

Technical note for PTPlot

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1 Input parameters

There are five input parameters when plotting a single parameter point at ptplot.org.

Wall velocity v_w is the bubble wall velocity, or more accurately, the velocity of the phase interface after nucleation in the rest frame of the plasma far away from the wall. It can be given values between $0 < v_w \leq 1$.

Phase transition strength α_θ is defined as the ratio of the trace of the energy momentum tensor and the energy density in the symmetric phase. It can be set freely, but is usually of the order of $\alpha_\theta \sim 10^{-2} - 10^0$.

Inverse phase transition duration β/H_* is the inverted duration of the phase transition β , which is often expressed relative to the Hubble rate H_* as a dimensionless quantity. We assume the duration to be short enough that the expansion of the universe can be neglected, $\beta/H_* \gtrsim 1$, but otherwise it can be set freely. We assume the Hubble rate to be the same right after the phase transition and at nucleation, $H_n \simeq H_*$.

Transition temperature T_* is the temperature of the universe right after the phase transition. We expect this to be around $T_* \sim 150 - 200$ GeV, close to the Standard Model crossover temperature of ~ 160 GeV [4], but it can be set freely based on the chosen model. We assume this to be the same as the nucleation temperature, $T_n \simeq T_*$.

Degrees of freedom g_* is the number of relativistic degrees of freedom at the time of the phase transition. It is related to the temperature and energy density of the universe at that time, not the dynamics of the transition itself. It is highly constrained to the Standard Model value of $g_* = 106.75$ [4] when we are close to the critical temperature, but it can be varied according to the chosen model.

It should be noted that currently the user is responsible for making sure that the parameters given to PTPlot are self-consistent and produce physical results.

2 Energy budget computations in `espinosa.py`

Analytical fits for the efficiency coefficient. We compute the efficiency coefficient $\kappa(v_w, \alpha)$ using analytical fits split into three regions. The approximations for the fits and corresponding four boundary cases are taken from appendix A of [1]. We define κ as a function of the wall velocity v_w and phase transition strength α , with the speed of sound $c_s = \frac{1}{\sqrt{3}}$. It describes the fraction of kinetic energy transformed into vacuum energy.

The four boundary cases for small velocities (A), the speed of sound (B), Jouguet detonations (C), and large velocities (D) are

$$\kappa_A \simeq v_w^{6/5} \frac{6.9\alpha}{1.36 - 0.037\sqrt{\alpha} + \alpha} \quad (v_w \ll c_s) \quad (1)$$

$$\kappa_B \simeq \frac{\alpha^{2/5}}{0.017 + (0.997 + \alpha)^{2/5}} \quad (v_w = c_s) \quad (2)$$

$$\kappa_C \simeq \frac{\sqrt{\alpha}}{0.135 + \sqrt{0.98 + \alpha}} \quad (v_w = \xi_J), \text{ where } \xi_J = \frac{\sqrt{\frac{2}{3}\alpha + \alpha^2} + \sqrt{\frac{1}{3}}}{1 + \alpha} \quad (3)$$

$$\kappa_D \simeq \frac{\alpha}{0.73 + 0.083\sqrt{\alpha} + \alpha} \quad (v_w \rightarrow 1). \quad (4)$$

The analytical fits for subsonic deflagrations, detonations and supersonic deflagrations (hybrids), respectively, are

$$\kappa(v_w \lesssim c_s) \simeq \frac{c_s^{11/5} \kappa_A \kappa_B}{(c_s^{11/5} - v_w^{11/5}) \kappa_B + v_w c_s^{6/5} \kappa_A} \quad (5)$$

$$\kappa(v_w \gtrsim \xi_J) \simeq \frac{(\xi_J - 1)^3 \xi_J^{5/2} v_w^{-5/2} \kappa_C \kappa_D}{[(\xi_J - 1)^3 - (v_w - 1)^3] \xi_J^{5/2} \kappa_C + (v_w - 1)^3 \kappa_D} \quad (6)$$

$$\kappa(c_s < v_w < \xi_J) \simeq \kappa + (v_w - c_s) \delta\kappa + \frac{(v_w - c_s)^3}{(\xi_J - c_s)^3} [\kappa_C - \kappa_B - (\xi_J - c_s) \delta\kappa], \quad (7)$$

$$\text{where } \delta\kappa \simeq -0.9 \log \frac{\sqrt{\alpha}}{1 + \sqrt{\alpha}} \text{ is the first derivative of } \kappa \text{ at } v_w = c_s. \quad (8)$$

We note that [1] uses the parameter α_N when defining the analytical fits. We use $\alpha \equiv \alpha_\theta$ instead, which at the current state of knowledge is sufficient for estimating the gravitational wave power spectrum [2].

