

# **The Probe of Inflation and Cosmic Origins**

A Space Mission Study Report  
December, 2018

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# 1 Executive Summary

Recent theoretical developments and measurements of the cosmic microwave background (CMB) have uncovered tremendous potential for new discoveries over the next 10-20 years. The new discoveries are promising to be no less revolutionary than those attained to date. Many of the potential new discoveries are based on deeper measurements of the spatial pattern of the CMB's polarization. **would like to make this broader to not shortchange T-science, but still connect to  $E$ ,  $B$  in the next paragraph.**

# 2 Science

## 2.1 Introduction

The angular power spectra of sky-based  $Q$  and  $U$  polarization Stokes parameters are commonly recast in terms of curl-free  $E$  mode and gradient-free  $B$  mode patterns.  $E$  modes are generated by either scalar fluctuations, such as density perturbations in the early Universe or by tensor perturbations, such as gravitational waves.  $B$  modes are only generated through tensor perturbations. The Probe of Inflation and Cosmic Origins (PICO) is an imaging polarimeter designed to survey the entire sky at frequencies between 21 and 800 GHz with a polarization sensitivity that is 57 or 82 times that of the *Planck* mission for the PICO baseline and current best estimate (CBE) configurations, respectively. This sensitivity surpasses any other current or planned CMB instrument.

Fluctuations of the space-time metric during the epoch of inflation, near the Planck time, have generated gravitational waves that embed a unique  $B$ -mode signature on the polarization of the CMB. A detection of this inflationary gravity wave (IGW) signal "would be a watershed discovery", a quote from the 2010 decadal panel report [1], as it would be our first signal from the epoch of quantum gravity at the beginning of the Universe. The signal would also give strong clues about the nature of inflation, as the  $B$ -mode signal is proportional to the energy scale of inflation through a parameter commonly labeled  $r$ , the tensor-to-scalar ratio. The combination of data from *Planck* and the BICEP/Keck Array give the strongest constraint to date  $r < 0.06$  (95%) [2]. But the measurements have also revealed that emission within our own galaxy is a source of confusion that must be separated with high fidelity before definitive discovery, or stronger upper limits, can be claimed [? ]. For the levels of  $r$  targeted in the next decade, PICO has the frequency coverage and sensitivity to measure and separate sources of foreground confusion and is thus poised to detect or place unprecedented constraints on the physics of inflation. **Its measurements of the spectral index of primordial fluctuations will give the strongest constraints yet on specific models of inflation.**

A few hundred million years after the Big Bang, the neutral hydrogen gas permeating the Universe was reionized by photons emitted by the first luminous sources to have formed. The nature of these sources (e.g., star-forming galaxies or high-redshift quasars) and the exact history of this epoch are key missing links in our understanding of structure formation. Various measurements, including *Planck*'s measurement of the optical depth to reionization  $\tau = 0.054 \pm 0.007$ , have indicated that reionization concluded by  $z \approx 6$ , but its onset at higher redshift is poorly constrained. PICO will yield a breakthrough in this context via a cosmic-variance-limited<sup>1</sup> measurement of  $\tau$ , with  $\sigma\tau = 0.002$ , which can only be directly measured in large-scale CMB polarization fluctuations. The only proven method to date for measuring this signal, which requires exquisite control of systematics and foreground contamination, is a space-based platform.

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<sup>1</sup>The cosmic variance limit is the statistical limit arising from observing a single universe.

Lensing of the CMB photons by structures as they traverse the Universe provides a projected map of all the matter in the universe from the epoch of decoupling until today. The non-zero mass of neutrinos affects the clustering of matter and thus can be inferred from maps of the projected matter distribution. The quantity that can specifically be inferred is the sum of the neutrino masses. The current constraint from the combination of *Planck* and large-scale structure data is  $\sum m_\nu < 0.12$  eV (95%). This is approaching the minimum summed mass allowed in the inverted neutrino hierarchy of  $\approx 0.1$  eV and is within a factor of two of the minimal mass allowed in the normal hierarchy of  $\approx 0.06$  eV. A detection thus appears imminent. However, the precision of determining the neutrino mass scale, using the CMB or *any* other cosmological probe, is limited by knowledge of  $\tau$ , due to the strong degeneracy between  $\tau$  and the amplitude of matter fluctuations. PICO’s map of the projected matter, with signal to noise ratio (SNR) exceeding 500, *and* its own cosmic-variance-limited measurement of  $\tau$  will give a  $4\sigma$  detection of  $\sum m_\nu$  in the normal hierarchy, rising to  $\sim 7\sigma$  for the inverted hierarchy.

