List of parametrs used in Mcluster

- **Initial number of objects for each sub-population** n_1 , n_2 , ... with n_i = $n_{singles} + n_{binaries}$, where n_i is the total number of stars in the simulation

- **Primoridal binary fraction** fb_1 , fb_2 , ... with $n_{binaries} = fb_i \cdot n_i$

- **Initial density distribution** (could be different for sub-populations) *initialModel*

    - 0 - **Homogeneous sphere**
    - 1 - **Plummer**
    - 2 - **King**
    - 3 - **Subr**

- **King model parameter** (could be different for sub-populations) w_0 for w_0 in range of $(1.0 - 12.0)$ - DEFAULT $w_0 = 5.0$

- **Mass segregation parameter** S in range of $(0.0 - 1.0)$ - $S = 0.0 \rightarrow$ unsegregated, $S = 1.0 \rightarrow$ completely segregated
  For King model take $S = 0.99$ maximally - DEFAULT $S = 0.0$

- **Fractal dimensions** *fractal* (2^D children per parent) (could be different for sub-populations) - DEFAULT $fractal = 3.0$ (unfractal)

- **Virial ratio** *qvir*, $qvir > 0.5 \rightarrow$ expanding, $qvir > 0.5 \rightarrow$ equilibrium, $qvir > 0.5 \rightarrow$ collapsing - DEFAULT $q = 0.5$

- **Stellar mass function** *mfunc*

    - 0 - **equal masses equivalent to a set to parameter single mass**
        - A - DEFAULT single_mass= 1.0
    - 1 - **[Kroupa (2001)] IMF** - DEFAULT model
        - A - lower mass limit - default mlow m_low = 0.08
        - B - upper mass limit - default mup m_up = 100.00
    - 2 - **multi-power law (Subr)**
        - A - alpha slopes - DEFAULT alpha_imf = $[1.35, 2.35, 2.7, 0.0, 0.0]$
        - B - mass limits - DEFAULT mlim_imf = $[0.08, 0.5, 4.0, 100, 0.0, 0.0]$
    - 4 - **L3 IMF [Maschberger (2012)]**
        - A - lower mass limit - default mlow m_low = 0.08
        - B - upper mass limit - default mup m_up = 100.00
        - C - $\alpha-$slope for mass function - default alpha_L3 = 2.3
        - D - $\beta-$slope for mass function - default beta_L3 = 1.4
        - E - $\mu-$parameter - default mu_L3 = 0.2

- **Pairing of binary components** *pairing*

    - 0 - **random pairing**
    - 1 - **ordered pairing for components with masses** $M > 5M_\odot$
    - 2 - **random but separate pairing for components with masses** $M > 5M_\odot$
    - 3 - **uniform distribution of mass ratio** $(0.1 < q < 1.0)$ **for** $M > 5M_\odot$ **and random pairing for remaining** [Kiminki & Kobulnicky (2012); Sana et al., (2012); Kobulnicky et al., (2014)] - DEFAULT

- **Semi-major axis distribution** *adis*

    0 - uniform distribution in log(a), between amin and amax

    1 - lognormal distribution distribution for a

    2 - based on [**Kroupa (1995)**] period distribution

    3 - based on [**Kroupa (1995)**] period distribution for $M < 5M_\odot$; based on [**Sana et al.(2012); Oh, S., Kroupa, P., & Pflamm-Altenburg, J. (2015)**] period distribution for $M > 5M_\odot$ - DEFAULT distribution

    4 - flat uniform distribution in a ranging from amin to amax

    5 - based on [**Duquennoy & Mayor (1991)**] period distribution

    6 - uniform distribution in log(a), between amin and amax for $M < 5M_\odot$; **Sana et al. (2012)** period distribution for $M > 5M_\odot$ ALERT: *this will generate uniform distribution of mass ratio ($0.1 < q < 1.0$) for $M > 5M_\odot$ and also for $M < 5M_\odot$ indipendently from pairing !*

- **Eigenevolution** *eigen*

    0 - **Off** - DEFAULT

    1 - [**Kroupa (1995)**] eigenevolution for pre-main sequence short-period binaries and feeding algorithm
    ALERT: *This will set adis = 2 and pairing = 1 !*

    2 - **New eigenevolution and feeding algorithm [Kroupa (2013), rewieved in Belloni et al. (2017)]**
    ALERT: *This will set adis = 3 and pairing = 3 !*

- **Minimum binary semi-major axis** *amin* - defined in Solar radii [$R_\odot$]:

    A - $amin > 0 \rightarrow amin = amin$

    B - $amin < 0 \rightarrow amin = -amin(R_1 + R_2)$, where R_1, R_2 are the stellar radii of the composite stars

    C - DEFAULT $amin = -1.0$

- **Maximum binary semi-major axis** *amax* - defined in Solar radii [$R_\odot$]:

    A - $amax > 0 \rightarrow amax = amax$

    B - $amax < 0 \rightarrow amax = -amax \cdot 2.5 \cdot R_h/N$, where R_h is the total half-mass radius

    C - DEFAULT $amax = 50 \, [a.u.] = 10747.0$

- **Tidal field** *tf*

    0 - **no tidal field**

    1 - **point mass galaxy** - DEFAULT tidal field

- **Tidal radius** *rbar* defined in parsec [pc] - DEFAULT $rbar = 35.8$

- **Half mass radius** [pc] (for the whole system) *rh_mcl*

    A - if $rh\_mcl > 1.d9$ the model is not underfilled (standard King model, for Plummer model it will set $rh\_mcl = 0.1 \cdot rbar$)

    B - if $rh\_mcl < 1.d9$ the model is underfilled by a rplum factor, defined as $rplum = rbar/rh\_mcl$

    C - DEFAULT $rh\_mcl = 1.0$

- **Concentration radius parameter** *conc_pop* defined as Rh_i/Rh_1, the ratio between the half-mass radii of the i-th and the first generation (it will be skipped for single stellar population)

  - the first value is Rh_2/Rh_1 (the ratio between the half-mass radii of the second and the first generation)

  - the second value would be Rh_3/Rh_1 (the ratio between the half-mass radii of the third and the first generation)

  - DEFAULT conc_pop= 0.5

- **Potential energy evaluation** *potential_energy*

  0 **- potential energy evaluated as sum of gravitational potential for every object** (N^2)

  1 **- potential energy evaluated in spherical symmetry (N)** - DEFAULT energy evaluation

- **Age of population [Myr]** - DEFAULT $epoch = 0.0\ Myr$

- **Initial metallicity** - DEFAULT $zini = 0.001$

- **Initial integer number to start random number generator for Mcluster initial conditions** - DEFAULT seedmc = 10

- **Output format**

  0 **- single nbody.dat and binary nbody.dat files are generated.** The structure of those files are:
    * **single.dat** - mass $[M_\odot]$, x, y, z, Vx, Vy, Vz [N-body units], age, metallicty, index of the population
    * **binary.dat** - e, a [log10(Ro)], m1 $[M_\odot]$, m2 $[M_\odot]$, x, y, z, Vx, Vy, Vz [binary center of mass, N-body units], age, metallicity, index of the population

  1 **- standard dat.10 file is generated.** The structure of the file is: binaries at the beginning of the file (i.e. $2 \cdot NBIN$ lines with the binary individual masses,positions and velocities in the cluster frame) and remaining single star lines (with mass, positions and velocities in the cluster frame)

  - DEFAULT outputf = 0

- **Make energy check at end of Mclusters**

  0 **- off**

  1 **- on**

  - DEFAULT check_en = 1