

# The Probe of Inflation and Cosmic Origins



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## Report from a Probe-Scale Mission Study

January, 2019

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The research was funded by a NASA grant NNX17AK52G to the University of Minnesota / Twin Cities, and by the Jet Propulsion Laboratory, California Institute of Technology, under a contract with the National Aeronautics and Space Administration, and by Lockheed Martin Corporation. Substantial contributions to this project were volunteered by scientists at many institutions world-wide. They are gratefully acknowledged.

The information presented about the PICO mission concept is pre-decisional and is provided for planning and discussion purposes only.

The cost information contained in this document is of a budgetary and planning nature and is intended for informational purposes only. It does not constitute a commitment on the part of JPL and/or Caltech.

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# 1 Executive Summary

The cosmic microwave background (CMB) comes to us from the furthest reaches of the observable Universe, and its photons experience all of cosmic history. Created when the Universe was a hotter, simpler place, CMB photons probe fundamental physics, provide exquisite measurements of the constituents of the cosmos, and tests of relativity. On their journey they feel the impact of the gravitational potentials formed from the assembling cosmic web of superclusters, clusters and galaxies. They interact with the ionized gas in the inter- and circum-galactic medium, gas that eventually fuels star and galaxy formation. Superposed upon the CMB is the emission from multiple extragalactic sources and from our Galaxy. All of this leaves an imprint which sensitive measurements can disentangle so that CMB studies impact essentially every aspect of cosmology and astrophysics.

Building upon a long legacy of successful measurements, the next decade holds tremendous potential for new, exciting CMB discoveries. Such discoveries, delivered by the Probe of Inflation and Cosmic Origins (PICO, Fig. 1.1), promise to be revolutionary, affecting physics, astrophysics, and cosmology. PICO is an imaging polarimeter that will scan the sky for 5 years with 21 frequency bands spread between 21 and 800 GHz; see Table 2.1. It will produce 10 independent full sky surveys of intensity and polarization with a final combined-map noise level equivalent to 3250 *Planck* missions for the baseline required specifications, though in our current best-estimate (CBE) it would perform as 6400 *Planck* missions.

With these capabilities, unmatched by any other existing or proposed platform, PICO will have compelling and broad science deliverables. The mission will respond to seven science objectives (SOs), which are listed in Table 2.2. Delivering this set was the basis for selecting PICO’s design and for setting instrument requirements. But PICO’s science reach is broader than the baseline set.

PICO could determine the energy scale of inflation and give a first, direct probe of quantum gravity; this is SO1 (§ 2.2). The mission will attempt to detect the signal, which arises from gravitational waves sourced by inflation and parametrized by the tensor to scalar ratio  $r$ , at a level of  $r = 5 \times 10^{-4}$  ( $5\sigma$ ). This level is 100 times lower than current upper limits, and 10 tens lower than limits forecast by funded foreseeable experiments. If the signal is not detected, PICO will constrain broad classes of inflationary models and exclude at  $10\sigma$  models for which the characteristic scale in the potential is the Planck scale (SO1 and SO2). The combination of data from PICO and LSST can constrain features in the inflationary potential, the field content during inflation and could rule out all models of slow-roll single-field inflation, marking a watershed in studies of inflation.

The mission will have a deep impact on particle physics by measuring the expected sum of the neutrino masses with  $4\sigma$  confidence, rising to  $7\sigma$  if the sum is near 0.1 eV (SO3). Reaching the  $4\sigma$  level can only be done with an instrument that can measure the polarization of the CMB on the largest angular scales, a measurement optimally done from space, which gives access to the full sky and with a broad band of frequencies to remove foreground contaminants. Cluster counts provided by PICO in combination with followup red-shift measurements, and PICO’s map of the projected gravitational potentials along the line of site in combination with LSST gold sample of galaxies, will give two additional independent and equally competitive constraints on the sum of neutrino masses.

The measurements will either detect or strongly constrain deviations from the standard model of particle physics by counting the number of light particle species  $N_{eff}$  in the early universe. The constraint of  $\Delta N_{eff} < 0.06$  ( $2\sigma$ ) will move the allowed decoupling temperature of a hypothetical new vector particle to temperatures that are 400 times higher than currently determined by *Planck* (SO4). The data will enable a



Figure 1.1: The PICO spacecraft

search for the existence of primordial magnetic fields with sufficient sensitivity to rule them out as the sole source for the largest observed galactic magnetic fields; they will improve by  $\sim 300$  constraints on polarization rotation arising from early universe fields that lead to cosmic birefringence, and will thus constrain string theory-motivated axions; and they will constrain generic models of dark matter candidates.

PICO will elucidate the processes affecting the evolution of cosmic structures. It will measure the optical depth to reionization  $\tau$  with an error  $\sigma(\tau) = 0.002$  limited only by the statistics of the small number of modes available to observe the largest angular scale CMB polarization (SO5). The measurement will be used to constrain models of the formation of the first luminous sources, and is a key input to all astrophysical attempts to improve the determination of the sum of neutrino masses. The data will give a map of the projected gravitational potential due to all structures with a signal-to-noise ratio (SNR) 14 times higher than *Planck*, and a catalog of 150,000 clusters extending to their earliest formation redshift. Each of these datasets will be used in combination with other data – from LSST and from future optical and infrared surveys – to independently constrain the evolution of the amplitude of linear fluctuations  $\sigma_8(z)$ , with sub-percent accuracy.

Cross-correlating PICO’s map of the thermal Sunyaev–Zel’dovich effect with LSST’s gold sample of galaxies, a correlation that is forecast to have an SNR exceeding 1000, will give precise tracing of the evolution of thermal pressure with  $z$ . These will be used to place constraints on models of energetic feedback, which are the most uncertain ingredient in models of galaxy formation.

PICO’s maps of the Milky Way will be used to resolve long-standing questions about our own Galaxy. Galactic interstellar dust grains are a link between atoms and molecules and planetary objects, yet their composition and their role in Galactic chemistry is still under debate. Galactic magnetic fields are known to play a key role in the dynamics of gas in the Galaxy, and in determining the efficiency of star formation, but their quantitative contribution relative to turbulence is yet to be determined. With the mission’s Galactic dust polarization maps we will constrain dust properties, including composition, temperature, and emissivities (SO6). The derived, detailed 1' resolution maps of magnetic field will be used to quantify the relative roles of turbulence and magnetic fields in the dynamics of the Galaxy and in the observed low star formation efficiency (SO7).

PICO will give 21 full-sky maps of intensity and polarization, each much deeper compared to *Planck*’s seven; five will have polarization information in frequencies between 385 and 800 GHz that *Planck* did not have. At 385 GHz PICO’s noise is 100 times lower than *Planck*’s at 353 GHz. And PICO’s highest resolution is five times finer than *Planck*’s. Only PICO will provide such full sky legacy maps. Six are at frequencies not accessible to ground-based experiments, with which we will: constrain the early phases of galaxy evolution by discovering 4500 strongly lensed dusty galaxies with  $z$  up to 5; investigate the early phases of cluster evolution by discovering 50,000 proto-clusters out to  $z \sim 4.5$ ; give census of cold dust in 30,000 low  $z$  galaxies; give cosmic infrared background maps of the anisotropies due to dusty star-forming galaxies; map magnetic fields in 70 nearby galaxies; and, with 3,000-fold increase relative to *Planck* in the number of independent measurements of magnetic field in our own Galaxy, study how magnetic fields are generated through a combination of turbulence and large scale gas motion.

now here; not complete

With its broad frequency coverage PICO is better equipped than any other current or foreseeable instrument to separate the detected signals to their original sources of emission. This capability is most important for the faintest of signals, the telltale of inflation, which is already known to be dominated by Galactic foregrounds. Our simulations indicate that PICO’s combination of low noise and multitude of bands is most likely necessary and sufficient to separate the inflationary signal from the foregrounds at the required level. But current uncertainties on the parameters characterizing Galactic foregrounds are large and we recommend support for (1) modeling, simulation, and algorithm development for effective foreground separation, and (2) improved Galactic emission measurements with sub-orbital experiments.

one that may to with 21 frequency bands

Table 2.1: Mission Specifications

Bands	GHz	21	25	30	36	43	52	62	75	90	108	129	155	186	223	268	321	385	462	555	666	799	Cost: \$ 958M	Launch mass: 2147 kg
FWHM	arcmin	38.4	32.0	28.3	23.6	22.2	18.4	12.8	10.7	9.5	7.9	7.4	6.2	4.3	3.6	3.2	2.6	2.5	2.1	1.5	1.3	1.1	Mission length: 5 years	Total power: 1526 W
Polarization map depth																					Combined polarization map depth			
Baseline	$\mu\text{K}_{\text{CMB}}$ arcmin	23.9	18.4	12.4	7.9	7.9	5.7	5.4	4.2	2.8	2.3	2.1	1.8	4.0	4.5	3.1	4.2	4.5	9.1	45.8	177.2	1047	0.87 $\mu\text{K}_{\text{CMB}}$ equivalent to 3250 <i>Planck</i> Missions	
Baseline	Jy/sr	8.3	10.9	11.8	12.9	19.5	23.8	45.4	58.3	59.3	77.3	96.0	119.1	433.1	604.2	433.4	577.8	429.1	551.1	1580	2075	2884	0.61 $\mu\text{K}_{\text{CMB}}$ equivalent to 6400 <i>Planck</i> Missions	
CBE <sup>a</sup>	$\mu\text{K}_{\text{CMB}}$ arcmin	16.9	13.0	8.7	5.6	5.6	4.0	3.8	3.0	2.0	1.6	1.5	1.3	2.8	3.2	2.2	3.0	3.2	6.4	32.4	125.3	740.3	0.61 $\mu\text{K}_{\text{CMB}}$ equivalent to 6400 <i>Planck</i> Missions	
CBE <sup>a</sup>	Jy/s	5.9	7.7	8.3	9.2	13.8	16.8	32.1	41.3	41.8	53.5	69.3	83.7	301.5	436.3	303.5	411.1	303.1	387.3	1117	1467	2040	0.61 $\mu\text{K}_{\text{CMB}}$ equivalent to 6400 <i>Planck</i> Missions	

<sup>a</sup> Current best estimate

Table 2.2: Science Traceability Matrix (STM)

Science Goals from NASA Science Plan	Science Objectives	Scientific Measurement Requirements						Instrument (single instrument, single mode)				Mission Functional Requirements	
		Model Parameters		Physical Parameters		Observables		Functional Requirements	Projected Performance				
<i>Explore how the Universe began (Inflation)</i>	SO1. Probe the physics of the big bang by detecting the energy scale at which inflation occurred if it is above $4 \times 10^{15}$ GeV, or place an upper limit if it is below (§ 2.2)	Tensor-to-scalar ratio $r$ : $\sigma(r) = 1 \times 10^{-4}$ at $r = 0$ ; $r < 5 \times 10^{-4}$ at 5 $\sigma$ confidence level <sup>a</sup>		CMB polarization <i>BB</i> power spectrum for modes $2 < \ell < 300$ to cosmic-variance limit, and CMB lensing power spectrum for modes $2 < \ell < 1000$ to cosmic-variance limit		Linear polarization across $60 < v < 300$ GHz over entire sky; foreground separation requires $21 < v < 799$ GHz		Frequency coverage: central frequencies $v_c$ from 21 to 799 GHz.		Sun-Earth L2 orbit with Sun-Probe-Earth $< 15^\circ$ . 5 yr survey with $\geq 95\%$ survey efficiency.			
	SO2. Probe the physics of the big bang by excluding classes of potentials as the driving force of inflation (§ 2.2, Fig. 2.2)	Spectral index ( $n_s$ ) and its derivative ( $n_{\text{run}}$ ): $\sigma(n_s) < 0.0015$ ; $\sigma(n_{\text{run}}) < 0.002$		CMB polarization <i>BB</i> power spectrum for modes $2 < \ell < 1000$ to cosmic-variance limit		Intensity and linear polarization across $60 < v < 220$ GHz over the entire sky		Frequency resolution: $\Delta v/v_c = 25\%$ . Sensitivity: See Table 2.1. Combined instrument noise: $< 0.61 \mu\text{K}_{\text{CMB}} \sqrt{s}$ .		Frequency coverage: See Table 2.1. 21 bands with $v_c$ from 21 to 799 GHz.		Full sky survey: Spin instrument at 1 rpm; boresight $69^\circ$ off spin axis; spin axis $26^\circ$ off anti-Sun line, precessing $360^\circ / 10$ hr.	
<i>Discover how the Universe works (neutrino mass and <math>N_{\text{eff}}</math>)</i>	SO3. Determine the sum of neutrino masses. (§ 2.2, Fig. 2.5)	Sum of neutrino masses ( $\Sigma m_\nu$ ): $\Sigma m_\nu < 15$ meV with DESI or Euclid <sup>b</sup>		CMB polarization <i>BB</i> power spectrum for modes $2 < \ell < 4000$ to cosmic-variance limit; CMB intensity maps (to give Compton <i>Y</i> map from which we extract clusters)		60 < $v < 220$ GHz over the entire sky		Angular resolution [for delensing and foreground separation]: $\text{FWHM} = 6.2' \times (155 \text{ GHz}/v_c)$ . Effective aperture: 1.4 m. Sampling rate: $(3/\text{BeamFWHM}) \times (336'/\text{s})$ .		Angular resolution: $\Delta v/v_c = 25\%$ . Sensitivity: See Table 2.1. Combined instrument noise: $0.43 \mu\text{K}_{\text{CMB}} \sqrt{s}$ . Angular resolution: See Table 2.1.		Pointing control: Spin axis $60'$ ( $3\sigma$ , radial). Spin $1 \pm 0.1$ rpm ( $3\sigma$ )	
	SO4. Tightly constrain the thermalized fundamental particle content of the early Universe (§ 2.2, Fig. 2.4)	Number of neutrino effective relativistic degrees of freedom ( $N_{\text{eff}}$ ): $\sigma(N_{\text{eff}}) < 0.03$		CMB temperature and <i>EE</i> polarization power spectra $2 < \ell < 4000$ to cosmic-variance limit		Linear polarization across $60 < v < 300$ GHz over entire sky; foreground separation enveloped by SO1		Sampling rate: $(3/\text{BeamFWHM}) \times (336'/\text{s})$ . FWHM = $6.2' \times (155 \text{ GHz}/v_c)$ ; $1.1'$ for $v_c = 799$ GHz.		Sampling rate: See Table 3.1. $(3/\text{BeamFWHM}) \times (336'/\text{s})$ . Pointing knowledge (telescope boresight): $10''$ ( $3\sigma$ , each axis) from spacecraft attitude; $1''$ ( $1\sigma$ , total) final reconstructed;		Pointing stability: Drift of spin axis $< 1'/1\text{min}$ ( $3\sigma$ , radial); jitter $< 20''/20$ ms ( $3\sigma$ , radial).	
<i>Explore how the Universe evolved (reionization)</i>	SO5. Distinguish between models that describe the formation of the earliest stars in the Universe (§ 2.3, Fig. 2.6)	Optical depth to reionization ( $\tau$ ): $\sigma(\tau) < 0.002$		CMB polarization <i>EE</i> power spectrum for modes $2 < \ell < 20$ to cosmic-variance limit		Linear polarization across $60 < v < 300$ GHz over entire sky; foreground separation enveloped by SO1		Enveloped by SO1–4, except: Angular resolution: $\leq 1.1'$ (at highest frequency) Sensitivity at 799 GHz: 27.4 kJy/sr		Return and process instrument data: 1.5 Tbits/day (after 4 $\times$ compression)		Thermally isolate instrument from solar radiation and from spacecraft bus	
	SO6. Constrain the temperatures and emissivities characterizing the Milky Way's interstellar diffuse dust (§ 2.4)	Intrinsic polarization fractions of the components of the diffuse interstellar medium to accuracy better than 3% when averaged over $10'$ pixels		Fractional polarization and intensity as a function of frequency		Intensity and linear polarization maps in 12 frequency bands between 108 and 799 GHz.							
<i>Explore how the Universe evolved (Galactic structure and dynamics)</i>	SO7. Determine if magnetic fields are the dominant cause of low Galactic star-formation efficiency (§ 2.4)	Ratio of cloud mass to maximum mass that can be supported by magnetic field ("Mass to flux ratio" $\mu$ ); ratio of turbulent energy to magnetic energy (Alfvén Mach number $\mathcal{M}_A$ ) on scales 0.05–100 pc		The turbulence power spectrum on scales 0.05–100 pc; magnetic field strength ( $B$ ) as a function of spatial scale and density; hydrogen column density; gas velocity dispersion		Intensity and linear polarization with $< 1$ pc resolution for thousands of molecular clouds and with $< 0.05$ pc for the 10 nearest molecular clouds; maps of polarization with $1'$ resolution over the entire sky							

<sup>a</sup> The values predicted include delensing and foreground subtraction; see Section 2.2.<sup>b</sup> Using the PICO  $\tau$  and *BB* lensing power spectrum, and BAO from DESI or Euclid. The same strength constraint can be derived independently using clusters detected by PICO, if their redshift becomes available, and using  $\sigma_8$  from LSST.

include: single instrument, simple scan; need funding for foregrounds and systematics

## 2 Science

### 2.1 Introduction

The Probe of Inflation and Cosmic Origins (PICO) has seven science objectives (SOs) that derive from NASA’s three science strategic goals: to explore how the Universe began; to discover how the Universe works; and to explore how the Universe evolved. The SOs, which include probing the inflationary epoch near the big bang, constraining the properties of fundamental particles, probing the structure and evolution of the Universe, and understanding the structure of our own Galaxy, require measurements in and around frequency bands in which the cosmic microwave background is most intense. The SOs and the measurement requirements derived from them are given in Table 2.2 and define the PICO ‘baseline design’. This design gives rise to a mission that delivers a much broader set of science deliverables. This report describes the broad array of science deliverables, but focuses primarily on the primary SOs listed in the Table.

The PICO mission consists of a single instrument: an imaging polarimeter that surveys the entire sky at 21 frequency bands spread between 21 and 799 GHz. The telescope has an aperture of 1.4 meter giving diffraction limited resolution between  $38'$  and  $1'$ . The instrument incorporates a 0.1 K-cooled focal plane that hosts 12,996 transition edge sensor (TES) bolometric detectors. The baseline design contains a margin of 40% in detector noise. Experience with past space missions, most recently with *Planck*, shows that pre-mission calculated detector performance is in fact achieved in space [1, 2]. We therefore include throughout this report performance estimates that are based also on our current estimate for the actual performance. Those are labeled current best estimate (CBE). Table 2.1 gives the frequency bands, resolution, both baseline and CBE noise levels, as well as other key mission parameters.

Mission operations throughout the 5-year duration of the mission are simple. PICO executes one specific sky scan pattern that repeats itself for the entire duration of the mission. With this scan pattern the entire sky is scanned every 6 months giving 10 independent full-sky maps of intensity and polarization, giving the three Stokes parameters  $T$ ,  $Q$  and  $U$  of the detected electromagnetic radiation. The scan pattern is optimized for polarization measurements as each sky pixel is scanned along multiple orientations. Therefore full-sky maps of the Stokes parameters can be reconstructed from the data of each of the 12,996 polarization sensitive detectors.

Some of the PICO polarization science goals are more appropriately described in terms of  $E$  and  $B$  polarization maps rather than  $Q$  and  $U$  [3, 4]. This is because sources of polarization signatures that are scalar in nature, such as primordial density perturbations, can only produce  $E$ -mode polarization. Sources that are tensor in nature, such as gravitational waves, can produce both  $E$ - and  $B$ -mode polarization. The angular power spectra of  $E$  and  $B$  maps will be denoted as  $EE$  and  $BB$ .

This report assumes that PICO’s phase A will start in 2023. The science outcomes are expected to break new ground, and to be complementary to data sets available at the end of 2020s and the beginning of the following decade. Therefore we are including performance comparisons to funded projects that are in implementation and for which final design specifications and projections exist in the literature. Such next generation US-based CMB experiments are collectively denoted as Stage-III (S3) [5–11].

This Section describes PICO’s SOs, places them in context of current knowledge, and provides performance forecasts. It also includes a discussion of PICO’s complementarity with ground-based measurements; of our estimates of the efficacy of separating the detected radiation to the several astrophysical sources of emission; of anticipated systematic uncertainties; and of the measurement requirements that derive from the combination of these topics. Section 3 describe the instrument, which consists of the telescope, the focal plane, the detector readout, and the thermal design. The section also describes plans for integration and testing. Section 4 describes the operations, including the instrument’s survey of the sky, and the spacecraft. In Section 5 we discuss the path to maturing any of small number of technologies that are not yet at TRL-6, and potential descopes. Project management, assessment of risk, and cost are presented in Section 6.

## 2.2 Fundamental Physics

### 2.2.1 Gravitational waves and Inflation

• **Targets** Measurements of the CMB together with Einstein’s theory of general relativity imply that the observed density perturbations must have been created long before the CMB was released, and rather remarkably even before the Universe became filled with a hot and dense plasma of fundamental particles. Understanding the mechanism generating these perturbations, which evolved to fill the Universe with structures, is one of the most important open questions in cosmology. In addition to density perturbations, this mechanism may have also produced gravitational waves that would have left a *B*-mode polarization signature in the CMB [12, 13]. Any detection of primordial *B*-mode polarization by PICO will constitute evidence for gravitational waves from the same primordial period that created the density perturbations and will open a new window onto this early epoch. Because the dynamics of gravitational waves is essentially unaffected by the plasma, they would be a pristine relic from the earliest moments of our Universe, and their properties would shed light on the mechanism that created the primordial perturbations.

Inflation, a period of nearly exponential expansion of the early Universe [14–17], is the leading paradigm explaining the origin of the primordial density perturbations [18–22]. It predicts a nearly scale invariant spectrum of primordial gravitational waves originating from quantum fluctuations [23]. Measurements of the CMB are the only foreseeable way to detect these gravitational waves.

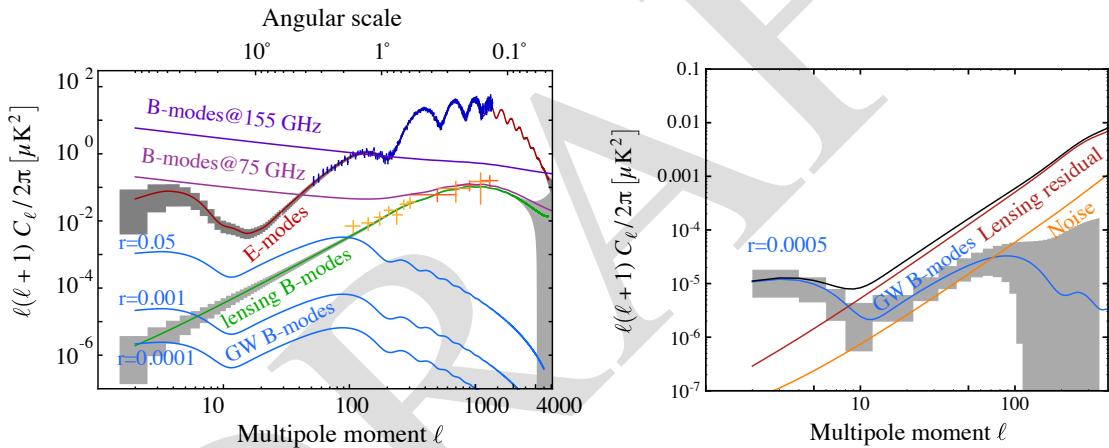


Figure 2.1: With PICO’s **baseline** configuration we will measure the *EE* (left, red) and lensing *BB* (green) angular power spectra with high precision (grey). PICO’s goal is to detect  $r = 5 \times 10^{-4}$  ( $5\sigma$ ) (right, grey). This forecast includes PICO’s 80% delensing (red) and foreground separation. The **baseline** noise level (right, orange) allows detection of even lower levels; we expect foreground separation to limit performance. As an example we show the *BB* spectra of Galactic emission on the cleanest 60% of the sky at 75 and 155 GHz (left, purple). They largely dominate the cosmological signals. Also shown are measurements of lensing from current experiments (left, orange) [24–27], *Planck*’s *EE* measurements (left, dark blue) [28], and the *BB* spectrum produced by IGW with different values of  $r$  (cyan).

The strength of the signal, quantified by the tensor-to-scalar ratio  $r$ , is a direct measure of the expansion rate of the Universe during inflation. Together with the Friedmann equation, it reveals one of the most important characteristics of inflation: its energy scale.<sup>1</sup> A detection of  $r$  “would be a watershed discovery”, a quote from the 2010 decadal panel report [30]. The combination of data from *Planck* and the BICEP/Keck Array give the strongest constraint to date,  $r < 0.06$  (95%) [31]. Next decade S3 efforts strive to reach  $\sigma(r) = 2 \times 10^{-3}$  [9, 32].

PICO’s goal is to detect primordial gravitational waves if inflation occurred at an energy scale of at

<sup>1</sup>In some models of inflation the one-to-one correspondence between  $r$  and the energy scale of inflation does not hold because there are additional sources of gravitational waves [29]. However, in these models the signal is highly non-Gaussian and could be distinguished from quantum fluctuations.

least  $5 \times 10^{15}$  GeV, or equivalently  $r = 5 \times 10^{-4}$  ( $5\sigma$ ) (SO1 in Table 2.2 and Figure 2.1). A detection will have profound implications for fundamental physics because it will provide evidence for a new energy scale tantalizingly close to the energy scale associated with grand unified theories, probe physics at energies far beyond the reach of terrestrial colliders, and be the first observation of a phenomenon associated with quantum gravity [33].

