

# **X-ray Morphology of the Galaxy Group NGC 1550 with eROSITA**

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Unterschrift

1. Gutachter: Prof. Dr. Thomas Reiprich
2. Gutachterin: Prof. \_\_ \_\_

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## Acknowledgements

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I would like to thank ...

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*Context:* Galaxy groups like NGC 1550 play an important role in understanding the Universe's large-scale structure. They contain a significant fraction of the Universe's baryonic mass and are characterized by hot, ionized intragroup medium emitting X-rays, making X-ray observations essential for their study.

*Aims:* This thesis aims to characterize the X-ray morphology of the galaxy group NGC 1550 through surface brightness analysis in different directions, utilizing data from the *extended ROentgen Survey with an Imaging Telescope Array* (eROSITA).

*Methods:* Data from eRASS:1 was corrected and reduced for instrumental and background contamination. A  $\beta$ -model was employed to fit the full azimuthal surface brightness profile, and sectoral analyses were conducted to compare the profiles across different regions of the galaxy group.

*Results:* NGC 1550 appears relaxed and spherically symmetrical. The  $\beta$ -model fit yielded a  $\beta$ -value of  $0.478 \pm 0.008$  and a core radius of  $r_c = 60'' \sim (15 \pm 2)$  kpc. No significant deviations from spherical symmetry were found in the full azimuthal profile or in the sectoral analyses beyond  $390''$ . Discrepancies between observational and fitted background levels were noted, suggesting the need for improved background estimation.

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# CHAPTER 1

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## Introduction

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Unlike the larger and more massive galaxy clusters, galaxy groups are smaller systems with typical masses around  $3 \times 10^{13} M_{\odot}$  (Schneider, 2006). Despite their relative modest size, the study of galaxy groups is fundamental to our understanding of the Universe's large-scale structure (LSS), as they likely contain a significant fraction of the total baryonic mass in the Universe (Fukugita, Hogan, and Peebles, 1998). The space between galaxy clusters and groups is characterized by a hot, ionized gas known as the intracluster medium (ICM) or the intragroup medium (IGM), which fills the space between the galaxies. This gas emits copious amounts of X-rays making X-ray observations an essential tool for identifying and studying these structures (Kravtsov and Borgani, 2012).

The galaxy group NGC1550, first linked to the extended X-ray source RX J0419+0225 through the ROSAT All-Sky Survey (Böhringer et al., 2000), has been the subject of extensive X-ray observations. Analysis have been conducted using data from various instruments, including ASCA (Kawaharada, Makishima, Takahashi, et al., 2003), XMM-Newton (Kawaharada, Makishima, Kitaguchi, et al., 2009), Chandra (Sun et al., 2003), and Suzaku (Sato et al., 2010). In this thesis, data gathered with the *extended ROentgen Survey with an Imaging Telescope Array* (eROSITA) shall be utilized to characterize the X-ray surface brightness profile of NGC1550 and compare the results with previous studies.

Following a brief overview of the theoretical background in Chapter 2, Chapter 3 will focus on reducing and correcting the data for instrumental and background contamination sources. Chapter 4 will analyze the surface brightness profile of NGC1550, including a beta model fitting of the full azimuthal profile. Additionally, the analysis will compare the north, south, east, and west sectors against each other and against the full azimuthal profile. Chapter 5 will summarize all results, offer suggestions for improvement and for future works.

# CHAPTER 2

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## Theoretical Background

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### 2.1 Clusters and Groups of Galaxies

Throughout the Universe, galaxies are not distributed homogeneously but are instead aggregated into large cosmic structures known as galaxy groups or galaxy clusters, which are the largest virialized structures in the Universe. Galaxy clusters typically have masses exceeding  $M \gtrsim 3 \times 10^{14} M_{\odot}$ , whereas galaxy groups have masses around  $M \sim 3 \times 10^{13} M_{\odot}$  (Schneider, 2006). Advancements in X-ray astronomy have demonstrated that these structures are significant sources of X-ray radiation (A. G. Cavaliere, Gurkaynak, and Tucker, 1971). This emission is well understood to originate from a hot intergalactic gas known as the intracluster medium (ICM)<sup>1</sup>, which is characterized by temperatures in the range of  $10^7$  to  $10^8$  K and constitutes the primary baryonic component of galaxy clusters and groups (Schneider, 2006).

#### 2.1.1 The Intracluster Medium

Within the deep gravitational wells of galaxy clusters and groups, the temperature becomes sufficiently high to fully ionize lighter elements and partially ionize heavier elements, resulting in the formation of a plasma. This hot, diffuse, and optically thin plasma, known as the intracluster medium (ICM), emits a significant amount of X-ray radiation. X-ray observations of the ICM have enabled a wide variety of cosmological studies including advances in understanding the large-scale structure formation in the Universe (Kravtsov and Borgani, 2012).

#### 2.1.2 Emission Processes within the ICM

A key principle of electrodynamics is that accelerated charges radiate energy. This radiation is referred to as bremsstrahlung or “free-free” when an unbound charged particle, typically an electron, is accelerated by the electric field of other charges, usually ions. In the ICM, this process predominates at temperatures above  $k_B T_e \gtrsim 2$  keV, where the total emissivity  $\epsilon_{\text{ff}}$  at solar metallicity scales

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<sup>1</sup> The intercluster medium is also called the intragroup medium (IGM) in the case of galaxy groups. In the context of this thesis, the terms ICM and IGM will be used interchangeably, as a distinction between them is not necessary.

approximately as  $\epsilon_{\text{ff}} \propto T_e^{0.5} n_e^2$ , with  $n_e$  and  $T_e$  as the electron number density and temperature, respectively. At lower temperatures ( $k_B T \lesssim 2 \text{ keV}$ ), line emission becomes significant, with the emissivity being roughly described by  $\epsilon \propto T_e^{-0.6} n_e^2$ . Thus, for the low energy band 0.1 to 2.4 keV, one has a weak temperature dependence, i.e.  $\epsilon \propto n_e^2$ . (Reiprich and Pacaud, 2019).

### 2.1.3 The $\beta$ -model of the ICM

In the mid-1970s, (A. Cavaliere and Fusco-Femiano, 1976) laid the groundwork for the current understanding of the ICM. The X-ray morphology of clusters and groups is typically described by the beta model. This model assumes gas and galaxies share the same gravitational potential. Given a constant temperature  $T$  and velocity dispersion  $\sigma_r$ , the radial distributions of the gas  $n_{\text{gas}}$  and galaxies  $\rho_{\text{gal}}$  obey

$$\frac{n_{\text{gas}}(r)}{n_{\text{gas}}(0)} = \left[ \frac{\rho_{\text{gal}}(r)}{\rho_{\text{gal}(0)}} \right]^\beta ; \quad \beta = \frac{\mu m_p \sigma_r^2}{k_B T}$$

where  $\mu$  is the mean molecular mass,  $m_p$  is the proton mass and  $k_B$  is the Boltzmann constant. Using the King approximation  $\rho_{\text{gal}}(r) = \rho_{\text{gal}(0)} [1 + (r/r_c)^2]^{-3/2}$  (King, 1962), where  $r_c$  is the core radius and  $\epsilon \propto n_e^2$ , the X-ray surface brightness  $S_X$ , obtained by integrating over the emissivity, follows the profile

$$S_X(r) = S_X(0) \left[ 1 + \left( \frac{r}{r_c} \right)^2 \right]^{-3\beta+1/2}.$$

The beta model is widely used to describe cluster and group emissions, despite the assumptions listed above often being unmet. Excess core emission, for example, is typically addressed by adding a second beta model component:  $S_X^{\text{tot}} = S_{X,1} + S_{X,2}$  (Reiprich and Pacaud, 2019).

### 2.1.4 The galaxy group NGC1550

The galaxy group NGC1550 is located at a right ascension of  $64.9066^\circ$  and a declination of  $2.4151^\circ$ , with a redshift of  $z = 0.0123$  (Reiprich and Böhringer, 2002). Although the lenticular galaxy NGC 1550 has long been observable in the optical, it was first associated with the extended X-ray source RX J0419+0225 through the ROSAT All-Sky Survey (Böhringer et al., 2000). The presence of this extensive X-ray halo, along with its considerable mass ( $\sim 2 \times 10^{13} M_\odot$ ), led to its classification as a galaxy group (Kawaharada, Makishima, Takahashi, et al., 2003).

Recent studies indicate that NGC1550 does not meet the formal criteria for a fossil group (Sun et al., 2003; cf. Jones et al., 2003 for the criteria), but it shares several features with such groups. One notable characteristic is its high X-ray bolometric luminosity ( $\gtrsim 4.8 \times 10^{41} \text{ erg}$ ), which predominantly originates from a central, dominant, early-type galaxy. Fossil groups are thought to be formed when member galaxies in a regular galaxy group merge into a central dominant galaxy, and evidence for this process has been found in the metal distribution of NGC1550 (Kawaharada, Makishima, Kitaguchi,

et al., 2009; Sato et al., 2010). Additionally, a recent study has identified signs of AGN (active galactic nucleus) feedback and sloshing, suggesting a minor merger occurred around 33 million years ago (Kolokythas et al., 2020). While the group appears to be relaxed overall, a slight east-west elongation in the central region has been noted by multiple authors (Kolokythas et al., 2020; Sun et al., 2003).

## 2.2 eROSITA

The *extended ROentgen Survey with an Imaging Telescope Array* (eROSITA) is a highly sensitive, wide-field X-ray telescope designed to capture deep images across large areas of the sky. Mounted on the Spektrum-Roentgen-Gamma (SRG) observatory in a halo orbit around the second Lagrange Point (L2), eROSITA operates within the 0.2 to 10.0 keV energy range. It is the first instrument to perform an all-sky imaging survey in the hard X-ray band (2.0 to 10.0 keV). In the soft X-ray band (0.5 to 2.0 keV), eROSITA boasts a sensitivity greater than that of its predecessor, the ROSAT All-Sky Survey. eROSITA features seven identical mirror modules, known as Telescope Modules (TMs), with 54 mirror shells in Wolter-I geometry and a 1.6-meter focal length. Five TMs (TM1, TM2, TM3, TM4 and TM6) have aluminum on-chip optical light filters and are collectively referred to as TM8. The remaining two TMs (TM5, TM7), designed for low-energy spectroscopy, lack these filters and are referred to as TM9. Collectively, TM8 and TM9 are referred to as TM0. (Predehl et al., 2021). In this thesis, data from the first all-sky survey (eRASS:1) will be utilized. The data is provided in the form of event lists, which essentially detail the spatial position of X-ray photon arrivals on the detector, their time of arrival, and their energy. The event lists are made available in  $3.6^\circ \times 3.6^\circ$  regions known as skytiles.

## 2.3 Sky background and contamination sources

For a thorough analysis of X-ray photons, it is essential to carefully consider both external background and internal instrumental contamination effects. The following section will provide a brief overview of the most important factors relevant to this analysis.

**Cosmic X-ray Background (CXB):** The cosmic X-ray background (CXB) comprises multiple sources, including diffuse, unabsorbed thermal emissions from the Local Hot Bubble, a plasma cavity surrounding the Sun, and absorbed thermal emissions from the Galactic halo (Galeazzi et al., 2006). Additionally, it includes discrete extragalactic sources, predominantly unresolved AGNs (Brandt and Hasinger, 2005). The diffuse component is more prominent in the lower energy band  $\sim 1$  keV, while the extragalactic sources dominate at higher energies. In this analysis, background emission from the nearby Orion-Eridanus superbubble, which spans roughly  $40^\circ$  on the sky, shall also be present, as it emits in the soft X-ray regime (Krause et al., 2014). Figure 2.1, taken from (Krause et al., 2014), shows an X-ray map of the Orion-Eridanus superbubble in the 0.5 to 2.0 keV band, illustrating its overlap with the emission from the galaxy group NGC1550.

