

# Collision strengths and transition probabilities for Co III forbidden lines

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## ABSTRACT

In this paper we compute the collision strengths and their thermally averaged Maxwellian values for electron transitions between the 15 lowest levels of doubly ionized cobalt,  $\text{Co}^{2+}$ , which give rise to forbidden emission lines in the visible and infrared region of spectrum. The calculations also include transition probabilities and predicted relative line emissivities. The data are particularly useful for analysing the thermodynamic conditions of supernova ejecta.

**Key words:** atomic data – atomic processes – radiation mechanisms: non-thermal – supernovae: general – infrared: general.

## 1 INTRODUCTION

Cobalt is an iron group element but is the least abundant of this group with a solar abundance of about 300 times less than Fe. However, in supernova (SN) ejecta it is much more abundant. For example, in SN 1987A the ratio of Co to Fe, 255 days after outburst, is approximately 0.2 by number (Varani et al. 1990). The spectral lines of Co are therefore valuable investigative tools in analysing the chemical and thermodynamic conditions of SNe where these emissions are mostly found. These lines are also useful in investigating the evolutionary history and chemical development by nucleosynthesis and decay processes within the SN explosions (Colgate & McKee 1969; Axelrod 1980; Kuchner et al. 1994; Bowers et al. 1997; Liu et al. 1997; Churazov et al. 2014; Childress et al. 2015). The lines of cobalt have also been observed in the spectral emissions of astronomical objects with more normal Co abundances such as planetary nebulae (Baluteau et al. 1995; Pottasch & Surendiranath 2005; Zhang et al. 2005; Wang & Liu 2007; Fang & Liu 2011).

Little computational and experimental work has been done previously to generate essential atomic data for Co III and none of the previous work deals with excitation of  $\text{Co}^{2+}$  levels by electron impact. Hansen, Raassen & Uylings (1984) calculated magnetic dipole and electric quadrupole transition probabilities in the  $3d^7$  ground configuration of Co III using parametric fitting to the observed energy levels and Hartree–Fock values for the electric quadrupole moments. In their investigation of the forbidden transition probabilities relevant to the analysis of infrared lines from SN 1987A, Nussbaumer & Storey (1988) provided a few transition probabilities for low levels of Co III assuming *LS* coupling. Tankosić, Popović & Dimitrijević (2003) calculated Stark broadening data for a number of Co III spectral lines as a function of temperature by using a semi-empirical approach. Experimental investigations have also been conducted by Sugar & Corliss (1981, 1985) where atomic data related to Co III transitions, mainly energy levels of  $\text{Co}^{2+}$ , have been

collected. Very recently, Fivet, Quinet & Bautista (2016) calculated radiative probabilities of Co III forbidden transitions between low-lying levels of doubly ionized cobalt as part of a larger investigation of the radiative rates in doubly ionized iron-peak elements.

We have recently reported a calculation of atomic parameters for energetically low-lying levels of  $\text{Co}^+$  (Storey, Zeippen & Sochi 2016). In this paper we present a similar calculation of atomic parameters related to forbidden transitions in  $\text{Co}^{2+}$ , which includes lines ranging from the visible to the three mid-infrared lines which arise from transitions within the ground term at 11.88, 16.39 and 24.06  $\mu\text{m}$ . The paper primarily addresses a shortage in collisional atomic data which forced some researchers (Dessart et al. 2014; Childress et al. 2015) to adopt collision strengths generated for Ni IV (Sunderland et al. 2002) as a substitute for corresponding data of Co III justifying this by the fact that the two ions possess similar electronic and term structures. Our principal result is collision strengths and their thermally averaged Maxwellian values for electron excitation and de-excitation between the 15 lowest levels of  $\text{Co}^{2+}$ . The study also includes the most important radiative transition probabilities for the same levels. The main tools used in generating these data are the R-matrix atomic scattering code (Berrington et al. 1974, 1987; Hummer et al. 1993; Berrington, Eissner & Norrington 1995)<sup>1</sup> and the general purpose AUTOSTRUCTURE code (Eissner, Jones & Nussbaumer 1974; Nussbaumer & Storey 1978; Badnell 2011).<sup>2</sup> The scattering calculations were performed using a 10-configuration atomic target within a Breit–Pauli intermediate coupling approximation, as will be detailed in Section 2.

The paper is structured as follow. In Section 2 the  $\text{Co}^{2+}$  model is presented and the resulting transition probabilities are given, whereas in Section 3 the Breit–Pauli R-matrix  $\text{Co}^{2+} + e$  scattering calculation is described. Results and general analysis related to the

<sup>1</sup> See Badnell: R-matrix write-up on [http://amdpp.phys.strath.ac.uk/UK\\_RmaX/codes/](http://amdpp.phys.strath.ac.uk/UK_RmaX/codes/)

<sup>2</sup> See Badnell: AUTOSTRUCTURE write-up on <http://amdpp.phys.strath.ac.uk/autos/>

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diagnostic potentials of some transitions appear in Section 4, and Section 5 concludes the paper.

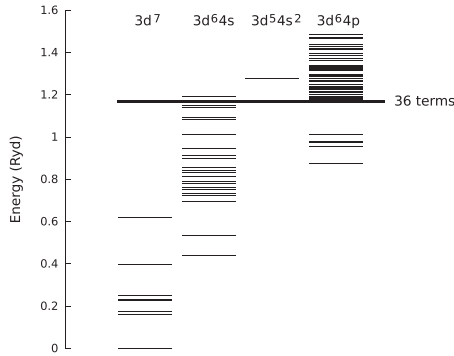
## 2 CO<sup>2+</sup> ATOMIC STRUCTURE

### 2.1 The scattering target

A schematic diagram of the term structure of Co III up to 1.5 Ryd is shown in Fig. 1. The extent of our target is shown by the heavy solid line in that figure and includes 36 terms and 109 levels. The lowest 21 terms of this ion are of even parity from the configurations 3d<sup>7</sup> and 3d<sup>6</sup>4s. Transitions from higher terms give rise to lines that should be weaker at the typical temperatures of SN ejecta and hence they will be ignored. The odd-parity terms of the 3d<sup>6</sup>4p configuration are expected to give rise to resonances that affect the collision strengths for excitation of the low-lying even-parity levels and hence they are included in the target for the scattering calculations.

A set of 10 electron configurations, listed in Table 1, was used to expand the target states. The target wavefunctions were generated with the *AUTOSTRUCTURE* program (Eissner et al. 1974; Nussbaumer & Storey 1978; Badnell 2011), using radial functions computed within scaled Thomas-Fermi-Dirac statistical model potentials. The scaling parameters were determined by minimizing the sum of the energies of all the target terms, computed in *LS* coupling, i.e. by neglecting all relativistic effects. The resulting scaling parameters,  $\lambda_{nl}$ , are given in Table 2.

In Table 3 a comparison is made between the term energies calculated using our scattering target with experimental values for the 36 terms of the target. The term energies are computed with the in-



**Figure 1.** Schematic term energy diagram of Co III. The heavy solid line shows the extent of the close-coupled target states.

**Table 1.** The 10 target configuration basis where the core electronic structure ([Ar]) is suppressed. The bar indicates a correlation orbital.

