

Revised Bolometric Corrections and Interstellar Extinction Coefficients for the ACS and WFPC2 Photometric Systems

LÉO GIRARDI

Osservatorio Astronomico di Padova—INAF, Padova, Italy

JULIANNE DALCANTON AND BENJAMIN WILLIAMS

University of Washington, Seattle, WA

ROELOF DE JONG

Space Telescope Science Institute, Baltimore, MD

CARME GALLART AND MATTEO MONELLI

Instituto de Astrofísica de Canarias, La Laguna, Tenerife, Spain

MARTIN A. T. GROENEWEGEN

Katholieke Universiteit Leuven, Belgium

JON A. HOLTZMAN

New Mexico State University, Las Cruces, NM

KNUT A. G. OLSEN

Gemini Science Center, NOAO, Tucson, AZ

ANIL C. SETH

Harvard Center for Astrophysics, Boston, MA

AND

DANIEL R. WEISZ

University of Minnesota, Minneapolis, MN

(FOR THE ANGST/ANGRRR COLLABORATION)

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ABSTRACT. We present extensive tables of bolometric corrections and interstellar extinction coefficients for the WFPC2 and ACS (both WFC and HRC) photometric systems. They are derived from synthetic photometry applied to a database of spectral energy distributions covering a large range of effective temperatures, surface gravity, and metal content. Carbon stars are also considered. The zero points take into consideration the new high-accuracy Vega fluxes from Bohlin. These tables are employed to transform Padova isochrones into WFPC2 and ACS photometric systems using interstellar extinction coefficients on a star-to-star basis. All data are available either in tabular form or via an interactive web interface in the case of the isochrones. Preliminary tables for the WFC3 camera are also included in the database.

Online material: color figures

1. INTRODUCTION

The Advanced Camera for Surveys (ACS) aboard the *Hubble Space Telescope* (*HST*) has provided some of the deepest photometric optical images ever taken. At the same time, the Wide Field and Planetary Camera 2 (WFPC2) camera aboard *HST* has a rich photometric data archive dating back more than a

decade. Both cameras have taken many images of galaxies within ~ 5 Mpc, providing reliable photometry for individual stars within these nearby galaxies. This resolved stellar photometry contains detailed information about the history of star formation and chemical enrichment. Extracting this information requires stellar evolution isochrones that track the photometric properties

of stars of different ages and metallicities. Only with precise models can these rich samples of stellar photometry provided by ACS and WFPC2 be translated into coherent histories of star formation and chemical evolution.

Given the extreme importance of the stellar photometric data provided by ACS and WFPC2, many authors have described the conversion from theoretical stellar models into these systems. For the WFPC2 system, transformations and isochrones have been provided by Holtzman et al. (1995), Chiosi et al. (1997), Salasnich et al. (2000), Girardi et al. (2002), Lejeune (2002), and Dotter et al. (2007). For ACS, bolometric correction (BC) tables and isochrones have been provided at <http://pleiadi.oapd.inaf.it> since 2003, as an addition to the Girardi et al. (2002) database, and later by Bedin et al. (2005) and Dotter et al. (2007).

With this paper, our goal is to provide revised and homogeneous tables of BCs and interstellar extinction coefficients in the ACS and WFPC2 systems, updating as much as possible the files distributed by Girardi et al. (2002) in precedence. The motivation for an overall revision of the BC tables reside mainly on the definition of better spectrophotometric standards by Bohlin and collaborators (Bohlin 2007, and references therein), and more complete libraries of stellar spectra, whereas filter transmission curves for these instruments remain essentially the same as in the original papers by the instrument teams. A consistent set of interstellar extinction coefficients for these systems would be particularly important because the unique resolution of *HST* allows useful optical photometry to be obtained even for heavily reddened regions (with say $A_V \gtrsim 3$; see, e.g., Recio Blanco et al. 2005), for which interstellar extinction coefficients being used so far may not be accurate enough.

The BC and interstellar extinction tables will be used to transform theoretical isochrones to these systems, and will be extensively applied in the interpretation of ANGST/ANGRRR data.¹

We proceed as follows: § 2 describes the steps involved in the derivation of bolometric corrections via synthetic photometry, from the assembly of a spectral library to the definition of zero points. Section 3 describes the interstellar extinction coefficients. The results are used in § 4 to convert theoretical isochrones into the ACS and WFPC2 systems, which are made publicly available. They are briefly illustrated in § 5, by means of a few comparisons with observations.

2. BOLOMETRIC CORRECTIONS

Given the spectral fluxes at the stellar surface, F_λ , bolometric corrections for any set of filter transmission curves S_λ ² are given by (see Girardi et al. 2002 for details)

¹ ANGST (ACS Nearby Galaxy Survey) is a Treasury *HST* proposal for the imaging of a complete sample of galaxies up to a distance of 3.5 Mpc. Together with archival data re-reduced by the ANGRRR (Archive of Nearby Galaxies—Reduce, Recycle, Reuse) collaboration, it will allow the measurement of their star formation histories with a resolution of 0.3 dex in $\log(\text{age})$, up to the oldest ages.

$$\text{BC}_{S_\lambda} = M_{\text{bol},\odot} - 2.5 \log[4\pi(10\text{pc})^2 F_{\text{bol}}/L_\odot] + 2.5 \log\left(\frac{\int_{\lambda_1}^{\lambda_2} \lambda F_\lambda 10^{-0.4A_\lambda} S_\lambda d\lambda}{\int_{\lambda_1}^{\lambda_2} \lambda f_\lambda^0 S_\lambda d\lambda}\right) - m_{S_\lambda}^0, \quad (1)$$

where f_λ^0 represents a reference spectrum (incident on the Earth's atmosphere) that produces a known apparent magnitude $m_{S_\lambda}^0$, $F_{\text{bol}} = \int_0^\infty F_\lambda d\lambda = \sigma T_{\text{eff}}^4$ is the total emerging flux at the stellar surface, and A_λ is the interstellar extinction curve in magnitude units. Once BC_{S_λ} are computed, stellar absolute magnitudes follow from

$$M_{S_\lambda} = M_{\text{bol}} - \text{BC}_{S_\lambda}, \quad (2)$$

where

$$\begin{aligned} M_{\text{bol}} &= M_{\text{bol},\odot} - 2.5 \log(L/L_\odot) \\ &= M_{\text{bol},\odot} - 2.5 \log(4\pi R^2 F_{\text{bol}}/L_\odot). \end{aligned} \quad (3)$$

As in Girardi et al. (2002), we adopt $M_{\text{bol},\odot} = 4.77$, and $L_\odot = 3.844 \times 10^{33} \text{ erg s}^{-1}$ (Bahcall et al. 1995).

