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1. RMTABLE STANDARD COLUMN DEFINITIONS

The following text has been mostly taken from the standard definition paper, Van Eck et al. (submitted).

The RMTable standard contains a set of proposed standard columns to guide the development of future RM catalogs. The list of columns is given in Table 1, and specific information about each column is given below. Not all of these columns will be appropriate for all catalogs, and individual catalogs will almost always need to have additional columns with survey- or catalog-specific information, so these columns serve as a starting point for authors to consider while designing their catalogs and pipelines. The columns defined in this standard, and particularly their internal names, should be treated as reserved names for future catalogs to avoid confusion.

All columns have standardized internal names, which are used in the data representation to assign machine-readable names to columns. Relevant columns also have assigned standard units, to avoid problems with propagating units through appropriate metadata, and to ease comparisons between different catalogs. A few columns (those defining the source sky coordinates) are labelled as essential, meaning that they must be present for a catalog to be sensibly interpreted in terms of the standard. All other columns are assigned a default value for missing data, which is usually a floating-point NaN or blank string as appropriate. To encourage standardization across catalogs, certain columns (particularly those defining methods used) come with lists of suggested values. The currently defined lists are given in another document on this repository.

For columns specified as being floating point numbers, it is not specified whether these should be 32- or 64-bit floats. The exception to this is the coordinate columns, which are specified as 64-bit (double precision) floating point in order to accommodate RMs derived from Very Long Baseline Interferometry observations. All uncertainties/errors are expected to be 1σ standard deviations. All string columns should take care not to contain special characters such as tabs or newlines, to avoid causing problems with text storage formats. Where columns are expected to refer to published papers, the suggested format is to use 'bibcodes' created by the SAO/NASA Astrophysics Data System (ADS), with the second preferred option being digital object identifiers (DOIs); these are suggested because they provide an identifier that is guaranteed to be unique and can easily lead to the publication using either the ADS or an internet search engine, with the ADS bibcodes producing typically fewer false hits when searching and being partially human-readable. For cases where the papers are not yet published a short descriptive name can be used as a temporary measure but should be replaced after publication. All columns are defined the same way for both RM synthesis techniques and QU-fitting techniques, although the methods of deriving some quantities may be significantly different.

1.0.1. *Right Ascension, Declination, Galactic longitude and latitude*

The sources' coordinates, in both equatorial and Galactic coordinate systems, in decimal degrees (for ease of coordinate-based selection). Both sets of coordinates are considered essential in order to make it easy to select sources based on location. For equatorial coordinates, the International Celestial Reference System (ICRS) is used as it is the currently adopted standard of the International Astronomical Union.

1.0.2. *Position error*

The 1-dimensional uncertainty in the position, in decimal degrees. Commonly, position errors are reported on both position coordinates, but this neglects the possible covariance between the uncertainties (or equivalently, the orientation of the uncertainty ellipse, if it is not aligned with one of the coordinate axes). This makes it impossible to accurately transform the uncertainties between coordinate systems, or to assess the significance of source position differences if not aligned with the coordinate axes. There is also an ambiguity as to whether errors in RA or Galactic longitude are great-circle distances or coordinate distances. To avoid all of these issues, the standard uses a single value to describe a circular position uncertainty (although this comes with the disadvantage of being less accurate where the position error ellipse is significantly elongated). In cases where the position error ellipse is significantly non-circular, we suggest using the semi-major axis (i.e. the larger error) for this column in order to be conservative with the uncertainty values. The standard does not prevent catalog authors from including either the full position uncertainty ellipse or the errors in the each coordinate, but we suggest that in such cases authors also include this column to maximize ease of combination with other RMTable2023 catalogs.

1.0.3. *RM and error*

The RM and corresponding error, in rad m^{-2}. For Faraday-simple polarized sources, this is simply the measured Faraday depth of the polarized emission. As described above, in the case of Faraday-thick features this is the centroid in Faraday depth or a similar quantity. For partially spatially-resolved sources this is the average over the source, or whatever other characteristic value the authors decide applies to the source. If multiple polarization components are present, each is recorded as a separate row in the table with its own RM, error, and other polarization properties. The catalog value should always represent the astrophysical RM of the full line of sight to the source as measured; RMs with any component or foreground subtraction (e.g., removing the Milky Way contribution) or redshift correction should not be reported. If a catalog author wishes to report a corrected RM, this should be as a separate column; including the 'raw' RM values allows catalog users to perform other types of analysis (e.g., Galactic magnetism studies) or apply a correction of their choice (e.g., using a newer foreground model). The exception to this is the subtraction of ionospheric Faraday rotation (discussed below) as such subtraction is often done early in data processing (to allow time-dependent corrections), preventing any calculation of a non-corrected RM.