3d <sup>7</sup>
3d <sup>6</sup> 4s, 4p, $\bar{4}d$
3d <sup>5</sup> 4s <sup>2</sup> , 4p <sup>2</sup> , $\bar{4}d^2$ , 4s4p, 4s $\bar{4}d$ , 4p $\bar{4}d$

**Table 2.** Potential scaling parameters. The bar over the principal quantum number signifies a correlation orbital.

1s	1.42912				
2s	1.13799	2p	1.08143		
3s	1.06915	3p	1.05203	3d	1.04962
4s	1.03440	4p	1.02977	$\bar{4}d$	1.51187

**Table 3.** Energies of the 36 target terms in cm<sup>-1</sup>, ordered according to the experimental energy. The calculated values include only the spin-orbit contribution to the fine-structure energies. Core electronic structure ([Ar]) is suppressed from all configurations.

Config.	Term	Term energy	
		Exp. <sup>a</sup>	Calc.
3d <sup>7</sup>	a <sup>4</sup> F	0	0
	a <sup>4</sup> P	14 561	17 891
	a <sup>2</sup> G	16 510	19 120
	a <sup>2</sup> P	19 618	25 103
	a <sup>2</sup> H	22 227	25 205
	a <sup>2</sup> D	22 712	27 507
	a <sup>2</sup> F	36 372	43 416
	a <sup>6</sup> D	46 230	48 501
	a <sup>4</sup> D	55 448	58 817
	b <sup>4</sup> P	70 965	79 599
3d <sup>6</sup> 4s	a <sup>4</sup> H	71 096	76 483
	b <sup>4</sup> F	72 717	80 163
	a <sup>4</sup> G	76 219	83 370
	b <sup>2</sup> P	76 521	85 780
	b <sup>2</sup> H	76 690	82 428
	b <sup>2</sup> F	78 323	86 408
	b <sup>2</sup> G	81 793	89 400
	b <sup>4</sup> D	83 031	92 162
	a <sup>2</sup> I	84 676	91 484
	c <sup>2</sup> G	85 485	93 867
3d <sup>6</sup> 4p	b <sup>2</sup> D	90 897	98 436
	z <sup>6</sup> D <sup>o</sup>	97 807	97 268
	z <sup>6</sup> F <sup>o</sup>	102 620	102 460
	z <sup>6</sup> P <sup>o</sup>	104 861	104 906
	z <sup>4</sup> D <sup>o</sup>	106 074	106 802
	z <sup>4</sup> F <sup>o</sup>	106 676	107 272
	z <sup>4</sup> P <sup>o</sup>	109 902	111 225
	z <sup>2</sup> F		111 250
	z <sup>2</sup> P		119 049
	z <sup>4</sup> S <sup>o</sup>	122 305	125 226
3d <sup>6</sup> 4s	z <sup>4</sup> G <sup>o</sup>	124 219	125 937
	z <sup>2</sup> F		129 103
	z <sup>2</sup> P		129 937
	z <sup>4</sup> S <sup>o</sup>		129 103
	z <sup>4</sup> G <sup>o</sup>		127 494
	z <sup>2</sup> F		125 226
	z <sup>2</sup> P		125 937
	z <sup>4</sup> S <sup>o</sup>		129 103
	z <sup>4</sup> G <sup>o</sup>		127 494
	z <sup>2</sup> F		125 226
	z <sup>2</sup> P		125 937

<sup>a</sup>Experimental energies are from NIST ([www.nist.gov](http://www.nist.gov)).

clusion of one-body relativistic effects, the Darwin and mass terms, and the spin-orbit interaction. This is the type of approximation that we applied for the scattering calculations in the R-matrix code. In Table 4 the calculated energies of the 15 lowest levels are compared with the corresponding experimental values. The table also shows the values obtained by including the two-body fine structure interactions as described by Eissner et al. (1974). The calculated fine-structure splittings of these levels are improved by this inclusion. For the total fine-structure splitting of the six terms, the average absolute difference from experiment drops from 7.3 to 4.6 per cent.

A widely accepted measure for the quality of the scattering calculations is the degree of agreement between weighted oscillator strengths, *gf*, calculated in the velocity and length formulations, where good agreement is regarded as necessary but not sufficient condition for the quality of the target wavefunctions. Table 5 provides this comparison where it shows an average difference in the absolute values of *gf* of about 5.8 per cent between the two formulations, which in our view is acceptable for an open d-shell atomic system.

**Table 4.** Energies in  $\text{cm}^{-1}$  of the 15 lowest levels of  $\text{Co}^{2+}$ , ordered according to the experimental energy, where the configuration of all levels is  $[\text{Ar}] 3d^7$ .

Index	Level	Exp. <sup>1</sup>	Calc. <sup>2</sup>	Calc. <sup>3</sup>
1	$a^4F_{9/2}$	0.	0.	0.
2	$a^4F_{7/2}$	841	810	824
3	$a^4F_{5/2}$	1451	1408	1428
4	$a^4F_{3/2}$	1867	1819	1842
5	$a^4P_{5/2}$	15 202	18 481	18 502
6	$a^4P_{3/2}$	15 428	18 770	18 785
7	$a^4P_{1/2}$	15 811	19 125	19 118
8	$a^2G_{9/2}$	16 978	19 565	19 581
9	$a^2G_{7/2}$	17 766	20 348	20 357
10	$a^2P_{3/2}$	20 195	25 601	25 633
11	$a^2P_{1/2}$	20 919	26 486	26 474
12	$a^2H_{11/2}$	22 720	25 690	25 687
13	$a^2D_{5/2}$	23 059	27 795	27 804
14	$a^2H_{9/2}$	23 434	26 367	26 379
15	$a^2D_{3/2}$	24 237	29 058	29 033

<sup>1</sup>Sugar & Corliss (1985).<sup>2</sup>Calculated with only spin-orbit interaction.<sup>3</sup>As 2 plus two-body fine-structure interactions for the first four configurations of Table 1.

## 2.2 Transition probabilities

The forbidden transition probabilities between the even parity low-lying terms are calculated using the aforescribed target wavefunctions, with empirical adjustments to the computed energies to ensure more reliable calculation of the fine-structure interactions and accurate energy factors connecting the ab initio calculated line

**Table 5.** Weighted  $LS$  oscillator strengths,  $gf$ , in the length and velocity formulations from the two energetically lowest terms of the  $3d^7$  and  $3d^64s$  configurations.

Transition				$gf_L$	$gf_V$	
3d <sup>7</sup>	<sup>4</sup> F	—	3d <sup>6</sup> 4p	<sup>4</sup> D <sup>o</sup>	2.34	2.48
		—		<sup>4</sup> F <sup>o</sup>	1.16	1.21
		—		<sup>3</sup> G <sup>o</sup>	2.38	2.30
3d <sup>6</sup> 4s	<sup>6</sup> D	—	3d <sup>6</sup> 4p	<sup>6</sup> D <sup>o</sup>	9.45	9.75
		—		<sup>6</sup> F <sup>o</sup>	13.7	13.5
		—		<sup>6</sup> P <sup>o</sup>	5.77	4.83

strengths to the transition probabilities. The results for the lowest 15 levels are given in Table 6 where the values represent the sum of the electric quadrupole and magnetic dipole contributions for each transition. This table includes only those probabilities from a given upper level which exceed 1 per cent of the total probability from that level.