Notice that equation (1) uses photon-count integration instead of energy integration, as appropriate to most modern photometric systems that use photon-counting detectors instead of energy-amplifiers. The equation also includes the interstellar extinction curve A_λ , which will be discussed later in § 3. $A_\lambda = 0$ is assumed for the moment.

2.1. A Library of Spectral Fluxes

Girardi et al. (2002) assembled a large library of spectral fluxes, covering wide ranges of initial metallicities ($[M/H]$ from -2.5 to $+0.5$), T_{eff} (from 500 to 50000 K), and $\log g$ (from -2 to 5). This grid is wide enough to cover most of the stellar types that constitute the bulk of observed samples in resolved galaxies. The grid is composed of ATLAS9 “NOVER” spectra from Castelli et al (1997; see also Bessell et al. 1998) for most of the stellar types, Allard et al. (2000a) for M, L, and T dwarfs, Fluks et al. (1994) empirical spectra for M giants, and, finally, pure blackbody spectra for stars exceeding 50000 K. Since then, we have updated part of this library, namely:

1. As part of the TRILEGAL project (Girardi et al. 2005), we have included spectra for DA white dwarfs from Finley et al. (1997) and Homeier et al. (1998) for T_{eff} between 100,000 and 5000 K.

2. We now use Castelli & Kurucz (2003) ATLAS9 “ODF-NEW” spectra, which incorporate several corrections to the atomic and molecular data, especially regarding the molecular lines of CN, OH, SiO, H₂O and TiO, and the inclusion of HI – H⁺ and HI – H⁺ quasi-molecular absorption in the UV.

² Throughout this paper, S_λ is the total system throughput, i.e., including the telescope, camera, and filters.

These new models are considerably better than the previous NOVER ones, especially in the UV region of the spectrum, and over most of the wavelength range for T_{eff} between ~ 4500 and 3500 K (see Figs. 1 to 3 in Castelli & Kurucz 2003).

3. We have also incorporated in the library the Loidl et al. (2001) spectral fluxes of solar-metallicity carbon (C) stars derived from static model atmospheres. They are computed for T_{eff} in the range from 2600 to 3600 K, and for C/O = 1.1 and 1.4. The C/O = 1.1 spectra are now used for all C/O > 1 models, replacing the Fluks et al. (1994) spectra of M giants that were previously (and improperly) used for all cool giants including C-rich ones.

Regarding this last aspect, it is to be noticed that all Padova isochrones published since 2001 (Marigo & Girardi 2001; Cioni et al. 2006; Marigo et al. 2008) do take into consideration the third dredge-up events and hence the conversion to a C-type phase during the TP-AGB evolution. This phase, characterized by the surface C/O > 1, has to be represented with either proper C-rich spectra or, alternatively, with empirical T_{eff} -color relations as in Marigo et al. (2003). The use of Loidl et al. (2001) spectra frees us definitely from using empirical relations for C stars. Note that such empirical relations would not even be available for the ACS and WFPC2 filters we are discussing. In fact, empirical T_{eff} -color relations for C stars are usually limited to red and near-IR filters, and mainly to *IJHK*.

Although the Loidl et al. (2001) models are quite limited in their coverage of the parameter space of C stars (i.e., [M/H], C/O, T_{eff} , and $\log g$) they represent a useful first approach toward obtaining realistic colors for them. Work is in progress to extend significantly the parameter space of such models (Aringer et al. 2008, in preparation).

2.2. Filters and Zeropoints

As can readily be seen in equation (1), in the formalism adopted here, the photometric zero points are fully determined by a reference spectral energy distribution f_{λ}^0 , which provides a given set of reference magnitudes $m_{S_{\lambda}}^0$.

In this work, we adopt the latest Vega spectral energy distribution from Bohlin (2007) as f_{λ}^0 . This Vega spectrum represents a great improvement over previous spectra distributed and used in the past. It is the result of an effective effort using Space Telescope Imaging Spectrograph (STIS) to provide a set of stellar spectrophotometric standards with an absolute calibration to better than 1% at near-UV-to-blue wavelengths. The best-fitting ATLAS9 model for Vega in this spectral range is then used to complement its spectral energy distribution at red and near-IR wavelengths. From Bohlin (2007), one can conclude the new Vega spectrum is likely accurate to about 1–2 percent over the complete 3200–10000 Å interval.

Having adopted Bohlin’s (2007) Vega spectrum, we have just to specify the Vega magnitudes in ACS and WFPC2 filters to completely define our zero points. Following Holtzman et al.’s

(1995) definition of the WFPC2 synthetic system, we define that Vega has the same magnitudes in WFPC2 filters as in the Johnson-Cousins filters that are closest to them in wavelength. This choice means a magnitude of 0.02 for all filters from F336W to F450W, 0.03 for F555W and F569W, 0.035 from F606W to F850LP, but for F702W which has magnitude 0.039. For filters blueward of F300W, Vega magnitudes are assumed to be zero.

Then, for ACS we adopt the same definition as Sirianni et al. (2005), that Vega has zero magnitudes in all ACS filters. For a question of completeness, we provide tables also for the AB and ST magnitude systems.

These are, of course, synthetic definitions of WFPC2 and ACS systems, that may present small systematic offsets with respect to the real systems (such as the flight WFPC2 system defined by Holtzman et al. 1995) to which observations are commonly transformed.