1.0.4. *Width in Faraday depth and error*

The width of the polarized feature in the FDF, and corresponding error, in rad m^{-2}. This can be either measured (from RM synthesis) or inferred from a model (QU-fitting). This quantity is zero when an explicitly Faraday-thin model is fit or assumed, and can take other values when either a Faraday thick model is fit (e.g., Ma et al. 2019) or the dispersion in RM-synthesis clean components is measured (e.g., Livingston et al. 2021). It must be noted that there are many (incompatible) measures of width in Faraday depth: the ΔRM of a uniform 'Burn slab' model (Burn 1966) is not equivalent to the Gaussian σ_{RM} of an external Faraday dispersion model (Sokoloff et al. 1998) or to a σ_ϕ^2 derived from RM-synthesis clean components. This is further complicated by cases where two width parameters are included in the same component fit (e.g., O'Sullivan et al. 2017); in such cases authors may need to determine a method of combining these parameters into a single value. In general this quantity will not be comparable between models and between different catalogs, but may be useful to compare sources within the same catalog. Authors should ensure that the width parameter is explicitly defined, and catalog users should refer back to the original publications to find these definitions.

1.0.5. *Faraday complexity flag and metric*

A single parameter for whether the source shows Faraday complexity, which for the purposes of this standard is anything other than being compatible with a single-component Faraday-thin model. An additional value, for sources where the complexity has not been assessed or reported, is also needed. Thus three possible values are needed, and a single character is used with the possible values of 'Y','N', or 'U' for 'Yes', 'No', and 'Unknown', respectively. For sources with two or more RM components or a component with a significant non-zero width in Faraday depth, this flag should automatically become 'Y' (for each row, where multiple components are present). The motivation for including this column is to provide users with an easy way to identify the Faraday simple or complex sources present in the catalog without needing to check for additional components or significant Faraday width.

The metric or test used to assess Faraday complexity is also recorded as a short string. This string could be a reference to a paper describing the metric, or a short description (preferrably 80 characters or less) of the method. A list of currently used or suggested values appears in an accompanying document and will continue to be updated in the online documentation as new values come into use, in order to encourage authors to use the same values when using the same methods.

1.0.6. *RM determination method*

A short string describing the method used to determine the RM from the polarized spectra. This can be useful when using a consolidated catalog to assess differences between measurements of the same source; with QU-fitting this column is used to specify the exact model used. A list of currently used or suggested values appears in an accompanying document and will continue to be updated in the online documentation as new values come into use, in order to encourage authors to use the same values when using the same methods.

1.0.7. *Ionospheric correction method*

A string describing the method used to correct for ionospheric Faraday rotation. This could take the form of the software title, a paper reference, or a short algorithm name. A list of currently used or suggested values appears in an accompanying document and will continue to be updated in the online documentation as new values come into use, in order to encourage authors to use the same values when using the same methods. If no correction is applied, the value should be 'None'; if it is not known if a correction was applied the default value is 'Unknown'. This column can be useful for catalog users to assess the possible presence of systematic RM offsets in the catalog due to ionospheric effects.

1.0.8. *Number of RM components*

An integer giving the number of measured RM components identified in the source. This is intended to indicate when the same source will have additional rows in the table (containing the information on the other components). The default value is '1'.

1.0.9. *Stokes I, Q, U, V and errors*

The four Stokes parameters and their corresponding errors, at a given reference frequency. Stokes I may have a different reference frequency, as it may be derived from multi-frequency synthesis imaging or some other algorithm that is different than how the other Stokes parameters are derived. The Stokes parameters may be either intensities (Jy/beam) or flux densities (solid-angle integrated intensities, Jy). Brightness temperatures are strongly discouraged, as such values are rarely appropriate for discrete sources. For Stokes Q and U, these are the values derived from the Faraday rotation model (whether that model is a QU-fitting model, an RM-Clean (Heald et al. 2009) component model, a fit to the peak in RM-synthesis, or something analogous to these), at the corresponding reference frequency and for only this polarized component, and not the actual channel values.

