Notes on curved-sky quadratic estimation

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This document supplements the release of the Planck 2018 CMB lensing [1] pipeline. It collects the formulae relevant to curved-sky quadratic estimators in the spin-weight, position-space formalism, including in particular estimator cross-responses and Gaussian noise biases between arbitrary pairs of quadratic estimators. JC: Document to be included with the pipeline release after submission of the revised L08. in progress

CONTENTS

JC: too many αsThe purpose of this document is to collect the formulae relevant to CMB lensing and other quadratic estimators in the spin-weight, position-space formalism.

The gradient (g) and curl (c) modes of definite parity of a complex spin-r field ${}_r\alpha(\hat n)$ (defined by the condition that it transforms under a rotation of angle ψ of the local axes through ${}_r\alpha(\hat n) \rightarrow e^{ir\psi}\,{}_r\alpha(\hat n)$) are defined through

$$g^r_{LM} = -\tfrac{1}{2}\left({}_{|r|}\alpha_{LM} + (-1)^r\,{}_{-|r|}\alpha_{LM}\right) \quad (0.1)$$
$$c^r_{LM} = -\tfrac{1}{2i}\left({}_{|r|}\alpha_{LM} - (-1)^r\,{}_{-|r|}\alpha_{LM}\right) \quad (0.2)$$

where ${}_{\pm r}\alpha_{LM} \equiv \int d^2n\, {}_{\pm r}\alpha(\hat n)\, {}_{\pm r}Y^*_{LM}(\hat n)$ (we adopt the convention, standard in CMB lensing, to write quadratic estimator multipoles with L, M and use ℓ, m for the CMB fields from which they are built). The inverse relation is

$${}_{\pm|r|}\alpha_{LM} = -(\pm)^r\left(g^r_{LM} \pm i c^r_{LM}\right). \quad (0.3)$$

Prior to projection onto gradient and curl modes, and prior to proper normalization, quadratic estimators can be written in the form

$${}_r\hat\alpha(\hat n) \equiv \left(\sum_{\ell m} w_\ell^{s_{\rm o}s_{\rm i}}\,{}_{s_{\rm i}}\bar X_{\ell m}\,{}_{s_{\rm o}}Y_{\ell m}(\hat n)\right) \cdot \left(\sum_{\ell m} w_\ell^{t_{\rm o}t_{\rm i}}\,{}_{t_{\rm i}}\bar X_{\ell m}\,{}_{t_{\rm o}}Y_{\ell m}(\hat n)\right) \quad (0.4)$$

where $s_{\rm i}, t_{\rm i}$ are input spins, $s_{\rm o}, t_{\rm o}$ outputs spins, and $w_\ell^{s_{\rm o}s_{\rm i}}, w_\ell^{t_{\rm i}t_{\rm o}}$ associated weights. Obviously, $s_o + t_o = r$, and by consistency with ${}_{-r}\alpha(\hat n) = {}_r\alpha^*(\hat n)$ the weights have symmetry $w_\ell^{-s_{\rm o}-s_{\rm i}} = (-1)^{s_{\rm o}+s_{\rm i}} w_\ell^{*s_{\rm o}s_{\rm i}}$.

The maps ${}_s\bar X_{lm}$ are the inverse variance filtered CMB maps; the filtered scalar temperature

$${}_0\bar X_{\ell m} = \bar T_{\ell m} \quad (0.5)$$

and filtered spin ± 2 Stokes polarization ${}_{\pm 2}P = \bar Q \pm i\bar U$,

$${}_{\pm 2}\bar X_{\ell m} = {}_{\pm 2}\bar P_{\ell m} = -\left(\bar E_{\ell m} \pm i\bar B_{\ell m}\right), \quad (0.6)$$

(for the purposes of the analytical calculations in this document) are isotropically related to the (beam-deconvolved) data maps ${}_sX^{\rm dat}$ through a matrix F,

$${}_s\bar X_{\ell m} \equiv \sum_{s_2 = 0,2,-2} F_\ell^{ss_2}\,{}_{s_2}X_{\ell m} \quad (0.7)$$

