



An overview of Cosmological Inflation and Supergravity

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Abstract

[1]

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1 Prologue

1.1 Introduction to Cosmology

On large scales the universe follows The Cosmological Principle which states that the universe is both homogenous and isotropic in every direction [1]. This is a rather simple idea, equating to the universe appearing to be the same everywhere on large scales, however, the reason as to why the universe exists in this manner is not so simple.

In the years following Einstein's publication of 'The Foundations of the General Theory of Relativity' [2], 2 models were put forward, the Einstein universe and the de Sitter universe. Both were static models though they differed in their composition with the Einstein model containing a non-zero density for matter, for which he was required to include a cosmological constant, Λ to prevent the universe from collapsing in on itself, while the de Sitter model neglects ordinary matter and is dominated by this cosmological constant. The model that Einstein described is static due to his inclusion of the cosmological constant, and the de Sitter model is one that undergoes rapid expansion, and unlike the Einstein model it satisfies the perfect cosmological principle, assuming homogeneity and isotropy both spatially and temporally.

The universe was first discovered to be expanding by Edwin Hubble [3] in 1929, though this was a known prediction derived from General Relativity [4], first done so by Alexander Friedmann in his 1922 paper 'On the Curvature of Spacetime' [5], in which he derived, what would come to be known as the Friedmann-Lemaître equations [6], as in 1927 Georges Lemaître independently derived them [7]. Friedmann demonstrated that the energy density of matter, $\rho(t)$, and the cosmological constant, Λ , were linked with the time evolution of the universe.

1.2 Mathematical Background

1.2.1 Cosmological Key

$$H(t) : \text{Hubble parameter} \quad (1.2.1)$$

$$a(t) : \text{Scale factor} \quad (1.2.2)$$

$$G : \text{Newtons Gravitational Constant} \quad (1.2.3)$$

$$\rho(t) : \text{Energy density} \quad (1.2.4)$$

$$p(t) : \text{Pressure} \quad (1.2.5)$$

$$w : \text{Barotropic parameter} \quad (1.2.6)$$

$$m_p : \text{Reduced Planck Mass} \quad (1.2.7)$$

$$k : \text{Curvature parameter} \quad (1.2.8)$$

$$\text{Flat Universe} : k = 0 \text{ with } w \neq -1 \quad (1.2.9)$$

$$H \propto a^{-1} \quad (1.2.10)$$

$$H \propto \rho^{\frac{1}{2}} \quad (1.2.11)$$

$$\rho \propto a^{-3(1+w)} \quad (1.2.12)$$

$$a \propto t^{\frac{2}{3(1+w)}} \quad (1.2.13)$$

$$(1.2.14)$$

1.2.2 Universe Models

Einstein Universe

De Sitter Universe

1.2.3 A brief crash course

The Hubble Parameter, H , is defined as the expansion rate of the universe and is given by the formula

$$H = \frac{\dot{a}}{a} \quad (1.2.15)$$

where a is the scale factor of the universe. This is a key part of the 2 non-linear ordinary differential equations derived by Friedmann which take the form

$$H^2 \equiv \left(\frac{\dot{a}}{a}\right)^2 = \frac{1}{3m_p^2}\rho - \frac{k}{a^2} \quad (1.2.16)$$

known as the Friedmann equation, which for a flat universe reduces to

$$\rho = 3m_p^2 H^2 \quad k \approx 0, \quad (1.2.17)$$

and

$$\dot{H} + H^2 = \frac{\ddot{a}}{a} = -\frac{1}{6m_p^2}(\rho + 3p) \quad (1.2.18)$$

known as the acceleration equation where the overdots, \dot{x} , denote derivatives with respect to time, t . In an expanding universe, ($\dot{a} > 0$), filled with matter that satisfies $\rho + 3p \geq 0$, means Eq. (1.2.3) implies $\ddot{a} < 0$. The indication of this condition is that there was a singularity in the finite past, $a(t \equiv 0) = 0$ known as the Big Bang singularity. Combining Eq. (1.2.3) and Eq. (1.2.3), results in the continuity equation

$$\dot{\rho} + 3H(\rho + p) = 0. \quad (1.2.19)$$

Defining

$$w = \frac{p}{\rho} \quad (1.2.20)$$

where, w , is the barotropic parameter for the universe, means that the continuity equation, Eq. (1.2.3) may be integrated to find the relation of the energy density, ρ , to the scale factor, a , taking the form

$$\rho \propto a^{-3(1+w)}. \quad (1.2.21)$$

Combining this with the Friedmann equation, Eq. (1.2.3), demonstrates the relation between the scale factor and time,

$$a(t) \propto \begin{cases} t^{\frac{2}{3(1+w)}}, & \text{for } w \neq -1, \\ e^{Ht}, & \text{for } w = -1 \end{cases} \quad (1.2.22)$$

where the scale factor relation for $w = -1$ is the relation present in the de Sitter universe model hence inflation where $w \approx -1$ is called quasi-de Sitter inflation.

ρ is the total energy density of the universe, though it can be used for the individual species of the universe, denoted by ρ_x , where the energy density of each species is determined by its barotropic parameter. Depending on what species is dominant in the universe, the relations vary as follows:

i	w	$\rho(a)$	$a(t)$
matter ¹ , m	0	a^{-3}	$t^{\frac{2}{3}}$
radiation, γ	$\frac{1}{3}$	a^{-4}	$t^{\frac{1}{2}}$
dark energy, Λ	-1	a^0	e^{Ht}

If there are multiple species contributing significantly then the pressure, p , and energy density, ρ , are

$$\rho \equiv \sum_i \rho_i \quad p \equiv \sum_i p_i. \quad (1.2.23)$$

For each species, i , the present ratio of the energy density to the critical energy density is defined as

$$\Omega_i \equiv \frac{\rho_i}{\rho_{crit}} \quad \rho_{crit} \equiv 3H_0^2 \quad (1.2.24)$$

and this corresponds to the barotropic parameter

$$w_i \equiv \frac{p_i}{\rho_i}, \quad (1.2.25)$$

and for a flat universe, $k \approx 0$,

$$\sum_i \Omega_i \approx 1 \quad (1.2.26)$$

For an accelerated rate of expansion, (i.e. $\ddot{a} > 0$), requires $w < -\frac{1}{3}$ and from table 1.2.3 it can be inferred that the dominant species must be dark energy. Dark energy is a substance with the interesting

¹Matter refers to both baryonic matter and dark matter as they have the same relations

property of exerting a negative pressure, and is responsible for the expansion of the universe. It is modelled as a fluid, but it expands the universe by expanding the space between non-gravitationally bound objects. It is denoted by Λ and in currently accepted universe model, the Λ CDM model, it is dominant, causing the accelerated expansion of the universe observed. The observed composition ratio of the universe at present from Planck [8] is

Ω_i	Relative Quantity
Ω_m	0.3111 ± 0.0056
Ω_Λ	0.6847 ± 0.0073

using Planck TT,TE,EE+lowE+lensing+BAO 68% limits.

2 Problems with the Hot Big Bang

2.1 The Horizon Problem

2.2 The Flatness Problem

3 Cosmic Inflation

3.1 How Inflation solves the Horizon and Flatness Problems

3.2 Slow-Roll Inflation

4 Starobinsky Inflation

4.1 Starobinsky Potential

4.2 Results

5 α -attractors and Supergravity

5.1 Kinetic Lagrangian

5.2 E-model

5.3 T-model

5.4 The η -problem

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