The CMB offers a unique window into the thermal history of the universe, from the time of reheating through today. It is during these eras that the matter and radiation that fill the universe were produced and evolved to form the structures observed at low redshifts. Measurements of the CMB on small angular scales are sensitive to the many components that make up the universe including the baryons, cosmic neutrinos, dark matter, and a wide variety of particles motived by extensions of the Standard Model.

The Standard Model of particle physics posits three neutrino families, but it also allows for additional light, relativistic particles, if they existed early enough during the evolution of the Universe. We count the total number light particles thermalized in the early universe using  $N_{\text{eff}}$ . Light particles thermalized in the early universe leave a universal contribution to  $N_{\text{eff}}$  that is sensitive to the freeze-out temperature and the spin of the particle. A mission like PICO holds the promise to reach back to times when the temperature of the universe was orders of magnitude hotter than we have probe today. Such a measurement would shed light on the history of the universe at those very early times and can address important questions about the particles and forces in the Standard Model and Beyond. The history of the universe prior to a few seconds is still largely unexplored observationally and an PICO could reveal important clues to the nature of the fundamental laws and our cosmic origins.

The current measurement of  $N_{\text{eff}} = 2.99 \pm 0.17$  from *Planck* is sensitive to particles thermalized after the QCD phase transitions. Reaching much earlier time is possible with PICO because of much lower noise levels in polarization. Larger sky coverage further improves the statistics and compensates for the lower resolution compared to ground based measurements. These features are advantageous not only for  $N_{\text{eff}}$  for any new physics present in the primary CMB and/or lensing potential. Of particular interest is the nature of dark matter and its interactions, which can be manifest itself in any or all of these probes, depending on the details physics of the dark sector.

the paragraph below covers a lot of ground, but the anticipated impact is not clear. do we want to mention cluster counts? the impact of source counts? Gianfranco thinks pico is unique, but this is not clear. Need to say what PICO will do for these topics - Nick, thoughts? NB: I gave this a shot

Secondary anisotropies in the CMB provide a wealth of information on the growth and evolution of structure in our universe. CMB lensing, the thermal and kinematic Sunyaev-Zel’dovich (SZ) effects, and extragalactic point sources all contribute significantly to the CMB intensity fluctuations on small angular scales (note that lensing is also present in polarization fluctuations). The all-sky, projected mass map reconstructed from CMB lensing that PICO will provide can be correlated with

tracers of large-scale structure to tomographically probe the growth of structure at unprecedented signal-to-noise. The thermal SZ effect provides a map of the integrated free electron pressure along the line of sight, and the peaks of this map trace the locations of all galaxy clusters in the universe. PICO will find all the massive, virialized, galaxy clusters at any redshift. The epoch of reionization imprints information in the statistical moments of the kinematic SZ signal. The combination of these statistical moments with its cosmic variance limited  $\tau$  measurement, PICO will provide information on the nature of the sources responsible for reionization.

Pico will provide a full sky catalog of tens of thousands of extragalactic millimeter and sub-millimeter point sources, which are beacons for active galactic nuclei (in the radio) and dust emission from vigorously star-forming galaxies at  $z \sim 2$  and earlier (in the far-IR). Cosmic magnetism is an outstanding puzzle of fundamental importance to astrophysics. Magnetic fields are ubiquitous, and their evolution is critically interwoven with the dynamics of the universe. Hence, it is crucial to understand their origin and the dynamo processes that must have amplified weak primordial seed fields and maintained their strength across cosmic time [3].