**Non-X-ray Background (NXB):** The non-X-ray background consists of two main components: highly variable soft protons flares (SPF) from the solar corona and Earth's magnetosphere, which

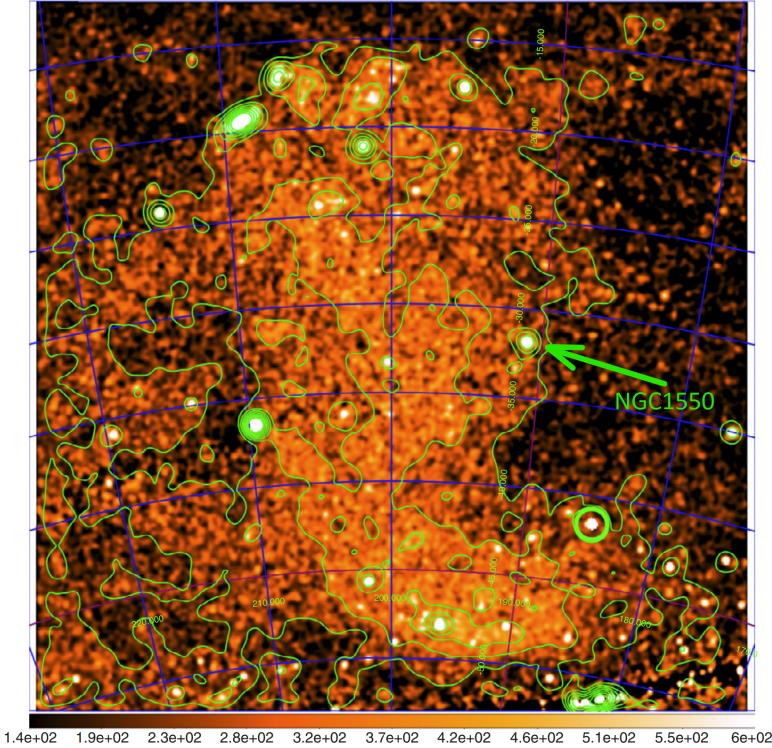


Figure 2.1: X-ray map in galactic coordinates of the Orion-Eridanus superbubble in the 0.5 to 2.0 keV band from the ROSAT all-sky survey. The colorbar units are  $10^{-6} \text{ cts} \cdot \text{arcmin}^{-2} \text{s}^{-1}$  and the contour levels are 220, 300, 380, 460, 540, 620, 700. The image was taken from (Krause et al., 2014). The green arrow and label were added to highlight the position of NGC1550.

can be focused onto detectors, and energetic Galactic Cosmic Ray (GCR) primaries, which interact with the detector to produce secondary particles. The secondary particles can deposit charge in the detector, making it challenging to distinguish them from true X-ray events. This is generally referred to as the particle-induced background (PIB) (Bulbul et al., 2020).

**eROSITA light leak:** Shortly after the launch of eROSITA, it was observed that CCDs lacking an on-chip filter (TM9) recorded a notably higher number of events. This was attributed to optical and ultraviolet light from the Sun entering the CCD through an unidentified gap in the detector shielding, and was subsequently termed “light-leak” (Predehl et al., 2021).

**$N_{\text{H}}$  absorption:** As X-rays travel to the detector, they undergo photoelectric absorption in the interstellar and intergalactic medium. The cross-section  $\sigma \propto E^{-3}$  is inversely proportional to energy, causing a bias toward harder X-rays, as they interact less. Additionally, the cross-section is proportional to the atomic number,  $\sigma \propto Z^5$ , making metal abundance crucial for energies  $\gtrsim 0.2 \text{ keV}$  (Willingale et al., 2013).

# CHAPTER 3

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## Data Reduction

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In the following section, the underlying data shall be reduced and corrected for the various effects and contamination sources explained in Section 2.3. Data from eRASS:1<sup>1</sup> and all TMs (1-7) is utilized, which were processed with the eROSITA pipeline processing version c010. The galaxy group NGC1550 is located in skytile 065087. In addition, the surrounding skytiles 062084, 062087, 062090, 065084, 065090, 068084, 068087, and 068090 are used to encompass regions up to  $\sim 3R_{200}$  (cf. value below). The data reduction is performed with the software HEASoft version<sup>2</sup> 6.29 and the extended Science Analysis Software System<sup>3</sup> (eSASS 4DR1).

Values for  $R_{200}$  and  $R_{500}$  were taken from Reiprich and Böhringer, 2002 and converted to  $R_{200} = 58.58'$  and  $R_{500} = 37.15'$  based on the cosmology used in this paper. The cosmology assumes  $\Omega_m = 0.3$ ,  $\Omega_\Lambda = 0.7$  and  $h = 0.7$ , where  $1'' = 0.252$  kpc at the redshift  $z = 0.0123$  of the galaxy group. For image visualization, the Matplotlib (Hunter, 2007) and Astropy (Collaboration et al., 2022) Python-packages were utilized.

### 3.1 Raw photon images

Before the data reduction process, raw photon images for all combined skytiles and TMs are presented across the following energy bands: 0.2 to 2.3 keV, 2.3 to 6.0 keV, and 6.0 to 9.0 keV. The skytiles were combined using the *eSASS* task `evtool` with no additional parameters applied to show all inherent deficiencies. The raw photon images are presented in Figure 3.1. The counts are noticeably lower on the right half of each image, which shall be addressed in detail through the absorption and exposure correction in Section 3.6. As observed, the group emission is mainly concentrated in the lower energy band, making the low energy band the focus for detecting emission structures. In this analysis, the 0.2 to 2.3 keV energy band is used. Due to the light-leak in TM9, however, the energy range is restricted to 0.8 to 2.3 keV, while it remains 0.2 to 2.3 keV for TM8. Hereafter, this

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<sup>1</sup> [https://erosita.mpe.mpg.de/dr1/AllSkySurveyData\\_dr1/](https://erosita.mpe.mpg.de/dr1/AllSkySurveyData_dr1/) (Last accessed: 05.08.2024)

<sup>2</sup> <https://heasarc.gsfc.nasa.gov/docs/software/heasoft/> (Last accessed: 04.08.2024)

<sup>3</sup> <https://erosita.mpe.mpg.de/dr1/eSASS4DR1/> (Last accessed: 04.08.2024)

TM-dependent energy band will be referred to as the “soft-band” and the 6.7 to 9.0 keV energy range as the “hard-band”.

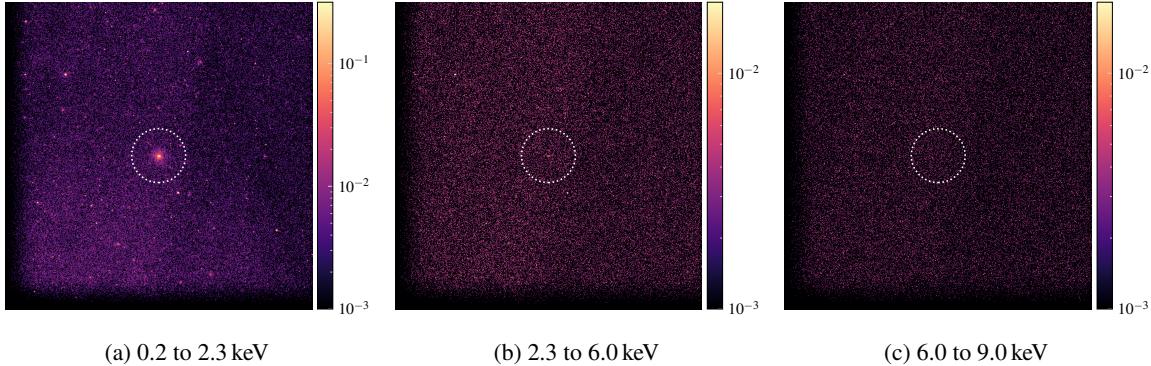


Figure 3.1: Raw photon images from TM0 of all combined skytiles centered around NGC1550, displayed in the energy bands 0.2 to 2.3 keV (a), 2.3 to 6.0 keV (b), and 6.0 to 9.0 keV (c), with Gaussian smoothing of 4 pixels applied. The colorbar represents (smoothed) photon counts. For visualization purposes, different cuts were used for the middle and right image and  $R_{200}$  is shown (white dashed circle). Most of the group emission is visible in the lower energy band (0.2 to 2.3 keV).

## 3.2 Image filtering

Each skytile is cleaned individually using `evtool` with `pattern=15` to select all event patterns and `flag=0xe00fff30` to remove bad pixels and CCD corners. Subsequently, soft proton flares are identified and mitigated through the following process: the `flaregti` task is used to generate light curves with 20 s time bins in the energy range of 5 to 10 keV. A  $3\sigma$  threshold is determined; time intervals where count rates exceed this threshold are likely due to soft proton flares. The task `flaregti` is then rerun using this threshold to establish good-time-intervals (GTIs) excluding these flare periods, which are applied using `evtool` with the `gti="FLAREGTI"` parameter. All SPF-filtered and cleaned skytiles are combined into a single TM0 photon image using `evtool`. Hereafter, these combined images shall be referred to as “filtered”.

## 3.3 Exposure map and image creation

The `evtool` task, along with the `telid` parameter, is used to split the TM0 filtered photon images into individual filtered images for each TM. These images are then resized to the soft-band and the hard-band using `evtool` with the parameters `emin` and `emax`. For subsequent PIB and exposure corrections, both vignetted and flat (non-vignetted) exposure maps are generated for each TM in the soft band using the `expmap` task with the `withvignetting` parameter. Moreover, the `withdetmaps` parameter is specified to crop the images at the edge of the field of view for each TM. By keeping `withmergedmaps` at its default value of YES, merged all-telescope exposure maps are created. This means the exposure maps consider the 7 TMs as one TM with on-chip TM, assuming the combined

effective area of the 7 on-chip TMs. The vignetted and flat exposure maps can be found in Appendix TODO.

### 3.4 PIB Correction

It is necessary to subtract the particle-induced background (PIB) from the combined filtered photon image. The following approach is based on extensive studies of the eROSITA filter wheel closed (FWC) data conducted by Dr. F. Pacaud, as utilized for example in (Reiprich, Veronica, et al., 2021). The PIB is modeled for each TM using FWC data. Due to the minimal spatial variation of the PIB, this modeling utilizes the flat exposure map created in Section 3.3. Furthermore, given the negligible spectral variation, the counts  $H_{\text{obs}}$  in the hard band, where PIB counts dominate, is used to estimate the PIB contribution in the soft bands by multiplication with the ratio  $R$  of the number of FWC counts in the soft band  $S_{\text{FWC}}$  to the hard band  $H_{\text{FWC}}$ . The PIB-map for each TM is then generated by applying this factor to the flat exposure map, normalized to 1 by dividing by the sum of all pixel values (norm. flat exposure map). Hence, the PIB map of a given TM is given by

$$\text{PIB map}_{\text{TM}} = H_{\text{obs}} R \cdot (\text{norm. flat exposure map}).$$

Soft-band, PIB-corrected image are obtained by subtracting the PIB map of each TM from the respective soft-band, filtered photon image. The complete image for TM0 is obtained by co-adding all PIB corrected images. To better visualize the PIB-correction, individual background maps are also co-added to form a complete PIB map, which can be found in Appendix A.1. Appendix ?? also contains the obtained  $H_{\text{obs}}$  and  $R$  values.