Finally, the filter transmission curves S_{λ} —including the telescope and camera throughputs—for WFPC2 are created with the STSDAS package SYNPHOT v3.7 using the most up-to-date throughput and sensitivity reference files available on 2008 January 15. For WFPC2/F170W and WFPC2/F218W, instead, they are taken from Holtzman et al. (1995). All of the wide filters are considered. ACS filter throughputs are from Sirianni et al. (2005). They are almost identical to those provided by SYNPHOT v3.7.

2.3. Behavior of Bolometric Corrections

Figure 1 shows how BCs depend on T_{eff} for a series of ACS and WFPC2 filters. In practice, we illustrate it for a sequence of dwarfs of $\log g = 5$, and for a sequence of subgiants and giants (of M type at the coolest T_{eff}) following the relation $T_{\text{eff}} = 3250 + 500 \log g$ (in cgs units), which roughly describes the position of the Hayashi track in the Hertzsprung-Russell diagram of low-mass stars with near-solar metallicity. The central part of the T_{eff} interval reflects the behavior of ATLAS9 synthetic spectra. The small discontinuities in the figure occur as we change from one spectral library to another, namely:

1. The wiggles at $\log T_{\text{eff}} \approx 3.54$ occur when we pass from ATLAS9 to Fluks et al. (1994) spectra for M giants, and to Allard et al. (2000a) spectra for M dwarfs.
2. The step at $\log T_{\text{eff}} \approx 3.27$ occurs when we change between the “AMES” and “AMES-dusty” among Allard et al. (2000a) spectra. The latter include dust formation in the outer atmosphere and are computed for solar metallicity only (see Chabrier et al. 2000; Allard et al. 2000b, 2001 for more details).
3. A second, less pronounced step at $\log T_{\text{eff}} \approx 4.75$ occurs when we start using blackbody spectra instead of ATLAS9.

Among these discontinuities, the most important are the wiggles at $\log T_{\text{eff}} \approx 3.54$ because they occur at the T_{eff} in which many stars of Local Group galaxies are routinely observed in optical to near-IR passbands. In order to limit the effect of such

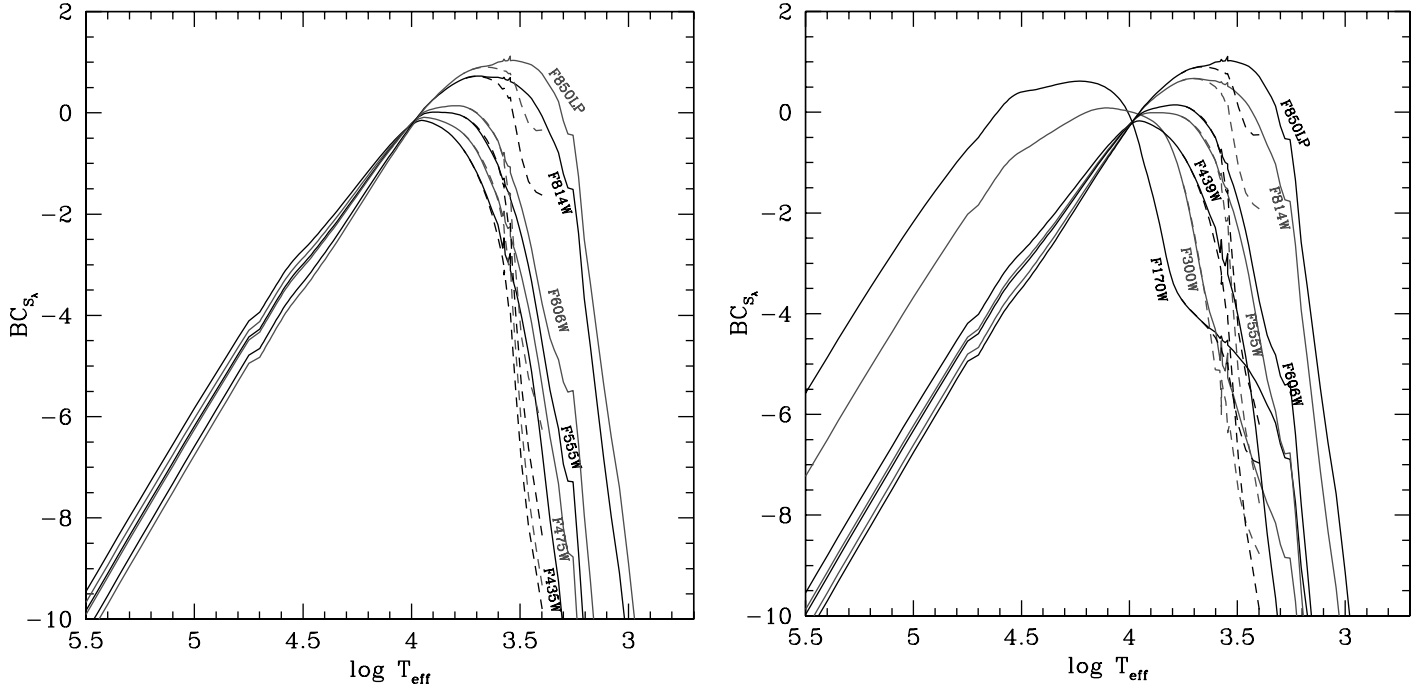


FIG. 1.— BC_{S_λ} as a function of T_{eff} , for both dwarfs (solid lines) and giants (dashed lines), for some of the ACS WFC and WFC2 filters. See the electronic edition of the PASP for a color version of this figure.

wiggles in the final isochrones, we adopt a smooth transition between the different sources of BC, for both dwarfs and giants, over the $\log T_{\text{eff}}$ interval from 3.542 and 3.590. This is the reason why isochrones using these transformations (see Fig. 5 later for an example) do not present significant discontinuities.

