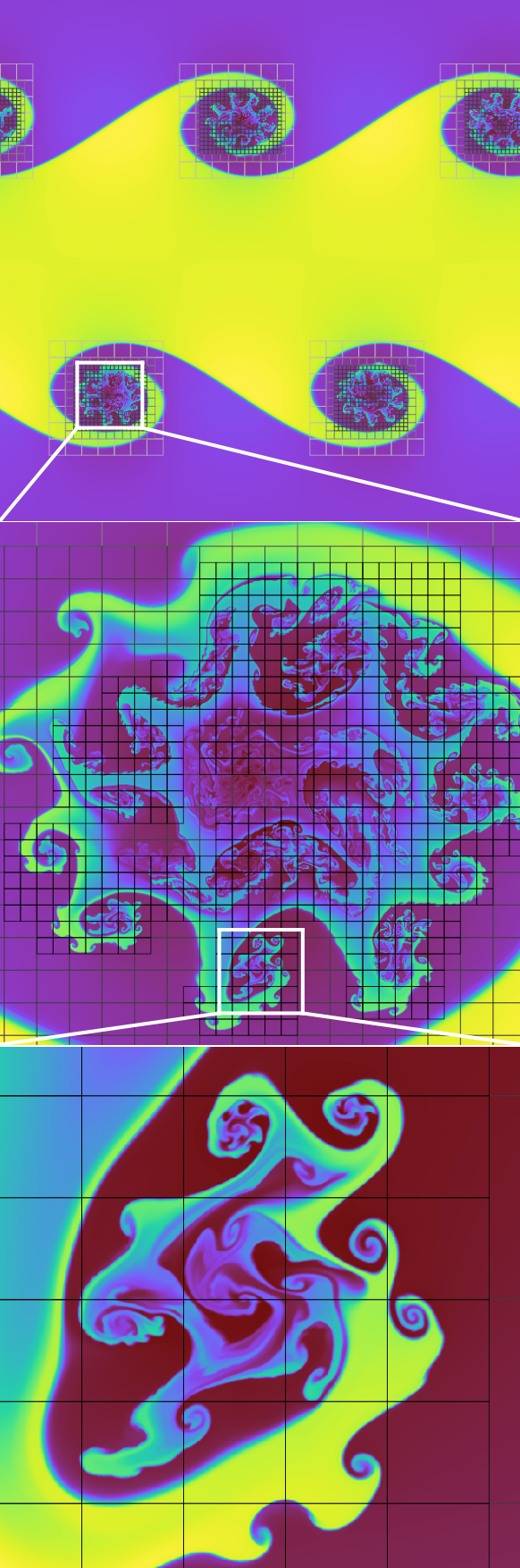


Figure 8. A Kelvin-Helmholtz problem with 4 levels of refinement.

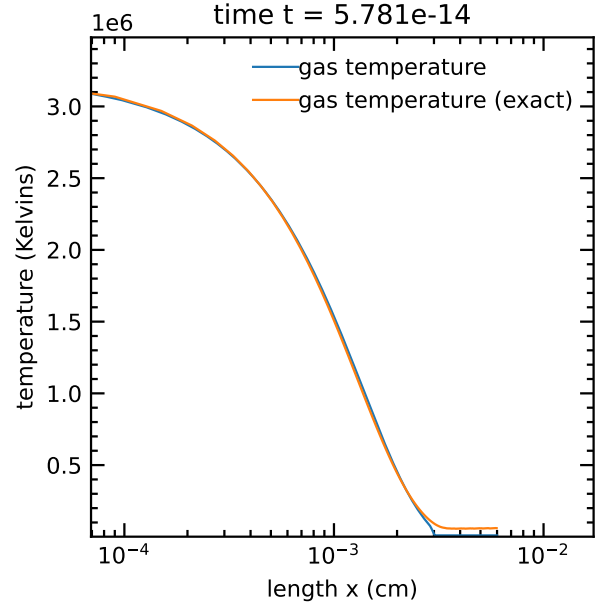


Figure 9. A Marshak wave test problem (Su & Olson 1996).