As often in astrophysics, our understanding is rooted in observations of the very local universe: the Milky Way and nearby galaxies. Magnetic fields are observed to be a foremost agent of the Milky Way's ecology. They hold keys for making progress on some exciting issues in the astrophysics of galaxies: the dynamics and energetics of their multiphase interstellar medium (ISM), the efficiency of star formation, the acceleration and propagation of cosmic rays and the impact of feedback on their evolution. Magnetic fields are not only critical for understanding galaxies. The magnetized ISM in the Solar Neighborhood presents a challenge for the investigation of cosmological signals. Cosmological signals of interest, such as CMB B-mode polarization, CMB spectral distortions, and 21cm line emission from cosmic dawn and reionization are obscured by Galactic dust and synchrotron emission that can be orders of magnitude brighter. Discovery of these signals will require understanding and correcting for the magnetized ISM with unprecedented precision.

A broad range of science topics call for progress in our modeling of Galactic magnetic fields, which in turn motivates ambitious efforts to obtain relevant data. As a result, today Galactic magnetism is a dynamic research field, driven by major advances in observational capabilities. PICO will provide a transformative data set to our understanding of this exciting area of exploration.

## 2.2 Science Objectives (17.5 pgs)

The Science Traceability Matrix can be found in Tables 1 and 2.

### 2.2.1 Fundamental Physics

#### Inflation and Gravitational waves

Measurements of the CMB together with Einstein's theory of general relativity imply that the observed density perturbations must have been created long before the CMB was released, and rather remarkably even before the universe became filled with a hot and dense plasma of fundamental particles. Understanding the mechanism generating these perturbations, which evolved to fill the universe with structures, is one of the most important open questions in cosmology.

PICO's precision measurements of temperature and  $E$ -mode polarization anisotropy would provide additional information about the statistical properties of the primordial density perturbations generated during this epoch. In addition, PICO would be exquisitely sensitive to gravitational waves. Any gravitational waves present during recombination leave a faint imprint in the polarization of the CMB, and unlike density perturbations, they not only generate primordial temperature