1.0.10. *Stokes I spectral index and error*

The flux density spectral index (α) in total intensity, following the $I \propto \nu^{+\alpha}$ convention (such that most synchrotron sources will have negative values), and corresponding error. This can be the in-band spectral index or that derived with an additional band. Higher order (curvature) terms are not included in this standard, but where authors choose to include them we suggest that the mathematical definition of those terms should be explicitly described. Providing the spectral index can be very useful for users looking to classify sources or select certain source populations.

1.0.11. *Reference frequency for Stokes I*

The frequency, in Hz, for which the Stokes I intensity or flux density applies, as well as for the spectral index if a spectral curvature model was fit. This value is necessary for users to understand and compare the Stokes I values to observations made at other frequencies.

1.0.12. *Polarized intensity and error*

The polarized intensity (in either intensity or flux density units) of the polarized component, at the polarization reference frequency. If polarization bias correction is used, this is the corrected polarized intensity. In sources where multiple polarized components have been identified, this is the polarized intensity of only the component with the corresponding RM. In the Faraday-complex case polarized intensity is not well defined, as there are multiple ways it could be defined depending on the method being used. In RM-synthesis methods this could be the amplitude of the peak of the FDF or some form of integrated measurement over the Faraday depth range of the component. In QU-fitting methods, this is most commonly either the polarized intensity after all depolarizing effects are removed (which is a direct fitting parameter in most models) or the polarized intensity extrapolated to zero wavelength (which is the same as the former in most models). In polarization angle linear fitting, this may be the average of polarized intensity across all the channels. Catalog authors should describe how they defined the polarized intensity when publishing their catalogs; catalog users should carefully check this definitions before using the values, especially if comparing values across multiple catalogs.

1.0.13. *Polarization bias correction method*

A string containing the method used to correct the polarized intensity for bias (e.g., Wardle & Kronberg 1974; Simmons & Stewart 1985). If no correction was applied, this should be 'None'; if it is not known if a correction was applied, this should be 'Unknown'. A list of currently used or suggested values appears in an accompanying document and will continue to be updated in the online documentation as new values come into use, in order to encourage authors to use the same values when using the same methods. Reporting this value in a catalog is important to allow users to assess the possible effects of polarization bias on the polarized intensity values in the catalog.

1.0.14. *Stokes extraction method*

A string describing the method used to extract the source spectra, for example whether they were intensities derived from peak-pixel values, aperture-integrated flux densities, or intensities or flux densities derived from point-source or Gaussian fitting. If the method is not known, the default value is 'Unknown'. A list of currently used or suggested values appears in an accompanying document and will continue to be updated in the online documentation as new values come into use, in order to encourage authors to use the same values when using the same methods. Including this column in a catalog can be helpful for users when trying to evaluate possible causes of differences between different measurements/catalogs or to interpret results from partially resolved sources.

1.0.15. *Integration aperture*

The linear size of the integration aperture over which the spectra have been integrated or averaged, in decimal degrees. If only peak/single pixel values are extracted from the images, this should be zero. If a circular or square aperture is used, the diameter or side length should be given. If a Gaussian-fit or similar process was used, then the FWHM of the fitted area would be appropriate. This information can be useful when trying to analyze a partially resolved source or to reproduce a catalog result.

1.0.16. *Fractional (linear) polarization and error*

The fractional polarization of the polarized component, and corresponding error. This parameter suffers from the same ambiguities and multiple possible definitions as polarized intensity; see the description of the polarized intensity column for a description of the complications. Values for this column should be fractional, and not percentages; nominally values should be less than 1, but this is not enforced because of the complications that can potentially be introduced by interferometer response and other effects. These columns can be very useful for studies of depolarization in sources (e.g., by comparing fractional polarization across different frequencies) as well as evaluating possible effects of polarization leakage.

1.0.17. *Electric vector polarization angle and error*

The electric vector position angle (EVPA, or simply polarization angle) and corresponding error, in degrees and at the polarization reference frequency, following the IAU standard convention (Contopoulos & Jappel 1974): increasing east from north, with zero degrees being towards the north celestial pole. We note that cosmology data sometimes use a different convention (de Serego Alighieri 2017), west-from-north, so care should be used in correcting the convention when using such data. Where Stokes Q and U follow the IAU convention, the EVPA is defined as $\frac{1}{2}\tan^{-1}\frac{U}{Q}$ Polarization angles should not be reported relative to the Galactic coordinate frame; the angles should also be explicitly the EVPA and not the plane-of-sky magnetic field. Polarization angle is only defined over a 180° span; we have chosen [0°,180°) rather than the sometimes-used (−90°,90°].

