

# Multimessenger constraints on outflows from neutron star mergers

A thesis submitted  
in partial fulfillment for the award of the degree of

**Master of Science**

in

**Astronomy and Astrophysics**

by

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June 4, 2021



# Certificate

This is to certify that the thesis titled *Multimessenger constraints on outflows from neutron star mergers* submitted by **B.S.Bharath Saiguhan**, to the Indian Institute of Space Science and Technology, Thiruvananthapuram, in partial fulfillment for the award of the degree of **Master of Science in Astronomy and Astrophysics** is a bona fide record of the original work carried out by him/her under my supervision. The contents of this thesis, in full or in parts, have not been submitted to any other Institute or University for the award of any degree or diploma.

Dr. Resmi Lekshmi  
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**Place:** Thiruvananthapuram  
**Date:** June 4, 2021



# Declaration

I declare that this thesis titled *Multimessenger constraints on outflows from neutron star mergers* submitted in partial fulfillment for the award of the degree of **Master of Science in Astronomy and Astrophysics** is a record of the original work carried out by me under the supervision of **Dr. Resmi Lekshmi**, and has not formed the basis for the award of any degree, diploma, associateship, fellowship, or other titles in this or any other Institution or University of higher learning. In keeping with the ethical practice in reporting scientific information, due acknowledgments have been made wherever the findings of others have been cited.

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(SC16B123)



*This thesis is dedicated to ...*



# Acknowledgements

I acknowledge ...

B.S.Bharath Saiguhan



# Abstract

Abstract here.



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# Chapter 1

## Introduction

### 1.1 Outflows from BNS Mergers

Due to the joint electromagnetic and gravitational wave detection of the binary neutron star merger event GW170817 (see B. P. Abbott et al. 2018), there has been renewed interest in SGRBs. Specifically, this detection gives credence to the claim that the central engines of SGRBs are binary neutron star mergers (see Narayan, Paczynski, and Piran 1992). However, there are other aspects of the process that are not as clear. Specifically, several aspects of GW170817 have not been completely explained. The main concerns are as follows (Lazzati 2020):

- The outflow of GRB170817A was lower in energy than a typical cosmological SGRB, by a factor of  $10^4 - 10^5$ , even though the event was one of the closest GW events recorded, at a distance of  $\sim 40$  Mpc. This could be due to two factors :
  - The structured jet was off-axis with respect to the observer.
  - The internal engine powering this SGRB was intrinsically less energetic, and differs from the one observed in other typical SGRBs.
- A clear consensus has not been reached on how the gamma-ray prompt emission was produced. The uncertainty partly comes from the fact that the various delays involved before the prompt emission is observed are not accurately constrained. Models which have been considered include :
  - The structured outflow model, characterised by functions for the Lorentz factor and the energy per unit solid angle, both of which vary with the

angle made with the polar jet axis ( $\theta$ ). This model produces detectable signals even at moderately large off-axis angles.

- The shock breakout model, wherein the leading edge of the wind emits the prompt emission as it breaks out of the cocoon of nuclear matter ejected before the jet was launched. This model has been shown to explain the energetics and spectrum of the prompt emission, although it does require a setup in which the wind is fast enough so that it can reach a large enough distance at breakout.

More light can be shed on these questions by observing more such SGRBs, using both the gravitational wave (GW) and electromagnetic (EM) windows. However, the possibility of joint detections are slim, due to the fact that the EM observations are highly dependent on the viewing angle of the system with respect to the observer (due to relativistic beaming), whereas GW signal amplitudes depend on the distance to the event (see ).

Given that this is the case, it would be expedient to look for constraints on the structure parameters of various models. Furthermore, it would be useful to develop models which are resilient to non-detections, and can produce constraints on the parameters using even upper limits on the flux/fluence observed by the various EM follow-up satellites, such as INTEGRAL, Fermi or Swift.

### 1.1.1 Modelling outflows from BNS Mergers

As mentioned before, the electromagnetic follow-up of the binary neutron star merger event GW170817 helped measure the various time delays between the time of the GW signal trigger (which roughly is the merger time itself) and the time the gamma-ray signals were picked up. This time delay is denoted  $\Delta t_{GW-\gamma}$ , and was around 1.75 seconds for this event. The components which make up this delay are as follows (Lazzati 2020):

- **Engine Delay** – this is the delay due to a transition mechanism in the central engine which powers the jet (such as a metastable, fast spinning neutron star collapsing into a black hole when its rotation period increases; this process can take years) or due to the time elapsed in amplifying the magnetic field to a

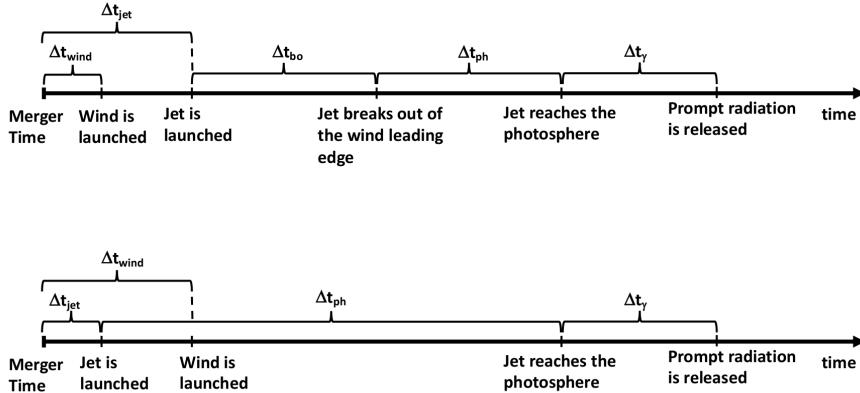
value large enough for jet launching (this process is significantly faster, taking only seconds). This is denoted by  $\Delta t_{eng}$ .

- **Wind Delay** – this is simply a delay in the launching of a non-relativistic wind due to the neutron-rich matter from the progenitor(s) being tidally shredded. For this reason, it can be *negative* as well, since the tidal shredding can occur before the merger itself. This is denoted as  $\Delta t_{wind}$ .
- **Breakout Delay** – if the wind is ejected before the jet, then the latter will have to propagate through the former. This happens at a sub-relativistic speed, whereas the GW signal travels at a relativistic speed. The delay due to this crossing is the breakout delay, and is denoted  $\Delta t_{bo}$ . During this time, jet-wind interactions cause the development of a structured outflow that maintains a bright core but also has energetic wings at large polar angles.
- **Photospheric Delay** – once the jet has crossed the wind, it still needs to propagate out to the photospheric radius, where the outflow becomes transparent and the prompt gamma-ray emission is radiated. The delay from the breakout radius to the photospheric radius is  $\Delta t_{ph}$ . This is given by (for GW170817):

$$\Delta t_{ph} \sim \frac{R_{ph}}{c\Gamma^2} = 1.4 \frac{R_{ph}}{2 \times 10^{12} \text{ cm}} \left( \frac{7}{\Gamma} \right)^2 \text{ s} \quad (1.1)$$

- **Dissipation Delay** – this is a requirement in some models, such as the internal shock synchrotron model, wherein the outflow needs to travel to the internal shock radius before the bulk energy of the flow is dissipated and turned into radiation. The time required to get to this point after crossing the photospheric radius is the dissipation delay, denoted  $\Delta t_\gamma$

Several attempts to constrain the various time delay components have been made. However, except for relative comparisons, no conclusions have been arrived upon. For example, one can only say that the photospheric delay is the major component out of all the delays, and that wind delay (if non-zero) has to be lesser than the jet delay, so that the jet catches up to the wind and the jet-wind interaction generates the structured outflow. Fig. 1.1 summarises these delays in the broader context.



**Figure 1.1:** Two possible scenarios for the relative positioning of the delays in time, which contribute to  $\Delta t_{GW-\gamma}$ . Owing to the requirement of a structured outflow, GW170817 possibly follows the top timeline. The relative contributions of the various delays are debated, but it is agreed that  $\Delta t_{wind} < \Delta t_{jet} \ll 1$  s,  $\Delta t_{bo} \ll 1$  s,  $\Delta t_\gamma \sim 0$  and  $\Delta t_{ph} \sim \Delta t_{GW-\gamma}$ .

Due to the uncertainties in the delay terms, several models for the jet can explain the energetics and observed structure. Numerical simulations are also unequivocal about their favouring of one model over the other (see Shibata and Hotokezaka 2019). Some models try to explain the *apparent* structure of the jet, which are the observables seen by a particular observer at a particular viewing angle. Other models are used to explain the *intrinsic* structure, such as the polar angle variation of the bulk Lorentz factor and the energy across the solid angle, in the jet co-moving frame. See Salafia et al. 2015 for a detailed discussion of the differences between the two structures. Some of the models considered are described below (see also Figs. 1.2 and 1.3):

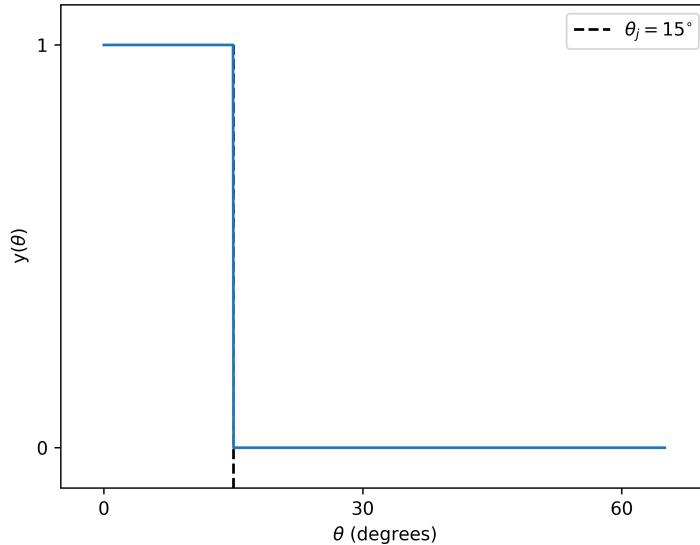
- Top-hat – This model, as used in M. Saleem et al. 2020, assumes that the bulk Lorentz and energy functions drop to zero beyond some cutoff angle,  $\theta_j$ . Below this threshold, the functions are at their respective on-axis values.
- Gaussian – This model is widely used, in some contexts to explain the apparent jet structure (by Hayes et al. 2020), and in others the intrinsic jet structure (by M. Saleem et al. 2020). The former is simply given by  $y_{GJ}(\theta) = e^{-\frac{1}{2}\left(\frac{\theta}{\theta_\sigma}\right)^2}$ , since the authors consider only the apparent jet structure, as explained above and  $\theta_\sigma$  is a structure parameter which is inferred by the authors' Bayesian inference. In the latter, as the authors consider the intrinsic jet structure, they assume

that  $\Gamma\beta(\theta) = \Gamma_0\beta_0 \exp(-\theta^2/2\theta_c^2)$  and that  $\epsilon(\theta) \propto \exp(-\theta^2/\theta_c^2)$ <sup>1</sup>, and derive the observed properties (see below).

- Power Law – This model is used by Hayes et al. 2020 to explain the apparent structure of the jet, assuming that any variation in the energy is simply because of relativistic beaming and the jet being viewed off-axis. It is given using the shape function  $y(\theta)$  (which is multiplied with the on-axis isotropic equivalent energy  $E_{iso,0}$  to give  $E_{iso}(\theta)$ )<sup>2</sup>:

$$y(\theta) = \begin{cases} 1, & 0 \leq \theta \leq \theta_c, \\ (\theta/\theta_c)^{-2}, & \theta_c < \theta \leq \theta_j, \\ 0, & \theta_j < \theta \end{cases} \quad (1.2)$$

Here  $\theta_c$  and  $\theta_j$  are simply structure parameters, inferred using Bayesian methods.

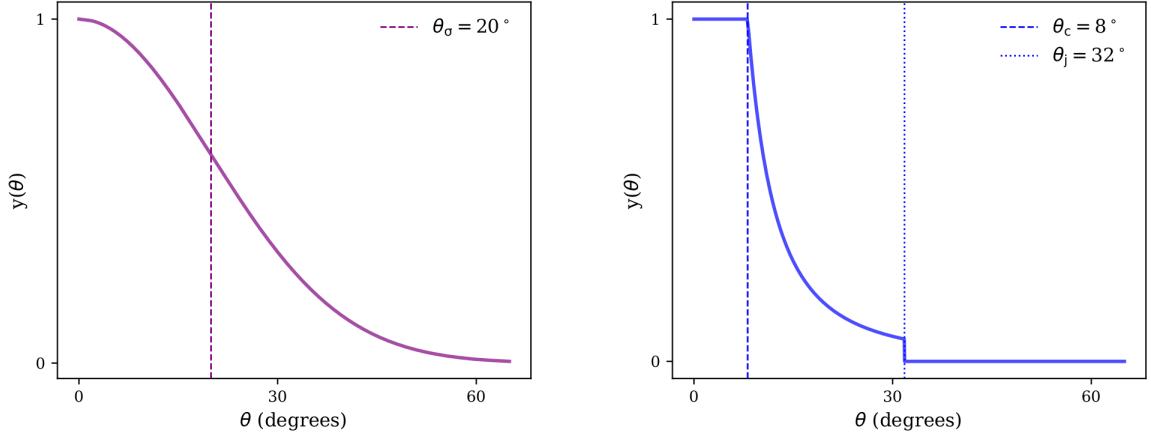


**Figure 1.2:** Functional form of the tophat jet structure model, as considered in M. Saleem et al. 2020. The dashed line denotes the jet angle  $\theta_j = 15^\circ$ .

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<sup>1</sup>This is the normalised energy profile function. The normalisation constant is estimated by the condition  $2\pi \int d(\cos \theta) \epsilon(\theta) = E_{tot.,\gamma}$ , where  $E_{tot.,\gamma}$  is the total energy in gamma-rays.