### 3.5 Absorption Correction

As discussed in Section 2.3, X-ray absorption by the ISM must be considered. To correct for this absorption, the methodology outlined in (Reiprich, Veronica, et al., 2021) is followed. Additionally, scripts for the absorption correction were provided by Angie Veronica (as utilized in Veronica, 2020). At energies  $\gtrsim 0.2$  keV, as is relevant for this analysis, absorption from metals play a significant role (cf. 2.3). Assuming solar metallicity, the hydrogen column density ( $N_{\text{H}}$ ) can be used to trace the absorbing material. A cutout of the HI4PI all-sky survey (HI4PI Collaboration et al., 2016) is reprojected onto field of view of the combined tile image to create a neutral atomic hydrogen map ( $N_{\text{HI}}$ -map). Additionally, as outlined in (Willingale et al., 2013), a molecular hydrogen map ( $N_{\text{H}_2}$ -map) is constructed by dividing the full sky image into  $52 \times 52$  pixel cells, querying  $N_{\text{H}_2}$  values from the Swift homepage<sup>4</sup>, and distributing these values across each cell. The total  $N_{\text{Htot}}$ -map, constructed by  $N_{\text{Htot}} = N_{\text{H}} + 2N_{\text{H}_2}$ , is shown in Figure ??, with  $N_{\text{H}}$  ranging from  $N_{\text{Htot, min}} = 2.91 \times 10^{20} \text{ cm}^{-2}$  to  $N_{\text{Htot, max}} = 1.61 \times 10^{21} \text{ cm}^{-2}$ . Next, for each individual  $N_{\text{H}}$ -value in the  $N_{\text{Htot}}$ -map, the expected soft band count rates for TM1 and TM5 – which serve as proxies for TM8 and TM9, respectively –

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<sup>4</sup> <https://www.swift.ac.uk/analysis/nhtot/index.php> (Last accessed: 25.07.2024)

are simulated for the model

$$\text{apc}_{\text{LHB}} + \text{tb}_{\text{abs.}} \cdot (\text{apc}_{\text{MWH}} + \text{pow})$$

using the X-ray spectral fitting package XSPEC<sup>5</sup> (Arnaud, 1996) with the `fakeit` command and their respective area (ARF) and response files (RMF). Here,  $\text{apc}_{\text{LHB}}$  represents unabsorbed Local Hot Bubble emission,  $\text{tb}_{\text{abs.}}$  the absorption along the line of sight,  $\text{apc}_{\text{MWH}}$  the absorbed Milky Way Halo emission, and  $\text{pow}$  the absorbed emission from unresolved point sources (e.g., AGNs). A correction factor  $A_{\text{corr}}$  is determined for each value of  $N_{\text{H}}$  by dividing the simulated count rate for the  $N_{\text{Htot}}$ -map median ( $\overline{N_{\text{Htot}}}$ ) by the simulated count rate of the  $N_{\text{H}}$  of interest, hence

$$A_{\text{corr}}(N_{\text{H}}) = \frac{\text{simulated count rate}(\overline{N_{\text{Htot}}})}{\text{simulated count rate}(N_{\text{H}})}.$$

Finally, each  $N_{\text{H}}$  in the  $N_{\text{Htot}}$ -map is replaced by the corresponding correction factor  $A_{\text{corr}}$  to create an absorption correction map. Absorption correction maps are created for both TM1 and TM5.

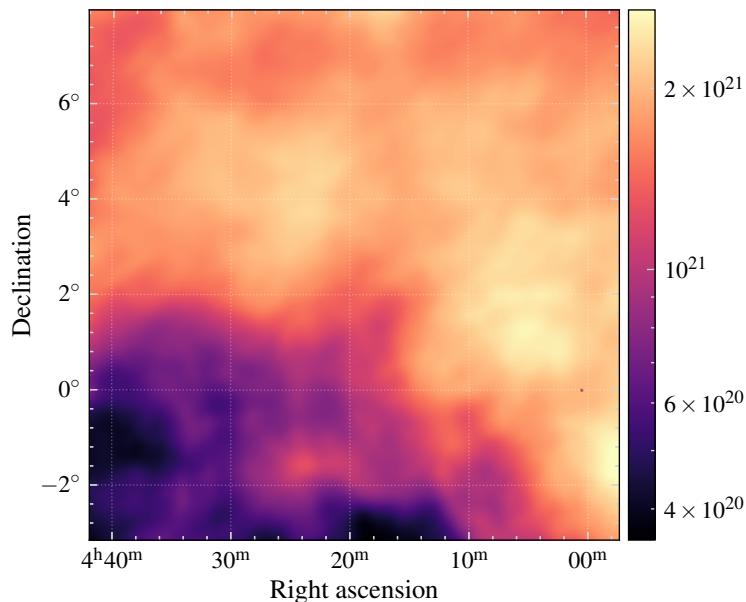


Figure 3.2:  $N_{\text{H}}$ -map

### 3.6 Exposure Correction

The count image of any given X-ray observation depends inherently on a detector's effective area and the telescope's pointing motion throughout the observation. To obtain meaningful flux units

<sup>5</sup> <https://heasarc.gsfc.nasa.gov/xanadu/xspec/> (Last accessed: 25.07.2024)

(e.g.  $\text{cts} \cdot \text{arcsecond}^{-2}\text{s}^{-1}$ ), these effects must be considered. This is achieved by dividing the count image by the vignetted exposure map created in Section 3.3, thereby rescaling all segments of the count image to the same relative exposure (Davis, 2001). However, an exposure map must be created separately for TM8 and TM9 due to three reasons: first, TM9 uses a narrower energy band, which lowers its expected count rate; second, this narrower energy band necessitates a different absorption correction map; third, TM8 and TM9 have very distinct area response files because of their different filter configurations. Thus, both exposure maps cannot simply be combined but must be corrected for these effects. The procedure outlined in (Reiprich, Veronica, et al., 2021) will be followed. First, the vignetted exposure map of TM8 and TM9 are divided by their respective absorption correction maps to obtain an absorption-corrected exposure map ( $\text{exmap}_{\text{TM8, corr}}$ ,  $\text{exmap}_{\text{TM9, corr}}$ ). Second, the ratio of PIB-corrected count rates of TM8 and TM9 is used to define a correction factor

$$E_{\text{corr}} = \frac{\text{PIB corr. count rate(TM9)}}{\text{PIB corr. count rate(TM8)}}.$$

The total absorption-corrected exposure map for TM0 is then given by

$$\text{exmap}_{\text{TM0, corr}} = \text{exmap}_{\text{TM8, corr}} + E_{\text{corr}} \cdot \text{exmap}_{\text{TM9, corr}}$$

A correction factor of  $E_{\text{corr}} = 0.4628$  is obtained. Dividing the filtered and PIB-corrected soft-band TM0 image by its absorption-corrected exposure map results in the final filtered, PIB-corrected, absorption-corrected, and exposure-corrected soft-band image. Henceforth, this will be referred to as the “corrected” soft-band image. Figure 3.3 compares the filtered, full energy-band count rate image (the filtered full band count image divided by its respective exposure map) with the corrected soft-band image, both smoothed using a Gaussian kernel of 12 pixels. It is evident that group emission can be much better distinguished from the background in the corrected soft-band image. Additionally, the count , as seen in Figure 3.1, has been addressed through the absorption and exposure correction.

### 3.7 Wavelet filtering and Point Source Removal

It is necessary to remove the emission from point-like sources (e.g. AGNs) to prevent interference with the group emission under study. This is achieved using the wavelet filtering pipeline as described in (Pacaud et al., 2006). Wavelet transformation decomposes an image  $I(x, y)$  into coefficients  $(w_1, \dots, w_n, c_n)$

$$I(x, y) = c_n(x, y) + \sum_{j=1}^n w_j(x, y).$$

Here,  $c_n(x, y)$  is the smoothed image, and  $w_j(x, y)$  essentially represents the contribution of a wavelet function at a scale  $j$  and position  $(x, y)$  to the total image. By retaining only the coefficients that satisfy

$$|w_j(x, y)| > k\sigma_j,$$

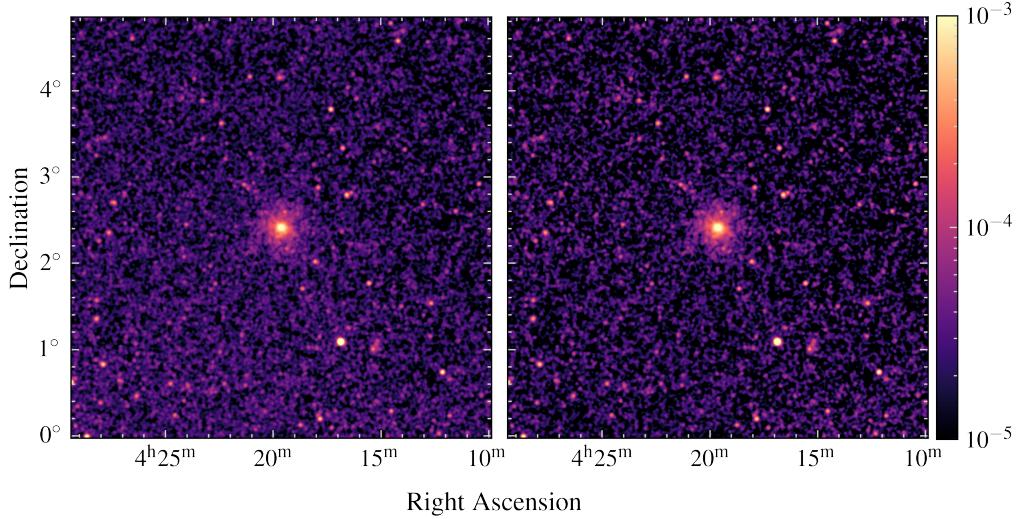


Figure 3.3: Comparison between the filtered, full energy-band count rate image (left) and the final corrected count rate image (right), both smoothed using a Gaussian kernel of 12 pixels. The corrected image clearly shows improved differentiation of group emission from the background in comparison to the filtered, full band image.

where  $\sigma_j$  is the standard deviation at scale  $j$  and  $k$  is a clipping factor, and then applying the inverse wavelet transformation, an image is obtained that includes only significant scales, i.e. those not due to noise (Starck, J.-L. and Pierre, M., 1998). The resultant image is referred to as a wavelet-filtered image. In the pipeline being used, the absorption-corrected exposure map,  $\text{exmap}_{\text{TMO, corr}}$ , and the filtered photon image are used to account for previous data reduction and to statistically handle the Poisson noise. After wavelet filtering, a source catalogue is obtained by SExtractor (Bertin and Arnouts, 1996). This is enabled by the significant noise reduction and smoothed background achieved by the wavelet filtering. The extended X-ray emission near NGC1550 is manually removed from the catalog, and a cheese-mask is created from the catalog and applied to  $\text{exmap}_{\text{TMO, corr}}$ . Emission around unrelated, but projectionally nearby, clusters and groups is also manually added to the cheese-mask. Rerunning wavelet filtering with the cheese-masked map reduces ringing artifacts, which typically occur near bright sources with steep flux gradients. Figure 3.4 compares the wavelet filtering before and after application of the cheese-mask to the exposure map.

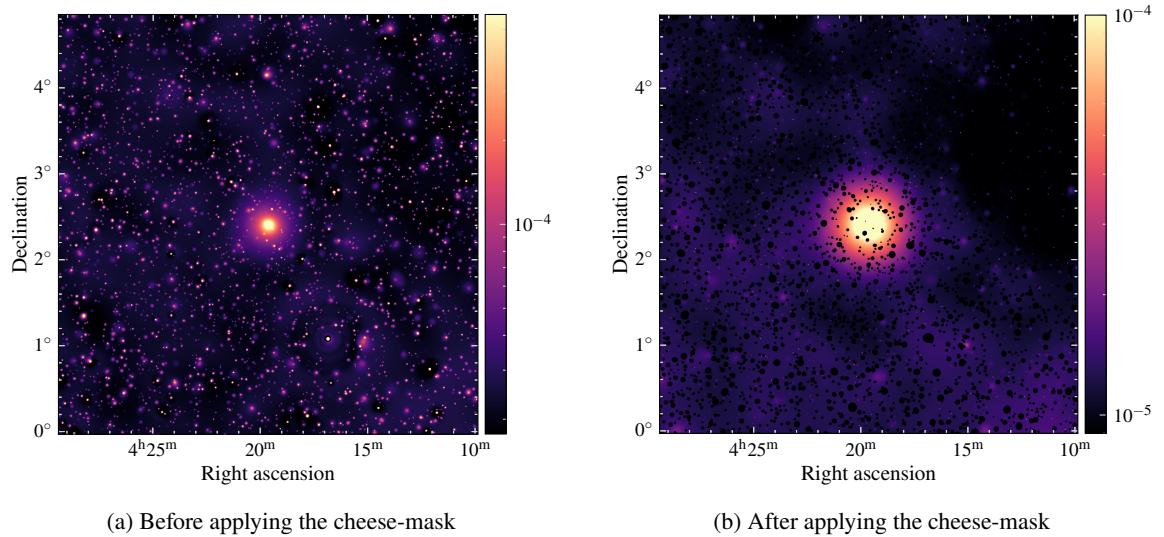


Figure 3.4: Corrected wavelet filtered image before (a) and after (b) applying the cheesemask created with SExtractor. The left image is normalized such that the ringing artifacts are more prominent. The cheese-masked image has significantly reduced ringing artifact and group emission is well visible.