(isotropic approximation of $\bar X = \mathcal{B}^\dagger {\rm Cov}^{-1} X^{\rm dat}$ in the notation of Ref. [1]

$$\begin{aligned}\xi^{st}_+(\beta) &\equiv \left\langle e^{-is\alpha}\,{}_sX(\hat n_1)\,\left({}_tX(\hat n_2)e^{-it\gamma}\right)^*\right\rangle \\ \xi^{st}_-(\beta) &\equiv \left\langle \left(e^{-is\alpha}\,{}_sX(\hat n_1)\right)^*\,\left({}_tX(\hat n_2)e^{-it\gamma}\right)^*\right\rangle\end{aligned} \quad (0.8)$$

γ is the angle at $\hat n_1$ that aligns the local x-axis to the geodesic connecting $\hat n_1$ and $\hat n_2$ (with the axis pointing towards $\hat n_2$), β the angle between $\hat n_1$ and $\hat n_2$, and α is defined just as γ but at $\hat n_2$.[2, 3] JC: Mixups with n1 and n2 defs...fix this!

$$
\begin{aligned}
\xi_{+}^{st}(\beta) &= (+1)^{s}\sum_{L}\left(\frac{2L+1}{4\pi}\right)\left[C_L^{g^s g^t} + C_L^{c^s c^t} - i\left(C_L^{g^s c^t} + C_L^{c^s g^t}\right)\right] d_{st}^{L}(\beta)\\
\xi_{-}^{st}(\beta) &= (-1)^{s}\sum_{L}\left(\frac{2L+1}{4\pi}\right)\left[C_L^{g^s g^t} - C_L^{c^s c^t} - i\left(C_L^{g^s c^t} - C_L^{c^s g^t}\right)\right] d_{-st}^{L}(\beta)
\end{aligned}
\tag{0.9}
$$

The formulae exposed in this document can be derived through simple application of this formal relation,

$$
\begin{aligned}
&\sum_{m_1,m_2}\int d^2n_1\, {}_{s_1}Y_{\ell_1 m_1}(\hat n_1)\, {}_{s_2}Y_{\ell_2 m_2}(\hat n_1)\, {}_{r_1}Y_{LM}(\hat n_1)\int d^2n_2\, {}_{t_1}Y_{\ell_1 m_1}(\hat n_2)\, {}_{t_2}Y_{\ell_2 m_2}(\hat n_2)\, {}_{r_2}Y_{L'M'}(\hat n_2)\\
&= \delta_{LL'}\delta_{MM'}\frac{2\ell_1+1}{4\pi}\frac{2\ell_2+1}{4\pi}2\pi\int_{-1}^{1} d\beta\, d^{\ell_1}_{s_1,t_1}(\beta) d^{\ell_2}_{s_2 t_2}(\beta) d^{L}_{r_1 r_2}(\beta) \quad (\text{whenever } s_1+s_2+r_1 = 0 = t_1+t_2+r_2).
\end{aligned}
\tag{0.10}
$$

where $d^{\ell}_{mm'}$ are Wigner small d-matrices.

