

# Absolute power calibration using gravity field and photon calibrators for gravitational wave experiment

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**Abstract.**

## 1. Introduction

The discovery of the gravitational wave give us the new probe for observing our universe. Instruments for measuring the gravitational wave must be designed for accuracy as well as precession. The typical strain sensitivity of 2nd generation interferometric detectors(IFO), such as Advanced LIGO and Advanced Virgo, and KAGRA, is around  $10^{-23}/\sqrt{Hz}$ . In order to estimate the parameters of gravitational parameters, Understanding of the systematic and statistical noise sources are critical.

To reduce the systematic error of the gravitational wave reconstruction, we need to calibrate the response of IFO. LIGO and Virgo employ the photon calibrator for the calibration of time-dependent optical gain of IFO. The first generation of the photon calibrator is developed by the Gragow. The frequency range of photon calibrator is between 10 and 3000Hz. However, it has a challenging issue of the absolute calibration due to the accuracy of the absolute laser power of laser standard between each country. Current uncertainty of absolute power is a few percent.

The dynamic gravity field generator is one of the candidates of the solutions. Related techniques using photon calibrator and gravity field calibrator are described as Karuki. et. al and Matone et al.

## 2. Photon calibrator

Photon calibrator (Pcal) relies on the photon radiation pressure from the power modulated laser beams reflecting on the test mass to apply periodic force via the recoil

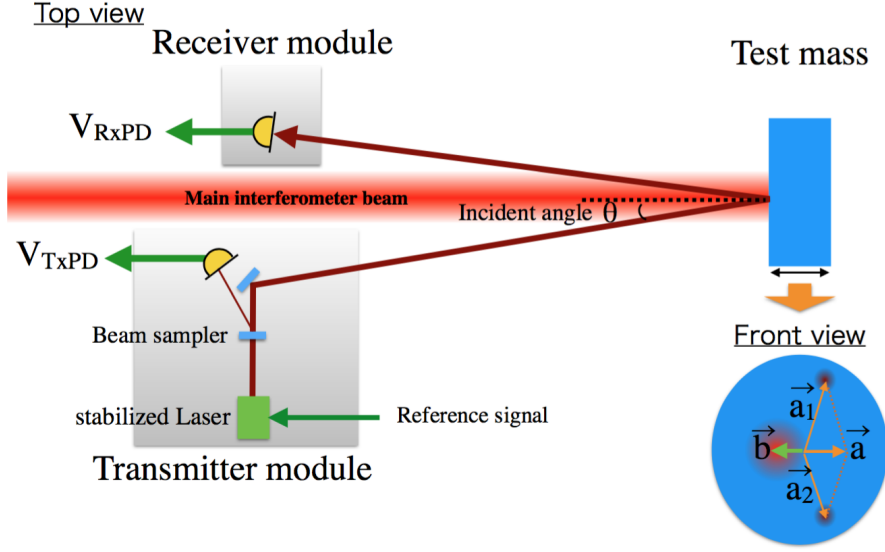


Figure 1.

of photons. LIGO, Virgo and KAGRA employ the photon calibrator for the calibration of the interferometer response. The displacement of the test mass can be described as

$$x = \frac{P \cos \theta}{2c} s(\omega) \left( 1 + \frac{M}{I} \vec{a} \cdot \vec{b} \right), \quad (1)$$

where  $P$  is absolute laser power,  $\theta$  is incident angle of the Pcal laser,  $M$  is mass of test mass,  $\omega$  is angular frequency,  $I$  is inertia,  $\vec{a}$  and  $\vec{b}$  are position vector of Pcal laser beams.  $s(\omega)$  is transfer function between force and displacements. The laser power is stabilized less than design sensitivity. The stabilized laser is mounted on the transmitter module. The power is monitored by the average response of the photo detectors at the transmitter module and receiver module. The uncertainty of displacement corresponds to uncertainty of laser power. The largest uncertainty of photon calibrator is uncertainty of laser power. LIGO and KAGRA use the working standard to cross-calibrate the relative response of each interferometer. The uncertainty of each calibrator is XXXX %. For the absolute calibration, we compare the NIST laser power standard. We send Gold standard to NIST and calibrate the response of detector. After that, we bring the gold standard to LHO and calibrate the working standard Hanford, Livingston and KAGRA. The uncertainty of laser standard of each institute of standard is XXX %. The second largest uncertainty of photon calibrator is optical efficiency of optical path. the estimated uncertainty of optical efficiency in LIGO is XXXX%. We need to estimate the laser power response through the response of photo detector at the outside of vacuum chamber. The reduction of these uncertainties are essential for the gravitational wave observation.

Table 1.