The infrared lines of principal interest here arise from transitions between the levels of the ground  $^4F$  term and are predominantly of magnetic dipole type. There is therefore a stepwise decay through the levels and only three relevant transition probabilities, for  $a^4F_{3/2} - a^4F_{5/2}$ ,  $a^4F_{5/2} - a^4F_{7/2}$  and  $a^4F_{7/2} - a^4F_{9/2}$ . We are aware of only two previous calculations of transition probabilities for  $\text{Co III}$ , one by Hansen et al. (1984) and one by Nussbaumer & Storey (1988), as well as one contemporary calculation by Fivet et al. (2016). Nussbaumer & Storey (1988) only give values for these three probabilities and these differ by less than 1 per cent from our values. Hansen et al. (1984) give more extensive results which we compare with the present values in Table 6. We find excellent agreement with Hansen et al. (1984) for the magnetic dipole transitions between the

**Table 6.** Transition probabilities in  $\text{s}^{-1}$  among the energetically lowest 15 levels of  $\text{Co}^{2+}$  as obtained from the current work (CW), from Hansen et al. (1984) (HRU), from Fivet et al. (2016) using HFR (FQB1) and from Fivet et al. (2016) using AUTOSTRUCTURE (FQB2). The transition indices  $i$  and  $j$ , which refer to the lower and upper levels respectively, are as in Table 4. Only the CW transition probabilities that are at least 1 per cent of the total probability from a given upper level are listed. The powers of 10 by which the numbers are to be multiplied are given in brackets.

Transition		A-value				Transition		A-value			
$j$	$i$	CW	HRU	FQB1	FQB2	$j$	$i$	CW	HRU	FQB1	FQB2
2	1	2.00(–2)	2.0(–2)	2.01(–2)	2.00(–2)	11	3	2.23(–3)	2.2(–3)		
3	2	1.31(–2)	1.3(–2)	1.32(–2)	1.31(–2)	11	4	2.69(–3)	2.4(–3)		
4	3	4.63(–3)	4.7(–3)	4.65(–3)	4.63(–3)	11	7	1.77(–1)	2.0(–1)	1.98(–1)	2.01(–1)
5	1	5.55(–2)	4.8(–2)	6.59(–2)	6.65(–2)	11	10	6.42(–3)	6.4(–3)		
5	2	1.51(–2)	1.35(–2)	1.74(–2)	1.78(–2)	12	1	6.02(–4)	6.2(–4)		
5	3	3.14(–3)	2.68(–3)			12	8	3.94(–2)	4.2(–2)	4.29(–2)	4.69(–2)
6	2	3.14(–2)	2.7(–2)	3.69(–2)	3.73(–2)	13	2	7.34(–1)	7.5(–1)	7.44(–1)	8.27(–1)
6	3	1.85(–2)	1.63(–2)	2.18(–2)	2.21(–2)	13	3	7.94(–2)	8.1(–2)		
6	4	5.14(–3)	4.41(–3)			13	4	3.65(–2)	3.5(–2)		
7	3	2.30(–2)	2.0(–2)	2.71(–2)	2.73(–2)	13	5	4.74(–2)	4.7(–2)		
7	4	3.02(–2)	2.6(–2)	3.57(–2)	3.60(–2)	13	6	2.38(–2)	2.4(–2)		
7	6	2.45(–3)	2.5(–3)			13	10	1.87(–2)	1.8(–2)		
8	1	3.71(–1)	4.0(–1)	3.91(–1)	4.34(–1)	14	1	3.61(–3)	4.32(–3)		
8	2	1.17(–1)	1.2(–1)	1.22(–1)	1.36(–1)	14	2	1.90(–3)	2.24(–3)		
9	1	1.38(–2)	1.4(–2)			14	8	1.23(–1)	1.3(–1)	1.33(–1)	1.46(–1)
9	2	1.40(–1)	1.5(–1)	1.50(–1)	1.67(–1)	14	9	3.70(–2)	3.9(–2)	4.03(–2)	4.41(–2)
9	3	1.04(–1)	1.1(–1)	1.12(–1)	1.24(–1)	14	12	5.26(–3)	5.3(–3)		
9	8	7.19(–3)	7.2(–3)			15	3	6.93(–1)	7.3(–1)	7.34(–1)	8.02(–1)
10	2	5.36(–3)	5.1(–3)			15	4	3.67(–1)	3.9(–1)	3.86(–1)	4.19(–1)
10	3	6.52(–2)	6.43(–2)	6.20(–2)	8.08(–2)	15	6	1.52(–2)	1.4(–2)		
10	4	4.64(–2)	4.46(–2)	4.27(–2)	5.52(–2)	15	10	1.49(–1)	1.5(–1)	1.41(–1)	1.67(–1)
10	5	1.41(–1)	1.5(–1)	1.55(–1)	1.58(–1)	15	11	2.71(–2)	2.7(–2)		
10	6	7.26(–2)	8.0(–2)	8.04(–2)	8.08(–2)	15	13	2.43(–2)	2.5(–2)		
10	7	3.01(–2)	3.3(–2)								

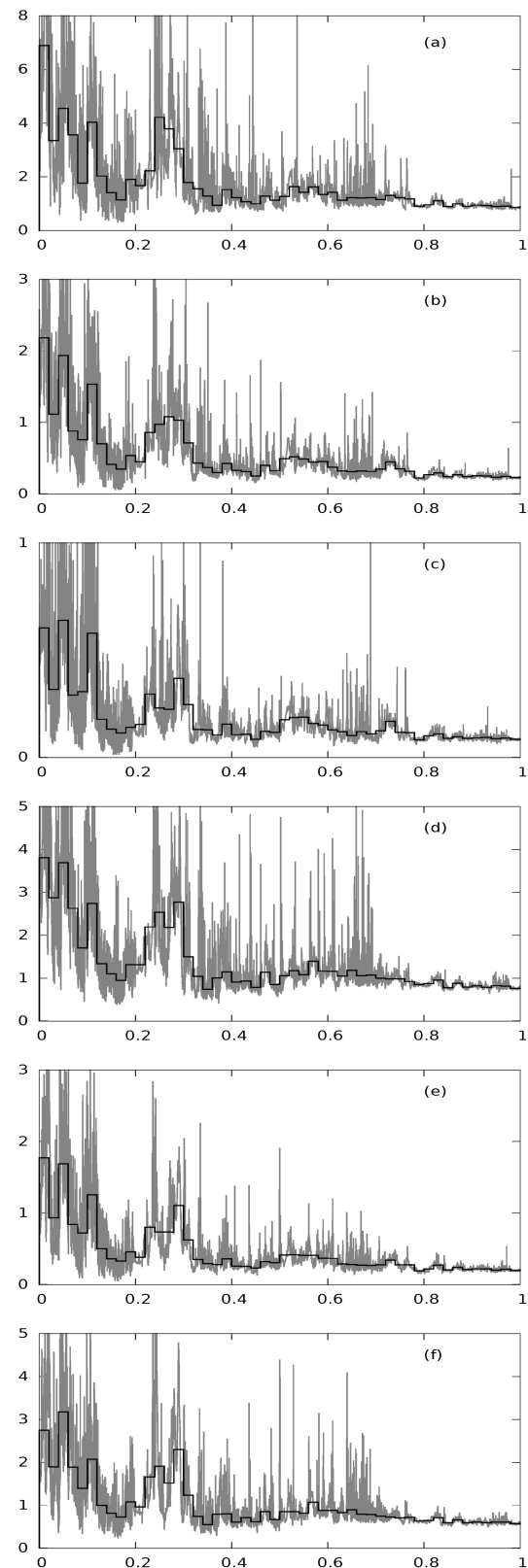
levels of individual terms with differences of a few per cent or less. There are larger differences for the electric quadrupole transition probabilities between terms. For example the probabilities for the principal transitions between the  $a^4F$  and  $a^4P$  terms, the 5-1, 5-2 and 5-3 probabilities, are all larger, by on average 13 per cent, in our calculation than in Hansen et al. (1984). The fact that all three transitions differ by approximately the same factor suggests that the cause of the difference lies in the radial quadrupole integrals used in the two calculations. There is configuration interaction between the terms of the  $3d^7$  electron configuration and the  $3d^6\bar{4}d$  configuration in our calculation and not in the single configuration calculation of Hansen et al. (1984). With this interaction included, the quadrupole line strength involves both the  $3d$  radial quadrupole integral and the  $\bar{4}d$  integral which is significantly larger than for the  $3d$ .