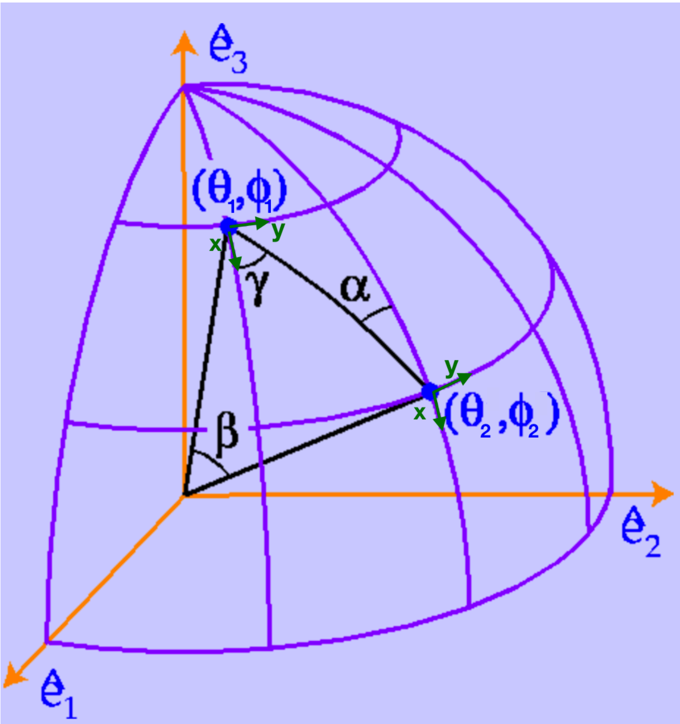

FIG. 1. The geometry and angles in Eq. (0.8), with the local axes in green. It holds $\alpha(\hat n_2,\hat n_1) = \pi - \gamma(\hat n_1,\hat n_2)$ and $\gamma(\hat n_2,\hat n_1) = \pi - \alpha(\hat n_1,\hat n_2)$. Figure originally from Wayne Hu tutorials, `http://background.uchicago.edu/~whu/tamm/webversion/node5.html`.

I. GAUSSIAN COVARIANCE CALCULATIONS

Q.E. noise covariance can be evaluated with a series of one-dimensional integrals as was first demonstrated by Ref. []. For two generic estimators as defined in Eq. (0.4), we now obtain their gradient (g) and curl (c) covariances with four integrals as follows. JC: describe here the matrix FFor an isotropy estimator ${}_r\hat\alpha$ let $s = (s_{\rm i}, s_{\rm o}, w^{s_{\rm i}s_{\rm o}})$ collectively describes the in and out spins and weight function of the left leg, and similarly with t for the right leg (with $s_{\rm o} + t_{\rm o} = r$). In the same way, let u and v describes another estimator ${}_{r'}\hat\alpha$ (with $u_{\rm o} + v_{\rm o} = r'$). Then, their Gaussian correlation functions are

$$
\boxed{\xi^{rr'}_{\pm}(\beta) = \xi^{\pm s,u}(\beta)\xi^{\pm t,v}(\beta) + \xi^{\pm s,v}(\beta)\xi^{\pm t,u}(\beta)}, \tag{1.1}
$$

where $\xi^{s,t}$ is

$$
\boxed{\xi^{s,t}(\beta) \equiv \sum_{\ell}\left(\frac{2\ell+1}{4\pi}\right) w_{\ell}^{s_{\rm o}s_{\rm i}} w_{\ell}^{*t_{\rm o}t_{\rm i}} \bar C_{\ell}^{s_{\rm i}t_{\rm i}} d^{\ell}_{s_{\rm o}t_{\rm o}}(\beta)} \tag{1.2}
$$

and $\bar C_\ell^{s_{\rm i}t_{\rm i}} \equiv \left\langle {}_{s_{\rm i}}\bar X_{\ell m}\, {}_{t_{\rm i}}\bar X^{*}_{\ell m}\right\rangle$. Projecting onto gradient and curl modes results in

$$
\begin{aligned}
\left\langle \hat g^{r}_{LM}\hat g^{*,r'}_{L'M'}\right\rangle\Big|_{\rm G.} &= \delta_{LL'}\delta_{MM'}\frac{1}{2}\Re\left[C_L^{rr'} + (-1)^r C_L^{-rr'}\right]\\
\left\langle \hat c^{r}_{LM}\hat c^{*,r'}_{L'M'}\right\rangle\Big|_{\rm G.} &= \delta_{LL'}\delta_{MM'}\frac{1}{2}\Re\left[C_L^{rr'} - (-1)^r C_L^{-rr'}\right]\\
\left\langle \hat g^{r}_{LM}\hat c^{*,r'}_{L'M'}\right\rangle\Big|_{\rm G.} &= \delta_{LL'}\delta_{MM'}\frac{1}{2}\Im\left[-C_L^{rr'} - (-1)^r C_L^{-rr'}\right]\\
\left\langle \hat c^{r}_{LM}\hat g^{*,r'}_{L'M'}\right\rangle\Big|_{\rm G.} &= \delta_{LL'}\delta_{MM'}\frac{1}{2}\Im\left[C_L^{rr'} - (-1)^r C_L^{-rr'}\right]
\end{aligned}
\tag{1.3}
$$

where

$$
C_L^{\pm rr'} \equiv 2\pi\int_{-1}^{1} d\mu\, d^{L}_{\pm rr'}(\mu)\xi^{rr'}_{\pm}(\beta) \tag{1.4}
$$