# CHAPTER 4

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## Data Analysis and Discussion

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### 4.1 Image Inspection and galaxy distribution

A contour plot of the corrected and cheese-masked wavelet-filtered image is shown in Figure 4.1 to visualize emission beyond  $R_{500}$  more clearly. Furthermore, the galaxy distribution using data from NASA/IPAC Extragalactic Database (NED) is extracted within a  $1^\circ$  radius of the galaxy group NGC 1550 and overlayed. There are 26 galaxies with similar redshifts  $0.010 \leq z \leq 0.015$  including the galaxy NGC1550 ( $z = 0.0123$ ) itself. From Figure 4.1 it is clear that the group appears spherically symmetrical and relaxed, with a distinct emission cutoff visible at  $R_{500}$ . Beyond this radius, emission becomes difficult to distinguish from the background, which is fairly complex due to the Orion-Eridanus superbubble (cf. Section 2.3). There is a noticeable gradient in the background emission, with a lower background to the north compared to the south. The galaxy distribution is loosely clustered to the northeast, but no clear correlation with the X-ray contours is observed. Usually, this helps indicate whether excess emission along a certain direction is significant.

To highlight emission structures within  $R_{500}$ , Figure 4.2 presents a contour plot of the corrected soft-band image smoothed with a Gaussian kernel of 8 pixels. Within this radius, some deviations from perfect azimuthal symmetry are noticeable. A slight east-west elongation within approximately  $390''$  (inner circle in the figure) can indeed be inferred as mentioned in Section 2.1.4. Beyond  $390''$ , features are harder to distinguish: the emission in the north, south and east appear slightly more irregular. Emission in the west, falls more uniformly and there seems to be slight emission dip between  $390''$  and  $810''$  (outer circle in the figure) compared to the other sectors.

### 4.2 Full Azimuthal Surface Brightness Profile

To more accurately identify or exclude potential deviations from spherical symmetry, surface brightness (SB) analysis will be performed. First, the emission center is estimated by constructing a  $2'$  aperture around the apparent center (right ascension  $64.9066^\circ$  and a declination  $2.4151^\circ$ ). This aperture size is chosen to capture a statistically significant number of photons without leaving the group center. The flux-weighted average coordinates of the fully corrected image within this aperture

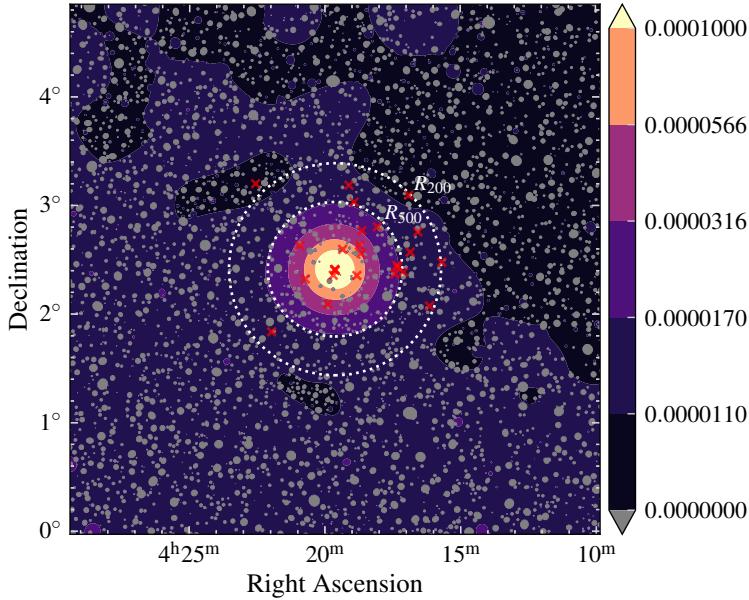


Figure 4.1: Contour plot of the corrected, cheese-masked, wavelet-filtered image with overlaid galaxy distribution (red crosses) within a  $1^\circ$  radius of NGC 1550. The color bar represents counts per second. The plot shows a distinct emission boundary at  $R_{500}$ , with emission beyond this radius becoming difficult to distinguish from the background. The galaxy distribution is clustered to the northeast, though no clear correlation with the X-ray contours is observed.

yields a right ascension of  $64.909^\circ$  and a declination of  $2.414^\circ$  and shall be referred to as the SB-center. Concentric annulli of  $1'$  width are constructed from the SB-center up to an angular distance of  $80' \approx 1.5R_{200}$ . The counts  $C$  within each annulus are determined using the `funcnts` task from the `funtools` software, with a Poisson error of  $\sqrt{C}$ . The surface brightness  $S$  for each annulus is calculated by

$$S = \frac{C_{\text{filt}} - C_{\text{PIB}}}{C_{\text{expmap}} \cdot A},$$

where  $C$  denotes the counts in the filtered photon image, total PIB map, and exposure map, respectively, and  $A$  is the annulus area. Errors are calculated using Gaussian error propagation. Furthermore, background estimation is performed using 10 circular regions with a  $30'$  radius, each centered at a distance of  $100'$  from the SB-center (Figure 4.3). The average background surface brightness is evaluated for all circles combined and separately for the northern (Circles 1-5) and southern region (Circles 6-10) to account for the background gradient. Table 4.1 lists the average background for all 3 cases. As is clear from the table, the total background level is consistent with the other two within  $1\sigma$ . Therefore, unless otherwise specified, the total background value derived from all circles will be used for the subsequent analysis and interpretation.

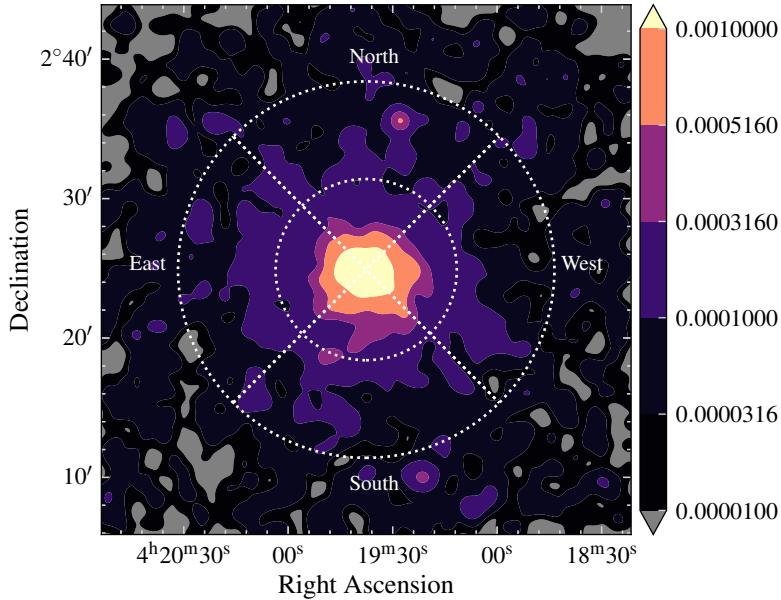


Figure 4.2: Contour plot of the corrected soft-band image smoothed with a Gaussian kernel of 8 pixels. The color bar indicates counts per second. Some deviations from azimuthal symmetry are observed within the outer circle ( $\sim 810''$ ). The inner circle denotes a radius of  $\sim 390''$ . Within this region, an east-west elongation of the core is noticeable. The four wedges labeled north, south, east, and west correspond to the angular (but not radial) ranges used to extract surface brightness.

Region	Background / $\text{cts s}^{-1} \text{arcsec}^{-2}$
North	$(1.03 \pm 0.10) \times 10^{-7}$
South	$(1.27 \pm 0.04) \times 10^{-7}$
Total	$(1.15 \pm 0.14) \times 10^{-7}$

Table 4.1: Surface brightness and associated errors for the different background regions.

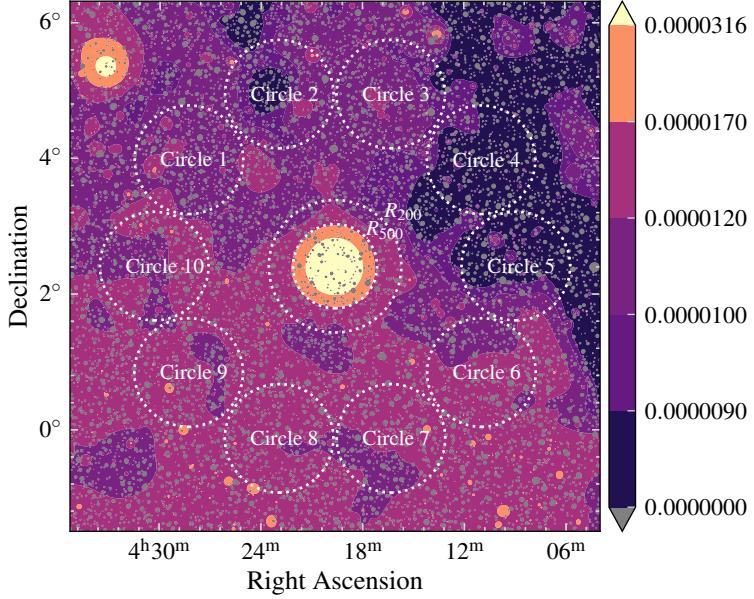


Figure 4.3: Background estimation. Depicted are 10 circular regions utilized to estimate the observed background. Circles 1-5 correspond to the northern background, and Circles 1-6 to the southern background. The total background corresponds to all circular regions. The obtained values can be found in Table 4.1.

A  $\beta$ -model (cf. Section 2.1.3) is employed to characterize the surface brightness profile, given by

$$S_X(r) = S_X(0) \left( 1 + \left( \frac{r}{r_c} \right)^2 \right)^{-3\beta+0.5} + d$$

where  $d$  represents the background emission. The center of each annulus is taken to be  $r$  and the corresponding surface brightness values are utilized. The optimized parameters and  $\chi^2/\text{d.o.f}$  (chi squared per degrees of freedom) are listed in Table 4.2. Figure 4.4 illustrates the surface brightness as a function of radial distance from the center, including both the fitted  $\beta$ -model and the observationally estimated background level.

Parameter	$S_0 / \text{cts s}^{-1} \text{arcsec}^{-2}$	$\beta$	$r_c / ''$	$d / \text{cts s}^{-1} \text{arcsec}^{-2}$
	$(4.1 \pm 0.4) \times 10^{-5}$	$0.478 \pm 0.008$	$60 \pm 5$	$(9.3 \pm 0.4) \times 10^{-8}$
$\chi^2/\text{d. o. f}$	0.96			

Table 4.2: Fit parameters, their errors, and the reduced chi-squared value for the full azimuthal surface brightness profile.

