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Journal:	<i>Monthly Notices of the Royal Astronomical Society</i>
Manuscript ID	Draft
Manuscript type:	Main Journal
Date Submitted by the Author:	n/a
Complete List of Authors:	Kasliwal, Vishal; University of Pennsylvania, Department of Physics & Astronomy; Princeton University, Department of Astrophysical Sciences Vogeley, Michael; Drexel University, Physics Richards, Gordon; Drexel University, Physics
Keywords:	galaxies: active < Galaxies, galaxies: Seyfert < Galaxies, (galaxies:) quasars: general < Galaxies, accretion, accretion discs < Physical Data and Processes

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# Extracting Information from AGN Variability

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Accepted XXX. Received YYY; in original form ZZZ

## ABSTRACT

AGN exhibit rapid, high amplitude stochastic flux variability across the entire electromagnetic spectrum on timescales ranging from hours to years. The cause of this variability is poorly understood. We present a new method for using variability to (1) measure the time-scales on which flux perturbations evolve and (2) characterize the driving flux perturbations. We model the observed light curve of an AGN as a linear differential equation driven by stochastic impulses. Physically, the impulses could be local ‘hot-spots’ in the accretion disk—the linear differential equation then governs how the hot spots evolve and dissipate. The impulse-response function of the accretion disk material is given by the Green’s function of the linear differential equation. The timescales on which the hot-spots radiate energy is characterized by the powerspectrum of the driving stochastic impulses. We analyze the light curve of the *Kepler* AGN Zw 229-15 and find that the observed variability behavior can be modeled as a damped harmonic oscillator perturbed by a colored noise process. The model powerspectrum turns over on time-scale 385 d. On shorter time-scales, the log-powerspectrum slope varies between 2 and 4, explaining the behavior noted by previous studies. We recover and identify both the 5.6 d and 67 d timescales reported by previous work. These timescales represent the time-scale on which flux perturbations grow, and the time-scale on which flux perturbations decay back to the steady-state flux level respectively. We make the software package used to study light curves using our method, KĀLĪ, available to the community.

**Key words:** galaxies: active – galaxies: Seyfert – quasars: general – accretion, accretion discs

## 1 INTRODUCTION

Active Galactic Nuclei (AGN) exhibit variability on timescales ranging from minutes and hours to years over the full electromagnetic spectrum. The underlying cause of the variability is not clear. Proposed models range from local viscosity fluctuations to oscillatory modes in the accretion disk (Ulrich et al. 1997). AGN are powered by the accretion of matter onto a central supermassive black-hole (Rees 1984). Matter infalling onto the supermassive black hole must lose angular momentum during inspiral. The loss of angular momentum results in the conversion of gravitational binding energy to the kinetic energy of the flow. A portion of this kinetic energy is radiated away resulting in the characteristic appearance of the AGN: high luminosity ( $\sim 10^{47}$  erg s<sup>-1</sup> for the brightest quasars) emanating from a small volume (radius  $r \sim 10^{14}$  cm) of space (Edelson et al. 1996).

It is generally accepted that the accretion inflow results in the formation of an accretion disk (Koratkar & Blaes 1999; Pringle 1981). At least three mechanisms may be responsible for the extraction of angular momentum from the accretion flow in a standard disk: (1) Magneto-rotational Instability (MRI) generated turbulence (Balbus & Hawley 1991, 1997); (2) Large scale magneto-hydrodynamic (MHD) outflow caused external stresses (Blandford & Payne 1982); and (3) Shocks produced by non-axis-symmetric waves (Fragile & Blaes 2008). MRI prescriptions for angular momentum transport are succinctly quantified by the alpha-prescription of Shakura & Sunyaev (1973)—see Balbus & Papaloizou (1999). Under the Shakura-Sunyaev  $\alpha$ -prescription, accretion flows are modelled as solutions of the vertically integrated hydrodynamic equations for the conservation of mass, radial and angular momentum, and energy. Stability criteria then dictate that solutions be of one of three types as characterized by location in surface density-mass accretion rate phase space (Blaes 2014). Low accretion rate coupled with

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5 low surface density results in the formation of advection-  
 6 dominated accretion flows (Narayan & Yi 1994; Chen et al.  
 7 1995) where the cooling time of the accretion flow is much  
 8 greater than the in-fall time. Such disks are optically thin  
 9 and may exist in low-luminosity AGN. Some quasars, a few  
 10 Seyfert 1s and the intermediate states of black-hole X-ray bi-  
 11 naries (BHXR) are thought to possess radiation pressure  
 12 dominated ‘slim’ disks— $H/r < 1$  where  $H$  is the local disk  
 13 height at radius  $r$ —with medium surface mass densities and  
 14 very high accretion rates (Abramowicz et al. 1988). Such  
 15 disks also have shorter inflow time as compared to the cool-  
 16 ing time-scale, making the flow advective. The third stable  
 17 solution occurs at high surface mass density but low accre-  
 18 tion rate and results in the formation of classic ‘thin’ ac-  
 19 cretion disks that are optically thick and geometrically thin  
 20 ( $H/r \ll 1$ ) (Shakura & Sunyaev 1973; Frank, King & Raine  
 21 2002). Most AGN and BHXR in the high and soft states  
 22 are thought to possess such disks (Blaes 2014).

23 Numerous mechanisms have been posited for producing  
 24 variability in accretion disks—see Done (2014), Mac-  
 25 carone (2014), and Uttley & Casella (2014) for comprehen-  
 26 sive overviews. While we are concerned primarily with  
 27 accretion onto supermassive black holes (SMBH) in AGN, a  
 28 fair fraction of the theory is developed with BHXR accre-  
 29 tion in mind. Although similarities exist between accretion  
 30 in BHXR and AGN, the difference in energy densities is  
 31 considerable, with AGNs having considerably higher lumi-  
 32 nosities and exhibiting variability on longer time-scales. For  
 33 instance, quasi-periodic oscillations (QPO) are routinely ob-  
 34 served in BHXR, whereas in AGN there are few reports  
 35 of QPOs (Gierliński et al. 2008; Andrae et al. 2013; Grind-  
 36 lay et al. 2014). While BHXR are commonly observed to  
 37 shift between various spectral states, such transitions have  
 38 not been observed in AGN to date (Kelly et al. 2011).

39 AGN optical variability may arise due to fluctuations  
 40 in the Shakura & Sunyaev (1973) viscosity parameter  $\alpha$ .  
 41 Lyubarskii (1997) showed that local stochastic fluctuations  
 42 in the accretion disk are capable of producing the power-  
 43 law power spectral densities (PSD) observed in AGN and  
 44 BHXR. Such fluctuations are thought to propagate in-  
 45 wards over time, suggesting that a powerful observational  
 46 test is the observation of blue-lags (Uttley & Casella 2014),  
 47 i.e., light curve variations at shorter wavelengths lagging  
 48 those at longer wavelengths. Evidence for the propagating  
 49 fluctuation model was provided by Miyamoto et al. (1988)  
 50 who observed that hard X-ray variations in Cyg X-1 lagged  
 51 behind soft X-ray variations with an inverse scaling between  
 52 the time lag and the variability time-scale. Starling et al.  
 53 (2004) used the viscosity fluctuation model to impose lower  
 54 limit of  $\alpha \sim 0.01$  on the disk viscosity. A later study by  
 55 King et al. (2007) using AGN variability triggered by vis-  
 56 cosity fluctuations suggests that  $\alpha \sim 0.1\text{--}0.4$ . Wood et al.  
 57 (2001) examined a variant of the Lyubarskii (1997) propa-  
 58 gating fluctuations model in which episodic mass deposition  
 59 at the outer edge of the accretion disk triggers viscosity fluc-  
 60 tuations in the accretion disk. This model was revisited by  
 Titarchuk et al. (2007) who generalized the treatment of  
 Wood et al. (2001) to include time-dependent fluctuation  
 sources located randomly on the disk face. Titarchuk et al.  
 (2007) follow the evolution of perturbations to the mass sur-  
 face density and luminosity arising from viscosity that is a  
 power-law function of disk radius. Subsequently, Kelly et al.

(2011) studied solutions to the surface mass density equa-  
 tion perturbed by stochastic fluctuations and showed that  
 the resulting light curve can be expressed as a linear combi-  
 nation of multiple Ornstein-Uhlenbeck processes (Gillespie  
 1996; Kelly et al. 2009).

Accretion disk variability may also be caused by the  
 existence of ‘hot-spots’ in the accretion disk (e.g., Maccarone  
 2014). Hot spots are thought to arise when accretion streams  
 impact the disk resulting in shock heating. Such hot spots  
 have already been detected during periods of quiescence in  
 BHXR (Froning et al. 2011; McClintock et al. 1995).

Thermal viscous instabilities or a variable mass transfer  
 rate may be responsible for accretion disk variability (La-  
 sota 2001; Coriat et al. 2012), though the amplitude of the  
 variability should be much smaller and the time-scale much  
 longer for AGN than for BHXR (Hameury et al. 2009).

Other sources of variability include magnetic flares in  
 the corona of the accretion disk (Poutanen & Fabian 1999)  
 or from the hot accretion flow itself (Veledina et al. 2013).  
 Alternatively, dynamo processes due to the magnetic fields  
 threading the disk may be responsible for producing vari-  
 ability (Livio et al. 2003; King et al. 2004; Mayer & Pringle  
 2006). Janiuk & Czerny (2007) have suggested that the vari-  
 ability may be produced in the hot X-ray corona surrounding  
 the accretion disk from the evaporation of mass perturba-  
 tions in the disk. Misra & Zdziarski (2008) suggested that  
 the X-ray variability of Cyg X-1 can be modeled as a damped  
 harmonic oscillator perturbed by noise. Quasi-periodic sig-  
 nals in the light curve may be produced by weak shocks in  
 the innermost region of the accretion caused by the presence  
 of a tilt in the accretion disk of the black hole (Fragile et al.  
 2007; Fragile & Blaes 2008).

Quasi-periodic signals may be caused by global oscil-  
 lation modes or parametric resonance instability (Reynolds  
 & Miller 2009; O’Neill et al. 2009, 2011). Variability may  
 also be produced by standing shocks interacting with fluc-  
 tuations in tilted accretion disks (Henisey et al. 2012).

Numerical simulations of light curves from accre-  
 tion disks began appearing in the literature starting with  
 Schnittman et al. (2006). A ray-tracing and radiative trans-  
 fer code was used to create images of a simulated accretion  
 disk. Using images generated during multiple time dumps  
 in the simulation, Schnittman et al. (2006) were able to cre-  
 ate mock light curves resulting from accretion flows. Noble & Krolik (2009) performed fully General-Relativistic  
 (GR) magneto-hydrodynamic (MHD) simulations of accre-  
 tion flows around black holes and showed that while the  
 variability was generated by accretion rate variances caused  
 by MHD turbulence, time delay effects steepened the PSD of  
 the variations. Other GRMHD simulations have also shown  
 that MHD turbulence results in accretion rate variations  
 that can produce variability (Mościbrodzka et al. 2009; Dex-  
 ter et al. 2009, 2010). Recently, Schnittman & Krolik (2013)  
 and Schnittman et al. (2013) have even managed to produce  
 full (X-ray) spectra demonstrating the effect of variability  
 on the spectral slope.

Several mechanisms may be responsible for the variabil-  
 ity observed in AGN light curves. Light curves in the opti-  
 cal are thought to be driven by X-ray variability on shorter  
 time-scales but exhibit a local source of variability on longer  
 time-scales (Uttley & Casella 2014). This implies that it is  
 unlikely that a single mechanism drives variability across

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the entire electromagnetic spectrum. Rather, the nature of the underlying mechanism responsible for variability may be different at different wavelengths.

Studies of AGN variability model the observed light curve as stochastic processes i.e. mathematically the light curve as a function of time is a random variable that correlates with previous values of the same light curve. Prescriptions for stochastic models require two properties to be specified: (1) the distribution from which to draw the random values of the light curve; and (2) the strength of the correlation of a given value of the light curve with previous values of the same light curve. Most studies of AGN variability (implicitly) invoke some form of the central limit theorem to argue that property 1 should be a Gaussian distribution which then lets them specify property 2 using a second-order statistic such as the power spectral density (PSD), or the auto-covariance function (ACVF), or the structure function (SF). The most straightforward method of fitting such a stochastic model to a light curve is by direct estimation of the PSD of the observed light curve. Using periodogram estimators of the PSD, the observed PSD can be compared to the predicted theoretical PSD as advocated in Uttley, McHardy & Papadakis (2002). Alternatively, it is also possible to estimate the ACVF of the observed light curve and fit it to the theoretical ACVF of a given stochastic process. Equivalently, following Kasliwal, Vogeley & Richards (2015a), estimates of the structure function (SF) of the observed light curve can be fit to the theoretical SF computed either directly from the theoretical ACVF, or from Monte-Carlo simulations of light curves generated using the prescribed model PSD.

The principle drawback with the methods above is that while stochastic models provide expressions for the analytic forms of the PSD, ACF, and SF, typically no analytic results exist about the *distributions* of the estimates of these statistics about the theoretical curve (Brockwell & Davis 2010). Even simple assumptions about the Gaussianity of the distribution of the estimates may be incorrect. For example, as shown in Emmanoulopoulos, McHardy & Uttley (2010) and in Kasliwal et al. (2015a), it is the estimates of the logarithm of the SF rather than the SF that are Gaussian distributed. This means that the application of the methods above rely on Monte-Carlo techniques to estimate the distribution of the property, be it the PSD, ACVF or SF that is being fit. Usually, some assumption of normality is made and typically the estimates are treated as independent purely for computational practicality.

Recently, Kelly et al. (2014) suggested that stochastic variability in a large class of astrophysical sources may be modelled using Continuous-time AutoRegressive Moving Average (C-ARMA) processes. The attractiveness of using C-ARMA processes to model stochastic phenomenon stems from two factors: (1) C-ARMA processes directly specify property 2, i.e., the correlation structure of the light curve by providing a differential equation that governs how the light curve evolves over time. This alternative to providing a second-order statistic is powerful because it makes the light curve evolution more amenable to physical interpretation as is shown in this work. It also provides some idea of how to replace Gaussian distributions with more generalized  $\alpha$ -stable distributions that may be truer to reality (Brockwell & Marquardt 2005; Brockwell & Lindner 2009); and

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(2) flexibility of the shape of their PSD. Since the PSD of a C-ARMA process is a rational function, it can be made arbitrarily complex to model a very wide range of phenomena. This flexibility is achieved by allowing C-ARMA processes to have multiple Autoregressive (quantified by  $p$ ) and Moving Average (quantified by  $q$ ) terms under the constraint  $q < p$ . If  $\chi$  is a small flux perturbation from the mean flux level, then a C-ARMA( $p,q$ ) process that models the time evolution of  $\chi$  has the general form

$$\begin{aligned} d^p \chi + \alpha_1 d^{p-1} \chi + \dots + \alpha_{p-1} d \chi + \alpha_p \chi = \\ \beta_0 dW + \beta_1 d(dW) + \dots + \beta_{p-1} d^{p-1}(dW) \end{aligned} \quad (1)$$

where  $W$  is a *Wiener* process and is responsible for introducing stochasticity into the behavior of  $\chi$ . The  $\alpha$  and  $\beta$  are constants that govern the behavior of the resulting process and have dimensions of various powers of time. In this paper, we demonstrate how C-ARMA processes can arise from the physical processes in accretion disks and present a Green's function-based method for understanding what a C-ARMA process tells us about accretion physics.

## 2 OUTLINE

We propose a Green's function-based method for analysing observed light curves, particularly in the UV through IR where the fluctuations are smaller than they tend to be in the X-ray. Such a method facilitates comparison with analytic and simulation work on variability. Our method builds on the C-ARMA modelling technique of Kelly et al. (2014) by proposing that the C-ARMA process is a direct result of the physics driving the variability.

We discuss our proposed interpretation of AGN variability as modelled by C-ARMA processes in section 3 and present the underlying mathematical details in appendices A, B, and C. A simple example of a C-ARMA process—the C-ARMA(2,1) process—is presented in section 4. We have written a fast but easy-to-use software package, KALI, for inferring C-ARMA model parameters from observed light curves. We present an overview of the operation of KALI in section 5 and defer mathematical details to appendix D.

Appendices D1 through D5 present details of inferring the properties of the C-ARMA process from noisy observations using the Kalman filter. Alternatives exist to using the Kalman filter for inferring the parameters of a C-ARMA process. We discuss why the Kalman filter is the optimal choice of fitting techniques in appendix D1 and suggest an alternative to the traditional *controllable canonical form* state-space representation used by Jones & Ackerson (1990) and Brockwell (2014). Our choice of state-space representations is (1) better suited for extension to more complex models in the future and (2) computationally cheaper for regularly sampled data. The Kalman filter is presented in appendix D4. C-ARMA models can have an arbitrary number of autoregressive ( $p$ ) and moving average ( $q$ ) terms. In appendix D5, we discuss information criteria that can be used to select the model order.

We present an example of the application of C-ARMA modelling to the *Kepler* light curve of the Seyfert 1 galaxy Zw 229-15 in section 6. In this section, we demonstrate how the C-ARMA model can be used to explain the multiple

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slopes and time-scales found in the light curve of this AGN by previous work. We present our conclusions in section 7.