($\Re$ and $\Im$ stands for real and imaginary parts respectively). Ref. [1] calculates the covariance matrix based on these equations using the empirical, realisation dependent power spectra $\bar{C}_\ell^{s_i,t_i}$. A gradient-curl mode cross-covariance may be sourced by gradient-curl couplings in the inverse-variance filtered CMB fields (i.e., non-zero $C_\ell^{\bar{T}\bar{B}}$ or $C_\ell^{\bar{E}\bar{B}}$). In most relevant situations there is no such couplings and the gradient to curl and curl to gradient covariance vanish.

II. RESPONSE AND CROSS-RESPONSES CALCULATIONS

We now turn to the calculation of the response of the estimator to a source of anisotropy. Anisotropy can sometimes be parametrized at the level of the CMB maps, (for example for lensing), with

$$ {}_s\delta X(\hat{n}) = \sum_{a=\pm r} {}_a\alpha(\hat{n}) \left(\sum_{\ell m} R_\ell^{a,s} \, {}_sX_{\ell m} \, {}_{s-a}Y_{\ell m}(\hat{n}) \right) \tag{2.1} $$

for response kernel functions $R_\ell^{r,s}$. More generally, let the covariance of the CMB data respond as follows to a spin-r anisotropy source α:

$$ \delta \left\langle {}_sX(\hat{n}_1) \, {}_tX^*(\hat{n}_2) \right\rangle = \sum_{\ell m, a=\pm r} {}_a\alpha(\hat{n}_1) W_\ell^{a,st} \, {}_{s-a}Y_{\ell m}(\hat{n}_1) \, {}_tY^*_{\ell m}(\hat{n}_2) + W_\ell^{*a,ts} \, {}_sY_{\ell m}(\hat{n}_1) \, {}_{t-a}Y^*_{\ell m}(\hat{n}_2) \, {}_a\alpha^*(\hat{n}_2) \tag{2.2} $$

for some weights functions $W_\ell^{a,st}$. For map-level descriptions in Eq. (2.1) then holds

$$ W_\ell^{a,st} = R^{a,s} C_\ell^{st}. \tag{2.3} $$

However, Eq. (2.2) is more general. Section **??** lists weights functions of some relevant cases.

Let as before s,t denote collectively the QE spins and weight functions for an estimator ${}_r\hat{\alpha}(\hat{n})$ of spin $r = s_o + t_o$, and let r' be the spin of anisotropy source ${}_{r'}\beta(\hat{n})$ with covariance response kernel $W^{r'}$ as above. Let $\mathcal{R}_L^{g_r g_{r'}} \delta_{LL'}\delta_{MM'}$ be defined as the response of the gradient mode of α_{LM} to the gradient mode of $\beta_{L'M'}$, and similarly for the curl. It holds:

$$ \begin{aligned} \mathcal{R}_L^{g_r g_{r'}} &= \Re\left[R_L^{st,r'} + (-1)^{r'} R_L^{st,-r'} \right] \\ \mathcal{R}_L^{c_r c_{r'}} &= \Re\left[R_L^{st,r'} - (-1)^{r'} R_L^{st,-r'} \right] \\ \mathcal{R}_L^{g_r c_{r'}} &= \Im\left[-R_L^{st,r'} + (-1)^{r'} R_L^{st,-r'} \right] \\ \mathcal{R}_L^{c_r g_{r'}} &= \Im\left[R_L^{st,r'} + (-1)^{r'} R_L^{st,-r'} \right] \end{aligned} \tag{2.4} $$

where

$$ R_L^{st,r'} = 2\pi \int_{-1}^{1} d\mu \, d^L_{rr'}(\mu) \sum_{\tilde{s}_i, \tilde{t}_i = 0,2,-2} \left[\xi^{s_o s_i \tilde{s}_i}(\mu) \psi^{t_o t_i \tilde{t}_i \tilde{s}_i, r'}(\mu) + \xi^{t_o t_i \tilde{t}_i}(\mu) \psi^{s_o s_i \tilde{s}_i \tilde{t}_i, r'}(\mu) \right] \tag{2.5} $$