The  $\beta$ -model fit for the surface brightness profile, shown in Figure 4.4, yields  $\beta = 0.478 \pm 0.008$  and a core radius  $r_c = (60 \pm 5)'' \approx (15 \pm 2) \text{kpc}$  with  $\chi^2/\text{d.o.f.} = 0.96$ . This result is consistent

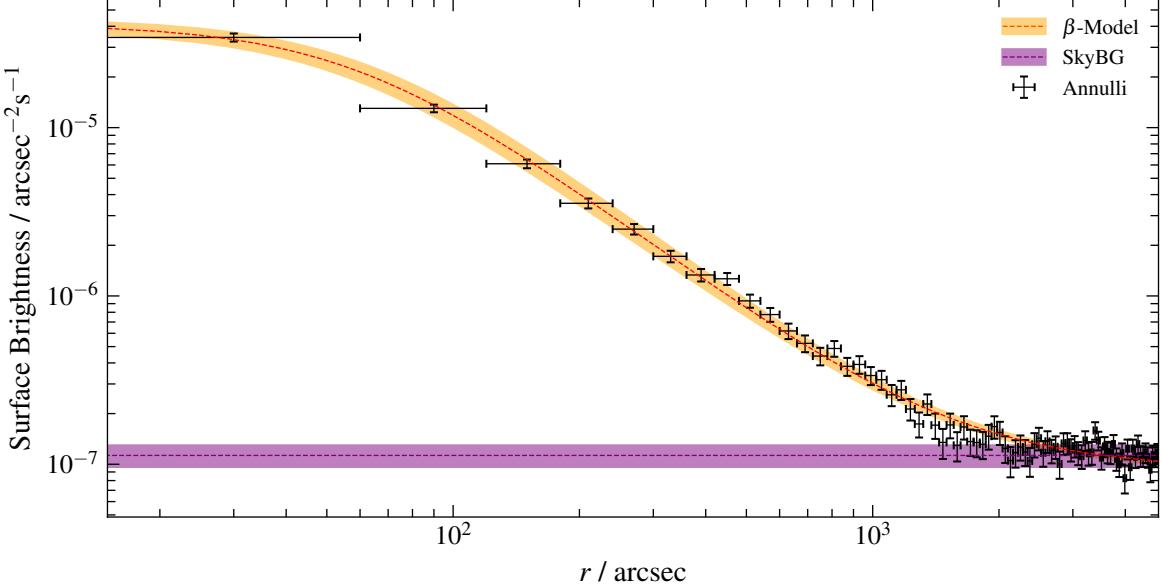


Figure 4.4: Surface Brightness within each annulus (black crosses) as a function of distance from the determined SB-center including error bars. The dashed orange line represents the best  $\beta$ -model fit, while the purple line indicates the observationally determined background. The shaded regions correspond to their respective  $1\sigma$  intervals.

with previous studies: (Kawaharada, Makishima, Takahashi, et al., 2003) report  $\beta = 0.47$  and  $r_c \approx (16 \pm 1)$  kpc using ASCA data; (Kawaharada, Makishima, Kitaguchi, et al., 2009) use a double  $\beta$ -model to find  $\beta = 0.50 \pm 0.05$  and  $r_c \approx (20 \pm 3)$  kpc for the outer component from XMM-Newton data; (Sun et al., 2003) also derive  $\beta \sim 0.48$  and a slightly higher  $r_c \approx 26$  kpc for the outer component using Chandra data; (Reiprich and Böhringer, 2002), although, report a somewhat higher  $\beta = 0.554^{+0.049}_{-0.037}$  using ROSAT. The surface brightness profile also highlights that the group emission decreases to the background level at approximately  $R_{500}$ , as observed in Section 4.1. The observationally estimated total background in Table 4.1, however, is approximately 24% higher than the background obtained by the  $\beta$ -model fit (cf. Table 4.2) at  $1.5\sigma$ . Although this discrepancy is not too alarming, it does indicate challenges associated with the complex background, implying that they could have been chosen more carefully.

A double  $\beta$ -model fit was not successful. This can be attributed to the angular resolution of the 1' width annulli, limiting the ability to resolve the inner core emission. Performing a  $\beta$ -model fit with finer annuli was attempted, but this did not significantly improve the description of the surface brightness profile within the core and resulted in poor error estimation in the outskirts due to low count rates. Indeed, (Kawaharada, Makishima, Kitaguchi, et al., 2009) find an inner core component of  $r \sim 3$  kpc corresponding to  $\approx 12''$  which is comparable to the eROSITA angular resolution  $\sim 15''$  (Predehl et al., 2021). A successful two- $\beta$  model, however, might be possible by fixing the inner core radii or convolving the model with the eROSITA point spread function, but this was not attempted within the scope of this thesis.

### 4.2.1 $\beta$ -Model and Residual Image

Using the parameters from Table 4.2, a  $\beta$ -model image (Fig. 4.5) and a residual image (Fig. 4.6) are created. The  $\beta$ -model image is generated by distributing the one-dimensional  $\beta$ -model across the image. This involves computing the distance from the surface brightness center to each coordinate  $(x, y)$  in the image and then scaling the output of the  $\beta$ -model by the eROSITA pixel area of  $16\text{arcsec}^2$ . The residual image (res. img) was obtained by subtracting the scaled  $\beta$ -model from the corrected soft-band image (corr. img). Hence,

$$\text{res. img}(x, y) = \text{corr. img}(x, y) - S \left( \sqrt{x^2 + y^2} \right) \cdot \text{pix. area.}$$

The residual image indicates that the  $\beta$ -model overestimates emission within  $R_{500}$ , while underestimating emission at the center significantly (highlighted by a green arrow in Figure 4.6). The latter can be attributed to the absence of a second  $\beta$ -component, as a single  $\beta$ -model often inadequately describes the inner core emission, which is often higher-than-expected. (Sun et al., 2003) report, however, that even a two  $\beta$ -fit is not sufficient to describe the surface brightness within the central 1 kpc. Disregarding the excess emission within the core, no significant substructure or features can be noticed from the residual image.

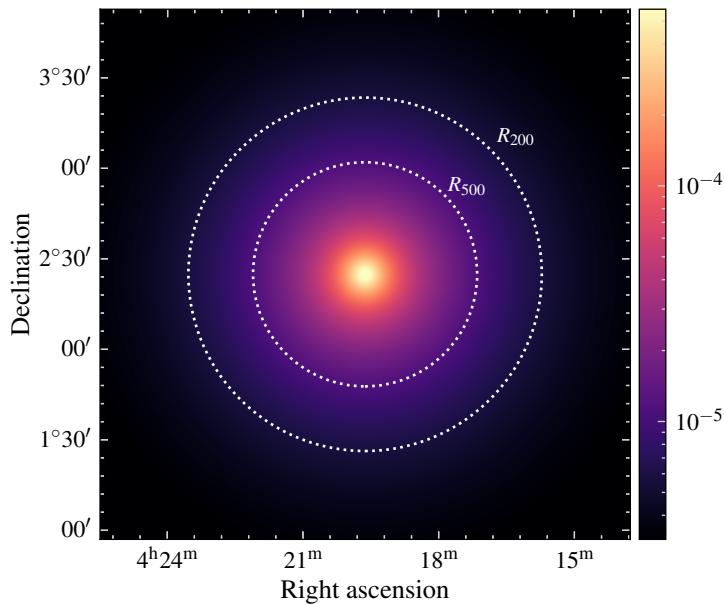


Figure 4.5:  $\beta$ -model image of NGC1550 utilizing the parameters in 4.2. The color bar represents counts per second.

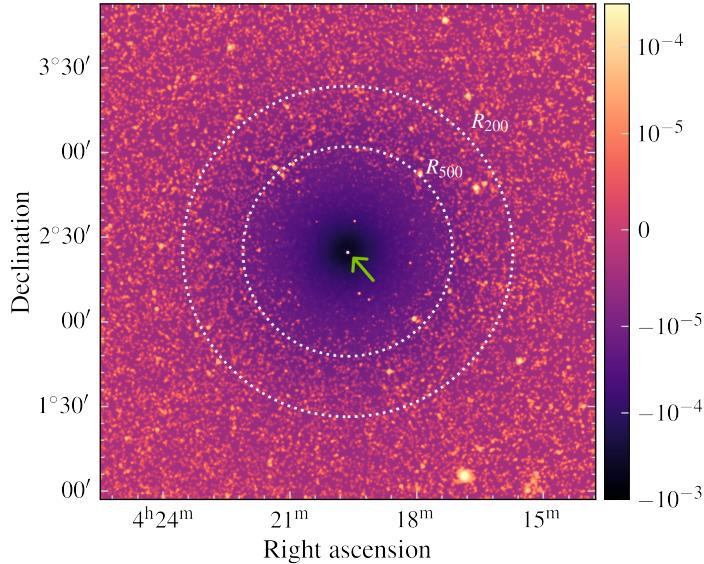


Figure 4.6: Residual image of NGC1550 derived by subtracting  $\beta$ -model image from the corrected soft-band image. The color bar represents counts per second and the image has been smoothed with a Gaussian-kernel of 5 pixels. The residual image indicate that the  $\beta$ -model slightly overestimates the emission around the core, due to the inadequacy of a single  $\beta$ -model in capturing the core's emission.

### 4.3 Sectoral Surface Brightness Analysis

Thus far, the surface brightness analysis has encompassed the entire azimuthal extent of the group. By analyzing surface brightness within specific sectors and comparing them to the full azimuthal profile, elongation in certain directions can be revealed. Various authors have reported an east-west elongation at the center of NGC1550 (Sun et al., 2003, Kolokythas et al., 2020), making it particularly interesting to determine whether this elongation extends into the outer regions of the group. To investigate this, four sectoral regions (north, south, east, and west) are divided into annuli, as shown in Figure 4.2. The center, width and extension of the annuli remain as described in Section 4.1, and surface brightness calculations and  $\beta$ -model fitting follow the same methodology.

Each region is initially fitted individually. Thereafter, the north and south sectors are combined for fitting, as are the east and west sectors. Figure 4.7(a) illustrates the comparison between the east and west profiles and Figure 4.7(b) compares the north and south profiles. The top graph in each figure shows the surface brightness profiles, including the observationally determined sky background and the corresponding  $\beta$ -model. The lower two graphs display the relative surface brightness differences of each region ( $SB_{sec}$ ) compared to the full azimuthal surface brightness profile ( $SB_{an}$ ). The error bars represent the error of the relative difference computed by Gaussian error propagation. The fit parameters for each  $\beta$ -model are listed in Table 4.3.

The fits are of lower quality when compared to the full azimuthal  $\beta$ -model fit ( $\chi^2/\text{d.o.f.} \lesssim 0.8$ ), suggesting possible overfitting or, more likely, an overestimation of the errors as attempts using wider  $2'$  annuli resulted in significantly improved  $\chi^2/\text{d.o.f.} \gtrsim 0.95$  for the west, north, and combined

east-west profiles, and a modest improvement for the remaining sectors. However, since the full azimuthal analysis was only performed with the smaller annuli and the resulting fit parameters were similar across both cases, the smaller annuli will be utilized for the following comparison.

When inspecting the western region, the emission dip observed in the western region between approximately  $390''$  and  $810''$  – as noted in section 4.1 – is noticeable in relation to the full azimuthal surface brightness profile. However, these deviations are not quantitatively significant and may be attributed to statistical fluctuations. For the northern and southern sectors, a significant deviation is noted at the first data point. This discrepancy is likely due to a suboptimal choice of the SB-center and is probably not physically significant, which shall be elaborated further below. In all sectors, however, no notable deviation from the full azimuthal profile is observed beyond  $390'$ . The combined east-west and north-south emissions are compared in Figure 4.8 with the  $\beta$ -model parameters being listed in Table 4.3. Here, too, no significant deviation from the full azimuthal profile is observed beyond the first data points, indicating that the east-west elongation observed within  $390''$  in Section 4.1 does not extend towards the outskirts of the group. The east-west elongation within  $390''$  can not be quantitatively confirmed, as only the third annulus shows an appreciable deviation compared to the combined north-south profile. This is complicated further by the potentially misplaced SB-center (*vide infra*) and the wide annuli, which set an arbitrary data resolution.

Parameter	$S_0 / \text{cts s}^{-1} \text{arcsec}^{-2}$	$\beta$	$r_c / ''$	$d / \text{cts s}^{-1} \text{arcsec}^{-2}$	$\chi^2/\text{d. o. f}$
East-West	$(2.88 \pm 0.27) \times 10^{-5}$	$(5.08 \pm 0.12) \times 10^{-1}$	$(8.60 \pm 0.80) \times 10^1$	$(1.02 \pm 0.05) \times 10^{-7}$	0.83
North-South	$(6.20 \pm 0.71) \times 10^{-5}$	$(4.58 \pm 0.07) \times 10^{-1}$	$(4.01 \pm 0.43) \times 10^1$	$(8.45 \pm 0.46) \times 10^{-8}$	0.71
West	$(3.42 \pm 0.42) \times 10^{-5}$	$(5.20 \pm 0.16) \times 10^{-1}$	$(8.11 \pm 0.98) \times 10^1$	$(9.63 \pm 0.60) \times 10^{-8}$	0.77
East	$(3.18 \pm 0.43) \times 10^{-5}$	$(4.87 \pm 0.15) \times 10^{-1}$	$(7.22 \pm 0.97) \times 10^1$	$(1.02 \pm 0.07) \times 10^{-7}$	0.82
North	$(2.23 \pm 0.35) \times 10^{-5}$	$(4.49 \pm 0.14) \times 10^{-1}$	$(6.6 \pm 1.1) \times 10^1$	$(7.89 \pm 0.78) \times 10^{-8}$	0.78
South	$(1.10 \pm 0.20) \times 10^{-4}$	$(4.69 \pm 0.10) \times 10^{-1}$	$(3.32 \pm 0.53) \times 10^1$	$(8.69 \pm 0.64) \times 10^{-8}$	0.86

Table 4.3: Fit parameters, their errors, and the reduced chi-squared value for the  $\beta$ -model fit of the north, south, east and west sectors and the combined north-south and east-west sectors.