There are only two classes of slow-roll inflation in agreement with current data that naturally explain the observed value of the spectral index of primordial fluctuations  $n_s$  [34]. The first class is characterized by potentials of the form  $V(\phi) \propto \phi^p$ . This class includes many of the simplest models of inflation, some of which have already been strongly disfavored by existing observations. Select models in this class are shown as blue lines in Fig. 2.2. If the constraints on  $n_s$  tighten by about a factor of two with the central value unchanged, and the upper limit on  $r$  improves by an order of magnitude, this class would be ruled out.

The second class is characterized by potentials that approach a constant as a function of field value, either like a power law or exponentially. Two representative examples in this class are shown as the green and gray bands in Fig. 2.2. This class also includes  $R^2$  inflation, which predicts a tensor-to-scalar ratio of  $r \sim 0.004$ . All models in this class with a characteristic scale in the potential that is larger than the Planck scale predict a tensor-to-scalar ratio of  $r \gtrsim 0.001$ .

Many microphysical models in this class possess a characteristic scale that is super-Planckian. PICO will either detect gravitational waves with high confidence or will exclude all models with a Planckian characteristic scale with high significance. But not all models have a super-Planckian characteristic scale. The Goncharov-Linde model is an example with a somewhat smaller characteristic scale. It predicts a tensor-to-scalar ratio of  $r \sim 4 \times 10^{-4}$  [35], still within reach of PICO. In addition, there are models with significantly smaller values that are out of reach, but distinguishing between models with sub- and super-Planckian characteristic scales would provide much needed guidance to discriminate between classes of ideas for the earliest moments of our universe.

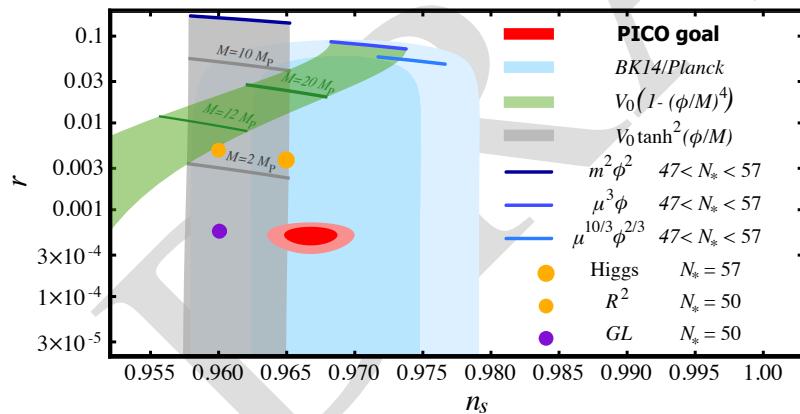


Figure 2.2: Current  $1\sigma$  and  $2\sigma$  limits on  $r$  and  $n_s$  (cyan) and forecasted constraints for a fiducial model with  $r = 0.0005$  for PICO together with predictions for selected models of inflation. Characteristic super-Planckian scales in the potentials are marked with darker lines.

- **Observational Considerations** The  $BB$  angular power spectrum measured by PICO will have contributions by Galactic sources of emission and ‘lensing’  $B$ -modes, created by gravitational lensing of  $E$ -modes as the CMB photons traverse the gravitational potentials throughout the Universe (Fig. 2.1 and § 2.3.2). In case of an  $r$  detection, there will be two additional features due to the inflationary signal. One is the ‘recombination peak’ at  $\ell \sim 80$  and the other is the ‘reionization peak’ at multipoles of  $\ell \lesssim 10$ . PICO’s strong constraints on  $r$  derive from using all available  $\ell$  modes.

The Galactic signals act as foregrounds, and uncertainty in the characterization of these foregrounds already now limits our ability to constrain  $r$ . PICO’s goal for reaching  $\sigma(r) = 1 \times 10^{-4}$  is driven by estimates of the efficacy of foreground separation, not noise. An analytic performance forecast accounting for PICO’s statistical noise level and a foreground model that has polarized emission from two components of dust,

synchrotron radiation, and correlations between synchrotron and dust emission, gives  $\sigma(r) = 2 \times 10^{-5}$ , five times lower than our baseline goal. Our goal is set based on end-to-end, map-based simulations, which include physical effects known to exist but not captured in the analytic forecasts; see Section 2.6.

When the tensor-to-scalar ratio  $r \simeq 0.01$ , the  $BB$  lensing and inflation spectra are comparable in magnitude at the recombination peak ( $\ell \sim 80$ ). For lower levels of  $r$ , the lensing  $B$ -mode dominates, but the  $B$ -mode maps can be ‘delensed’ if the polarization maps are measured with few-arcmin resolution and sufficient depth [36, 37]. Forecasts for PICO show that at least 73% of the lensing  $B$ -mode power can be removed for the baseline configuration, after accounting for Galactic foreground separation. As much as 84% will be removed for the CBE and for milder foreground contamination. For measuring the recombination peak, delensing is essential to reach PICO’s limits on  $r$ , and this was a driver in choosing the resolution of the instrument.

For the levels of  $r$  targeted by PICO, the  $BB$  reionization signal ( $\ell \sim 5$ ) has somewhat higher level than the lensing spectrum, but the map-level foregrounds at this angular scale are at least two orders of magnitude brighter. PICO’s instrument temporal stability, absence of atmospheric noise, full-sky coverage, and unmatched capability to characterize and separate foregrounds make it the most suitable instrument to measure these lowest multipoles. No S3 experiments have measured or plan to measure  $B$ -modes at  $\ell < 40$  that reach to  $\sigma(r) < 0.006$  at the lowest multipoles [11, 38].

If an inflationary  $B$ -mode signal is detected, it is important to characterize its entire  $\ell$  dependence, in both the predicted reionization and recombination peaks. Furthermore, the PICO full sky coverage will enable detection of the recombination peak in several independent patches of the sky, giving an important systematic cross-check.

- **Scalar Spectral Index and Non-Gaussianity** Models of the early Universe differ in their predictions for the scalar spectral index  $n_s$  and its scale dependence  $n_{\text{run}}$ . PICO will improve  $n_s$  and  $n_{\text{run}}$  constraints by a factor of three relative to *Planck* and achieve  $\sigma(n_s) = 0.0015$  and  $\sigma(n_{\text{run}}) = 0.002$ . For many models of inflation how reheating occurred is unknown. This translates into an uncertainty for the theoretical predictions for these models of inflation. PICO’s precision is sufficient to distinguish between different possible reheating scenarios at  $> 3\sigma$ .

The simplest models of inflation, in which there is a single inflaton field, predict primordial fluctuations that are very nearly Gaussian with  $|f_{NL}^{\text{local}}| < 1$ , where  $f_{NL}^{\text{local}}$  is a parameter quantifying the level of local non-Gaussianity. A detection of  $|f_{NL}^{\text{local}}| > 1$  points exclusively to models of inflation with multiple fields (Fig. 2.3), making  $\sigma(f_{NL}^{\text{local}}) \simeq 1$  a compelling target. *Planck* gives a constraint of  $f_{NL}^{\text{local}} = 0.8 \pm 5(1\sigma)$  [39], and further measurements of the CMB alone can not improve on this constraint by more than a factor of 2-3. However, it has been shown that correlating large-scale structure tracers that have different clustering bias factors can enhance the signature of non-Gaussianity [40–42]. Fig. 2.3 shows expected constraints from correlations between the PICO lensing potential maps (§ 2.3.2) and LSST galaxies. For  $f_{NL}^{\text{local}} = 2$ ,  $3\sigma$  evidence will be reached if large angular scale ( $L \geq 8$ ) auto- and cross-correlation spectra can be used. If LSST’s auto-correlation can only be used on smaller angular scales  $L \geq 20$ , the  $3\sigma$  evidence weakens to  $2\sigma$ .

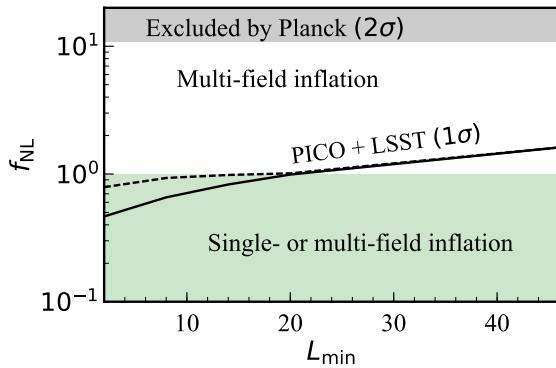


Figure 2.3: Cross-correlating PICO’s lensing potential map with LSST galaxies will allow detecting or excluding  $f_{NL}^{\text{local}} = 2$  with  $3\sigma$  evidence if the data can be used at angular scales  $L \geq 8$  (solid black). A detection above  $|f_{NL}^{\text{local}}| = 1$  indicates that inflation is driven by multiple fields; single-field inflation has  $|f_{NL}^{\text{local}}| < 1$  (green region). The *Planck* constraint is  $f_{NL}^{\text{local}} < 10.8(2\sigma)$ . The cross-correlations will allow excluding or detecting  $f_{NL}^{\text{local}} = 2(2\sigma)$  if LSST data is used only for  $L \geq 20$  (dash). Figure 2.8 gives the assumptions for the LSST data.

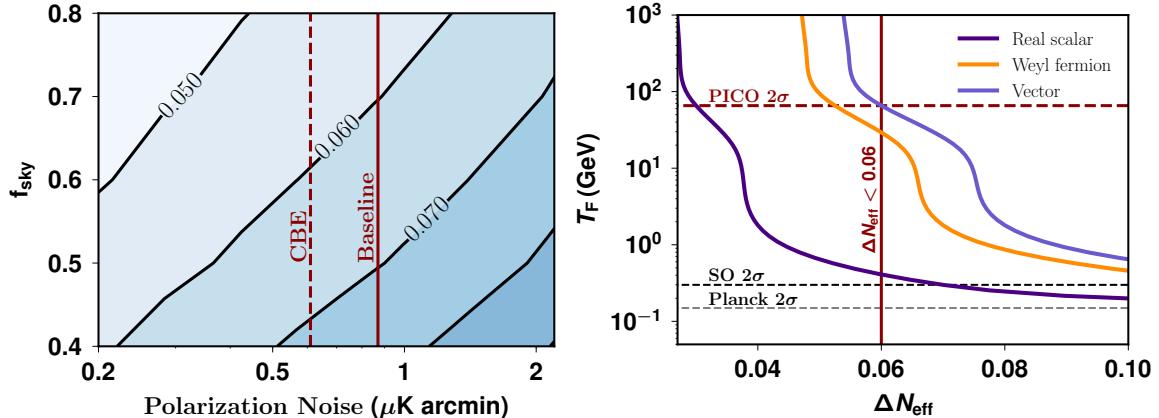


Figure 2.4: PICO will achieve a constraint  $\Delta N_{\text{eff}} < 0.06$  (95%) (left, 2 $\sigma$  contours shown) in the baseline configuration (vertical solid) using its cosmic-variance-limited measurement of  $EE$  for  $\ell \leq 2300$ , and 21 frequency bands to utilize data over 70% of the sky (5' resolution assumed). This constraint translates to moving up the lowest decoupling temperature  $T_F$  for vector, Weyl Fermion, and scalar particles by a factor of 400, 200, and 6, respectively, relative to *Planck* (right, dash black, only  $T_F$  for vector particle is shown). We also show the projected vector particle limit for the Simons Observatory [32].

### 2.2.2 Fundamental Particles: Light relics, Dark Matter, and Neutrinos

- **Light Relics** In the inflationary paradigm, the Universe was reheated to temperatures of at least 10 MeV and perhaps as high as  $10^{12}$  GeV. At these high temperatures, even very weakly interacting or very massive particles, such as those arising in extensions of the Standard Model of particle physics, can be produced in large abundances [43, 44]. As the Universe expands and cools the particles fall out of equilibrium, an event referred to as ‘decoupling’, and characterized by a decoupling temperature  $T_F$ . The decoupling leaves observable signatures in the CMB power spectra. Through these effects the CMB is a sensitive probe of neutrino and of other particles’ properties.

One particularly compelling target is the effective number of light relic particle species  $N_{\text{eff}}$ . The canonical value with three neutrino families is  $N_{\text{eff}} = 3.046$ . Additional light particles contribute a change  $\Delta N_{\text{eff}}$  that is a function only of the decoupling temperature and the effective degrees of freedom of the particle,  $g$ . The magnitude of  $\Delta N_{\text{eff}}$  is quite restricted, even for widely varying decoupling temperatures  $T_F$ . A range  $0.027g \leq \Delta N_{\text{eff}} \leq 0.07g$  corresponds to a range in  $T_F$  spanning decoupling during post-inflation reheating ( $0.027g$ ) down to lower  $T_F$  with decoupling occurring just prior to the QCD phase transition ( $0.07g$ ).

Information about  $N_{\text{eff}}$  is gleaned from the  $TT$  and  $EE$  power spectra. For an experiment like PICO, which has sufficient resolution to reach cosmic-variance-limited measurement<sup>2</sup> of  $EE$  up to  $\ell = 2300$ , the two additional most important parameters for improving constraints are the fraction of sky observed  $f_{\text{sky}}$  and the noise (Fig. 2.4, left). The PICO baseline will use data from 70% of the sky to constrain  $\Delta N_{\text{eff}} < 0.06$  (95%).<sup>3</sup> This constraint, which is a factor of 4.7 improvement relative to *Planck* ( $\Delta N_{\text{eff}} < 0.28$ , 95%) and will not be matched by any currently funded effort, opens up a new range of temperatures in which to detect the signature of light relic species. If no new species are detected, then the lowest temperature  $T_F$  at which any particle with spin could have fallen out of equilibrium will move up by a factor of 400 (Fig. 2.4, right).

While our theoretical target for  $N_{\text{eff}}$  is defined by particles that decoupled long before neutrinos did, there are a number of well-motivated scenarios in which the thermal evolution of the Standard Model is altered after the time of neutrino decoupling. These scenarios will change the relationship between  $N_{\text{eff}}$  as measured

<sup>2</sup>A measurement is cosmic-variance-limited when the measurement uncertainty is dominated by the statistics of observing the finite number of Fourier modes available in our Universe.

<sup>3</sup>The CMB  $EE$  and the Galactic foregrounds  $EE$  and  $BB$  spectra are comparable in level (Fig. 2.1). With 21 frequency bands PICO should be able to separate signals at the mild levels necessary for 70% of the sky.



Figure 2.5: **Left:** PICO will give a factor of 25 more stringent constraint on spin-independent velocity-independent dark matter scattering cross-section (dash) relative to current *Planck* 95% confidence limit (red) [50]. Terrestrial direct detection experiments are expected to give complementary and stronger constraints, but only for the higher dark matter masses (grey) [58]. **Right:** Using cosmic-variance-limited (CVL) measurement of  $\tau$ ,  $\sigma(\tau) = 0.002$ , BAO information from DESI, and 21 frequency bands to separate foregrounds over 70% of the sky, PICO will reach  $\sigma(\Sigma m_\nu) = 14 \text{ meV}$  (contours) giving at least  $4\sigma$  detection of the minimal expected sum of neutrino masses  $\Sigma m_\nu = 58 \text{ meV}$ .

in the CMB relative to the value of  $N_{\text{eff}}$  that affects BBN and the production of helium, which is quantified through the helium fraction  $Y_p$ . For example, the decay of a thermal relic into photons after nucleosynthesis would reduce  $N_{\text{eff}}$  in the CMB but would leave  $Y_p$  unaltered from its Standard Model value. PICO will make a simultaneous measurement of  $N_{\text{eff}}$  and  $Y_p$  with  $\sigma(N_{\text{eff}}) = 0.08$  and  $\sigma(Y_p) = 0.005$ , respectively, giving a 2% uncertainty on the value of  $Y_p$ . These uncertainties are equivalent to those available with other astrophysical measurements, but the systematic uncertainties are entirely different. Systematic uncertainties currently limit our knowledge of  $Y_p$ .

• **Dark Matter** Cosmological measurements have already confirmed the existence of one relic that lies beyond the Standard Model: dark matter. CMB experiments are effective in constraining dark matter candidates in the lower mass range, which is not available for terrestrial direct detection experiments [45, 46].

Interactions between dark matter and protons in the early Universe create a drag force between the two cosmological fluids, damping acoustic oscillations and suppressing power in density perturbations on small scales. As a result, the CMB temperature, polarization, and lensing power spectra are suppressed at high multipoles relative to a Universe without such drag forces. This effect has been used to search for evidence of dark matter-proton scattering over a range of masses, couplings, and interaction models [47–54], to test the possibility of an interacting dark-matter sub-component [53], and to provide consistency tests of dark matter in the context of the anomalous 21-cm signal reported by the EDGES collaboration [53, 55–57].

PICO’s constraining power comes primarily from making high SNR maps of the lensing-induced deflections of polarized photons, which are discussed in Section 2.3. For a spin-independent velocity-independent contact-interaction, chosen as our fiducial model, PICO will improve upon *Planck*’s dark matter cross-section constraints by a factor of 25 over a broad range of candidate masses that will not be probed by direct detection experiments (Fig. 2.5, right).

The axion is another dark matter candidate that is well motivated by string theory [59] and that is consistent with straightforward extensions of the standard model of particle physics [60–62]. For an axion mass in the intermediate range  $10^{-30} < m_a < 10^{-26} \text{ eV}$ , current measurements constrain its fraction to be equal to or less than 2% ( $1\sigma$ ) of the total dark matter density. If 2% of the total dark content is made of axions, PICO’s measurement of the  $TT$ ,  $TE$  and  $EE$  spectra with additional constraints from the lensing reconstruction will detect this species at between  $7 - 13\sigma$  depending on the mass range. This is an average improvement of a factor of 10 relative to *Planck*.

- **Neutrino Mass** The origin and structure of the neutrino masses is one of the great outstanding questions about the nature of the Standard Model particles. Cosmology offers a measurement of the sum of the neutrino masses  $\sum m_\nu$  through the gravitational influence of the non-relativistic cosmic neutrinos. The current measurement of  $N_{\text{eff}} = 2.99 \pm 0.17$  [63] already confirms the existence of these neutrinos at  $> 10\sigma$  and their mass implies that they will contribute to the matter density at low redshifts. The best current mass constraint arises from a combination of *Planck* and BOSS BAO giving  $\sum m_\nu < 0.12$  eV (95%) [63].

Cosmological measurements are primarily sensitive to the suppression of power on small scales after the neutrinos become non-relativistic, which can be measured via CMB lensing (§ 2.3.2), or weak lensing in galaxy surveys. However, these measurements are limited by our knowledge of the amplitude of the primordial fluctuation power spectrum  $A_s$  because they only constrain the combination  $A_s e^{-2\tau}$ . Although many surveys hope to detect  $\sum m_\nu$ , any detection of the minimum value expected from particle physics,  $\sum m_\nu = 58$  meV, at more than  $2\sigma$  will require a better measurement of  $\tau$ .

The best constraints on  $\tau$  come from the  $EE$  spectrum at  $\ell < 10$ , which require measurements over the largest angular scales, and good separation of Galactic foreground sources of emission. The current limit of  $\sigma(\tau) = 0.007$  is from *Planck* [63]. With the current uncertainty in  $\tau$  one is limited to  $\sigma(\sum m_\nu) \gtrsim 25$  meV, after including forthcoming BAO information (Fig. 2.5, right); no other survey or cosmological probe will improve this constraint, unless a more accurate measurement of  $\tau$  is made. One of the S3 experiments is attempting to measure the lowest  $\ell$ 's [38]. A space mission with its access to the entire sky and broad frequency coverage is the most suitable platform for the measurement. PICO will reach the cosmic-variance limit uncertainty on  $\tau$ ,  $\sigma(\tau) \sim 0.002$ , and using its deep CMB lensing map (§ 2.3.2) will therefore reach  $\sigma(\sum m_\nu) = 14$  meV when combined with measurements of BAO from DESI or Euclid [64] (without BAO data the constraint is  $\sigma(\sum m_\nu) = 43$  meV). This measurement will give a  $4\sigma$  detection of the minimum sum.

### 2.2.3 Fundamental Fields: Primordial Magnetic Fields and Cosmic Birefringence

- **Primordial Magnetic Fields** One of the long-standing puzzles in astrophysics is the origin of observed 1-10  $\mu$ G galactic magnetic fields [65]. Producing such fields through a dynamo mechanism requires a primordial seed field [66]. Moreover,  $\mu$ G-strength fields have been observed in proto-galaxies that are too young to have gone through the number of revolutions necessary for the dynamo to work [67]. A primordial magnetic field (PMF), present at the time of galaxy formation, could provide the seed or even eliminate the need for the dynamo altogether. Specifically, a 0.1 nG field in the intergalactic plasma would be adiabatically compressed in the collapse to form a  $\sim 1$   $\mu$ G galactic field [68]. PMFs could have been generated in the aftermath of phase transitions in the early Universe [69], during inflation [70, 71], or at the end of inflation [72]. A detection of PMFs with the CMB would be a major discovery as it would establish the magnetic field's primordial origin, signal new physics beyond standard models of particle physics and cosmology, and discriminate among different theories of the early Universe [73–75].

The current CMB bounds on PMF strength are  $B_{1\text{Mpc}} < 1.2$  nG at 95% CL for the scale-invariant PMF spectrum [76], based on measurements of the  $TT$ ,  $TE$ ,  $EE$  and  $BB$  spectra.<sup>4</sup> The much more accurate measurement of  $BB$  by PICO would only marginally improve the PMF bound because CMB spectra scale as  $B_{1\text{Mpc}}^4$ . However, Faraday rotation provides a signature that scales linearly with the strength of PMF [77]. It converts CMB  $E$  modes into  $B$  modes, generating mode-coupling  $EB$  and  $TB$  correlations. So far this signature was out of reach because prior experiments did not have sufficient sensitivity. Using Faraday rotation, PICO will probe PMFs as weak as 0.1 nG ( $1\sigma$ ), a limit that already includes the effects of imperfect lensing subtraction, Galactic foregrounds [78–80], and other systematic effects. With this limit, which is a factor of five stronger than achievable with S3 experiments, PICO will conclusively rule out the purely primordial (i.e. no-dynamo driven) origin of the largest galactic magnetic fields.

- **Cosmic Birefringence** A number of well-motivated extensions of the Standard Model involve (nearly)

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<sup>4</sup>It is conventional to quote limits on the PMF strength smoothed over a 1 Mpc region in comoving units, *i.e.* rescaled to  $z = 0$ :  $\mathbf{B}_{\text{today}} = a^2 \mathbf{B}(a)$ .



Figure 2.6: Contours of 1 and  $2\sigma$  constraints on the mean redshift and duration of reionization using PICO and CMB-S3 data (solid dark blue), and comparison with *Planck* and CMB-S3 (dash light blue). Source efficiency and IGM opacity (dark lines) are two physical parameters controlling the reionization process in current models. The PICO measurements, together with higher resolution data of the kSZ effect, will significantly constrain the range of models allowed. We also include other constraints from *Planck*, EDGES, the Gunn-Peterson (GP) trough, and *Planck*+ the South Pole Telescope [63, 96–98].

massless axion-like pseudo-scalar fields coupled to photons via the Chern-Simons interaction term [81–84]. These couplings also generically arise within quintessence models for dark energy [83], chiral-gravity models [85], and models that produce parity-violation during inflation [86]. Regardless of the source of the parity-violating coupling, its presence may cause cosmic birefringence—a rotation of the polarization of an electromagnetic wave as it propagates across cosmological distances [83, 87, 88]. Cosmic birefringence converts primordial  $E$ -modes into  $B$ -modes, producing  $TB$  and  $EB$  cross-correlations whose magnitude depends on the statistical properties of the rotation field in the sky [89–91]. Previous studies have constrained both a uniform rotation angle as well as anisotropic rotation described by a power spectrum [91]. The current bound on a uniform angle is 30 arcmin (68%) [92], and the bound on the amplitude of a scale-invariant rotation angle spectrum is  $0.11 \text{ deg}^2$  (95%) [93]. Using the combination of 5 bands in the  $70 - 156 \text{ GHz}$  range, PICO will reduce the 95% CL bound on the uniform rotation angle by a factor of 300 to  $0.1'$ . The 95% CL bound on the amplitude of a scale-invariant rotation spectrum will be reduced by a factor of 275 to  $4 \times 10^{-4} \text{ deg}^2$ , giving important constraints on string-theory-motivated axions [94, 95].

## 2.3 Cosmic Structure Formation and Evolution

### 2.3.1 The Formation of the First Luminous Sources

A few hundred million years after the Big Bang, the neutral hydrogen gas permeating the Universe was reionized by photons emitted by the first luminous sources to have formed. The nature of these sources and the exact history of this epoch are key missing links in our understanding of structure formation (SO5).

The reionization of the Universe imprints multiple signals in the temperature and polarization of the CMB. In polarization, the most important signature is an enhancement in the  $EE$  power spectrum at large angular scales  $\ell \lesssim 10$  (Fig. 2.1). This signal gives a direct measurement of the optical depth to the reionization epoch  $\tau$  and thus to the mean redshift of reionization  $z_{re}$ , with very little degeneracy with other cosmological parameters (Fig. 2.6).<sup>5</sup> *Planck*'s determination of the optical depth to reionization  $\tau = 0.054 \pm 0.007(1\sigma)$  has indicated that reionization concluded by  $z \approx 6$ , but the measurement uncertainty leaves many unanswered questions including: were the ionizing sources primarily star-forming galaxies or more exotic sources such as supermassive black holes or annihilating dark matter? What was the mean free path of ionizing photons during this epoch? What was the efficiency with which such photons were produced by ionizing sources? Did the reionization epoch extend to  $z \approx 15-20$ , as has been claimed recently? [99] PICO's cosmic-variance-limited measurement of the large-scale  $E$  modes reaching  $\sigma(\tau) = 0.002$  will settle some of these questions and will significantly constrain the others.