and

$$ \begin{aligned} \xi^{s_o s_i \tilde{s}_i}(\mu) &\equiv \sum_\ell \left(\frac{2\ell+1}{4\pi} \right) w_\ell^{s_o s_i} F_\ell^{s_i \tilde{s}_i} d^\ell_{s_o, \tilde{s}_i}(\mu) \\ \psi^{s_o s_i \tilde{s}_i \tilde{t}_i, r'}(\mu) &\equiv \sum_\ell \left(\frac{2\ell+1}{4\pi} \right) w^{s_o s_i} F_\ell^{s_i \tilde{s}_i} W_\ell^{*-r', -\tilde{t}_i \tilde{s}_i} d^\ell_{s_o, -\tilde{t}_i + r'}(\mu) \end{aligned} \tag{2.6} $$

Again, in most relevant cases, the gradient to curl and curl to gradient responses do vanish. If there is a unique source of anisotropy, properly normalized gradient and curl estimators are then given by $\hat{g}^r_{LM}/\mathcal{R}_L^{g_r g_r}$ and $\hat{c}^r_{LM}/\mathcal{R}_L^{c_r c_r}$.

III. DERIVATION OF OPTIMAL QE WEIGHTS

Optimal (in the sense of minimal Gaussian variance) QE weights are easily gained from the representation in Eq. 2.2 of the anisotropy. Let the CMB likelihood gradients be

$$ {}_{\pm r}\hat{\alpha}(\hat{n}) = \frac{\delta}{\delta_{\mp r}\alpha(\hat{n})} - \frac{1}{2}\, {}_{s_1}X \mathrm{Cov}^{-1}_{s_1 s_2}\, {}_{s_2}X \Big|_{\alpha\equiv 0} \tag{3.1} $$

where $\mathrm{Cov}_{s_1 s_2}(\hat{n},\hat{n}') \equiv \langle {}_{s_1}X(\hat{n})\, {}_{s_2}X(\hat{n}')\rangle$, and where ${}_r\alpha(\hat{n})$ and ${}_{-r}\alpha(\hat{n})$ are treated as independent variables. Using Eq. (2.2) and comparing to Eq. (0.4), we find

$$ \boxed{w_\ell^{st} = \delta_{st} \text{ (1st leg)} \quad w_\ell^{-s+r,t} = 2W_\ell^{-r,-st} \text{ (2nd leg)}} \tag{3.2} $$

JC: why 2 again? JC: FIXME: The right expression is

$$ {}_r\hat{g}(\hat{n}) = \sum_s {}_{-s}\bar{X}(\hat{n}) \cdot \left(2W_\ell^{-r,st}\, {}_t\bar{X}_{\ell m}\, {}_{s+r}Y_{\ell m}(\hat{n})\right) \tag{3.3} $$

where $\bar{X}$ has the $(0,2,-2)$ elements (note the additional factor of 2! in pol w.r.t. to naive spin defs.)

$$ \begin{pmatrix} \bar{T} \\ -\frac{1}{2}\left(\bar{E}+i\bar{B}\right) \\ -\frac{1}{2}\left(\bar{E}-i\bar{B}\right) \end{pmatrix} \tag{3.4} $$

Factor of 2 in front of W comes from 2 $\delta/\delta\ {}_{-r}\alpha(\hat{n})$ to get $d/dre + d/dim$ (?).