<sup>2</sup>Using the equation  $E_{iso}(\theta_v) = E_{iso,0} \cdot y(\theta_v)$



**Figure 1.3:** Functional forms of the jet structure models, as considered by Hayes et al. 2020. (Left) The gaussian jet structure with a width  $\theta_\sigma = 20^\circ$ , also marked by the dashed line. (Right) The power-law structure with a core angle  $\theta_c = 8^\circ$  and a jet angle  $\theta_j = 32^\circ$ .

In order to compute the observed structure given the intrinsic structure, following Granot et al. 2002 we start off by considering the emission profile of a point source moving at some angle with the observer, essentially rendering this scenario off-axis. This will affect the prompt jet emission, as well as the initial afterglow, and so warrants careful analysis. Now, let the initial jet opening angle be  $\theta_0$  and let the observer be at an angle  $\theta_{obs}$ . In general, for a point source moving at any angle  $\theta$  with respect to the observer, the observed flux is given by :

$$F_\nu = \frac{L'_{\nu'}}{4\pi d_L^2} \left( \frac{\nu}{\nu'} \right)^3 = \frac{1+z}{4\pi d_L^2} \frac{L'_{\nu'}}{\gamma^3 (1-\beta \cos \theta)^3} \quad (1.3)$$

Here,  $L'_{\nu'}$  and  $\nu'$  are the jet comoving frame spectral luminosity and frequency,  $d_L$  is the luminosity distance,  $\gamma = (1-\beta^2)^{-1/2}$  is the jet Lorentz factor. If  $t$  and  $\nu$  are the observed time and frequency for an observer at  $\theta$ , and  $t_0$  and  $\nu_0$  are those for an observer on the axis, then:

$$\frac{t_0}{t} = \frac{\nu}{\nu_0} = \frac{(1-\beta)}{(1-\beta \cos \theta)} \equiv a \approx \frac{1}{(1+\gamma^2 \theta^2)} \quad (1.4)$$

And finally putting Eq. 1.4 into Eq. 1.3 and expanding using a Taylor series approximation upto the leading order:

$$F_\nu(\theta_{obs}, t) = a^3 F_{\nu/a}(0, at) \quad (1.5)$$

This gives us a handle on how to relate observed off-axis quantities to the on-axis ones. Furthermore, this enables us to go from an intrinsic structure to an observed one, which is what was required.

## 1.2 Outflows from NSBH Mergers

The main difference in the NSBH merger pathway to SGRBs, compared to the case of BNS mergers, is that though there is theoretical and simulational support for the launching of SGRB jets from the merger of a neutron star and a black hole of appropriate mass (see for example Ruiz et al. 2020, Shibata and Hotokezaka 2019, Francois Foucart 2020), there has not been strong observational evidence for the same. In the first half of the third observing run of the LVC (also known as O3a), there have been several triggers which have been reportedly confident NSBH triggers. However, there were no counterpart EM signals picked up, which decreases the credibility of NSBH mergers as the progenitors of SGRBs.

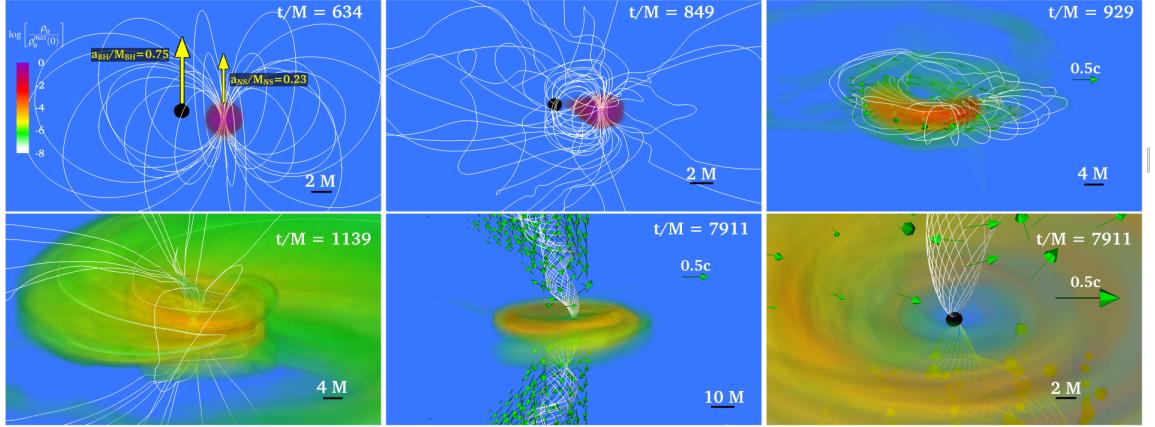
The electromagnetic component from NSBH mergers, is largely decided based on the amount of mass left post-merger, outside the horizon of the black hole. This decides how much matter participates in the subsequent processes, which may be the rapid neutron-capture process which gives rise to the kilonova signal or the magnetic field amplification process via the Magneto-Rotational Instability (MRI) which leads to a SGRB jet.

Qualitatively, for a binary where the neutron star is treated as a test mass and the black hole's spin is aligned with the orbital angular momentum of the binary, the innermost-stable circular orbit radius  $r_{ISCO}$  scales as  $r_{ISCO} \sim f(\chi_{BH})GM_{BH}/c^2$  (where  $f$  is a function ranging from 1 to 9, decreasing for increasing (prograde) spins; see Bardeen, Press, and Teukolsky 1972) and the radius at which the tidal disruption of the neutron star occurs,  $r_{dis}$  scales as  $r_{dis} \sim k(M_{BH/M_{BNS}})^{1/3}R_{NS}$  (where  $k$  is a constant with a dependence on the black hole spin and the equation of state). Only requiring that  $r_{dis} \gtrsim r_{ISCO}$ , as a rough requirement for disruption to occur before the neutron star plunges into the black hole, leads to the conclusion that (a.) low-mass black holes (b.) larger NS radii (c.) higher prograde black hole spins favour disruption. This is seen from Fig. 1.7 as well. However, for actual quantitative results simulations need to be performed such the effect of the various components in the problem are correctly taken into account. As seen from the literature, wherein such general-relativistic magnetohydrodynamic simulations are carried out, the matter left

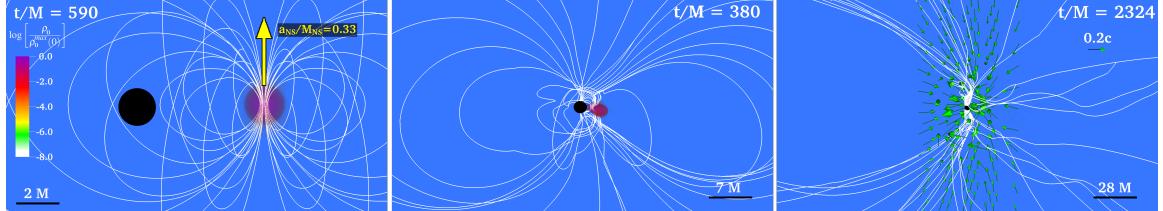
over post-merger heavily depends on (for a summary, see Fig. 1.6):

- **The mass ratio of the system.** This is defined as  $q = M_{BH}/M_{NS}$  so that  $q > 1$  always. Fully general relativistic, magnetohydrodynamic simulations (such as Ruiz et al. 2020) show that in cases where the mass ratio is 3:1, regardless of the neutron spin, a collimated outflow is observed, whereas the same is not realised in cases where the mass ratio is 5:1 or higher.
- **The spin of the components of the system.** In geometrized units (where  $G = c = 1$ ), these are prescribed in terms of  $a_{BH}/M_{BH}$  or  $a_{NS}/M_{NS}$ , and whether these two spins align (prograde) or are anti-aligned (retrograde) decides whether the neutron star would be tidally disrupted, and hence participate in the processes mentioned previously, or not. Via simulations, it is seen that the more the prograde spin of the neutron star, the farther out the neutron star is tidally disrupted, albeit this is only observed for the case of  $q = 3:1$  (comparing say, Figs. 1.4 and 1.5). Also, this leads to long tidal tails, which produces a baryon-loaded environment and thus the magnetic field of the tidally disrupted matter must overcome the baryon ram pressure to launch the jet. This process hence delays the launching of the jet.

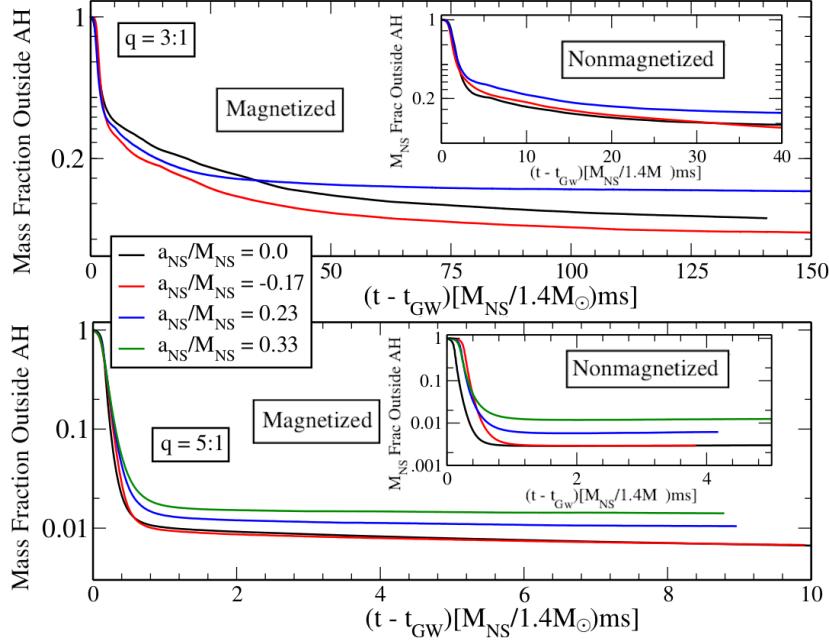
Aside from the SGRB jet, which requires magnetic field amplification (via MRI) as well as thermal pair production (from the disk remnant) followed by the Blandford-Znajek process, there is a possibility that NSBH mergers can produce kilonovae signatures. For this, the dynamically ejected mass has to be between  $10^{-4.5} - 10^{-2}(M_{NS}/1.4M_{\odot})M_{\odot}$  (see Ruiz et al. 2020 for more details), and this will lead to kilonovae potentially detectable by the Large Synoptic Survey Telescope (LSST).



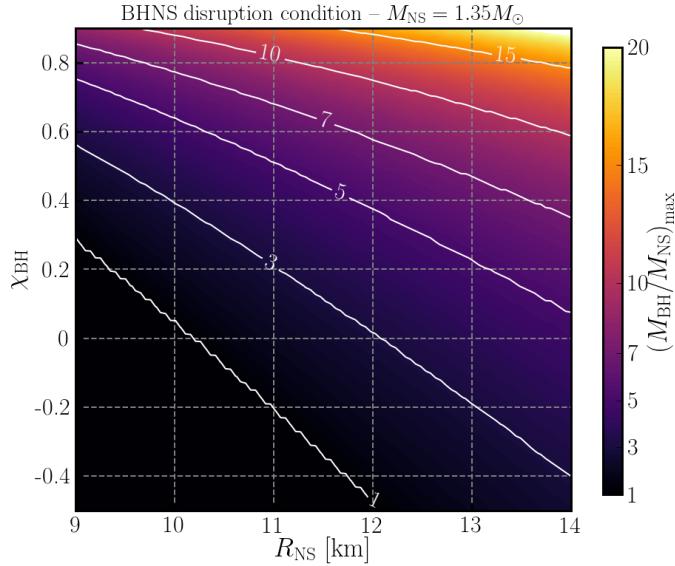
**Figure 1.4:** Volume rendering of the rest mass density ( $\rho_0$ ) (in log scale), normalized to the NS maximum value  $\rho_0 = 8.92 \times 10^{14} (1.4M_\odot/M_{NS})^2 \text{ g/cm}^3$ , for particular times for a magnetized neutron star, with  $q = 3:1$  and a prograde NS spin of 0.23. Top three panels highlight the inspiral and the tidal disruption, whereas the bottom three panels highlight the appearance of the magnetically-driven jet. White lines denote the magnetic field, arrows denote the fluid velocity and the BH's apparent horizon is the black sphere. Here  $M = 2.5 \times 10^{-2} (M_{NS}/M_{1.4M_\odot}) \text{ ms} = 7.58 (M_{NS}/M_{1.4M_\odot}) \text{ km}$  (in geometrized units). From Ruiz et al. 2020.



**Figure 1.5:** Similar to Fig. 1.4, however with the NS spin being 0.33, the BH spin being 0 and  $q = 5:1$ . In this case, no strong collimation of the magnetic field is observed from the merger remnant, and so a magnetically-driven jet is also not observed. From Ruiz et al. 2020.



**Figure 1.6:** Fraction of rest-mass of the NS outside the apparent horizon of the black hole as a function of coordinate time, for the various configurations considered in Ruiz et al. 2020. The inset figures report the same for non-magnetized cases, and the coordinate time is shifted such that the merger time coincides with 0.



**Figure 1.7:** Maximum value of the mass-ratio ( $M_{BH}/M_{NS}$ ) for which a NSBH system disrupts, as a function of the neutron star radius  $R_{NS}$ , and the aligned component of the dimensionless black hole spin  $\chi_{BH}$ , assuming  $M_{NS} = 1.35M_\odot$ . Results for other neutron star masses can be obtained by rescaling considering the disruption condition at constant compaction  $C_{NS} = GM_{NS}/R_{NS}c^2$ . From Francois Foucart 2020.