1.0.18. *Reference frequency for polarization*

The frequency, in Hz, at which the relevant polarization properties are applicable. This applies to the polarized intensity, EVPA, and the Stokes Q, U, and V values. For RM synthesis techniques, this frequency typically corresponds to the parameter λ_0^2 (Brentjens & de Bruyn 2005); for QU-fitting there is no equivalent value and authors can choose a suitable value (we suggest to make this equal to the Stokes I reference frequency) or leave the corresponding columns blank.

1.0.19. *De-rotated EVPA and error*

The EVPA with the effects of Faraday rotation removed (i.e. the polarization angle at the location of emission), and associated error, in degrees. This is sometimes called the 'zero wavelength' or 'intrinsic' polarization angle. This value can be useful in cases when users are interested in the magnetic field orientation (for synchrotron-emitting sources).

1.0.20. *Beam major axis, minor axis, and position angle*

The three parameters describing the shape of the synthesized beam at the reference frequency, as a Gaussian: the major axis, minor axis, and position angle, all in degrees. The major and minor axes are the FWHM of the Gaussian beam model, along the major and minor axes respectively, and the position angle is the angle of the major axis measured east from north similarly to the polarization angle, and is similarly defined in the range [0° 180°). This information can be very useful when comparing results with different studies that have differing resolution, if a source may be (partially) resolved in one or more of the observations.

1.0.21. *Reference frequency for beam shape*

The reference frequency for the beam shape parameters, in Hz, if the beam size follows a typical $1/\nu$ frequency dependence. If the individual channels have been convolved to a common size, this frequency should be set to zero to indicate that the beam has no frequency dependence. If the frequency dependence of the beam is not known, this should be left blank (NaN). This value can be useful as an indicator of whether frequency-dependent resolution could be producing any effects on the results, as well as allowing users to determine beam effects at different frequencies.

1.0.22. *Lowest and highest frequencies*

The center frequencies of the lowest and highest frequency channels used in determining the RM, in Hz. Channels that have been flagged out or otherwise not used should not be used to determine these values. These values can be useful for users who wish to assess the expected Faraday depth resolution and sensitivity to Faraday-complex signals that are broad in Faraday depth for a given observation.

1.0.23. *Typical channel width*

The bandwidth of the channels used to determine the RM, in Hz. If channels were averaged before being used to compute the RM, the width of the averaged channels should be used. If channels of different widths have been used together, this should be the most common channel bandwidth (the mode). Channel bandwidth is the key determining factor in bandwidth depolarization, so this value is helpful for assessing whether an observation may be biased against large RMs.

1.0.24. *Number of channels*

The number of frequency channels used to determine the RM. Channels that have been flagged out or otherwise were not used should not be included in this count. Since integers cannot use NaN values to represent missing data, any negative number can be used to represent missing values. This column can be useful in conjunction with the channel width to estimate total bandwidth, or to assess the ability of the data to constrain QU-fitting models with many parameters.

1.0.25. *Full width at half max of the rotation measure spread function*

The FWHM of the rotation measure spread function (RMSF) calculated during RM-synthesis, in rad m^{-2}. If RM-synthesis is not used, this column can be ignored or set to the theoretical RMSF FWHM (as defined in Brentjens & de Bruyn 2005) given the frequency coverage of the data. This column can provide a convenient metric for users to assess the Faraday depth resolution and sensitivity to Faraday complexity of an observation.

1.0.26. *Typical per-channel noise in Stokes Q, U*

An estimate for the noise in Stokes Q and U for a typical channel, in the same units as the Stokes parameters. The exact method of determining this is left to the catalog authors, but a mean or median of the channel noise values would be reasonable. This value can be useful as a metric of data quality, and for users planning deeper observations of interesting sources.

1.0.27. *Name of Telescope(s)*

A string containing the names or acronyms of all telescopes from which data were used, as a comma separated list. A list of currently used or suggested values appears in an accompanying document and will continue to be updated in the online documentation as new values come into use, in order to encourage authors to use the same values. The default value, if the data origin is not known, is ‘Unknown’.