Figure 2.6 presents forecasts for reionization constraints in the  $z_{re} - \Delta z_{re}$  parameter space. These are obtained from PICO's measurement of  $\tau$  in combination with S3 experiments measurements of the “patchy”

<sup>5</sup>The mean redshift to reionization is the redshift when 50% of the cosmic volume was reionized.

kinematic Sunyaev-Zel'dovich (kSZ) effect, due to the peculiar velocities of free electron bubbles around ionizing sources [100]. The figure includes curves of constant efficiency of production of ionizing photons in the sources, and of intergalactic-medium opacity, two parameters that quantify models of reionization. The curves shown are illustrative; families of models, that would be represented by parallel ‘source efficiency’ and ‘IGM Opacity’ lines, are allowed by current data. PICO’s data will give simultaneous constraints on these physical parameters, yielding important information on the nature of the first luminous sources. For example, galaxies and quasars predict significantly different values for their IGM opacities and source efficiencies.

The process of reionization leaves specific non-Gaussian signatures in the CMB. In particular, patchy reionization induces non-trivial 4-point functions in both temperature and polarization [101, 102]. The temperature 4-point function can be used to separate reionization and late-time kSZ contributions. Combinations of temperature and polarization data can be used to build quadratic estimators for reconstruction of the patchy  $\tau$  field, analogous to CMB lensing reconstruction (next Section). These estimators generally require high angular resolution, but also rely on foreground-cleaned CMB maps. PICO’s data in its high-frequency bands — which have better than 2 arcmin resolution and cover frequencies that are not suitable for observations from the ground — will enable these estimators to be robustly applied to high resolution ground-based CMB data, a strong example of ground-space complementarity.

With ten independent maps of the entire sky, multiple frequency bands and ample sensitivity to remove foregrounds, PICO is uniquely suited to make the low  $\ell$   $EE$ -spectrum measurements that are needed to elucidate the formation of the first luminous sources. No such measurements have yet been done from the ground. Of the currently operating and planned S3 experiments, one is targeting the lowest  $\ell$   $EE$  spectrum [8].

Lowering the uncertainty on  $\tau$  is crucial for many cosmological observables of the growth of structure. As discussed in Section 2.2.2, these observables require knowledge of the amplitude of the initial fluctuations power spectrum  $A_s$ ; however, the signatures of  $A_s$  and  $\tau$  are degenerate in the CMB power spectra, except in the low- $\ell$   $EE$  power spectrum. PICO’s cosmic-variance-limited polarization measurements will break this degeneracy and thus improve constraints on the sum of neutrino masses, on dark energy, and modified gravity coming from *any* low-z growth measurement such as CMB and galaxy lensing, velocity-field measurements, redshift-space distortions, and galaxy surveys.

### 2.3.2 Probing the Evolution of Structures via Gravitational Lensing and Cluster Counts

The particle content of the Universe, gravitational collapse, the effects of dark energy, and energetic feedback processes that recycle energy determine the evolution of structures in the Universe. The amplitude of linear fluctuations as a function of redshift, parameterized through  $\sigma_8(z)$ , is thus a sensitive probe representing the effects of physical process affecting growth. CMB photons are affected by, and thus probe  $\sigma_8(z)$  as they traverse the entire Universe. PICO will tightly constrain  $\sigma_8(z)$  through measurements of gravitational lensing and cluster counts.

- **Gravitational Lensing** Matter between us and the last-scattering surface deflects the path of photons through gravitational lensing, imprinting the 3-dimensional matter distribution across the volume of the Universe onto the CMB maps. The specific quantity being mapped by the data is the projected gravitational potential  $\phi$  that is lensing the photons. From the lensing map, which receives contributions from all redshifts between us and the CMB, with the peak of the distribution at  $z \simeq 2$ , we infer the angular power spectrum  $C_L^{\phi\phi}$  (Fig. 2.7). Both the temperature and polarization maps of the CMB, and by extension the angular power spectra, are affected by lensing.

*Planck*'s  $\phi$  map had SNR of  $\sim 1$  per  $L$  mode over a narrow range of scales,  $30 < L < 50$ .<sup>6</sup> PICO’s map would represent true mapping, with  $\text{SNR} \gg 1$  for each mode in the range  $2 \leq L \lesssim 1000$ . While *Planck* had an SNR of 40 integrated across the entire  $C_L^{\phi\phi}$  power spectrum [103], PICO will give SNR of 560 and 644 for the baseline and CBE configurations, respectively; both values already account for foreground separation

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<sup>6</sup> $L$  refers to multipoles in the CMB lensing and galaxy clustering fields, in contrast to the use of  $\ell$  for the CMB itself.



Figure 2.7: PICO will make a high SNR full sky map of the projected gravitational potential  $\phi$  due to all matter between us and the last scattering surface at all angular scales  $L$  for which its noise (red and blue) is below the theoretically predicted power spectrum  $C_L^{\phi\phi}$  (black). Noise predictions are given with (dash) and without (solid) foregrounds separation, and we include values for the SNR for the measurement of  $C_L^{\phi\phi}$  for  $2 \leq L \leq 1000$ .

(Fig. 2.7).

PICO’s  $\phi$  map is a key ingredient in the delensing process that improves constraints on  $r$  (§ 2.2.1) and in extracting neutrino mass constraints (§ 2.2.2). It will also be used to constrain the properties of quasars and other high-redshift astrophysics. For example, cross-correlations with quasar samples from DESI will yield a precise determination of the quasar bias (and hence host halo mass) as a function of the quasar properties, such as (non-)obscuration. Such studies are not possible with any other lensing techniques, due to their sensitivity to lower redshifts.

- **$\sigma_8(z)$  from Gravitational Lensing** Cross-correlations between the PICO lensing potential map and wide-field samples of galaxies and quasars provide a powerful technique to measure the time dependence of the amplitude of matter fluctuations  $\sigma_8(z)$  in tomographic redshift bins. This is achieved by overcoming the limitations of auto-correlations of these data sets: The lensing  $\phi$  map is sensitive to the projection of all matter back to the last scattering surface, so it cannot resolve the time dependence of fluctuations, while galaxies and quasars trace matter in an unknown biased way so that the matter amplitude cannot be determined. Cross-correlations of the two data sets, broken down to several tomographic redshift bins, will constrain how galaxies in each bin trace the dark matter, which will yield strong constraints on  $\sigma_8(z)$  and thereby on structure formation and models of dark energy and modified gravity [40, 41].

In Fig. 2.8 we show projected  $1\sigma$  errors on  $\sigma_8(z)$  when using cross-correlations with LSST’s gold sample of galaxies [104]. Sub-percent accuracy is obtainable with PICO’s resolution which will give information extending to  $L = 1000$ .<sup>7</sup> This accuracy will be used to constrain dark energy or modified gravity, in the context of specific models, and to give a neutrino mass constraint that is independent from and competitive with that inferred from the CMB lensing auto-power spectrum (§ 2.2.2) [105].

- **Cluster Counts** The distribution of galaxy clusters in redshift is one consequence of the evolution of structures and is thus a sensitive measure of  $\sigma_8(z)$ . The observational quantity of interest is  $dN/dz dm$ , the number of observed clusters per redshift and per mean mass, from which constraints on  $\sigma_8$  can be derived. Galaxy clusters found by PICO via the thermal Sunyaev–Zel’dovich (tSZ) effect (§ 2.3.3) provide a catalog with a selection function that is simple to model and thus straightforward to use for cosmological inference. PICO’s catalog will provide all clusters with masses above  $\sim 3 \times 10^{14} M_\odot$ <sup>8</sup> check values out to redshifts  $z \sim 3$ , as long as the clusters have started to virialize. We forecast that PICO will find  $\sim 150,000$  galaxy clusters, assuming the cosmological parameters from *Planck* and using the 70% of sky not obscured by the Milky Way. Redshifts will be provided by optical and infrared follow-up surveys. Cluster masses will be inferred by optical weak lensing for clusters with  $z < 1.5$  and by PICO’s own CMB halo lensing data at higher redshifts (see discussion of halo lensing below). The catalog will give the most massive clusters over the full sky. Sub-percent determination of  $\sigma_8$  will be provided with this catalog for  $0.5 < z < 2$  (Fig. 2.8),

<sup>7</sup>PICO’s resolution is sufficient to give information for  $L > 1000$ , but at these scales structures are non-linear and will not be used to constrain  $\sigma_8(z)$ .

<sup>8</sup> $M_\odot$  is one solar mass



Figure 2.8: Sub-percent constraints on the evolution of  $\sigma_8$  as a function of redshift will come from two independent PICO products: Correlations between PICO’s deep gravitational lensing map (Fig. 2.7) and LSST’s gold sample of galaxies (left) and cluster counts (right). Fractional uncertainties in  $\sigma_8$  relative to fiducial  $\Lambda$ CDM values are given as a function the finest angular scale  $L_{max}$  of the correlation analysis for seven redshift bins (left). The baseline and CBE configurations give essentially the same fractional errors of  $\sigma_8(z)$  using cluster counts (right). LSST assumptions: 10 years, 50% sky fraction, 55 galaxies per arcmin<sup>2</sup> at redshift  $z < 3$  with magnitude limit  $i < 25.3$  [104], and dropout galaxies at  $z > 3$  [106] extrapolating recent Hyper Suprime-Cam observations [107–109], with linear bias  $b(z) = 1 + z$ .

and a neutrino mass constraint of  $\sim 10$  meV that is independent from the one coming from the lensing measurements (§ 2.2.2).

Calibrating the masses of clusters, that is determining  $m(z)$ , is the most uncertain step in inferring  $\sigma_8$  and other cosmological parameters using cluster counts. PICO will provide calibration using ‘CMB halo lensing’, an approach that uses the small-scale effects of gravitational lensing due to dark matter halos around clusters and proto-clusters [110–112]. The technique is particularly effective for measuring halo masses out to high redshifts where gravitational lensing of background objects no longer works because there are no background sources. The approach is illustrated in Fig. 2.9, which gives the  $1\sigma$  uncertainty in a halo mass measurement as a function of the object’s redshift. PICO will measure the mass of individual low-mass clusters ( $\sim 10^{14} M_\odot$ ) over a wide redshift range, and by stacking will determine the mean mass of smaller halos, with masses of  $\sim 10^{13} M_\odot$ , which include those hosting individual galaxies. Because the vast majority of clusters have masses that are larger than  $\sim 10^{14} M_\odot$ , the PICO data will provide mass calibration for all objects of interest. The flattening at high redshift reflects the fact that the technique is sensitive over a broad range of redshifts. The high-frequency PICO data, at which the resolution matches ground-based instruments’ resolution at lower frequencies, will play an essential role in cleaning foregrounds particularly those derived from the temperature-based estimator, which is most contaminated by foregrounds.

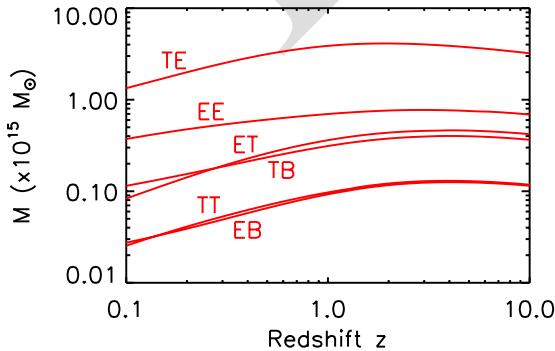


Figure 2.9: PICO will provide mass calibration for individual clusters and proto-clusters with mass as low as  $10^{14}$  solar masses at  $z > 2$  using ‘halo lensing’. Curves for different CMB signal correlations (red) give the  $1\sigma$  sensitivity of an optimal mass filter [113] for a given mass as a function of  $z$ . The curves are flat at high redshift, demonstrating that the technique probes a broad range of redshifts. For PICO, the  $EB$  and  $TT$  estimators are equivalent, offering important cross-validation of measurements because the systematics are very different for temperature and polarization.

Beyond its role in calibrating masses for cluster counts, PICO’s halo lensing measurements will also be

a unique tool for measuring the relation between galaxies and their dark matter halos during the key epochs of cosmic star formation at  $z \geq 2$ , which is not reachable by other means. This will provide valuable insight into the role of environment on galaxy formation during the rise to and fall from the peak of cosmic star formation at  $z \sim 2$ .

### 2.3.3 Constraining Feedback Processes through the Sunyaev–Zel'dovich Effect

Not all CMB photons propagate through the Universe freely; about 6% are Thomson-scattered by free electrons in the inter-galactic medium (IGM) and inter-cluster medium (ICM). These scattering events leave a measurable imprint on CMB temperature fluctuations, which thereby contain a wealth of information about the growth of structures and the thermodynamic history of baryons. A fraction of these photons are responsible for the thermal and kinetic Sunyaev–Zel'dovich effects (tSZ and kSZ) [114, 115]. The amplitudes of the tSZ and kSZ signals are proportional to the integrated electron pressure and momentum along the line of sight, respectively. They thus contain information about the thermodynamic properties of the IGM and ICM, which are highly sensitive to astrophysical feedback. Feedback is the process of energy injection into the IGM and ICM from accreting supermassive black holes, supernovae, stellar winds, and other sources. Feedback processes are the most uncertain, yet crucial, ingredient in modern theories of galaxy formation; they are required in order to match observations of the stellar properties of galaxies, but the underlying details of the physical processes involved are still highly uncertain.

Multifrequency CMB data also allow the reconstruction of full-sky maps of the tSZ signal, commonly called ‘Compton-y map’. With low noise and broad frequency coverage, which is essential for separating out other signals, PICO will yield a definitive Compton-y map over the full sky, with a total SNR of 1270 for the CBE and  $\approx 10\%$  lower for the baseline configurations (Fig. 2.10). This is nearly two orders of magnitude higher SNR than *Planck*, which already gave data with much higher SNR than ground-based experiments.

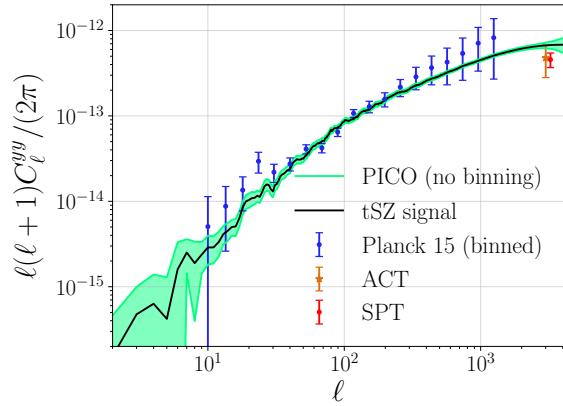


Figure 2.10: The PICO y-map will give a tSZ power spectrum with an SNR of 1270 (green,  $1\sigma$  per  $\ell$  mode), which is nearly 100 times larger than from *Planck* (blue). Binning the data (not shown) as was done for *Planck* would further increase the SNR. We include current measurements by the ground-based SPT and ACT [116, 117]. In these forecasts we reconstruct the Compton-y field from maps that include Galactic foregrounds, CMB fluctuations, and PICO CBE noise using the needlet internal linear combination algorithm [118]. The input maps use the *Planck* sky model [119].

Strong constraints on models of astrophysical feedback will be obtained from the analysis of the PICO y-map, both from its auto-power spectrum and from cross-correlations with galaxy, group, cluster, and quasar samples. As an example, we forecast the detection of cross-correlations between the PICO y-map and galaxy weak lensing maps constructed from LSST and WFIRST data. Considering the LSST ‘gold’ sample with a source density of 26 galaxies/arcmin<sup>2</sup> covering 40% of the sky, we forecast a detection of the tSZ-weak lensing cross-correlation with SNR = 3000. Cross-correlations with the galaxies themselves will be measured at even higher SNR. At this immense significance, the signal can be broken down into dozens of tomographic redshift bins, yielding a precise tracing of the evolution of thermal pressure over cosmic time. For PICO and WFIRST (assuming 45 galaxies/arcmin<sup>2</sup> covering 5.3% of the sky), we forecast SNR = 1100 for the tSZ-weak lensing cross-correlation. The WFIRST galaxy sample extends to higher redshift, and thus this high-SNR measurement will allow the evolution of the thermal gas pressure to be probed to  $z \approx 2$  (the peak of the cosmic star formation history) and beyond. These measurements will revolutionize our understanding of galaxy formation and evolution by distinguishing between models of feedback energy injection

at high significance. Additional cross-correlations of the PICO  $y$ -map with quasar samples, filament catalogs, and other large-scale structure tracers will provide valuable information on baryonic physics that is complementary to inferences from the lensing cross-correlations described earlier.

## 2.4 Galactic Structure and Star Formation

*Planck* enabled an immense step forward in Galactic astrophysics [120]. With 7 full sky polarization maps at frequencies between 30 and 353 GHz and a highest resolution of 5', *Planck* provided entirely new and surprising data about the structure of the interstellar matter (ISM); the data have a lasting legacy for the foreseeable future. PICO will provide an even greater leap forward. It will produce 21 polarization maps of Galactic emissions, and for similar frequency bands to *Planck* they will be much deeper; for example, at 321 (30) GHz, PICO's map is 105 (17) times deeper than *Planck*'s mean map depth at 353 (30) GHz. At 799 GHz PICO will have 5 times the resolution of *Planck*'s highest resolution map (Fig. 2.11). Such a data set can only be obtained from space. These data will complement a rich array of other polarization observations forthcoming in the next decade, including stellar polarization surveys to be combined with Gaia astrometry, and synchrotron observations with the SKA (and its precursors) to measure Faraday rotation at radio wavelengths.

While the PICO data will likely provide many new insights and surprises, we focus here on two particularly important science objectives that are integral to NASA's science goal to explore how the Universe evolved; they relate to the structure and evolution of the Milky Way. These objective can only be addressed by the PICO dataset.

(1) *Test Composition Models of Interstellar Dust:* Less than one thousandth of a millimeter in size, dust grains are intermediate in the evolution from atoms and molecules to large solid bodies such as comets, asteroids, and planets. Encoded in the composition of dust are the pathways through which grains formed and grew. Dust grains also participate directly in interstellar chemistry, for example by catalyzing the formation of H<sub>2</sub> and organic molecules on their surfaces, in ways that depend upon their chemical makeup. Thus, the composition of dust grains is an essential aspect of the chemical evolution of interstellar matter from the formation of complex molecules in space to the growth of planets. Through vastly improved spectral characterization of Galactic polarization, the PICO data will discriminate among models of Galactic dust composition to elucidate the chemical evolution of the Galaxy (SO6). The data will also inform methods to separate diffuse dust emission from cosmological signals of interest .

(2) *Determine how magnetic fields affect the processes of molecular cloud and star formation:* Stars are formed through interactions between gravitational and magnetic fields, turbulence, and gas over more than four orders of magnitude of spatial scales. However, the role magnetic fields play in the large scale structure of the diffuse ISM and in the observed low star-formation efficiency has eluded answer because of dearth of data. By virtue of the strong dynamical coupling of dust and gas and the systematic alignment of dust grains with magnetic fields, PICO's dust polarization measurements will for the first time probe the large scale Galactic magnetic field with resolution to trace the role of magnetic fields through the entirety of the star formation process (SO7).

### 2.4.1 Test Composition Models of Interstellar Dust

Strong extinction features at 9.7 and 18  $\mu\text{m}$  indicate that much of interstellar dust is in the form of amorphous silicates while features at 2175 Å, 3.3  $\mu\text{m}$ , and 3.4  $\mu\text{m}$  attest to abundant hydrocarbons. It is unknown, however, whether the silicate and carbonaceous materials coexist on the same grains or whether grains of each composition grow through distinct, parallel pathways dictated by their surface chemistry.

Some data suggest that the populations are distinct. Spectropolarimetry of dust extinction reveals robust polarization in the 9.7  $\mu\text{m}$  silicate feature [e.g., 123], indicating that the silicate grains are aligned with the interstellar magnetic field. In contrast, searches for polarization in the 3.4  $\mu\text{m}$  carbonaceous feature have yielded only upper limits, even along sightlines where silicate polarization is observed [124, 125]. These

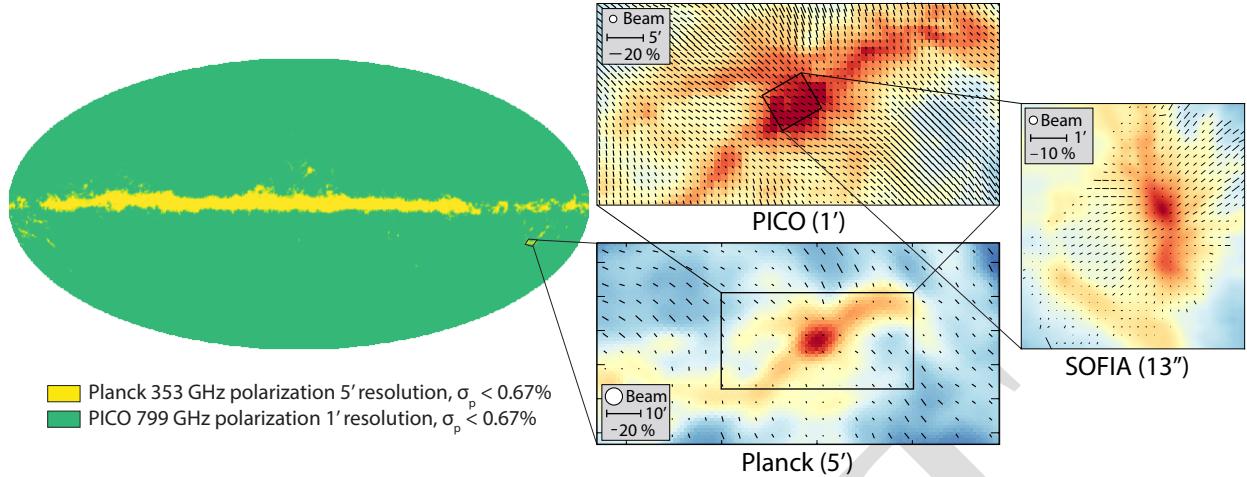


Figure 2.11: *Planck*'s 353 GHz polarization map gave a resolution of 5' and sensitivity to polarization intensity of  $\sigma_p < 0.67\%$  over a small portion of the sky (left, yellow). At 799 GHz, the PICO baseline mission will give a polarization map of the *entire* sky and with 5 times higher resolution (left, green). The *Planck* map of the Orion region overlaid with vectors that are aligned with the inferred magnetic field (lower panel), and a simulated PICO observation (upper panel) illustrate the leap in information content (vector lengths are proportional to polarization fraction). With this map, and maps at many other frequencies that *Planck* did not have, PICO will characterize Galactic magnetized turbulence at scales spanning the diffuse ISM down to dense star forming cores, which will be mapped with high-resolution polarimetry by instruments such as HAWC+/SOFIA [121] (right panel) and the Atacama Large Millimeter Array [122].

data are consistent with silicate and carbonaceous materials existing on separate grains that have different alignment properties.

At odds with the spectropolarimetric evidence from dust extinction are current measurements of the polarization fraction of the far-infrared dust emission with *Planck* [126] and BLASTPol [127]. They show little to no frequency dependence, whereas substantial frequency dependence would be expected if two components with distinct polarization properties were contributing to the total emission.

With excellent polarization sensitivity, even in diffuse regions, PICO will provide a definitive test of the two component paradigm [128]. In this case, the PICO baseline mission will determine the intrinsic polarization fractions of each of the two components to a precision of 3%. With this level of precision the data will validate or reject state-of-the-art dust models [e.g. 129, Hensley & Draine, in prep], test for the presence of additional grain species with distinct polarization signatures, such as magnetic nanoparticles [130], and will be used as a crucial input for the foreground separation necessary to extract cosmological B-mode science.

'Anomalous Microwave Emission (AME)' is a component of Galactic emission peaking in the 20-30 GHz range that has been tentatively identified with small, rapidly-spinning dust grains [131]. As only upper limits have been placed on its polarization, its role as a foreground for cosmological B-mode science remains unclear. Given the present uncertainty on its physical origin and SED variability, even small levels of polarization could prove challenging. PICO will finely sample the AME SED with its bands at 21, 25, 30, 36, and 43 GHz. Combined with ground-based maps at lower frequencies, for example C-BASS at 5 GHz [132], PICO will be used to efficiently separate the AME from synchrotron and free-free emission and either detect or place stringent upper limits on its polarization. Further, the enhanced frequency coverage will enable characterization of systematic changes in the AME SED with interstellar environment and thus elucidate its underlying physics.

## 2.4.2 Determine How Magnetic Fields Affect the Processes of Molecular Cloud and Star Formation

Stars form out of dense, gravitationally unstable regions within molecular gas clouds, which themselves form through the flow of diffuse, atomic-phase gas to denser regions. Magnetic fields play an important role throughout this process.

On the largest scales, magnetized turbulence mediates the flow of the gaseous ISM from the atomic to the denser, molecular phase. Recent observations suggest that the structure of the diffuse medium is highly anisotropic, and strongly coupled to the local magnetic field [133–136]. As molecular gas clouds collapse to form stars, magnetic fields can slow the process of star formation by inhibiting movement of gas in the direction perpendicular to the field lines. Observations to date suggest that the outer envelopes of clouds can be supported against gravity by magnetic fields and turbulence, but in dense cores gravity tends to dominate, and so these dense structures can collapse to form stars [137]. The degree to which magnetic fields affect the formation of molecular clouds, as well as stars within these clouds, is poorly constrained, in large part due to the difficulty of making detailed maps of magnetic fields in the ISM.