### 3 PHYSICS FROM THE C-ARMA PROCESS EQUATION

Motivated by approaches to systems modeling in the engineering literature (Pandit & Wu 2001), we propose that equation (1), the C-ARMA process equation, be interpreted literally; we view the Left-Hand Side (LHS) as characterizing the dissipative forces at work in the accretion disk while the Right-Hand Side (RHS) specifies the distribution of flux perturbations. For example, if the dissipation equation is found to be of 1<sup>st</sup>-order, then small perturbations of the flux appear as soon as they are triggered and then exponentially decay back to the mean level (assuming that the steady-state corresponds to constant flux). On the other hand, if the dissipation equation is of 2<sup>nd</sup>-order, then flux perturbations will rapidly grow to a peak amplitude after they are triggered before decaying back to the mean level. In this case, the decay will not be a simple exponential decay and can occur sinusoidally depending on the exact solution.

In appendix A, we show how the LHS of eq (1) can arise from the non-linear equations that govern how flux perturbations evolve. The LHS of eq (1) is given by

$$d^p \chi + \alpha_1 d^{p-1} \chi + \dots + \alpha_{p-1} d \chi + \alpha_p \chi = u, \quad (2)$$

where the  $u$  are the input stochastic impulses. We can understand how accretion physics works by studying the Green's function of equation (2) (Pandit & Wu 2001). The Green's function quantifies how a unit flux perturbation evolves after it is triggered. Conceptually, the Green's function of a system characterizes the impulse response of the system. Particular solutions are found by linear superposition (integration) of the Green's function over all the impulses that drive the system. By comparing the Green's function to the expected behaviour of perturbations, we can gain insight into accretion physics. The form of the Green's function for equation (2) is given by

$$G(t) = \sum_{k=1}^p c_k e^{\rho_k t}, \quad (3)$$

where the  $\rho_k$  are the roots of the autoregressive polynomial of equation (A16) and the  $c_k$  are constants that can be determined as stated in appendix B. Using Euler's formula, it is clear that the real part of each  $\rho_k$  corresponds to an exponential decay timescale  $\tau_k = 1/\Re(\rho_k)$  while the complex part of each  $\rho_k$  produces modulating oscillations with period  $T_k = 2\pi/\Im(\rho_k)$ .

The cause of perturbations, i.e., variability, in the flux emanating from an AGN is unknown, but we may construct a very general mathematical model for the form of the flux perturbations based on physical and statistical arguments. The region of space surrounding the central engine is vast. The inner edge of the accretion disk is thought to begin a few  $r_G \sim 20$  AU ( $M_{BH} = 10^9 M_\odot$ ) away from the event horizon, while the outer edge is thought to lie at a distance of a few parsec from the event horizon. Contributions to the deviations in the flux most likely simultaneously arise from

many individual locations in this system. Under the assumption that the distribution of energy released by each individual event is drawn from a distribution with finite second-moment, the central-limit-theorem assures us that the total energy introduced into the system at any instant will be drawn from a normal distribution, i.e., some form of noise. Due to the non-linearity of the physics involved, it is quite likely that correlations exist between the energy released into the disk at different moments in time. Therefore, we shall model the correlations by treating the driving process  $u$  in equation (2) as a correlated Gaussian noise process. In appendix C, we argue why the natural mathematical form of the RHS should be given by

$$u = \beta_0(dW) + \beta_1 d(dW) + \dots + \beta_{p-2} d^{p-2}(dW) + \beta_{p-1} d^{p-1}(dW). \quad (4)$$

This form for the input stochastic impulses can be characterized by the power spectral density (PSD) of the impulses given by

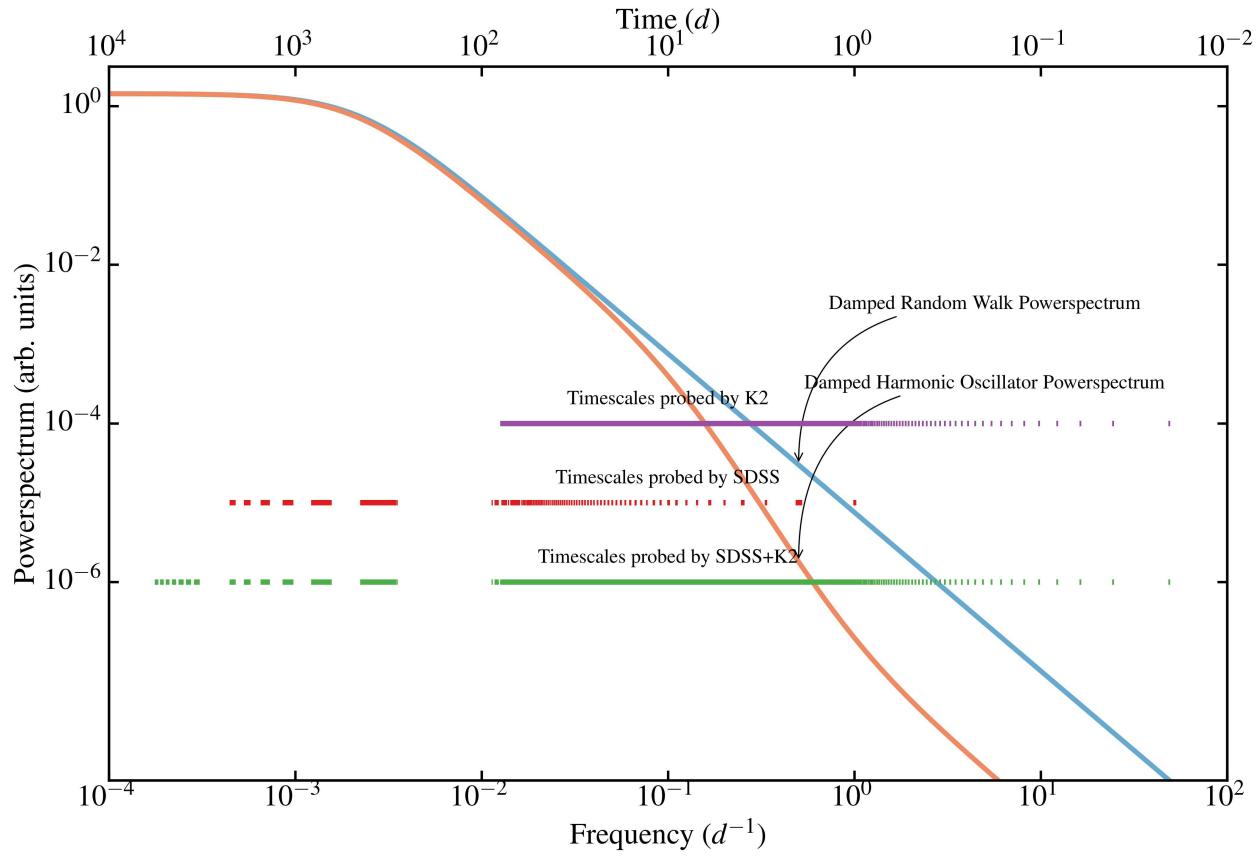
$$S_{uu}(\nu) = \frac{1}{2\pi} |\beta(2\pi\nu)|^2, \quad (5)$$

where  $\beta(z)$  is the moving average polynomial of equation (C4). The PSD of the stochastic impulses tells about the timescales on which the stochastic impulses occur. Combining equation (2) with equation (4) yields the original C-ARMA process equation (1). A C-ARMA process is a linear differential equation perturbed/driven by a random process. This driving process may be viewed as disturbances generated by MRI instabilities in the accretion disk. *Thus we see that we can describe AGN luminosity variations in terms of a linear system driven by colored noise.*

Fitting the observed light curve of an AGN to the CARMA process of equation (1) tells us how flux perturbations evolve via the Green's function equation (3). Insight into the flux perturbation producing mechanism may be gained from the impulse-PSD equation (5) by comparing it to predictions from simulations of turbulence (Miesch et al. 2015) or from thermodynamic considerations of the accretion disk material (Mellen 1952). Alternatively, the values of characteristic time-scales implied by the forms of the Green's function equation (3) and impulse-PSD equation (5) may be matched to theoretical time-scales pertinent to the generation and evolution of accretion disk instabilities by the MRI process. In the next section we discuss a simple C-ARMA process—the C-ARMA(2,1) process—in detail and derive the characteristic time-scales present in this process as an example.

### 4 AN EXAMPLE OF A C-ARMA PROCESS: THE C-ARMA(2,1) PROCESS

We shall discuss the C-ARMA(2,1) process in detail because it has strong physical motivations and is relatively simple to interpret. Notice that in the process of linearising the original non-linear differential equation of equation (A1), we have not changed the order of the highest derivative of the flux that appears in that equation; if the original equation governing the flux emitted by the accretion disk is of  $p^{\text{th}}$ -order, the corresponding linearised C-ARMA process is also of  $p^{\text{th}}$ -order. Most physical phenomena are governed by 1<sup>st</sup>-



**Figure 1.** Why do AGN poorly-sampled AGN light curves look like a damped random walk (DRW)? We plot two PSD models - the DRW of Kelly et al. (2009) in cyan and a damped harmonic oscillator (DHO) in orange. The DRW is a C-ARMA(1,0) model i.e. the simplest possible C-ARMA model. The DHO is C-ARMA(2, $q$ ) model, i.e., it is a 2<sup>nd</sup>-order C-ARMA model. The vertical dashes are located at the time-intervals probed by NASA's *K2* mission (purple), the SDSS survey (red), and an idealized combination of the two (green). On timescales longer than  $\sim 10$  d, both PSDs are identical, explaining why the simpler DRW model is popular for light curves that do not probe short timescales. The availability of better data in the future will necessitate more sophisticated C-ARMA models.

and 2<sup>nd</sup>-order equations. Hence, we may expect to find that in the linear regime, AGN variability is well described a C-ARMA(1,0) process or a C-ARMA(2, $q$ ) process with  $q < 2$ .

The general form of the C-ARMA(2,1) process is

$$d^2\chi + \alpha_1 d\chi + \alpha_2 \chi = \beta_0(dW) + \beta_1 d(dW), \quad (6)$$

where  $[\alpha_1] = T^{-1}$ ,  $[\alpha_2] = T^{-2}$ ,  $[\beta_0] = [\chi] T^{-3/2}$ , and  $[\beta_1] = [\chi] T^{-1/2}$ . The LHS determines how individual perturbations evolve over time. The RHS determines the shape of the driving process PSD. If  $\beta_1 = 0$ , the driving noise process PSD is flat over all frequencies, i.e., it is a white noise process. If  $\beta \neq 0$ , the disturbance process will have different amounts of power at different frequencies, i.e., it is a ‘colored’ noise process. Equation (6) is of the same mathematical form as a *Damped Harmonic Oscillator* (DHO) that is driven/perturbed by a colored noise process. Light curves that obey the C-ARMA(2,1) process equation have perturbations from the mean flux level that individually behave like damped harmonic oscillators. Given the identification of the C-ARMA(2,1) process with a DHO,  $\alpha_2 = \omega^2$  gives the angular frequency of the HO, and  $\alpha_1 = 2\omega\zeta_{\text{DHO}}$  gives the damping ratio  $\zeta_{\text{DHO}}$  (Pandit & Wu 2001, see chapter 7

for notation). The time-period of the harmonic oscillator is  $T_{\text{HO}} = 2\pi/\omega$ . The roots of the LHS of equation (6) are given by

$$\rho_1, \rho_2 = -\omega\zeta_{\text{DHO}} \pm \omega\sqrt{\zeta_{\text{DHO}}^2 - 1} = -\frac{\alpha_1}{2} \pm \sqrt{\frac{\alpha_1^2}{4} - \alpha_2}. \quad (7)$$

The behavior of the flux perturbations is non-physical in this context if  $\zeta_{\text{DHO}} = 0$ , i.e., the flux perturbations are undamped and never die out. For  $0 < \zeta_{\text{DHO}} < 1$ , the flux perturbations are under-damped and will oscillate in value around the steady-state flux value with angular frequency  $\omega' = \omega\sqrt{1 - \zeta_{\text{DHO}}^2}$  but gradually decrease in amplitude. If  $\zeta_{\text{DHO}} = 1$ , the perturbations are critically damped and will rapidly die down to the steady-state flux. If  $\zeta_{\text{DHO}} > 1$ , flux perturbations are over-damped and exponentially decay to the steady-state flux. The rate of decay decreases with  $\zeta_{\text{DHO}}$ ; over-damped flux perturbations with larger  $\zeta$  take longer to decay, increasing the correlation timescale of the light curve. From equation (C11), the PSD of the C-ARMA(2,1) process

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6 is given by  
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$$S_{\chi\chi}(\nu) = \frac{1}{2\pi} \frac{\beta_0^2 + 4\pi^2\beta_1^2\nu^2}{16\pi^4\nu^4 + 4\pi^2(a_1^2 - 2\alpha_2)\nu^2 + \alpha_2^2}, \quad (8)$$

8 which has frequency-dependent behavior.  
9

10 Figure 1 shows the dependence of a DHO PSD on  
11 frequency. At very high- and low-frequencies, the DHO PSD  
12 looks like the damped random walk PSD of Kelly et al.  
13 (2009) with  $S_{\chi\chi} \sim 1/\nu^2$ . There is a mid-range of frequencies  
14 where the  $\nu^4$  term in the denominator dominates. At the  
15 lowest frequencies, the PSD is dominated by the frequency  
16 independent terms and asymptotically approaches  $\beta_2^2/2\pi\alpha_2^2$ .  
17 This frequency-dependent PSD behavior may explain why  
18 ground-based observations of AGN light curves suggest that  
19 AGN variability is well-modelled by a DRW whereas Kepler  
20 observations of AGN suggest that the PSD is steeper than  
21 a DRW: the ground based observations are not sampling at  
22 high enough frequency to be able to see the  $\nu^4$  dominated  
23 regime.

24 Individual accretion disk perturbations are smeared by  
25 processes such as turbulent dissipation, rotation and thermal  
26 dissipation. The nature of these processes can be inferred via  
27 the Green's function approach as discussed in section 3. The  
28 Green's function of the DHO is given by  
29

$$G(t) = \frac{e^{\rho_1 t} - e^{\rho_2 t}}{\rho_1 - \rho_2}, \quad (9)$$

30 where  $\rho_1$  and  $\rho_2$  are the roots of the autoregressive  
31 characteristic polynomial of equation (7). If  $\zeta_{\text{dHO}} > 1$ , i.e., flux  
32 perturbations are constrained to be damped out by the  
33 dissipative processes in the accretion disk, then the roots are  
34 purely real-valued. A perturbation to the accretion disk flux  
35 in the form of an impulse at  $t = t_0$  will result in the flux  
increasing from  $t_0$  to  $t_{\max}$  given by

$$t_{\max} = \frac{1}{\rho_2 - \rho_1} \ln \left| \frac{\rho_1}{\rho_2} \right|. \quad (10)$$

36 After  $t_{\max}$ , the effect of the original perturbation decays to  
37 the steady-state flux level. Although the decay is not a pure  
38 exponential, it is useful to characterize the decay rate by an  
39 e-folding time-scale  $t_{\text{e-fold}}$ . It may be computed by solving  
40

$$G(t_{\text{e-fold}}) = \frac{G(t_{\max})}{e}, \quad (11)$$

41 numerically for  $t_{\text{e-fold}}$ .

42 From the PSD of the C-ARMA(2,1) process in equation  
43 (8), we see that on very long time-scales, both the nu-  
44 merator and the denominator should approach a constant  
45 value. The PSD begins to flatten on time-scale  $t_{\text{flat}}$  where  
46

$$t_{\text{flat}} \sim \max \left( \frac{2\pi\beta_1}{\beta_0}, \frac{2\pi}{\sqrt{\alpha_2}}, \frac{2\pi\sqrt{\alpha_1^2 - 2\alpha_2}}{\alpha_2} \right). \quad (12)$$

47 On the other hand, on very short time-scales, the  $\nu^4$  term in  
48 the denominator dominates. The PSD  $\sim \nu^{-4}$  on time-scale  
49  $t_{\text{steep}}$  given by  
50

$$t_{\text{steep}} \sim \min \left( \frac{2\pi}{\sqrt{\alpha_1^2 - 2\alpha_2}}, \frac{2\pi}{\sqrt{\alpha_2}} \right). \quad (13)$$

51 We see that the C-ARMA(2,1) model contains sev-  
52 eral time-scales of interest: the peak of the Green's  
53

54 function( $t_{\max}$ ), the e-folding time-scale of the Green's function-  
55 tion ( $t_{\text{e-fold}}$ ), the time-scale on which the PSD of the driving  
56 flux disturbances turns over ( $t_{\text{turn}}$ ), the time-scale on which  
57 the PSD begins to flatten ( $t_{\text{flat}}$ ), and finally the time-scale on  
58 which the PSD  $\sim \nu^4$  ( $t_{\text{steep}}$ ). All of these time-scales should  
59 be detectable in the PSD of the light curve in the form of  
60 various bends and turnovers. Similar time-scales can eas-  
61 ily be derived for higher-order C-ARMA models with the  
62 number of significant time-scales increasing with the model  
63 order, i.e., more complex C-ARMA models will have more  
64 'features' in the PSD indicative of still more complex pro-  
65 cesses at work in the accretion disk.