Fivet et al. (2016) have made calculations of forbidden transition probabilities for the twice ionized iron-peak elements from Sc to Ni, including Co, and we compare with their results in Table 6. Their calculations were made with two different methods which we label as FQB1 and FQB2. The FQB2 values were computed with AUTOSTRUCTURE as in the present work. Apart from the magnetic dipole transitions between the levels of the ground term, which agree to all tabulated figures, the FQB2 results for the electric quadrupole transitions between levels of different terms are systematically larger than the present work by 15–20 per cent with half of them differing by the same fixed amount of 19 per cent. As discussed above in the comparison with the work of Hansen et al. (1984), the systematic nature of the difference suggests that it is due to a different value for the  $3d$  radial quadrupole integral rather than details of the wavefunction expansions of individual terms. The configuration expansions in the present work and in that of Fivet et al. (2016) are very similar but differ in one key aspect. We use a somewhat contracted  $\bar{4}d$  orbital to allow for the differences in the  $3d$  orbital between the  $3d^7$  and  $3d^64s$  configurations, while Fivet et al. (2016) employ a spectroscopic  $4d$  orbital but a contracted  $5s$  orbital which provides flexibility to the spectroscopic  $4s$ . These two different expansions give broadly similar energy levels and fine structure but result in differences in the quadrupole radial integrals. It is not clear that either approach is necessarily superior, so the approximately 15–20 per cent differences are probably a realistic measure of the uncertainty in the results for the electric quadrupole line strengths. We note that the results for the electric quadrupole transition probabilities from the FQB1 calculation of Fivet et al. (2016) agree better with their FQB2 for some transitions and better with the current work for others.

### 3 SCATTERING CALCULATIONS

In this work we used the Breit–Pauli R-matrix method, which is detailed in Hummer et al. (1993), Berrington et al. (1995) and the references therein, to perform the scattering calculations. The calculations were made using the R-matrix codes<sup>3</sup> where the serial version of the codes were used in some stages and the parallel version in others. An R-matrix boundary radius of 11.3 au defining the inner region was applied so that the most extended orbital ( $4p$ ) of our target is covered. Each one of the partial waves of the scattered electron was expanded over 12 basis functions within the R-matrix boundary, and the expansion extends to a maximum of  $J = 9$ .

<sup>3</sup> See Badnell: R-matrix write-up on [http://amdpp.phys.strath.ac.uk/UK\\_RmaX/codes/](http://amdpp.phys.strath.ac.uk/UK_RmaX/codes/)



**Figure 2.** Collision strength (vertical axis) versus final electron energy in Ryd (horizontal axis) for the (a) 1–2, (b) 1–3, (c) 1–4, (d) 2–3, (e) 2–4 and (f) 3–4 transitions, where the grey smooth line represents the continuous function while the black discrete line represents binned plot of the same function to show the magnitude of resonance contributions more clearly. For level indexing refer to Table 4.

**Table 7.** Thermally averaged collision strengths among the 15 energetically lowest levels of  $\text{Co}^{2+}$  as a function of  $\log_{10}$  of temperature in Kelvin where  $i$  and  $j$  refer to the index of the lower and upper level, respectively (see Table 4 for indexing).