It is also interesting to note, in Figure 1, the anomalous behavior of F170W in WFC2 which for cool stars provides brighter magnitudes than redder filters. This behavior is a consequence of its red leak (see also Holtzman et al. 1995). A similar effect is also present, but much less pronounced, in the F300W filter.

3. INTERSTELLAR EXTINCTION COEFFICIENTS

As already mentioned, equation (1) can be used together with an interstellar extinction curve, A_λ , to derive bolometric corrections already including the effect of interstellar extinction. Alternatively, we can derive relative interstellar extinction coefficients, A_{S_λ}/A_V , by simply using

$$A_{S_\lambda}/A_V = [BC_{S_\lambda}(0) - BC_{S_\lambda}(A_V)]/A_V, \quad (4)$$

where the 0 stands for the $A_\lambda = 0$ case.

The interstellar extinction curve A_λ we use in equation (1) is taken from Cardelli et al. (1989) with $R_V = A_V/E_{B-V} = 3.1$, and integrated with the O'Donnell (1994) correction for $909 < \lambda/\text{\AA} < 3030$. Other extinction curves and R_V values can be easily implemented whenever necessary (e.g., Vanhollebeke et al. 2008). Our calculations are essentially identical to

those by Holtzman et al. (1995) and Sirianni et al. (2005), but are now performed for a more extended range of stellar parameters and extinction values.

Before proceeding, it is interesting to note that the two parameters in Cardelli et al. (1989) interstellar extinction curve, A_V and R_V , refer to the Johnson's B and V bands. However, this does not imply that synthetic realizations of the Johnson system will recover the same values of R_V and A_V when the extinction curve is applied. As examples, we find that for a G2V star the Bessell (1990) version of the Johnson system provide $A_{V,\text{Bessell}} = 1.006A_V$ and $R_{V,\text{Bessell}} = 3.457$, whereas the Maíz Apellániz (2006) version gives essentially the same numbers, with differences appearing just at the fifth decimal. These numbers change somewhat as a function of spectral type and A_V itself (as will be illustrated below for *HST* systems). Therefore, A_V should be considered as just “the parameter that describes the total amount of interstellar extinction in Cardelli et al. (1989) law,” rather than a precise physical measure of the V -band extinction. A similar comment applies also to R_V : in the context of this paper it is “the parameter that describes the shape of the interstellar extinction curve,” rather than the total-to-selective interstellar extinction of Johnson-like systems.

The interstellar extinction coefficients derived from equation (4) depend on the F_λ under consideration, and hence on parameters such as T_{eff} , $\log g$, and $[M/H]$. Figure 2 illustrates the variation of A_{S_λ}/A_V as a function of T_{eff} , for several ACS Wide Field Channel (WFC) filters, and for both dwarfs (contin-

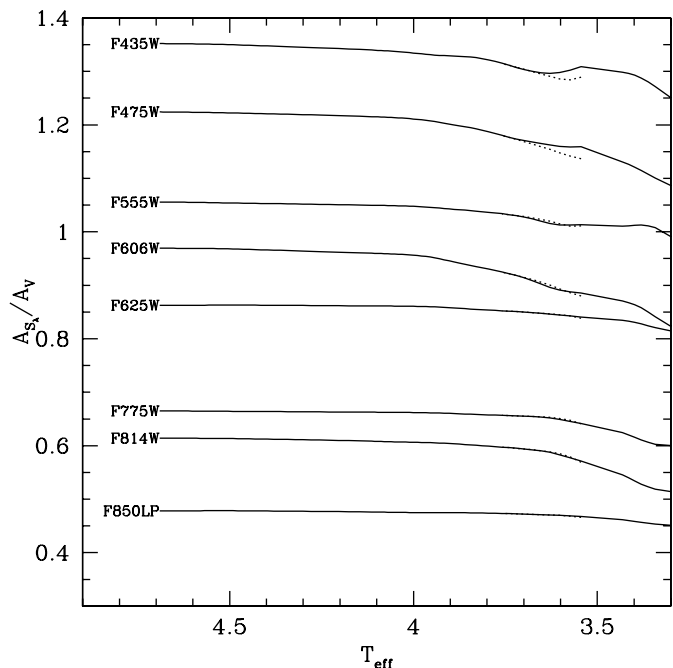


FIG. 2.—Variation of A_{S_λ}/A_V as a function of T_{eff} , for both dwarfs (solid lines) and giants (dashed lines), for some of the ACS and WFC filters.

ous lines) and giants (dashed lines). Giants are again defined by the relation $T_{\text{eff}} = 3250 + 500 \log g$. The figure shows that A_{S_λ}/A_V can change by as much as ~ 0.3 between very cool and very hot stars, especially for the bluest (F435W, F475W) and widest (F606W, F814W) ACS filters. This variation will inevitably change the morphology of a reddened isochrone in the color-magnitude diagram (CMD), as compared to the unreddened case. This is actually the main reason for taking star-to-star interstellar extinction into consideration.

A_{S_λ}/A_V variations are more sizeable in the regime of cool temperatures and blue filters, which, however, is not the regime targeted by most *HST* observations of stellar populations. The most relevant T_{eff} interval for A_{S_λ}/A_V variations is instead between ~ 4000 and $15,000$ K, which comprehends the main sequence turn-off, horizontal branch, and tip of the red giant branch (TRGB) of old metal-poor populations. Inside this interval, A_{S_λ}/A_V variations amount to ~ 0.15 , and hence become potentially important—causing ~ 0.05 mag effects in colors and magnitudes—at A_V higher than ~ 3 mag.

Another effect that can be accessed by our calculations is the “Forbes effect,” or interstellar nonlinear heterochromatic extinction (Forbes 1842). It consists of the variation of the relative extinction A_{S_λ}/A_V as the total extinction A_V increases, and results from extinction being more effective for the spectral regions with the higher flux. The A_{S_λ}/A_V variations are expected to occur whenever the flux F_λ varies within the wavelength interval of a passband, and they are obviously higher for wider passbands and at higher total extinction. This effect was originally conceived to describe atmospheric extinction, but has

been shown to apply also for interstellar extinction by Grebel & Roberts (1995).