66 In the case of the C-ARMA(2,1) process, the driving  
67 noise is best interpreted as the 'acceleration' of flux per-  
68 turbations away from the steady-state flux level. The PSD of  
69 the driving noise process is given by  
70

$$S_{uu}(\nu) = \frac{1}{2\pi} (\beta_0^2 + 4\pi^2\beta_1^2\nu^2), \quad (14)$$

71 and can be described (at higher frequencies) as a violet- or  
72 purple-noise spectrum. At low frequencies, the  $\beta_0^2$  term in  
73 the RHS dominates: the log-PSD slope is 0 and the PSD is  
74 similar to that of white-noise. The transition from the  $\beta_0^2$ -  
75 dominated to the  $4\pi^2\beta_1^2\nu^2$ -dominated behaviour occurs at  
76

$$t_{\text{turn}} = \frac{2\pi\beta_1}{\beta_0}. \quad (15)$$

77 Thermal motion within a fluid results in the formation  
78 of sound-waves in the fluid with characteristic PSD  $\sim \nu^2$   
79 (Mellen 1952) suggesting that the driving noise PSD of  
80 CARMA(2,1) models can be ascribed to the thermal noise of  
81 the accretion disk material. Alternatively, it may be caused  
82 by the presence of eddies in the turbulent flow of the accre-  
83 tion disk (Miesch et al. 2015).

84 In the next section, we discuss how to infer the values of  
85 the coefficients in the C-ARMA process describing accretion  
86 disk fluctuations in an AGN from observations of the light  
87 curve of the object.

## 5 KĀLĪ: SOFTWARE FOR THE C-ARMA ANALYSIS OF A LIGHT CURVE

88 We have implemented a fully parallelized and vectorized  
89 software package, KĀLĪ (named jointly after the Hindu god-  
90 dess of time, change, and power and also as an acronym  
91 for KArma LIBrary), to analyze stochastic light curves us-  
92 ing the C-ARMA process presented in equation (1). KĀLĪ  
93 is implemented in the C++ programming language with  
94 PYTHON language bindings for ease of use. KĀLĪ may be ob-  
95 tained from <https://github.com/AstroVPK/libcarma> and  
96 is installable on Linux variants and Mac OSX using the  
97 install scripts provided. Appendix D describes the mathe-  
98 matics implemented in KĀLĪ that are used to perform the  
99 Bayesian Markov Chain Monte-Carlo (MCMC) inferencing  
of a stochastic light curve.

100 Before performing a C-ARMA analysis, it should first  
101 be determined if a C-ARMA model is appropriate. A C-  
102 ARMA process may be unsuitable if the light curve ex-  
103 hibits non-stationary behaviour, i.e., (1) the appearance of  
104 the light curve changes dramatically at some point in the

light curve or (2) the light curve has a marked trend. A C-ARMA analysis may also be unsuitable if the light curve has flares (as in the case of the blazar of [Edelson et al. 2013](#)). C-ARMA models are also unsuitable if the variations in the flux are very large compared to the flux. A C-ARMA model is suitable if the sample autocorrelation and partial autocorrelation functions of the light curve exponentially decay to below the significance level for a pure white noise process ([Brockwell & Davis 2010](#)).

Once it is decided that a C-ARMA process may be an appropriate choice of model for the light curve, we must estimate the model-order of the C-ARMA process. Some insight may be available from the autocorrelation and partial autocorrelation functions (see [Brockwell & Davis 2010](#) for details). We automate the process by sampling parameters from a range of models with increasing numbers of C-ARMA parameters and using the *Deviance Information Criteria* (DIC) of equation (D42) to select the best model. For a given model order  $p$  and  $q$ , we use MCMC to sample the space of model parameters. Before beginning the Kalman recursions of appendix D4, we construct a mask matrix to keep track of missing observations and subtract the mean of the light curve from every observation to make the light curve a zero mean process. For a given set of model parameters ( $\alpha_k$  and  $\beta_k$ ), we check for the validity of the parameter set, i.e., are real parts of the roots of the autoregressive and moving average polynomials less than zero? If they are not, we reject the model outright by setting the prior likelihood to 0. If the model parameters are permissible, we set the prior likelihood to 1 and compute the transition matrix  $\mathbf{F}$  using equation (D9) and the disturbance variance-covariance matrix  $\mathbf{Q}$  using equation (D10). We then use equations (D32) and (D33) to initialize the state  $\hat{\mathbf{x}}_0^+$  and determine the initial state uncertainty  $\mathbf{P}_0^+$ . To compute the likelihood of the light curve, we iterate through the observations using the Kalman filtering equations (D34) through (D40) to compute the innovation and the uncertainty of the innovation for every observation. The computed innovations and uncertainties are used to calculate the likelihood using equation (D30) which is then equal to the model likelihood of equation (D29) given our choice of the prior. The MCMC algorithm can then sample the parameter space and obtain draws of the C-ARMA model parameters from the posterior probability distribution of the parameters. Using the DIC with the full set of draws (Bayesian approach), or the *corrected Akaike Information Criteria* (AICc) with the parameter set that has the highest likelihood (frequentist approach), we can pick the best fitting values of  $p$  and  $q$  simultaneously with the parameter estimation process. Once we have selected the model order, we can use the inferred parameter values to compute estimates of (1) the Green's function; (2) the PSD of the driving impulses; and (3) time-scales such as  $t_{\max}$ .

In the next section, we present a study of the accretion physics of the Seyfert 1 AGN Zw 229-15 using data from the *Kepler* mission.

## Extracting Information from AGN Variability 7

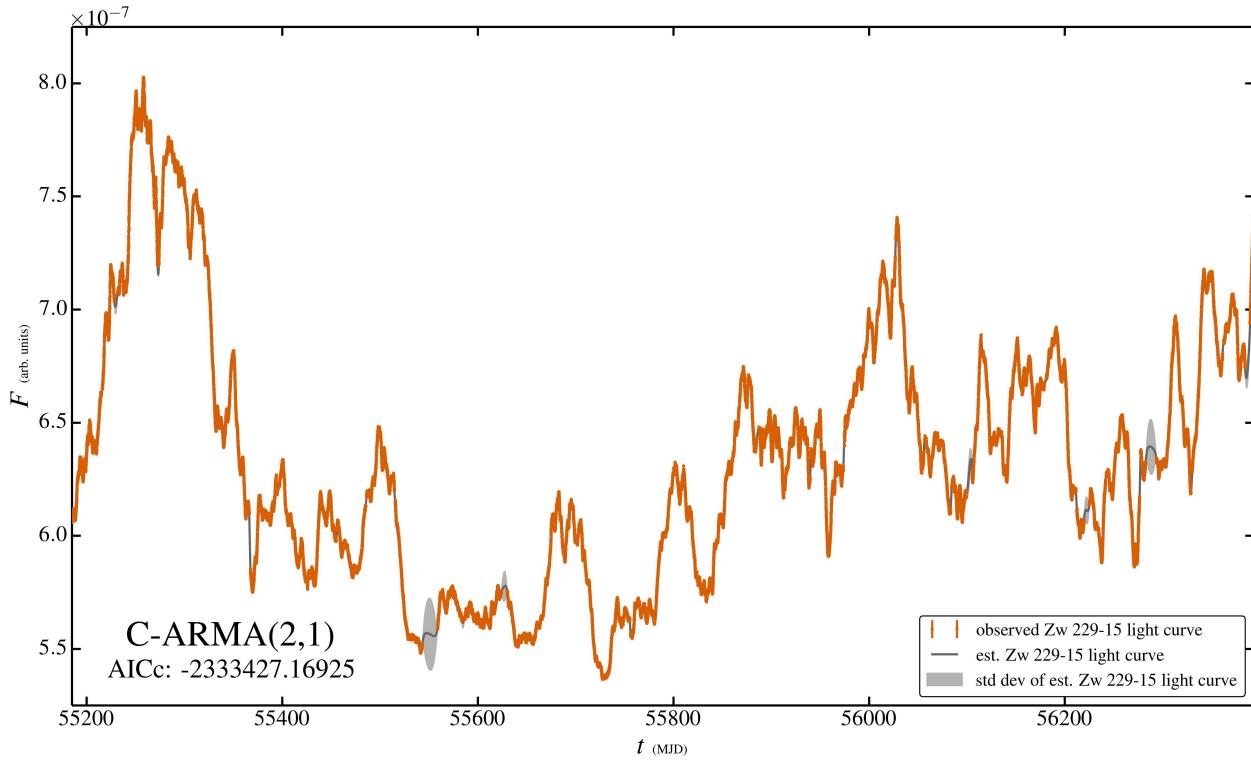
### 6 PROBING ACCRETION PHYSICS WITH THE *Kepler* SEYFERT 1 ZW 229-15: C-ARMA ANALYSIS OF THE LIGHT CURVE

We apply the C-ARMA inferencing techniques described in sections D1 through D5 to the re-processed *Kepler* light curve of the Seyfert 1 AGN Zw 229-15 (*Kepler* catalog name KIC 006932990) located at  $z = 0.028$ . Details of the re-processing applied to the light curve can be found in ([Williams & Carini 2015](#)) and in ([Kasliwal et al. \(2015b\)](#)).

#### 6.1 Determination of a C-ARMA Model for the Light Curve of Zw 229-15

We use our package, KALI to fit the observed light curve of Zw 229-15 to C-ARMA( $p,q$ ) processes with  $1 \leq p \leq 10$  and  $0 \leq q < p$ . In addition to using our package KALI, we performed the same analysis using the C++ and PYTHON library CARMA\_PACK of [Kelly et al. \(2014\)](#). While both codes produced similar numerical results, KALI runs about 50× faster due to our extensive usage of the high performance Intel Math Kernel Library (MKL). Prior to the analysis, the light curve of Zw 229-15 was corrected to the rest-frame of the galaxy at  $z = 0.0275$  by scaling  $\delta t$  by  $1/(1+z)$ . No binning of the light curve was performed. Instead of adopting the *Kepler* error estimates, CARMA\_PACK assumes that the variance of the observation noise may be inaccurate and sets  $\sigma_{N,\text{true}}^2 = M_{\text{Err}}\sigma_{N,\text{stated}}^2$  for each observation. The multiplicative scaling factor  $M_{\text{Err}}$  is treated as a fit parameter, i.e., all purely white noise in the data is assumed to originate in process of observation rather than be intrinsic to the signal. We find that the median value of the multiplicative scaling factor is  $M_{\text{Err}} = 1.30$ . The observation noise levels inferred from the pixel noise statistics of the *Kepler* CCDs appear to underestimate the true observation noise level in the signal by ∼14 percent at the 95-percent confidence level.

Although we searched an extensive space of C-ARMA models ( $1 \leq p \leq 10$  and  $0 \leq q < p$ ), the relatively simple C-ARMA(2,1) model achieved the lowest AICc value, i.e., the best fit. The AICc value of  $-2333427.16925$  for the C-ARMA(2,1) model v/s the value  $-2333328.70173$  for the C-ARMA(3,2) model (the next best model) suggests that the C-ARMA(3,2) is very unlikely (∼ $10^{-22}$  times as likely) to explain the observed data ([Burnham & Anderson 2003](#)). Figure 2 shows the full light curve of Zw 229-15 (orange) along with the RTS smoothed realization (grey) of section D6. In addition to the expectation value of the smoothed realization, we show (light grey) the  $1\sigma$  limits on the smoothed light curve. Notice how this  $1\sigma$  limit is usually too small to be visible but increases dramatically when the observed light curve has missing values. Figure 3 shows the median estimated PSD of the light curve of Zw 229-15 (grey) along with the 95<sup>th</sup>-percentile confidence region (light grey). From the reported *Kepler* noise estimates and the estimates of  $M_{\text{Err}}$ , we compute the noise level in the PSD using  $S_{NN} = 2\delta t M_{\text{Err}} \sigma_N^2$  where we use median values for  $M_{\text{Err}}$  and  $\sigma_N^2$ . We estimate the frequency  $\nu_{\text{noise}} \sim 2.9 \text{ d}^{-1}$  at which the PSD of the light curve drops below the noise level by linearly interpolating the PSD between points. The corresponding time-scale is  $t_{\text{noise}} \sim 8.3 \text{ h}$ . Since the PSD of the underlying flux variations drops below the PSD of the *Kepler* noise level at  $t_{\text{noise}}$ , we have no information about the



**Figure 2.** Light curve of the Seyfert 1 AGN Zw 229-15. This light curve is best fit by the C-ARMA(2,1) process discussed in section 4.

behaviour of the light curve on time-scales shorter than  $\sim 8$  h and any features that appear on time-scales of under  $\sim 16$  h should be regarded as dubious given the noise properties of *Kepler*. We mark the noise-dominated region of the PSD plot in the inset in red.

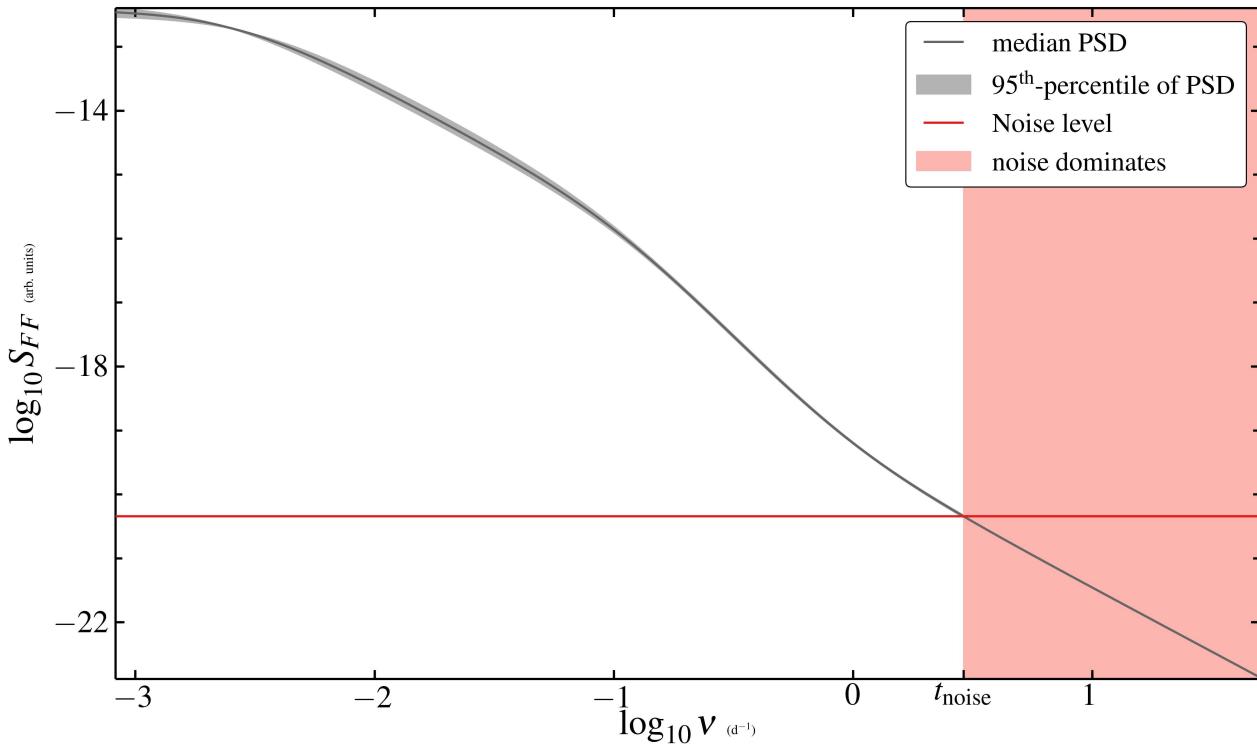
## 40 41 42 43 44 45 46 47 48 49 50 51 52 53 54 55 56 57 58 59 60 6.2 The Damped Harmonic Oscillator-like Behavior of the Light Curve of Zw 229-15

Since the Zw 229-15 light curve is well-described by a C-ARMA(2,1) process, the differential equation governing the behavior of flux perturbations may be interpreted as the damped harmonic oscillator driven by a colored stochastic process as described in section 4. Figure 4 shows the distribution of the un-damped-oscillator time period  $T_{\text{dHO}}$  against the damping ratio  $\zeta_{\text{dHO}}$ . The damping ratio is very large with median value 3.30 with 95-percent CI [2.78, 3.85] while the time period of the harmonic oscillator in the absence of damping is 59.9 d with 95-percent CI [56.9, 62.8] d. Physically this implies that flux perturbations slowly return to the steady-state flux level after reaching peak intensity rather than oscillating about the mean flux level. This observation supports the idea that flux perturbations may occur due to local Magneto Rotational Instability-generated hot- and cold-spots in the accretion disk that gradually disperse over time.