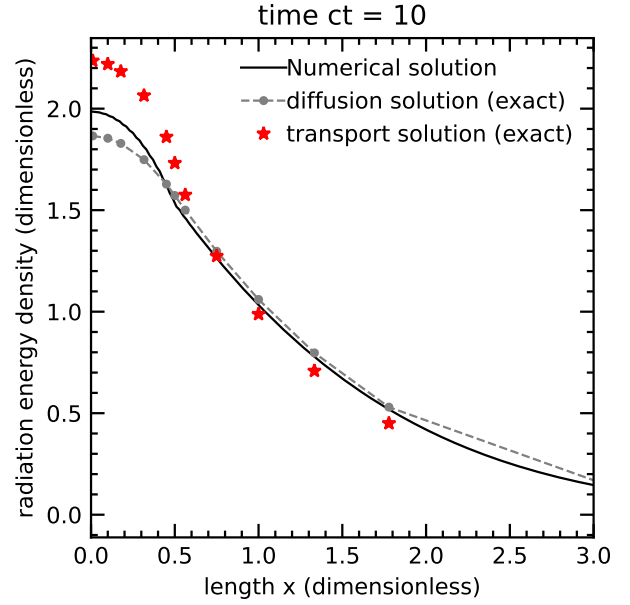


Figure 10. The Su-Olson test problem (Su & Olson 1997).

DATA AVAILABILITY

The source code and entire commit history for QUOKKA is hosted in this public [GitHub repository](#). The version of the source code used to produce the results in this paper as well as the output files at the final timestep for the simulations shown in the Figures are permanently archived at Zenodo DOI:XXX.

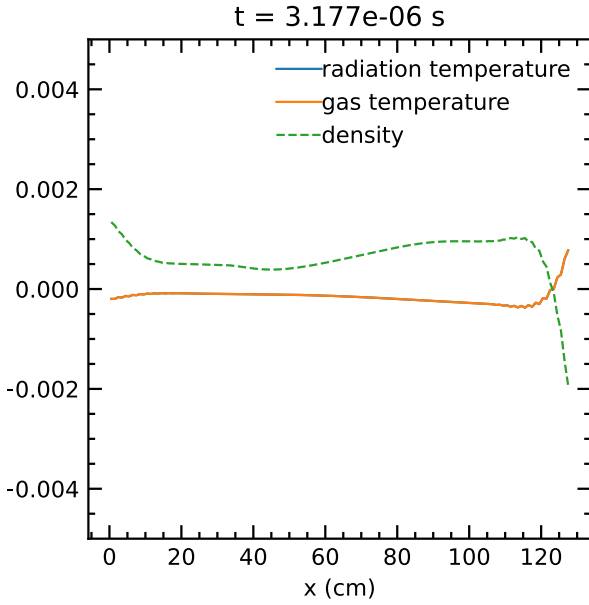


Figure 11. The radiation pressure tube (Krumholz et al. 2007a).

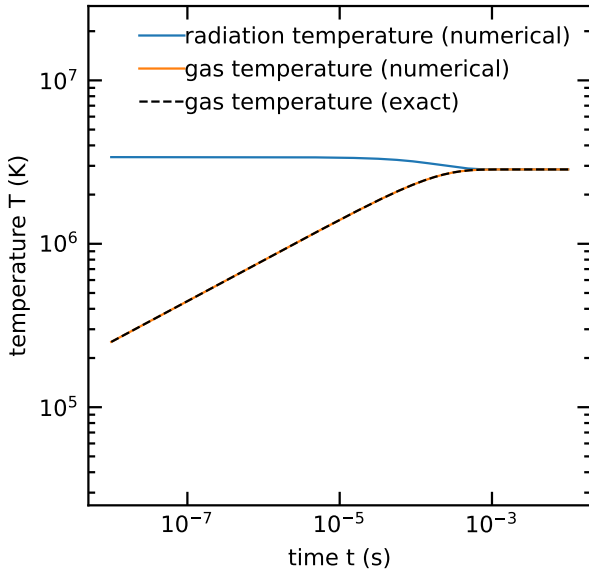


Figure 12. The radiation-matter coupling test (Turner & Stone 2001).

REFERENCES

- AMReX Development Team et al., 2021, AMReX-Codes/amrex: AMReX 21.09, [doi:10.5281/zenodo.5363443](https://doi.org/10.5281/zenodo.5363443), <https://doi.org/10.5281/zenodo.5363443>
- Abel T., Wandelt B. D., 2002, *MNRAS*, **330**, L53
- Adams M. L., 1997, *Trans. Theory Stat. Phys.*, **26**, 385
- Alme M. L., Wilson J. R., 1973, *ApJ*, **186**, 1015
- Almgren A., et al., 2020, *Journal of Open Source Software*, **5**, 2513
- Audit E., Charrier P., Chièze J. P., Dubroca B., 2002, arXiv e-prints, [pp](https://arxiv.org/abs/2002.00000)