Figure 3 illustrates the Forbes effect for some of the ACS and WFPC2 passbands. It plots A_{S_λ}/A_V as a function of A_V (solid lines), for a few values of T_{eff} and for the $0 < A_V < 12$ interval. It can be easily noticed that A_{S_λ}/A_V decreases with A_V , and that this decrease is more marked for wider filters. In fact, A_{S_λ}/A_V is almost constant for the narrow F660N filter. Among the wide filters, the hottest spectra ($\sim 40,000$ K) present the most moderate variations in A_{S_λ}/A_V , which is of approximately $0.05 A_V^{-1}$ for the most widely used filters F555W and F814W.

Figure 4 instead plots A_{S_λ}/A_{F555W} as a function of A_{F555W} , in the ACS or WFC system. This changes somewhat the description of the Forbes effect, which becomes increasing for F435W and F628N and decreasing for F814W. Actually, the most striking manifestation of the Forbes effect is expected to be on the stellar colors, and this will be independent of the wavelength taken as a reference for the total extinction.

Finally, we note that for stars cooler than 3000 K and for the bluest ACS/WFC filters (for instance F435W) the Forbes effect becomes dramatic already at $A_V \gtrsim 4$. These cases, however, are of little practical interest because observations of heavily reddened cool stars would hardly be done with such blue filters.

Figure 5 presents a practical application of the star-to-star interstellar extinction coefficients derived in this work, as compared to the usual approximation of applying a single set of coefficients for all stars and independently of the actual amount of total extinction. In the figure, one can notice that a set of

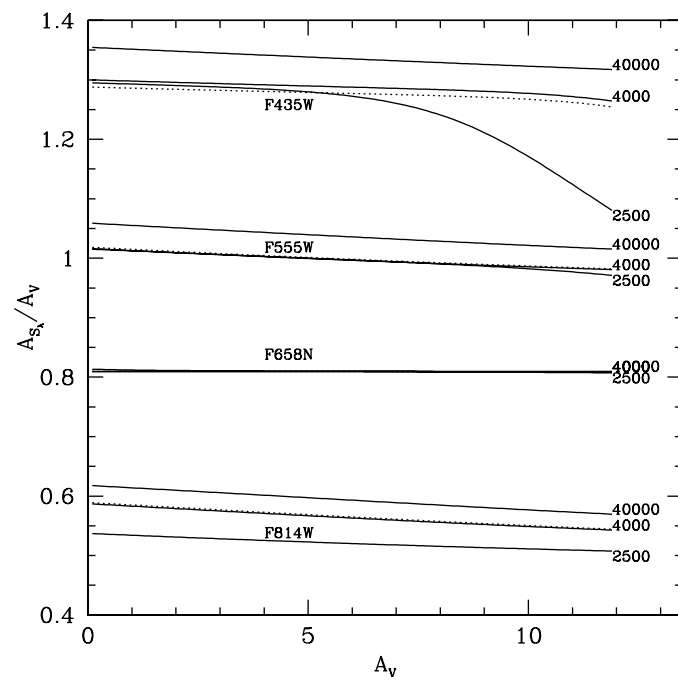


FIG. 3.—Forbes effect for a few ACS filters, for dwarfs of $T_{\text{eff}} = 2500$ K, 4000 K and $40,000$ K (solid lines), and giants with $T_{\text{eff}} = 4000$ K (dashed lines). The T_{eff} values in K are indicated at the right extremity of their curves.

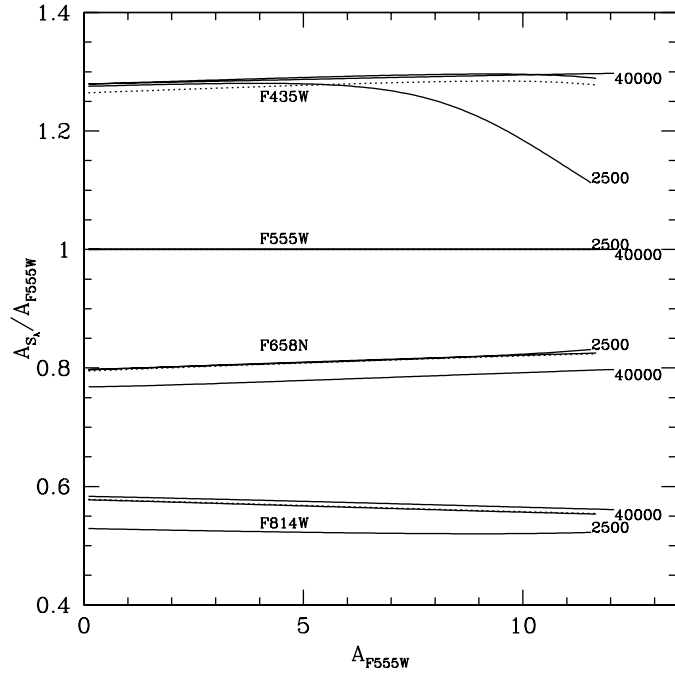


FIG. 4.—Same as Fig. 3, but now as a function of A_{F555W} , in the ACS or WFC system.

isochrones extinguished by $A_V = 6$ mag in the correct (star-to-star) way, is narrower in color (by about 0.3 mag in $F475W-F84W$) than an isochrone extinguished by the same A_V but using a single set of coefficients derived from a low-extinction yellow dwarf. Moreover, using the right interstellar extinction coefficients produces brighter red giants than in the case of constant coefficients. This latter effect is striking in the comparison between the AGB sequences at the top-right corner of Figure 5.

Although the $A_V = 6$ mag case is rather extreme, it well illustrates the effect of star-to-star interstellar extinction. Whether this effect is important depends on the photometric accuracy of the data and models available, and on the passbands being considered. A very approximate rule is that, if one aims at reproducing data with an accuracy better than 0.05 mag in optical bands, star-to-star interstellar extinction should start being considered already at $A_V > 2$ mag. A more consistent and safer rule, however, is to *always* take star-to-star extinction into consideration.