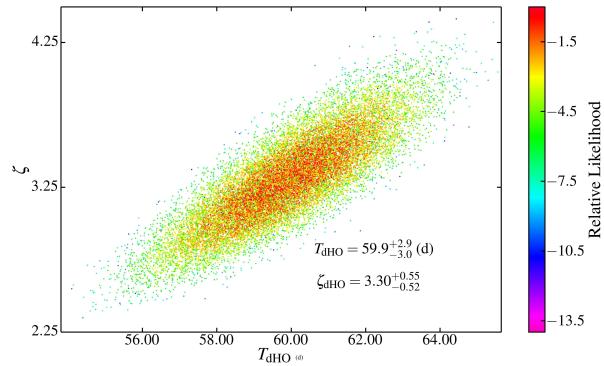
The corresponding Green's function is shown in figure 5 along with the 95 percent CI. Recall that the Green's function quantifies the evolution of a 'unit-impulse'. In this case,

the impulses consist of perturbations of the flux away from the steady-state level. The effect of an impulse is to drive the flux away from the mean-level until the rate at which the flux is changing drops to zero. This occurs  $t_{\text{max}} = 5.65$  d with 95-percent CI [5.50, 5.80] d after the original impulse (orange dashed line). On the left hand side of the figure, we see that on short time-scales (red-shaded region) the observation noise makes it impossible to observe structure in the Green's function. However, the  $t_{\text{max}} \sim 5.6$  d time-scale that we detect here is safely above  $t_{\text{noise}} \sim 8.3$  h and is therefore unlikely to be an artifact of the instrumentation-noise. Due to the very large damping ratio, flux perturbations do not dissipate rapidly—perturbations drop in intensity from the peak by a factor of e by  $t_{\text{e-fold}} = 68.7$  d with 95-percent CI [62.38, 75.08] d. On longer time-scales, the correlation between the total integrated flux from the accretion disk and the original perturbation will decrease as the flux becomes dominated by the increasingly larger numbers of intervening perturbations. It is plausible that conventional approaches such as PSD-fitting and structure function analysis detect the  $t_{\text{e-fold}}$  time-scale as the de-correlation time-scale.

Figure 6 shows the PSD of the disturbance properties inferred from the estimates of  $\beta_1$  and  $\beta_2$ , i.e., the power-spectrum of the impulses that drive the flux perturbations. On very long time-scales, the disturbance have equal power over  $\sim 3.5$  decades of frequency. On time-scales of  $\sim 1$  d, the disturbance PSD begins to rise with log-slope 2. Impulses with such PSD resemble 'violet noise'. Such PSD are produced by the thermal noise of the medium (Mellen 1952)



**Figure 3.** Power spectral density of the light curve of Zw 229-15. Using the *Kepler* noise properties  $\sigma_{N,\text{stated}}^2$  and  $M_{\text{Err}}$ , we compute the observation noise level (red line). The inferred PSD crosses the noise level at  $t_{\text{noise}} \sim 8.3$  h. On time-scales comparable to  $t_{\text{noise}}$ , i.e.,  $< 1$  d, the noise level does not permit us to infer anything about the behaviour of this light curve. Compare this PSD to the DHO PSD in figure 1.

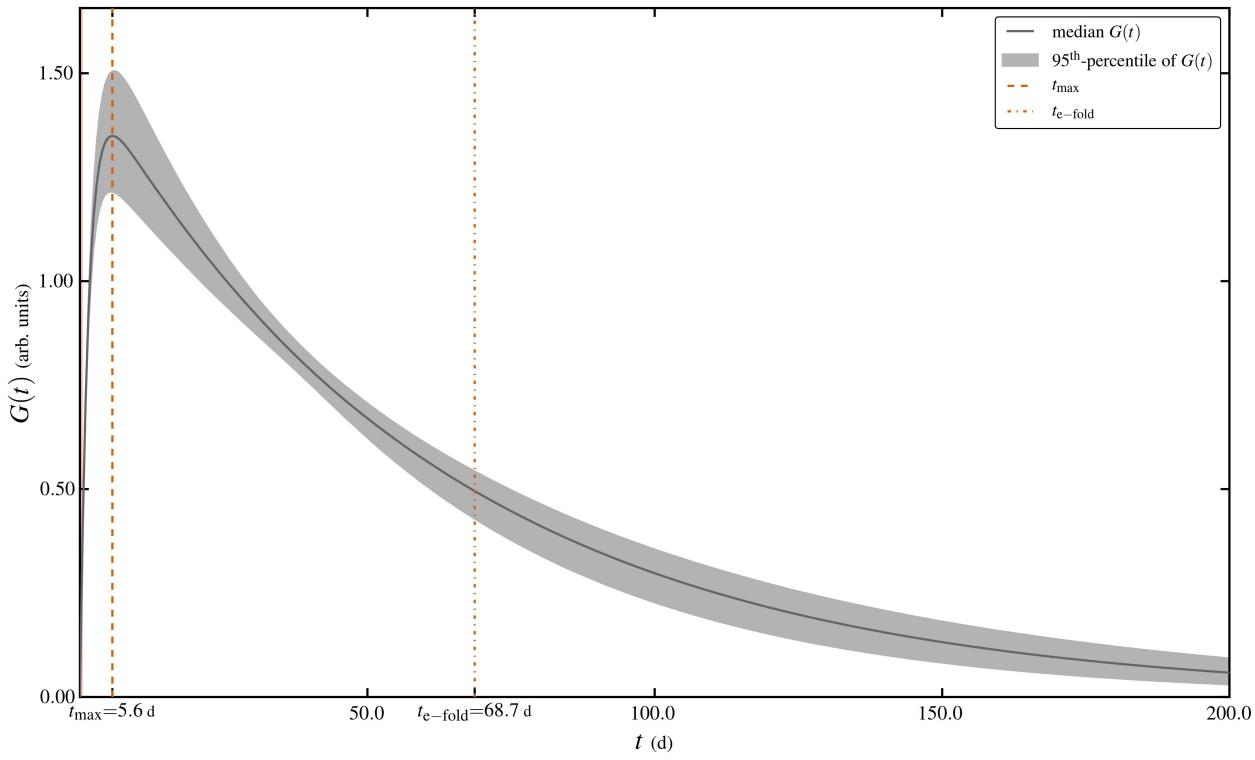


**Figure 4.** Damped Harmonic oscillator parametrization of the light curve of Zw 229-15. The time period of the oscillator is  $T$  and the damping ratio is  $\zeta$ . The perturbations in the accretion disk of Zw 229-15 are over-damped as indicated by the mean value of  $\zeta > 1$ . This indicates that flux perturbations steadily decay to the mean flux level without oscillating.

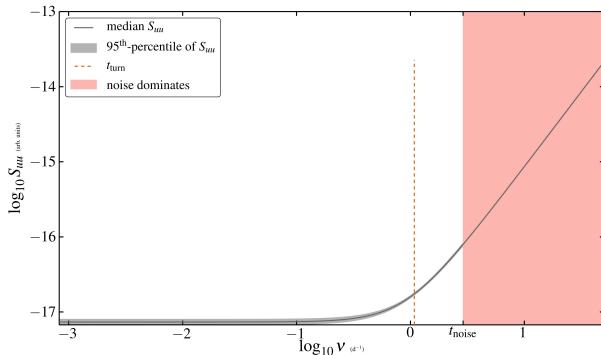
suggesting that on these very short time-scales, we may be probing the accretion disk matter directly. It may also be the case that we are detecting very low wavenumber, large lengthscale eddies in the turbulent flow (Miesch et al. 2015). Unfortunately, the proximity of the inferred turnover time-

scale  $t_{\text{turn}} = 1.081$  d with 95-percent CI [1.05, 1.11] d to the noise-dominated time-scale  $t_{\text{noise}} \sim 8.3$  h makes it difficult to draw conclusions about the disturbance PSD behaviour on these time-scales.

Using equation (12), the PSD of the Zw 229-15 light curve flattens on time-scales greater than  $t_{\text{flat}} = 386.78$  d with 95-percent CI [308.36, 472.36] d. Evidence of this flattening is barely visible at the lowest frequencies in figure 3. A visual indication of the flattening of the PSD would require that the sampled light curve be more than  $\sim 10\times$  the detected flattening time-scale. Is it possible to determine the optimal duration to sample the light curve in real-time, i.e., in the middle of obtaining the observations? During the course of the observations, we can use the Kalman filter to analyse the existing data and update the existing best estimate of  $t_{\text{flat}}$  as each new observation is made. Our prediction is that if the light curve is stationary, the estimates of  $t_{\text{flat}}$  will eventually converge to the true value. At this point we would have an estimate of the duration that we should sample the light curve if we desire to see the PSD flatten from the observations themselves. We also estimate  $t_{\text{steep}} = 9.29$  d with 95-percent CI [8.20, 10.66] d. This is *longer* than  $t_{\text{turn}}$ —we see that the effect of the moving average parameter  $\beta_1$  is to introduce short time-scale power into the light curve on time-scales under  $\sim 1$  d. Figure 7 shows cross-correlation plots of the various time-scales. Note that the estimates of  $t_{\text{e-fold}}$  and  $t_{\text{flat}}$  are poorly constrained. Future time-domain surveys



**Figure 5.** The Green’s function for the light curve of Zw 229-15 quantifies how a unit impulse flux perturbation evolves as a function of time. A tiny unit impulse perturbation at  $t = 0$  causes the flux from the perturbation to peak at  $t = t_{\max}$  ( $\sim 5.65$  d). The flux perturbation then decays exponentially with  $t_{\text{e-fold}} \sim 68.7$  d, dropping to under  $\sim 1$  percent the peak intensity  $\sim 320$  d later.



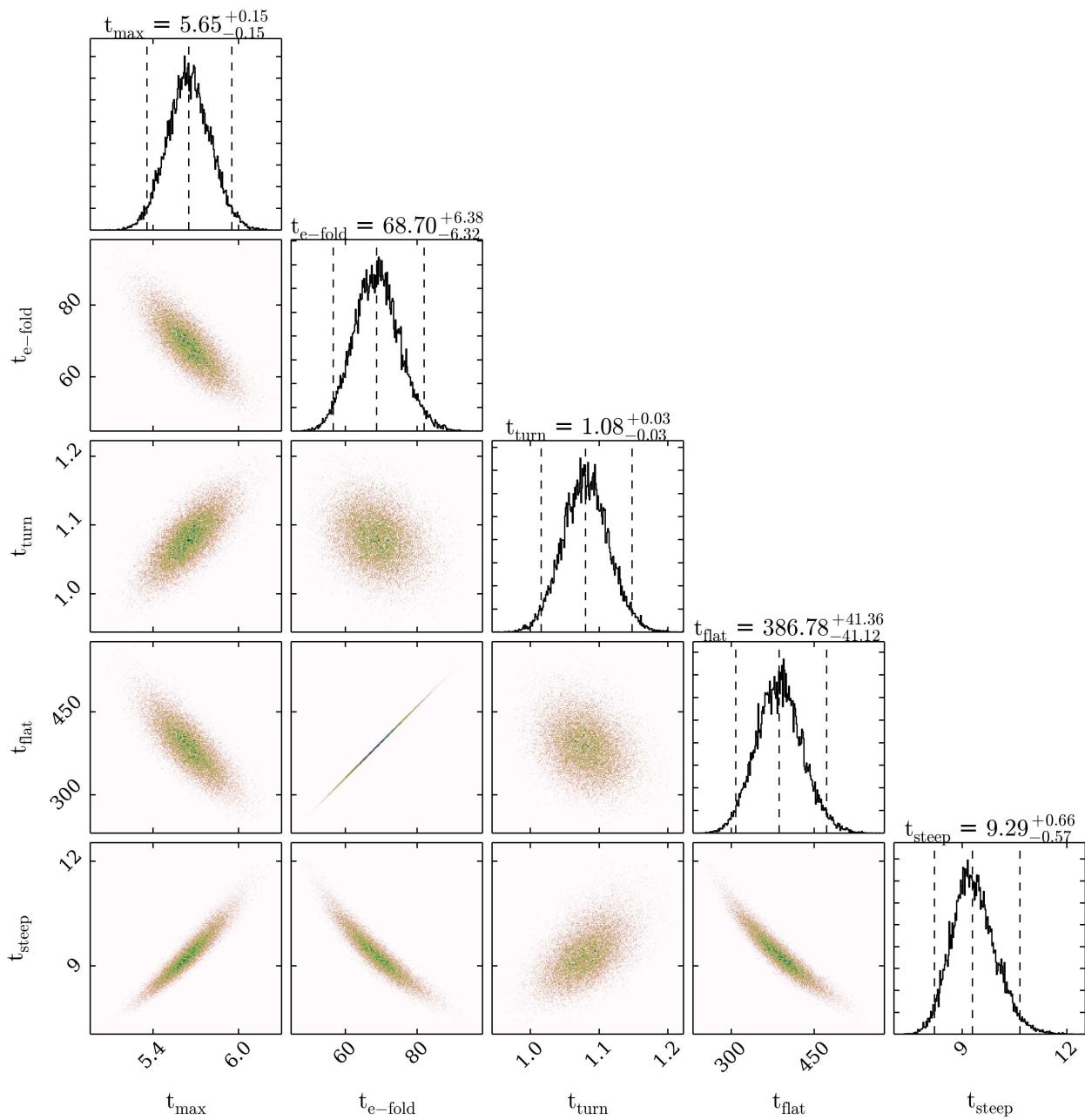
**Figure 6.** PSD of the flux perturbation impulses in the accretion disk of Zw 229-15. The PSD begins to steepen at  $t = t_{\text{turn}}$  ( $\sim 1$ ) d i.e. on time-scales longer than  $\sim 1$  d, the log PSD has no slope implying that the flux perturbation impulses behave like white noise. On shorter time-scales, the log PSD increases linearly with slope 2.

such as LSST will probe a much longer temporal baseline. This should help constrain both  $t_{\text{e-fold}}$  and  $t_{\text{flat}}$  more tightly.

If the C-ARMA model is not estimated correctly, mock light curves generated using the model will not resemble the observed light curve. To visually test the C-ARMA(2,1) model inferred by us, we randomly selected sets of model

parameter values, i.e., sets of values of  $\alpha_1$ ,  $\alpha_2$ ,  $\beta_1$ , and  $\beta_2$ , and generated simulated light curves using these parameters with the same sampling and noise levels as the true light curve. Figure 8 shows a sample of three such mock light curves along with the true light curve (orange). Visually all four light curves look very similar suggesting that the C-ARMA(2,1) is capturing most of the information in the light curve.

The light curve of Zw 229-15 observed by *Kepler* is well described by a C-ARMA(2,1) process suggesting that a 2<sup>nd</sup>-order (probably non-linear) differential equation may be responsible for smoothing out flux perturbations over time. The *Kepler* data suggest that the impulses to the flux that drive the perturbations have equal power on time-scales ranging from days to years. Perturbations to the flux behave like over-damped harmonic oscillators—the perturbation peaks in intensity at  $\sim 5.6$  d after which it gradually decays in intensity with an e-folding time-scale of  $\sim 69$  d. On very short time-scales ( $\sim 1$  d) there are indications that the impulse powerspectrum begins to rise with log-PSD slope 2 which may be due to the thermal motion of the accretion disk plasma. However, the observation noise level makes it difficult to probe accretion physics on these time-scales. We discuss the effect of the *Kepler* noise characteristics on our results in the next section.

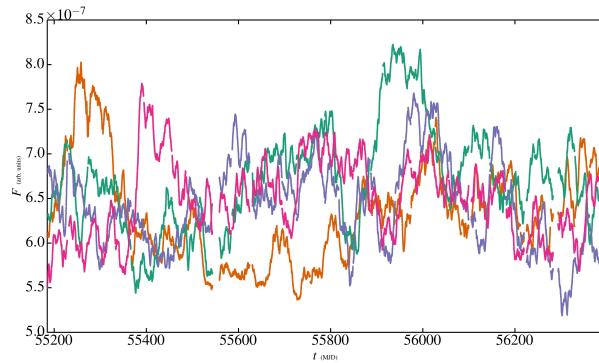


**Figure 7.** Scatter plots of  $t_{\text{max}}$  v/s  $t_{\text{e-fold}}$  v/s  $t_{\text{turn}}$  in units of d. The diagonal shows histograms of the estimates of these time-scales. The time taken to achieve peak output flux,  $t_{\text{max}}$ , is anti-correlated with the e-folding time of the flux perturbation  $t_{\text{e-fold}}$  but correlated with  $t_{\text{turn}}$ .

### 52 6.3 Moire Pattern Noise in the Light Curve of Zw 53 299-15

54 Instrument-induced artifacts were expected in the *Kepler* data prior to launch based on extensive ground-testing  
 55 ([Kolodziejczak et al. 2010](#)). Further testing was performed  
 56 during the commissioning of the telescope. Based on the testing,  
 57 it was determined that several artifacts did not require  
 58 mitigation for the primary science mission of *Kepler* : exo-  
 59

60 planet detection. [Kolodziejczak et al. \(2010\)](#) recommended  
 61 strategies for correcting or flagging several of the artifacts  
 62 generated by the *Kepler* detector electronics. Some of these  
 63 recommendations have been implemented in the new *Dynablock*  
 64 module ([Clarke et al. 2014](#)) that has been added to  
 65 the *Kepler* SOC pipeline for Data Release 24 ([Thompson  
66 et al. 2015](#)). The light curve that we analyze here is *not* created  
 67 using the per-pixel data provided by DR24—therefore



**Figure 8.** Simulated light curves with parameters drawn from the MCMC estimates of the C-ARMA(2,1) parameters of Zw 229-15. The true light curve of Zw 229-15 is also plotted. Can the true light curve be distinguished by eye?

we perform a test to quantify the effects of the various *Kepler* detector-induced artifacts on the results presented in section 6.2.