It is notable that the  $\beta$  parameter and the core radius  $r_c$  are not always consistent when comparing the sectors against each other and with the full azimuthal profile. For example, the core radius  $r_c$  for the combined east-west region is more than twice as large as that for the combined north-south region, and the  $\beta$  value is approximately 10% larger, with discrepancies at  $> 5\sigma$  and  $> 3\sigma$ , respectively. In fact, most of the  $\beta$ -values and core radii are inconsistent within  $1\sigma$  when comparing the sectors against each other. This discrepancy may indicate a poor SB-center choice, which is biasing the surface brightness towards the south. For instance, the southern sector's normalization factor  $S_0$  shows a relative difference of almost 80% compared to the north sector, with a  $4.32\sigma$  deviation. Consequently, the estimated SB-center at  $64.909^\circ$  and  $2.414^\circ$  is likely too far north, leading to an overestimation of southern emission in the initial data points, which is clearly visible in Figure 4.7(b). This issue is probably exacerbated by the wide  $1'$  annuli, meaning that excess emission of the first data point might be significantly affecting the fits.

Regarding the background, the fitted background level is in agreement with the total background level for the east, west and combined east-west regions. However, for the north and south regions, the fitted background level underestimates the observationally determined north and south backgrounds

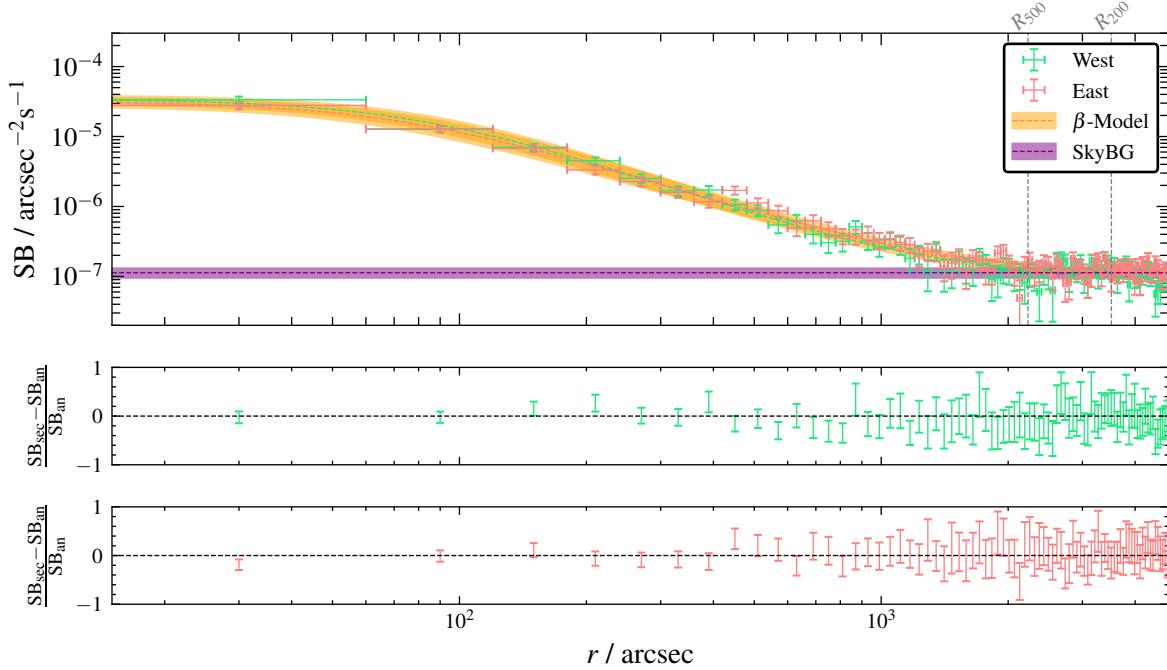
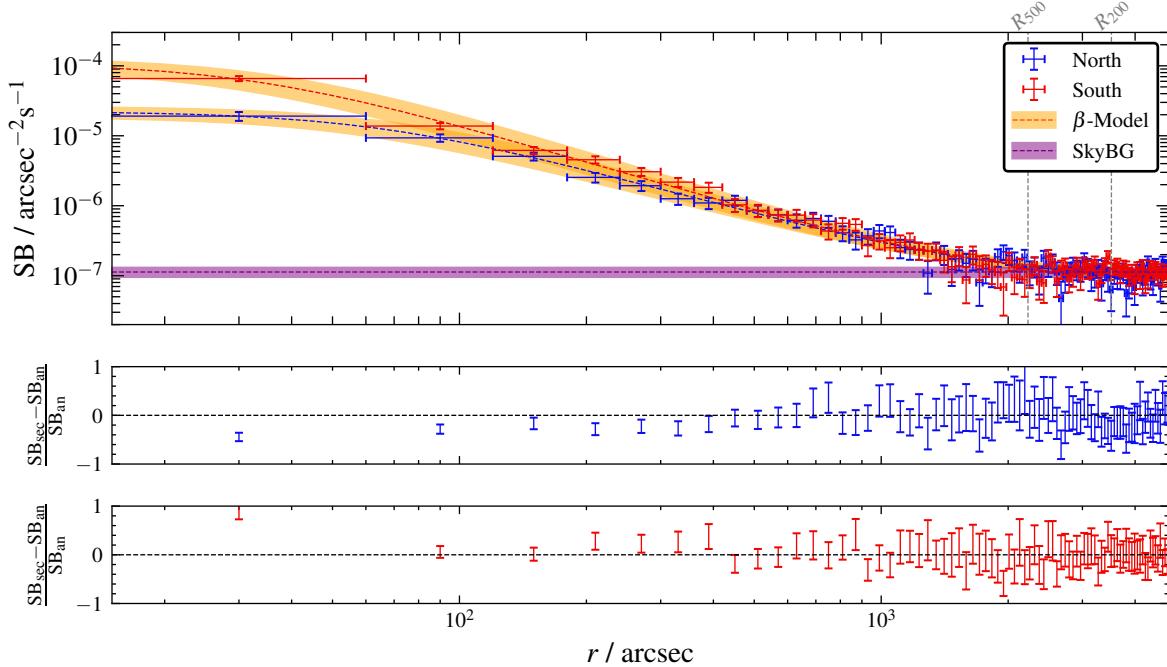

 (a)  $\beta$ -model fits for the West and East regions.

 (b)  $\beta$ -model fits for the North and South regions.

Figure 4.7:  $\beta$ -model fits for the West-East and North-South regions. *Top of each panel:* Surface brightness profiles, including the observationally determined sky background. *Bottom of each panel:* Relative surface brightness differences between each region ( $\text{SB}_{\text{sec}}$ ) and the full azimuthal surface brightness profile ( $\text{SB}_{\text{an}}$ ).

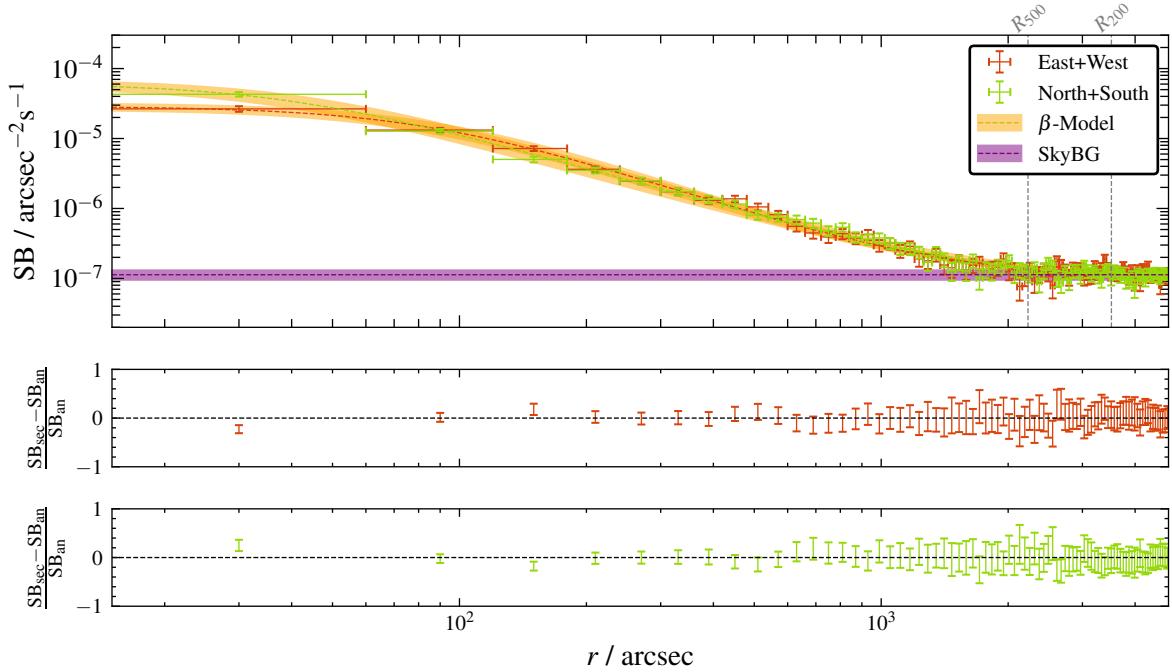


Figure 4.8:  $\beta$ -model fits for the combined east-west and north-south regions. The surface brightness profile is shown in the top graph and the lower two graphs display the relative surface brightness differences between each region ( $\text{SB}_{\text{sec}}$ ) and the full azimuthal surface brightness profile ( $\text{SB}_{\text{an}}$ ). Error bars are derived from Gaussian error propagation. Excluding the first data point, both combined surface brightness profiles show no significant deviation from the full azimuthal profile and are entirely consistent with each other, indicating no significant east-west elongation in the outskirts of the galaxy group.

by 31% and 46%, at  $1.9\sigma$  and  $0.98\sigma$ , respectively. This discrepancy might be simply due to the inherently complex background of the Orion-Eridanus superbubble, though it can also be taken as an indication that the circles chosen for background estimation in Figure 4.3 are not far enough from the SB-center.

In general, such deviations in fit parameters would align well with the common notion that galaxy groups never exhibit perfect azimuthal symmetry, since infalling matter inevitably causes irregularities. Due to the aforementioned issue regarding the SB-center, one can not conclusively take the fit deviations of the sectoral surface brightness analysis as confirmation for this view. Regardless, the sectoral surface profiles confirm that no significant deviations from azimuthal symmetry can be established in the outskirts, were the above mentioned issues of the SB-center and annuli width are largely insignificant, and contamination from the background dominates.

# CHAPTER 5

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## Conclusion

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The aim of this thesis was the characterization of the X-ray morphology of the galaxy group NGC1550 through detailed surface brightness analysis. After correcting and reducing the raw photon data obtained from eRASS:1 in Chapter 3, Chapter 4 began with qualitative inspection of the corrected data. A contour plot of the corrected, cheese-masked wavelet filtered image in the soft-band with an overlaid galaxy distribution was provided (Figure 4.1). From this figure, it was visually established that the galaxy group NGC1550 appears relaxed and spherically symmetrical. Furthermore, using a contour plot of the corrected soft-band image (Figure 4.2), some indications of asymmetry were visually identified within  $810''$ . Emission was found to drop off to background levels at  $R_{500}$  and a complex background was noted, which can be attributed to the Orion-Eridanus superbubble. No obvious correlation between the X-ray contours and the galaxy distribution could be ascertained.