4. APPLICATION TO ISOCHRONES AND DATA RELEASE

Extensive tables with BC and interstellar extinction coefficients have been computed and are provided in the static web repository <http://stev.oapd.inaf.it/dustyAGB07>. The BC tables contain entries for all stars in our spectral library. The extinction coefficients instead have been computed for solar metallicity only. In fact, the effective temperature, and to a lower extent $\log g$, is the main parameter driving the changes in A_{S_λ}/A_V .

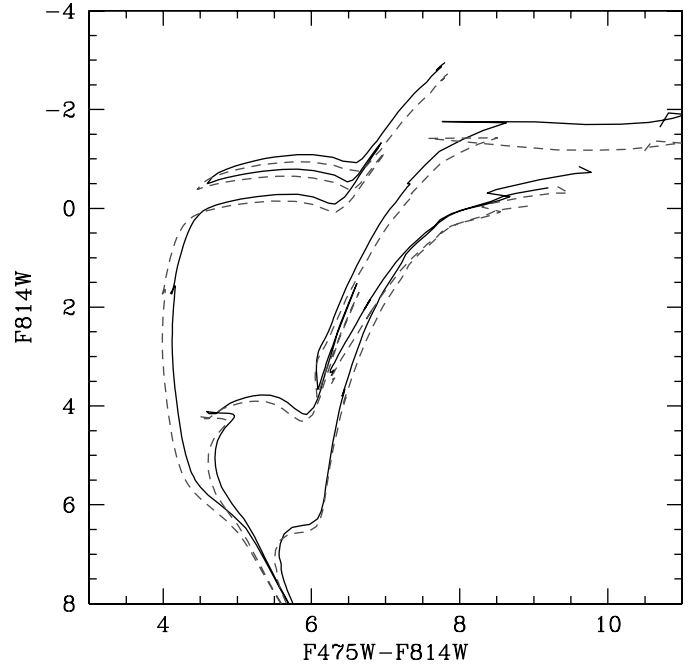


FIG. 5.—Set of $Z = 0.008$ isochrones from Marigo et al. (2008) in the ACS WFC system, extinguished by $A_V = 6$ mag according to two different prescriptions: adopting the star-to-star interstellar extinction coefficients as described in this paper (solid lines) and adopting the coefficients derived from a yellow dwarf at the regime of low interstellar extinction (dashed lines). The isochrone ages are, from top left to bottom right, 0.1, 1, and 10 Gyr. See the electronic edition of the PASP for a color version of this figure.

We have verified that A_{S_λ}/A_V values change very little as a function of $[M/H]$: as an example, for a dwarf star of $T_{\text{eff}} = 4000$ K extinguished by $A_V = 1$, A_{F475W}/A_V values differ by just 0.17% between the $[M/H] = 0$ and $[M/H] = -2.5$ cases.

Moreover, the interactive web interface <http://stev.oapd.inaf.it/cmd> allows the present tables to be applied to the theoretical isochrones from Padova that are based on Girardi et al. (2000) tracks. Taking into account the Bertelli et al. (1994) isochrone models for initial masses higher than $7 M_\odot$, interpolated isochrones can be constructed for any age between 0 and 17 Gyr, and for any metallicity value between $Z = 0.0001$ and $Z = 0.03$ (see Girardi et al. 2002). A recent paper by Marigo et al. (2008) presents a substantial revision of these isochrones, with the replacement of the old TP-AGB tracks by much more detailed ones (cf. Marigo & Girardi 2007). The same isochrones will soon be extended to the planetary nebulae and white dwarfs stages.

A novelty in this latter web interface is the possibility of producing isochrones already incorporating the effect of interstellar extinction and reddening. This is done via interpolations in the tables of $BC_{S_\lambda}(A_V)$, using $\log T_{\text{eff}}$, $\log g$, and A_V as the independent parameters. The dependence in metallicity is ignored because it is really minimal (with $\lesssim 0.2\%$ differences in A_{S_λ}/A_V , going from the solar-metallicity to the very metal-poor

cases). Therefore, the $BC_{S_\lambda}(A_V)$ tables are prepared for $[M/H] = 0$ only, and the $BC_{S_\lambda}(A_V) - BC_{S_\lambda}(0)$ corrections derived from these tables are applied at all metallicities. This conceptually simple procedure is completely equivalent to alternative ones based on the interpolation of A_{S_λ}/A_V tables.

In this way, users of the web interface just need to specify the total interstellar extinction A_V , to apply the proper A_{S_λ} value at every point in the isochrone in each filter passband. We remind readers that the traditional approach is to apply a single value of extinction $A_{S_{\lambda_1}}$ and reddening $A_{S_{\lambda_1}} - A_{S_{\lambda_2}}$ to all points along an isochrone— $A_{S_{\lambda_1}}$ and $A_{S_{\lambda_2}}$ being related to the total A_V via constant multiplicative factors—then occasionally seeking for the A_V value that best fits a certain kind of observations. Instead, by applying the proper A_{S_λ} value to each isochrone point, we take into account its dependence on T_{eff} , on $\log g$, and on the total A_V itself. Our procedure will lead to both (1) the traditional overall translation of isochrones in the CMD, and (2) moderate changes in the isochrone shapes as interstellar extinction increases, especially for colors involving blue passbands and in the regime of high extinction. The changes in the isochrone shapes will not be noticeable for low extinctions (for $A_V \lesssim 2$ mag, in optical passbands) and/or for near-infrared passbands.

Moreover, the present tables of bolometric corrections, for the no-extinction case only, are being applied to Bertelli et al. (2008) isochrones, which cover a wide region of the helium-metal content (Y, Z) plane.