• **Formation of Magnetized Molecular Clouds from The Diffuse Interstellar Medium** A comprehensive understanding of the magnetized diffuse ISM is challenging because of its diverse composition, its sheer expanse, and the multi-scale nature of the physics that shapes it. To understand how matter and energy are exchanged between the diffuse and dense media, it is essential to measure the properties of the magnetic field over more than four orders of magnitude in column density. PICO is unique in its ability to provide the necessary data. *Planck* achieved measurements of the diffuse sky at  $60'$  resolution, resulting in  $\sim 30,000$  independent measurements of the magnetic field direction. With  $1.1'$  resolution PICO will expand the number of independent polarization measurements to  $86,000,000$ . The data will thus robustly characterize turbulent properties like the Alfvén Mach number,  $\mathcal{M}_A$ , across a previously unexplored regime of parameter space.

PICO’s observations will complement recently completed high dynamic range neutral hydrogen surveys, such as HI4PI [138] and GALFA-HI [139], as well as planned surveys of interstellar gas, most prominently with the Square Kilometer Array (SKA) and its pathfinders. One of the open questions in diffuse structure formation is how gas flows within and between phases of the ISM. A planned all-sky absorption line survey with the forthcoming SKA-1 will increase the number of measurements of the ISM gas temperature by several orders of magnitude [140]. Quantitative comparisons of the ISM temperature distribution from SKA-1 and estimates of the magnetic field strength and coherence length scale from PICO will elucidate the role of magnetized turbulence in the flow of matter in the ISM from diffuse regions to regions of denser molecular gas.

• **Formation of Stars within Magnetized Molecular Clouds** The role of magnetic field in star-formation is quantified by the ratio of energy stored in magnetic and gravitational fields, and the ratio of energy stored in magnetic field and that stored in turbulence. The first ratio is parameterized through a mass-to-flux ratio  $\mu$ , and the second through  $\mathcal{M}_A$ .

With full-sky coverage and a resolution of  $1.1'$ , PICO will map all molecular clouds with better than  $1\text{ pc}$  resolution, out to a distance of  $3.4\text{ kpc}$ . Extrapolating from the Bolocam Galactic Plane Survey [BGPS, 141], PICO is expected to make highly detailed magnetic field maps of over  $2,000$  molecular clouds with thousands to hundreds of thousands of independent polarization measurements per cloud. These are the *only foreseeable* measurements that will give  $\mu$  and  $\mathcal{M}_A$  over a statistically significant sample of molecular clouds. *Planck*, for example, mapped only  $10$  nearby clouds to a similar level of detail [142]. A large sample of clouds is crucial because (1) dust polarization observations are sensitive to only the magnetic field projected on the plane of the sky, and therefore polarization maps will look very different for molecular clouds observed at different viewing angles; and (2) the relative importance of the magnetic field will likely be a function of cloud age and mass. By observing thousands of molecular clouds PICO will determine  $\mu$  and  $\mathcal{M}_A$  for different sub-classes of cloud age and mass.

### 2.4.3 Galactic Legacy Science

PICO will also produce legacy datasets that will revolutionize our understanding of how magnetic fields influence physical processes ranging from planet formation to galaxy evolution. For 10 nearby clouds, which have distances of less than 500 pc, PICO will resolve magnetic fields on scales of 0.1 pc. This is the scale of dense cores and filaments for these clouds, and thus the observations will constrain how magnetic fields on these scales influence the formation of cloud cores. By comparing the orientation of the core-scale magnetic fields with the orientation and sizes of proto-planetary disks, PICO will probe whether magnetic braking influences the growth of such disks [143, 144] and provide complementarity to higher angular resolution instruments such as ALMA and SOFIA [122, 145].

Key processes in the diffuse ISM, including heat transport, streaming of cosmic rays, and magnetic reconnection strongly depend on the level of the environment's magnetization [146–148]. PICO will give information about these processes with tens of millions of independent measurements of magnetic field orientation over the entire Galaxy. The measurements will also enable studying the physics of how magnetic fields are generated through a combination of turbulence and large-scale gas motions [149].

Finally, PICO observations will create detailed magnetic field maps of approximately 70 nearby galaxies, with more than 100 measurements of magnetic field directions per galaxy. These observations will determine how interaction between large-scale magnetic fields, turbulence, and feedback from previous generations of star formation affect galaxy evolution and star formation efficiency.

## 2.5 Legacy Surveys

PICO was designed to respond to requirements posed by the 7 SOs listed in Table 2.2. It will also generate a rich catalog of hundreds of thousands of new sources consisting of proto-clusters, strongly lensed galaxies, and polarized radio and dusty galaxies. An abundance of information about galaxy and cluster evolution, dark matter, the physics of jets of active galactic nuclei, and magnetic fields of dusty galaxies will be stored in this catalog (Table 2.3). This catalog will be mined in future years through subsequent analysis and follow-up observations.

### 2.5.1 Early Phases of Galaxy Evolution

PICO's catalog of high- $z$  strongly-lensed galaxies will provide answers to major, still open issues in galaxy formation and evolution. What are the main physical mechanisms shaping the properties of galaxies [150, 151]: in situ processes, interactions, mergers, or cold flows from the intergalactic medium? And how do feedback processes work? To settle these issues we need direct information on the structure and dynamics of high- $z$  galaxies. But these are compact, with typical sizes of 1–2 kpc [152]), corresponding to angular sizes of 0.1–0.2 arcsec at  $z \simeq 2$ –3. Thus they are hardly resolved, even by ALMA or by HST. If they are resolved, high enough SNRs per resolution element are only achieved for the brightest galaxies, which are probably not representative of the general population.

Strong gravitational lensing provides a solution to these problems. Since lensing conserves the surface brightness, the effective angular size is stretched on average by a factor of  $\mu^{1/2}$ , where  $\mu$  is the gravitational magnification, thus substantially increasing the resolving power. A spectacular example is ALMA observations of the *Planck*-discovered, strongly lensed galaxy PLCK\_G244.8+54.9 at  $z \simeq 3.0$  with  $\mu \simeq 30$  [153]. ALMA observations with a 0.1'' resolution reached an astounding spatial resolution of 60 pc, substantially smaller than the size of a Milky Way giant molecular clouds. CO spectroscopy of this object, measuring the kinematics of the molecular gas, gave an uncertainty of 40–50 km/s. Such precision allows a high SNR detection of the predicted  $\sim 1000 \text{ km s}^{-1}$  outflows capable of sweeping the galaxy clear of gas that would otherwise be available for star formation [154]. In this specific case, there were no clear indications that mergers or cold flows shaped the galaxy, but similar spectroscopy of another strongly lensed galaxy at  $z = 5.3$  detected a fast (800 km/s) molecular outflow due to feedback.

PICO will detect thousands of early forming galaxies whose flux densities are boosted by large factors

Table 2.3: Legacy Surveys

Catalog	Impact	Science
Strongly lensed galaxies	<p>Discover <math>4500^a</math> strongly lensed and highly magnified dusty galaxies across redshift.</p> <p>Current knowledge: 13 sources confirmed in <i>Planck</i> data; few hundred candidates in <i>Herschel</i>, SPT and ACT data.</p>	Gain information about the physics governing early, $z \simeq 5$ , galaxy evolution, taking advantage of magnification and extra resolution enabled by gravitational lensing; learn about dark matter sub-structure in the lensing galaxies.
Proto-clusters	<p>Discover <math>50,000^a</math> mm/sub-mm proto-clusters distributed over the sky out to <math>z \sim 4.5</math>.</p> <p>Current knowledge: <i>Planck</i> ACT/SPT data expected to yield a few tens.</p>	Probe the earliest phases of cluster evolution, well beyond the reach of other instruments; test the formation history of the most massive virialized halos; investigate galaxy evolution in dense environments.
Nearby galaxies	<p>Detect 30,000 galaxies at <math>z \lesssim 0.1</math> at frequencies above 300 GHz.</p> <p>Current knowledge: 3400 (280) source candidates with <i>Planck</i> 857 (353) GHz band.</p>	Using frequencies that match cold (15 – 25 K) dust emission, give its spectral energy distribution as a function of galaxy properties to enable correlations with star formation activity
Polarized point sources	<p>Detect <math>2000^b</math> radio and several thousand dusty galaxies in polarization.</p> <p>Current knowledge: about 200 radio sources up to 100 GHz; one polarization measurement of a dusty galaxy.</p>	Study the physics of jets of extragalactic sources, close to their active nuclei; determine the large-scale structure of magnetic fields in dusty galaxies; determine the importance of polarized sources as a foreground for CMB polarization science.
Cosmic infrared background	<p>Provide eight maps of the anisotropy from dusty star-forming galaxies for frequencies <math>\nu &gt; 200</math> GHz, and with <math>1'</math> resolution at 800 GHz.</p> <p>Current knowledge: Three <i>Planck</i> (higher noise) maps between 300 and 900 GHz with <math>5'</math> resolution</p>	Improve constraints on the parameters describing universal star-formation history. Construct a tracer of large-scale structure for CMB de-lensing. Cross-correlate with galaxy surveys and CMB lensing map.

<sup>a</sup> Confusion (not noise) limited

<sup>b</sup> Noise and confusion limited

(Fig. 2.12, right). Currently there are reports of just a few other high- $z$  galaxies that are spatially resolved thanks to gravitational lensing, albeit with less extreme magnifications [155–157]. PICO’s catalog will be transformative as it will probe the spectral energy distribution (SED) of the lensed galaxies at their peaks. Two examples of known sources are shown in the left panel of Figure 2.12. While ground-based instruments observe at frequencies up to  $\log \nu = 11.45$ , PICO’s data will extend to the peak of the SED, up to  $\log \nu = 11.9$  (Fig. 2.12, left).

A straightforward extrapolation of the *Herschel* counts to the 70% non-Galactic sky gives a detection of 4,500 strongly-lensed galaxies with a redshift distribution peaking at  $2 \lesssim z \lesssim 3$  [158], but extending up to  $z > 5$  (Fig. 2.12, left panel). If objects like the  $z = 5.2$  strongly lensed galaxy HLSJ091828.6+514223 exist at higher redshifts, they will be detectable by PICO out to  $z > 10$ . At the 600 GHz detection limit, about 25% of all detected extragalactic sources will be strongly lensed; for comparison, at optical/near-IR and radio wavelengths, where intensive searches have been carried out for many years, the yield is only about 0.1%, that is more than two orders of magnitude lower [159]. To add to the extraordinary sub-mm lensing bonanza, the selection of PICO-detected strongly lensed galaxies will be extremely easy because of their peculiar sub-mm colors (Fig. 2.12, left panel), resulting in a selection efficiency close to 100% [160]. The survey will give the brightest objects over the entire sky, maximizing the efficiency of selecting sources for follow-up observations.

The intensive high spectral and spatial resolution follow-up campaign of this large sample will enable a leap forward in our understanding of the processes driving early galaxy evolution and open up other exciting

prospects, both on the astrophysical and on the cosmological side (see for example Treu [159]).

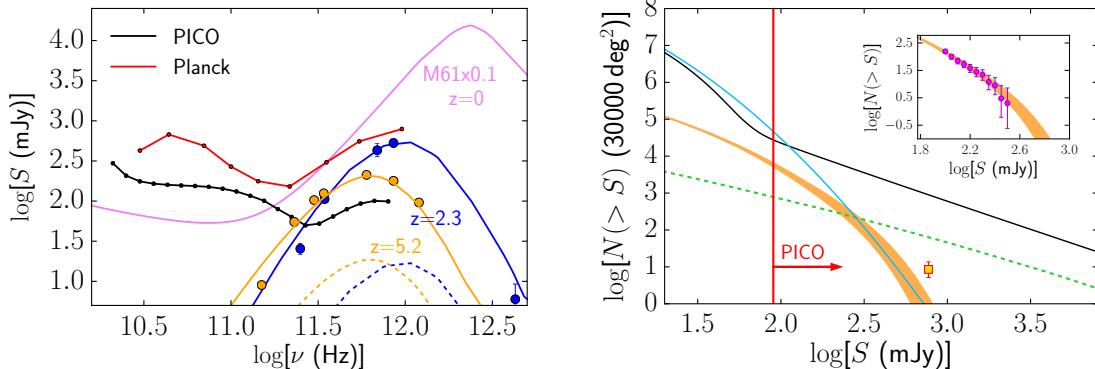


Figure 2.12: **Left:** PICO will detect thousands of new strongly lensed galaxies near the peak of their spectral energy distributions (SEDs), such as SMM J2133-0102 (blue) at  $z = 2.3$  [161] and HLSJ091828.6+514223 (orange) at  $z = 5.2$  [162]. The dashed lines are the SEDs pre-lensing-induced magnification. PICO’s higher resolution gives point-source detection limits (black line) that are up to 10 times fainter compared to *Planck*’s 90% completeness limits (red line [163]). High frequency measurements ( $\nu > 300$  GHz) of 30,000 low- $z$  galaxies, like M61 (magenta, SED was scaled down by a factor of ten.), will give a census of their cold dust. **Right panel.** Integral counts of unlensed (black) and strongly lensed, high- $z$  (orange) star-forming galaxies for 70% of the sky away from the Galactic plane at 600 GHz based on fits of *Herschel* counts over 1000 deg $^2$  (inset [158]). The PICO detection region (right of vertical red line) will yield a factor of 1000 increase in strongly lensed galaxies relative to *Planck* (yellow square), as well as about 50,000 proto-clusters (blue) and 2,000 radio sources (green) [164].

### 2.5.2 Early Phases of Cluster Evolution

PICO will open a new window for the investigation of early phases of cluster evolution, when their member galaxies were actively star forming (and dusty), but the hot IGM was not necessarily in place. In this phase, traditional approaches to cluster detection (X-ray and SZ surveys, and searches for galaxy red sequences) work only for the more evolved clusters, which do include hot IGM; indeed these methods have yielded only a handful of confirmed proto-clusters at  $z \gtrsim 1.5$  [165].<sup>9</sup> *Planck* has demonstrated the power of low-resolution surveys for the study of large-scale structure [166], but its resolution was too poor to detect individual proto-clusters [164]. Studies of the high- $z$  2-point correlation function [164, 167] and *Herschel* images of the few sub-mm bright protoclusters detected so far, at  $z$  of up to 4 [168–170], all of which will be detected by PICO, indicate sizes of  $\simeq 1'$  for the proto-cluster cores, nicely matching the PICO FWHM at the highest frequencies.

PICO will detect 50,000 proto-clusters as peaks in the high frequency maps, which are not available for ground-based instruments (Table 2.3; blue line in the right-hand panel of Fig. 2.12). The redshift distribution will extend out to  $z \sim 4.5$ . This catalog will be augmented by 150,000 evolved clusters, detected by the SZ effect. This will constitute a breakthrough in the observational validation of the formation history of the most massive dark-matter halos, traced by clusters, a crucial test of models for structure formation. Follow-up observations will characterize the properties of member galaxies, probing galaxy evolution in dense environments and shedding light on the complex physical processes driving it.

### 2.5.3 Additional Products of PICO Surveys

PICO will yield a complete census of cold (15 – 25 K) dust, available to sustain star formation in the nearby Universe, by detecting tens of thousands of galaxies mostly at  $z \lesssim 0.1$ ; the SED of M61 is a typical example (Fig. 2.12, left). With a statistical population, and information only available using data at frequencies above 300 GHz, we will investigate the spectral energy distribution of the dust as a function of galaxy properties, such as morphology, and stellar mass.

<sup>9</sup>More high- $z$  proto-clusters have been found by targeting the environment of tracers of very massive halos, such as radio-galaxies, QSOs, sub-mm galaxies. These searches are, however, obviously biased.

PICO will increase by an order of magnitude the number of blazars selected at sub-mm wavelengths and will determine the SEDs of many hundreds of them up to 800 GHz and up to  $z > 5$ . Blazar searches are the most effective way to sample the most massive black holes at high  $z$  because of the Doppler boosting of their flux densities. PICO’s surveys of the largely unexplored mm/sub-mm spectral region will also offer the possibility to discover new transient sources or events, such as blazar outbursts [171].

PICO will make a leap forward in the determination of the polarization properties of both radio sources and of dusty galaxies over a frequency range where ground-based surveys are impractical or impossible. At 320 (800) GHz it will find 1,200 (500) radio sources and 350 (15,000) dusty galaxies at flux limit down to 4 (6) mJy. These data will give information on the structure and ordering of dusty-galaxies’ large-scale magnetic fields. In the case of radio sources emission at higher frequencies comes from regions closer to the central engine, providing information on the innermost regions of the jets, close to the active nucleus.

The anisotropy of the cosmic infrared background (CIB), produced by dusty star-forming galaxies over a wide redshift range  $0 < z \lesssim 5$ , is an excellent probe of the history of star-formation across time. The *Planck* collaboration derived values for parameters describing the rate of star-formation out to  $z \sim 4$  [172–174]). PICO’s lower noise and twice the number of frequency bands will give an order of magnitude improvement on the statistical errors for these parameters [175]. Similar improvement will be achieved in constraining  $M_{\text{eff}}$ , the galaxy halo mass that is most efficient in producing star-formation activity. PICO’s increased sensitivity to Galactic dust polarization will enhance the separation of signals coming from the largely unpolarized CIB and polarized Galactic dust; an effective separation of signals currently limits making reliable, legacy-quality CIB maps. By providing a nearly full-sky map of matter fluctuations traced by dusty star-forming galaxies, such a set of maps could be used for delensing the CMB [176], for measuring local primordial non-Gaussianity from CIB auto-correlations [177], or for cross-correlations with CMB lensing and with galaxy surveys [107].

## 2.6 Signal Separation

### 2.6.1 The Signal Separation Challenge

In the frequency bands that PICO operates there are Galactic intensity emissions due to free-free, synchrotron, and dust, which arise respectively from free electron-proton scattering, free electrons spiraling around Galactic magnetic field lines, and from  $\sim 20$  K elongated interstellar dust grains partially aligned with the local magnetic field. The emission from synchrotron and dust are linearly polarized, and data from *Planck* indicates that their level of  $E$ -mode is approximately twice that of  $B$  [178]. Extra-Galactic sources of emission include the CMB, which has both  $E$  and  $B$ -modes, and point sources of various types whose polarization level and type are not well constrained. The task of ‘signal separation’ is to decompose the detected signal to its constituent sources. Signal separation precision is determined by the requirement to detect or set an upper limit on the inflationary  $B$ -mode, which is the faintest among PICO’s targeted signals. In that context, the terms ‘foreground separation’ and ‘foreground cleaning’ are used as equivalents to ‘signal separation’.

Galactic emission dominates the sky’s polarized intensity on large angular scales  $\ell \lesssim 10$ , it dominates the cosmological  $B$ -modes signals for  $\ell \lesssim 150$  for all allowed levels of  $r$ , and it is expected to be significant even at  $\ell \simeq 1000$ , posing challenge for reconstructing the  $B$ -mode signal from lensing. This is illustrated in Fig. 2.1 and 2.13, which show Galactic emission power spectra calculated based on 60% of the cleanest – that is, the least Galactic-emission-contaminated – portion of the sky. But even in the cleanest, smaller patches of the sky, far from the Galactic plane, Galactic emission levels are known to be substantial relative to an inflationary signal of  $r \sim 0.01$ , and overwhelm it for  $r \lesssim 0.001$  [178]. Separating the cosmological and Galactic emission signals (together with control of systematic uncertainties) are the challenges facing any next-decade experiment attempting to reach these levels of constraints on  $r$ .

A challenge in foreground-separation arises because we do not have sufficiently accurate information about the spectral emissions of the sources anywhere across the sky. To a first approximation, the spectrum

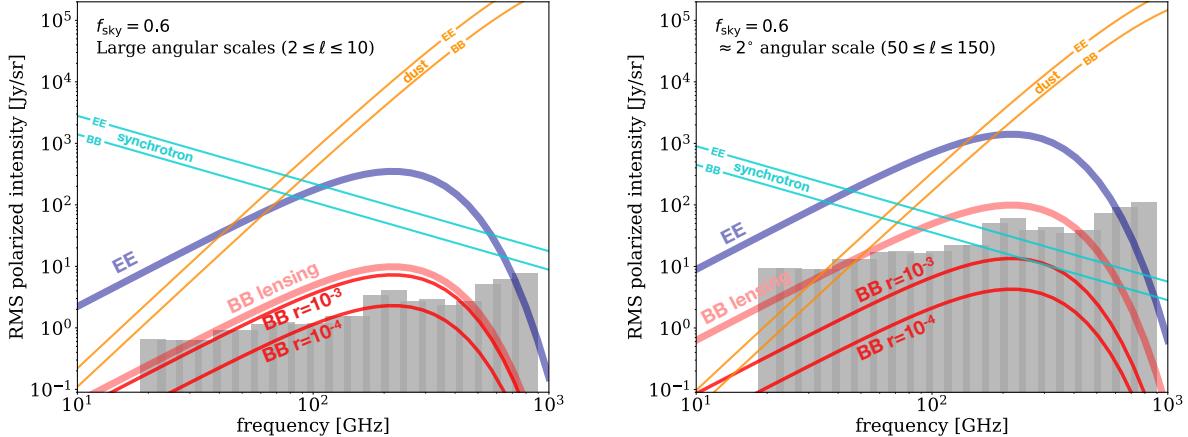


Figure 2.13: Polarization  $BB$  spectra of Galactic synchrotron and dust, compared to CMB polarization  $EE$  and  $BB$  spectra of different origins for two values of  $r$  and for two ranges of angular scales: large  $\ell \leq 10$ , corresponding to the reionization peak (left panel); and intermediate  $50 \leq \ell \leq 150$ , corresponding to the recombination peak (right panel). The locations and sensitivities of the 21 PICO frequency channels are shown as vertical gray bands.

of synchrotron emission is a power law  $I_{\text{sync}} \propto v^\alpha$ , with  $\alpha \simeq -1$ . The spectrum of dust emission is  $I_{\text{dust}} \propto v^\beta B_v(T_{\text{dust}})$ , where  $\beta \simeq 1.6$ ,  $T_{\text{dust}} \simeq 20$  K, and  $B_v(T)$  is the Planck function; this is referred to as ‘modified blackbody emission’. If those models exactly reflected the properties of emitting sources, then in principle an experiment that had 6 frequency bands could determine the three emission parameters, as well as the three amplitudes corresponding to that of dust, synchrotron, and the CMB. However, recent observations have shown that neither emission law is universal, that spectral parameters vary with the region of sky [179–181], and thus that the analytic forms and parameter values given above are valid only as averages across the sky. Also, while both emission laws are well-motivated phenomenological descriptions, the fundamental physics of emissions from grains of different materials, sizes, and temperatures, and of electrons spiraling around magnetic fields implies that these laws are expected to be neither exact, nor universal.

At the low levels of  $r$  targeted by PICO and by other next-decade experiments, even small inaccuracies in foreground modeling and characterization lead to biases and false detections. For example, several publications have demonstrated that fitting complicated dust temperature profiles using a simple one- or two-temperature model will bias the fitted CMB signal at levels  $\delta r \lesssim 10^{-3}$ , which is significant compared to PICO’s goal [182–186].

Further complicating the foreground-separation challenge is the fact that additional polarized foregrounds may exist. ‘Anomalous microwave emission’ (AME) has been observed at mm wavelengths, spatially correlated with thermal dust emission, but with intensity peaking at frequencies near 30 GHz. It has been tentatively identified with small, rapidly-spinning dust grains [181]. Very few measurements of AME polarization exist, and there are only loose constraints on its fractional polarization; it is less than 3% ( $2\sigma$ ) at 18 GHz in one 0.5% region of the sky [187]. Astrophysical emission from CO lines at mm wavelengths is expected to be 0.1 - 1% polarized [188, 189]. Extragalactic radio sources show a median polarization of 2% [190–192], and next to no information is available about the polarization of dusty-galaxies emitting in the PICO wave-bands. Initial quantitative estimates show that ignoring CO lines and radio sources may each lead to a bias  $\delta r \simeq 1 \times 10^{-3}$  [193, 194]. If AME is 1% polarized, left uncorrected it would give rise to a bias of  $\delta r \simeq 5 \times 10^{-4}$  [195]. These levels are appreciable compared to PICO’s and other next-decade experiments’ goals.

## 2.6.2 Foreground Separation Assessment and Methodology

Two broad approaches are used for foreground separation and for assessment of its precision. In the parametric approach the foregrounds are assumed to follow emission laws that are quantified in terms of a number of

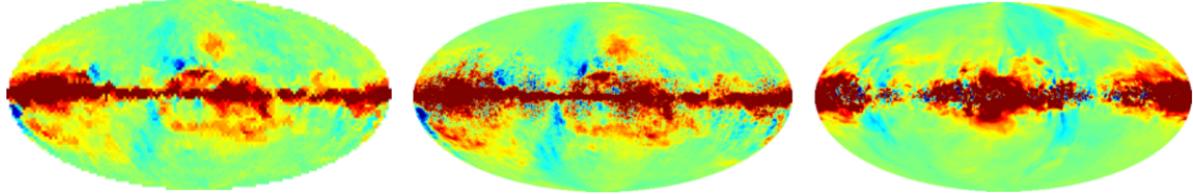


Figure 2.14: Foreground maps: *Planck* real sky (left) at 143 GHz, models at 155 GHz from PySM (middle) [200] and Galactic MHD simulations (right).

parameters. Parametric models use the frequency dependence of the data in each line of sight to determine the values of the parameters [196]. Since the CMB spectrum is well determined, measurements with a sufficient number of frequency bands and appropriately broad frequency coverage, can distinguish foreground emission from the CMB using their different spectral dependences. Non-parametric techniques, in contrast, rely on the fact that CMB emission is uncorrelated with the foregrounds and thus correlations within a given spatial/frequency data-cube can be used to separate the two distinct sources of emission [197, 198]. Simulated data are used to assess the efficacy of both techniques as a function of increasing complexity for the assumed foreground emission. For the parametric models, analytic models have been developed that give the uncertainty on emission parameters as a function of instrument noise, but specific assumptions must be made about the underlying nature of the emission laws [199].