$i$	$j$	$\log_{10} T$												
		2.0	2.2	2.4	2.6	2.8	3.0	3.2	3.4	3.6	3.8	4.0	4.2	4.4
1	2	4.037	4.171	4.321	4.573	5.001	5.470	5.699	5.586	5.229	4.732	4.177	3.636	3.135
1	3	1.490	1.471	1.511	1.630	1.795	1.926	1.957	1.893	1.769	1.607	1.419	1.228	1.045
1	4	0.429	0.448	0.473	0.502	0.528	0.545	0.544	0.529	0.505	0.473	0.429	0.378	0.325
1	5	1.285	1.328	1.379	1.409	1.404	1.364	1.299	1.226	1.177	1.182	1.221	1.243	1.224
1	6	0.578	0.611	0.626	0.618	0.590	0.551	0.509	0.471	0.456	0.472	0.497	0.506	0.490
1	7	0.232	0.215	0.201	0.188	0.178	0.171	0.164	0.156	0.155	0.160	0.167	0.168	0.162
1	8	0.963	0.944	0.937	0.949	0.973	0.997	1.027	1.071	1.115	1.154	1.193	1.225	1.235
1	9	0.312	0.322	0.315	0.299	0.282	0.270	0.267	0.274	0.283	0.292	0.302	0.309	0.308
1	10	0.361	0.377	0.404	0.422	0.421	0.407	0.393	0.384	0.373	0.361	0.349	0.340	0.332
1	11	0.180	0.166	0.147	0.127	0.109	0.094	0.084	0.077	0.073	0.072	0.073	0.077	0.081
1	12	4.532	3.989	3.412	2.884	2.444	2.093	1.805	1.570	1.407	1.330	1.316	1.327	1.328
1	13	0.375	0.369	0.374	0.393	0.417	0.430	0.431	0.428	0.431	0.442	0.463	0.486	0.500
1	14	0.374	0.374	0.364	0.342	0.313	0.284	0.259	0.242	0.238	0.244	0.254	0.261	0.260
1	15	0.070	0.066	0.065	0.066	0.068	0.070	0.073	0.075	0.078	0.081	0.086	0.090	0.092
2	3	3.301	3.280	3.245	3.264	3.382	3.536	3.617	3.581	3.442	3.209	2.905	2.578	2.258
2	4	0.732	0.760	0.831	0.962	1.139	1.315	1.428	1.457	1.415	1.319	1.184	1.034	0.884
2	5	1.089	1.064	1.046	1.026	0.987	0.930	0.864	0.804	0.772	0.785	0.820	0.838	0.816
2	6	0.658	0.682	0.695	0.690	0.668	0.639	0.606	0.571	0.544	0.540	0.552	0.559	0.551
2	7	0.292	0.274	0.254	0.233	0.215	0.203	0.192	0.183	0.185	0.202	0.223	0.232	0.227
2	8	0.617	0.610	0.602	0.601	0.603	0.605	0.609	0.625	0.648	0.673	0.698	0.716	0.719
2	9	0.490	0.487	0.474	0.461	0.451	0.447	0.454	0.472	0.489	0.501	0.512	0.521	0.524
2	10	0.231	0.230	0.240	0.258	0.275	0.284	0.284	0.281	0.278	0.277	0.281	0.282	0.278
2	11	0.177	0.188	0.195	0.191	0.178	0.160	0.143	0.132	0.125	0.121	0.120	0.120	0.118
2	12	1.319	1.343	1.329	1.257	1.137	0.998	0.862	0.753	0.686	0.664	0.673	0.690	0.697
2	13	0.354	0.349	0.340	0.335	0.338	0.344	0.347	0.351	0.358	0.372	0.388	0.398	0.397
2	14	0.593	0.590	0.585	0.573	0.551	0.523	0.495	0.472	0.464	0.475	0.497	0.523	0.539
2	15	0.163	0.157	0.153	0.150	0.151	0.155	0.161	0.166	0.171	0.177	0.184	0.193	0.200
3	4	1.591	1.611	1.692	1.855	2.071	2.278	2.413	2.462	2.436	2.327	2.143	1.923	1.696
3	5	1.016	0.953	0.886	0.822	0.752	0.677	0.607	0.550	0.518	0.518	0.533	0.538	0.518
3	6	0.544	0.574	0.594	0.592	0.568	0.535	0.501	0.469	0.448	0.448	0.463	0.474	0.467
3	7	0.301	0.283	0.264	0.248	0.240	0.238	0.236	0.231	0.231	0.243	0.262	0.274	0.273
3	8	0.343	0.344	0.342	0.340	0.336	0.329	0.323	0.322	0.326	0.332	0.342	0.350	0.351
3	9	0.509	0.493	0.476	0.466	0.462	0.464	0.475	0.498	0.520	0.537	0.553	0.565	0.569
3	10	0.129	0.130	0.140	0.155	0.171	0.181	0.184	0.182	0.179	0.181	0.188	0.193	0.192
3	11	0.151	0.159	0.172	0.184	0.185	0.176	0.161	0.150	0.142	0.139	0.138	0.136	0.132
3	12	0.312	0.358	0.393	0.401	0.384	0.350	0.312	0.280	0.262	0.261	0.272	0.285	0.292
3	13	0.251	0.244	0.232	0.220	0.213	0.212	0.216	0.223	0.232	0.243	0.255	0.262	0.259
3	14	0.627	0.625	0.640	0.646	0.628	0.591	0.550	0.519	0.507	0.516	0.541	0.572	0.595
3	15	0.179	0.177	0.175	0.175	0.177	0.184	0.193	0.202	0.211	0.221	0.234	0.245	0.250
4	5	0.910	0.818	0.721	0.631	0.546	0.467	0.398	0.342	0.306	0.293	0.291	0.288	0.275
4	6	0.373	0.394	0.404	0.394	0.369	0.340	0.314	0.291	0.281	0.291	0.311	0.324	0.320
4	7	0.261	0.248	0.233	0.223	0.220	0.223	0.226	0.225	0.224	0.228	0.240	0.250	0.251
4	8	0.163	0.166	0.166	0.164	0.159	0.151	0.143	0.138	0.134	0.131	0.131	0.133	0.133
4	9	0.376	0.365	0.356	0.355	0.359	0.365	0.376	0.396	0.416	0.431	0.444	0.456	0.461
4	10	0.059	0.067	0.078	0.089	0.099	0.104	0.105	0.102	0.100	0.102	0.109	0.116	0.118
4	11	0.106	0.105	0.117	0.135	0.147	0.146	0.137	0.127	0.122	0.119	0.118	0.116	0.111
4	12	0.053	0.056	0.062	0.067	0.069	0.069	0.067	0.065	0.065	0.068	0.075	0.081	0.085
4	13	0.161	0.154	0.143	0.130	0.119	0.111	0.108	0.109	0.111	0.116	0.124	0.130	0.132
4	14	0.494	0.501	0.525	0.538	0.525	0.492	0.453	0.420	0.403	0.407	0.429	0.458	0.480
4	15	0.145	0.146	0.146	0.147	0.150	0.157	0.167	0.177	0.187	0.200	0.213	0.223	0.226
5	6	1.391	1.531	1.673	1.759	1.760	1.676	1.526	1.343	1.171	1.043	0.961	0.910	0.875
5	7	0.900	0.835	0.777	0.732	0.689	0.636	0.570	0.499	0.439	0.403	0.386	0.381	0.378
5	8	0.518	0.490	0.473	0.473	0.486	0.493	0.480	0.449	0.407	0.367	0.340	0.328	0.324
5	9	0.339	0.343	0.330	0.304	0.272	0.239	0.210	0.184	0.163	0.146	0.136	0.132	0.133
5	10	0.205	0.197	0.205	0.225	0.242	0.250	0.249	0.245	0.243	0.249	0.263	0.281	0.291
5	11	0.143	0.131	0.118	0.106	0.096	0.089	0.084	0.083	0.083	0.084	0.087	0.089	0.089
5	12	0.265	0.273	0.297	0.332	0.364	0.379	0.376	0.362	0.348	0.345	0.356	0.374	0.388
5	13	0.519	0.495	0.453	0.407	0.364	0.328	0.302	0.287	0.282	0.284	0.289	0.299	0.309
5	14	0.144	0.134	0.121	0.108	0.097	0.089	0.083	0.079	0.079	0.084	0.092	0.101	0.106
5	15	0.116	0.115	0.115	0.114	0.116	0.124	0.135	0.145	0.153	0.158	0.162	0.167	0.169