5. A FEW COMPARISONS WITH DATA

The present bolometric corrections and interstellar extinction coefficients for ACS and WFPC2 are part of a large database of transformations to many photometric systems (see Marigo et al. 2008 for the most updated version). Regarding the BC and color- T_{eff} relations for non-C stars, they have just slightly changed – a few hundredths of magnitude at most – with respect to those presented by Girardi et al. (2002). Therefore, the transformations presented in this paper are, in fact, already largely tested in the literature, in many papers that compared data and models seeking accuracies to within ~ 0.05 mag in colors. Most of these previous comparisons were, however, done in Johnson-Cousins, 2MASS, and SDSS systems.

In this paper we limit ourselves to a couple of simple comparisons between present models and observations in the ACS WFC F475W, F606W, and F814W passbands. These filters, together with F555W, are those in which most observations of resolved galaxies are available. These are also the filters used by the ANGST project.

The two panels of Figure 6 compare $Z = 0.0001$ and $Z = 0.0004$ isochrones superimposed on the CMD of the globular cluster M92 in three band combinations: $(V - I)$, $(F475W - F814W)$, and $(F606W - F814W)$. The VI data are from the Rosenberg et al. (2000) database, and has kindly been provided by A. Rosenberg. The F475W, F814W CMD has been obtained

using archival data from the ACS LCID project³ and the F606W, F814W CMD using ACS archival data of P.ID. 9453 (see Brown et al. 2005). The isochrones have $\log(\text{age/yr}) = 10.10$ and 10.15 , and are plotted from the lower main sequence up to the TRGB; additionally, we plot the zero-age horizontal branch from Girardi et al. (2000) tracks as derived by Palmieri et al. (2002). In all cases we have assumed a distance modulus $(m - M)_0 = 14.62$ (Del Principe et al. 2005), and interstellar extinction $A_V = 0.068$ following the Schlegel et al. (1998) maps. The cluster measured metallicity is $[\text{Fe}/\text{H}] = -2.24$ (Zinn & West 1984); or -2.16 (Carretta & Gratton 1997), which would correspond to $Z = 0.0001$ of Girardi et al. (2000) scaled-solar isochrones. This case is plotted in the upper panel. Note, however, that the $Z = 0.0004$ isochrones (*lower panel*) seem to fit the cluster sequences best, especially the RGB slope and the position of the main sequence turnoff. Note that the different mismatches that can be appreciated go in the same direction in the ground-based and the ACS fit. These mismatches are likely to be associated with the theoretical isochrones themselves, rather than to the color transformations.

This comparison with a nearby globular cluster illustrates the usefulness of the isochrones in the interpretation of star clusters in general, in which the photometry goes from the red giants down to the unevolved section of the main sequence. Most of the *HST* archival data refers instead to the composite stellar populations of nearby galaxies. Figure 7 shows the case of field WIDE1 located about $13'$ along the major axis of the spiral galaxy NGC 253, observed with ACS WFC for the ANGST project. Surprisingly, this region shows no sign of young stellar populations. A set of 10-Gyr-old isochrones overimposed on the data CMDs would indicate a range of $[M/H]$ going from about -2 to 0 . The observed CMD presents a striking decrease of stellar density above a certain level of the red giants, which obviously corresponds to the TRGB position. In Figure 7, the green line marks the TRGB for a series of 10-Gyr-old isochrones of varying $[\text{Fe}/\text{H}]$, displaced by a distance modulus of 27.6 mag. It can be noticed that this line describes remarkably well the shape of observed boundary between the stellar-rich RGB and the less populated AGB. Although the precise position of this TRGB line is determined also by the quality of the underlying stellar models, its reproduction over a wide range in color (covering more than 2 mag in F475W–F814W, and more than 1 mag in F606W–F814W) is clearly suggesting a correct behavior of the set of BCs that has been applied.

6. CONCLUDING REMARKS

In this paper we present tables of bolometric corrections and interstellar extinction coefficients, as a function of T_{eff} , $\log g$,

³The LCID (Local Cosmology from Isolated Dwarfs project, *HST* P.ID. 10505 and 10590) has obtained deep CMDs reaching the oldest main sequence turnoffs in five isolated Local Group galaxies and two clusters, M92 and NGC 1851.