To more accurately characterize the morphology of NGC1550 in the X-ray, surface brightness analysis was performed. A beta model was employed to characterize the full azimuthal profile, as shown in Figure 4.4. The model fitting yielded a  $\beta$ -value of  $0.478 \pm 0.008$  and a core radius of  $r_c = 60'' \approx (15 \pm 2) \text{ kpc}$  with a reduced chi-squared value of 0.96, indicating a good fit. This result was consistent with the findings of previous studies, showing that the eROSITA view of NGC1550 aligns with previous findings. The full azimuthal surface brightness profile showed no significant deviations from spherical symmetry beyond. Moreover, the sectoral surface brightness analysis across the north, south, east, and west sectors, as well as the combined north-south and east-west sectors, also revealed no significant deviations from azimuthal symmetry beyond  $390''$ .

The residual image indicated that the beta model underestimates the core emission, as is common, but no other features were detected. Attempts to use a two beta model were unsuccessful, likely due to the large width of the annuli and the small inner core component, which, from previous studies (Kawaharada, Makishima, Kitaguchi, et al., 2009), is of order of eROSITA's angular resolution.

Within  $390''$ , the qualitatively observed east-west elongation could not be quantitatively confirmed by the surface brightness analysis. Although the sectoral surface brightness profiles did not differ significantly from the azimuthal profile, their fit parameters often varied significantly and the individual beta fits of each sector were of slightly lower quality ( $\chi^2 \lesssim 0.8$ ). It was speculated that these issues were caused by the fairly wide  $1'$  annulli and due to a poor choice of the SB-center, resulting in a significant discrepancy between the north and south emission for the first data point at  $\sim 60''$ .

The surface brightness analysis also highlighted a slight discrepancy between the observationally estimated backgrounds and the fitted background levels, suggesting that background estimation could have been performed with greater care.

## 5.1 Outlook

After analyzing the results of this thesis, some suggestions for improvements can be made. The SB-center could have been more precisely determined, either manually or, better, by a two-dimensional beta fit of the full azimuthal profile, allowing it to be freely fit.

Furthermore, given that only eRASS:1 data was utilized, several features that were qualitatively observed, such as a slight emission dip in the western sector between  $\sim 390''$  and  $\sim 810''$ , could not be quantitatively verified. With additional data, such as from eRASS:4, these features could be more accurately interpreted and their significance better assessed. The increased data would also allow for finer annuli in the surface brightness analysis due to the higher count rates in each annulus bin, facilitating feature identifications.

Furthermore, future work could leverage eROSITA to perform detailed spectroscopy of the outskirts of NGC1550. X-ray spectroscopic measurements can provide insights into the enrichment of the intracluster medium (ICM) and the historical processes of galaxy groups or clusters Liu et al., 2020. Understanding the distinctions in enrichment processes between galaxy groups and clusters remains an ongoing challenge. eROSITA's wide field of view presents a unique opportunity to study these processes in greater detail, enhancing our knowledge of the chemical composition and evolutionary history of NGC 1550 and similar groups.

# **Appendix**

# APPENDIX A

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## Appendix

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### A.1 PIB Map

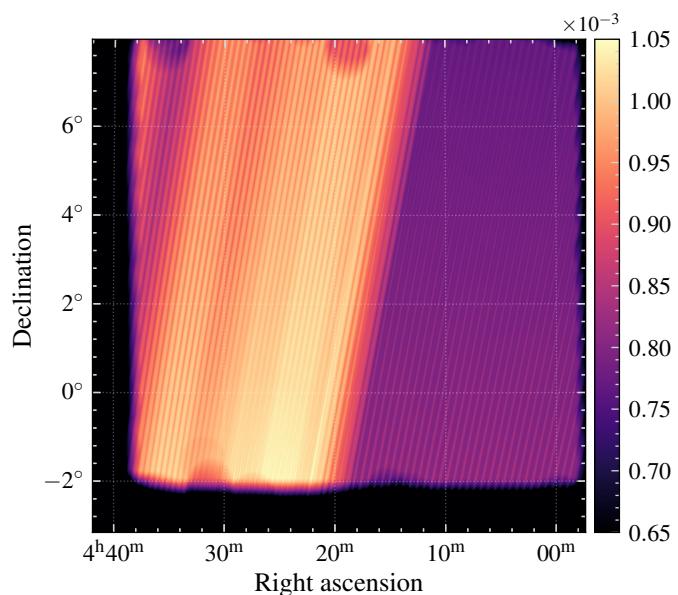


Figure A.1: PIB map created by co-adding individual background maps for each telescope module (TM).

### A.2 Exposure maps

TM	R	$H_{\text{obs}}$
1	1.30	9512
2	1.36	8362
3	1.35	9187
4	1.24	10177
5	0.85	10817
6	1.29	9144
7	0.82	10214

Table A.1: Values of  $R$ , and  $H_{\text{obs}}$  utilized for the PIB-correction in Section

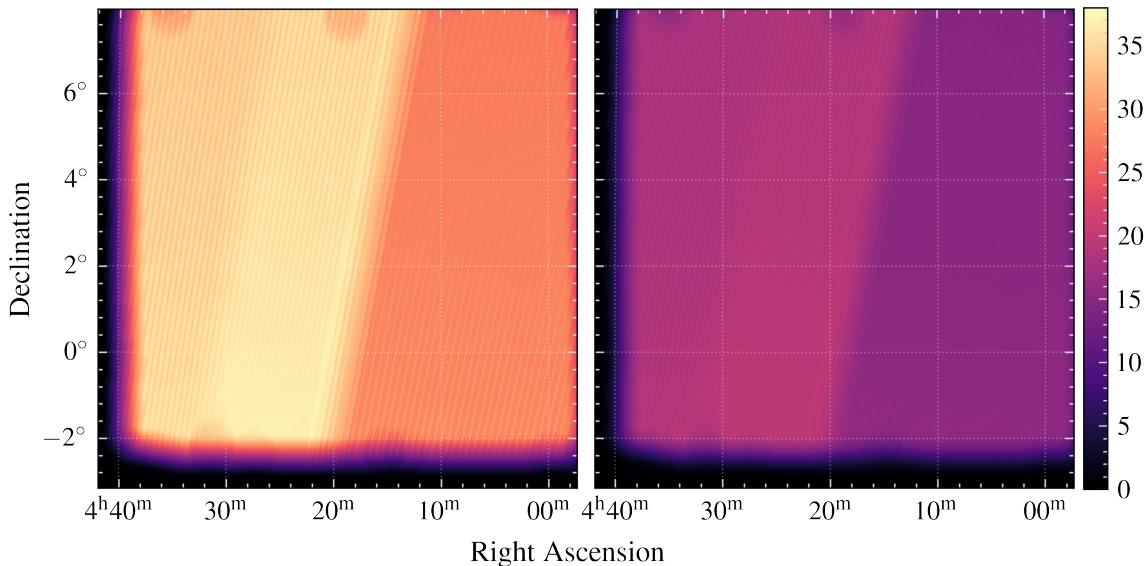
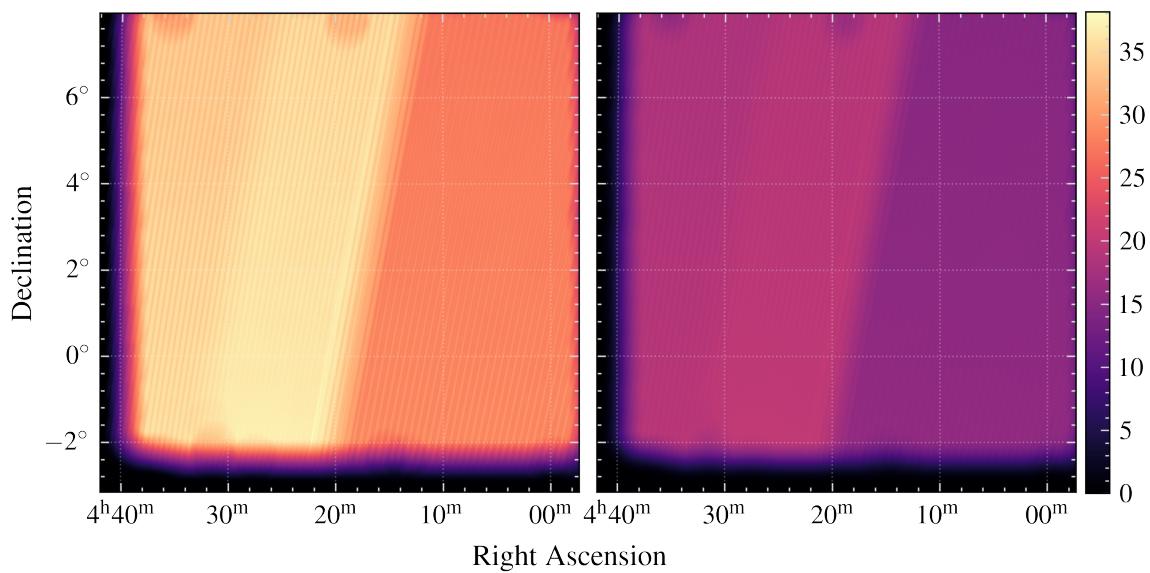
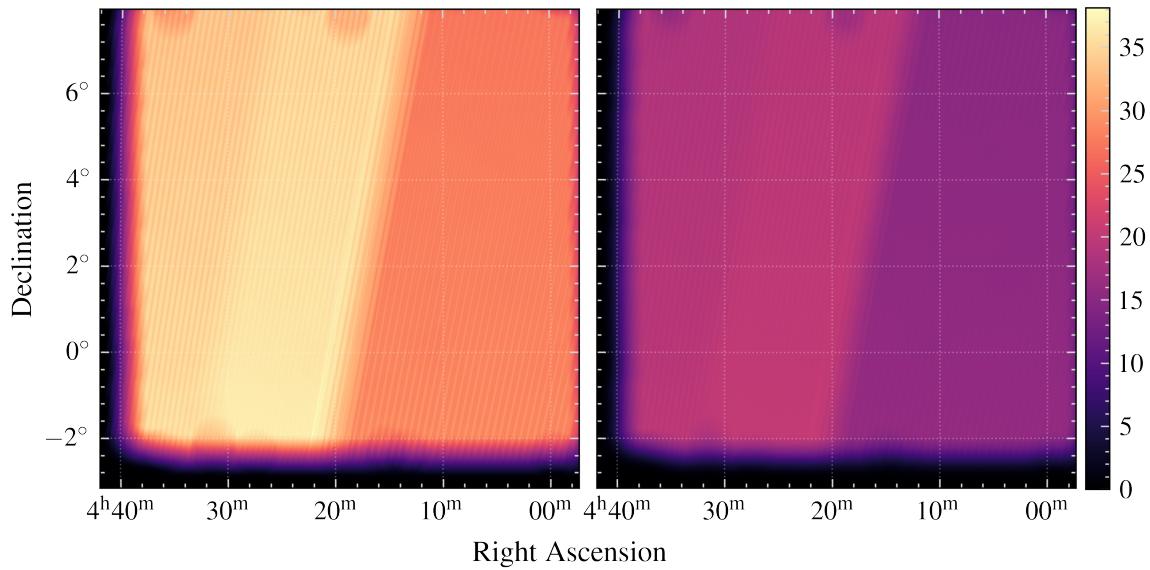