To investigate the efficacy of PICO to address the foreground-separation challenge, we used an analytic forecast and map-domain simulations.

- **Analytic Forecast** The analytic forecast relies on an established, documented, publicly available, cosmological parameters forecasting code [199]. The code uses *Planck*-reported Galactic emissions; it assumes that the foreground spectral indices are constant across patch sizes of  $\sim 15$  deg on a side; it employs a parametric maximum-likelihood approach to remove the foregrounds and to forecast  $\sigma(r)$ ; and it uses the cleanest 60% of the sky.  $B$ -mode due to lensing and a process of delensing are included, but the input for the inflationary signal is  $r = 0$ . Result from the publicly available code, have been verified using a second, entirely independent code that uses similar analytic calculations.

- **Map-Domain Simulations** Map-domain simulations have become the ‘gold standard’ in the community. In this approach, we simulate sky maps that are constrained by available data, but otherwise have a mixture of foreground properties. We ‘observe’ these maps just like a realistic experiment will do, and then apply foreground separation techniques – both parametric and non-parametric – to separate the Galactic and CMB emissions.

what exactly do you want to say here We use eight different sky models. All models are consistent with available data and with uncertainties from WMAP and *Planck*. The range of models include varying degree of complexity, including spectral parameters varying spatially and along the line of sight, anomalous microwave emission up to 2% polarized, dust polarization that rotates slightly as a function of frequency because of projection effects, or dust SEDs that depart from a simple modified blackbody. All foreground maps were generated at native resolution of 6.8 arcmin pixels [201]. They were generated with widely-used and thoroughly-tested map-generation codes [119, 200]. Fig. 2.14 shows three of the eight models, illustrating that distinctly different realizations of the sky are allowed by current data, and highlighting the level of current Galactic emission uncertainties.

For each of the eight models, we added CMB signals in both intensity and polarization, matching a  $\Lambda$ CDM universe. The input inflationary signal was  $r = 0$ , namely, no signal, and the  $BB$ -lensing matched the level after 85% delensing as forecasted for PICO. Each of these sky models had 50 realizations of the PICO noise level. The sky models were analyzed with a variety of techniques, which were based on the two broad categories described above.

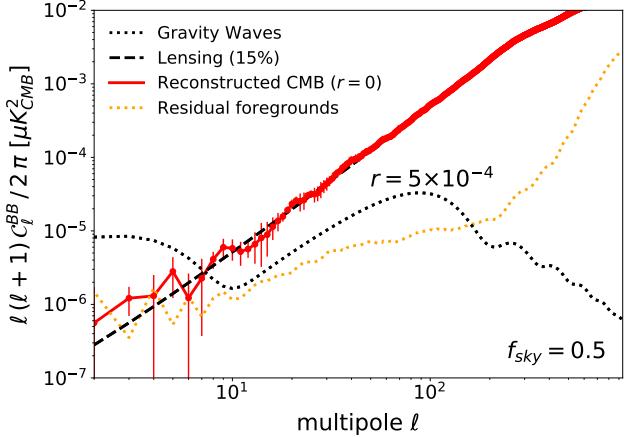


Figure 2.15: Angular power spectra of  $BB$  due to the CMB and of residual foregrounds after an end-to-end map-based foreground-separation exercise. The PICO low noise levels and breadth in frequency coverage enable suppression of foregrounds by a factor of ?? (??) at  $\ell = 4$  (80) relative to their original levels (not shown), such that the residual foreground spectrum (yellow dots) is a factor of ten (four) below a  $BB$  inflationary signal with  $r = 5 \times 10^{-4}$  (black dots). The recovered CMB (red) matches well the input CMB, which consists of only lensing  $B$ -modes, over most angular scales. In this exercise we used 50% of the sky. Lower foreground residual levels are obtainable with smaller, cleaner patches of  $\sim 10\%$  of sky, which would reduce the residual foregrounds at  $\ell \simeq 80$ .

### 2.6.3 Results and Discussion

When using the PICO baseline noise levels with the analytic forecasts we find that  $\sigma(r) = 2 \times 10^{-5}$ , a level that is five times lower than required ( $\sigma(r) = 1 \times 10^{-4}$ , see SO1). We consider this forecast optimistic because it assumes strictly white noise, it assumes a specific model for the underlying foregrounds that has only five parameters per  $15 \times 15 \text{ deg}^2$  pixel, and it is based on analytic calculations of the width of the parameter likelihood functions that are a-priori assumed to be Gaussian. The foregrounds may be more complex, requiring more than five parameters, and may have stronger spatial variations. The parameter likelihoods may not be Gaussian.

The ‘gold-standard’ map-based simulations give initial evidence that the combination of PICO’s sensitivity and broad frequency coverage are effective in foreground removal and that PICO will reach the requirement of  $r = 5 \times 10^{-4}$  ( $5\sigma$ ). Fig. 2.15 shows the result of a foreground-separation exercise over 50% of the sky using GNILC, one of the non-parametric techniques [198]. In this exercise GNILC was tuned to give low foregrounds on the largest angular scales, that is, the lowest  $\ell$  modes. The reconstructed CMB  $BB$  signal, consisting of only lensing  $B$ -modes, is reconstructed within errors for all  $\ell \geq 5$ . The residual foreground  $BB$  power spectrum, encoding the levels of remaining foreground emission after foreground-separation, is a factor of 50 (5) below the CMB at  $\ell = 100$  (10). Most importantly, the residual foreground is a factor of ten below an inflationary  $BB$  signal for  $r = 5 \times 10^{-4}$  at  $\ell \simeq 4$ , demonstrating a suppression of foregrounds by a factor of ?? at these angular scales. These are the angular scales at which the inflationary signal is stronger than the signal from lensing.

At intermediate angular scales  $\ell \simeq 80$  the residual foreground is a factor of four lower than the inflationary signal. We expect lower residuals when GNILC is optimized for this  $\ell$  range. And for reconstructing signals at this  $\ell$  range, it is sufficient to analyze data from smaller  $\sim 3\%$  regions of the sky. These will have lower mean foregrounds levels, making the foregrounds-separation exercise easier, and pushing residuals to levels lower than already demonstrated here for 50% of the sky. With its full sky coverage, PICO will have access to several such independent patches, and will thus make several independent detections of its  $r$  target.

Our results validate the need for a broad frequency coverage with a strong lever arm on Galactic emissions outside of the primary CMB bands. Fig. 2.16 shows that removing several of PICO’s frequency bands, particularly those that monitor dust and synchrotron at high and low frequencies, respectively, significantly biases the extracted  $BB$  power spectrum, especially at the lowest multipoles. In this exercises the input CMB contained the lensing signal *and* an inflationary signal with  $r = 0.001$ .

**now here** While our results are encouraging, as they suggest that PICO’s frequency coverage and sensitivity will be adequate for this level of  $r$ , more work should be invested to gain complete confidence. This work includes: running numerous realizations of different sky models, and analyzing them with various approaches; optimizing sky masks; and using combination of techniques to handle large, intermediate, and small angular scale foregrounds differently.

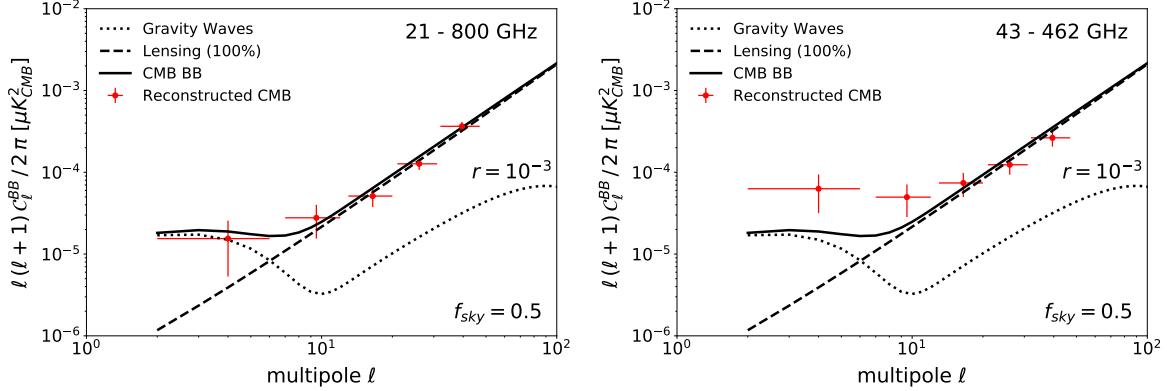


Figure 2.16: Foreground separation with all of PICO’s 21 frequency bands recovers the input CMB *BB* power spectrum (solid black, left and right) without bias (left, red). The input CMB spectrum has contribution from lensing (dash, left and right) and an inflationary signal with  $r = 0.001$  (dots, left and right). This exercise uses a parametric approach [196] with foregrounds varying on  $4^\circ$  pixels, and using 50% sky fraction. Running the same algorithm on the same sky but using only PICO’s bands between 43 and 462 GHz produces an output spectrum (right, red) that is biased at low multipoles relative to the input. With real data, the bias would be erroneously interpreted as higher value of  $r$ .

## 2.7 Systematic Uncertainties

Just as requirements on signal separation (§ 2.6) are determined by the need to reach the faintest inflationary signal so are the requirements on control of systematic uncertainties. Since an inflationary *BB* power spectrum **with  $r = 5 \times 10^{-4}$  has a peak signal level of  $\sim 10$  nK**, systematic effects need to be controlled to a level of few nK at most. It has long been recognized that exquisite control of systematic uncertainties will be required from any experiment attempting to reach levels of  $r \lesssim 1 \times 10^{-3}$ , and it is widely accepted that the stability provided aboard a space platform makes it best suited to control systematic uncertainties compared to other platforms. This is one of the most compelling reasons to observe the CMB from space. As WMAP and *Planck* demonstrated, an L2 orbit offers excellent stability, as well as flexibility in the choice of scan strategy. PICO takes advantage of an L2 orbit, using a rotating spacecraft (at 1 rpm) whose spin axis precesses with a 10 hour period, thus scanning the sky in a way that is crosslinked on many time scales and at many angles, without interference from the Sun, Earth, or Moon; this reduces the effects of low frequency noise without the need for additional internal-to-the-instrument mechanical modulation. The redundancy of observations allows the checking of consistency of results and an improved ability to calibrate and to correct systematic errors in post-processing analysis.

A rich literature investigates the types of systematic errors due to the environment, the instrumentation, observation strategies, and data analysis that could confound polarization measurements by creating a bias or an increased variance [202–204]. Every measurement to date has reached a systematic error limit, **what do you mean?** and many sophisticated techniques have been advanced to mitigate systematics, finding both new technological solutions and new analysis techniques. As an example, BICEP systematics limited it to  $r=0.1$  [205] while through additional effort within the program, BICEP2 achieved a systematics limit of  $r=6\times 10^{-3}$  [206]. In the near term, the ground-based and suborbital CMB community will continue to develop new techniques in handling systematics, particularly in developing the CMB-S4 project.

**what does paragraph try to say?** All prior on-orbit measurements of CMB polarization were limited by systematic errors until an in-depth study of the systematics was performed allowing a post-processing data analysis to suppress them [28, 207, 208]. Particularly we note figure 3 of Ref.[28] which quantifies *Planck*’s systematic error limits on the polarization power spectral measurements. Recently studied space missions, such as EPIC-IM, LiteBird and *CORE*, have placed systematic error mitigation at the forefront of the case for their mission and have developed tools and strategies for estimating and mitigating these [209–211].

Systematic effects are coupled with the spacecraft scan strategy, and the details of the data analysis pipeline. Thus, end-to-end simulation of the experiment is an essential tool, including realistic instabilities and non-idealities of the spacecraft, telescope, and instrument, as well as folding in data post-processing techniques used to mitigate the systematic effects.

### 2.7.1 List of Systematics

The systematic errors faced by PICO can be categorized into three broad categories: 1) Coupling of the intensity into the polarization signal; 2) stability; and 3) straylight. These are listed in Table 2.4 and were prioritized for further study using a risk factor incorporating the working group’s assessment of how mission-limiting the effect is, how well these effects are understood by the community and whether mitigation techniques exist.

The three highest risk systematic errors were studied further and are discussed in subsections below. The PICO team used simulation and analysis tools developed for *Planck* [212] and *CORE*, adapting them for PICO.

Name	Risk	Effect	
<b>Leakage</b>			
Polarization angle calibration	5	$E \rightarrow B$	Sect. 2.7.2.
Bandpass mismatch .....	4	$T \rightarrow P, E \rightarrow B$	
Beam mismatch .....	4	$T \rightarrow P, E \rightarrow B$	Sect. 2.7.2
Time response accuracy and stability .....	4	$T \rightarrow P, E \rightarrow B$	
Readout cross-talk .....	4	spurious P	
Chromatic beam shape .....	4	spurious P	
Gain mismatch .....	3	$T \rightarrow P$	
Cross-polarization .....	3	$E \rightarrow B$	
<b>Stability</b>			
Gain stability .....	5	$T \rightarrow P, E \rightarrow B$	Sect. 2.7.3
Pointing jitter .....	3	$T \rightarrow P, E \rightarrow B$	
<b>Straylight</b>			
Far sidelobes .....	5	spurious P	Sect. 2.7.4.
<b>Other</b>			
Residual correlated noise (1/f, cosmic ray hits) .....	3	increased variance	

Table 2.4: Enumeration of potential systematic errors anticipated in PICO’s measurements together with their assessed risk level, their effect on the measurements, and, in few cases, subsections with further discussion. The symbol  $T \rightarrow P$  denotes coupling of the intensity (=temperature) signal into polarization; similar meaning holds for the  $E$  and  $B$  modes. Risk level rises with value. Level 5 indicates a highly significant, design-driving effect; it may have limited past measurements, and/or isn’t well understood. Level 4 indicates an effect that is either known to be large but is understood reasonably well, or a smaller effect that requires precise modeling. Level 3 indicates that we expect the effect to be small, but it is not sufficiently well understood and detailed modeling will be done during Phase A study. We used simulations to investigate effects with risk level 5.

### 2.7.2 Absolute polarization angle calibration

The measured CMB polarization can be rotated due to (1) a birefringent primordial Universe, or a Faraday rotation due a primordial magnetic field [213]; (2) birefringent foregrounds, or interaction with the Galactic magnetic field; (3) systematic effects in the instrument, and in particular an error in the direction of polarization measured by each detector. While the first two sources create a rotation that may depend on scale, position and/or frequency, the latter depends mainly on the detector itself.

A rotation  $\alpha$  of the direction of polarization mixes the  $Q$  and  $U$  Stokes parameters via  $Q \pm iU \rightarrow e^{\mp i2\alpha} (Q \pm iU)$  and thus mixes the power spectra and their correlations, as illustrated in Fig. 2.17.

The most recent constraints on cosmological birefringence [92] were limited by uncertainties on the detector orientations. In *Planck*, the detectors were characterized pre-launch to  $\pm 0.9^\circ$  (rel.)  $\pm 0.3^\circ$  (abs.) [214]. For PICO, the relative rotation of the detectors will be measured to a few  $0.1'$  using the CMB, but the overall rotation is unlikely to be known pre-launch to better than *Planck*. Known polarized sources, such as the Crab Nebula, are not characterized well enough independently to serve as calibrators; Aumont et al. [215] show that the current uncertainty of  $0.33^\circ$  on the Crab polarization orientation, limits a  $B$ -mode measurement to  $r \sim 0.01$ , far from PICO’s target.

In the absence of other systematics and foregrounds, a polarization rotation error  $\alpha$  of  $10'$  degrades the uncertainty of  $r$  by 30%, while  $EB$ ,  $TB$ , and  $BB$  spectra can measure a rotation  $\alpha$  at  $3\sigma$  when  $\alpha = 0.07, 0.2$  and  $0.9'$ , respectively, on perfectly delensed maps, and  $0.25, 0.9$ , and  $4.5'$  on raw maps.

In principle, the technique of using the  $TB$  and  $EB$  spectra can measure a global polarization rotation

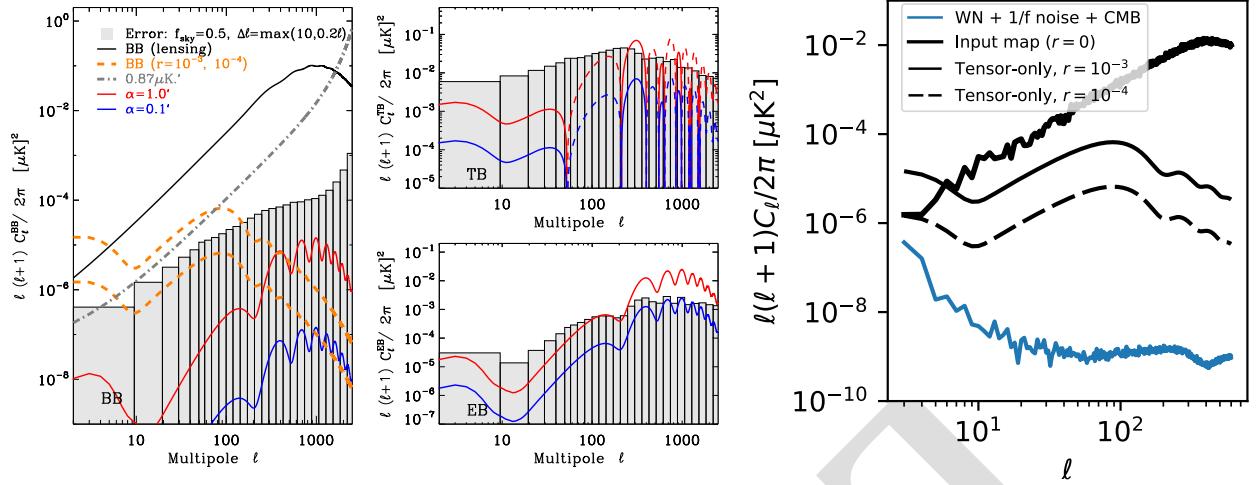


Figure 2.17: Left and center: Spurious power due to a rotation of the angle of polarization, assuming the *Planck* 2018  $\Lambda\text{CDM}$  model [63] with  $\tau = 0.054$  and expected PICO noise performance, assuming removal of 85% of the lensing power. Right: residual power after removal of temporal gain drifts using the dipole.

error at levels (around  $0.1'$ ) below those affecting  $r$  measurements in *BB* ( $> 1'$ ). However, a future mission should simulate additional aspects of this issue, such as delensing, the interaction with foregrounds, and  $1/f$  noise in simulating and assessing the impact of an angle-calibration error.

### 2.7.3 Gain Stability

Photometric calibration is the process of converting the raw output of each detector to physical units via the characterization of the ‘gain factor’  $G(t)$ , which is a function of time. Statistical or systematic errors in the determination of  $G$  translate to increased uncertainties, or biased results, respectively. We investigated whether the anticipated error on  $G$  is adequate for PICO’s requirements on measuring the faint inflationary signal.

The inflationary signal will be extracted from data in the primary CMB bands between 60 and 300 GHz. Detectors in these bands will be calibrated using measurements of the CMB dipole, a signal that will be measured once per minute as the telescope scans the sky (§ ??). We evaluated the combined impact of the scan strategy and white- and  $1/f$ -noise in the estimation of  $G(t)$ .<sup>10</sup> We simulated the time stream of four detectors over a year of observations, including PICO baseline noise levels and  $1/f$  noise parameters (§ ??) where?. We fit the an end-to-end simulation using now here

instrument and the CORE mission proposal. The quality of the estimate depends on the noise level of the receivers, but also on the details of the scanning strategy. To analyze the impact of calibration uncertainties on PICO, we performed the following analysis. Firstly, we simulated the observation of the sky, assuming four receivers, the nominal scanning strategy, and  $1/f$  noise. The simulated sky contained CMB anisotropies, plus the CMB dipole. Then we ran the calibration code to fit the dipole against the raw data simulated during the first step. We again simulated the observation of the sky, this time using the values of  $G$  computed during the second step, which contain errors due to the presence of noise and the CMB signal.

The presence of large-scale Galactic emission features can bias the estimation of calibration factors. Ideally, a full data-analysis pipeline would pair the calibration step with the component-separation step, following a schema similar to *Planck/LFI*’s final data processing [216]: the calibration code is followed by a component-separation analysis, and these two steps are iterated until the solution converges.

Results of the simulation (neglecting foregrounds) are shown as power spectrum residuals in Fig. 2.17. We estimate the gain fluctuations to better than  $10^{-4}$  when solving for the gain every 40 hours (4 precession

<sup>10</sup>We use software tools developed for and validated in the context of analysis of the *Planck/LFI* data.

periods). The scanning strategy employed by PICO allows for a much better calibration than was possible for *Planck*, thanks to the much faster precession.

#### 2.7.4 Far-Sidelobe Pickup

Measurement of each detector's response to signals off axis, which tends to be weak (~80dB less than the peak response) but spread over a very large solid angle, is difficult to do pre-launch, and may not even be done accurately after launch. Nonetheless, this far sidelobe can couple bright Galactic signals from many tens of degrees off-axis and confuse it with polarized signal from the CMB off the Galactic plane. To evaluate this systematic error, GRASP software<sup>11</sup> was used to compute the PICO telescope's response over the full sky. The computed full-sky beams showed features peaking at about -80 dB of the on-axis beam. This full-sky beam was convolved with a polarized Galactic signal and a one-year PICO mission scan using the simulation pipeline; this preliminarily shows that the far sidelobe pickup must be reduced by 90% to keep the sidelobe signal from appreciably increasing the variance on the B-mode power measurement. This reduction can be achieved by adjusting the instrument design or by modeling, measuring, and removing the sidelobe pickup during data analysis.

#### 2.7.5 Key Findings

Properly modeling, engineering for, and controlling the effects of systematic errors in a next-generation CMB probe is critical. As of today, we conclude that there is a clear path to demonstrate that state-of-the-art technology and data processing can take advantage of the L2 environment and control systematic errors to a level that enables the science goals of PICO. In particular we note the following points:

- The raw sensitivity of the instrument should include enough margin that data subsets can independently achieve the science goals. This allows testing of the results in the data analysis and additional data cuts, if needed.
- For PICO mission, a physical optics model of the telescope should be developed, enabling full-sky beam calculations, which should be validated as much as possible on the ground. This will be needed to characterize and remove far-sidelobe pickup seen during the mission.
- NASA's support of ground-based and suborbital CMB missions will mitigate risk to a future space mission such as PICO by continuing to develop analysis techniques and technology for the reduction of systematic errors.
- In the PICO mission's phase A, a complete end-to-end system-level simulation software facility should be developed to assist the team in setting requirements and conducting trade-offs between subsystem requirements while realistically accounting for post-processing mitigation. Any future CMB mission is likely to have similar orbit and scan characteristics to those of PICO, thus there is an opportunity for NASA and the CMB community to invest in further development of this capability now.
- Low frequency excess noise (also called 1/f noise) should be studied in detail as part of a simulation effort to set detailed detector and systems level requirements. The systematics simulations performed here show that PICO's science goals can be achieved with no additional modulation and assuming current state-of-the-art levels of low frequency noise (a total knee frequency of 20 mHz) based on demonstrated TES performance, and system-level residuals achieved by *Planck*.

### 2.8 Complementarity with Sub-Orbital Measurements

Since the first CMB measurements, more than 50 years ago, important observations have been made from the ground, from balloons, and from space. Each of the CMB satellites flown to date - COBE, WMAP, and *Planck*- has relied crucially on technologies and techniques that were first proved on ground and balloon flights, making these also crucial to the success of PICO. The phenomenal success of, and the immense science outcomes produced by, past space missions is a direct consequence of their relative advantages, as

<sup>11</sup><https://www.ticra.com>

Table 2.5: Relative characteristics of ground, balloon, and space platforms for experiments in the CMB bands.

Characteristic	Ground	Balloon	Space
Sky coverage . . . . .	Partial from single site	Partial from single flight	Full
Frequency coverage . . . . .	70 GHz inaccessible, <sup>a</sup> $v \geq 300$ GHz unusable, limited atmospheric windows	70 GHz inaccessible, <sup>a</sup> otherwise, almost unlimited	Unrestricted
Angular resolution at 150 GHz <sup>b</sup> . .	1'5 with 6 m telescope	6' with 1.5 m telescope	6' with 1.5 m telescope
Detector noise <sup>c</sup> . . . . .	$265 \mu\text{K}_{\text{CMB}} \sqrt{\text{s}}$	$162 \mu\text{K}_{\text{CMB}} \sqrt{\text{s}}$	$36 \mu\text{K}_{\text{CMB}} \sqrt{\text{s}}$
Integration time . . . . .	Unlimited, with interruptions	Weeks, continuous	Several years, continuous
Repairability, Upgradeability . . . .	Good	None; multiple flights possible	None

<sup>a</sup> 70 GHz is the frequency at which large angular scale  $B$ -mode Galactic emissions have a minimum (Fig. 2.13).

<sup>b</sup> We give representative telescope apertures. Significantly larger apertures for balloons and in space result in higher mass, volume, and cost.

<sup>c</sup> The noise-equivalent temperatures given are illustrative of general capabilities. Detailed comparisons depend on detector heat sink temperatures, bandwidths, and other factors that differ among specific implementations. Ground – median detector noise at 95 GHz from BICEP3 [217]; balloon – median detector noise at 94 GHz from SPIDER [?]; space – 90 GHz from PICO CBE use req.

listed in Table 2.5. In every respect, with the exception of repairability and upgradeability, space has the advantage.