The *Kepler* focal plane science CCDs are read out over 84 channels that are referred to using the module#.output# convention presented in Van Cleve & Caldwell (2009) and in Kasliwal et al. (2015a). Zw 229-15 lands on the module#.output# combinations 14.48, 8.24, 12.40, and 18.64 as the *Kepler* focal plane performs a 90 degree rotation after every quarter, i.e., the first quarter of Zw 229-15 data, obtained during quarter 4 (Q4) of the *Kepler* mission were read out over 14.48. Zw 229-15 landed on 14.48 again during quarters Q8, Q12, and Q16. Of these module#.output# combinations, 14.48 exhibits moderate levels of rolling band and moire artifacts while 12.40 has an out of specification undershoot—see Van Cleve & Caldwell (2009) for a detailed explanation of the sources of these artifacts. We assess the impact of the rolling-band and moire artifacts by removing the affected quarters (Q4, Q8, Q12, and Q16) from the light curve and re-doing the C-ARMA model fit.

We find that the C-ARMA(2,1) model still provides the best AICc value, i.e., the order of the fit was not sensitive to the moire and rolling band issues. The inferred excess instrumentation noise is  $M_{\text{Err}} \sim 1.24$  which implies that the measurement errors are  $\sim 11.3$  percent higher than the quoted value. This suggests that there is a systematic underestimation of the *Kepler* observation noise level that cannot be attributed to the abnormal channel. From the inferred observation noise level, we find that the noise confusion limit drops slightly to  $t_{\text{noise}} = 7.9$  h. The largest change is manifest in the value of the damping ratio of the damped harmonic oscillator  $\zeta_{\text{dHO}} = 2.97$  with 95-percent CI [2.40, 3.57] while the time period of the harmonic oscillator remains  $T_{\text{dHO}} = 59.0$  d with 95-percent CI [55.4, 62.6] d. The Green's function peaks  $t_{\text{max}} = 5.89$  d after the initial perturbing impulse with 95-percent CI [5.54, 6.26] d, after which it begins to slowly decay with an e-folding time of  $t_{\text{e-fold}} = 61.7$  d with 95-percent CI [48.8, 76.0] d. The turnover time-scale on which the driving impulse PSD is observed to begin rising with log-PSD slope 2 occurs at  $t_{\text{turn}} = 1.16$  d with 95-percent CI [1.08, 1.24] d.

All of these estimates are consistent (to within the 95-percent confidence intervals computed) with the corre-

sponding values for the full light curve. This suggests that while the data obtained over module and output 14.48 have higher levels of observation noise and spacecraft-induced systematics, the observed behaviour of Zw 229-15 cannot be attributed to observational issues and has physical significance.

#### 6.4 Discussion

The optical variability properties Zw 229-15 has been extensively studied by multiple groups with various analysis techniques. Mushotzky et al. (2011) discovered the inconsistency of light curve of Zw 229-15 with the damped random walk (DRW) of Kelly et al. (2009) by analysing the PSD of the first 4 quarters of the light curve. They reported values for the short time-scale log-PSD slope between  $-2.96$  (Q8) and  $-3.31$  (Q6) by fitting the estimates of the PSD directly. Similarly high values were reported by Carini & Ryle (2012) who fit two PSD models to quarters 4 through 7 of the re-processing of light curve analysed here and found log-PSD slopes  $-2.88$  and  $-2.83$  for a knee and broken power-law model respectively.

The full *Kepler* light curve of Zw 229-15 was analysed using PSD and C-ARMA analysis methods by Edelson et al. (2014). Unlike the re-processing of Williams & Carini (2015) used by us and in Kasliwal et al. (2015b), the analysis in Edelson et al. (2014) used a much larger pixel-aperture (32-pixels) to minimize spacecraft-induced variability due to thermally-induced focus variations and differential velocity aberration—known issues with *Kepler* light curves (Kinemuchi et al. 2012). At the same time, using such large masks results in crowding problems—faint background sources bleed flux into the aperture because of the large *Kepler* PSF contaminating the light curve with their variability. The simple PSD analysis of the re-processed light curve suggests that on time-scales longer than  $\sim 5.6$  d, the PSD is well modelled by a power law with log-slope  $-2$ . On shorter time-scales, the log-PSD slope is much higher ( $-4.51$ ). There is an extra PSD component with log-PSD slope  $-1.28$  that contributes mostly at the 1 d time-scale. The C-ARMA analysis uses the same software package (CARMA\_PACK) used by us and found a bend time-scale of  $\sim 4$  d with log-PSD slopes  $-1.99$  and  $-3.65$  on longer and shorter time-scales. Edelson et al. (2014) tested the light curve for moire pattern noise and found evidence to suggest that moire pattern noise is a significant contaminant in both the flagged and un-flagged quarters.

Kasliwal et al. (2015a) analysed several *Kepler* MAST light curves that had been de-trended of spacecraft-induced variability by the *Kepler* SOC pipeline using co-trending basis vectors to quantify light curve features common to targets across the *Kepler* FOV. Using a structure function method to fit a bent power-law PSD model, Kasliwal et al. (2015a) found log-PSD slope  $-2.7$  with characteristic de-correlation time-scale 27.5 d.

Most recently, Williams & Carini (2015) applied a PSD analysis to the full light curve studied here and found log-PSD slope  $2.80 \pm 0.43$  with turnover time-scale  $66.94^{+14.3}_{-11.8}$  d. The same log-PSD slope was found by Kasliwal et al. (2015b) using the structure function approach of Kasliwal et al. (2015a) suggesting that the slope and time-scale found by

Williams & Carini (2015) is significant and not sensitive to the analysis technique.

Regardless of the mechanism for variability, certain time-scales of interest may be associated with a thin disk. The shortest time-scale is the light crossing time-scale

$$t_{lc} \sim \frac{r_{lc}}{c} < 11 \text{ d}, \quad (16)$$

where  $r_{lc}$  is the maximum radius that can be ascribed to  $t_{lc}$  (Peterson 1997), i.e., this time-scale places a (very conservative) upper limit on the size of the region that variations could originate from. The shortest time-scale on which variations should occur is the dynamical time-scale given by

$$t_{dyn} \sim \frac{r}{v_{Kep}} = \left( \frac{r^3}{GM_{BH}} \right)^{\frac{1}{2}} \sim 1.5 - 1500 \text{ d}, \quad (17)$$

where  $v_{Kep}$  is the Keplerian rotation velocity. Perturbations to hydrostatic equilibrium in the vertical ( $z$ ) direction are smoothed out on time-scale

$$t_z \sim \frac{H}{c_s} \sim t_{dyn}, \quad (18)$$

which may be observable in the form of quasi-periodic variations in the light curve. Deviations from thermal equilibrium, caused by fluctuations in the local dissipation rate, are damped out on the time-scale

$$t_{therm} \sim \frac{c_s^2 r^2}{v_{Kep}^2 \nu} \sim \frac{t_{dyn}}{\alpha} \sim 2 \text{ d} - 5 \text{ yr}. \quad (19)$$

Under the effect of viscous torques, matter diffuses through the disc on the viscous time-scale,

$$t_{visc} \sim \frac{r^2}{\nu} \sim \frac{t_{dyn}}{\alpha(H/r)^2} \sim 1 - 10 \text{ yr}. \quad (20)$$

The Barth et al. (2011) reverberation mapping study of Zw 229-15 suggests that  $M_{BH} = 1.0^{+0.19}_{-0.24} \times 10^7 M_\odot$  with  $r_{Sch} = 2GM/c^2 \sim 3 \times 10^{10} \text{ m}$  or 0.2 au. The bolometric luminosity was found to be  $L_{bol} = 6.4 \times 10^{43} \text{ erg s}^{-1}$  corresponding to  $L/L_{Edd} \sim 0.05$ . Assuming disk radii of  $100r_{Sch}$  (20 au) to  $10000r_{Sch}$  (2000 au), Shakura Sunyaev  $\alpha$  parameter value between 0.1 and 0.5, and disc height to radius ratios between 0.01 and 0.1, we can compute the various time-scales of equations (16) through (20).

The smallest viscous time-scale is just under 1 yr, suggesting that viscous fluctuations generated by the  $\alpha$  parameter variations of Lyubarskii (1997) are unlikely to be responsible for the observed variability.  $\alpha$  parameter variations in the Lyubarskii (1997) theory result in inward propagating fluctuations characterized by correlations between variability in different bands with hard lags, i.e., shorter wavelength bands lag longer wavelength bands. This phenomenon is typically observed only in the X-rays (Vaughan & Fabian 2004; McHardy et al. 2004; Arévalo et al. 2006), whereas in the UVOIR the situation is reversed with longer wavelength bands lagging shorter wavelength bands (Wanders et al. 1997; Sergeev et al. 2005).

Estimates of the thermal time-scales can range between  $\sim 3 \text{ d}$  ( $r = 10^2 r_{Sch}$  &  $\alpha = 0.5$ ) and  $\sim 5 \text{ yr}$  ( $r = 10^4 r_{Sch}$  &  $\alpha = 0.1$ ), making it plausible to associate the  $t_{e-fold} = 69 \text{ d}$  time-scale with the thermal time-scale of the disc. This may imply that the variability observed by *Kepler* in the optical is driven by the unobserved X-ray variability of Zw 229-15

(Krolik et al. 1991). However, it is now known that while X-ray variability drives small scale fluctuations in optical variability on short time-scales, longer time-scale large amplitude fluctuations exist in the optical that are too energetic to be driven by re-processed X-ray emission (Uttley et al. 2003; Arévalo 2009). There is strong evidence to suggest that there are two sources of variability in the optical-reprocessing of X-ray emissions that drives short time-scale low-amplitude variability, and a process local to the optical emitting region of the disk that drives larger amplitude, longer time-scale changes (Gaskell 2008). While the *Kepler* light curve of Zw 229-15 is long enough to probe the days-weeks time-scales associated with the re-processing of the X-ray variability, it may not be sensitive to the longer time-scale variability intrinsic to the optical.

The dynamical time-scale of the disk ranges between  $\sim 1.6$  and  $\sim 1600 \text{ d}$  making it plausible to associate the observed  $t_{e-fold}$  with this quantity as well. The dynamical time-scale characterizes states that are perturbed from dynamical equilibrium such as by global g-modes (Reynolds & Miller 2009; O'Neill et al. 2009). The light crossing time-scale is smaller than  $\sim 11 \text{ d}$  making it unsuitable to associate with any of the time-scales that we have measured.

We have successfully recovered the  $\sim 5.6 \text{ d}$  time-scale found by Edelson et al. (2014). We interpret it as the time lag at which the Green's function peaks, i.e., the duration between the initial impulse driving a flux perturbation and the peak in emitted flux. We have found that there is some indication of excess power at very high frequencies ( $> 1 \text{ d}^{-1}$ ) though the observation noise level makes the detection dubious. We treat this excess power as belonging purely to the disturbance process that generates the stochastic variability, in which case it may be caused by the thermal noise of emitting medium. We identify the 67 d time-scale reported by Williams & Carini (2015) with the e-folding time-scale  $t_{e-fold}$  of a flux perturbation in the accretion disk of the AGN, i.e., it is the duration over which flux perturbations decay by a factor of e. Finally, we point out that the exact slope of the log-PSD is a function of frequency. From the C-ARMA(2,1) PSD in equation (8), we see that over a considerable range of frequencies, one may expect the log-PSD to have slope close to  $-2$ . At very high frequencies, the log-PSD approaches  $-4$  as the  $\nu^4$  term in the denominator of equation (8) begins to dominate. At very low frequencies, we expect the PSD to flatten. This is exactly the behavior inferred by all the studies of Zw 229-15 and is predicted of the C-ARMA(2,1) process. The reason why different studies have fit a range of slopes at different frequencies may be attributable to variations between the exact processing used but is likely primarily due to sensitivity to different parts of the PSD.

## 7 CONCLUSIONS

We demonstrate that the Continuous-time AutoRegressive Moving Average (C-ARMA) model of Kelly et al. (2014) results from linearization of arbitrary non-linear differential equations for the flux emitted by the AGN accretion disk. Small perturbations of the total flux emitted by an AGN can be modelled in the linear-regime as a linear differential equation driven by noise or a C-ARMA process. Such pro-

cesses consist of an  $n^{\text{th}}$ -order linear differential equation for the flux perturbation (LHS) stochastically driven by a linear combination of Wiener increments (RHS). The driving impulses may characterize complex MHD processes such as MRI turbulence. We show how insights into the variability-driving physics can be obtained by examining the PSD of the flux impulses. We propose that the homogenized linear differential equation corresponding to the C-ARMA process governs how flux perturbations evolve after they have been generated. We show how the evolution of the flux perturbations can be characterized by the Green's function of the linear-differential equation.

We propose the use of a new representation of the C-ARMA process in state-space form based on the observable companion form of a linear time-invariant system. This representation puts the dynamics of the flux variations into the state-equation and possesses a very simple form for the observation equation. We argue that this representation is ideally suited for analyzing well-sampled flux light curves with constant sampling rate, especially in the presence of missing observations. Using the representation that we propose, the Kalman filtering and RTS smoothing equations can be applied to infer the values of the parameters the C-ARMA model and provide an estimate of the true light curve sans observation error.

We present a detailed analysis of the C-ARMA(2,1) model or damped harmonic oscillator driven by violet noise. The C-ARMA(2,1) process may be well suited for modelling AGN light curves based on light curves from the *Kepler* mission and the Sloan Digital Sky Survey. We demonstrate how the Green's function increases steeply over a characteristic time-scale  $t_{\max}$  before decaying exponentially with e-folding time-scale  $t_{\text{e-fold}}$ . The driving disturbances of the C-ARMA(2,1) model possess a flat PSD on long time-scales. On time-scales shorter than  $t_{\text{turn}}$ , the PSD of the driving disturbances increases as  $v^2$ . We suggest that this short time-scale behaviour of the PSD may arise from the thermal motion of the material in the accretion disk.

We analyze a custom re-processing of the optical light curve of the Seyfert 1 galaxy Zw 229-15 presented in Williams & Carini (2015) using the C-ARMA formalism. We find that the C-ARMA(2,1) model of section 4 well characterizes the observed flux variations. Fluctuations in the light curve of Zw 229-15 are strongly damped with damping ratio  $\zeta = 3.3$ . Thus, perturbations to the flux result in a smooth decay to the steady-state flux level with no oscillatory behavior. The Green's function of the flux perturbations is observed to peak on time-scale  $t_{\max} = 5.6$  d which is consistent with the PSD time-scale reported by Edelson et al. (2014). After the initial peak, flux perturbations decay with e-folding time-scale  $t_{\text{e-fold}} = 69$  d that is consistent with the PSD time-scale reported by Williams & Carini (2015). We find that on time-scales shorter than  $\sim 1$  d, the PSD of the disturbances rises steeply as  $v^2$  as reported by Edelson et al. (2014), however the measurement noise level of *Kepler* makes this finding tentative. We hunt for non-physical spacecraft-induced contributions to the observed behavior by examining the effect of moire-pattern noise on our results. While the moire-pattern noise affected quarters have slightly higher measurement noise than the clean quarters, the overall model-fit does not change significantly when excluding the noisy quarters. We conclude that the variability behav-

ior observed in Zw 229-15 by *Kepler* is consistent with an over-damped harmonic oscillator driven by a colored-noise process. We demonstrate how breaks and features found in conventional PSD analysis of light curves may be interpreted in the context of characteristic time-scales that arise from the C-ARMA(2,1) model.

C-ARMA analysis techniques offer tremendous promise for studying AGN & BHXR variability. The C-ARMA formalism can be extended to model different accretion states using formalisms such the Threshold C-ARMA processes of Brockwell (1994). C-ARMA processes can also be extended to non-simultaneous multi-wavelength observations to probe the connection between the variability observed in different bands. Numerical simulations of accretion disks are already able to produce time-domain synthetic spectra (Schmitman et al. 2013) and will soon robustly include radiative transfer in the modelling (Fragile 2014). C-ARMA models may be useful in comparing the synthetic spectra and light curves produced by numerical models of the accretion disk and winds to observational results. Moving forward, more sophisticated techniques such as the particle filter may be able to directly infer the underlying non-linear stochastic differential equations governing variability without the need for linearization (Hanif & Protopapas 2015). Tools for performing C-ARMA analysis are provided in the C++ and PYTHON package KÄLÍ and can be obtained from <https://github.com/AstroVPK/libcarma>.

## ACKNOWLEDGEMENTS

We acknowledge support from NASA grant NNX14AL56G. This paper includes data collected by the Kepler mission. Funding for the Kepler mission is provided by the NASA Science Mission directorate. Some of the data presented in this paper were obtained from the Mikulski Archive for Space Telescopes (MAST). STScI is operated by the Association of Universities for Research in Astronomy, Inc., under NASA contract NAS5-26555. Support for MAST for non-HST data is provided by the NASA Office of Space Science via grant NNX09AF08G and by other grants and contracts. KÄLÍ uses the optimization libraries provided by Steven G. Johnson, The NLOpt nonlinear-optimization package, <http://ab-initio.mit.edu/nlopt>. We wish to thank Jack O'Brien, Jackie Moreno, Nate B. Lust and Adam Lidz for their input and discussions.