### 1.2.1 Modelling outflows from NSBH Mergers

As mentioned before, the outflows from NSBH mergers depend on the amount of mass left outside the event horizon post-merger. In the work done by Francois Foucart, Hinderer, and Nissanke 2018, the authors consider a suite of 75 numerical relativity simulations of NSBH mergers over the parameter space  $\mathcal{Q} \in [1, 7]$ ,  $\chi_{BH} \in [-0.5, 0.97]$ ,  $\mathcal{C}_{NS} \in [0.13, 0.182]^3$  and fit the results for the *remnant mass*  $M_{\text{rem}}$  (sometimes also denoted as  $M_{\text{out}}$ ), as a function of the binary parameters (masses, spins of the components, tidal deformability of the neutron star etc). This fit is given as follows:

$$M_{\text{out}} = M_{NS}^b \cdot \max \left( \alpha \frac{1 - 2\rho}{\eta^{1/3}} - \beta \hat{R}_{\text{ISCO}} \frac{\rho}{\eta} + \gamma, 0 \right)^\delta \quad (1.6)$$

where:

- The baryonic mass of the neutron star is given by the equation

$$M_{NS}^b = M_{NS} \left( 1 + \frac{0.6C_{NS}}{1 + 0.5C_{NS}} \right)$$

- The tidal deformability of the neutron star is given by  $\Lambda_{NS}$  and  $\rho = (15\Lambda_{NS})^{-1/5}$ . It is also related to the compactness of the neutron star via the C-Love relation (see Yagi and Yunes 2017):

$$C_{NS} = \sum_{k=0}^2 a_k (\ln \Lambda_{NS})^k \quad (1.7)$$

where  $a_0 = 0.360$ ,  $a_1 = -0.0335$ ,  $a_2 = 0.000705$ .

- $\eta$  is the symmetric mass ratio, given by  $\eta = \frac{\mathcal{Q}}{(1 + \mathcal{Q})^2}$ .
- $\hat{R}_{\text{ISCO}} = c^2 R_{\text{ISCO}} / GM_{BH}$  is the normalized ISCO radius for a spinning black hole, given in Bardeen, Press, and Teukolsky 1972 as :

$$\begin{aligned} \hat{R}_{\text{ISCO}} &= 3 + Z_2 - \text{sgn}(\chi_{BH}) \sqrt{(3 - Z_1)(3 + Z_1 + 2Z_2)} \\ \hookrightarrow Z_1 &= 1 + (1 - \chi_{BH}^2)^{1/3} [(1 + \chi_{BH})^{1/3} + (1 - \chi_{BH})^{1/3}] \\ \hookrightarrow Z_2 &= \sqrt{3\chi_{BH}^2 + Z_1^2} \end{aligned} \quad (1.8)$$

---

<sup>3</sup>Here,  $\mathcal{Q} = M_{BH}/M_{NS}$  is the mass-ratio,  $\chi_{BH} = c|\mathbf{S}|/GM_{BH}^2$  is the effective spin of the black hole and  $C_{NS} = GM_{NS}/R_{NS}c^2$  is the compactness of the neutron star.

- $(\alpha, \beta, \gamma, \delta) \equiv (0.308, 0.124, 0.283, 1.536)$  are the fit coefficients.

Similarly, Kawaguchi et al. 2016 fit to the results of 45 numerical relativity simulations over the parameter space  $\mathcal{Q} \in [3, 7]$ ,  $\chi_{\text{BH}} \in [0, 0.90]$ ,  $C_{\text{NS}} \in [0.138, 0.180]$ . This fit produces a formula for the *dynamic* mass  $M_{\text{dyn}}$ , which is the unbound mass ejected at the time of disruption, in terms of the binary parameters. This fit is given as follows:

$$\frac{M_{\text{dyn}}}{M_{\text{NS}}^b} = \max \left\{ a_1 Q^{n_1} (1 - 2C_{\text{NS}}) C_{\text{NS}}^{-1} - a_2 Q^{n_2} \hat{R}_{\text{ISCO}}(\chi_{\text{BH}}) + a_3 \left( 1 - \frac{M_{\text{NS}}}{M_{\text{NS}}^b} \right) + a_4, 0 \right\} \quad (1.9)$$

where the symbols have their usual meanings, and additionally:

$$\begin{aligned} a_1 &= 4.464 \times 10^{-2} & a_2 &= 2.269 \times 10^{-3} \\ a_3 &= 2.431 & a_4 &= -0.4159 \\ n_1 &= 0.2497 & n_2 &= 1.352 \end{aligned}$$

From these two quantities, we can derive the disc mass  $M_{\text{disc}}$  as :

$$M_{\text{disc}} = \max\{M_{\text{out}} - M_{\text{dyn}}, 0\} \quad (1.10)$$

However, due to the fact that these two fits are derived from simulations over different regions of the input parameter space, care must be taken while applying them together. This is to ensure that  $M_{\text{dyn}} \leq M_{\text{out}}$  always, so that the disc mass is non-negative. This validation is performed by considering the ratio  $M_{\text{out}}/M_{\text{dyn}}$ . Another constraint is imposed, which is motivated by the fact that NSBH simulations carried out by F. Foucart et al. 2019 in the near-equal mass ratio regime found an unbound component no more massive than roughly 30% of the total remnant mass (note that one expects maximal tidal disruption in this regime, given a fast spinning black hole). Thus we set:

$$M_{\text{dyn,max}} = f \cdot M_{\text{rem}} = 0.3 \cdot M_{\text{rem}} \quad (1.11)$$

Additionally, the masses of the other wind ejecta, namely the neutrino-driven and

viscosity-driven wind ejecta, are derived from that of the disc mass:

$$\begin{aligned} M_{\text{vis}} &= \xi_{\text{vis}} M_{\text{disc}} = 0.2 M_{\text{disc}} \\ M_{\nu} &= \xi_{\nu} M_{\text{disc}} = 0.01 M_{\text{disc}} \end{aligned} \quad (1.12)$$

To model the SGRB jet, the procedure of Zhu et al. 2020 is followed. The kinetic energy of the jet is decided by the disc mass and the black hole spin as follows:

$$E_{\text{K,jet}} = \epsilon(1 - \xi_{\text{vis}} - \xi_{\nu})M_{\text{disc}}c^2\Omega_H^2f(\Omega_H) \quad (1.13)$$

where:

- The dimensionless angular velocity at the horizon of the black hole is given by:

$$\Omega_H = \frac{\chi_{BH}}{2(1 + \sqrt{1 + \chi_{BH}^2})} \quad (1.14)$$

- $f(\Omega_H)$  is a high-spin correction factor given by:

$$f(\Omega_H) = 1 + 1.38\Omega_H^2 - 9.2\Omega_H^4 \quad (1.15)$$

- $\epsilon$  is a fudge factor which depends on the large-scale geometry of the magnetic field, disc aspect ratio and the ratio of the magnetic field energy density to disc pressure at saturation.

In order to set it to a definite value, it is noted that the maximum disc mass cannot exceed the total NS baryonic mass i.e.  $M_{\text{disc}} \lesssim 2M_{\odot}$ . Also, the spin-dependent factor  $\Omega_H^2f(\Omega_H)$  cannot exceed 0.2 (since  $\chi_{BH} \in [-1, 1]$ ). Furthermore, the most energetic of SGRBs has had a  $E_{\gamma,\text{iso}} \sim 7.4 \times 10^{52}$  erg, and if one assumes a 10% conversion efficiency of kinetic to gamma-ray energy along-with a typical jet opening angle of  $5^\circ$ , this corresponds to a kinetic energy of  $E_{\text{K,jet}} \sim 3 \times 10^{52}$  erg.

Based on this, once can calculate  $\boxed{\epsilon \approx 0.015}$ .

Using this we can define the structure of the SGRB jet, given by the following equations:

$$\frac{dE(\theta)}{d\Omega} = \frac{E_{\text{k,jet}}}{\pi\theta_{\text{c,E}}^2} e^{-(\theta/\theta_{\text{c,E}})^2} \quad (1.16)$$

$$\Gamma(\theta) = (\Gamma_c - 1)e^{-(\theta/\theta_{c,E})^2} + 1 \quad (1.17)$$

$$E_{\text{iso}}(\theta_v) = \eta \int \frac{\delta^3}{\Gamma} \frac{dE}{d\Omega} d\Omega \quad (1.18)$$

where:

- $\Gamma_c = 100$ ,  $\theta_{c,E} = 0.1$ ,  $\theta_{c,\Gamma} = 0.2$ . See Salafia et al. 2015 and Barbieri et al. 2019.
- $\eta$  is the conversion efficiency of gamma-ray energy to kinetic energy, which is traditionally taken to be 10%.
- $\delta$  is the Doppler factor, given by

$$\delta = \frac{1}{\Gamma[1 - \beta \cos \alpha_v]}$$

where  $\alpha_v$  is the angle between the jet element at  $(\theta, \phi)$  and the observer's direction.

### 1.3 NS mergers in GW regime

Consider any astrophysical source emitting gravitational waves, which come in two polarizations, namely the *plus*  $h_+(t; \Theta_{GW})$  and the *cross*  $h_\times(t; \Theta_{GW})$  polarizations. Here  $\Theta_{GW}$  is the parameter vector, and is typically  $\{m_1, m_2, \chi_1, \chi_2, D_L, \iota, t_c, \phi_c\}$  which are the component masses, component spins, the binary's luminosity distance, the inclination angle of the orbital plane with respect to the line of sight, and two constants of integration: the time and phase of coalescence.

A detector's response is recorded as the GW strain such waveforms produce, but in the frequency domain and so the input waveforms are Fourier transformed before processing. Also the detector's antenna patterns (sensitivity as a function of the source location on the sky) and location phase factor (effect of the earth's rotation) play a role in the response. Thus, the detector response to these gravitational waves is of the form:

$$H(f; \Theta) = F_{lp}(f; \alpha, \delta) \cdot [H_+(f; \Theta_{GW})F_+(f; \alpha, \delta, \psi) + H_\times(f; \Theta_{GW})F_\times(f; \alpha, \delta, \psi)] \quad (1.19)$$

where  $F_{lp}$  is the location phase factor as a function of the frequency and the source RA and declination,  $H_{+/\times}$  are the frequency domain waveforms, and  $F_{+/\times}$  are the detector antenna patterns for each polarization. Also  $\Theta = \{\Theta_{GW}, \alpha, \delta, \psi\}$ .

The sensitivity of a detector is given by the detector's noise  $n(t)$  and its autocorrelation  $\kappa = \overline{n(t_1)n(t_2)}$ . Usually, the noise is assumed to be stationary, zero-mean and Gaussian. Thus, one can define the one-sided power spectral density  $S_n(f)$  as the Fourier transform of the autocorrelation.

From this, one can define the ‘overlap’ between two GW signals (for eg.: the detector responses for two different waveforms) using the noise-weighted scalar product:

$$\langle H, G \rangle = 2 \int_0^\infty \frac{H(f)G^*(f) + H^*(f)G(f)}{S_n(f)} df \quad (1.20)$$

And using this definition of the scalar product, the signal-to-noise ratio is defined as :

$$\rho^2 = \langle H, H \rangle = 4 \int_0^\infty \frac{|H(f)|^2}{S_n(f)} df \quad (1.21)$$

Now, since the noise  $n(t) = s(t) - h(t)$  is assumed to a zero-mean Gaussian, its Fourier transform also behaves the same way, and thus the probability of noise can be written down as :

$$p(\Theta) = p^0(\Theta) e^{-\frac{1}{2} \langle S - H(\Theta), S - H(\Theta) \rangle} \quad (1.22)$$

where  $p^0$  is the prior on the parameter vector of the detector response. Assuming that an event signal  $S$  has a high SNR, the value of  $\Theta$  at peak probability is a good estimate of the true value  $\Theta^*$ . Additionally peak probability occurs when the exponential  $E = \langle S - H, S - H \rangle$  is the largest. Expanding it around the maximum value:

$$E(\Theta) = E(\Theta^*) + \frac{1}{2} \left. \frac{\partial^2 E(\Theta)}{\partial \Theta_i \partial \Theta_j} \right|_{\Theta=\Theta^*} \Delta \Theta_i \Delta \Theta_j + \dots \quad (1.23)$$

where  $\Delta \Theta_i = \Theta_i - \Theta_i^*$ . The Hessian given by:

$$\frac{\partial^2 E(\Theta)}{\partial \Theta_i \partial \Theta_j} = 2 \langle \partial_{\Theta_i} H(\Theta), \partial_{\Theta_j} H(\Theta) \rangle + \langle \partial_{\Theta_i} \partial_{\Theta_j} H(\Theta), N \rangle \quad (1.24)$$

can be simplified for large SNR, where second-order differentials become negligible. This leads to the definition of the Fisher Information Matrix  $\Gamma$ :

$$\Gamma_{ij} = \langle \partial_{\Theta_i} H(\Theta), \partial_{\Theta_j} H(\Theta) \rangle \quad (1.25)$$

And hence, Eqn. 1.22 becomes:

$$p(\Theta) \sim \exp\left(-\frac{1}{2}\Gamma_{ij}\Delta\Theta_i\Delta\Theta_j\right) \quad (1.26)$$

which implies that the assumption of Gaussian noise helps associate the FIM to the inverse of the covariance matrix  $\Sigma \equiv \Gamma^{-1}$ . This also means that the diagonal and off-diagonal elements of  $\Gamma^{-1}$  denote the variances and covariances of the parameters, respectively, with  $1\sigma$  estimates of the error are given as  $\sigma_{\Theta_i} = \sqrt{\Sigma_{ii}}$ .

This Fisher Information Matrix (FIM) formalism<sup>4</sup> is a method of rapid GW data analysis, which approaches the accuracy of traditional Bayesian parameter estimation for events with large SNR.

**GWBENCH** (see Borhanian 2020) is a GW benchmarking tool, which can compute the Fisher matrix for a particular NS merger, given the network configuration and binary parameters. In this way, it enables rapid calculations to benchmark detector upgrades as well as forecast the confidence with which parameters may be estimated for the NS mergers in question.

## 1.4 Summary

NS mergers can present an ideal testing environment for physical theories under extreme gravity, and by observing them in both the EM and GW windows, current theories can be better understood and refined. Several questions also remain about the exact mechanisms which power the outflows from these NS mergers.