When the entire sky is needed, as for fluctuations on the largest angular scales, space is by far the most suitable platform. When broad frequency coverage is needed, space will be required to reach the ultimate limits set by astronomical foregrounds. As Figures 2.1 and 2.13 demonstrate, Galactic emissions overwhelm the inflationary signal on large and intermediate angular scales ( $\ell \leq 150$ ), and they are significant even at high  $\ell$ , potentially limiting the process of delensing that is necessary for reaching levels of  $r \lesssim 0.001$ . The stability offered in space can not be matched on any other platform, and it translates to superb control of systematic uncertainties. There is a broad consensus within the CMB community that for levels of  $r \lesssim 0.001$  the challenges in the measurement are the ability to control systematic uncertainties and to remove Galactic emissions; modern focal plane arrays, like the one employed by PICO have ample raw sensitivity. The PICO goal of reaching  $\sigma(r) = 5 \times 10^{-4}$  ( $5\sigma$ ) is beyond the reach of sub-orbital observations in the foreseeable future.

For science requiring higher angular resolution, such as observations of galaxy clusters with  $\sim 1$  arcmin resolution at 150 GHz, the ground has an advantage. An appropriately large aperture on the ground will also provide high resolution information at lower frequencies, which may be important for separating Galactic emissions at high  $\ell$ .

The relative advantages of a space mission used to come with higher costs relative to sub-orbital experiments. However, with the advent of massive ground-based efforts this balance shifts; the costs for a next generation ground-based CMB experiment planned for the next decade are squarely within the cost window of this Probe. A recommended plan for the next decade is therefore to pursue a space mission, and complement it with an aggressive ground program that will overlap in  $\ell$  space, and will add science at the highest angular resolution, beyond the reach of a space mission.

Balloon observations have been exceedingly valuable in the past. They co-led discoveries of the temperature anisotropy and polarization, provided proving grounds for the technologies enabling the success of COBE, WMAP and *Planck*, and trained the scientists that then led NASA’s space missions. There are specific areas for which balloon missions can continue to play an important role, despite their inherently limited observing time. Balloon payload can access frequency bands above 280 GHz; currently there are no plans for any ground program to conduct observations at higher frequencies. These frequency bands will provide important, and perhaps critical information about polarized emission by Galactic dust, a foreground

that is currently known to limit knowledge of the CMB signals. With flights above 99% of the atmosphere, balloon-borne observations are free from the noise induced by atmospheric turbulence, making them good platforms for observations of the low  $\ell$  multipoles, and for characterizing foregrounds on these very large angular scales. From a technology point of view, the near-space environment is the best available for elevating detector technologies to TRL6; and balloon-platforms continue to be an excellent arena for training the scientists of tomorrow.

[say something about other complementarities?](#)

## 2.9 Measurement Requirements

The set of physical parameters and observables that derive from the PICO science objectives place requirements on the depth of the mission, the fraction of sky the instrument scans, the frequency range the instrument probes and the number of frequency bands, the angular resolution provided by the reflectors, and the specific pattern with which PICO will observe the sky. We discuss each of these aspects in turn now.

- **Depth** We quantify survey depth in terms of the RMS fluctuations that would give a signal-to-noise ratio of 1 in a sky pixel that is  $1'$  on a side. Depth in any frequency band is determined by detector sensitivity, the number of detectors in the focal plane, the sky area covered, and the duration of the mission. The science objective driving the depth requirement is SO1, the search for the IGW signal, which requires a combined depth of  $0.87 \mu\text{K} \cdot \text{arcmin}$ . This requirement is a combination of the low-level of the signal, the need to separate the various signals detected in each band, and the need to detect and subtract systematic effects to anticipated levels. The CBE value is  $0.61 \mu\text{K} \cdot \text{arcmin}$  coming from a realistic estimate of detector noise, and giving 40% margin on mission performance.

- **Sky Coverage** There are several science goals driving a full-sky survey for PICO. The term ‘full-sky’ here refers to the entire area of sky available after separating other astrophysical sources of confusion. In practice this implies an area of 50-70% of the full sky for probing non-Galactic signals, and the rest of the sky for achieving the Galactic science goals. There are five main goals, each driving different mission requirements.

- (1) Probing the optical depth to the epoch of reionization (STM SO5) requires full sky coverage as the signal peaks in the  $EE$  power spectrum on angular scales of  $20^\circ$  to  $90^\circ$ . Measuring this optical depth to limits imposed by the statistics of the small number of available  $\ell$  modes is crucial for minimizing the error on the neutrino-mass measurement.

- (2) If  $r \neq 0$ , the  $BB$  power spectrum due to IGW (STM SO1) has local maxima on large angular scales ( $20^\circ$  to  $90^\circ$ ,  $2 \leq \ell \leq 10$ ), and around  $1^\circ$  ( $\ell \simeq 80$ ). A detection would strongly benefit from confirmation at *both* angular scales – a goal that is beyond the capabilities of ground-based instruments – and, for the  $\ell = 80$  peak, *in several independent patches of the sky*, a goal PICO will achieve, but that is currently not planned for any next-decade instrument.

- (3) The PICO constraint on  $N_{\text{eff}}$  (STM SO4) requires a determination of the  $EE$  power spectrum to limits imposed by the statics of available  $\ell$  modes. Full sky coverage is required to achieve this limit.

- (4) Achieving the targeted neutrino mass limits (STM SO3), giving two independent  $4\sigma$  constraints on the minimal sum of 58 meV, requires a lensing map, and cluster counts from as large a sky fraction as possible.

- (5) PICO’s survey of the Galactic plane and regions outside of it is essential to achieving its Galactic structure and star-formation science goals (SO6, 7).

- **Frequency Bands** The multitude of astrophysical signals that PICO will characterize determine the frequency range and number of bands that the mission uses. The IGW signal peaks in the frequency range between 30 and 300 GHz. However, Galactic signals, which are themselves signals PICO strives to characterize, are a source of confusion for the IGW. The Galactic signals and the IGW are separable using their spectral signatures. Simulations indicate that 21 bands, each with 25% bandwidth, that are spread across the range of 20 - 800 GHz can achieve the separation at the level of fidelity required by PICO.

Additionally, characterization of the Galactic signals, specifically the make-up of Galactic dust (SO6),

requires spectral characterization of Galactic dust in frequencies between 100 and 800 GHz.

• **Resolution** Several science objectives require an aperture of 1.5 m and the resolution per frequency listed in Table 2.2. To reach  $\sigma(r) = 1 \times 10^{-4}$  we will need to ‘delens’ the *E*- and *B*-mode maps, as described in Sections 2.2 and 2.3. Delensing is achieved with a map that has a native resolution of 2-3 arcminutes at frequencies between 100 and 300 GHz. Similar resolution is required to meet the constraints on the number of light relics (SO4), which will be extracted from the *EE* power spectrum at multipoles  $100 \lesssim \ell \lesssim 2500$ . The process of delensing may be affected by other signals, primarily that due to Galactic dust. It is thus required to map Galactic dust to at least the same resolution as at 300 GHz. Higher resolution is mandated by SO6 and 7, which require resolution of 1' at 800 GHz. We have thus chosen to implement diffraction-limited resolution between 20 and 800 GHz.

• **Sky Scan Pattern** Control of polarization systematics uncertainties at anticipated levels is enabled by: (1) making *I*, *Q*, and *U* Stokes-parameter maps of the entire sky from each independent detector; and (2) by enabling sub-percent absolute gain calibration of the detectors through observations of the CMB dipole. With these requirements we chose a sky scan pattern that enables each detector to scan a given pixel of the sky in a multitude of directions, satisfying requirement (1). The scan we chose also gives strong CMB dipole signals in every rotation of the spacecraft throughout the lifetime of the mission, satisfying requirement (2).

### 3 Instrument

PICO meets all of its science-derived instrument requirements (Table 2.2) with a single instrument: an imaging polarimeter with 21 logarithmically spaced frequency bands centered between 21 and 799 GHz. The instrument is built around a two-reflector Dragone-style telescope (§ 3.1 and Fig. 3.1) with an internal aperture stop between the primary and secondary. The focal plane is populated by 12,996 TES bolometers (§ 3.2) read out using a time-domain multiplexing scheme (§ 3.3). The instrument employs a single science observing mode: fixed rate imaging while scanning the sky. [add section?](#)

The instrument is configured inside the shadow of a V-grooves assembly that thermally and optically shields it from the Sun (Fig. 3.1). The V-groove assembly consists of 4 nested radiation shields that provide passive cooling (§ 3.4.3). The sun shadow cone depicted in Fig. 3.1 is 29°. The angle to the Sun during the survey,  $\alpha = 26^\circ$  (§ 4.1.2 and Fig. 4.2), is supplemented with a margin of 3° to account for the radius of the sun (0°25), pointing control error, design margin, and alignment tolerances.

The V-groove assembly is attached to the bipod struts that support the instrument structural ring. The ring supports the primary reflector and telescope box. The telescope box contains the actively cooled components (§ 3.4.1, § 3.4.2), including the secondary reflector, the focal plane and sub-kelvin adiabatic refrigerator structures. Just inside the box, a thermal liner serves as a cold optical baffle and aperture stop. Instrument integration and test (I&T) are described in § 3.5.

During the survey, the instrument is spun at 1 rpm (§ 4.1.2). Spacecraft control is simplified by mounting the instrument on a spinning spacecraft module, while a larger non-spinning module houses most spacecraft subsystems (§ 4.3). Instrument elements that act as heat sources are accommodated on the spinning module of the spacecraft.

#### 3.1 Telescope

The PICO telescope design is driven by a combination of science requirements and physical volume limits. The science requirements are: a large diffraction-limited field of view (DLFOV) sufficient to support  $\sim 10^4$  detectors, arcminute resolution at 800 GHz, low spurious polarization, and low sidelobe response. All requirements are met with PICO’s 1.4 m aperture modified open-Dragone design. There are no moving parts in the PICO optical system.

The PICO optical design was selected following a trade study examining cross-Dragone, Gregorian Dragone, and open-Dragone designs [218]. The open-Dragone and crossed-Dragone offer more diffraction-limited focal plane area than the Gregorian Dragone [219], and are able to support enough detectors to

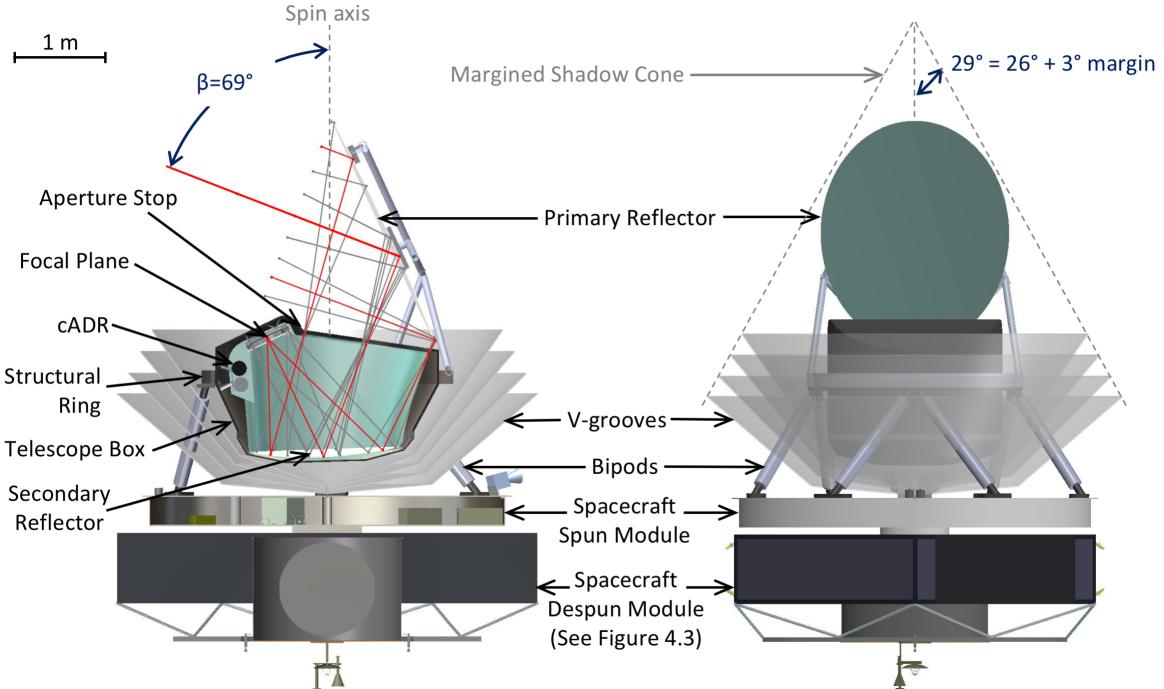


Figure 3.1: Detailed PICO instrument configuration. *Left:* Side view in cross section. *Right:* Front view with V-Groove assembly shown semi-transparent. The spacecraft spun module accommodates warm instrument components: the 4 K cooler compressor and drive electronics, the sub-K cooler drive electronics, and the detector warm readout electronics.

provide the required sensitivity. The open-Dragone does not require the massive and voluminous baffles that the cross-Dragone does, and hence can satisfy the aperture size requirement within the shadow cone.

PICO’s initial open-Dragone design has been modified by adding an aperture stop and adding corrections to the primary and secondary reflectors to enlarge the DLFOV. The detailed geometric parameterization of the PICO optical design is described in [218]. The primary reflector ( $270\text{ cm} \times 205\text{ cm}$ ) is passively cooled and the secondary reflector ( $160\text{ cm} \times 158\text{ cm}$ ) is actively cooled. The highest frequency (900 GHz) sets the surface accuracy requirement of the reflectors at  $\sim \lambda/14 = 24\text{ }\mu\text{m}$ . The focal ratio is 1.42. The slightly concave focal surface, which has a radius of curvature of 4.55 m, is telecentric to within  $0.12^\circ$  across the entire FOV.

**what about the reflector material?** An actively cooled circular aperture stop between the primary and secondary reflectors reduces detector noise and shields the focal plane from stray radiation. Stray-light analysis of the PICO open-Dragone design using GRASP confirms that the focal plane is protected from direct view of the sky, and that spillover past the primary is suppressed by 80 dB relative to the main lobe for both co-pol and cross-pol beams. Detailed baffle design will be performed during mission formulation.

### 3.2 Focal plane

PICO’s focal plane is populated by an array of TES bolometers operating in 21 frequency bands, each with 25% fractional bandwidth, and band centers ranging from 21 to 799 GHz. Polarimetry is achieved by differencing the signals from pairs of two co-pointed bolometers that are sensitive to two orthogonal polarization states. **polarimetry is incomplete; need to elaborate more** A conceptual layout of the PICO focal plane is shown in Fig. 3.2 and detailed in Table 3.1.

Bolometers operating in the mm/sub-mm wave band are photon-noise limited. Therefore, increase in sensitivity is achieved through an increase in detector count. The PICO focal plane has 12,996 detectors,

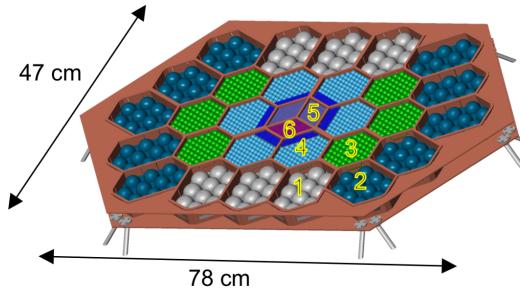


Figure 3.2: PICO focal plane. Detectors are fabricated on six types of tiles (shown numbered and colored as in Table 3.1). The wafers are located on the focal plane such that higher frequency bands, which require better optical performance, are placed nearer to the center.

Table 3.1: PICO makes efficient use of the focal area with multichroic pixels (three bands per pixel, § 3.2.1). The sampling rate is based on the smallest beam (Table 3.2), with 3 samples per FWHM at a scan speed ( $360^\circ/\text{min}$ )  $\sin(\beta = 69^\circ) = 336^\circ/\text{min}$ . Scaling from suborbital experience, we anticipate that TES bolometers can support these sampling rates with a factor of  $\sim 4 \times$  margin.

Tile type	$N_{\text{tile}}$	Pixels/tile	Pixel type	Bandcenters [GHz]	Sampling rate [Hz]
1	6	10	A	21, 30, 43	45
2	10	10	B	25, 36, 52	55
3	6	61	C	62, 90, 129	136
4	6	85	D	75, 108, 155	163
		80	E	186, 268, 385	403
5	2	450	F	223, 321, 462	480
6	1	220	G	555	917
		200	H	666	
		180	I	799	

175 times the number flown aboard *Planck* thereby providing a breakthrough increase in sensitivity with a comparably sized telescope. This breakthrough is enabled by development and demonstration in suborbital projects, which now commonly operate arrays of  $10^3$ – $10^4$  detectors (§ 5).

### 3.2.1 21–462 GHz bands

Several optical coupling technologies have matured over the past ten years to efficiently use focal plane area: horns with ortho-mode transducers (OMTs) [220], lithographed antenna arrays [221], and sinuous antennas under lenslets [222]. Horn-coupling and sinuous antenna/lenslet-coupling deliver quantum efficiency  $> 70\%$  over more than an octave of bandwidth, which can be partitioned into two or three colors per pixel. Alternatively, antenna-array coupling has only demonstrated single-color pixels to date, but uniformly illuminated feeds enable much smaller pixels and therefore more densely packed focal planes.

The PICO baseline focal-plane design employs three-color sinuous antenna/lenslet-coupling [223] for the 21–462 GHz bands. Niobium microstrips mediate the signals between the antenna and detectors, and partition the feed’s wide continuous bandwidth into three narrow channels using integrated micro-machined filter circuits [224]. The technology maturation required for PICO is described in § 5.1.

### 3.2.2 555–799 GHz bands

PICO’s highest three frequency channels are beyond the niobium superconducting band-gap, rendering microstrip filters a poor solution for defining the optical passband. In this regime, PICO instead measures a single band with each pixel using feedhorn-coupled polarization-sensitive bolometers. Radiation is coupled through horns directly to absorber in the throat of a waveguide. TES bolometers detect the incident power.

Table 3.2: PICO has 21 partially overlapping frequency bands with band centers ( $v_c$ ) from 21 GHz to 799 GHz and each with bandwidth  $\Delta v/v_c = 25\%$ . The beams are single mode, with FWHM sizes of  $6.2 \times (155 \text{ GHz}/v_c)$ . The CBE per-bolometer sensitivity is photon-noise limited (§ 3.2.3). The total number of bolometers for each band is equal to (number of tiles)  $\times$  (pixels per tile)  $\times$  (2 polarizations per pixel), from Table 3.1. Array sensitivity assumes 90% detector operability. The map depth assumes 5 yr of full sky survey at 95% survey efficiency, except the 25 and 30 GHz frequency bands, which are conservatively excluded during 4 hr/day Ka-band (26 GHz) telecom periods (§ 4.2).

Band Center [GHz]	Beam FWHM [arcmin]	CBE Bolo NET [ $\mu\text{K}_{\text{CMB}} \text{s}^{1/2}$ ]	$N_{\text{bolo}}$	CBE Array NET [ $\mu\text{K}_{\text{CMB}} \text{s}^{1/2}$ ]	Baseline Array NET [ $\mu\text{K}_{\text{CMB}} \text{s}^{1/2}$ ]	Baseline polarization map depth [ $\mu\text{K}_{\text{CMB}} \text{ arcmin}$ ] [ $\text{Jy sr}^{-1}$ ]
21 . . . . .	38.4	112.2	120	12.0	17.0	23.9 8.3
25 . . . . .	32.0	103.0	200	8.4	11.9	18.4 10.9
30 . . . . .	28.3	59.4	120	5.7	8.0	12.4 11.8
36 . . . . .	23.6	54.4	200	4.0	5.7	7.9 12.9
43 . . . . .	22.2	41.7	120	4.0	5.6	7.9 19.5
52 . . . . .	18.4	38.4	200	2.8	4.0	5.7 23.8
62 . . . . .	12.8	69.2	732	2.7	3.8	5.4 45.4
75 . . . . .	10.7	65.4	1020	2.1	3.0	4.2 58.3
90 . . . . .	9.5	37.7	732	1.4	2.0	2.8 59.3
108 . . . . .	7.9	36.2	1020	1.1	1.6	2.3 77.3
129 . . . . .	7.4	27.8	732	1.1	1.5	2.1 96.0
155 . . . . .	6.2	27.5	1020	0.9	1.3	1.8 119.1
186 . . . . .	4.3	70.8	960	2.0	2.8	4.0 433.1
223 . . . . .	3.6	84.2	900	2.3	3.3	4.5 604.2
268 . . . . .	3.2	54.8	960	1.5	2.2	3.1 433.4
321 . . . . .	2.6	77.6	900	2.1	3.0	4.2 577.8
385 . . . . .	2.5	69.1	960	2.3	3.2	4.5 429.1
462 . . . . .	2.1	132.6	900	4.5	6.4	9.1 551.1
555 . . . . .	1.5	657.8	440	23.0	32.5	45.8 1580
666 . . . . .	1.3	2212	400	89.0	125.8	177.2 2075
799 . . . . .	1.1	10430	360	525.9	743.8	1047 2884
Total . . . . .		12 996		0.43	0.61	0.87

The waveguide cut-off defines the lower edge of the band, and quasi-optical metal-mesh filters define the upper edge. Numerous experiments have successfully used similar approaches [225–227]. The technology maturation required for PICO is described in § 5.2.

### 3.2.3 Sensitivity

PICO’s Current Best Estimate (CBE) sensitivity meets the requirements of the baseline mission with  $> 40\%$  margin (Table 3.2). [this is an odd statement, need to rework](#)

We developed an end-to-end noise model of the PICO instrument to predict mission sensitivity and provide a metric by which to evaluate mission design trades. The model includes four noise sources per bolometer: photon, phonon, Johnson, and readout (from both cold and warm readout electronics). To validate our calculations, we compared two independent software packages that have been validated with several operating CMB instruments. The calculations agreed within 1% both for individual noise terms and for overall mission noise. A detailed description of the PICO noise model and its inputs is available in Young et al. [218]; small differences between that publication and Table 3.2 are due to refinements of the primary mirror and stop temperatures.

Laboratory experiments have demonstrated that TES bolometers can be made background-limited in the low loading environment they would experience at L2 [228]. For PICO, the primary contributor to noise is the optical load. The sources of optical load are the CMB, reflectors, aperture stop, and low-pass filters. The CMB and stop account for at least 50% of the optical load at all frequencies up to and including 555 GHz. At higher bands emission from the primary mirror dominates.

The sensitivity model assumes white noise at all frequencies. Sub-orbital submillimeter experiments

have demonstrated TES detectors that are stable to at least as low as 20 mHz [229], meeting the requirements for PICO’s scan strategy (§ 4.1.2). are we saying anything about 1/f? are we saying anything about yield? Is the observing time consistent with Young et al?

### 3.3 Detector readout

Suborbital experiment teams over the past ten years have chosen to use voltage-biased TESs because their current readout scheme lends itself to Superconducting Quantum Interface Device (SQUID) based multiplexing. Multiplexing reduces the number of wires to the cryogenic stages and thus the total thermal load that the cryocoolers must dissipate. This approach also simplifies the instrument design.

In the multiplexing circuitry, SQUIDs function as low-noise amplifiers and cryogenic switches. The current baseline for PICO is to use a time-domain multiplexer (TDM), which assigns each detector’s address in a square matrix of simultaneously read columns, and sequentially cycles through each row of the array [230]. The PICO baseline architecture uses a matrix of 128 rows and 102 columns. The thermal loading on the cold stages from the wire harnesses is subdominant to conductive loading through the mechanical support structures.

Because SQUIDS are sensitive magnetometers, suborbital experiments have developed techniques to shield them from Earth’s magnetic field using highly permeable materials and superconducting materials [231]. Total suppression factors better than  $10^7$  have been demonstrated for dynamic magnetic fields [232]. PICO will use these demonstrated techniques to shield SQUID readout chips from the ambient magnetic environment, which is 20,000 times smaller than near Earth, as well as from fields generated by on-board components including the cADR (§ 3.4.1). The cADR is delivered with its own magnetic shielding which reduces the field to less than 0.1 G (less than that experienced by suborbital experiments).

SQUIDS are also sensitive to radio-frequency interference (RFI). Several suborbital experiments have demonstrated RFI shielding using aluminized mylar wrapped at cryogenic stages to form a Faraday cage around the SQUIDS [233–235]. Cable shielding extends the Faraday cage to the detector warm readout electronics.

Redundant warm electronics boxes perform detector readout and instrument housekeeping using commercially available radiation hardened analog-to-digital converters (ADCs), requiring 75 W total. The readout electronics compress the data before delivering them to the spacecraft, requiring an additional 15 W. PICO detectors produce a total of 6.1 Tbits/day assuming 16 bits/sample, sampling rates from Table 3.1, and bolometer counts from Table 3.2. *Planck* HFI typically achieved  $4.7\times$  compression in flight with information loss increasing noise by  $\sim 10\%$  [236, 237]. Suborbital work has demonstrated  $6.2\times$  lossless compression [238]. PICO assumes  $4\times$  lossless compression.

## 3.4 Thermal

Like the *Planck*-HFI instrument, PICO’s focal plane is maintained at 0.1 K to enable detector operation (§ 3.4.1). sentence not quite correct To minimize detector noise due to instrument thermal radiation, the aperture stop and reflectors are cooled using both active and radiative cooling (§ 3.4.2, § 3.4.3, Fig. 3.3). All thermal requirements are met with robust margins (Table 3.3).