Table 7 – continued

$i$	$j$	2.0	2.2	2.4	2.6	2.8	3.0	$\log_{10}T$ 3.2	3.4	3.6	3.8	4.0	4.2	4.4
6	7	0.656	0.614	0.589	0.580	0.570	0.546	0.507	0.458	0.411	0.373	0.348	0.333	0.323
6	8	0.247	0.239	0.233	0.232	0.236	0.235	0.227	0.211	0.192	0.173	0.160	0.154	0.152
6	9	0.333	0.315	0.295	0.276	0.258	0.240	0.222	0.204	0.187	0.173	0.165	0.162	0.163
6	10	0.105	0.110	0.123	0.141	0.157	0.167	0.172	0.175	0.178	0.184	0.195	0.205	0.210
6	11	0.176	0.161	0.144	0.128	0.114	0.102	0.094	0.090	0.091	0.094	0.099	0.106	0.108
6	12	0.096	0.113	0.138	0.158	0.164	0.159	0.148	0.137	0.127	0.122	0.123	0.127	0.131
6	13	0.270	0.259	0.242	0.222	0.203	0.188	0.178	0.173	0.173	0.175	0.180	0.187	0.192
6	14	0.193	0.187	0.181	0.175	0.169	0.163	0.159	0.156	0.154	0.156	0.164	0.176	0.184
6	15	0.088	0.088	0.088	0.087	0.087	0.090	0.095	0.103	0.111	0.119	0.125	0.131	0.136
7	8	0.077	0.077	0.080	0.086	0.092	0.094	0.090	0.082	0.072	0.062	0.053	0.048	0.045
7	9	0.164	0.157	0.152	0.149	0.147	0.144	0.138	0.130	0.120	0.111	0.105	0.104	0.103
7	10	0.035	0.037	0.043	0.053	0.063	0.070	0.074	0.076	0.077	0.079	0.082	0.084	0.085
7	11	0.105	0.101	0.095	0.088	0.079	0.071	0.065	0.062	0.063	0.066	0.071	0.076	0.078
7	12	0.025	0.042	0.061	0.071	0.070	0.060	0.049	0.039	0.031	0.026	0.024	0.023	0.024
7	13	0.114	0.113	0.108	0.101	0.094	0.086	0.081	0.077	0.076	0.075	0.075	0.077	0.079
7	14	0.138	0.135	0.132	0.129	0.125	0.121	0.118	0.115	0.113	0.113	0.118	0.125	0.132
7	15	0.043	0.043	0.043	0.044	0.046	0.048	0.051	0.055	0.059	0.063	0.067	0.071	0.075
8	9	1.248	1.372	1.481	1.552	1.573	1.557	1.531	1.501	1.448	1.388	1.369	1.406	1.445
8	10	0.419	0.446	0.483	0.516	0.543	0.570	0.594	0.606	0.601	0.581	0.561	0.550	0.544
8	11	0.411	0.384	0.358	0.340	0.332	0.324	0.316	0.306	0.292	0.273	0.254	0.242	0.235
8	12	1.282	1.265	1.281	1.321	1.362	1.385	1.383	1.380	1.434	1.596	1.839	2.063	2.164
8	13	0.735	0.717	0.692	0.664	0.643	0.628	0.615	0.607	0.607	0.629	0.689	0.775	0.841
8	14	0.956	0.989	1.015	1.033	1.039	1.025	1.000	0.978	0.990	1.058	1.160	1.247	1.278
8	15	0.358	0.372	0.386	0.391	0.387	0.376	0.363	0.349	0.339	0.341	0.362	0.395	0.418
9	10	0.493	0.429	0.384	0.367	0.376	0.401	0.427	0.446	0.453	0.445	0.437	0.437	0.440
9	11	0.411	0.390	0.386	0.402	0.417	0.414	0.395	0.372	0.346	0.318	0.297	0.288	0.286
9	12	0.426	0.516	0.639	0.763	0.846	0.874	0.862	0.838	0.830	0.856	0.909	0.962	0.988
9	13	0.422	0.410	0.402	0.396	0.395	0.398	0.405	0.415	0.428	0.450	0.490	0.539	0.574
9	14	0.990	0.967	0.940	0.909	0.878	0.856	0.847	0.856	0.915	1.055	1.248	1.413	1.482
9	15	0.293	0.315	0.342	0.364	0.376	0.381	0.382	0.382	0.383	0.392	0.420	0.464	0.498
10	11	0.357	0.338	0.334	0.349	0.371	0.387	0.403	0.427	0.446	0.451	0.450	0.449	0.440
10	12	0.158	0.174	0.192	0.199	0.190	0.172	0.154	0.143	0.144	0.156	0.173	0.186	0.189
10	13	0.587	0.573	0.558	0.547	0.546	0.557	0.580	0.609	0.635	0.660	0.695	0.731	0.747
10	14	0.099	0.098	0.097	0.095	0.095	0.098	0.102	0.106	0.115	0.129	0.144	0.154	0.156
10	15	0.423	0.454	0.495	0.534	0.558	0.571	0.578	0.575	0.562	0.549	0.542	0.540	0.530
11	12	0.053	0.064	0.075	0.080	0.076	0.068	0.059	0.052	0.050	0.054	0.058	0.060	0.058
11	13	0.319	0.296	0.275	0.259	0.255	0.265	0.284	0.302	0.313	0.321	0.332	0.344	0.350
11	14	0.042	0.041	0.040	0.040	0.042	0.046	0.050	0.055	0.061	0.068	0.076	0.083	0.087
11	15	0.178	0.198	0.227	0.258	0.285	0.304	0.313	0.310	0.297	0.283	0.276	0.274	0.268
12	13	0.343	0.332	0.316	0.299	0.286	0.283	0.292	0.310	0.336	0.368	0.400	0.425	0.434
12	14	2.767	2.722	2.622	2.476	2.301	2.119	1.958	1.841	1.805	1.880	2.050	2.241	2.361
12	15	0.073	0.072	0.073	0.079	0.088	0.098	0.106	0.114	0.130	0.151	0.169	0.179	0.178
13	14	0.164	0.168	0.171	0.171	0.171	0.171	0.174	0.182	0.197	0.220	0.247	0.269	0.277
13	15	0.295	0.302	0.314	0.328	0.346	0.378	0.427	0.481	0.526	0.564	0.610	0.659	0.688
14	15	0.160	0.162	0.170	0.185	0.200	0.212	0.222	0.232	0.246	0.261	0.277	0.293	0.301

Collision strengths were computed over two non-overlapping energy meshes: a fine mesh consisting of 20 000 evenly divided intervals which goes from zero up to the highest target threshold (about 1.2 Ryd), and a coarse mesh consisting of 2000 evenly divided intervals which reach 1 Ryd above the highest target threshold. The purpose of the first mesh is to cover the main resonance region while the second mesh is intended to cover the region where all scattering channels are open, up to an incident electron energy of about 2.2 Ryd. Our results demonstrate that these meshes have achieved these purposes. In Fig. 2 we illustrate our results with the computed collision strengths between the lowest four levels of the ground  $3d^7\ ^4F$  term as a function of final electron energy up to 1 Ryd above threshold. Dense and complex resonance structure can be seen in these plots due to the multiple close lying thresh-

olds. We also show the collision strength averaged over 0.02 Ryd intervals.

To ensure that the computed collision strengths have converged in partial wave for all the levels for which data are given, the contribution of partial wave  $J = 9$  was compared to the sum for all transitions and energies. This comparison showed that in almost all cases the contribution from  $J = 9$  is negligible. Specifically, the largest contribution from  $J = 9$  is for the transition 8–12 at about 1 per cent and the next largest is about 0.1 per cent of the total. However, we note that it is certain that the collision strengths from the lower levels to the levels of the  $4p$  configuration are not converged because they are allowed transitions which have significant high partial wave contributions. We therefore do not provide collision strengths for any of these transitions.

## 4 RESULTS AND DISCUSSION

The thermally averaged collision strengths between the 15 lowest energy levels are given in Table 7 as a function of electron temperature. These values were calculated using the full energy range, as described above. In the energy region where all scattering channels are open there are some small irregular features in the collision strengths that are almost certainly non-physical and caused by the correlation orbital in the target representation. We computed thermally averaged collision strengths for the transitions and temperature range given in Table 7 both including and excluding the contribution from the region of all channels open, and found the largest change for any transition is 0.3 per cent at  $\log_{10}T = 4.0$ , 2.4 per cent at  $\log_{10}T = 4.2$  and 9.4 per cent at  $\log_{10}T = 4.4$ . The values tabulated in Table 7 were computed using the full energy range.