In this report, we focus mainly on SGRB jets from NSBH mergers and perform population synthesis studies to infer the conditions for and implications of observing a SGRB jet from NSBH mergers.

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<sup>4</sup>Sometimes also referred to as the Fisher Information Formalism (FIF), in which case the Fisher information is not a matrix but is instead the variance of the partial derivative with respect to the parameter vector, of the natural logarithm of the likelihood function for the random variable whose parameters are to be estimated.

# Chapter 2

## Population Synthesis

In order to derive meaningful conclusions about the outflows from NS mergers, given the models described in 1, it was necessary to carry out population synthesis studies. In these, population models which are physically or observationally motivated are taken from the literature and used to compute the statistical properties of the outflows from NS mergers created with these properties. If no confident or relevant model exists in literature, physically motivated ansätze are used (see for example, §§2.1.2).

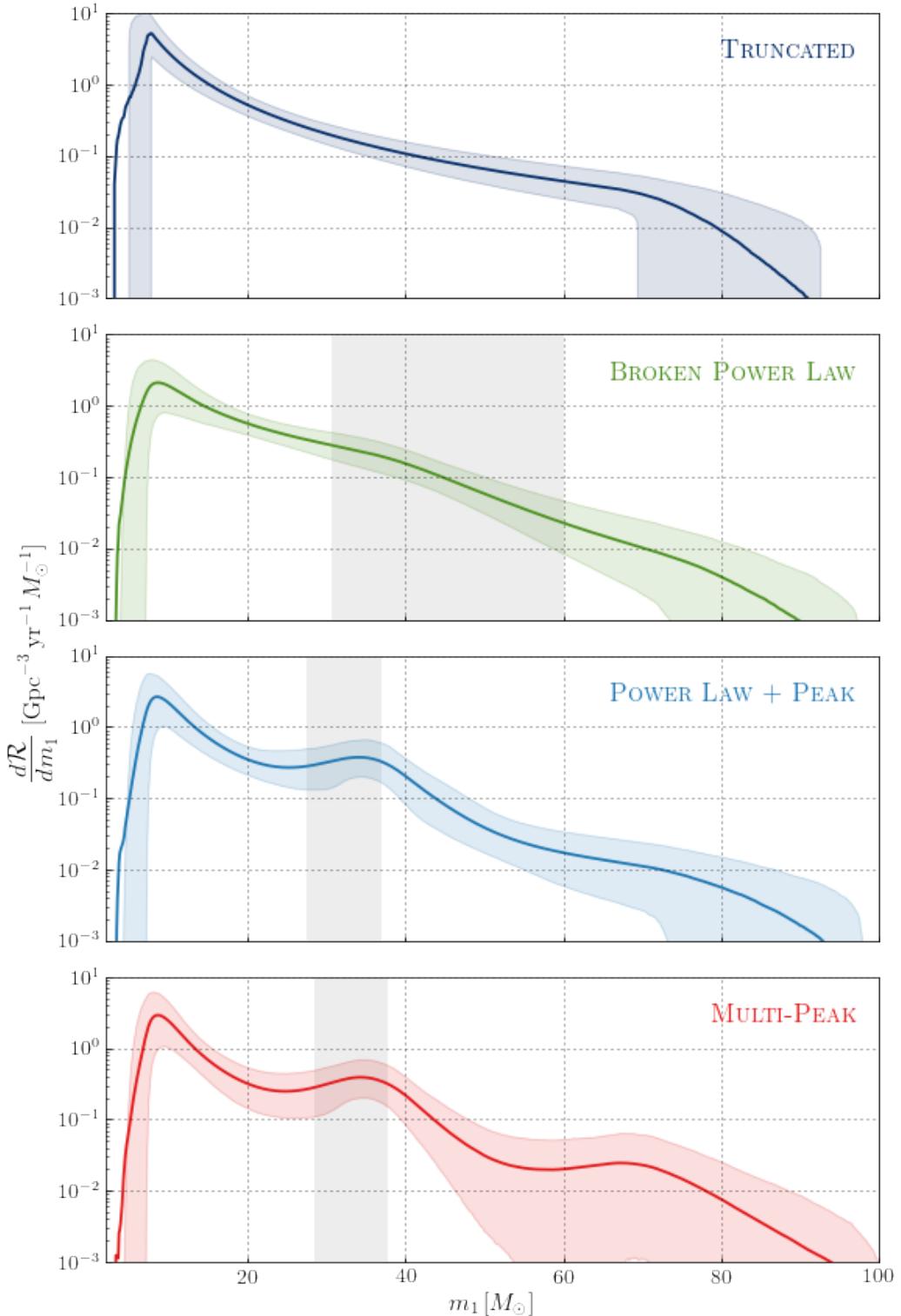
In each study which uses (slightly) different population models,  $10^5$  samples were drawn from the relevant parameter distributions. Using these as inputs the expected outflows were calculated using the fit formulae described in 1.

In this chapter, the various population models are briefly described along with their pertinence. Furthermore, preliminary checks for the population synthesis code are also discussed, which help verify the consistency of the code with theoretical results.

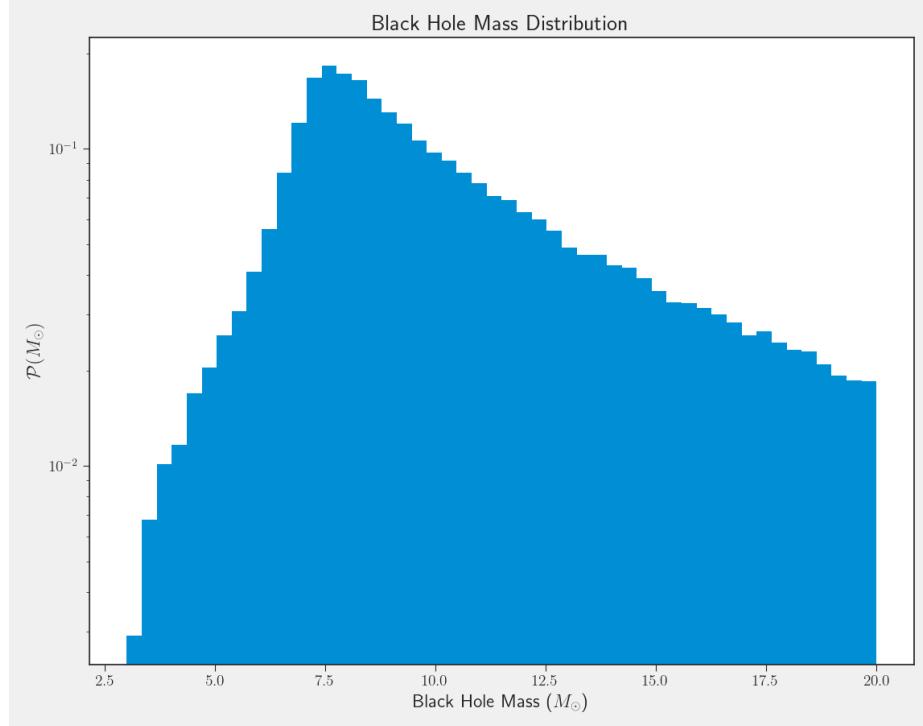
### 2.1 Black Hole Population Models

#### 2.1.1 Mass, $M_{BH}$

The masses of the black holes was sampled from the TRUNCATED mass distribution from The LIGO Scientific Collaboration et al. 2020. The distribution ‘produces’ black holes with masses between  $3\text{--}100 M_\odot$  (as can be seen from Fig. 2.1). However, NS binaries with extremely massive ( $M_{BH} > 20M_\odot$ ) black holes will not produce any appreciable EM emission due to the NS companion plunging into the black hole directly without significant tidal disruption. For this reason the population synthesis code imposes an upper limit to the black hole masses sampled.



**Figure 2.1:** Black Hole Mass Distributions from The LIGO Scientific Collaboration et al. 2020. In the current study the TRUNCATED mass distribution with an upper limit at  $20 M_\odot$ , since more massive black holes would not produce significant EM emission when in a merging NSBH binary.

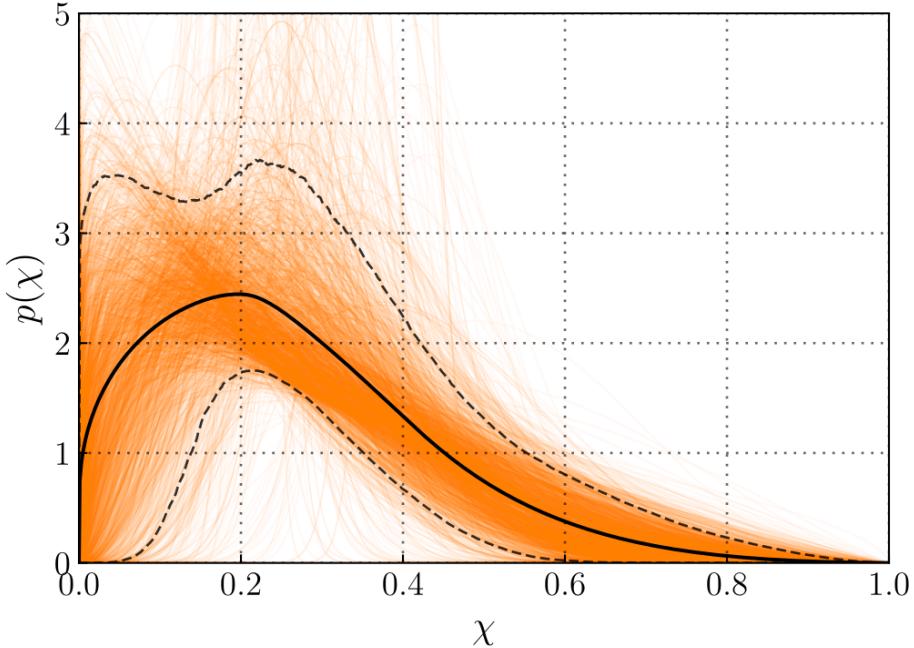


**Figure 2.2:** Black hole mass distribution as used in the current report, with an upper limit of  $20 M_{\odot}$ .

### 2.1.2 Spin, $\chi_{BH}$

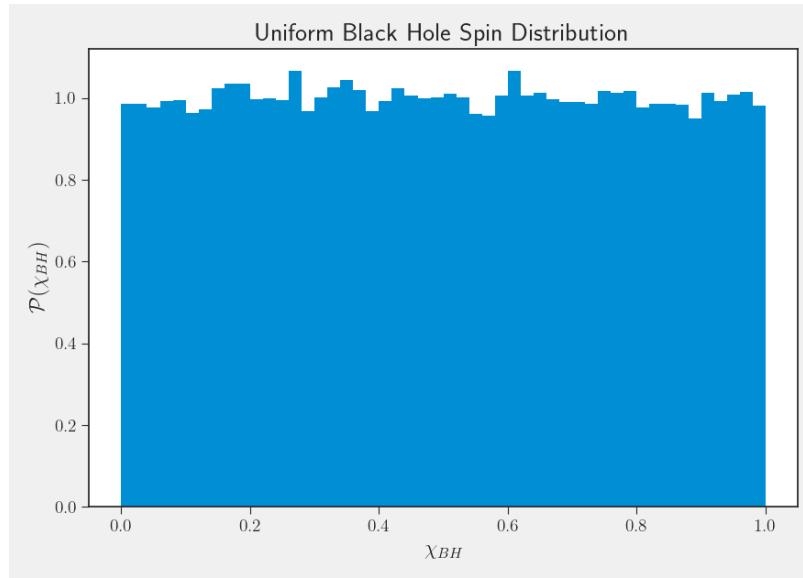
The spins of the black holes are sampled from the DEFAULT distribution given in The LIGO Scientific Collaboration et al. 2020. This distribution is essentially a Beta distribution , which ensures that the spin parameters sampled from this distribution remain within  $[0, 1]$ .

However, since there have been no confident NSBH merger detections in the GW regime, this distribution is largely derived from looking at BBH mergers, and so may not represent the true spin distribution of a black hole in a NSBH binary. To circumvent this, the following ansätze are used to probe the effect of the spin distribution on the EM outflows:



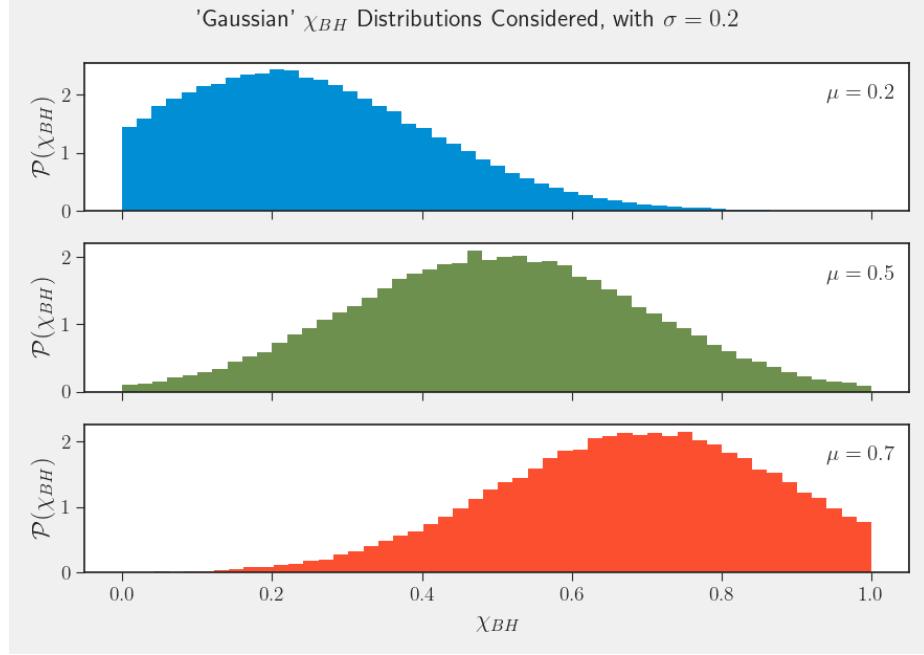
**Figure 2.3:** Beta distribution for the black hole spin from The LIGO Scientific Collaboration et al. 2020. Here the light traces are samples from the posterior distribution, whereas the solid black line is the posterior probability distribution for  $\chi_{BH}$ . Dashed black lines mark the 90% quantiles for the same.