	KAGRA	advanced LIGO	advanced Virgo
Mirror material	Sapphire	Silica	Silica
Mirror mass	22.8 kg	40 kg	40 kg
Mirror diameter	220 mm	340 mm	350 mm
Mirror thickness	150 mm	200 mm	200 mm
Distance of Pcal from ETM	36 m	8 m	1.5 m
Pcal laser power	20 W	2W	3 W
Laser frequency	1047 nm	1047 nm	1047 nm
Incident angle	0.72 deg	8.75 deg	30 deg

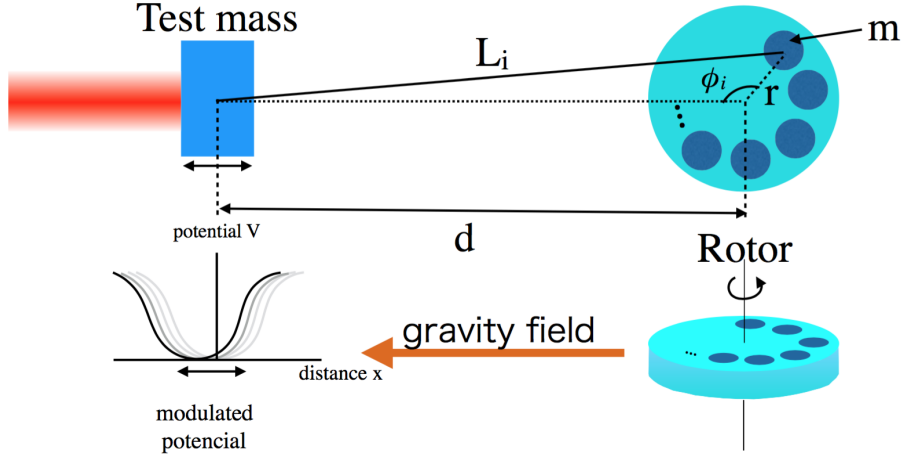


Figure 2.

### 3. Gravity field calibrator

To solve the calibration uncertainty, we propose the gravity field calibrator. The original idea is described in Matone. et.al. We calculated the displacement by changing dynamic gravity field of multipole moment with with  $N$  masses. The assumed model is the suspended test masses for the interferometer and disc with multipole mass as shown in Fig. 2. We put the masses  $m$  at the positions of the radius  $r$ . The distance between the center of mass of mirror and disc is assumed  $d$ . We rotate the disc at the angular frequency of  $\omega_{rot} = 2\pi f_{rot}$ .

Distance between  $i$ -th mass and center of test mass is

$$L_i = d \sqrt{1 + \left(\frac{r}{d}\right)^2 - 2 \left(\frac{r}{d}\right) \cos \phi_i}, \quad (2)$$

where the angle of  $i$ -th mass is assumed as  $\phi_i = \omega_{rot}t + 2\pi i/N$ . The gravity potential

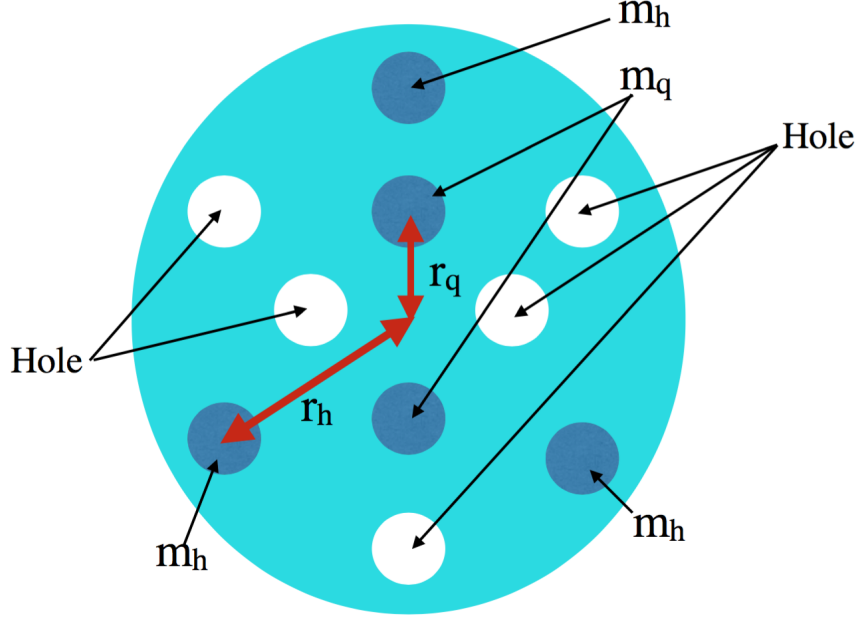


Figure 3.