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**APPENDIX A: HOW DO FLUX PERTURBATIONS EVOLVE?**

We assume that the variations in AGN luminosity may be ascribed to flux perturbations intrinsic to the accretion disk, corona, and winds. It is likely that variability originates from multiple mechanisms simultaneously. These mechanisms are almost certainly non-linear (Uttley et al. 2005). Hence the total flux emitted by the AGN evolves via a (probably non-linear) differential equation such as the integral of the surface mass density equation (see Lightman & Eardley 1974, eq. 4) that governs how perturbations evolve. We shall examine how linear perturbations of the total flux evolve. Let the total flux emitted by an AGN be  $F(t)$ . Suppose  $F$  obeys the  $p^{\text{th}}$ -order non-linear differential equation

$$\frac{d^p F}{dt^p} = g(F, \frac{dF}{dt}, \dots, \frac{d^{p-1} F}{dt^{p-1}}, \mathcal{F}, t), \quad (\text{A1})$$

where  $\mathcal{F}$  is the contribution to the total flux from the appearance of hot- and cold-spots in the accretion disk. We shall attempt to linearise this system by probing the behaviour of a small perturbation in flux about a solution to this equation (Wiberg 1971; Stengel 1994). We begin by re-writing this  $n^{\text{th}}$ -order non-linear differential equation for  $F(t)$  as a series of coupled 1<sup>st</sup>-order non-linear differential equations for  $F(t)$  and its derivatives. Define

$$F_k(t) = \frac{d^{k-1} F(t)}{dt^{k-1}}, \quad (\text{A2})$$

for  $1 \leq k \leq n$ . Since

$$\frac{dF_k}{dt} = \frac{d}{dt} \left( \frac{d^{k-1} F}{dt^{k-1}} \right) = \frac{d^k F}{dt^k} = F_{k+1}, \quad (\text{A3})$$

we have

$$\begin{aligned} \frac{dF_1}{dt} &= F_2, \\ &\vdots \\ \frac{dF_{p-1}}{dt} &= F_p. \end{aligned} \quad (\text{A4})$$

Then, we may rewrite equation (A1) as

$$\frac{dF_p}{dt} = g(F_1, F_2, \dots, F_p, \mathcal{F}, t). \quad (\text{A5})$$

Non-linear differential equations can be difficult to treat analytically. However, we are not interested in the behaviour of the total flux. Rather, we are interested in how small perturbations ( $\delta F/F \ll 1$ ) from a known solution to the non-linear equations behave. In order to understand how small perturbations behave, we shall linearize this system of equations by looking at the form of the equations in the presence of a small perturbation. Starting from initial conditions  $F_1(t_0), \dots, F_p(t_0)$  with input contributions  $\mathcal{F}_0(t)$ , suppose our system evolves as per the solutions  $\phi_1(t), \dots, \phi_p(t)$ , i.e., these solutions obey equation (A5) under the constraints of equation (A4).  $\phi_1(t)$  is the flux emitted by the object at time  $t$  while  $\phi_k$  with  $1 < k \leq p$  are the  $k^{\text{th}}$  derivatives of the flux. Now consider the evolution of the system starting from the (slightly perturbed) initial conditions  $F'_1(t_0) = F_1(t_0) + X_1(t_0), \dots, F'_p(t_0) = F_p(t_0) + X_p(t_0)$  where the various  $X_k$  are small. Furthermore, we shall also subject the system to a slightly perturbed input  $\mathcal{F}'_0(t) = \mathcal{F}_0(t) + X_0(t)$ . Then, the perturbed

solutions,  $\phi_1(t) + \chi_1(t), \dots, \phi_p(t) + \chi_p(t)$ , must continue to satisfy equation (A5) under the constraints of equation (A4). We have introduced  $\chi_k(t)$  with  $1 \leq k \leq p$  as small perturbations to the original solutions  $\phi_k$ , i.e.,  $\chi_1(t)$  is a small perturbation to the flux  $\phi_1(t)$  at time  $t$  while the  $\chi_k(t)$  are small perturbations of the derivatives of the flux. Starting with the constraint equations, we have

$$\begin{aligned} \frac{d(\phi_1 + \chi_1)}{dt} &= \frac{d\phi_1}{dt} + \frac{d\chi_1}{dt} = \phi_2 + \chi_2, \\ &\vdots \end{aligned} \quad (\text{A6})$$

But since the  $\phi_k$  satisfy equation (A4), we have

$$\begin{aligned} \frac{d\chi_1}{dt} &= \chi_2, \\ &\vdots \\ \frac{d\chi_{p-1}}{dt} &= \chi_p. \end{aligned} \quad (\text{A7})$$

To linearize equation (A5), we begin by noting

$$\frac{d(\phi_p + \chi_p)}{dt} = \frac{d\phi_p}{dt} + \frac{d\chi_p}{dt} = g(F'_1, F'_2, \dots, F'_p, \mathcal{F}'_0, t). \quad (\text{A8})$$

We may Taylor expand the RHS of this equation about the  $\phi_1, \dots, \phi_p$  and  $\mathcal{F}_0$  and cancel  $d\phi_p/dt = f(\phi_1, \phi_2, \dots, \phi_p, \mathcal{F}_0, t)$  to obtain

$$\frac{d\chi_p}{dt} = \frac{\partial g}{\partial F_1} \Big|_{\phi_1} \chi_1 + \dots + \frac{\partial g}{\partial F_p} \Big|_{\phi_p} \chi_p + \frac{\partial g}{\partial \mathcal{F}_0} \Big|_{\mathcal{F}_0} \mathcal{F}'_0. \quad (\text{A9})$$

Equations (A7) and (A9) are the linearised, 1<sup>st</sup>-order system of equations that describe the evolution of a tiny perturbation of the flux about an exact solution. In general, the coefficients of the  $\chi_k$  in equation (A9) are functions of time and change as the system evolves. Let us consider the special case where the exact solution is the steady-state solution (if it exists). Under this restriction, equation (A9) governing the evolution of a tiny perturbation in the flux simplifies to

$$\frac{d\chi_p}{dt} = a_1 \chi_1 + a_2 \chi_2 + \dots + a_n \chi_p + b \mathcal{F}'_0, \quad (\text{A10})$$

where

$$a_k = \frac{\partial g}{\partial F_k} \Big|_{\phi_k}, \quad (\text{A11})$$

and

$$b = \frac{\partial g}{\partial \mathcal{F}_0} \Big|_{\mathcal{F}_0}. \quad (\text{A12})$$

Therefore, we may re-write the linearised system of equations in equation (A10) along with the constraints in equation (A7) in the form of the n<sup>th</sup>-order linear differential equation

$$\frac{d^p \chi}{dt^p} + a_1 \frac{d^{p-1} \chi}{dt^{p-1}} + \dots + a_{p-1} \frac{d \chi}{dt} + a_p \chi = u, \quad (\text{A13})$$

where we have set  $a_k = -a_k$  and  $u = b \mathcal{F}'_0$  for brevity. Using differentials, this equation may be written as

$$d^p \chi + a_1 d^{p-1} \chi + \dots + a_{p-1} d \chi + a_p \chi = u, \quad (\text{A14})$$

where  $du$  is an increment of  $u$ . We see that the LHS of this equation is identical to the LHS of the C-ARMA process

equation (1) suggesting that a C-ARMA process may be a good model for small flux variations.

Dimensional consistency requires that

$$[\alpha_k] = T^{-k}, \quad (\text{A15})$$

where  $T$  is time, i.e., the  $\alpha_k$  have units that are powers of various frequencies. Since equation (A13) governs how flux perturbations evolve over time, we shall refer to this equation as the ‘Dissipation Equation’. We define the characteristic polynomial of the LHS of this equation as

$$\alpha(z) = z^p + \alpha_1 z^{p-1} + \dots + \alpha_{p-1} z + \alpha_p, \quad (\text{A16})$$

with roots  $\rho_k$ . The flux perturbations are stable, i.e., do not increase without bound, if  $\text{Re}(\rho_k) < 0$ . We shall posit a form for  $u$  in section C.

We find that if the total flux emitted by the accretion disk obeys a non-linear differential equation, we can linearize that equation to determine the behaviour of small perturbations of the flux from the steady-state solution. The linear differential equation that governs how a small flux perturbation evolves is identical to the C-ARMA process equation suggesting that such processes may be good models for the stochastic variability seen in AGN accretion disks. In the next section, we present a convenient physical interpretation of equation (A13) using the method of the Green’s function.

## APPENDIX B: HOW TO COMPUTE THE GREEN’S FUNCTION

We compute the Green’s function of equation (A13) by finding a solution of the homogenized version of the equation. Consider the homogeneous differential equation corresponding to equation (A13)

$$\frac{d^p f}{dt^p} + \alpha_1 \frac{d^{p-1} f}{dt^{p-1}} + \dots + \alpha_{p-1} \frac{df}{dt} + \alpha_p f = 0. \quad (\text{B1})$$

To obtain the Green’s function solution of this equation, we shall drive this equation with a unit impulse, i.e., a Dirac delta function, located at  $t = 0$

$$\frac{d^p G}{dt^p} + \alpha_1 \frac{d^{p-1} G}{dt^{p-1}} + \dots + \alpha_{p-1} \frac{dG}{dt} + \alpha_p G = \delta(t), \quad (\text{B2})$$

where  $G(t)$  is the desired Green’s function solution. To find this solution, we shall first solve equation (B1) by finding the roots of the characteristic polynomial of equation (A16) and then apply boundary conditions based on the properties of  $\delta(x)$ . From the definition of the characteristic polynomial in equation (A16), the (distinct) roots are  $\rho_k$  with  $\text{Re}(\rho_k) < 0$  for stability. Then the solution of the homogeneous equation (B1) is

$$f(t) = \sum_{k=1}^n c_k e^{\rho_k t}, \quad (\text{B3})$$

where the  $c_k$  are constants. Hence the solution of the inhomogeneous equation (B2) may be found by using the continuity properties of  $G(t)$  and its derivatives. For  $t < 0$ ,  $d^{p-k} G/dt^{p-k} = 0$  for  $1 \leq k < p$  while at  $t = 0$ ,  $d^p G/dt^p$  has the same type of discontinuity as  $\delta(t)$ . Therefore  $d^{p-1} G/dt^{p-1} = 1$  when  $t = 0$ , i.e., it behaves like the step function. Next, we have that  $d^{p-k} G/dt^{p-k} = 0$  at  $t = 0$  for all  $1 < k \leq p$  since each successive lower-order derivative is

a  $(k-1)^{\text{th}}$ -order polynomial located at the origin. Following this line of reasoning, we arrive at the following boundary conditions

$$\frac{d^{p-1}G}{dt^{p-1}} \Big|_{t=0} = 1, \quad (\text{B4})$$

and

$$\frac{d^{p-k}G}{dt^{p-k}} \Big|_{t=0} = 0, \quad (\text{B5})$$

for all  $2 \leq k \leq p$ . Computation of the derivatives in equations (B4) and (B5) results in

$$\sum_{k=1}^p c_k \rho_k^{p-1} = 1, \quad (\text{B6})$$

and

$$\sum_{k=1}^p c_k \rho_k^{k-1} = 0, \quad (\text{B7})$$

for all  $2 \leq k \leq p$ . The  $c_k$  may then be found by solving the matrix equation

$$\begin{pmatrix} \rho_1^{p-1} & \rho_2^{p-1} & \dots & \rho_{p-1}^{p-1} & \rho_p^{p-1} \\ \rho_1^{p-2} & \rho_2^{p-2} & \dots & \rho_{p-1}^{p-2} & \rho_p^{p-2} \\ \vdots & \vdots & \ddots & \vdots & \vdots \\ \rho_1 & \rho_2 & \dots & \rho_{p-1} & \rho_p \\ 1 & 1 & \dots & 1 & 1 \end{pmatrix} \begin{pmatrix} c_1 \\ c_2 \\ \vdots \\ c_{p-1} \\ c_p \end{pmatrix} = \begin{pmatrix} 1 \\ 0 \\ \vdots \\ 0 \\ 0 \end{pmatrix}. \quad (\text{B8})$$

In short, the Green's function for equation (A13) is given by

$$G(t) = \sum_{k=1}^p c_k e^{\rho_k t}, \quad (\text{B9})$$

where the  $c_k$  satisfy equation (B8).

## APPENDIX C: THE MATHEMATICS OF THE DRIVING STOCHASTIC IMPULSES

Fundamental to the treatment of the input stochastic perturbations, i.e., the  $u$  of equation (A14) as some form of noise process is the concept of the Wiener process  $W$ , also known as Brownian motion (Doob 1990; Davis 2002; Jacobs 2010; Øksendal 2014). Wiener processes and processes derived from it are of importance in a number of areas of physics, engineering and finance (Jacobs 2010). For example, the random motion of the electrons in a strip of metal produces a small, stochastically fluctuating potential known as Johnson noise (Gillespie 1996) that can be modelled using Wiener increments. The Wiener process cannot be used directly as the driving process because the likelihood of a given value of the Wiener process is time-dependent, i.e., Wiener processes are not *stationary*. The power spectral density (PSD) of the Wiener process decreases as the square of the frequency and so has excess power on long time-scales. Processes with a flat PSD are described in the engineering literature as ‘white’ noise because they have equal power at all frequencies in analogy with white light possessing equal amounts of light of all colors. Analogously, processes with

non-flat PSD are described as ‘colored’ with various adjectives such as ‘blue’, ‘red’, and ‘pink’, being used to describe the relative amplitude of the PSD at various frequencies. In this sense, the Weiner processes is known in the engineering literature as ‘red’-noise. A better choice for the driving process is some linear combination of the differentials of  $W$ , i.e., increments of  $W$ . We use differentials because mathematically, due to the discontinuous nature of  $W$ , the derivatives of  $W$  are undefined. Writing  $u$  as a linear combination of the differentials of  $W$ , we obtain

$$u = \beta_0(dW) + \beta_1 d(dW) + \dots + \beta_{p-2} d^{p-2}(dW) + \beta_{p-1} d^{p-1}(dW) \quad (\text{C1})$$

where  $dW$  is the differential of the Wiener process and is mathematically well-defined. By definition

$$W(t) = \int_0^t dW. \quad (\text{C2})$$

Dimensional consistency requires that

$$[\beta_k] = [\chi] T^{k+\frac{1}{2}-p}, \quad (\text{C3})$$

where  $\chi$  is in units of flux, e.g. erg cm $^{-2}$ . Analogous to equation (A16), we define the characteristic polynomial of equation (C1) to be

$$\beta(z) = \beta_0 + \beta_1 z + \dots + \beta_{p-2} z^{p-2} + \beta_{p-1} z^{p-1}, \quad (\text{C4})$$

with roots  $\mu_k$ . For equation (C1) to be invertible, we require that  $\text{Re}(\mu_k) < 0$ .

The power spectral density (PSD),  $S_{uu}(\nu)$ , of the process  $u(t)$  is defined to be

$$S_{uu}(\nu) = \langle |\tilde{u}(\nu)|^2 \rangle, \quad (\text{C5})$$

where

$$\tilde{u}(\nu) = \frac{1}{2\pi} \int_{-\infty}^{\infty} u(t) e^{-2\pi i \nu t} dt \quad (\text{C6})$$

is the Fourier transform of  $u(t)$ , i.e.,  $\tilde{u}(\nu)$  denotes the Fourier transform of  $u(t)$ . The PSD of the Wiener increments is constant at all frequencies (Brockwell 2001), i.e.,

$$S_{dWdW} = \frac{1}{2\pi}, \quad (\text{C7})$$

corresponds to infinite total power when integrated over all frequencies. Real processes only resemble the Wiener process and Wiener increments over an interval of frequencies known as the bandpass of the process. Outside this bandpass, the PSD is set to 0. The PSD of the driving process in equation (C1) may be computed using

$$\widetilde{\frac{d^k u}{dt^k}} = (2\pi i \nu)^k \tilde{u}, \quad (\text{C8})$$

in conjunction with equation (C7) yielding

$$S_{uu}(\nu) = \frac{1}{2\pi} |\beta(2\pi i \nu)|^2, \quad (\text{C9})$$

where  $\beta$  is the moving average polynomial. The PSD of the driving impulses of equation (C9) is a polynomial in  $\nu$ . This implies that at high frequencies, the power contributed by this process is unbounded. However, consider what happens to the PSD when  $dW$  is integrated. The integral of the Wiener increments, i.e., the Wiener process itself, has PSD

$S_{WW} \propto 1/\nu^2$ . Successive integrations decrease the logarithmic slope of the PSD at high frequencies by a factor of 2 each time. If  $u$  is supplied as input to equation (A13), it is integrated  $n$ -times. If the overall process  $\chi$  is to have finite power at high frequencies, we *must* limit the highest-order of the derivative of  $dW$  in equation (C1) to  $p - 1$  so that after  $p$ -integrations, the overall process has PSD with negative logarithmic slope corresponding to finite high-frequency power, i.e., the bandpass of the driving disturbance process rapidly decreases at high frequencies. At low frequencies, we expect the power in our system to drop to 0 as we approach the time-scale on which AGN are active  $\sim 10^7$  yr. On the more relevant time-scales that can be probed by us, the PSD should essentially flatten to a constant at low frequencies. Finally, we may formulate our system by combining equations (A14) and (C1) into the Itô differential equation

$$d^p \chi + \alpha_1 d^{p-1} \chi + \dots + \alpha_{p-1} d\chi + \alpha_p \chi = \\ \beta_0 (dW) + \beta_1 d(dW) + \dots + \beta_{p-1} d^{p-1} (dW), \quad (\text{C10})$$

with  $\text{Re}(\rho_k) < 0$  and  $\text{Re}(\mu_k) < 0$  where the  $\rho_k$  and  $\mu_k$  are the roots of the autoregressive- and moving average- polynomials of equations (A16) and (C4) respectively. Itô differential equations arise when the solution of the differential equation requires the integration of a stochastic function against Wiener increments. Such equations can be solved by using the methods of the Itô calculus (Øksendal 2014). Given the form of the C-ARMA process, we can write the PSD of the C-ARMA process (Brockwell 2001) as

$$S_{\chi\chi}(\nu) = \frac{1}{2\pi} \frac{|\beta(2\pi\nu)|^2}{|\alpha(2\pi\nu)|^2}. \quad (\text{C11})$$

In practice not all of the moving average terms,  $\beta_k$ , may be present. In particular, if the first  $k$  moving average terms are absent, i.e.,  $\beta_l = 0$  for  $l < k + 1$ , the RHS has  $q = p - k$  terms. A stochastic processes of this form is known in the statistical inferencing community as a Continuous-time AutoRegressive Moving Average process of order  $p$  and  $q$  or C-ARMA( $p,q$ ) process (Brockwell 1994, 2001, 2014; Kelly et al. 2014). The maximum allowed value of  $q$  is dictated by the value of  $p$ , i.e.,  $q < p$ .