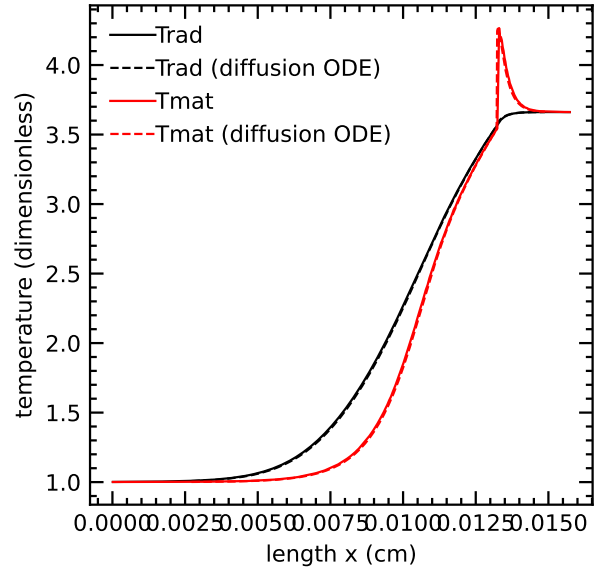


Figure 13. A $M = 3$ subcritical radiative shock test.

[astro-ph/0206281](https://arxiv.org/abs/2006.02062)

- Balsara D. S., 1999, *J. Quant. Spectrosc. Radiative Transfer*, **61**, 617
- Balsara D. S., 2017, *Living Reviews in Computational Astrophysics*, **3**, 2
- Benson D. J., 1992, *Journal of Computational Physics*, **100**, 143
- Berger M. J., Colella P., 1989, *J. Comp. Phys.*, **82**, 64
- Berger M. J., Olinger J., 1984, *J. Comp. Phys.*, **53**, 484
- Bloch H., Tremblin P., González M., Padialeau T., Audit E., 2021, *A&A*, **646**, A123
- Bryan G. L., et al., 2014, *ApJS*, **211**, 19
- Castor J. I., 2004, *Radiation Hydrodynamics*. Cambridge University Press, Cambridge, UK
- Clain S., Diot S., Loubère R., 2011, *J. Comp. Phys.*, **230**, 4028
- Colella P., 1990, *J. Comp. Phys.*, **87**, 171
- Colella P., Woodward P. R., 1984, *J. Comp. Phys.*, **54**, 174
- Commerçon B., Teyssier R., Audit E., Hennebelle P., Chabrier G., 2011, *A&A*, **529**, A35+
- Davis S. W., Stone J. M., Jiang Y.-F., 2012, *ApJS*, **199**, 9
- Dubroca B., Feugeas J., 1999, *Academie des Sciences Paris Comptes Rendus Serie Sciences Mathematiques*, **329**, 915
- Felker K. G., Stone J. M., 2018, *J. Comp. Phys.*, **375**, 1365
- Fleck Jr. J. A., Cummings Jr. J. D., 1971, *J. Comp. Phys.*, **8**, 313
- Fryxell B., et al., 2000, *ApJS*, **131**, 273
- Gittings M., et al., 2008, *Computational Science and Discovery*, **1**, 015005
- Gnedin N. Y., Abel T., 2001, *New Astron.*, **6**, 437
- Gol'din V., 1964, *USSR Computational Mathematics and Mathematical Physics*, **4**, 136
- González M., Audit E., Huynh P., 2007, *A&A*, **464**, 429
- Harris C. R., et al., 2020, *Nature*, **585**, 357
- Harten A., Lax P. D., Leer B. v., 1983, *SIAM Review*, **25**, 35
- Howell L. H., Greenough J. A., 2003, *J. Comp. Phys.*, **184**, 53
- Hunter J. D., 2007, *Computing in Science & Engineering*, **9**, 90
- Jameson A., Schmidt W., Turkel E., 1981, in *AIAA*.
- Janka H. T., 1992, *A&A*, **256**, 452
- Jiang Y.-F., Stone J. M., Davis S. W., 2012, *ApJS*, **199**, 14
- Jin S., Liu J.-G., 1996, *J. Comp. Phys.*, **126**, 373
- Koerner A., Janka H. T., 1992, *A&A*, **266**, 613
- Krumholz M. R., Klein R. I., McKee C. F., Bolstad J., 2007a, *ApJ*, **667**, 626
- Krumholz M. R., Stone J. M., Gardiner T. A., 2007b, *ApJ*, **671**, 518
- Lathrop K. D., Carlson B. G., 1964, [doi:10.2172/4666281](https://doi.org/10.2172/4666281)