- Uniform spin distribution: here  $\chi_{BH} \sim \mathcal{U}(0, 1)$ . Note also that this distribution will have a higher number of high spin samples as compared to the DEFAULT distribution considered above.



**Figure 2.4:** A realisation of the uniform black hole spin distribution considered here.

- Gaussian spin distribution: here  $\chi_{BH} \sim \mathcal{N}(\mu, \sigma)$ , but samples outside of  $[0, 1]$  are not considered. For simplicity and to cover a representative sample of the possible distributions, samples were taken from  $\mathcal{N}(0.2, 0.2)$ ,  $\mathcal{N}(0.5, 0.2)$  and  $\mathcal{N}(0.7, 0.2)$ . These represent distributions concentrated around low, medium and high spins respectively.



**Figure 2.5:** Realisations of the various ‘Gaussian’ black hole spin distributions considered here.

## 2.2 Neutron Star Population Models

In order to reduce the number of variables in the problem, the masses of all neutron stars in the population was set to  $1.4 M_\odot$ . This value corresponds to the median neutron star mass as inferred from GW170817 (see B. P. Abbott et al. 2018). Also, the spins of all neutron stars was set to 0, since it is assumed that sufficient amount of time would have passed between the formation of the binary and merger, allowing for the neutron star to spin down such that  $\chi_{NS} \sim 0$ .

Additionally, the tidal deformability of the neutron stars was set using the C-love relation (see Yagi and Yunes 2017). First, the radii of the neutron stars were set to 11 km, which is the median neutron star radius inferred for GW170817. Then, the

compactness of the neutron stars,  $C_{NS}$  was computed using the relation :

$$C_{NS} = \frac{GM_{NS}}{R_{NS}c^2} \quad (2.1)$$

where  $M_{NS}$ ,  $R_{NS}$  are the mass and radius of the neutron star, G is the universal gravitational constant and c is the speed of light. Finally, the tidal deformability is computed by solving the C-Love relation :

$$C_{NS} = \sum_{k=0}^2 a_k (\ln \Lambda_{NS})^k \quad (2.2)$$

where  $\Lambda_{NS}$  is the tidal deformability of the neutron star, and  $a_k$  are the fit coefficients as given in Yagi and Yunes 2017.

## 2.3 Spatial Distribution and Orientation of events

### 2.3.1 Constant comoving volume distribution

The events whose component mass and spin distributions were described in 2.1-2.2 are distributed in 3D space such that their number density is constant in comoving volume.

For this, firstly the latitudinal ( $\theta$ ) and longitudinal ( $\phi$ ) angles are sampled such that  $\cos \theta \sim \mathcal{U}(-1, 1)$  and  $\phi \in \mathcal{U}(0, 2\pi)$ , i.e. they are sampled such that they are uniform on a unit sphere. As for the comoving distance distribution  $\mathcal{P}(D_c)$ , consider the probability of an event lying in an infinitesimal shell of width  $dD_c$  at a comoving distance  $D_c$ . Then it can be seen that:

$$\mathcal{P}(D_c)dD_c \propto D_c^2 dD_c \Rightarrow \boxed{\mathcal{P}(D_c) = \alpha D_c^2} \quad (2.3)$$

where  $\alpha$  is the constant of proportionality. In the local universe, it can be safely assumed that  $D_c \approx D_L$ , where the latter is the luminosity distance. However, from the fact that the GW SNR  $\rho \propto D_L^{-1}$  it can also be seen that:

$$\mathcal{P}(\rho) = \mathcal{P}(D_L) \left| \frac{dD_L}{d\rho} \right| \quad (2.4)$$

$$= \mathcal{P}(D_c) \left| \frac{dD_c}{d\rho} \right| \quad (2.5)$$

$$\propto \frac{1}{\rho^2} \frac{1}{\rho^2} = \frac{1}{\rho^4}$$

$$\Rightarrow \mathcal{P}(\rho) = \frac{\beta}{\rho^4} \quad (2.6)$$

Once the SNR detection threshold<sup>1</sup> ( $\rho_{th}$ ) is set, the normalization constants can be computed as follows:

$$\int_0^{D_{c,max}} \mathcal{P}(D_c) dD_c = \int_0^{D_{L,max}} \mathcal{P}(D_L) dD_L \quad (2.7)$$

$$= \int_{\infty}^{\rho_{th}} \mathcal{P}(\rho) d\rho \quad (2.8)$$

$$= 1 \quad (2.9)$$

This gives:

$$\mathcal{P}(D_L) = 3 \frac{D_L^2}{D_{L,max}} \Leftrightarrow \mathcal{P}(\rho) = 3 \frac{\rho_{th}^3}{\rho^4} \quad (2.10)$$

where  $D_{L,max}$  is the luminosity distance corresponding to the detection threshold. As an example, for the Advanced LIGO/VIRGO configuration and a SNR threshold of 10,  $D_{L,max} \approx 1123$  Mpc.

Using Eq.2.10, samples are drawn from the luminosity distance distribution and using the previous samples for  $\theta$  and  $\phi$ , events are distributed in 3D space such that the number density of events is constant in the comoving volume.

### 2.3.2 Orientation of Events

The orientation of NSBH binaries with respect to the line-of-sight is prescribed using the inclination angle,  $\iota$  and the polarization angle of the incoming GW signal,  $\psi$ .

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<sup>1</sup>This is defined such that any event with a GW network SNR greater than the detection threshold will be considered a detection.

These are distributed for the population such that  $\cos \iota \sim \mathcal{U}(-1, 1)$  and  $\psi \sim \mathcal{U}(0, 2\pi)$ , from which samples are drawn for individual events.

## 2.4 Validation of Population Synthesis Code

With the population models for the binary component parameters as described in the previous section,  $10^5$  samples are drawn from each model to generate a population. These populations will majorly differ in the black hole spin distribution and are thus distinguished by that very factor.

Given these populations, several checks were carried out to confirm the consistency of the underlying models and the population generated, with the use of (both EM and GW) theoretical results. The rationale behind these checks and the results are given in the following sections.

### 2.4.1 GW Checks

Firstly, for the samples generated, the optimal GW SNR corresponding to the binary parameter combinations was computed. This was done by using the restricted post-Newtonian (PN) waveform (RWF; see Cutler and Flanagan 1994) as the GW template for the expected GW strain signal from an actual merger event. For this case, the template is given using the frequency domain representation:

$$\tilde{h}(f) = \mathcal{A} f^{-7/6} e^{i\phi(f)} \quad (2.11)$$

where  $\mathcal{A} = \sqrt{5/96}\pi^{-2/3}\mathcal{M}$  is the amplitude,  $\mathcal{M} = M\eta^{5/3}$  is the chirp mass,  $M$  is the total mass,  $\eta = \frac{m_1 m_2}{M}$  is the symmetric mass ratio, and  $\phi(f)$  is the frequency domain GW phase function. With the RWF template, the PN approximations made ensure maximum accuracy with respect to the phase but ignore the PN corrections to the amplitude of the gravitational waveform.

For such a template, the optimal SNR at a detector with a noise power spectral density (PSD)  $S_h(f)$ , is computed as:

$$\rho = \sqrt{4 \int_0^\infty \frac{|\tilde{h}(f)|^2}{S_h(f)} df} \quad (2.12)$$

Substituting the form of  $\tilde{h}(f)$  from Eq. 2.11, one obtains:

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$$\rho(m_1, m_2, D_L, \theta, \phi, \psi, \iota) = \sqrt{4 \frac{\mathcal{A}^2}{D_L^2} [F_+^2(\theta, \phi, \psi)(1 + \cos^2 \iota)^2 + 4F_\times^2(\theta, \phi, \psi) \cos^2 \iota]} I(M) \quad (2.13)$$

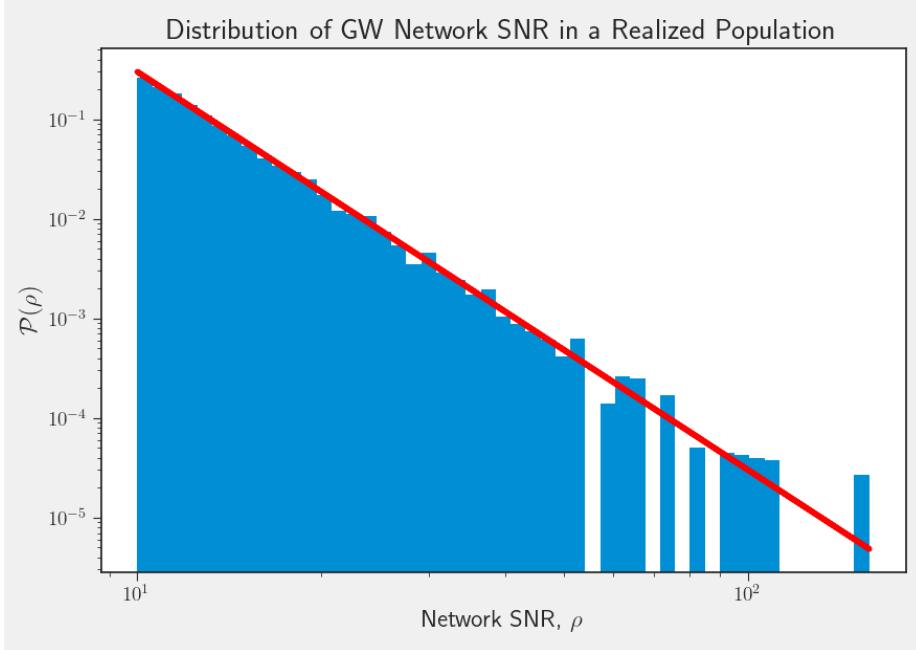
where  $F_{+, \times}(\theta, \phi, \psi)$  are the antenna pattern functions for the ‘plus’ and ‘cross’ GW polarizations, and the four angles  $\{\theta, \phi, \psi, \iota\}$  prescribe the location and orientation of the source with respect to the detector. Also,  $I(M)$  is the frequency integral given by:

$$I(M) = \int_0^\infty \frac{f^{-7/3}}{S_h(f)} df \quad (2.14)$$

$$\approx \int_{f_{low}}^{f_{LSO}} \frac{f^{-7/3}}{S_h(f)} df \quad (2.15)$$

where  $f_{low}$  corresponds to the seismic cutoff in the detector noise PSD curve, and  $f_{LSO}$  corresponds to the frequency at the last stable orbit, which under the PN approximation is  $f_{LSO} = (6^{3/2}\pi M)^{-1}$  for a given total mass  $M$ .

For each sample in the various populations, the optimal SNR was computed using the above formulae with the aid of the PyCBC (Nitz et al. 2021) and LALSimulation (LIGO Scientific Collaboration 2018) libraries. Using this, it was verified that the events *are* distributed in luminosity distance (and thus SNR) as expected from 2.10. This is given in Fig.2.6. Note that in order to verify this relation, one must only consider the events whose network SNR is above the threshold SNR chosen (in this case, 10).

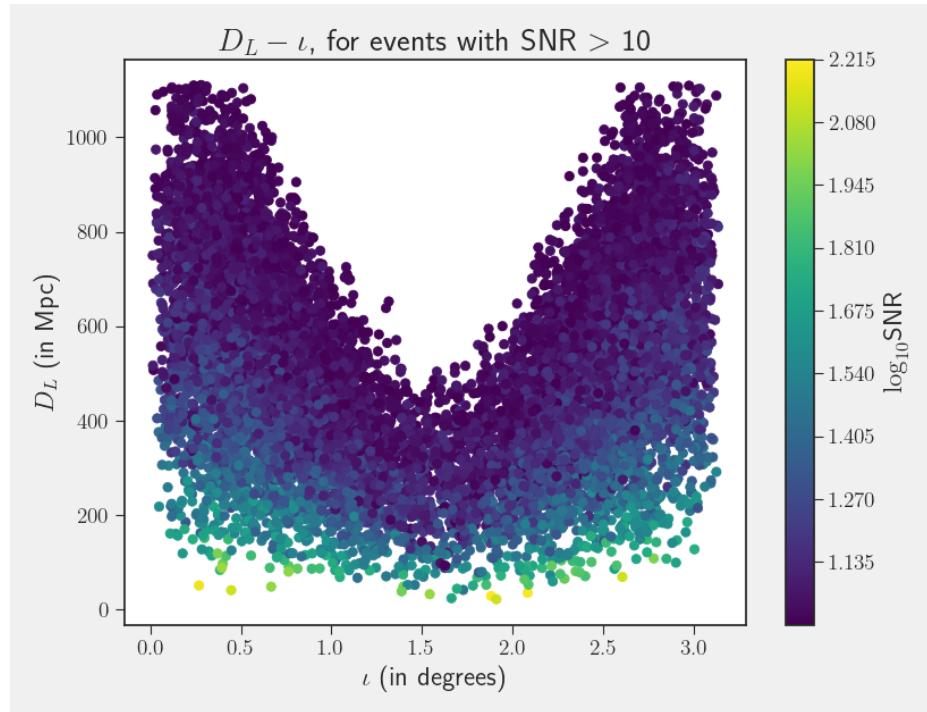


**Figure 2.6:** Distribution of GW network SNR, from a particular realisation of a population (specifically one with standard assumptions about binary parameters, except that the spin distribution is uniform). The red line is the curve given by theoretical considerations, which for a threshold SNR of 10, corresponds to  $f(\rho) = 3\rho_{th}^3/\rho^4 = 3000/\rho^4$ .

Furthermore, for these samples, the GW network SNR is higher in value if:

- The ‘event’ is closer by, i.e.  $D_L$  is low.
- The binary merges with the orbital plane face on, i.e.  $\iota$  is low
- Both  $D_L$  and  $\iota$  are low, although this is a small part of the  $D_L - \iota$  parameter space, and hence is less probable.