Table 1: Science Traceability Matrix

| Science Goals from NASA Science Plan   | Science Objectives   | Scientific Measurement Requirements   |  |   | Instrument (single instrument, single mode)  |  | Mission Functional Requirements  |
|--|--|---|--|---|--|--|--|
|  |  | Model Parameters  | Physical Parameters  | Observables   | Functional Requirements  | Projected Performance  |  |
| <i>Explore how the universe began (Inflation)</i>                                      | SO1. Probe the physics of the big bang by detecting the energy scale at which inflation occurred if it is above $4 \times 10^{15}$ GeV, or place an upper limit if it is below (§ 2.2.1) | Tensor-to-scalar ratio $r$ : $\sigma(r) < 5 \times 10^{-5}$ at $r = 0$ ; $r < 10^{-4}$ at 95% confidence level  | CMB polarization $BB$ power spectrum for modes $2 < l < 300$ to cosmic variance limit, and CMB lensing power spectrum for modes $2 < l < 1000$ to cosmic variance limit  | Linear polarization across $60 < v < 300$ GHz over entire sky; Foreground separation requires $20 < v < 800$ GHz                                  | Frequency coverage [for foreground separation]: $v_c$ from 30 to 500 GHz.<br>Frequency resolution:<br>$\Delta v/v_c = 25\%$ .  |  |  |
|  | SO2. Probe the physics of the big bang by excluding classes of potentials as the driving force of inflation (§ 2.2.1, Figure 2)  | Spectral index ( $n_s$ ) and its derivative ( $n_{\text{run}}$ ): $\sigma(n_s) < 0.0015$ ; $\sigma(n_{\text{run}}) < 0.002$   | CMB polarization $BB$ power spectrum for modes $2 < l < 1000$ to cosmic variance limit   |   | Sensitivity: See Table 3.2.<br>Combined instrument weight of $< 0.87 \mu\text{K}_{\text{CMB}} \sqrt{\text{s}}$ .<br>Angular resolution [for delensing and foreground separation]:<br>$\text{FWHM} = 6.2' \times (155 \text{ GHz}/v_c)$ .<br>Sampling rate:<br>$(3/\text{BeamFWHM}) \times (336'/\text{s})$ . | Frequency coverage: See Table 3.2.<br>21 bands with $v_c$ from 21 to 799 GHz.  | Sun-Earth L2 orbit with Sun-Probe-Earth $< 15^\circ$ .<br>5 yr survey with $\geq 95\%$ survey efficiency.  |
| <i>Discover how the universe works (Neutrino mass and <math>N_{\text{eff}}</math>)</i> | SO3. Determine the sum of neutrino masses. (§ 2.2.1, Figure 4)   | Sum of neutrino masses ( $\Sigma m_\nu$ ): $\Sigma m_\nu < 15$ meV with DESI or Euclid; $\Sigma m_\nu < X$ meV alone  | CMB polarization $BB$ power spectrum for modes $2 < l < 4000$ to cosmic variance limit; CMB intensity maps (to give Compton $Y$ map from which we extract clusters)  | Intensity and linear polarization across 60–400 GHz over entire sky   |  | Frequency resolution:<br>$\Delta v/v_c = 25\%$ .<br>Sensitivity: See Table 3.2.<br>Combined instrument weight of $0.43 \mu\text{K}_{\text{CMB}} \sqrt{\text{s}}$ . | Full sky survey: Spin instrument at 1 rpm;<br>Boresight $69^\circ$ off spin axis; Spin axis $26^\circ$ off anti-Sun line, precessing $360^\circ / 10\text{hr}$ . |
|  | SO4. Tightly constrain the thermalized fundamental particle content of the early Universe (§ 2.2.1, Figure 3)  | Number of neutrino effective relativistic degrees of freedom ( $N_{\text{eff}}$ ): $\sigma(N_{\text{eff}}) < 0.03$  | CMB temperature and $EE$ polarization power spectra $2 < l < 4000$ to cosmic variance limit  | Intensity and linear polarization across 60–300 GHz over entire sky   |  | Angular resolution: See Table 3.2.   | Pointing control: Spin axis $60'$ ( $3\sigma$ , radial). Spin $1 \pm 0.1$ rpm ( $3\sigma$ )  |
| <i>Explore how the universe evolved (reionization)</i>                                 | SO5. Distinguish between models that describe the formation of the earliest stars in the universe (§ 2.2.2, Figure 5)  | Optical depth to reionization ( $\tau$ ): $\sigma(\tau) < 0.002$  | CMB polarization $EE$ power spectrum for modes $2 < l < 20$ to cosmic variance limit   | Linear polarization across $60 < v < 300$ GHz over entire sky; Foreground separation enveloped by SO1 and less driving                            | Enveloped by SO1–4, and less driving: Angular resolution $< 1^\circ$ at XX GHz (role of intensity maps at high $\ell$ to be clarified). Combined instrument weight of $< 0.86 \mu\text{K}$ arcmin  | FWHM = $6.2' \times (155 \text{ GHz}/v_c)$ ; $1.1'$ for $v_c = 799$ GHz.<br>Sampling rate: See Table 3.1.<br>$(3/\text{BeamFWHM}) \times (336'/\text{s})$          | Pointing stability: Drift of spin axis $< 1'/1\text{min}$ ( $3\sigma$ , radial); Jitter $< 20''/20$ ms ( $3\sigma$ , radial).                                    |
|  | SO6. Determine if magnetic fields are the dominant cause of low star formation efficiency in our Galaxy. (§ 2.2.3)   |   | The turbulence power spectrum on scales 0.05–100 pc (from cores to diffuse cloud envelopes). Magnetic field strength ( $B$ ) as a function of spatial scale and density. Hydrogen column density. Gas velocity dispersion. | Intensity and linear polarization with $< 1$ pc resolution for thousands of molecular clouds and $< 0.05$ pc for the 10 nearest molecular clouds. | Enveloped by SO1–4, except:<br>Angular resolution: $\leq 1.1'$ (at highest frequency)<br>Sensitivity at 800 GHz: 27.4 kJy/sr   |  | Pointing knowledge (telescope boresight): $10''$ ( $3\sigma$ , each axis) from spacecraft attitude $1''$ ( $3\sigma$ , each axis) final reconstructed            |
| <i>Explore how the universe evolved (Galactic structure and dynamics)</i>              | SO7. Constrain the temperatures and emissivities characterizing Milky Way's interstellar diffuse dust. (§ 2.2.3)   | Intrinsic polarization fractions of the warm and cold components of the diffuse interstellar medium to accuracy better than 2% when averaged over 10 arcmin pixels. Temperatures and spectral indices of the two dust components to an accuracy better than ??% | Fractional polarization and intensity as a function of frequency   | Intensity and linear polarization maps in 12 frequency bands between 108 and 800 GHz.   |  |  | Return and process instrument data: 1.5 Tbits/day (after 4x compression)   |
|  | SO8. Determine the role of energy feedback in the evolution of Milky Way's interstellar medium. (§ 2.2.3)  | Ratio of turbulent energy to magnetic energy (Alfvén Mach number $Ma$ ) on scales 0.03–100 pc   | The turbulence power spectrum on scales 0.03–100 pc in the neutral ISM. Magnetic field strength ( $B$ ) as a function of spatial scale and density. Neutral hydrogen velocity dispersion.                                  | Maps of polarization with $1'$ resolution over the entire sky.  |  |  | Thermally isolate instrument from solar radiation and from spacecraft bus  |