The PSD of equation (C11) can be expressed as

$$S_{\chi\chi}(\nu) = \frac{1}{2\pi} \frac{\sum_{i,j=0}^p (2\pi\nu)^{2p-i-j} \beta_i \beta_j (-1)^{p-j}}{\sum_{m,n=0}^q (2\pi\nu)^{2q-m-n} \alpha_m \alpha_n (-1)^{q-n}}, \quad (\text{C12})$$

with  $\alpha_0 \equiv 1$ . Consider the terms with  $2p - i - j$  odd, i.e., complex terms. Since  $2p$  is always even, either  $i$  or  $j$  must be odd. While  $(2\pi\nu)^{2p-i-j}$  and  $\beta_i \beta_j$  are symmetric in  $i$  &  $j$ ,  $(-1)^{p-j}$  is not. Interchanging  $i$  &  $j$  leads us to the conclusion that odd (complex-valued) terms occur in pairs  $(2\pi\nu)^{2p-i-j} \beta_i \beta_j$  and  $-(2\pi\nu)^{2p-j-i} \beta_j \beta_i$  that cancel each other. We see that the numerator is purely real and consists of a sum of terms that are proportional to even powers of frequency. The same logic applies to the denominator.

Defining  $k = 2q - i - j$  &  $l = 2p - m - n$  allows us to rewrite equation (C13) as

$$S_{\chi\chi}(\nu) = \frac{N(\nu)}{D(\nu)} = \\ \frac{1}{2\pi} \frac{\sum_{i=0}^q \sum_{k=2q-i}^{3q-i} (2\pi\nu)^k r^k (-1)^{3q-i-k} \beta_i \beta_{k+i-2q}}{\sum_{m=0}^p \sum_{l=2p-m}^{3p-m} (2\pi\nu)^l l^l (-1)^{3p-m-l} \alpha_m \alpha_{l+m-2p}}, \quad (\text{C13})$$

from which we see that the  $\nu^{k-th}$  power term in the numerator is given by

$$N_{\nu^k}(\nu) = \sum_{i=0}^q (2\pi\nu)^k i^k (-1)^{3q-i-k} \beta_i \beta_{k+i-2q}, \quad (\text{C14})$$

and the  $\nu^{l-th}$  power term in the denominator is given by

$$D_{\nu^l}(\nu) = \sum_{m=0}^p (2\pi\nu)^l l^l (-1)^{3p-m-l} \alpha_m \alpha_{l+m-2p}. \quad (\text{C15})$$

The Wiener-Khintchine Theorem lets us compute the corresponding *Auto-Covariance Function* (ACVF) of the observed light curve (Brockwell 2001) as

$$\gamma_{\chi\chi}(\tau) = \int_{-\infty}^{\infty} \chi(t) \chi(t + \tau) dt = \sum_{k=1}^p \frac{\beta(\rho_k) \beta(-\rho_k)}{\alpha(\rho_k) \alpha(-\rho_k)} e^{\rho_k \tau}, \quad (\text{C16})$$

where the  $\rho_k$  are the roots of the autoregressive characteristic polynomial.

## APPENDIX D: HOW DOES KĀLĪ INFERENCE C-ARMA PARAMETERS

### D1 How to Fit a C-ARMA Process to Observations?

We wish to infer the values of the co-efficients in equation (1) that defines a particular C-ARMA process in order to compute the physically interesting Green's function and driving disturbance PSD for the observed light curve. While many techniques can be applied to perform this task, we use the Kalman filter (Kalman 1960; Simon 2006) since it requires trivial matrix inversions and determinant calculations ( $1 \times 1$  matrices in the case of single-filter observations). An added advantage of the Kalman filter is that it is designed to allow us to update the likelihood calculation when new measurements become available. Finally, the framework in which the Kalman filter operates, the state-space framework, is also easy to extend to the case of multi-band, non-simultaneous observations of AGN as will become available with next generation surveys such as PanSTARRS and the LSST.

The Kalman filter allows us to obtain Bayesian posteriors for the distribution function of the C-ARMA coefficients. These coefficients can then be used to compute physically meaningful time-scales that will yield insight into accretion physics. Inferencing linear systems using the Kalman filter has been successfully applied to control complex systems including the Apollo program for over 50 years (Grewal & Andrews 2010), and has since then come into vogue in fields as diverse as biostatistics (Jones 1993) and econometrics (Durbin & Koopman 2012; Harvey 1991).

Application of the Kalman filter requires that the system be represented in *state-space* form. The central concept underlying the application of the Kalman filter is that of the state of the system—an abstract quantity that completely characterizes the system at a given instant e.g. the position and velocity of a rocket. This state evolves via the state-equation—a differential (continuous-time) or difference (discrete-time) equation that tells us how to update the state of the system from time-step to time-step. Generically, the state-equation consists of three components—(1) a deterministic isolated time evolution component; (2) a deterministic

input response component; and (3) a random disturbance component. The input component can be used to control the system by driving it to a desired state. One of the key benefits of the Kalman approach is that the state is not observed directly—instead the system is monitored via the observation equation e.g. we observe the angular position of the rocket. The act of observation is necessarily imprecise and hence the observed quantities are contaminated with noise. The Kalman filter is used to recover estimates of the state of the system given the noisy observations of the state and the form of the state- and observation-equations.

We represent the C-ARMA process in continuous-time state-space form (Wiberg 1971; Friedland 2005) as

$$dx = Axdt + Bd\omega, \quad (D1)$$

and

$$\chi(t) = H(t)x + v, \quad (D2)$$

where  $x$  is the system-state vector,  $W$  is the Wiener process of section C,  $v \sim \mathcal{N}(0, \sigma_N^2)$  is the measurement noise, and  $\chi$  is the (noisy) observed flux at time  $t$ . The matrix  $A$  is known as the ‘state-transition matrix’, while the matrix  $H$  is known as the ‘observation matrix’. There are infinitely many ways of writing a C-ARMA process in state-space form—we shall discuss the two most relevant (Denham 1974; Dickinson et al. 1974) because of their use in common packages for C-ARMA inferencing such as KĀLÍ (this work) and CARMA\_PACK (Kelly et al. 2014). All other forms may be obtained from these ‘canonical’ forms via similarity transforms.

The ‘1<sup>st</sup> canonical form’ of Wiberg (1971) which is also the ‘2<sup>nd</sup> companion form’ of Friedland (2005), puts the dynamics of the C-ARMA process solely into the state-equation by representing equation (C10) using

$$A = \begin{pmatrix} -\alpha_1 & 1 & \dots & 0 & 0 \\ -\alpha_2 & 0 & \dots & 0 & 0 \\ \vdots & \vdots & \ddots & \vdots & \vdots \\ -\alpha_{p-1} & 0 & \dots & 0 & 1 \\ -\alpha_p & 0 & \dots & 0 & 0 \end{pmatrix}; \quad B = \begin{pmatrix} \beta_{p-1} \\ \beta_{p-2} \\ \vdots \\ \beta_1 \\ \beta_0 \end{pmatrix}; \quad H = \begin{pmatrix} 1 \\ 0 \\ \vdots \\ 0 \\ 0 \end{pmatrix}^\top. \quad (D3)$$

The ‘2<sup>nd</sup> canonical form’ or ‘phase-variable canonical form’ of Wiberg (1971) which is the ‘1<sup>st</sup> companion form’ of Friedland (2005) (see also Kelly et al. 2014 and references therein), breaks-up the dynamics of the C-ARMA process across the state- and observation-equations by representing

equation (C10) using

$$A = \begin{pmatrix} 0 & 1 & 0 & \dots & 0 \\ 0 & 0 & 1 & \dots & 0 \\ \vdots & \vdots & \vdots & \ddots & \vdots \\ 0 & 0 & 0 & \dots & 1 \\ -\alpha_p & -\alpha_{p-1} & -\alpha_{p-2} & \dots & -\alpha_1 \end{pmatrix}; \quad B = \begin{pmatrix} 0 \\ 0 \\ \vdots \\ 0 \\ 1 \end{pmatrix}; \quad H = \begin{pmatrix} \beta_0 \\ \beta_1 \\ \vdots \\ \beta_{p-2} \\ \beta_{p-1} \end{pmatrix}^\top. \quad (D4)$$

Note that the components of the state-vector  $x$  are not the same in the two representations but carry the same information, i.e., we are projecting the abstract state-vector into a particular coordinate system. Regardless of the choice of representations used, we can integrate the state-equation (Brockwell 1994; Øksendal 2014) to obtain an update equation for  $x$

$$x(t + \delta t) = x(t)e^{A\delta t} + \int_0^{\delta t} e^{A(\delta t-u)} BdW, \quad (D5)$$

where  $\int_0^{\delta t} e^{A(\delta t-u)} BdW$  is an Itô integral. The presence of the stochastic Itô integral implies that we cannot predict exactly what the future values of  $x$  will be because the integral in equation (D5) contributes randomly for every realization of the process. However, we can predict the expectation value of the integral and the  $1 - \sigma$  confidence limit on that expectation value (Davis 2002, chapter 6). These are given by

$$\left\langle \int_0^{\delta t} e^{A(\delta t-u)} BdW \right\rangle = 0, \quad (D6)$$

and

$$\left\langle \int_0^{\delta t} e^{A(\delta t-u)} BdW \int_0^{\delta t} e^{A(\delta t-u)} BdW \right\rangle = \int_0^{\delta t} e^{A\xi} BB^\top e^{A^\top \xi} d\xi, \quad (D7)$$

allowing us to rewrite equation D5 as

$$x(t + \delta t) = Fx(t) + w_k, \quad (D8)$$

with

$$F = e^{A\delta t}, \quad (D9)$$

and  $w_k \sim \mathcal{N}(0, Q)$  with

$$Q = \int_0^{\delta t} e^{A\xi} BB^\top e^{A^\top \xi} d\xi. \quad (D10)$$

$F$  is called the transition matrix and  $Q$  is the disturbance variance-covariance matrix.

Computing the matrices  $F$  and  $Q$  is readily accomplished using the eigen-decomposition of  $A$ . Let

$$A = v w v^{-1}, \quad (D11)$$

where  $v$  is the matrix whose  $i^{\text{th}}$  column is the eigenvector  $v_i$  of  $A$  and  $w$  is the diagonal matrix whose diagonal elements are the corresponding eigenvalues, i.e.,  $w_{i,i} = w_i$ . If the eigenvalues are distinct, then  $F$  may be computed as

$$F = e^{A\delta t} = v e^{w\delta t} v^{-1}, \quad (D12)$$

where we exploit the property that the matrix exponential of a diagonal matrix is simply the matrix of the exponentials of the diagonal elements, i.e.,

$$\mathbf{e}^{\mathbf{w}\delta t} = \begin{pmatrix} e^{w_1\delta t} & 0 & \dots & 0 & 0 \\ 0 & e^{w_2\delta t} & \dots & 0 & 0 \\ \vdots & \vdots & \ddots & \vdots & \vdots \\ 0 & 0 & \dots & e^{w_{p-1}\delta t} & 0 \\ 0 & 0 & \dots & 0 & e^{w_p\delta t} \end{pmatrix}. \quad (\text{D13})$$

To compute  $\mathbf{Q}$ , we rewrite equation D10 substituting in the expression for the matrix exponential of  $\mathbf{A}$  from equation D12.

$$\mathbf{Q} = \int_0^{\delta t} e^{\mathbf{A}\xi} \mathbf{B} \mathbf{B}^\top e^{\mathbf{A}^\top \xi} d\xi = \mathbf{v} \int_0^{\delta t} e^{\mathbf{w}\xi} \mathbf{v}^{-1} \mathbf{B} \mathbf{B}^\top (\mathbf{v}^{-1})^\top e^{\mathbf{w}\xi} d\xi \mathbf{v}^\top, \quad (\text{D14})$$

where we have factored the matrices  $\mathbf{v}$  and  $\mathbf{v}^\top$  out of the integral. Defining  $\mathbf{C} = \mathbf{v}^{-1} \mathbf{B} \mathbf{B}^\top (\mathbf{v}^{-1})^\top$  lets us rewrite the integral as

$$\mathbf{Q} = \mathbf{v} \int_0^{\delta t} e^{\mathbf{w}\xi} \mathbf{C} e^{\mathbf{w}\xi} d\xi \mathbf{v}^\top. \quad (\text{D15})$$

Let the integral be  $\mathbf{D}$  with  $i^{\text{th}}$  &  $j^{\text{th}}$  entries

$$\mathbf{D}_{i,j} = \int_0^{\delta t} e^{(w_i+w_j)\xi} \mathbf{C}_{i,j} d\xi = \frac{e^{(w_i+w_j)\delta t} - 1}{w_i + w_j} \mathbf{C}_{i,j}. \quad (\text{D16})$$

Then we calculate

$$\mathbf{Q} = \mathbf{v} \mathbf{D} \mathbf{v}^\top, \quad (\text{D17})$$

We also define

$$\mathbf{J}_{i,j} = \int_0^{\infty} e^{(w_i+w_j)\xi} \mathbf{C}_{i,j} d\xi = \frac{-1}{w_i + w_j} \mathbf{C}_{i,j}. \quad (\text{D18})$$

and use it to compute

$$\Sigma = \mathbf{v} \mathbf{J} \mathbf{v}^\top, \quad (\text{D19})$$

which can be used to compute the ACVF of the process via (Brockwell 1994)

$$\gamma_{XX}(T) = \mathbf{H} \mathbf{v} e^{\mathbf{w}T} \mathbf{v}^{-1} \Sigma \mathbf{H}^\top. \quad (\text{D20})$$

These equations for the computation of  $\mathbf{F}$ ,  $\mathbf{Q}$ , and  $\Sigma$  allow us to use the Kalman filter to perform C-ARMA simulations & perform inferencing.