This implies a connection between the two parameters, and indeed this  $D_L - \iota$  connection is well studied in the literature (see Schutz 2011 & Seto 2015 for discussions on the same). This relation was also verified to hold within the population generated, which is shown in figure 2.7. Note that some parts of the  $D_L - \iota$  parameter space are not at all populated by the samples, which reaffirms the fact that distant, edge-on mergers are disfavoured when such systems are being explored via GW.



**Figure 2.7:** Plot of the luminosity distance v/s the inclination angle for events in the population, color coded by the SNR. Note that only events whose  $\text{SNR} > 10$  (the SNR threshold) are considered here. This shows that the generated population also exhibits the relation as expected from theory.

### 2.4.2 EM Checks

The only EM component which has to be validated is the code which imposes the jet structure (Eq. 1.16-1.18), once the energetics of the jet are computed via Eq. 1.13-1.15.

In order do this, one extracts the  $\theta_v$  samples from the generated population, which can be done using the relation  $\theta_v = \min(\iota, \pi - \iota)$ . Then, a common baseline is achieved by setting the value of  $E_{kin,jet} = 10^{49}$  erg for all samples, regardless of the actual value if the binary parameters were to be taken into consideration. This is necessary to remove the variation of the structure with the intrinsic jet energetics, which is intimately related to the binary parameters.

Then using this constraint, the value of  $E_{iso}(\theta_v)$  is computed using Eq. 1.16-1.18. It is seen from Fig. **fig:jet-struct** that jet structure is being imposed correctly and according to what is expected from theory.

## 2.5 Summary

# Chapter 3

## Population Analysis

Using the population models described in Chapter 2, one can derive the properties of NSBH mergers as seen in the GW as well as EM regimes. Specifically, an important aspect of NSBH mergers is what fraction of mergers actually produce a prompt component which will be detectable by present-day Gamma-ray detectors, such as INTEGRAL. A related aspect is the dependence of this fraction on the priors which go into creating these populations. Specifically, different black hole spin prior distributions should affect the number of detectable prompt emissions, and so this behaviour is worth investigating.

Furthermore, this population synthesis framework also helps to investigate the rate density of NSBH mergers which lead to observable SGRBs, which can then be compared to the number derived currently from purely EM observations.

### 3.1 Summary

## Chapter 4

# NS Merger Candidates in GWTC-2

### 4.1 About GWTC-2

The first half of the third observing run (O3a) of the LVC started on the 1st of April 2019, and went on till 30th of September, 2019. Following this, instrumental upgrades were made during the month of October and the second half of the third observing run (O3b) was started on the 1st of November. Due to the global pandemic, the observing run had to be prematurely suspended on 30th of March, 2020.

Starting with O3, alerts were distributed via the public alerts section of Gravitational-wave Candidate Event Database (GRACE-DB). If triggers were registered that passed the detection threshold of the LVC network during observational runs, online parameter estimation was done using the raw GW data, and low latency estimates of the rough sky position, component masses and luminosity distance was made available for observers in the EM regime. This allowed for rapid follow-up in various bands of the electromagnetic spectrum, and these observations were reported and cross-verified using NASA’s GRB Coordinates Network (GCN). Using the circulars reported in the GCN for NSBH/BNS events of interest, along with the low-latency information from GRACE-DB, O3a’s non-BBH candidate events have been collected in table 4.3.

More detailed analysis of these events in the months following has led to at least 53 events in the third observing run alone. These can be classified as:

- 37 Binary Black Hole (BBH) merger candidates.
- 7 BNS merger candidates. Of these, only 1 corresponds to O3a, which is the event GW190425 discussed in the chapter 4.2.
- 4 events in the mass gap, which are events with compact objects with masses of  $3\text{-}5 M_{\odot}$ .

- 5 NSBH merger candidates. Of these, only 1 has been confirmed officially, which is the event GW190814.

Of these, 26 events were officially confirmed and 13 new events were reported for the first time in R. Abbott et al. 2020, and it is from there that the posterior distributions of the various parameters (such as inclination angle  $\iota$ , luminosity distance  $D_L$  etc) are used for further analysis. Note also that there are several marginal events that have been reported, in that they have non-negligible probability distributed between two classifications. For example, the event GW190426\_152155 has significant probability split between it being a terrestrial event (58 %) and a NSBH/BNS/Mass Gap event (cumulatively 42%). Without a significant EM counterpart, this event cannot be confidently placed in either of the classes, and thus warrants further analysis. This is the subject of chapter 4.3.

GRACE-DB					
Superevent	$\mathcal{P}(\text{BBH})$	$\mathcal{P}(\text{BNS})$	$\mathcal{P}(\text{MassGap})$	$\mathcal{P}(\text{NSBH})$	$\mathcal{P}(\text{Terr})$
ID					
<u>S190426c</u>	0	24	12	6	58
S190910h	0	61	0	0	39
S200213t	0	63	0	0	37
S191213g	0	77	0	0	23
S190901ap	0	86	0	0	14
<u>S190425z</u>	0	>99	0	0	0
S190930s	0	0	95	0	5
S190923y	0	0	0	68	32
S190930t	0	0	0	74	26
S191205ah	0	0	0	93	7
S190910d	0	0	0	98	2
S190814bv	0	0	0	>99	0

**Table 4.1:** List of candidate merger events and the probability of classification (reported in %) for each non-BBH event reported during O3a. Here the GRACE-DB superevent ID is used (instead of the GWTC-2 event ID), since for a particular GW event, the superevent collects both the EM followup as well as GW trigger information within GRACE-DB. The probabilities are assigned using the process described in Kapadia et al. 2020, and are reported in GRACE-DB. The events underlined are discussed in more detail in later chapters.

<b>UID</b>	<b>FAR</b>	$D_L$ (Mpc)	Error in $D_L$ (Mpc)
<u>S190426c</u>	1 per 1.6276 yr	377	100
S190910h	1.1312 per yr	230	88
S200213t	1 per 1.7934 yr	201	80
S191213g	1.1197 per yr	201	81
S190901ap	1 per 4.5093 yr	241	79
<u>S190425z</u>	1 per 69834 yr	158	43
S190930s	1 per 10.534 yr	709	191
S190923y	1.5094 per yr	438	133
S190930t	1 per 2.0536 yr	108	38
S191205ah	1 per 2.5383 years	385	164
S190910d	1 per 8.5248 years	606	197
S190814bv	1 per 1.559e+25 years	241	26

**Table 4.2:** List of candidate merger events and the probability of classification for each non-BBH event reported during O3a. FAR refers to the False Alarm Rate (in number of events per year or equivalently in Hz<sup>-1</sup>), and  $D_L$  is the luminosity distance. Both these values are those reported in GRACE-DB corresponding to each event. The events underlined are discussed in more detail in later chapters.

<b>UID</b>	<b>FERMI-LAT</b>	<b>FERMI-GBM</b>	<b>SWIFT/BAT</b>	<b>INTEGRAL</b>
S190426c	24342	24248	24255	24242
S190910h	25742	25714	25718	25709
S200213t	27062	<b>27056</b>	<b>27058</b>	27050
S191213g	26412	26409	26410	26401
S190901ap	25625	25610	25617	25605
<u>S190425z</u>	24266	24185	24184, 24296	24169, 24170 24178, 24181
S190930s	25895	25886	25889	25872
S190923y	25834	25823	25846	25815, 25825
S190930t	25898	25887	25888	25880
S191205ah	26363	26359	26365	26531
S190910d	25717	25699	25704	25698
S190814bv	25385	25326	25341	25323

**Table 4.3:** List of candidate merger events and the probability of classification for each non-BBH event reported during O3a. Here the GCN Circular number reporting the findings of the particular instrument in the column heading is reported, corresponding to each event. The events underlined are discussed in more detail in later chapters. GCNs marked with red should be ignored during analysis since they correspond to times when the respective instruments were in the South Atlantic Anomaly (SAA).

## 4.2 Analysis of GW190425

GW190425 (or S190425z in GRACE-DB) is a GW trigger which was recorded by the LVC on 25th April, 2019 at 08:18 UTC. At the time of this trigger, the Hanford site of LIGO (H1) was undergoing maintenance, whereas the Livingston site (L1) of LIGO and the VIRGO detector were both operational. However, at the VIRGO detector this event was sub-threshold, effectively making this event a single-detector trigger. As a consequence, the LVC sky localisation area was wider than as compared to GW170817, and EM follow-up was constrained to serendipitous observations by

satellites which happened to be observing the same area of the sky, or to diminished coverage by satellites due to observational schedules.

Initial work was carried out to reproduce the results of M. Saleem et al. 2020, where the authors use a frequentist approach to discuss the possibility of a relativistic jet from the binary neutron star (BNS) merger event GW190425<sup>1</sup>. In M. Saleem et al. 2020, three key ideas are developed which are described in detail in the following sections.

#### 4.2.1 Constraints on the $D_L - \iota$ Posterior

At the time of writing of M. Saleem et al. 2020, only the low-latency information was made public. Consequently, information about the inclination angle of this event was not released. However, both the luminosity distance ( $D_L$ ) and inclination angle ( $\iota$ ) of the event are required for the electromagnetic analysis of the event<sup>2</sup>. To solve this issue, a known correlation between the luminosity distance  $D_L$  and  $\iota$  (see Schutz 2011, Seto 2015) was used to infer the distribution of  $\iota$  for this event, which was possible since the former was publicly known.

The following are the other publicly released information relevant to the problem, and can be used to constrain the  $D_L - \iota$  joint distribution:

- The posterior probability of the event being a BNS merger is  $> 99\%$ ,
- The event was observed by the LIGO Livingston (L1), and Virgo (V1) detectors, whereas the LIGO Hanford (H1) detector was not observing. However, at V1 the signal-to-noise ratio (SNR) was below the threshold and thus this event was a single detector trigger.
- The preliminary luminosity distance estimate is given by  $D_L = 155 \pm 41$  Mpc.

Using these inputs the  $D_L - \iota$  space is constrained as follows.

1. A population of BNS mergers is simulated, such that they are uniformly distributed in the comoving volume, with the inclination angle of the binaries

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<sup>1</sup>At the time of the writing of the paper, this event was still not confirmed as a bona-fide GW event, and so it was denoted by GRACE-DB as S190425z. Consequently, several key pieces of information (such as the luminosity distance and inclination angle posteriors) from the GW inference was not yet public, and had to be worked around. See text for more details.

<sup>2</sup>The former is used to calculate the fluence, and the latter is directly related to the viewing angle, which in turn decides the observed isotropic equivalent energy (see Eq. 4.4)

being such that  $\cos \iota \in [-1, 1]$ . This also means that the luminosity distance is initially distributed such that  $\mathcal{P}(D) \propto D^2$  upto some threshold distance. For the purposes of the simulations, this threshold distance is set to be the distance corresponding to the 99% percentile of a Gaussian with a mean of 155 Mpc and a standard deviation of 41 Mpc. In practice, it is the maximum distance up till which the comoving and the luminosity distances can be used interchangeably, which corresponds to a redshift of roughly 0.1.

2. The NS masses are uniformly distributed between  $1\text{-}2 M_{\odot}$ . This enforces the constraint that from the analysis of the GW waveform, the event has  $> 99\%$  probability of being a BNS merger.
3. Then, the optimal SNR is computed for each realisation using the restricted post-Newtonian (PN) waveform (RWF) (see Cutler and Flanagan 1994, Kastha, Saleem, and Arun 2020). Usually, the process of matched filtering is done to detect GW signals amongst background noise, where various template waveforms are cross-correlated with the observed data. These templates all correspond to mergers with various signal parameters such as masses, spins etc., and the template which maximizes the SNR is the optimal template, with the corresponding SNR being the optimal SNR, defined as :

$$\rho = \sqrt{4 \int_0^\infty \frac{|\tilde{h}(f)|^2}{S_h(f)} df} \quad (4.1)$$

where  $S_h(f)$  is the detector's power spectral density (PSD) and  $\tilde{h}(f)$  is the frequency domain gravitational waveform. For the RWF, this is given as  $\tilde{h}(f) = \mathcal{A}f^{-7/6}e^{i\psi(f)}$ , where  $\mathcal{A}$  is the amplitude and  $\psi(f)$  is the frequency domain GW phase. In this scheme, the PN corrections to the amplitude of the waveform are ignored but the phase is accurately accounted for, since GW parameter estimations are most sensitive to the phase of the waveform (Poisson and Will 1995). Using the RWF, the optimal SNR for compact binary coalescences can be written as :

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$$\rho(m_1, m_2, D_L, \theta, \phi, \psi, \iota) = \sqrt{4 \frac{\mathcal{A}^2}{D_L^2} [F_+^2(\theta, \phi, \psi)(1 + \cos^2 \iota)^2 + 4F_\times^2(\theta, \phi, \psi)\cos^2 \iota] I(M)} \quad (4.2)$$

where  $F_{+,\times}(\theta, \phi, \psi)$  are the antenna pattern functions for the 'plus' and 'cross'