### 4.1 Principal spectral lines

We compute the predicted  $\text{Co}^{2+}$  fractional level populations using the results in Tables 6 and 7 with a 15 level model atom including electron collisional excitation and de-excitation and radiative decay. In Tables 8 and 9 we show the resulting 10 strongest lines of  $\text{Co III}$  in this model. We also ensure that the three  $\text{Co III}$  mid-infrared lines

**Table 8.** The emissivity ratio,  $\rho$ , of the 10 strongest lines of  $\text{Co III}$  in our 15-level model atom for electron temperature  $T_e = 10^4$  K and electron number density  $N_e = 10^4 \text{ cm}^{-3}$  typical of planetary nebulae (PNe). We also add the 24.06  $\mu\text{m}$  line to show the relative strength of all three mid-infrared lines. The powers of 10 of the  $\rho$  values are given in brackets and  $i$  and  $j$  refer to the lower and upper levels, respectively, as indexed in Table 4.

$j$	$i$	Transition	$\lambda$	$\rho$
2	1	$3d^7 \text{ } ^4\text{F}_{7/2} - 3d^7 \text{ } ^4\text{F}_{9/2}$	11.88 $\mu\text{m}$	5.85(+4)
8	1	$3d^7 \text{ } ^2\text{G}_{9/2} - 3d^7 \text{ } ^4\text{F}_{9/2}$	5888.48 $\text{\AA}$	2.83(+4)
5	1	$3d^7 \text{ } ^4\text{P}_{5/2} - 3d^7 \text{ } ^4\text{F}_{9/2}$	6576.31 $\text{\AA}$	2.27(+4)
3	2	$3d^7 \text{ } ^4\text{F}_{5/2} - 3d^7 \text{ } ^4\text{F}_{7/2}$	16.39 $\mu\text{m}$	1.26(+4)
8	2	$3d^7 \text{ } ^2\text{G}_{9/2} - 3d^7 \text{ } ^4\text{F}_{7/2}$	6195.45 $\text{\AA}$	8.47(+3)
6	2	$3d^7 \text{ } ^4\text{P}_{3/2} - 3d^7 \text{ } ^4\text{F}_{7/2}$	6853.53 $\text{\AA}$	6.93(+3)
5	2	$3d^7 \text{ } ^4\text{P}_{5/2} - 3d^7 \text{ } ^4\text{F}_{7/2}$	6961.53 $\text{\AA}$	5.83(+3)
13	2	$3d^7 \text{ } ^2\text{D}_{5/2} - 3d^7 \text{ } ^4\text{F}_{7/2}$	4499.67 $\text{\AA}$	4.07(+3)
6	3	$3d^7 \text{ } ^4\text{P}_{3/2} - 3d^7 \text{ } ^4\text{F}_{5/2}$	7152.69 $\text{\AA}$	3.91(+3)
12	8	$3d^7 \text{ } ^2\text{H}_{11/2} - 3d^7 \text{ } ^2\text{G}_{9/2}$	1.741 $\mu\text{m}$	3.82(+3)
4	3	$3d^7 \text{ } ^4\text{F}_{3/2} - 3d^7 \text{ } ^4\text{F}_{5/2}$	24.06 $\mu\text{m}$	1.94(+3)

**Table 9.** The emissivity ratio,  $\rho$ , of the 10 strongest lines plus the three mid-infrared lines of  $\text{Co III}$  with  $N_e = 10^7 \text{ cm}^{-3}$  typical of SN remnants. The other details are as in Table 8.

$j$	$i$	Transition	$\lambda$	$\rho$
8	1	$3d^7 \text{ } ^2\text{G}_{9/2} - 3d^7 \text{ } ^4\text{F}_{9/2}$	5888.48 $\text{\AA}$	1.26(+4)
13	2	$3d^7 \text{ } ^2\text{D}_{5/2} - 3d^7 \text{ } ^4\text{F}_{7/2}$	4499.67 $\text{\AA}$	5.39(+3)
9	2	$3d^7 \text{ } ^2\text{G}_{7/2} - 3d^7 \text{ } ^4\text{F}_{7/2}$	5906.78 $\text{\AA}$	3.82(+3)
8	2	$3d^7 \text{ } ^2\text{G}_{9/2} - 3d^7 \text{ } ^4\text{F}_{7/2}$	6195.45 $\text{\AA}$	3.78(+3)
9	3	$3d^7 \text{ } ^2\text{G}_{7/2} - 3d^7 \text{ } ^4\text{F}_{5/2}$	6127.67 $\text{\AA}$	2.74(+3)
5	1	$3d^7 \text{ } ^4\text{P}_{5/2} - 3d^7 \text{ } ^4\text{F}_{9/2}$	6576.31 $\text{\AA}$	2.62(+3)
15	3	$3d^7 \text{ } ^2\text{D}_{3/2} - 3d^7 \text{ } ^4\text{F}_{5/2}$	4387.52 $\text{\AA}$	2.38(+3)
15	4	$3d^7 \text{ } ^2\text{D}_{3/2} - 3d^7 \text{ } ^4\text{F}_{3/2}$	4469.02 $\text{\AA}$	1.24(+3)
6	2	$3d^7 \text{ } ^4\text{P}_{3/2} - 3d^7 \text{ } ^4\text{F}_{7/2}$	6853.53 $\text{\AA}$	9.23(+2)
14	8	$3d^7 \text{ } ^2\text{H}_{9/2} - 3d^7 \text{ } ^2\text{G}_{9/2}$	1.548 $\mu\text{m}$	7.49(+2)
2	1	$3d^7 \text{ } ^4\text{F}_{7/2} - 3d^7 \text{ } ^4\text{F}_{9/2}$	11.88 $\mu\text{m}$	6.76(+2)
3	2	$3d^7 \text{ } ^4\text{F}_{5/2} - 3d^7 \text{ } ^4\text{F}_{7/2}$	16.39 $\mu\text{m}$	2.17(+2)
4	3	$3d^7 \text{ } ^4\text{F}_{3/2} - 3d^7 \text{ } ^4\text{F}_{5/2}$	24.06 $\mu\text{m}$	3.26(+1)

at 11.88, 16.39 and 24.06  $\mu\text{m}$  are in the tables even if they are not among the 10 strongest. The 15 levels are all of even parity so all these lines are  $[\text{Co III}]$  forbidden transitions. The tabulated quantity  $\rho$  is the ratio of the energy emitted per unit time in a  $\text{Co III}$  line relative to  $\text{H}\beta$  for unit  $\text{Co}^{2+}$  and  $\text{H}^+$  ion number density. Hence for a downward transition of wavelength  $\lambda_{ij}$  between  $\text{Co}^{2+}$  levels  $j$  and  $i$ :

$$\rho = \frac{f_j A_{ji} \lambda_{\text{H}\beta}}{N_e \alpha_e^{\text{H}\beta} \lambda_{ij}}, \quad (1)$$

where  $f_j$  is the fraction of  $\text{Co}^{2+}$  in the upper state  $j$ ,  $A_{ji}$  is the Einstein  $A$  coefficient for the transition,  $\lambda_{\text{H}\beta}$  is the  $\text{H}\beta$  wavelength and  $\alpha_e^{\text{H}\beta}$  is the effective recombination coefficient for  $\text{H}\beta$  whose value is obtained from Storey & Hummer (1995). The values of  $\rho$  are tabulated for a temperature of  $10^4$  K and for two electron densities,  $N_e = 10^4 \text{ cm}^{-3}$  typical of PNe (Table 8), and  $N_e = 10^7 \text{ cm}^{-3}$  more typical of SN remnants in their nebular phase (Table 9). Thus, in typical PN conditions, assuming a Co abundance of  $10^{-7}$  with respect to  $\text{H}^+$  by number and assuming 20 per cent of Co is in the form of  $\text{Co}^{2+}$ , the brightest visible  $\text{Co III}$  line at 5888.5  $\text{\AA}$  would have an emissivity per unit volume  $5.7 \times 10^{-4}$  times that of  $\text{H}\beta$ . In principle this would be visible in deep spectra of bright PNe (e.g. Baluteau et al. 1995). In practice, Baluteau et al. (1995) do not identify this line in the spectrum of NGC 7027 which may reflect depletion of gas phase Co on dust grains.