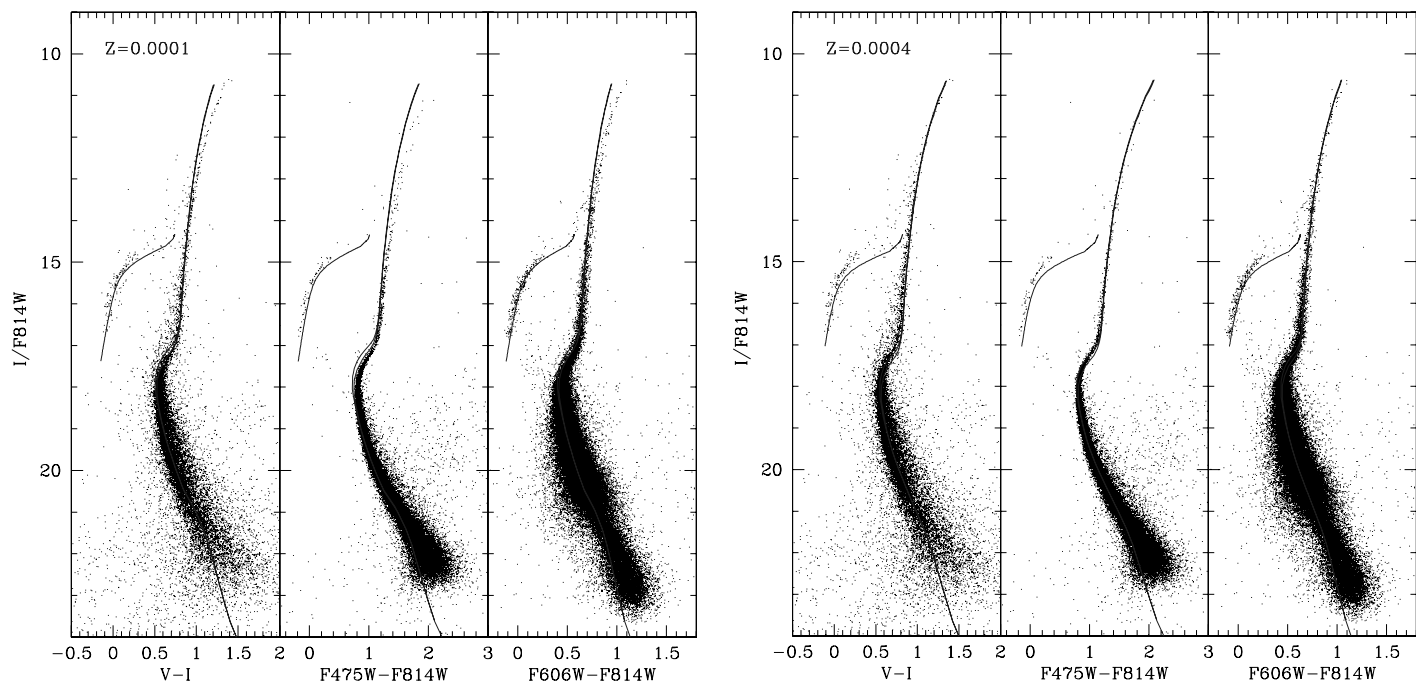


FIG. 6.—CMD of M92 in three different band combinations, compared with isochrones of metallicity $Z = 0.0001$ and 0.0004 as indicated in the labels. The isochrone ages are $\log(\text{age}/\text{yr}) = 10.10$ and 10.15 . See the text for details. See the electronic edition of the PASP for a color version of this figure.

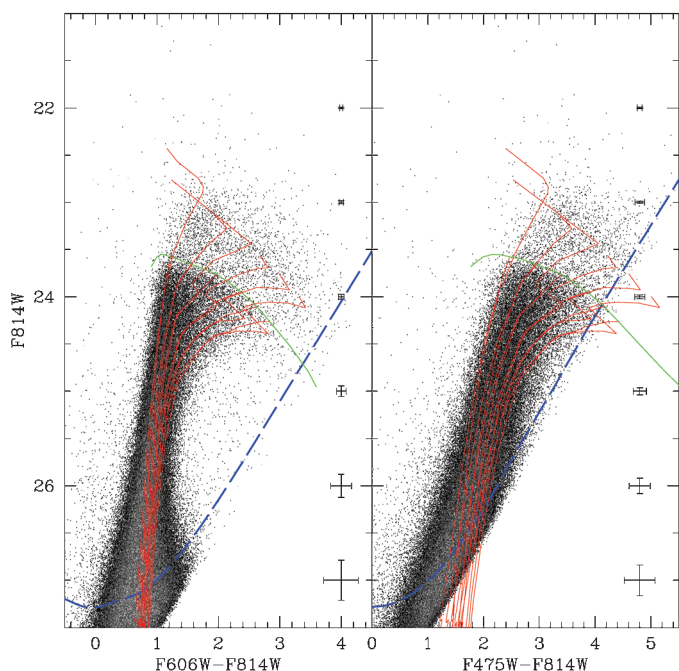


FIG. 7.—CMDs for the ACS Wide 1 pointing over NGC 253 (black and gray dots). The dashed line signals the 50% completeness limit of the ACS observations. The red lines are Marigo et al. (2008) 10-Gyr isochrones of metallicities $[M/H] = (-2.0, -1.5, -1.0, -0.5, 0.0)$, shifted by a distance modulus of 27.6. The green line marks the predicted locus of the TRGB for a more complete sequence of 10-Gyr isochrones with $[M/H]$ increasing from -2.3 to $+0.1$.

and $[M/H]$, for the WFPC2 and ACS systems. We keep the original definition of the Vega magnitudes by Holtzman et al. (1995) and Sirianni et al. (2005), while adopting the latest accurate measurements of the Vega fluxes by Bohlin (2007). The tables are applied to theoretical isochrones in <http://stev.oapd.inaf.it/cmd> and <http://stev.oapd.inaf.it/dustyAGB>, where similar data for many other photometric systems can be found. A major novelty of this work is that interstellar extinction coefficients are computed and applied star by star, so that isochrones present not only the overall shift in the CMDs as extinction increases, but also the subtle changes in their shape. We have made simple comparisons of these isochrones with real data, but a more careful testing of the transformations is left to future papers, and to the interested readers.

To the users of these *HST* isochrones, suffice it to recall a few general caveats, which apply to all photometric systems: First, our results are based on theoretical model atmospheres and spectral energy distributions, which are known not to be accurate at the hundredths-of-magnitude level, especially for the coolest stars. The same applies to Bohlin's (2007) Vega spectrum, which has uncertainties of 1%–2% in the optical. Moreover, the results for narrowband filters are particularly sensitive to errors in line opacity lists, and hence may be more uncertain. Overall, we expect the BCs for broadband filters to be accurate to within ~ 0.05 , with errors in colors being somewhat smaller. If we consider all the other uncertainties involved in the analysis of *HST* photometric data—including, for instance, the substantial uncertainties in the metallicity distribution and interstellar

extinction curves of nearby galaxies—errors of this magnitude may be considered as acceptable.

The data provided in this paper are not conceived as a static set of tables, but rather as a database that will be revised and/or extended as soon as there is a significant improvement on the input ingredients—comprising the spectral flux libraries, the reference Vega spectrum, filter transmission curves, the interstellar extinction curve, and the isochrone sets to which the transformations are applied. An important extension is the inclusion of tables for the WFC3 camera (Bond et al. 2006), which will soon replace WFPC2 after *HST* Servicing Mission 4. Preliminary tables for WFC3 are already provided in the present database because they may be useful in the preparation for *HST* Cycle 17. They were derived in the same way as WFPC2 and ACS tables, using filter throughputs from the STSDAS

package SYNPHOT v3.7 files available at 2008 January 15. The WFC3 data will be replaced as soon as revised throughputs are released by the WFC3 team.

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