- Lee D., 2011, in Pogorelov N. V., Audit E., Zank G. P., eds, *Astronomical Society of the Pacific Conference Series Vol. 444, 5th International Conference of Numerical Modeling of Space Plasma Flows (ASTRONUM 2010)*, p. 236
- Levermore C. D., 1984, *J. Quant. Spectrosc. Radiative Transfer*, 31, 149
- Levermore C. D., Pomraning G. C., 1981, *ApJ*, 248, 321
- Li P. S., et al., 2021, *J. Open Source Software*
- Lowrie R. B., Morel J. E., 2001, *J. Quant. Spectrosc. Radiative Transfer*, 69, 475
- Lowrie R. B., Morel J. E., Hittinger J. A., 1999, *ApJ*, 521, 432
- Mezzacappa A., Endeve E., Messer O. E. B., Bruenn S. W., 2020, *Living Reviews in Computational Astrophysics*, 6, 4
- Mignone A., Plewa T., Bodo G., 2005, *ApJS*, 160, 199
- Mihalas D., Klein R. I., 1982, *J. Comp. Phys.*, 46, 97
- Mihalas D., Mihalas B. W., 1984, *Foundations of radiation hydrodynamics*. Oxford University Press
- Miller G. H., Colella P., 2002, *J. Comp. Phys.*, 183, 26
- Minerbo G. N., 1978, *J. Quant. Spectrosc. Radiative Transfer*, 20, 541
- Modest M. F., 1989, *J. Thermophys. Heat Trans.*, 3, 3
- Olson G. L., Auer L. H., Hall M. L., 2000, *J. Quant. Spectrosc. Radiative Transfer*, 64, 619
- Pember R. B., Anderson R. W., 2001, [doi:10.2514/6.2001-2644](https://doi.org/10.2514/6.2001-2644)
- Pomraning G. C., 1973, *The equations of radiation hydrodynamics*. Pergamon Press
- Quirk J. J., 1994, *International Journal for Numerical Methods in Fluids*, 18, 555
- Reynolds D. R., Hayes J. C., Paschos P., Norman M. L., 2009, *Journal of Computational Physics*, 228, 6833
- Rijkhorst E., Plewa T., Dubey A., Mellema G., 2006, *A&A*, 452, 907
- Rosdahl J., Blaizot J., Aubert D., Stranex T., Teyssier R., 2013, *MNRAS*, 436, 2188
- Rosen A. L., Krumholz M. R., Oishi J. S., Lee A. T., Klein R. I., 2017, *J. Comp. Phys.*, 330, 924
- Sanders R., Morano E., Druguet M.-C., 1998, *J. Comp. Phys.*, 145, 511
- Schive H.-Y., Tsai Y.-C., Chiueh T., 2010, *ApJS*, 186, 457
- Schive H.-Y., Zuhone J. A., Goldbaum N. J., Turk M. J., Gaspari M., Cheng C.-Y., 2018, *MNRAS*, 481, 4815
- Schneider E. E., Robertson B. E., 2015, *ApJS*, 217, 24
- Schneider E. E., Robertson B. E., 2017, *ApJ*, 834, 144
- Shestakov A. I., Offner S. S. R., 2008, *J. Comp. Phys.*, 227, 2154
- Shu C.-W., Osher S., 1988, *J. Comp. Phys.*, 77, 439
- Shu C.-W., Osher S., 1989, *J. Comp. Phys.*, 83, 32
- Skinner M. A., Ostriker E. C., 2013, *ApJS*, 206, 21
- Skinner M. A., Dolence J. C., Burrows A., Radice D., Vartanyan D., 2019, *ApJS*, 241, 7
- Sod G. A., 1978, *Journal of Computational Physics*, 27, 1
- Stone J. M., Gardiner T. A., Teuben P., Hawley J. F., Simon J. B., 2008, *ApJS*, 178, 137
- Stone J. M., Tomida K., White C. J., Felker K. G., 2020, *ApJS*, 249, 4
- Su B., Olson G. L., 1996, *J. Quant. Spectrosc. Radiative Transfer*, 56, 337
- Su B., Olson G. L., 1997, *Annals of Nuclear Energy*, 24, 1035
- Toro E. F., 2013, *Riemann solvers and numerical methods for fluid dynamics: a practical introduction*. Springer Science & Business Media
- Tsang B. T.-H., Milosavljević M., 2015, *MNRAS*, 453, 1108
- Turner N. J., Stone J. M., 2001, *ApJS*, 135, 95
- Van Leer B., 1977, *J. Comp. Phys.*, 23, 276
- Zhang W., Howell L., Almgren A., Burrows A., Bell J., 2011, *ApJS*, 196, 20
- Zhang W., Howell L., Almgren A., Burrows A., Dolence J., Bell J., 2013, *ApJS*, 204, 7

APPENDIX A: ASYMPTOTIC DIFFUSION CORRECTION

This paper has been typeset from a \LaTeX file prepared by the author.