## D2 Equivalent Representations of C-ARMA Processes

We have introduced C-ARMA processes using equation (1). A particular C-ARMA process can be specified via the coefficients of the LHS and RHS of equation (1). That is to say that if  $\Theta$  is the set of C-ARMA process parameters, then  $\Theta$  can be written

$$\Theta = \{\alpha_1, \alpha_2, \dots, \alpha_{p-1}, \alpha_p, \beta_0, \beta_1, \dots, \beta_{q-1}, \beta_q\}. \quad (\text{D21})$$

A second, equivalent, representation can be obtained by factorizing the C-AR and C-MA polynomials of equation (A16) and equation (C4). However, the C-MA polynomial is scaled

by the co-efficient of the highest-order term  $\beta_q$ , suggesting that we define the *modified* C-MA polynomial to be

$$\beta'(z) = \frac{\beta_0}{\beta_q} + \frac{\beta_1}{\beta_q} z + \dots + \frac{\beta_{q-1}}{\beta_q} z^{q-1} + z^q, \quad (\text{D22})$$

with roots  $\mu_k$ . The corresponding component of the state-space representation is

$$\mathbf{B}' = \begin{pmatrix} 0 & \dots & 0 & 1 & \frac{\beta_{q-1}}{\beta_q} & \dots & \frac{\beta_1}{\beta_q} & \frac{\beta_0}{\beta_q} \end{pmatrix}^\top. \quad (\text{D23})$$

Define  $\mathbf{C}' = \mathbf{v}^{-1} \mathbf{B}' \mathbf{B}'^\top (\mathbf{v}^{-1})^\top$  and  $\mathbf{J}'_{i,j} = \frac{-1}{w_i + w_j} \mathbf{C}'_{i,j}$  to compute  $\Sigma' = \mathbf{v} \mathbf{J}' \mathbf{v}^\top$ . Clearly,  $\mathbf{C} = \beta_q^2 \mathbf{C}'$ , leading to the useful expression

$$\Sigma = \beta_q^2 \Sigma', \quad (\text{D24})$$

where  $\Sigma[0,0]$  is the true variance of the C-ARMA process and  $\Sigma'[0,0]$  is the variance of the C-ARMA process with unit  $\beta_q$ .  $\Sigma'[0,0]$  can be computed from the C-AR roots  $\rho_k$  and the C-MA roots  $\mu_k$  with no knowledge of  $\beta_q$  required. If we specify  $\Sigma[0,0]$ , we can recover  $\beta_q$  via equation (D24). Therefore, an equivalent representation,  $\mathbf{P}$  of the C-ARMA process is

$$\mathbf{P} = \{\rho_1, \rho_2, \dots, \rho_{p-1}, \rho_p, \mu_1, \dots, \mu_{q-1}, \mu_q, \Sigma[0,0]\}. \quad (\text{D25})$$

## D3 Sampling Patterns

Observation cadences can be broadly classified into two groups - those with regular sampling (RS) and those with irregular sampling (IRS). RS patterns occur when observations are made at fixed intervals to within the precision of the clock governing the schedule. We can relate the number of observations  $N$  made to the duration of the sampling  $T$  in terms of the sampling interval  $\delta t = T/N$ . The times at which the process is sampled  $t$  must then be given by  $t = n\delta t$  for  $1 \leq n \leq N$  and can therefore be indexed by  $n$ . Occasionally, natural or mission-induced events make it impossible to observe the object of interest for one or more simultaneous cadences. Such incidents are called ‘missing-observations’. We track missing-observations via a mask  $\mathcal{M}_n$ . If the  $n^{\text{th}}$  observation is missing,  $\mathcal{M}_n = 0$ . Otherwise,  $\mathcal{M}_n = 1$ . Suppose the light curve under consideration is regularly observed (e.g. with a space-based telescope such as Kepler) with missing-observations, then we shall write the integrated state-equation (D8) as

$$\mathbf{x}_{k+1} = \mathbf{F} \mathbf{x}_k + \mathbf{w}_k, \quad (\text{D26})$$

where  $\mathbf{w}_k \sim \mathcal{N}(\mathbf{0}, \mathbf{Q})$ . The original observation equation (D2) remains unchanged through this process but is now discrete

$$\chi_n = \mathbf{H}_n \mathbf{x}_n + v_n, \quad (\text{D27})$$

where  $v_n \sim \mathcal{N}(0, \sigma_{N,n}^2)$  are (heteroskedastic) measurement errors. We recognize equations (D26) and (D27) as a discrete-time state-space system and identify it with the *Auto-Regressive Moving Average* (ARMA) processes of Brockwell & Davis (2009), Brockwell & Davis (2010), Durbin & Koopman (2012), and in astronomical variability studies of Scargle (1981); Gaskell & Peterson (1987); Koratkar

& Gaskell (1991). If the light curve has missing observations, we recommend that the 1<sup>st</sup> canonical form from equation (D3) be chosen. The benefit of this form of the state-space equations is that the observation matrix is very simple. To correctly handle missing-observations, it suffices to set  $\mathbf{H}_n = \begin{pmatrix} \mathcal{M}_n & 0 & \dots & 0 & 0 \end{pmatrix}$  and the measurement uncertainty of the  $n^{\text{th}}$  observation  $\sigma_{N,n}^2 = \infty$  (Jones 1980; Jones 1993).

If the light curve is obtained with an IRS pattern as is usually the case with ground-based astronomical observations such as those carried out by the Sloan Digital Sky Survey (SDSS), we define  $\delta t_n = t_{n+1} - t_n$ . If the IRS pattern is fairly dense, it may be advantageous to simply determine an artificial fixed sampling interval  $\delta t$  as the greatest common divisor of the individual  $\delta t_n$ . Then we add missing-observations into the light curve so that the missing-observations correspond to the gaps in the original light curve. At this point, we may treat the formerly IRS light curve as a RS light curve with lots of missing observations.

If the IRS pattern is fairly sparse, it is computationally cheaper to recompute the transition matrix  $\mathbf{F}_n$  and the disturbance variance-covariance matrix  $\mathbf{Q}_n$  at every time step as proposed by Jones & Ackerson (1990), Jones (1993) and Kelly et al. (2014). The exact IRS pattern for which the two methods are equally fast depends on the machine architecture and hence our implementation of the C-ARMA modelling algorithm in KĀLĪ allows for both methods to be used at different time-steps within the same light curve for optimal speed, i.e., long gaps or sections can be treated as unevenly sampled data while occasional gaps are treated as missing values.

#### D4 Parameter Estimation with the Kalman Filter

Having obtained the discrete-time state- and observation-equations for the original C-ARMA process, we may now use the Kalman filter to compute the likelihood of having observed the light curve given the C-ARMA model. This likelihood can then be used to determine the likelihood of the C-ARMA model given the data using Bayes' rule. The future goal will then become to determine a suitable informative prior using a combination of previous observations and physical insight.

As stated earlier, one of the many attractive properties of the Kalman filter is that it operates sequentially on the light curve measurements—the Kalman filter uses the state-equation to *predict* the current state of the system based on the previous state before the observation is made available to the algorithm, i.e., it computes an estimate of the *a priori* state of the system  $\hat{\mathbf{x}}_n^-$ , where the ‘ $-$ ’ signifies that it is an *a priori* estimate. Once the  $n^{\text{th}}$ -observation is made available to the algorithm, the Kalman filter uses the observation to *correct* the *a priori* state of the system and thereby estimates the *a posteriori* state of the system  $\hat{\mathbf{x}}_n^+$ , where the ‘ $+$ ’ signifies that it is an *a posteriori* estimate. We can use the estimated *a priori* states of the system  $\hat{\mathbf{x}}_n^-$  to predict estimates of the observed flux  $\hat{y}_n = \mathbf{H}\hat{\mathbf{x}}_n^-$ . The difference between the *a priori* predictions of the observed flux and the actual values of the observed flux,  $n_n$ , are known as ‘innovations’

$$r_n = y_n - \hat{y}_n, \quad (\text{D28})$$

and may be shown to be independent and identically distributed normal deviates (Simon 2006, chapter 10.2) that are distributed according to  $\mathcal{N}(0, S_n)$  where  $S_n$  is the estimate of the variance of the  $n^{\text{th}}$ -innovation.

Suppose we wish to calculate the log-likelihood of the C-ARMA( $p,q$ ) model  $C$  for the set of model parameters  $\alpha_k$  and  $\beta_k$  given the observed light curve of an object,  $\{y_n\}$ . From Bayes' theorem,

$$\ln \mathcal{L}(C|\{y_n\}) = \ln \mathcal{L}(\{y_n\}|C) + \ln \mathcal{L}(C). \quad (\text{D29})$$

The log-likelihood of the observed light curve  $\{y_n\}$  given the model  $C$  may be computed using the innovations via

$$\ln \mathcal{L}(\{y_n\}|C) = \sum_{n=1}^N -\frac{\mathcal{M}_n}{2} (\ln 2\pi + \ln S_n + \frac{r_n^2}{S_n}), \quad (\text{D30})$$

where we include the value of the mask  $\mathcal{M}_n$  to ensure that missing-observations do not contribute to the log-likelihood. For the prior, i.e., the likelihood of the model itself, we use the piecewise function

$$\ln \mathcal{L}(C) = \begin{cases} 0 & \text{if } \operatorname{Re}(\rho_k) < 0 \text{ and } \operatorname{Re}(\mu_k) < 0 \text{ for all } k \\ -\infty & \text{otherwise} \end{cases}, \quad (\text{D31})$$

where  $\rho_k$  and  $\mu_k$  are the roots of the autoregressive and moving average polynomials. This prior, while uninformative for acceptable values of the roots, ensures that the C-ARMA process corresponding to the roots is both stable and invertible.

The innovations required for the calculation of the light curve likelihood in equation (D30) may be computed by applying the Kalman filter to the observed flux measurements. To do so, we require an estimate of the initial state of the system  $\hat{\mathbf{x}}_0^+$  as well as the uncertainty in the initial state of the system  $\mathbf{P}_0^+$ . The best estimate of the state of the system when initializing the Kalman filter is

$$\hat{\mathbf{x}}_0^+ = \mathbf{0} \quad (\text{D32})$$

and the best estimate of the state uncertainty is

$$\mathbf{P}_0^+ = \Sigma, \quad (\text{D33})$$

where  $\Sigma$  is computed using equations (D18) and (D19).

Then starting with the first observation, i.e.,  $n = 1$ , we iterate through the  $N$  observations using the Kalman filter. We begin by computing the expected value of the *a priori* state:

$$\hat{\mathbf{x}}_n^- = \mathbf{F}\hat{\mathbf{x}}_{n-1}^+, \quad (\text{D34})$$

and the expected value of the *a priori* state uncertainty:

$$\mathbf{P}_n^- = \mathbf{F}\mathbf{P}_{n-1}^+\mathbf{F}^\top + \mathbf{Q}. \quad (\text{D35})$$

At this point, the innovation for this observation and the uncertainty in the innovation can be computed by comparing the expected observation to the actual observation using

$$r_n = y_n - \mathbf{H}\hat{\mathbf{x}}_n^-, \quad (\text{D36})$$

and

$$S_n = (\mathbf{H}_n \mathbf{P}_n^- \mathbf{H}_n^\top + \sigma_{N,n}^2), \quad (\text{D37})$$

where  $\sigma_{N,n}^2$  is the (heteroskedastic) variance of the observation noise.

Before we may compute the innovation for the next observation, we must use our knowledge of the observed value of the AGN flux to update our knowledge of the state and the uncertainty in the state. Since the measurements are noisy, we must compute a gain to apply to the measurement. The gain should reflect the noisiness of the measurements, i.e., if the  $n^{\text{th}}$ -observation is very noisy, we would like the gain that we apply to be very small. More significance is given to the observation as compared to the a priori state estimate if the gain is large. The optimal gain is the Kalman gain given by

$$\mathbf{K}_n = \mathbf{P}_n^- \mathbf{H}_n^\top S_n^{-1}. \quad (\text{D38})$$

While any gain value will yield an unbiased estimate of the state, the Kalman gain yields an unbiased *minimum variance* a posteriori estimate of the state and state uncertainty (Simon 2006, chapter 3.3). The Kalman gain is used to compute the a posteriori state and state covariance via

$$\hat{\mathbf{x}}_n^+ = \hat{\mathbf{x}}_n^- + \mathbf{K}_n r_n, \quad (\text{D39})$$

and

$$\mathbf{P}_n^+ = (\mathbf{I} - \mathbf{K}_n \mathbf{H}_n) \mathbf{P}_n^- (\mathbf{I} - \mathbf{K}_n \mathbf{H}_n)^\top + \mathbf{K}_n \sigma_n^2 \mathbf{K}_n^\top. \quad (\text{D40})$$

The innovations computed using this procedure may then be used along with equations (D30) and (D31) to compute the log-likelihood of the C-ARMA model  $C$  given the observed light curve  $\{y_n\}$  using equation (D29). A multitude of numerical algorithms, including the Expectation-Maximization algorithm, can be used to maximize this likelihood. Alternatively, the full probability distribution of the C-ARMA model parameters can be probed by applying Markov Chain Monte Carlo (MCMC) techniques to the likelihood function.

We have discussed how the Kalman filter equations can be used to iterate through observations to compute the likelihood of having observed the light curve given a particular C-ARMA model. Bayes' rule allows us to use this likelihood to pick the most likely C-ARMA model. For now, we use a prior to select the portion of model space allowed by the rules of the C-ARMA process (the roots of the AR and MA polynomials must have negative real part). In the future, more informative priors can be used based on existing data, simulations and accretion theory.

## D5 How to Pick a Good C-ARMA Process?: Model Selection

Although we have reasons to expect AGN light curves to be well modelled by simple C-ARMA models of low order, it is prudent to search a large portion of the space of allowable models and use a statistically sound information criteria to pick the optimal model order. Information criteria such as the corrected Akaike Information Criteria (AICc) and the Deviance Information Criteria (DIC) balance model likelihood against model simplicity. The model that minimizes a chosen information criteria is the most-likely maximally-parsimonious model that fits the observed light curve (Burnham & Anderson 2003).

If a non-linear optimization algorithm is used to maximize the likelihood in equation (D29), we have point estimates of the model parameters and a single value for the likelihood of each model. To select the model order, we may

use the AICc

$$AICc = 2(p + q + 1) - 2 \ln \mathcal{L}(C|\{y_n\}) + \frac{2(p + q + 1)(p + q + 2)}{N - p - q - 2}, \quad (\text{D41})$$

where  $p$  and  $q$  are the orders of the C-ARMA( $p,q$ ) model and  $N$  is the number of observations in the light curve.

Markov Chain Monte Carlo (MCMC) can be used to sample the likelihood function of equation (D29). Sampling the likelihood function yields reliable estimates of the confidence intervals of the C-ARMA model parameters even if the likelihood function is highly non-Gaussian. If multiple samples from the likelihood function are available, it is possible to use the DIC to select model order. The DIC is given by

$$DIC = \mu_{D(C)} + \frac{1}{2} \sigma_{D(C)}^2, \quad (\text{D42})$$

where the deviance  $D(C)$  is computed as

$$D(C) = -2 \ln \mathcal{L}(C|\{y_n\}). \quad (\text{D43})$$

Either one of the DIC or the AICc may be used to select the optimum model order.

We have shown that regardless of the form of the equation governing the flux emitted by AGN and without a detailed analytic prescription for the form of accretion disk instabilities, it is possible to model small stochastic fluctuations in the observed light curve of AGN as a C-ARMA process. The physics of the accretion process may be probed by examining the Green's function and the Disturbance PSD of the inferred C-ARMA process. We have presented a method for inferring the model parameters of the C-ARMA process from observations of the light curve using the Kalman filter that is particularly suitable for light curves with regular sampling.

## D6 What was the Light Curve Actually Doing?: Light Curve Smoothing

After having selected a model and having determined the posterior probability distribution of the model parameters (or if using a point estimator, the maximum likelihood parameter estimates), the question arises: Can we estimate the true value of the light curve at any given instant regardless of the availability of an observation at the desired instant? This problem is known as the fixed-interval smoothing problem in the stochastic control and time series literature (Simon 2006; Durbin & Koopman 2012). Optimal smoothing can be performed with the the Rauch-Tung-Striebel (RTS) smoother (Rauch, Striebel & Tung 1965). The RTS smoother is based on the Kalman filter but requires additional steps after the standard Kalman filter has been applied. For astrophysical applications, smoothing can be important when comparing light curves from the same object obtained non-simultaneously in different filter bands. Smoothed light curves can be used to forecast the light curves of objects for short durations into the future. Such forecasts may be useful for survey optimization.

Assuming that the light curve has been regularly sampled (with missing observations), RTS smoothing is performed by first running the standard Kalman filter of section D4 in the forward direction (increasing time). In the application of the Kalman filter for computing the likelihood

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of the model, the a posteriori estimates of the state ( $\hat{x}_n^+$ ) and state covariance ( $\mathbf{P}_n^+$ ) for the  $n^{\text{th}}$ -observation are not required once the a priori estimates of the state ( $\hat{x}_{n+1}^-$ ) and state-covariance ( $\mathbf{P}_{n+1}^-$ ) of the next state have been computed. To perform RTS smoothing, all the a priori and a posteriori state and state-covariance estimates must be retained for later use during the smoothing phase of the algorithm, i.e., to provide the best estimate of the state and state uncertainty, the algorithm uses *all* the observations simultaneously instead of just the previous observations. After the Kalman filter has been used to estimate the state and state-covariance for all the observations, we have the a priori and a posteriori estimates of the state and state covariance for all the (regularly-spaced) observations, i.e.,  $\{\hat{x}_n^-\}$ ,  $\{\hat{x}_n^+\}$ ,  $\{\mathbf{P}_n^-\}$ , and  $\{\mathbf{P}_n^+\}$  for  $0 \leq n \leq N$ .

The RTS smoothed estimates ( $\hat{x}_n$  and  $\mathbf{P}_n$ ) can be obtained by initializing the RTS smoother with

$$\hat{x}_N = \hat{x}_N^+, \quad (\text{D44})$$

and

$$\mathbf{P}_N = \mathbf{P}_N^+. \quad (\text{D45})$$

Then iterating through the observations backwards (i.e., decreasing time), we first compute the Kalman gain

$$\mathbf{K}_n = \mathbf{P}_n^+ \mathbf{F}^\top (\mathbf{P}_{n+1}^-)^{-1}, \quad (\text{D46})$$

after which we can compute the smoothed state and state covariance using

$$\hat{x}_n = \hat{x}_n^+ + \mathbf{K}_n (\hat{x}_{n+1}^- - \hat{x}_{n+1}^+), \quad (\text{D47})$$

and

$$\mathbf{P}_n = \mathbf{P}_n^+ + \mathbf{K}_n (\mathbf{P}_{n+1}^- - \mathbf{P}_{n+1}^+) \mathbf{K}_n^\top. \quad (\text{D48})$$

The smoothed state vector ( $\hat{x}_n$ ) can then be used to calculate the best estimate of the underlying light curve using the noiseless-observation equation

$$\chi_n = \mathbf{H} \hat{x}_n, \quad (\text{D49})$$

where we have used a fixed observation matrix that *never* produces missing observations rather than the mask-dependent observation matrix  $\mathbf{H}_n$ .

In practice, RTS smoothing is only performed after the best C-ARMA model and model parameters have been determined, i.e., it is not performed as part of the estimation but instead is a product after the estimation has been completed.

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