The root-mean-square (RMS) fluid velocity. The kinetic energy fraction of the fluid K can be determined in terms of the RMS fluid velocity \bar{U}_f and the mean adiabatic index Γ , or in terms of κ and α [2] as

$$K = \Gamma \bar{U}_f^2, \quad K = \frac{\kappa \alpha}{1 + \alpha}. \quad (9)$$

We then compute the RMS fluid velocity, with $\Gamma \simeq \frac{4}{3}$ for an ultra-relativistic fluid, as

$$\bar{U}_f = \sqrt{\frac{\frac{3}{4} \kappa(v_w, \alpha)}{1 + \alpha}}. \quad (10)$$

3 Power spectra computations in calculate_powerspectrum.py

Ansatz for the sound wave power spectrum. The spectral shape of the sound waves is [2, 3]

$$C_{\text{sw}}(f) = \left(\frac{f}{f_{p,0}} \right)^3 \left[\frac{7}{4 + 3 \left(\frac{f}{f_{p,0}} \right)^2} \right]^{\frac{7}{2}}, \quad (11)$$

with the peak frequency today of

$$f_{p,0} = 26 \times 10^{-6} \left(\frac{1}{H_* R_*} \right) \left(\frac{z_p}{10} \right) \left(\frac{T_*}{10^2 \text{ GeV}} \right) \left(\frac{h_*}{100} \right)^{\frac{1}{6}} \text{ Hz}. \quad (12)$$

The peak angular frequency in units of the mean bubble separation $z_p = (k R_*)_{\text{max}} \simeq 10$ based on simulations [2, 3]. We assume that $H_n \simeq H_*$ and $T_n \simeq T_*$. The relation between the mean bubble separation R_* and the inverse phase transition duration β is [3]

$$R_* = (8\pi)^{\frac{1}{3}} \frac{v_w}{\beta}. \quad (13)$$

We compute the gravitational wave power spectrum from sound waves for a given frequency f as [3]

$$h^2\Omega_{\text{gw}} = h^2 2.061 F_{\text{gw},0} \Gamma^2 \bar{U}_f^4 (H_* R_*) \tilde{\Omega}_{\text{gw}} C_{\text{sw}}(f), \quad (14)$$

where

$$F_{\text{gw},0} = (3.57 \pm 0.05) \times 10^{-5} \left(\frac{100}{g_*} \right)^{\frac{1}{3}}, \quad (15)$$

$$\tilde{\Omega}_{\text{gw}} = 0.012. \quad (16)$$

We use the Planck best-fit value of $h = 0.678 \pm 0.009$ at the time [3]. This ansatz is in agreement with Eq. (2) in the erratum of [3], which is written without the customary factors of h^2 , and Eq. (33) in the erratum of [4]. It does not agree with Eq. (29) in [2] unless $c_s = \frac{1}{3}$, which could be considered a typing error.

For a conserved estimate of the sound wave power spectrum we take the shock time to be no larger than one,

$$H_* \tau_{\text{sh}} = \frac{H_* R_*}{\bar{U}_f} < 1. \quad (17)$$

Ansatz for the turbulence power spectrum. There is also an option to calculate the contribution from turbulence to the power spectrum, although it is not used by default, as further studies on turbulence are needed.

The spectral shape from turbulence is [4]

$$C_{\text{turb}}(f) = \frac{\left(\frac{f}{f_{\text{turb}}} \right)^3}{\left[1 + \left(\frac{f}{f_{\text{turb}}} \right) \right]^{\frac{11}{3}} \left(1 + \frac{8\pi f}{h_*} \right)}, \quad (18)$$

with the peak frequency today of

$$f_{\text{turb}} = 27 \times 10^{-6} \frac{1}{v_w} \left(\frac{\beta}{H_*} \right) \left(\frac{T_*}{100 \text{ GeV}} \right) \left(\frac{g_*}{100} \right)^{\frac{1}{6}} \text{ Hz}, \quad (19)$$

and the reduced Hubble rate at T_* ,

$$h_* = 16.5 \times 10^{-6} \left(\frac{T_*}{100 \text{ GeV}} \right) \left(\frac{g_*}{100} \right)^{\frac{1}{6}} \text{ Hz}. \quad (20)$$

We compute the gravitational wave power spectrum from turbulence as [4]

$$h^2\Omega_{\text{turb}} = 3.35 \times 10^{-4} \frac{H_*}{\beta} \left(\frac{\kappa_{\text{turb}} \alpha}{1 + \alpha} \right)^{\frac{3}{2}} \left(\frac{100}{g_*} \right)^{\frac{1}{3}} v_w C_{\text{turb}} \left(\frac{f}{f_{\text{turb}}} \right). \quad (21)$$

References

- [1] J.R. Espinosa, T. Konstandin, J. M. No, and G. Servant, *Energy budget of cosmological first-order phase transitions*, JCAP **06**, 028 (2010), arXiv:1004.4187 [hep-ph]
- [2] C. Caprini et al., *Detecting gravitational waves from cosmological phase transitions with LISA: an update*, JCAP **03** (2020) 024, arXiv:1910.13125 [astro-ph.CO]
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