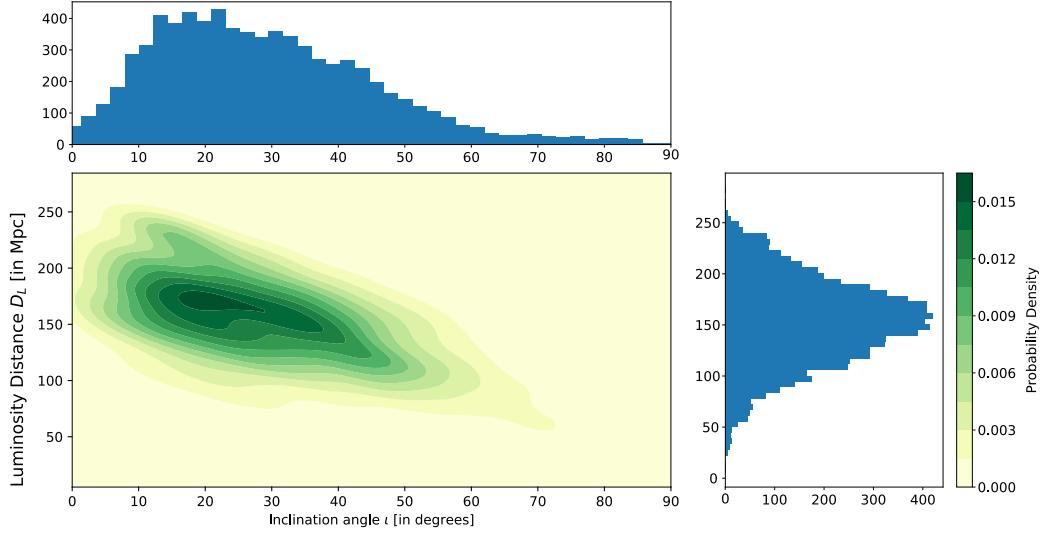
polarisations,  $\mathcal{A} = \sqrt{5/96}\pi^{-2/3}\mathcal{M}^{5/6}$ . Here,  $\mathcal{M}$  is the chirp mass, which is related to the total mass  $M$  by  $\mathcal{M} = M\eta^{3/5}$ , where  $\eta = \frac{m_1 m_2}{M^2}$  is the symmetric mass ratio of the system and  $m_1, m_2$  are the component masses. The four angles  $(\theta, \phi, \psi, \iota)$  describe the location and orientation of the source with respect to the detector.  $I(M)$  is the frequency integral defined as :

$$I(M) = \int_0^\infty \frac{f^{-7/3}}{S_h(f)} df \approx \int_{f_{low}}^{f_{LSO}} \frac{f^{-7/3}}{S_h(f)} \quad (4.3)$$

where,  $f_{low}$  is the lower seismic cut-off for the detectors and  $f_{LSO} = \frac{1}{6^{3/2}\pi M}$ , which is the GW frequency at the last stable orbit for a BBH merger with total mass  $M$ . To compute the optimal SNR in L1 and V1, the best reported O2 sensitivities were used as conservative O3 sensitivities, as an input for  $S_h(f)$ . As the trigger is an L1 single-detector trigger, the conditions that  $\text{SNR} < 4$  at V1 and the network  $\text{SNR} > 9$  are enforced. The former is motivated by the single-detector threshold of the GstLAL pipeline, whereas the latter is motivated by the fact that the network SNR of all O1/O2 events were  $> 9$ .

4. From the resulting population, a sub-population is extracted such that the luminosity distance follows a Gaussian distribution with a mean of 155 Mpc and a standard deviation of 41 Mpc. This is done so as to impose the constraint applied by the luminosity distance posterior distribution released by the LVC.

The resulting 2D distribution of  $D_L - \iota$  of this sub-population is shown in the figure below. This is used further on, as the prior for studying the possibility of a sGRB from GW190425.



**Figure 4.1:** Constraints on the  $D_L - i$  joint distribution, obtained from imposing the observed properties of S190425z/GW190425.

#### 4.2.2 Calculation of the Apparent structure

Assuming the intrinsic jet structure models described before (in Sec. 1.1), the apparent isotropic equivalent energy is calculated using the equation (as in Salafia et al. 2015 and Biscoveanu, Thrane, and Vitale 2020):

$$E_{iso}(\theta_v) = \frac{1}{2\pi} \int_0^{2\pi} d\phi \int_0^{\theta_{max}} d\theta \sin \theta \frac{\epsilon(\theta)}{\Gamma(\theta)^4 [1 - \beta(\theta) \cos \alpha_v]^3} \quad (4.4)$$

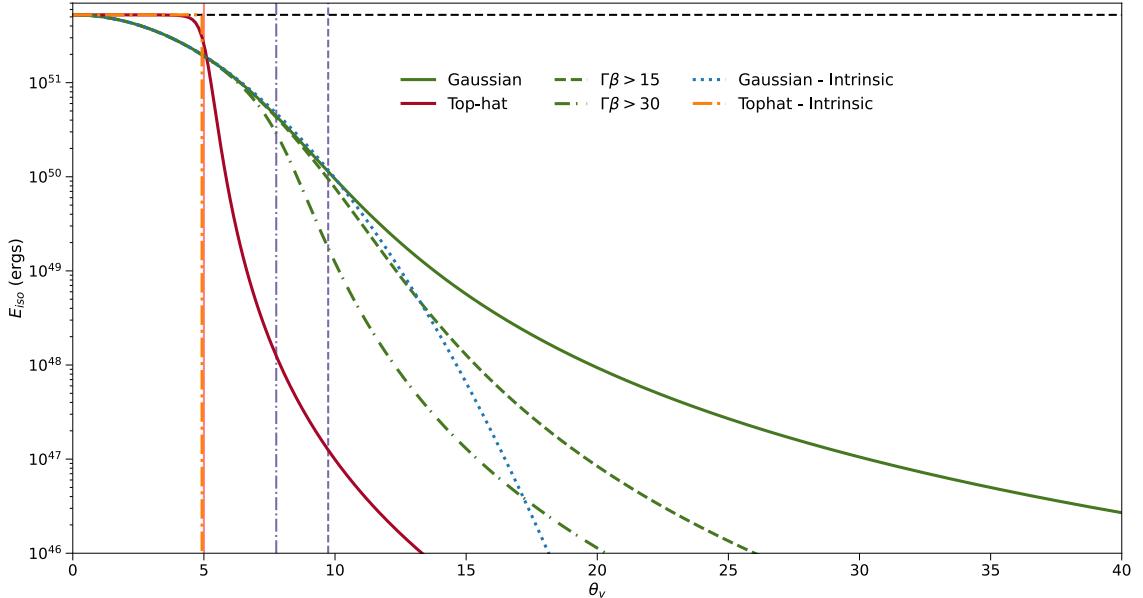
where:

- $\theta_v$  is the viewing angle of the observer.
- $\epsilon(\theta)$  is the normalised energy profile function.
- $\alpha_v$  is the angle between the line of sight and the direction to the jet element at  $(\theta, \phi)$ , given by  $\cos \alpha_v = \cos \theta_v \cos \theta + \sin \theta_v \sin \theta \cos \phi$ .
- $\theta_{max}$  is the upper cut-off of polar integration<sup>3</sup>.

Thus, depending on the underlying intrinsic structure assumed (be it Gaussian or tophat jet), using Eq. 4.4 one can infer the apparent structure of the jet. Note that

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<sup>3</sup>Such a cutoff can occur because the edge of the jet has been reached, or that the gamma-ray emission efficiency is lowered below a threshold, and so emission is negligible beyond  $\theta_{max}$ .



**Figure 4.2:** Variation of the apparent isotropic equivalent energy, for observers at different viewing angles. The figure shows both the top-hat (dark red) and Gaussian (dark green) jet structures, with  $E_{tot,\gamma} = 10^{49}$  ergs.  $\theta_j$  for the top-hat jet and  $\theta_c$  for the Gaussian jet are both  $5^\circ$  (marked with vertical solid bright red line), and  $\Gamma_0$  in both cases is 100. The horizontal dashed black line denotes  $E_{iso}(0)$ . The orange, dash-dotted line and the blue dotted lines are the tophat and Gaussian intrinsic jet structures, represented as  $4\pi\epsilon(\theta_v)$ . For the solid green curve, the entire jet emits gamma-rays, whereas for the dashed and dashed-dotted green curves the emission is restricted to regions where  $\Gamma\beta > 15$  and  $\Gamma\beta > 30$ , leading to limits in the polar integration of  $9.74^\circ$  (vertical dashed violet line) and  $7.76^\circ$  (vertical dash-dotted violet line) respectively.

this is the variation of the (apparent) isotropic equivalent energy ( $E_{iso}$ ) at various viewing angles ( $\theta_v$ ), and hence gives how the sGRB jet, if launched, would appear to an observer at an angle to the jet's axis. This variation is shown in Fig. 4.2, and is used a prior for further analysis.

### 4.2.3 Monte-Carlo simulations

Using the information from the previous analyses and further priors motivated from them, a Monte-Carlo simulation was run where  $10^5$  realisations of the Gaussian jet were made. Their model fluence was compared with what was reported by INTErnational Gamma-Ray Astrophysics Laboratory (INTEGRAL).

Around the time of the GW trigger, INTEGRAL was observing the entire Ad-

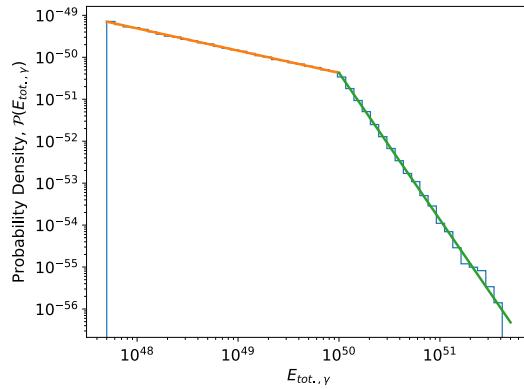
vLIGO/VIRGO localisation region and according to Minaev et al. 1925, saw a low SNR short duration ( $\sim 1$ s) excess roughly 6s after the merger. Further analysis reported a fluence of  $(1.6 \pm 0.4) \times 10^{-7}$  erg/cm $^2$ . The priors used for the various parameters of the Gaussian jets realised are as given in table 4.4.

	$E_{tot.,\gamma}$ (erg)	$\Gamma_0$	$\theta_c$
<b>Uniform Energy Prior</b>	$44 < \log_{10}(E_{tot.,\gamma}) < 51$	$5 < \Gamma_0 < 500$	$3^\circ < \theta_c < 20^\circ$
<b>Broken Power Law Energy Prior</b>	BPL { [5 $\times 10^{47}$ , 10 $^{50}$ ; -0.53], [10 $^{50}$ , 5 $\times 10^{51}$ ; -3.5] }	$100\Gamma_0 500$	$3^\circ < \theta_c < 20^\circ$

**Table 4.4:** Priors on the total energy emitted in gamma-rays ( $E_{tot.,\gamma}$ ), bulk on-axis Lorentz factor ( $\Gamma_0$ ) and core-angle ( $\theta_c$ ). The notation used for the broken power-law distribution is explained below.

Here, BPL{[5  $\times 10^{47}$ , 10 $^{50}$ ; -0.53], [10 $^{50}$ , 5  $\times 10^{51}$ ; -3.5]} is equivalent to:

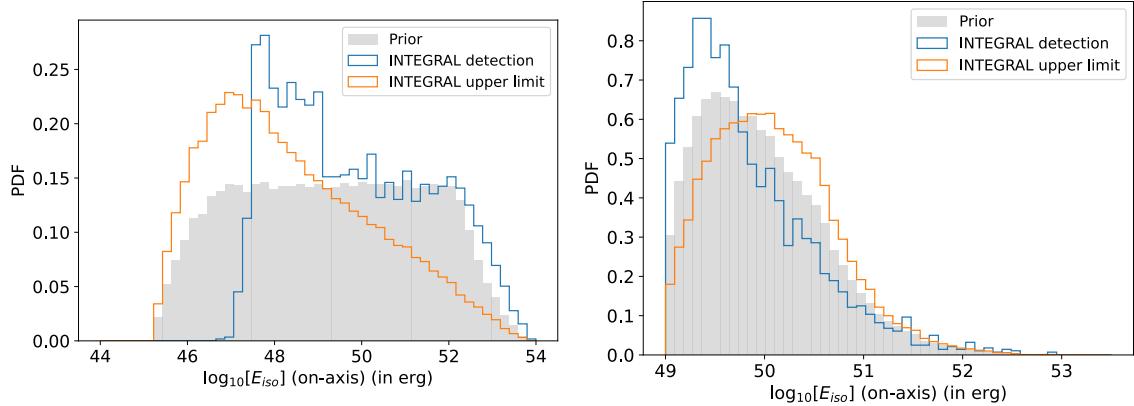
$$P(E_{tot.,\gamma}) \propto \begin{cases} E_{tot.,\gamma}^{-0.53}, & 5 \times 10^{47} \text{ ergs} < E_{tot.,\gamma} < 10^{50} \text{ ergs} \\ E_{tot.,\gamma}^{-3.5}, & 10^{50} \text{ ergs} < E_{tot.,\gamma} < 5 \times 10^{51} \text{ ergs} \end{cases} \quad (4.5)$$



**Figure 4.3:** The broken power law distribution as described in Eq. 4.5.

This particular distribution for the prior of  $E_{tot.,\gamma}$  is adopted since it is able to reproduce the fluence distribution observed, for values above the limiting fluence of  $2 \times 10^{-7}$  erg/cm $^2$  (see Mohan, Saleem, and Resmi 2019). Furthermore, the power-law indices are adopted from the luminosity function described in Ghirlanda et al. 2016.