at the center of test mass can be described as

$$V = \sum_{i=0}^N V_i \quad (3)$$

$$= -\frac{GMm}{d} \sum_{i=0}^N \sum_{n=0}^{\infty} \left(\frac{r}{d}\right)^n P_n \left( \cos \left( \omega_{rot} t + \frac{2\pi i}{N} \right) \right), \quad (4)$$

where  $P_n$  is Legendre polynomial. The equation of motion of test mass is

$$Ma = \left| \frac{\partial V}{\partial d} \right| = \frac{GMm}{d^2} \sum_{i=0}^N \sum_{n=0}^{\infty} (n+1) \left(\frac{r}{d}\right)^n P_n \left( \cos \left( \omega_{rot} t + \frac{2\pi i}{N} \right) \right), \quad (5)$$

where  $a$  is acceleration of test mass. We place the quadrupole and hexapole masses in same rotor as shown in Fig. ???. We will calculate the displacement of quadrupole and hexapole in the section 3.1

### 3.1. Displacement of test mass(Quadrupole)

We calculate the displacement of the quadrupole mass distribution, which corresponds to  $N = 2$ . The mass and radius of quadrupole are assumed as  $m_q, r_q$ . The equation of motion of test mass is described as

$$Ma = \frac{GMm_q}{d^2} \sum_{n=0}^{\infty} (n+1) \left(\frac{r_q}{d}\right)^n (P_n(\cos(\omega_{rot} t)) + P_n(\cos(\omega_{rot} t + \pi))). \quad (6)$$

If we assume  $r \ll d$ , the displacement of the time-dependent lower harmonics can be written by

$$x = \sum_{i=1}^{\infty} x_{ni} \cos(n\omega_{rot} t) \sim x_{2f} \cos(2\omega_{rot} t) = x_{2f} \cos \omega t, \quad (7)$$

where amplitude of 2-f is

$$x_{2f} = \frac{9}{8} \frac{GMm_q r_q^2}{d^4} s(\omega). \quad (8)$$

### 3.2. Displacement of test mass(Haxapole)

We calculate the displacement of the quadrupole mass distribution, which corresponds to  $N = 3$ . The mass and radius of quadrupole are assumed as  $m_h, r_h$ . The equation of motion of test mass is described as

$$Ma = \frac{GMm_h}{d^2} \sum_{n=0}^{\infty} (n+1) \left( \frac{r_h}{d} \right)^n \times \left( P_n(\cos(\omega_{rot}t)) + P_n\left(\cos\left(\omega_{rot}t + \frac{2\pi}{3}\right)\right) + P_n\left(\cos\left(\omega_{rot}t + \frac{4\pi}{3}\right)\right) \right). \quad (9)$$

If we assume  $r \ll d$ , the displacement of the time-dependent lower harmonics can be written by

$$x = \sum_{i=1}^{\infty} x_{nf} \cos(n\omega_{rot}t) \sim x_{3f} \cos(3\omega_{rot}t) = x_{3f} \cos\omega t, \quad (11)$$

where amplitude of 3-f is

$$x_{3f} = \frac{5}{6} \frac{GMm_h r_h^3}{d^5} s(\omega). \quad (12)$$

## 4. Absolute power calibration by using Gcal and Pcal

In this section, we will discuss about absolute laser power calibration using interferometer. Figure 4 shows the configuration of absolute laser power calibration. First, we modulate the mirror position using gravity field calibrator. We can measure the peaks of  $x_{2f}$  and  $x_{3f}$  in the response of interferometer. Second, we send the interferometer signal to the excitation port of photon calibrator as a reference and performs feedback control. Third, we measure the absolute laser power at the detector of transmitter module and receiver module. By using Eq (1),(13), and (14), the modulated powers are

$$P_{2f} = \frac{9}{4} \frac{Gcm_q Mr_q^2}{d^4 \cos\theta} \frac{1}{1 + \frac{M}{I} \vec{a} \cdot \vec{b}} \quad (13)$$

$$P_{3f} = \frac{5}{3} \frac{Gcm_h Mr_h^3}{d^5 \cos\theta} \frac{1}{1 + \frac{M}{I} \vec{a} \cdot \vec{b}} \quad (14)$$

Fourth, we demodulate the signal of transmitter module and receiver module using the encoder signal of gravity field calibrator. The demodulated signals are