Figure A.2: TM1



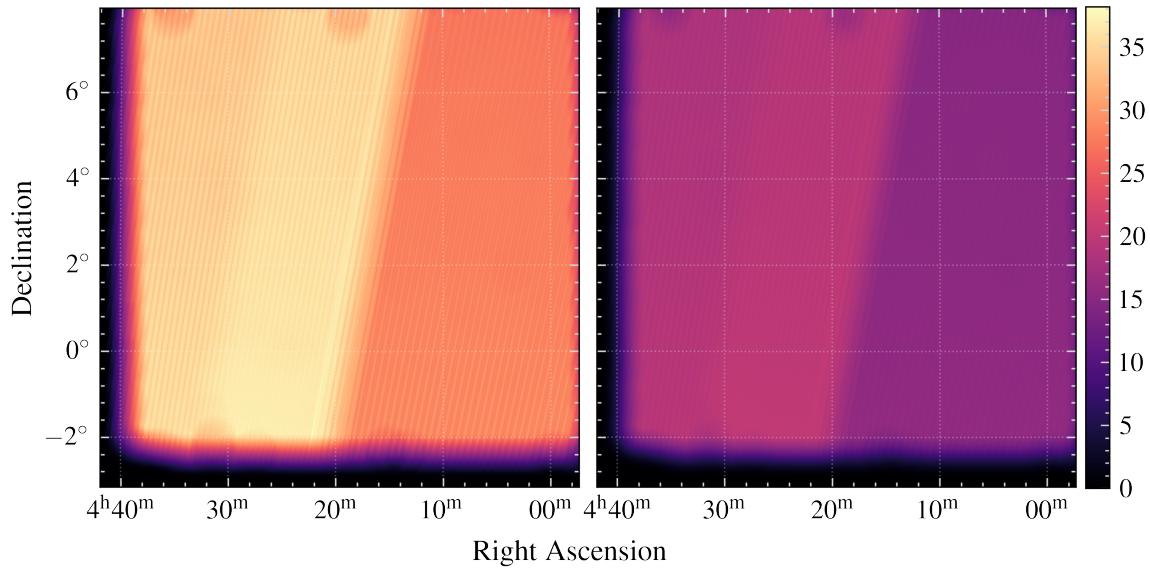


Figure A.5: TM4

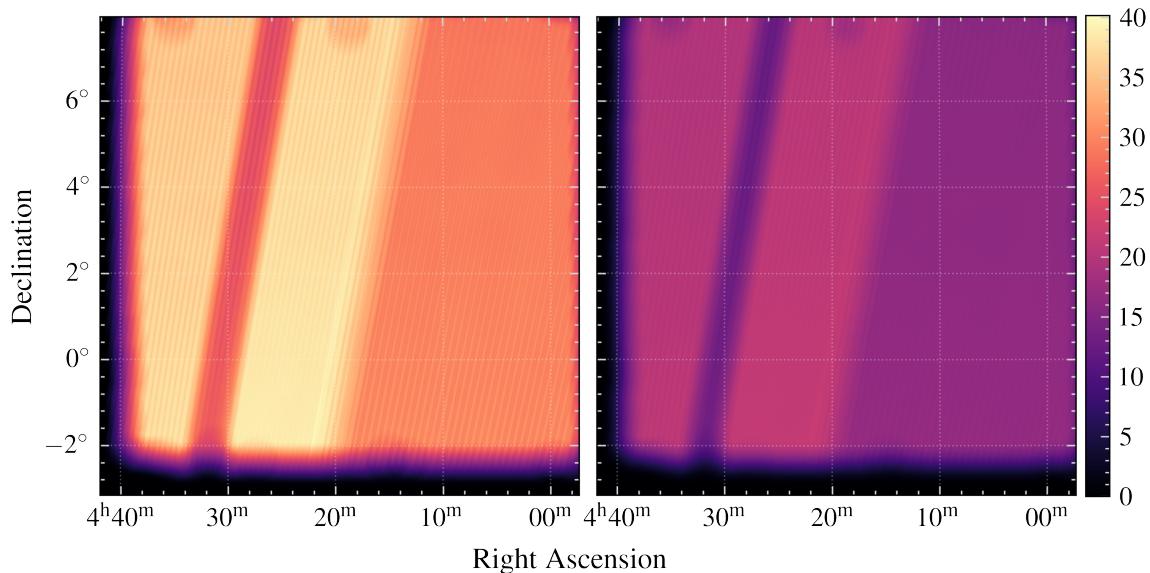


Figure A.6: TM5

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## Appendix A Appendix

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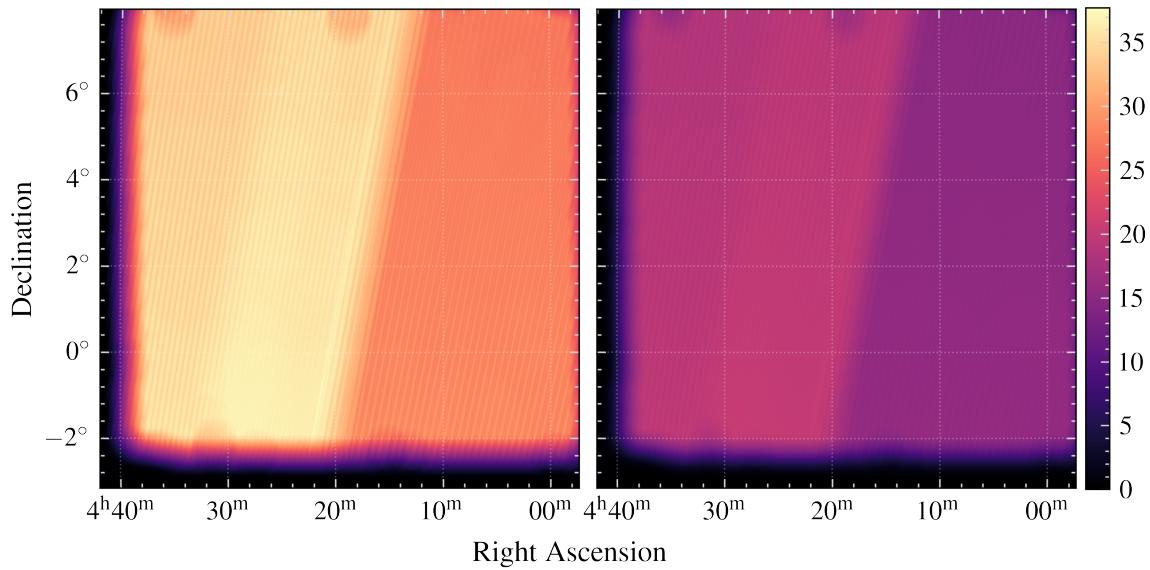


Figure A.7: TM6

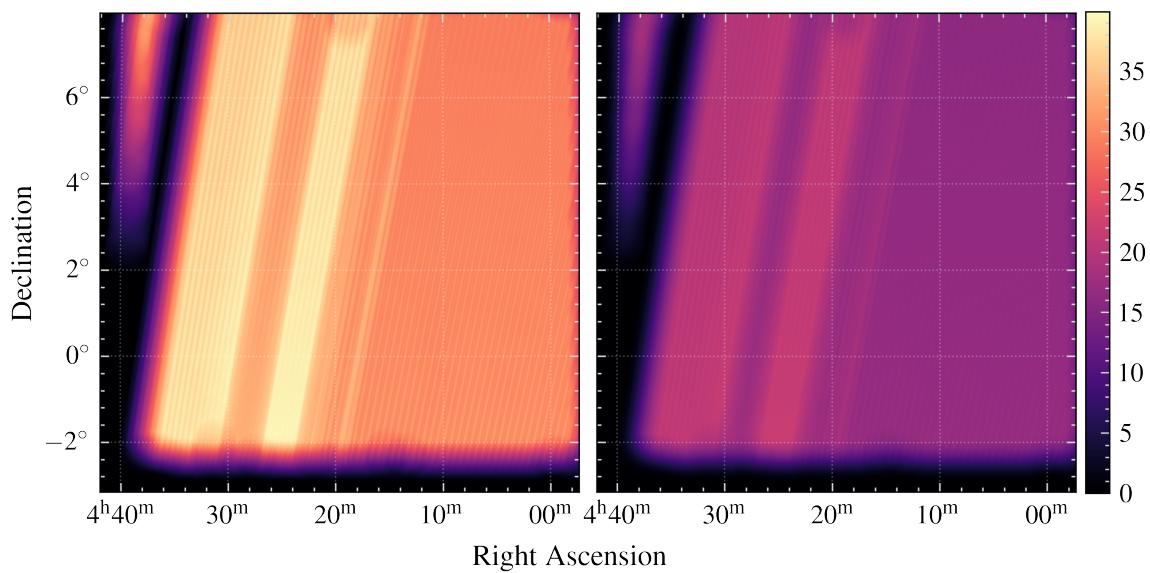


Figure A.8: TM7

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## Appendix A Appendix

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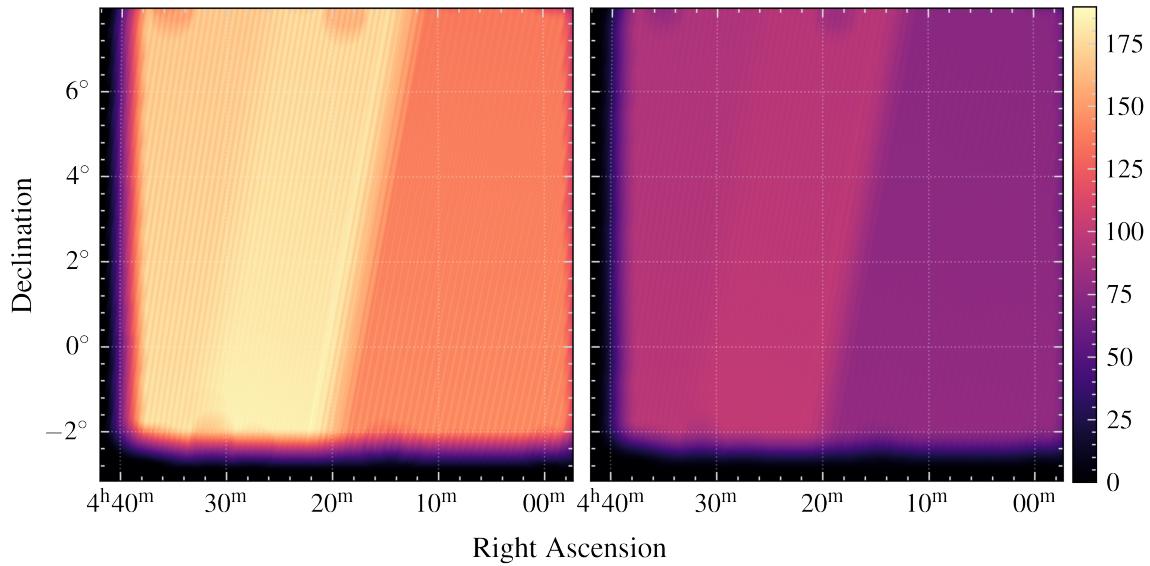


Figure A.9: TM8

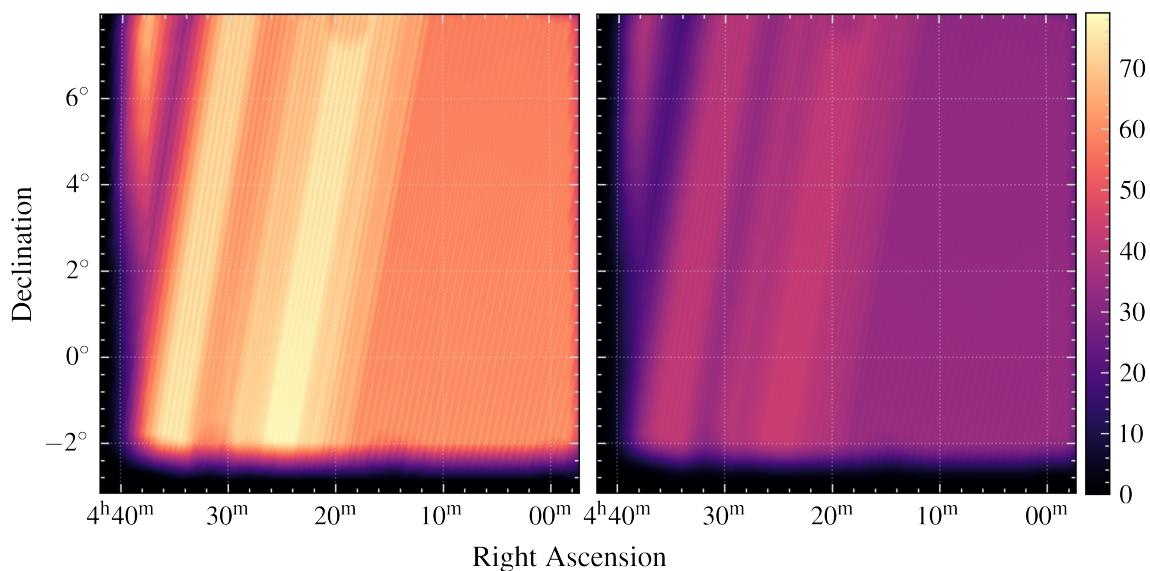


Figure A.10: TM9

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