### 3.4.1 cADR sub-kelvin cooling

A multi-stage continuous adiabatic demagnetization refrigerator (cADR) maintains the PICO focal plane at 0.1 K and the surrounding enclosure, filter, and readout components at 1 K. The cADR employs three refrigerant assemblies operating sequentially to absorb heat from the focal plane at 0.1 K and reject it to 1 K. Two additional assemblies, also operating sequentially, absorb this rejected heat at 1 K, cool other components to 1 K, and reject heat at 4.5 K. This configuration provides continuous cooling with small temperature variations at both the 0.1 K and 1 K. Heat straps connect the two cADR cold sinks to multiple points on the focal plane assembly was ‘focal plane assembly’ defined? isn’t it at 0.1K?, which has high thermal conductance paths built in, to provide spatial temperature uniformity and stability during operation.

Table 3.3: Projected cooler heat lift capabilities offer more than 100 % heat lift margin, complying with space technology best practices [239]. The cADR lift capability at 1 K and 0.1 K is from a Goddard quote. Both NGAS and Ball project > 100 mW lift capability at 4.5 K using higher compression-ratio compressors currently in development (§ 3.4.2). The required loads were calculated using Thermal Desktop. Reference [240] was used to estimate the thermal conductive loads through mechanical supports. In addition to the listed components, the total 4.5 K heat load includes the intercept on the focal plane mechanical supports.

Component	$T$ (req'd)	$T$ (CBE)	Req'd per model	Active heat lift Capability today	Projected capability
Primary reflector . . . . .	< 40 K	17 K		N/A (radiatively cooled)	
Secondary reflector . . . . .	< 8 K				
Aperture stop . . . . .	4.5 K	4.5 K	42 MW total at 4.5 K	–	> 100 mW (§ 3.4.2, Fig. 3.4)
cADR heat rejection . . . . .	4.5 K				
Focal plane enclosure and filter . . . . .	1.0 K	1.0 K	0.36 mW	1.0 mW	–
Focal plane . . . . .	0.1 K	0.1 K	5.7 $\mu$ W	32 $\mu$ W	–

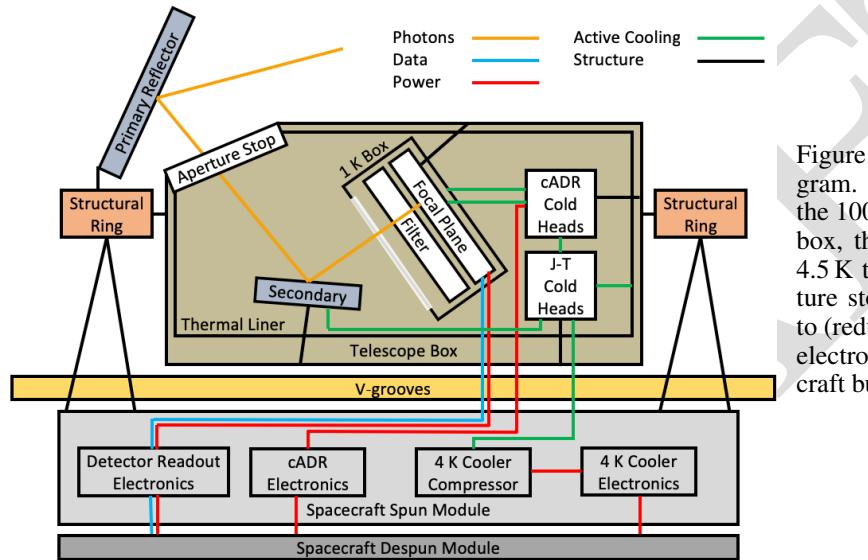


Figure 3.3: PICO instrument block diagram. Active coolers provide cooling to the 100 mK focal plane, the surrounding 1 K box, the 4.5 K secondary reflector, and the 4.5 K thermal liner that acts as a cold aperture stop. Data from the focal plane flows to (redundant, cross-strapped) warm readout electronics on the spun module of the spacecraft bus.

The detector arrays are thermally sunk to the mounting frame. Heat loads in the range of  $30 \mu\text{W}$  at 0.1 K and 1 mW at 1 K (time-average) are within the capabilities of current cADRs developed by GSFC [241, 242]. The PICO sub-kelvin loads are estimated at less than half of this capability (Table 3.3). *this paragraph doesn't say anything about the maturity of the multiple stage cADR, and about flight heritage. It should, even if to point to later paragraphs*

### 3.4.2 4 K Cooler

A cryocooler system similar to that used on JWST to cool the MIRI detectors [243, 244] removes the heat rejected from the cADR and cools the aperture stop and secondary reflector to 4.5 K. Both NGAS (which provided the MIRI coolers) and Ball Aerospace have developed such coolers under the NASA-sponsored Advanced Cryocooler Technology Development Program [245]. *is there a demonstration of the ball cooler? NGAS and Ball use slightly different but functionally-equivalent hardware approaches. A 3-stage precooler (acoustic Stirling by NGAS or mechanical Stirling by Ball) provides  $\sim 16$  K precooling to a separate circulated-gas loop driven by a similar compressor modified for DC flow very technical; is the reader supposed to follow this? is the difference between the ngas and ball relevant?* The circulated-gas loop utilizes Joule–Thomson (J-T) expansion, further cooling the gas to 4.5 K. The entire precooler assembly and the J-T circulator compressor are located on the warm spacecraft spun module (Fig. 3.3), with relatively short tubing lengths conducting the gas flow from the precooling point to the J-T expansion point.



Figure 3.4: Projected performance of the NGAS cooler using a multi-stage compressor and  $^4\text{He}$  [244] meets PICO’s requirements with  $> 100\%$  margin. PICO conservatively carries additional input power contingency on the efficiency of the cooler.

All waste heat rejected by the cooler compressors and drive electronics is transferred to the spacecraft heat rejection system. Unlike JWST, the PICO cooler does not require deployment of the remote cold head.

The J-T expansion point is located close to the cADR heat rejection point, thereby providing the lowest temperature to the cADR. Subsequent to cooling the cADR, the gas flow intercepts conducted heat to the focal plane enclosure, then cools the aperture stop and the secondary reflector before returning in counter-flow to the circulation compressor. Model-based projections indicate that the coolers delivered for MIRI could meet the PICO 4.5 K heat lift requirement with more than 100 % margin with minimal changes: the replacement of the  $^4\text{He}$  gas used for MIRI with  $^3\text{He}$ , plus resizing of the gas counterflow heat exchangers to take advantage of the  $^3\text{He}$  properties. **have we thought through about the change in gas? is it obvious that no other changes will be required as a consequence of the change in gas type?**

It is highly likely that a better solution will be available before Phase A. **why do we need a better solution? do you mean a cooler with even larger margin?** NGAS and Ball are actively working on increasing the flow rate and compression ratio of the J-T compressor, which should result in significantly higher system efficiency, and in greater heat-lift margin above the PICO requirement. These improvements entail the implementation of well-known techniques (standard thermal engineering) **what does this sentence mean?** The NGAS multi-stage J-T compressor has completed PDR-level development, and is expected to reach CDR level well before needed for PICO. Projected performance is shown in Fig. 3.4. The Ball approach started with a larger compressor, required less modification to achieve comparable performance, and eliminated the cold bypass-precooling valve that was problematic for MIRI. **must we mention problematic valves?** The Ball approach uses  $^3\text{He}$ , while the NGAS approach uses  $^4\text{He}$ . Both employ re-optimized gas heat exchangers (trivial engineering changes). **what does this mean?**

### 3.4.3 Radiative cooling

A set of four V-groove radiators provides passive cooling. The outermost of the four V-groove shields shadows the interior shields from the Sun. The V-grooves radiate to space, each reaching successively cooler temperatures. **can we be quantitative?** The V-groove assembly provides a cold radiative environment to the primary reflector, structural ring, and telescope box, so radiative loads on those elements are smaller than the conductive loads through the mechanical support structures. **shouldn’t we show evidence that say radiative cooling has been worked out successfully before.?**

## 3.5 Instrument integration and test

PICO instrument I&T planning benefits greatly from heritage experience with the *Planck* HFI instrument [246].

PICO screens detector wafer performance prior to selection of flight wafers and focal plane integration.

The cADR and 4 K cryocooler are qualified prior to delivery. The relative alignment of the two reflectors under thermal contraction is photogrammetrically verified in a thermal vacuum (TVAC) chamber.

PICO integrates the flight focal plane assembly and flight cADR in a dedicated sub-kelvin cryogenic testbed. Noise, responsivity, and focal-plane temperature stability are characterized using a representative optical load for each frequency band (temperature-controlled blackbody). Polarimetric and spectroscopic calibration are performed.

The focal plane is integrated with the reflectors and structures, and alignment verified photogrammetrically at cold temperatures in a TVAC chamber. The completely integrated observatory (instrument and spacecraft bus) is tested in TVAC to measure parasitic optical loading from the instrument, noise, microphonics, and radio-frequency interference (RFI). The observatory is 4.5 m in diameter and 6.1 m tall, with no deployables.

## 4 Design reference mission

The PICO design reference mission is summarized in Table 4.1.

### 4.1 Concept of operations

The PICO concept of operations is similar to that of the successful *WMAP* [247] and *Planck* [248] missions. After launch, PICO cruises to a quasi-halo orbit around the Earth–Sun L2 Lagrange point (§ 4.1.1). A two-week decontamination period is followed by instrument cooldown, lasting about two months. After in-orbit checkout is complete, PICO begins the science survey.

PICO has a single science observing mode, surveying the sky continuously for 5 years using a pre-planned repetitive survey pattern (§ 4.1.2). Instrument data are compressed and stored on-board, then returned to Earth in daily 4-hr Ka-band science downlink passes (concurrent with science observations). Because PICO is observing relatively static Galactic, extragalactic, and cosmological targets, there are no requirements for time-critical observations or data latency. Presently, there are no plans for targets of opportunity or guest observer programs during the prime mission. The PICO instrument does not require cryogenic consumables (as the *Planck* mission did), permitting consideration of significant mission extension beyond the prime mission.

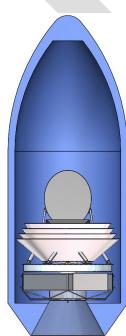


Figure 4.1: PICO is compatible with the Falcon 9.

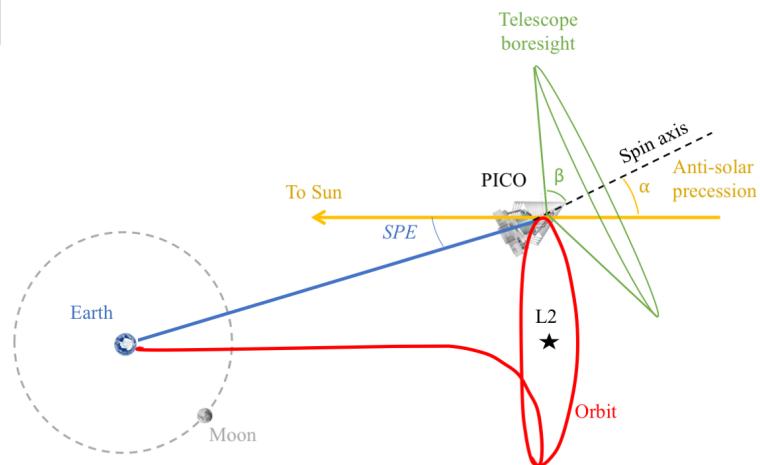


Figure 4.2: PICO surveys by continuously spinning the instrument about a precessing axis.

Table 4.1: PICO carries margin on key mission parameters. Maximum Expected Value (MEV) includes contingency.

Orbit type . . . . .	Sun-Earth L2 Quasi-Halo
Mission class . . . . .	Class B
Mission duration . . . . .	5 years
Propellant (hydrazine) . . . . .	213 kg (77 % tank fill)
Launch mass (MEV) . . . . .	2147 kg (3195 kg capability)
Max power (MEV) . . . . .	1320 W (with 125 % margin on available solar array area)
Onboard data storage . . . . .	4.6 Tb (3 days of compressed data, enabling retransmission)
Survey implementation . . . . .	Instrument on spin table
Attitude control . . . . .	Zero-momentum 3-axis stabilized

#### 4.1.1 Mission design and launch

PICO performs its science survey from a quasi-halo orbit around the Earth–Sun L2 Lagrange point. Predecessor missions *Planck* and *WMAP* both operated in L2 orbits.

L2 orbits provide favorable survey geometry (relative to Earth orbits) by mitigating viewing restrictions imposed by terrestrial and lunar stray light. The PICO orbit around L2 is small enough to ensure than the Sun–Probe–Earth (SPE) angle is less than  $15^\circ$ . This maintains the telescope boresight  $> 70^\circ$  away from the Earth (Fig. 4.2,  $70^\circ = 180^\circ - \alpha - \beta - \text{SPE}$ ).

High data rate downlink to the Deep Space Network (DSN) is available from L2 using near-Earth Ka bands. L2 provides a stable thermal environment, simplifying thermal control. The PICO orbit exhibits no post-launch eclipses.

NASA requires that Probes be compatible with an Evolved Expendable Launch Vehicle (EELV). For the purpose of this study, the Falcon 9 [249] is used as the reference vehicle. Figure 4.1 shows PICO configured for launch in a Falcon 9 fairing. The Falcon 9 launch capability for ocean recovery exceeds PICO’s 2147 kg total launch mass (including contingency) by a  $\sim 50\%$  margin.

Insertion to the halo manifold and associated trajectory correction maneuvers (TCMs) require  $150 \text{ m s}^{-1}$  of total  $\Delta V$  by the spacecraft. Orbit maintenance requires minimal propellant (statistical  $\Delta V \sim 2 \text{ m s}^{-1} \text{ year}^{-1}$ ). The orbital period is  $\sim 6$  months. There are no disposal requirements for L2 orbits, but spacecraft are customarily decommissioned to heliocentric orbit.

#### 4.1.2 Survey design

PICO employs a highly repetitive scan strategy to map the full sky. During the survey, PICO spins with a period  $T_{\text{spin}} = 1 \text{ min}$  about a spin axis oriented  $\alpha = 26^\circ$  from the anti-solar direction (Fig. 4.2). This spin axis is forced to precess about the anti-solar direction with a period  $T_{\text{prec}} = 10 \text{ hr}$ . The telescope boresight is oriented at an angle  $\beta = 69^\circ$  away from the spin axis (Fig. 3.1). This  $\beta$  angle is chosen such that  $\alpha + \beta > 90^\circ$ , enabling mapping of all ecliptic latitudes. The precession axis tracks with the Earth in its yearly orbit around the Sun, so this scan strategy maps the full sky (all ecliptic longitudes) within 6 months.

PICO’s  $\alpha = 26^\circ$  is chosen to be substantially larger than the *Planck* mission’s  $\alpha$  angle ( $7.5^\circ$ ) to mitigate systematic effects by scanning across each sky pixel with a greater diversity of orientations [250]. Increasing  $\alpha$  further would decrease the sun-shadowed volume available for the optics and consequently reduce the telescope aperture size. A deployable sun shade was considered but found not to be required, and was thus excluded in favor of a more conservative and less costly approach.

The instrument spin rate, selected through a trade study, matches that of the *Planck* mission. The study balanced low-frequency ( $1/f$ ) noise subtraction (improves with spin rate) against implementation cost and



Figure 4.3: Modular equipment bays provide easy access to all components in the spacecraft de-spun module and enable parallel integration of spacecraft subsystems.

heritage, pointing reconstruction ability (anti-correlated with spin rate), and data volume (linearly correlated with spin rate). The CMB dipole appears in the PICO data timestream at the spin frequency (1 rpm = 16.7 mHz). Higher multipole signals appear at harmonics of the spin frequency starting at 33 mHz, above the knee in the detector low-frequency noise (§ 3.2.3). A destriping mapmaker applied in data post-processing effectively operates as a high-pass filter, as demonstrated by *Planck* [251]. PICO’s spin axis precession frequency is  $> 400\times$  faster than that of *Planck*, greatly reducing the effects of any residual  $1/f$  noise by spreading the effects more isotropically across pixels.

## 4.2 Ground segment

The PICO Mission Operations System (MOS) and Ground Data System (GDS) can be built with extensive reuse of standard tools. The PICO concept of operations is described in § 4.1. All space-ground communications, ranging, and tracking are performed by the Deep Space Network (DSN) 34 m Beam Wave Guide (BWG). X-band is used to transmit spacecraft commanding, return engineering data, and provide navigation information (S-band is a viable alternative, and could be considered in a future trade). Ka-band is used for high-rate return of science data. The baseline 150 Mb/s transfer rate (130 Mb/s information rate after CCSDS encoding) is an existing DSN catalog service [252]. The instrument produces 6.1 Tb/day, which is compressed to 1.5 Tb/day (§ 3.3). Daily 4 hr DSN passes return PICO data in 3.1 hr, with the remaining 0.9 hr available as needed for retransmission or missed-pass recovery.

## 4.3 Spacecraft

The PICO spacecraft bus is Class B and designed for a minimum lifetime of 5 years in the L2 environment. Mission critical elements are redundant. Flight spares, engineering models and prototypes appropriate to Class B are budgeted.

The aft end of the spacecraft (the “de-spun module”) is comprised of six equipment bays that house standard components (Fig. 4.3). The instrument and V-grooves are mounted on bipods from the spacecraft “spun module,” which contains hosted instrument elements (Fig. 3.1). A motor drives the spun module at 1 rpm to support the science survey requirements (§ 4.1.2). Reaction wheels on the despun module cancel the angular momentum of the spun module and provide three-axis control (§ 4.3.1).

The bipods that mechanically support the instrument are thermally insulating. The passively radiating V-groove assembly thermally isolates the instrument from solar radiation and from the bus (§ 3.4.3). Like

*Planck* [248], the V-grooves are manufactured using honeycomb material. Additional radiators on the spun and despun spacecraft modules ( $\sim 1 \text{ m}^2$  each) reject heat dissipated by spacecraft subsystems and hosted instrument elements.

PICO's avionics are dual-string with standard interfaces. Solid state recorders provide three days of science data storage (4.6 Tbit), enabling retransmission of missed data.

PICO employs a fully redundant Ka- and X-band telecommunications architecture. The Ka-band system uses a 0.3 m high-gain antenna to support a science data downlink information rate of 130 Mb/s to a 34 m BWG DSN ground station with a link margin of 4.8 dB. The X-band system provides command and engineering telemetry communication through all mission phases using medium and low gain antennas. Amplifiers, switches, and all three antennas are on a gimballed platform, enabling Ka and X-band downlink concurrent with science observations.

The heritage power electronics are dual-string. A 74 A-hr Li-ion battery is sized for a 3 hr launch phase with 44 % depth of discharge. After the launch phase, the driving mode is telecom concurrent with science survey (1320 W including 43 % contingency). Solar cells on the aft side of the bus (5.8 m<sup>2</sup> array,  $\alpha = 26^\circ$  off-Sun) support this mode with positive power, and unused area in the solar array plane (7.4 m<sup>2</sup> more area by growing to 4.5 m diameter) affords 125 % margin (Fig. 4.3).

The propulsion design is a simple mono-propellant blow-down hydrazine system with standard redundancy. Two aft-pointed 22 N thrusters provide  $\Delta V$  and attitude control for orbit insertion and maintenance (§ 4.1.1), requiring 140 kg of propellant. Eight 4 N thrusters provide reaction wheel momentum management and backup attitude control authority (60 kg of propellant). Accounting for ullage (14 kg), the baseline propellant tank fill fraction is 77 %.

### 4.3.1 Attitude determination and control

PICO uses a zero net angular momentum control architecture with heritage from the SMAP mission (§ 6.3). PICO's instrument spin rate (1 rpm) matches that of the *Planck* mission, but the precession of the spin axis is much faster (10 hr vs 6 months), and the precession angle much larger ( $26^\circ$  vs  $7.5^\circ$ ). These differences make the spin-stabilized *Planck* control architecture impractical because of the amount of torque that would be required to drive precession.

The PICO 1 rpm instrument spin rate is achieved and maintained using a spin motor. The spin motor drive electronics provide the coarse spin rate knowledge used for controlling the spin rate to meet the  $\pm 0.1$  rpm requirement. Data and power are passed across the interface using slip rings.

Based on mass properties derived from the PICO CAD model, PICO requires  $\sim 220 \text{ N m s}$  to cancel the angular momentum of the instrument and spacecraft spun module (including mass contingency) at 1 rpm. Three Honeywell HR-16 reaction wheel assemblies (RWAs), each capable of  $150 \text{ N m s}$ , are mounted on the despun module parallel to the instrument spin axis, and spin opposite to the instrument to achieve zero net angular momentum. The despun module is three-axis stabilized. The spin axis is precessed using three RWAs mounted normal to the spin axis in a triangle configuration. Each set of three RWAs is sized such that two could perform the required function with margin, providing single fault tolerance.

Spin axis pointing and spin rate knowledge are achieved and maintained using star tracker and inertial measurement unit (IMU) data. The attitude determination system is single-fault tolerant, with two IMUs each on the spun and despun modules, and two star trackers each on the spun and despun modules. Two sun sensors on the despun module are used for safe-mode contingencies and instrument Sun avoidance. All attitude control and reconstruction requirements are met, including spin axis control  $< 60 \text{ arcmin}$  with  $< 1 \text{ arcmin/min}$  stability, and reconstructed pointing knowledge  $< 10 \text{ arcsec}$  (each axis,  $3\sigma$ ).

Additional pointing reconstruction is performed in post-processing using the science data. The PICO instrument will observe planets (compact, bright sources) nearly every day. By fitting the telescope pointing to the known planetary ephemerides, the knowledge of the telescope boresight pointing and the relative pointing of each detector will improve to better than 1 arcsec (each axis,  $3\sigma$ ). *Planck*, with fewer detectors,

making lower signal-to-noise ratio measurements of the planets, and observing with a scan strategy that acquired measurements of each planet only once every 6 months, demonstrated 0.8 arcsec ( $1\sigma$ ) pointing reconstruction uncertainty in-scan and 1.9 arcsec ( $1\sigma$ ) cross-scan [253].

## 5 Technology maturation

PICO builds off of the heritage of *Planck* HFI and *Herschel*. Since *Planck* and *Herschel*, suborbital experiments have used monolithically fabricated TES bolometers and multiplexing schemes to field instruments with thousands of detectors per camera (Fig. 5.1). By the time PICO enters Phase A, the Simons Observatory plans to be operating 60 000 detectors [254].

The remaining technology developments required to enable the PICO baseline design are:

1. Extension of three-color antenna-coupled bolometers down to 21 GHz and up to 462 GHz (§ 5.1).
2. Construction of high-frequency direct absorbing arrays and laboratory testing (§ 5.2).
3. Beam line and 100 mK testing to simulate the cosmic ray environment at L2 (§ 5.3).
4. Expansion of time-division multiplexing to support 128 switched rows per readout column (§ 5.4).

We recommend APRA and SAT support to complete development of these technologies through the milestones described in Table 5.1.

### 5.1 21–462 GHz bands

Suborbital teams have successfully demonstrated a variety of optical coupling schemes, including horns with ortho-mode transducers (OMTs), lithographed antenna arrays, and sinuous antennas under lenslets (Fig. 5.2). All have achieved background-limited performance in suborbital instruments with sufficient margin on design parameters to achieve this performance in the more demanding low background at L2. All have been packaged into modules and focal plane units in working cameras representative of the PICO integration. Experiments have covered many of PICO’s observing bands between 27 GHz and 270 GHz (Table 5.2). To date, suborbital experiments have achieved statistical map depths of  $3 \mu\text{K}_{\text{CMB}}$  arcmin on degree-scaled modes over small parts of the sky, within an order of magnitude of what PICO achieves over the entire sky (Table 3.2), and have demonstrated systematic control better than this level through full-pipeline simulations and null-test analysis (jackknife tests).

The baseline PICO instrument requires three-color dual-polarized antenna-coupled bolometers covering bands from 21 to 462 GHz (§ 3.2.1). The sinuous antenna has the bandwidth to service three colors per pixel, whereas horns and antenna arrays have only been used for two, so some version of the sinuous antenna will likely be needed to realize three-color pixels. SPT-3G has used the PICO-baselined three-color pixel design to deploy 16,260 detectors covering 90–150–220 GHz [255].

The extension to lower frequencies requires larger antennas and therefore control of film properties and lithography over larger areas. Scaling to higher frequencies forces tighter critical dimensions and materials tend to exhibit higher losses. Lenslets also need to be developed with multilayer AR coatings three times



Figure 5.1: SPT-3G operates a focal plane of three-color antenna-coupled pixels with  $> 16000$  total bolometers. [255]

Table 5.1: PICO technologies can be developed to TRL 5 prior to a 2023 Phase A start using the APRA and SAT programs, requiring a total of about \$ 13M. Per NASA guidance, these costs are outside the mission cost described in § 6.5.