$$V_{2f}^T = \rho_T P_{2f} \quad (15)$$

$$V_{2f}^R = \rho_R P_{2f} \quad (16)$$

$$V_{3f}^T = \rho_T P_{3f} \quad (17)$$

$$V_{3f}^R = \rho_R P_{3f} \quad (18)$$

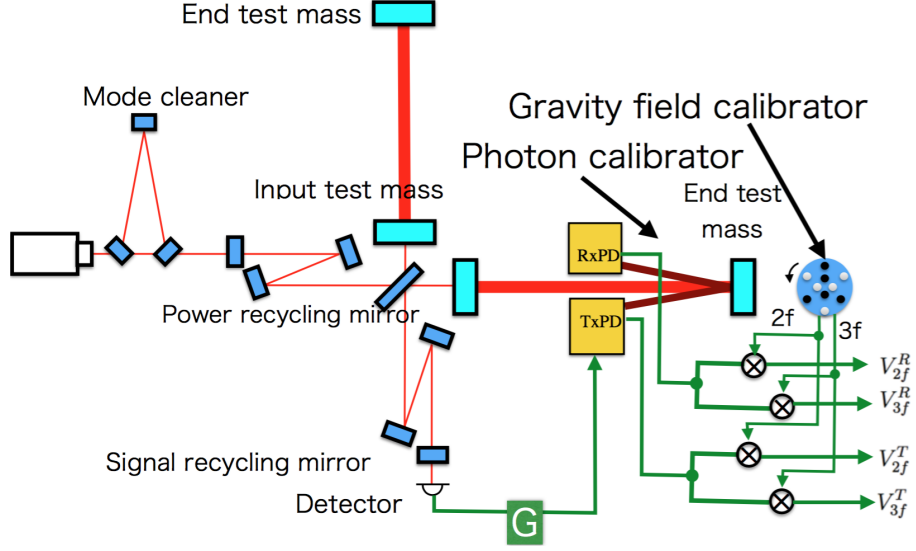


Figure 4.

Table 2. The assumed parameters.

$G$	Gravity constant	
$c$	speed of light[m/sec]	
$I$	Inertia[kgm <sup>2</sup> ]	
$\theta$	Incident andgle[deg]	0.7
$M$	Mass of test mass[kg]	23
$m_q$	Mass of quadrupole[kg]	3.777
$m_h$	Mass of hexapole[kg]	7.404
$r_q$	Radius of quadrupole[m]	0.130
$r_h$	Radius of hexapole[m]	0.060
$d$	Distance[m]	2.00

Therefore, we can estimate the distance using measured data.

$$d = \frac{20 V_{2f}^T m_h r_h^3}{27 V_{3f}^T m_q r_q^2} = \frac{20 V_{2f}^R m_h r_h^3}{27 V_{3f}^R m_q r_q^2} \quad (19)$$

Using Eq (19),(13) and (14), we can calibrate the absolute leaser power.

## 5. Estimation of uncertainty

We estimate the uncertainty of laser power. The assumed parameters are listed in Table XXX. We assumed the gaussian noise for each the parameters. We estimate the uncertainty of the laser power using Monte-Carlo simulation. The result is shown in Table XXXX. Table XXXX shows the uncertainty of photon calibrator system based on LIGO system. By using Gcal system, we can reduce the uncertainty less than XXXX.

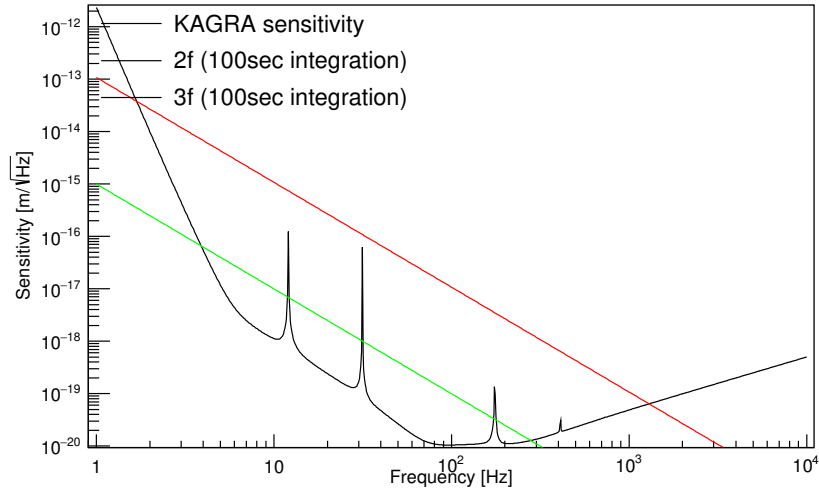


Figure 5.

## 6. Discussion

This method has an advantage of a direct comparison of the amplitude of injected power and gravity field. In previous study, we need to consider the uncertainty of the optical efficiency through the window and mirrors. This is because we put the working standard at the outside of the chamber. However, the compared gravity field can modulate the mirror directory.

We need to pay attention to the modulation effect of the intermediate mass because the gravity field act all the mass. 伝達関数がキャンセルする利点こと。低い周波数だと IM の効果を入れないといけなく、伝達関数のズレがあること。反跳効果と追従効果。いつでも絶対キャリブレーションができること。窓の効果率や積分球の効果を介さず計算できること。r/d の数字によっては、1 % の効果を考えるときにルジャンドル多項式の高次の効果を入れないといけない。有限要素法との比較。ローターのパラメータの最適化。

## 7. Summary