IV. EXAMPLES

Examples include:

A. Lensing

The source of anisotropy is the spin-1 deflection field ${}_1\alpha(\hat{n})$, with linear response (see Ref. [4]) $\delta_s X(\hat{n}) = -\frac{1}{2}\alpha_1(\hat{n})\bar{\eth}_s X(\hat{n}) - \frac{1}{2}\alpha_{-1}(\hat{n})\eth\, {}_sX(\hat{n})$ where $\eth$ and $\bar{\eth}$ are the spin raising and spin lowering operator respectively. From the action of these operator on the spherical harmonics follow immediately

$$ \begin{aligned} R_\ell^{-1,s} &= -\frac{1}{2}\sqrt{(l-s)(l+s+1)} \\ R_\ell^{1,s} &= +\frac{1}{2}\sqrt{(l+s)(l-s+1)} \end{aligned} \tag{4.1} $$

B. Modulation

The anisotropy source is a scalar, with response $\delta_s X(\hat{n}) = {}_0\alpha(\hat{n})_s X(\hat{n})$, hence

$$ R_\ell^{0,s} = 1 \tag{4.2} $$

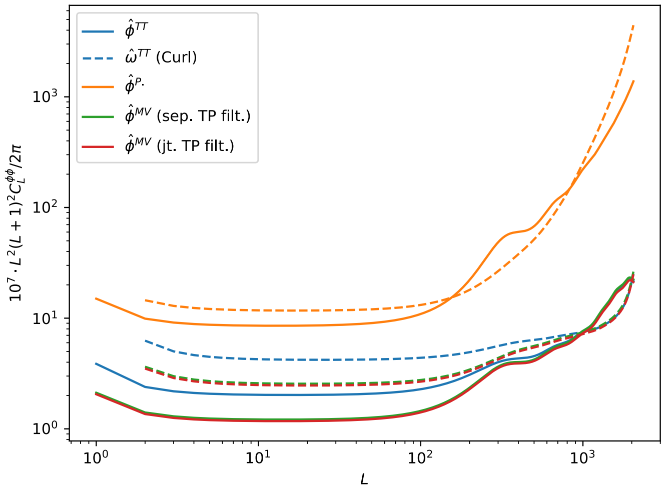

FIG. 2. Lensing gradient and curl reconstruction noise levels for a *Planck*-like experiment.

C. Polarization rotation

This is relevant in the case of systematic polarization angle miscalibration, or within more speculative ideas including cosmic birefringence. The observed polarization is rotated according to ${}_{\pm 2}P$ is $e^{\mp 2i\,{}_0\alpha}\, {}_{\pm 2}P$. Hence,

$$ R_\ell^{0,\pm 2} = \mp 2i \tag{4.3} $$

D. Point sources

Point sources in temperature (S^2, see Ref. [5]): here anisotropy is sought of the form $\delta\,\langle T(\hat{n})\,T(\hat{n}')\rangle = \delta_{\hat{n}\hat{n}'}S^2(\hat{n})$. Hence,

$$ W_\ell^{r,st} = \frac{1}{4}\delta_{r0}\delta_{s0}\delta_{t0} \tag{4.4} $$

E. Noise variance map anisotropies

This is conceptually the same as point sources but acting on beam-convolved maps

$$ W_\ell^{r,st} = \frac{1}{4}\delta_{r0}\delta_{s0}\delta_{t0}\frac{1}{b_\ell^2} \tag{4.5} $$

[1] N. Aghanim *et al.* (Planck), “Planck 2018 results. VIII. Gravitational lensing,” (2018), arXiv:1807.06210 [astro-ph.CO].

[2] Gayoung Chon, Anthony Challinor, Simon Prunet, Eric Hivon, and Istvan Szapudi, “Fast estimation of polarization power spectra using correlation functions,” Mon. Not. Roy. Astron. Soc. **350**, 914 (2004), arXiv:astro-ph/0303414 [astro-ph].

[3] Anthony Challinor and Antony Lewis, “Lensed CMB power spectra from all-sky correlation functions,” Phys. Rev. **D71**, 103010 (2005), arXiv:astro-ph/0502425 [astro-ph].

[4] Anthony Challinor and Gayoung Chon, “Geometry of weak lensing of CMB polarization,” Phys. Rev. **D66**, 127301 (2002), arXiv:astro-ph/0301064 [astro-ph].

[5] Stephen J. Osborne, Duncan Hanson, and Olivier Doré, “Extragalactic Foreground Contamination in Temperature-based CMB Lens Reconstruction,” JCAP **1403**, 024 (2014), arXiv:1310.7547 [astro-ph.CO].