Task	Current status	Milestone A	Milestone B	Milestone C	Current funding	Required funding	Date TRL5 achieved
1a. Three-color arrays $v < 90 \text{ GHz}$	2-color lab demos $v > 30 \text{ GHz}$	Field demo of 30–40 GHz (2020)	Lab demos 20–90 GHz (2022)	–	APRA & SAT funds	\$2.5M over 4 yr (1 APRA + 1 SAT)	2022
1b. Three-color arrays $v > 220 \text{ GHz}$	2-color lab demos $v < 300 \text{ GHz}$	Field demo of 150–270 GHz (2021)	Lab demos 150–460 GHz (2022)	–	APRA & SAT funds	\$3.5M over 4 yr (2 SATs)	2022
2. Direct absorbing arrays $v > 50 \text{ GHz}$	0.1–5 THz unpolarized	Design & prototype of arrays (2021)	Lab demo of 555 GHz (2022)	Lab demo of 799 GHz (2023)	None	\$2M over 5 yr (1 SAT)	2023
3. Cosmic ray studies	250 mK w/ sources	100 mK tests with sources (2021)	Beamline tests (2023)	–	APRA & SAT funds	\$0.5–1M over 5 yr (part of 1 SAT)	–
4a. Fast readout electronics	MUX66 demo	Engineering and Fab of electronics (2020)	Lab demo (2021)	Field demo (2023)	No direct funds	\$4M over 5 yr (1 SAT)	2023
4b. System engineering; 128× MUX demo	MUX66 demo	Design of cables (2020)	Lab demo (2021)	Field demo (2023)	No direct funds		

as thick and twice as thin as current designs to cover the lowest and highest frequency channels. These challenges require tight control of cleanliness and full understanding of process parameters. All developments require careful characterization of beam properties and studies of associated systematic challenges for PICO.

The sinuous antenna couples to states that “wobble” log-periodically with frequency. There are potential solutions to this in the focal plane design, analysis, and free parameters of the antenna geometry. Systematics studies for field demonstrations will be particularly important. The PICO concept is robust to any challenges in developing three-color pixels; § 5.5 describes an option to descope to two-color horn-coupled pixels.

## 5.2 555–799 GHz bands

The baseline PICO instrument requires single-color dual-polarized direct-absorbing bolometers from 555 to 799 GHz (§ 3.2.2). *Planck* and *Herschel* demonstrated the architecture of horns coupled to direct absorbing bolometers (Fig. 5.3). Ground experiments with similar designs have deployed focal planes with hundreds of horn-coupled spiderweb bolometers, replacing the *Planck* and *Herschel* NTD-Ge thermistors with TESs, and adjusting time constants as necessary (Table 5.3). *Planck* HFI and SPT-pol 90 demonstrated dual-polarized detectors. *Herschel* and SPT-SZ demonstrated monolithic unpolarized detectors. PICO will require detectors that merge these two designs in monolithic dual-polarized arrays. The greatest remaining challenge is the low risk development of a packaging design. Filled arrays of detectors such as Backshort Under Ground (BUG) bolometers are also an option [257].

## 5.3 Environmental testing

Laboratory tests and in-flight data from balloons suggest that TES bolometer arrays may be more naturally robust against cosmic rays than the individual NTD-Ge bolometers used in *Planck*. PICO will leverage lessons learned from *Planck* and ensure robust thermal sinking of detector array substrates. Cosmic ray

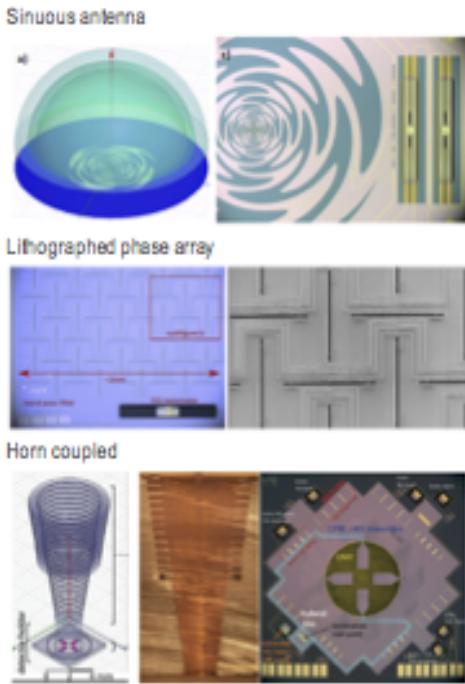


Figure 5.2: Multiple demonstrated optical coupling schemes are available to PICO. Images from CMB-S4 Technology Book [256].

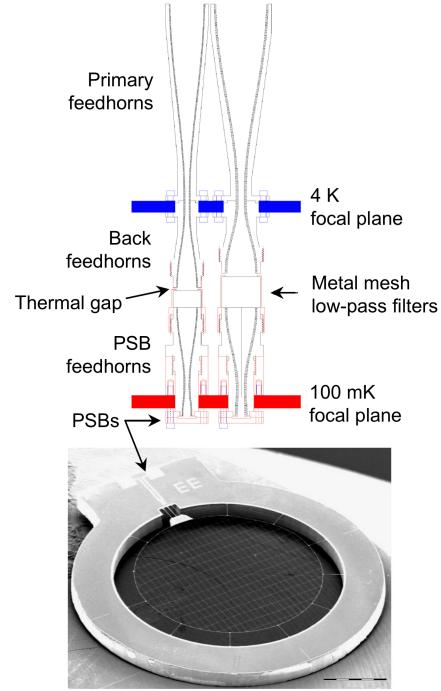


Figure 5.3: Direct-absorbing dual-polarized detectors and coupling horns used in *Planck* for 143–343 GHz bands.

glitches have fast recovery times and low coincidence rates [258, 259]. Residual risk can be retired with 100 mK testing where the array heat sinking may be weaker, and beam-line tests to simulate the expected flight environment.

#### 5.4 Multiplexing

More than ten experiments have used time-domain multiplexer (TDM) readout. SCUBA2 on JCMT has 10 000 pixels, nearly as many detectors as planned for PICO [260]. Most of these experiments have used 32-row multiplexing. Recently ACT has expanded this to 64-row multiplexing [230].

PICO’s sensitivity requirements dictate the use of  $\sim 13\,000$  transition-edge-sensor bolometers, requiring a highly multiplexed system. The PICO baseline design calls for TDM with 128 switched rows per readout column (TDM-128 $\times$ ). The leap to TDM-128 $\times$  requires:

- development of fast-switched room temperature electronics; and
- system engineering of room temperature to cryogenic row select cabling to ensure sufficiently fast row switch settling times.

The historical row revisit rate for bolometric instruments using 32 $\times$  TDM has been 25 kHz [e.g., 221]. However, x-ray instruments using TDM routinely switch between rows at a rate of 160 ns [261]. The PICO baseline assumes a 160 ns switch rate and TDM-128 $\times$ , which dictates a row revisit rate (effective sampling rate) of 48.8 kHz. To limit aliased noise, PICO implements L/R filters in each readout channel with a bandwidth of 6 kHz, dictated by detector stability considerations and the required  $\sim 1$  kHz signal bandwidth. With these parameters and using the same TDM multiplexer SQUID design, the increased total noise due to aliasing will be limited to 15 %. The system engineering study will culminate in a demonstration of TDM-128 $\times$  SQUID aliased noise below PICO detector sensitivity requirements.

Table 5.2: Multiple active suborbital efforts are advancing technologies relevant to PICO.

Project	Type	Optical Coupling	$v_c$ [GHz]	Colors per pixel	$N_{\text{bolo}}$	Significance	Reference
PICO baseline . . . . .	Probe		21 – 462	Three	11,796		§ 3.2.1
SPT-3G . . . . .	Ground	Sinuous	90 – 220	Three	16,260	Trichroic	[255]
Advanced ACT-pol . . . .	Ground	Horns	27 – 230	Two	3,072	Dichroic	[262]
BICEP/Keck . . . . .	Ground	Antenna arrays	90 – 270	One	5,120	50 nK-deg	[31]
Berkeley, Caltech, NIST	Lab	Various	30 – 270	Various	–	Band coverage	[231, 263, 264]
SPIDER . . . . .	Balloon	Antenna arrays	90 – 150	One	2,400	Stable to 10 mHz	[229]

Table 5.3: PICO high-frequency detectors leverage development and demonstration by *Planck*, *Herschel*, and SPT.

Project	Type	Polarized	Mono- lithic?	$v_c$ [GHz]	Colors per pixel	$N_{\text{bolo}}$	Significance	Reference
PICO baseline . . . . .	Probe	Yes	Yes	555 – 799	One	1,200		§ 3.2.2
<i>Planck</i> HFI . . . . .	Flight	143–343 GHz	No	143 – 857	One	48	TRL 9 polarized	[227]
<i>Herschel</i> . . . . .	Flight	No	Yes	570 – 1200	One	270	TRL 9 monolithic	[265]
SPT-SZ . . . . .	Ground	No	Yes	90 – 220	One	840	Monolithic array TESs	[225]
SPT-pol-90 . . . . .	Ground	Yes	No	90	One	180	Dual pol absorbing TESs	[266]

## 5.5 Technology desscopes

A descope from three-color sinuous antenna/lenslet-coupled pixels to two-color horn-coupled pixels remains a viable alternative should the three-color technology not mature as planned. Descope studies suggest that a PICO-size focal plane using two-color horn-coupled pixels at the lower frequencies and the baseline one-color pixels at the higher frequencies would contain 8 840 detectors (compared to the baseline 12 966) and map in 19 colors (baseline 21). Because horns have a 2.3 : 1 bandwidth, each of the two bands in a pixel has 35 % bandwidth (compared to the baseline 25 %), which compensates for pixel count, resulting in 0.61  $\mu\text{K}_{\text{CMB}}$  arcmin aggregate CBE map depth, which matches the PICO CBE map depth, and affords > 40 % margin against the 0.87  $\mu\text{K}_{\text{CMB}}$  arcmin baseline requirement (Table 3.2), but with coarser spectral resolution. Detailed analysis could be performed to assess the impact on signal component separation (§ 2.6).

## 5.6 Enhancing technologies

The following technologies are neither required nor assumed by the PICO baseline concept. They represent opportunities to extend scientific capabilities or simplify engineering.

PICO baselines TDM readout because of its relative maturity and demonstrated sensitivity and stability in relevant science missions. Lab tests of Frequency Domain Multiplexing (FDM) suggest comparable performance with higher multiplexing factors and lower loads on cryogenic stages relative to TDM. Suborbital experiments such as SPT-3G have used frequency division multiplexing (FDM) to readout focal planes comparable in size to PICO.

Microwave frequency SQUID multiplexing can increase the multiplexing density and reduce the number of lines between the 4 K and ambient temperature stages [267, 268]. Kinetic Inductance Detectors (KIDs) and Thermal KIDs (TKIDs) can further reduce the wire count, obviate the SQUIDs, and dramatically simplify integration by performing multiplexing on the same substrate as the detectors themselves [269, 270]. The cost to develop these technologies is \$3–4M/year, with a high chance of reaching TRL-5 before Phase A.

# 6 Project management, heritage, risk, and cost

## 6.1 PICO Study Participants

The PICO study was open to the entire mm/sub-mm science community. Seven working groups were led by members of PICO’s Executive Committee, which met weekly under the leadership of PI Shaul Hanany.

More than 60 people participated in-person in two community workshops (November 2017 and May 2018).

The PICO engineering concept definition package was generated by Team X (the JPL concurrent design lab). The Team X study was supported by inputs from a JPL engineering team and Lockheed Martin.

The full list of study report contributors and endorsers is on page [i](#).

## 6.2 Project management plan

PICO benefits from the experience of predecessor missions such as *Planck* and *WMAP*, as well as many years of investment in technology development and a multitude of suborbital experiments. In addition to demonstrated science and engineering capabilities, this heritage has developed a community of people with the expertise required to field a successful mission.

This study assumes mission management by JPL with a Principal Investigator leading a single science team. A Project Manager provides project oversight for schedule, budget, and deliverables. A Project Systems Engineer leads systems engineering activities and serves as the Engineering Technical Authority. A Mission Assurance Manager serves as the Independent Technical Authority. The PICO mission development schedule is shown in Fig. 6.1.

	FY24	FY25	FY26	FY27	FY28	FY29	FY30	FY31	FY32	FY33	FY34	FY35
CY 2023	CY 2024	CY 2025	CY 2026	CY 2027	CY 2028	CY 2029	CY 30	CY 31	CY 32	CY 33	CY 34	
PH A (12 mths)	PH B (12 mths)	PHASE C (22 mths)	PHASE D(18 mths)		PHASE E (5 yrs )							F 4 mths
♦ 10/23 KDP-A	♦ 10/24 KDP-B	♦10/25 KDP-C		♦ 8/27 KDP-D		♦ 2/29 PLAR (Start of Ph E)						KDP-F 2/34 ♦
Reviews	10/25 PDR ♦		♦ 7/26 CDR	♦ 7/27 ARR		Launch 1/29★						

Figure 6.1: The PICO baseline schedule is based on historical actuals from similarly-sized missions such as Juno and SMAP. Per NASA direction, Probe studies assume a Phase A start in October 2023.

Probes are medium-class missions, similar in cost scope to NASA’s New Frontiers missions, which are Category 1 and Risk Classification A or B, with Phase A–D costs capped at ~ \$850M (not including the launch vehicle). JPL is well-prepared to manage Probe missions, having managed the Juno New Frontiers mission (launched 2011) and also the development of the medium-class *Spitzer* Space Telescope (launched 2003). JPL delivered the bolometric detectors for the *Planck* HFI instrument (launched 2009). Presently, JPL is managing NEOCam, a Discovery class infrared space telescope.

The PICO spacecraft provider will be selected during mission formulation. Multiple organizations are capable of providing a spacecraft bus to meet PICO’s requirements. Lockheed Martin contributed to the PICO concept study, leveraging their experience with New Frontiers missions Juno and OSIRIS-REx.

## 6.3 Heritage

The successful *Planck* mission provides science heritage for PICO. Technical heritage traces to multiple missions.

Because PICO observes in the mm/sub-mm regime, the surface accuracy requirement for the reflectors is relatively easy to meet. PICO’s reflectors are similar to *Planck*’s, but somewhat larger (270cm × 205cm primary vs. 189cm × 155cm) [271]. *Herschel* observed at wavelengths more demanding than PICO’s and was larger (350cm diameter primary) [272].

The heritage of the PICO detectors and readout electronics (§ 3.2, § 3.3) is described in § 5.

PICO’s detectors are cooled by a cADR (§ 3.4.1) with requirements that are within the capabilities of current ADRs developed by Goddard Space Flight Center. These systems have been applied to several JAXA missions, including *Hitomi* [242].

PICO’s 4 K cryocooler (§ 3.4.2) is a direct extension of the JWST MIRI design [243, 244]. PICO benefits from a simpler and more reliable implementation of the J-T system than was required for MIRI, in that no

deployment of cooling lines is required, and all flow valving is performed on the warm spacecraft. Cooling multiple independent points with a J-T loop has been demonstrated on *Planck* with the JPL-supplied 18 K cooler [273].

Structures similar to PICO’s V-groove radiator assembly (§ 3.4.3) are a standard approach for passive cooling first described more than thirty years ago [274]. PICO has baselined a simple honeycomb material construction like that successfully flown by the *Planck* mission [273, 275].

Most requirements on the PICO spacecraft are well within typical ranges and can be met with standard high heritage systems (§ 4.3). PICO’s spin architecture and data volume requirements are less typical, and discussed below.

PICO’s spin system is generally less demanding than the successful SMAP spin system. PICO spins its instrument at 1 rpm, passes data and power across the spin interface (Fig. 3.3), and requires  $\sim 220$  N m s of spin momentum cancellation (§ 4.3.1). SMAP spins its 6-m instrument antenna at 14.6 rpm, successfully passes data and power across the spin interface, and requires 359 N m s of spin momentum cancellation [276].

Though PICO’s data volume is notable by current standards, it is already enveloped by missions in development. PICO produces 6.1 Tb/day of raw data which is compressed to 1.5 Tb/day (§ 3.3). PICO downlinks data daily, but baselines storage of 3 days of (compressed) data to mitigate missed telecom passes. This requires 4.5 Tb of onboard storage, in family with the 3.14 Tb solid state recorder currently in use by Landsat 8 and much smaller than the 12 Tb flash memory planned for NISAR [277]. The PICO baseline 150 Mb/s Ka-band data downlink is an existing DSN catalog service [252]. The baseline PICO mission generates  $\sim 2,200$  Tb of raw (uncompressed) data per year, less than the  $\sim 6,800$  Tb/year currently returned by Landsat 8 and  $\sim 9,300$  Tb/yr planned by NISAR [277].

## 6.4 Risk assessment

### 6.4.1 Pre-mission risks

Technology development (§ 5) is performed prior to the beginning of mission development, and is outside of the mission cost (per NASA direction), so associated risks do not represent threats to the cost of mission development. Rather, these technology development risks affect the availability of the described baseline mission. A technology-related mission descope is described in § 5.5.

### 6.4.2 Development risks

PICO’s healthy contingencies, margins, and reserves provide flexibility to address risks realized during mission development. PICO carries  $> 40\%$  instrument sensitivity margin (Table 3.2),  $> 100\%$  heat lift margin (Table 3.3), 43 % system power contingency, 31 % payload mass contingency, and 25 % spacecraft mass contingency. The Falcon 9 launch capability (assuming ocean recovery) exceeds PICO’s total launch mass (including contingency) by a  $\sim 50\%$  margin. The PICO budget includes 30 % cost reserves for Phases A–D (§ 6.5).

During mission development the Project Systems Engineer continually assesses risks, tracks progress toward retiring them, and updates mitigations. Mitigations for a few top risks identified during this study are described below.

- Thermal risk can be mitigated through extensive thermal modeling and review in Phase A, and design for early test verification.
- Risks associated with the instrument spin architecture can be mitigated by engaging JPL engineers who were involved in the SMAP mission.
- Detector delivery schedule risk can be mitigated by beginning fabrication early in the project life cycle and fabricating a generous number of detector wafers to ensure adequate yield. Multiple institutions (including, for example, JPL, GSFC, NIST, and ANL) would be capable of producing the PICO detectors. Suborbital programs generally achieve  $> 66\%$  detector wafer yield.

- Risks associated with the integration and test of a cryogenic instrument can be mitigated through advanced planning and allocation of appropriate schedule and schedule margin.

### **6.4.3 Operations risks**

The PICO design meets the requirements associated with the NASA Class B risk classification. For Class B missions, essential spacecraft and instrument functions are typically fully redundant. This increases mission cost, but significantly reduces the risk of mission failure.

The PICO mission utilizes a single instrument with a single observing mode mapping the sky using a repetitive survey pattern. The mission does not require any time-critical activities. The observatory fits in to the launch vehicle fairing in its operational configuration, so no hardware deployments are required. Because PICO observes at long wavelengths, the telescope does not require a dust cover (nor the associated mission-critical cover release).

The spacecraft incorporates a fault protection system for anomaly detection and resolution. The Sun-pointed, command receptive, thermally stable safe-mode attitude allows ground intervention for fault resolution without time constraints. PICO's high degree of hardware redundancy and onboard fault protection ensure spacecraft safety in the event of unforeseen failures and faults.

As described in § 2.6 and § 2.7, pre-Phase A simulation software maturation is recommended to mitigate the challenges associated with foreground separation and systematics control.

## **6.5 Mission cost**

We estimate PICO's total Phase A–E lifecycle cost between \$870M and \$960M, including the \$150M allocation for the Launch Vehicle (per NASA direction). These cost estimates include 30 % reserves for development (Phases A–D) and 13 % reserves for operations (Phase E). Pre-Phase-A technology maturation (§ 5) will be accomplished through the normal APRA and SAT processes, and is not included in the mission cost (per NASA direction).

Table 6.1 shows the mission cost breakdown, including the JPL Team X cost estimate as well as the PICO team cost estimate. Team X is JPL's concurrent design facility. Team X estimates are generally model-based, and were generated after a series of instrument and mission-level studies. Their accuracy is commensurate with the level of understanding typical to Pre-Phase-A concept development. They do not constitute an implementation or cost commitment on the part of JPL or Caltech.

The PICO team has generally adopted the Team X estimates, but also obtained a parametrically estimated cost range for the Flight System (WBS 6) and Assembly, Test and Launch Operations (ATLO, WBS 7) from Lockheed Martin Corporation to represent the cost benefits that might be realized by working with an industry partner. After adding estimated JPL overhead and Team X estimated V-groove assembly costs (not included in the Lockheed estimate), the PICO team cost is in-family with but lower than the Team X cost (Table 6.1).

Management, Systems Engineering, and Mission Assurance (WBS 1–3) development costs scale linearly with the WBS 4–12 development costs in the Team X model, and are adjusted accordingly in the PICO team estimate. Science team (WBS 4) costs are assessed by Team X based on PICO science team estimates of the numbers and types of contributors and meetings required for each year of PICO mission development and operations. These workforce estimates are informed by recent experience with the *Planck* mission.

Payload system (WBS 5) costs are discussed in detail in § 6.5.1. PICO's spacecraft (WBS 6) cost reflects a robust Class B architecture (§ 4.3). Mission-critical elements are redundant. Appropriate flight spares, engineering models and prototypes are budgeted. The V-groove assembly (§ 3.4.3) is costed in WBS 6. Mission operations (WBS 7), Ground Data Systems (WBS 9), and Mission Navigation and Design (WBS 12) costs reflect a relatively simple concept of operations (§ 4.1). PICO has a single instrument with a single science observing mode, surveying the sky continuously using a pre-planned repetitive survey pattern. Orbit maintenance activities are simple and infrequent.

Table 6.1: Detailed breakdown of Team X and PICO Team cost estimates (in FY18\$). Costs are based on the schedule in Fig. 6.1, which includes 5 years of operations.

Work Breakdown Structure (WBS) elements	Team X	PICO team
Development Cost (Phases A–D)	\$ 724M	\$ 634–677M
1.0, 2.0, 3.0 Management, Systems Engineering, and Mission Assurance	\$ 54M	\$ 47– 50M
4.0 Science	\$ 19M	
5.0 Payload System	\$ 168M	
6.0 Flight System	\$ 248M	\$ 210–240M
10.0 Assembly, Test, and Launch Operations (ATLO)	\$ 24M	
7.0 Mission Operations Preparation	\$ 16M	
9.0 Ground Data Systems	\$ 21M	
12.0 Mission and Navigation Design	\$ 7M	
Development Reserves (30%)	\$167M	\$ 146–156M
Operations Cost (Phase E)	\$ 84M	
1.0 Management	\$ 6M	
4.0 Science	\$ 20M	
7.0 Mission Operations	\$ 34M	
9.0 Ground Data Systems	\$ 14M	
Operations Reserves (13%)	\$ 10M	
Launch Vehicle Cost	\$ 150M	
Total Cost	\$ 958M	\$ 868–911M

### 6.5.1 Payload cost

The PICO payload consists of a single instrument: an imaging polarimeter. Payload costs are tabulated in Table 6.2.

The superconducting detectors require sub-kelvin cooling to operate. The active cooling system (the 0.1 K cADR and 4 K cryocooler, § 3.4.1 and § 3.4.2) comprises nearly half of the payload cost. The cADR cost for this study is an estimate from NASA Goddard Space Flight Center (GSFC), and assumes the provision of both a flight model and an engineering model. GSFC has produced ADRs for multiple spaceflight missions. The 4 K cryocooler cost for this study is based on the NASA Instrument Cost Model (NICM) VIII CER Cryocooler model [278], assuming a commercial build. PICO benefits greatly from recent and ongoing investment by commercial suppliers of 4 K coolers (as described in § 3.4.2). Team X used NICM VIII to model the cost of the focal plane and dual string readout electronics (§ 3.2, § 3.3). Team X estimated the telescope cost using the Stahl model [279]. The telescope is not a major cost driver, primarily because the reflectors only need to be diffraction limited at 330  $\mu\text{m}$  (900 GHz) (§ 3.1).

Based on JPL experience, 18 % of the instrument cost is allocated for integration and test. This includes integration and test of the flight focal-plane assembly with the flight cADR and then integration and test of the complete instrument including the focal-plane assembly, reflectors, structures, and coolers (§ 3.5). Integration and test of the instrument with the spacecraft is costed in WBS 10 (ATLO).

Table 6.2: Detailed breakdown of PICO instrument costs.

Instrument Elements	Cost
Management, Systems Eng., Assurance	\$ 18M
4 K Cooler and 0.1 K cADR . . . . .	\$ 71M
Focal plane and electronics . . . . .	\$ 27M
Mechanical, Thermal, Software . . . . .	\$ 17M
Telescope . . . . .	\$ 6M
Instrument integration and test . . . . .	\$ 29M
Total Instrument Cost . . . . .	\$ 168M

## NASA Standard Template Cost Table

### 2020 Astrophysical Decadal Survey - Probe Mission Preparatory Study Master Equipment List Based Parametric Total Lifecycle Cost Estimate

**Mission Name / Acronym:** PICO

**Cost Estimator:** JPL Team X

**Date of Cost Estimate:** October 9, 2018

**Cost Estimate Based On:** Final Master Equipment List

<u>PROJECT PHASE</u>		<u>COST [FY18 \$M]</u>
Phase A		(see Note 1)
Phases B-D	Mgmt, SE, MA	\$54
	Science	\$19
	Telescope	\$6
	Instrument	\$162
	Spacecraft, including ATLO	\$272
	MOS/GDS	\$44
	Launch Vehicle and Services	\$150
	Reserves	\$167
<b>Total Cost Phases B-D</b>		<b>\$874</b>
Phase E-F	Operations	\$74
	Reserves	\$10
	<b>Total Cost Phases E-F</b>	<b>\$84</b>
<b>TOTAL LIFECYCLE COST</b>		<b>\$958</b>

**Notes:**

- Team X estimates costs for Phase A-D. A break out of Phase A cost is not available. In this table, Phase A costs are included in Phase B-D.

- This parametric cost estimate is based on the Probe's Master Equipment List derived from the Final Engineering Concept Definition Package that accurately reflects the mission described in the Probe's Final Report. This estimate is to be used only for non-binding rough order of magnitude planning purposes.

- Team X estimates are generally model-based, and were generated after a series of instrument and mission level studies. Their accuracy is commensurate with the level of understanding typical to Pre-Phase-A concept development. They do not constitute an implementation or cost commitment on the part of JPL or Caltech.

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