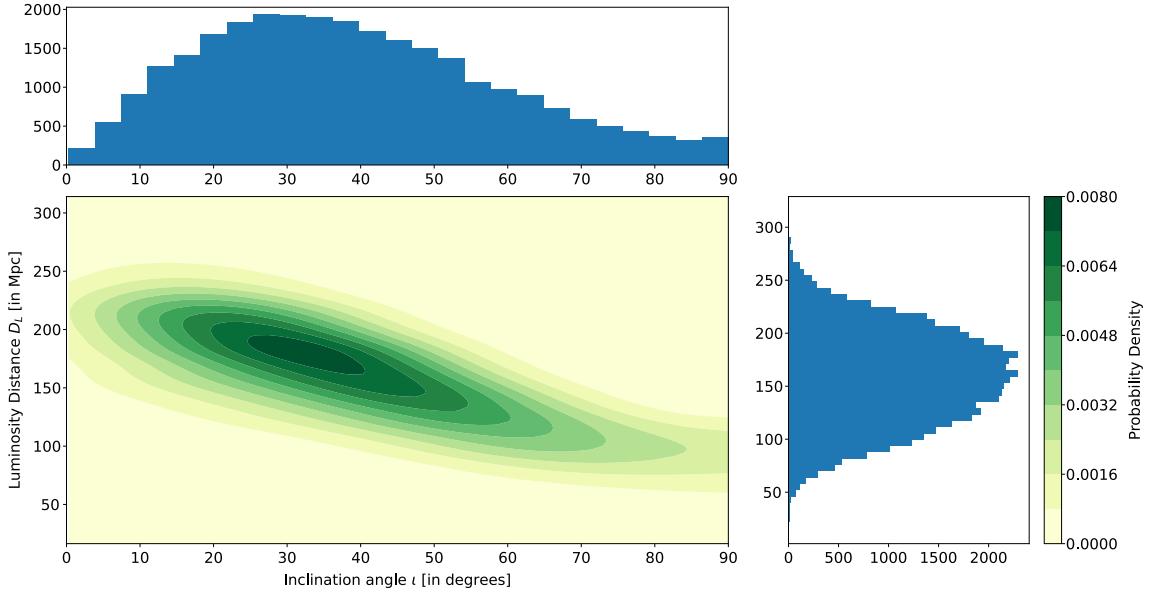
In applying these priors, along with the  $D_L - \iota$  prior and the fluence values of  $2 \times 10^{-7}$  and  $(1.6 \pm 0.4) \times 10^{-7}$  erg/cm<sup>2</sup> as upper limit and detections respectively, supplied by INTEGRAL, the marginalized posteriors for  $\theta_c, \Gamma_0, \theta_v$  and  $E_{tot,\gamma}$  are obtained. These are converted into  $E_{iso}(0)$ . The posterior distributions of the on-axis, apparent isotropic equivalent energy, for the two priors considered, is shown in Fig. 4.4. As is evident from seeing the figure, the INTEGRAL fluence is a good constraint for the priors. For the uniform prior case, considered as a detection, the posterior  $E_{iso}(0)$  is tightly constrained to be between  $3.51 \times 10^{47} - 6.26 \times 10^{52}$  ergs, which shows that for an on-axis observer, the event would have appeared as a typical sGRB along with the GW event. Even considering as an upper-limit constrains  $E_{iso}(0) \leq 1.48 \times 10^{51}$ , which is broadly in agreement with that observed for typical sGRBS. On the other hand, the narrower, broken power-law prior is not constrained well with the INTEGRAL fluence. Considered as a detection, the 90% credible posterior bounds on  $E_{iso}(0)$  are  $1.17 \times 10^{49} - 1.3 \times 10^{51}$  ergs, whereas considered as an upper limit,  $E_{iso}(0) \leq 7.69 \times 10^{50}$  erg. In both cases, the posteriors are sensitive to the choice of the prior, but nevertheless, one cannot rule out an sGRB jet which would have been seen by an on-axis observer.



**Figure 4.4:** Posterior distributions of the apparent on-axis isotropic equivalent energy  $E_{iso}(0)$ , for two assumed priors on the total energy emitted in gamma-rays,  $E_{tot.,\gamma}$ . These figures give constraints on the  $E_{iso}(0)$  of the sGRB associated with GW190425, assuming a Gaussian structured jet. **Left:** Grey histogram indicates the uniform priors on  $\log_{10}(E_{tot.,\gamma}/\text{erg})$  in the range of [44 — 51], and on  $\theta_c$  in [3, 20] degrees. **Right:** the same prior is used for  $\theta_c$  but a broken power law prior is used for  $E_{tot.,\gamma}$ . The orange histograms in both are a result of considering an INTEGRAL fluence upper limit of  $2 \times 10^{-7} \text{ erg/cm}^2$ , where the blue histograms in both are a result of considering an INTEGRAL fluence detection of  $(1.6 \pm 0.4) \times 10^{-7} \text{ erg/cm}^2$ . In both cases, the on-axis energy of a possible associated GRB is within the range of that of the cosmological sGRB population.

#### 4.2.4 Using LIGO posteriors

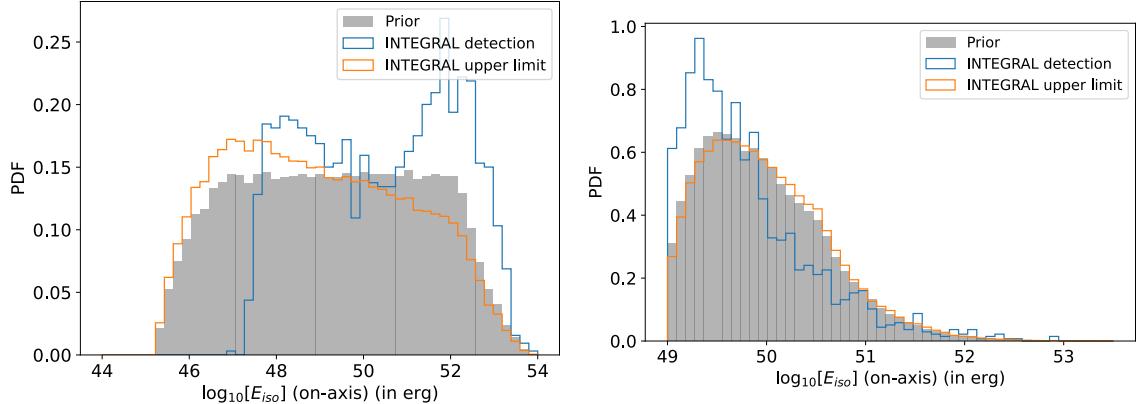
In the time since the analysis in M. Saleem et al. 2020 was carried out, the posteriors for the O3a events were released as part of the Gravitational Wave Transient Catalog (GWTC) 2 (see R. Abbott et al. 2020). These are now available as part of the event portal at the Gravitational Wave Open Science Center (GWOSC), which lists the files that store samples from posterior distributions for various parameters, for each event in GWTC-2. This allows one to use the actual posteriors for the inclination angle and the luminosity distance reported by the LVC. These are plotted in Fig. 4.5, below.



**Figure 4.5:** Similar as 4.1, but with samples for  $D_L$  and  $\iota$  from the posteriors released by the LVC, as part of the GWTC-2 data release.

These posteriors for the parameters  $\iota$  and  $D_L$  will be more accurate than the ones generated in §§4.2.1, since those posteriors are generated by approximating the O3a detector noise curves using the conservative best estimate of O2. Hence using the actual detector noise around the time of the event and performing parameter estimation for the parameters of interest (which is done by LVC), and using the resultant posteriors will be more accurate for further analysis.

Using these posteriors, the constraints on the energetics of a sGRB jet being powered by an event like GW190425 change slightly. This is shown in Fig. 4.6.



**Figure 4.6:** Posterior distributions of the apparent on-axis isotropic equivalent energy  $E_{iso}(0)$ , with similar priors as Fig. 4.4 but with LVC posterior samples as input.

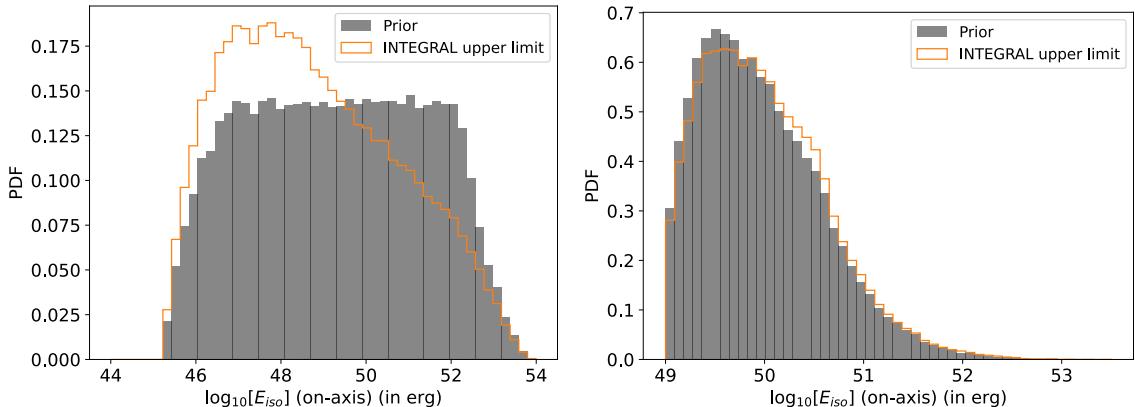
Considering the INTEGRAL fluence as a detection, the posterior bounds on  $E_{iso}(0)$  are  $5.61 \times 10^{47} - 8.48 \times 10^{52}$  ergs. This again leads to the conclusion that for an on-axis observer, the sGRB jet would have been detectable. Considering the INTEGRAL fluence as an upper limit instead, givens that  $E_{iso}(0) \leq 7.43 \times 10^{51}$  erg. Similar to the previous case, the narrow broken-power law prior is not well constrained by the INTEGRAL fluence limits, which places the bounds  $1.15 \times 10^{49} - 1.11 \times 10^{51}$  ergs (considered as a detection) and  $E_{iso}(0) \leq 6.74 \times 10^{50}$  erg (considered as an upper limit). In this case as well, the conclusion is the same, that the possibility of sGRB which would have been visible seen on-axis cannot be ruled out. However, this showcases the usefulness of the method described in §4.2.1, wherein the posterior for the parameter  $\iota$  can be approximated and used for further analysis, without having to wait for this information to be released officially.

### 4.3 Preliminary analysis of GW190426\_152155

The event GW190426\_152155 is listed in GWTC-2 (R. Abbott et al. 2020) as an event with a network matched filter signal-to-noise ratio of 10.1, a false-alarm rate of 1 event per 1.6276 yr and the component masses are  $5.7_{-2.3}^{+4.0} M_\odot$  and  $1.5_{-0.5}^{+0.8} M_\odot$ . From GRACE-DB, this event has probabilities 0.58, 0.24, 0.12, 0.06 respectively of being a Terrestrial, BNS merger, NSBH merger or MassGap merger event.

Although there were no significant excesses reported by any of the Gamma-ray satellites observing the LVC localisation area at the time of the GW trigger, the INTEGRAL satellite reported an upper limit fluence of  $1.7 \times 10^{-7}$  ergs/cm<sup>2</sup>. With the

priors as given in Table 4.4, this fluence upper limit reported by INTEGRAL is taken as the primary constraint. Similar to the process described in §4.2.3,  $10^5$  realisations of the jet are made, the apparent on-axis isotropic equivalent energy calculated, from which the fluence is computed and those realisations with a fluence beyond the upper limit are rejected. The resulting population has a distribution as given in Fig. 4.7. As can be seen from Fig. 4.7, at apparent on-axis isotropic equivalent energies below  $10^{48}$  erg, the uniform prior differs appreciably from the posterior, whereas at higher energies the posterior and prior exhibit the same behaviour. In the case of the broken power-law prior, at all energies considered, the posterior and prior distributions are the fluence upper limit offered by INTEGRAL doesn't offer tight constraints. This is because for both the priors, the posterior largely follows the same behaviour as the prior, thus indicating that the constraint applied is a poor one.



**Figure 4.7:** Posterior distributions of the apparent on-axis isotropic equivalent energy  $E_{iso}(0)$ , for two assumed priors on the total energy emitted in gamma-rays,  $E_{tot,\gamma}$ . These figures give constraints on the  $E_{iso}(0)$  of the sGRB associated with GW190426\_152155, assuming a Gaussian structured jet. **Left:** Grey histogram indicates the uniform priors on  $\log_{10}(E_{tot,\gamma}/\text{erg})$  in the range of [44 — 51], and on  $\theta_c$  in [3, 20] degrees. **Right:** the same prior is used for  $\theta_c$  but a broken power law prior is used for  $E_{tot,\gamma}$ . The orange histograms in both are a result of considering an INTEGRAL fluence upper limit of  $1.7 \times 10^{-7} \text{ erg/cm}^2$ .

This process sees whether the models used for the analysis of GW190425 also apply to the NSBH event GW190426\_152155. Although the constraint is not very good in either case of the prior, the conclusion that there is the possibility of a sGRB jet which could have been detected had the observer been on-axis, cannot be ruled out completely. In this case, considering the constraint as an upper limit, the bounds are  $E_{iso}(0) \leq 4.52 \times 10^{51} \text{ erg}$  and  $E_{iso}(0) \leq 6.81 \times 10^{50} \text{ erg}$  for the uniform prior

GRB ID	Relevant GCN Notices	Duration (ms)	Fluence (erg/cm <sup>2</sup> )	Redshift	Lum. Dist (Mpc)	$E_{iso}$ (ergs)
GRB20050509B	3385, 3390	30	$2.3^{+0.9}_{-0.9} \times 10^{-8}$	0.226	1133.2	$3.53^{+1.38}_{-1.38} \times 10^{48}$
GRB20130603B	14741, 14744	180	$(6.3^{+0.3}_{-0.3}) \times 10^{-7}$	0.356	1911.9	$2.76^{+0.044}_{-0.044} \times 10^{50}$
GRB20160821B	19844, 19846	480	$(1.0^{+0.1}_{-0.1}) \times 10^{-7}$	0.162	781.8	$7.31^{+0.73}_{-0.73} \times 10^{48}$

**Table 4.5:** List of typical cosmological sGRBs and their source parameters. The luminosity distance was calculated from the redshift using a standard cosmology of  $H_0 = 69.6$  km/s/Mpc,  $\Omega_M = 0.286$  and  $\Omega_\Lambda = 0.714$ .

and the broken power-law prior respectively. As an additional check, table 4.5 lists the parameters for some of the typical, well-studied sGRBs. Comparing what was obtained after the analysis of GW190426\_152155 to the observed isotropic equivalent energy for these sGRBs, it can be seen that this NSBH event is still in the ballpark of cosmological sGRBs.

## 4.4 Summary

# Chapter 5

# Results and Discussion

## 5.1 Summary

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