1.0.28. *Integration time*

The integration time is the amount of time the telescope spent observing the source, in seconds, for a typical channel. If multiple observations at the same frequency were combined, then this should be the sum of the individual observation integration times, but if the observations were for different frequencies then the mean, median, or mode of the integration time for the individual channels should be used. This value can be useful for users concerned about possible time-averaging effects.

1.0.29. *Median epoch and interval of observation*

The median epoch is the midpoint of time between the first and last observations used to determine the RM. If a single observation was used, it should be the time at which the observation was half-complete. This time is stored as the modified Julian date (MJD, JD-2,400,000.5). The interval of observation is the span of time between the beginning of the first observation and the end of the final observation used to determine the RM, in days. If only a single observation was used, this is the difference between the start and end times of that observation. These columns allow RMs to be used in analysis of the evolution of RM over time.

1.0.30. *Instrumental leakage estimate*

An estimate of the degree of instrumental leakage present in Stokes Q and U, expressed as a fraction of Stokes I. If a leakage correction has been applied, this should be an estimate of the residual leakage after correction. This information can be useful to assess the significance of a detection (i.e., the risk of a reported RM being due to instrumental leakage rather than the astrophysical source), as well as the possible degree of systematic error introduced by leakage (which is distinct from the random error which is usually reported for quantities like Stokes Q and U or polarized intensity).

1.0.31. *Distance from beam center*

The angular distance of the source from the primary beam center in the observations, in degrees. If multiple observations or a phased array feed are used, this should be the distance from the nearest beam center. Since this quantity can be an indication of the possible severity of off-axis leakage effects, its purpose is to allow users to qualitatively estimate the relative effect of leakage on different sources (especially if the instrumental leakage estimate column is not supplied).

1.0.32. *Name of catalog*

A string containing a unique name for the catalog. The first preference for this is the paper in which the catalog was published, following the usual preference for the ADS bibcode or DOI. If the paper is not yet published, a short descriptive name can be used as a temporary substitute.

1.0.33. *Data references*

A string containing references to the sources of data used in determining the RM, following the preferences for references described above. If the paper reporting the catalog also reports on the observations used, then this should be the same paper. If data from multiple papers are used, a comma-separated list should be used. This column can be critically important for determining if RMs published in different catalogs are not independent (i.e., calculated from the same observations).

1.0.34. *Source ID in catalog*

A string containing the name of the source used in the catalog, if any. It is left completely to the authors' discretion whether this is a name unique to their catalog, a source name/id from another catalog (e.g., a 3C number), or something else.

1.0.35. *Source Classification*

A string containing the source classification, specifically what kind of physical object the source is. If multiple classifications have been given, each should be separated by a comma. A list of currently used or suggested values appears in an accompanying document and will continue to be updated in the online documentation as new values come into use, in order to encourage authors to use the same values.

1.0.36. *Notes*

A string containing any short notes the authors have made about individual sources.

Table 1. Column definitions for the RMTable2023 standard.

Column Name	Internal name	Data format	Unit	Limits	Default/Missing
		Position columns:			
Right Ascension [ICRS]	ra	double	deg	[0,360)	Essential
Declination [ICRS]	dec	double	deg	[-90,90]	Essential
Galactic Longitude	l	double	deg	[0,360)	Essential
Galactic Latitude	b	double	deg	[-90,90]	Essential
Position uncertainty	pos_err	float	deg	$[0,\infty)$	NaN
		RM columns:			
Rotation measure	rm	float	rad m^{-2}	$(-\infty,\infty)$	NaN
Error in RM	rm_err	float	rad m^{-2}	$[0,\infty)$	NaN
Width in Faraday depth	rm_width	float	rad m^{-2}	$[0,\infty)$	NaN
Error in width	rm_width_err	float	rad m^{-2}	$[0,\infty)$	NaN
Faraday complexity flag	complex_flag	string	–	‘Y’/’N’/‘U’	‘U’
Faraday complexity metric[a]	complex_test	string	–	–	‘’
RM determination method[a]	rm_method	string	–	–	‘Unknown’