In Storey et al. (2016) we reported collision strengths and transition probabilities for low-lying transitions in  $\text{Co II}$  and discussed the spectroscopic uses of the three mid-infrared lines at 10.52, 14.74 and 15.46  $\mu\text{m}$ . There are also significant  $\text{Co II}$  visible and near-infrared lines which were not discussed by Storey et al. (2016), so in Tables 10 and 11 we show the strongest of these. The  $\text{Co II}$  model atom also comprises the energetically lowest 15 levels and the transition probabilities and thermally averaged collision strengths required are all from Storey et al. (2016).

## 5 CONCLUSIONS

In this study, the  $\text{Co III}$  forbidden lines arising from transitions between the 15 lowest energy levels of doubly ionized cobalt,  $\text{Co}^{2+}$ , have been investigated. Radiative transition probabilities and collision strengths for excitation and de-excitation by electron scattering, with their thermally averaged values based on a Maxwell–Boltzmann statistics, have been computed and reported. The scattering calculations used the R-matrix method in the Breit–Pauli approximation under an intermediate coupling scheme.

**Table 10.** The emissivity ratio,  $\rho$ , of the 10 strongest lines of  $\text{Co II}$  with  $N_e = 10^4 \text{ cm}^{-3}$  typical of PNe. We also add the 15.46  $\mu\text{m}$  line discussed by Storey et al. (2016). The other details are as in Table 8.

$j$	$i$	Transition	$\lambda$	$\rho$
9	1	$3d^7 4s \text{ } ^3\text{F}_4 - 3d^8 \text{ } ^3\text{F}_4$	1.019 $\mu\text{m}$	1.17(+5)
9	4	$3d^7 4s \text{ } ^3\text{F}_4 - 3d^7 4s \text{ } ^5\text{F}_5$	1.547 $\mu\text{m}$	6.53(+4)
2	1	$3d^8 \text{ } ^3\text{F}_3 - 3d^8 \text{ } ^3\text{F}_4$	10.52 $\mu\text{m}$	3.81(+4)
5	4	$3d^7 4s \text{ } ^5\text{F}_4 - 3d^7 4s \text{ } ^5\text{F}_5$	14.74 $\mu\text{m}$	2.86(+4)
12	2	$3d^8 \text{ } ^1\text{D}_2 - 3d^8 \text{ } ^3\text{F}_3$	9342.56 $\text{\AA}$	1.74(+4)
9	2	$3d^7 4s \text{ } ^3\text{F}_4 - 3d^8 \text{ } ^3\text{F}_3$	1.128 $\mu\text{m}$	1.27(+4)
9	6	$3d^7 4s \text{ } ^3\text{F}_4 - 3d^7 4s \text{ } ^5\text{F}_3$	1.903 $\mu\text{m}$	9.33(+3)
13	2	$3d^8 \text{ } ^3\text{P}_2 - 3d^8 \text{ } ^3\text{F}_3$	8121.13 $\text{\AA}$	9.04(+3)
12	3	$3d^8 \text{ } ^1\text{D}_2 - 3d^8 \text{ } ^3\text{F}_2$	9943.60 $\text{\AA}$	8.13(+3)
10	2	$3d^7 4s \text{ } ^3\text{F}_3 - 3d^8 \text{ } ^3\text{F}_3$	1.025 $\mu\text{m}$	7.27(+3)
3	2	$3d^8 \text{ } ^3\text{F}_2 - 3d^8 \text{ } ^3\text{F}_3$	15.46 $\mu\text{m}$	5.00(+3)

**Table 11.** The emissivity ratio,  $\rho$ , of the 10 strongest lines of Co II with  $N_e = 10^7 \text{ cm}^{-3}$  typical of SN remnants. We also add the 10.52, 14.74 and 15.46  $\mu\text{m}$  lines discussed by Storey et al. (2016). The other details are as in Table 8.

$j$	$i$	Transition	$\lambda$	$\rho$
12	2	$3d^8 \ ^1D_2 - 3d^8 \ ^3F_3$	9342.56 Å	4.25(+3)
9	1	$3d^7 4s \ ^3F_4 - 3d^8 \ ^3F_4$	1.019 $\mu\text{m}$	2.41(+3)
12	3	$3d^8 \ ^1D_2 - 3d^8 \ ^3F_2$	9943.60 Å	1.99(+3)
13	2	$3d^8 \ ^3P_2 - 3d^8 \ ^3F_3$	8121.13 Å	1.72(+3)
9	4	$3d^7 4s \ ^3F_4 - 3d^7 4s \ ^5F_5$	1.547 $\mu\text{m}$	1.35(+3)
10	2	$3d^7 4s \ ^3F_3 - 3d^8 \ ^3F_3$	1.025 $\mu\text{m}$	1.07(+3)
13	1	$3d^8 \ ^3P_2 - 3d^8 \ ^3F_4$	7539.01 Å	9.27(+2)
11	2	$3d^7 4s \ ^3F_2 - 3d^8 \ ^3F_3$	9639.21 Å	8.72(+2)
11	3	$3d^7 4s \ ^3F_2 - 3d^8 \ ^3F_2$	1.028 $\mu\text{m}$	8.70(+2)
10	1	$3d^7 4s \ ^3F_3 - 3d^8 \ ^3F_4$	9335.84 Å	7.67(+2)
2	1	$3d^8 \ ^3F_3 - 3d^8 \ ^3F_4$	10.52 $\mu\text{m}$	4.46(+2)
5	4	$3d^7 4s \ ^5F_4 - 3d^7 4s \ ^5F_5$	14.74 $\mu\text{m}$	1.41(+2)
3	2	$3d^8 \ ^3F_2 - 3d^8 \ ^3F_3$	15.46 $\mu\text{m}$	8.48(+1)

The emissivities of the Co III forbidden lines were calculated with a 15-level  $\text{Co}^{2+}$  model atom and the strongest lines listed with their expected strength relative to  $\text{H}\beta$  for conditions approximately representative of those in PNe and SN remnants. For comparison and completeness we also listed the strongest forbidden lines from Co II in the same conditions based on atomic parameters calculated and presented in a previous paper (Storey et al. 2016).

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## SUPPORTING INFORMATION

Additional Supporting Information may be found in the online version of this article:

### CoIII\_Collision\_Strength\_List.zip

(<http://www.mnras.oxfordjournals.org/lookup/suppl/doi:10.1093/mnras/stw776/-/DC1>).

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