Table 1 *continued on next page*

Table 1 *(continued)*

Column Name	Internal name	Data format	Unit	Limits	Default/Missing
Ionospheric correction method[a]	ionosphere	string	–	–	‘Unknown’
Number of RM components	Ncomp	integer	–	[1,∞)	1
		Polarization properties:			
Stokes I[b]	stokesI	float	Jy or Jy/beam	[0,∞)	NaN
Error in Stokes I[b]	stokesI_err	float	Jy or Jy/beam	[0,∞)	NaN
Stokes I spectral index	spectral_index	float	–	(-∞,∞)	NaN
Error in spectral index	spectral_index_err	float	–	[0,∞)	NaN
Reference frequency for Stokes I	reffreq_I	float	Hz	(0,∞)	NaN
Polarized intensity	polint	float	Jy or Jy/beam	[0,∞)	NaN
Error in Pol.Int.	polint_err	float	Jy or Jy/beam	[0,∞)	NaN
Polarization bias correction method[a]	pol_bias	string	–	–	‘Unknown’
Stokes extraction method[a]	flux_type	string	–	–	‘Unknown’
Integration aperture	aperture	float	deg	[0,∞)	NaN
Fractional (linear) polarization	fracpol	float	–	[0,∞)	NaN
Error in fractional polarization	fracpol_err	float	–	[0,∞)	NaN
Electric vector polarization angle	polangle	float	deg	[0,180)	NaN
Error in EVPA	polangle_err	float	deg	[0,∞)	NaN
Reference frequency for polarization	reffreq_pol	float	Hz	(0,∞)	NaN
Stokes Q[b]	stokesQ	float	Jy or Jy/beam	(-∞,∞)	NaN
Error in Stokes Q[b]	stokesQ_err	float	Jy or Jy/beam	[0,∞)	NaN
Stokes U[b]	stokesU	float	Jy or Jy/beam	(-∞,∞)	NaN
Error in Stokes U[b]	stokesU_err	float	Jy or Jy/beam	[0,∞)	NaN
De-rotated EVPA	derot_polangle	float	deg	[0,180)	NaN
Error in De-rotated EVPA	derot_polangle_err	float	deg	[0,∞)	NaN
Stokes V[b]	stokesV	float	Jy or Jy/beam	(-∞,∞)	NaN
Error in Stokes V[b]	stokesV_err	float	Jy or Jy/beam	[0,∞)	NaN

Table 1 *continued on next page*

Table 1 *(continued)*

Column Name	Internal name	Data format	Unit	Limits	Default/Missing
Observation properties:					
Beam major axis	beam_maj	float	deg	$[0,\infty)$	NaN
Beam minor axis	beam_min	float	deg	$[0,\infty)$	NaN
Beam position angle	beam_pa	float	deg	$[0,180)$	NaN
Reference frequency for beam	reffreq_beam	float	Hz	$[0,\infty)$	NaN
Lowest frequency	minfreq	float	Hz	$(0,\infty)$	NaN
Highest frequency	maxfreq	float	Hz	$(0,\infty)$	NaN
Typical channel width	channelwidth	float	Hz	$(0,\infty)$	NaN
Number of channels	Nchan	integer	–	$(0,\infty)$	Any negative integer[c]
Full-width at half maximum of the RMSF	rmsf_fwhm	float	rad m^{-2}	$[0,\infty)$	NaN
Typical per-channel noise in Q, U^b	noise_chan	float	Jy or Jy/beam	$[0,\infty)$	NaN
Name of Telescope(s)[a]	telescope	string	–	–	'Unknown'
Integration time	int_time	float	s	$[0,\infty)$	NaN
Median epoch of observation	epoch	float	days	$(-\infty,\infty)$	NaN
Interval of observation	interval	float	days	$[0,\infty)$	NaN
Instrumental leakage estimate	leakage	float	–	$[0,\infty)$	NaN
Distance from beam center	beamdist	float	deg	$[0,\infty)$	NaN
Miscellaneous:					
Name of catalog	catalog	string	–	–	Essential
Data references	dataref	string	–	–	''
Source ID in catalog	cat_id	string	–	–	''
Source classification[a]	type	string	–	–	''
Notes	notes	string	–	–	''

Table 1 *continued on next page*

Table 1 *(continued)*

Column Name	Internal name	Data format	Unit	Limits	Default/Missing

NOTE—Columns marked as essential are required to have a value and cannot be blank. See text for additional notes on some columns.

[a]: These columns have a list of currently used or suggested values to encourage standardization.

[b]: All Stokes parameters can be either flux densities or intensities.

[c]: Since NaN is not generally defined for integers, a negative integer should be used to represent missing data. The default behaviour in Python is to replace NaNs with -2147483648, so this value is generally used in RMTables generated by Python.

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