and  $E$ -mode polarization but primordial  $B$ -mode polarization [4, 5]. Any imprint of gravitational waves on degree scale polarization of the CMB detected by PICO would constitute evidence for gravitational waves from the same primordial period that created the density perturbations and open a new window on this early epoch.

Because the dynamics of gravitational waves is essentially unaffected by the plasma physics, they would be a pristine relic left over from the earliest moments of our universe, and their properties would shed light on the mechanism that created the primordial perturbations. Knowledge of the strength of the signal and its statistical properties would transform our understanding of many areas of fundamental physics.

Inflation, a period of nearly exponential expansion of the early universe [6–9], is the leading paradigm explaining the origin of the primordial density perturbations [10–14]. It predicts a nearly scale invariant spectrum of primordial gravitational waves originating from quantum fluctuations [15]. Thus, a detection of these gravitational waves would be the first detection of a phenomenon associated with quantum gravity [16]. Because the spectrum is scale-invariant, one may hope to detect primordial gravitational waves over a wide range of frequencies including, for example, at LIGO or LISA frequencies. However, as a consequence of the expansion of the universe, the energy density in the gravitational waves rapidly dilutes with increasing frequency, and observations of the CMB provide the easiest, and for the foreseeable future only way to detect these gravitational waves.

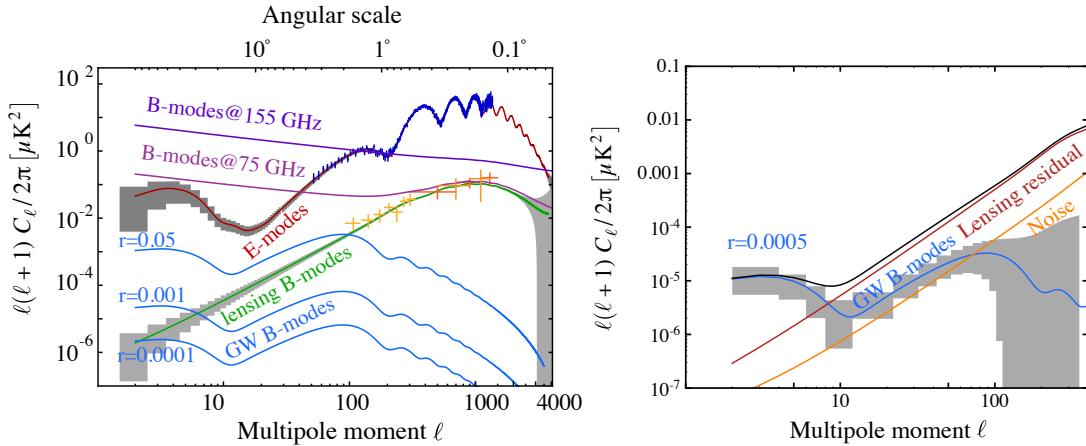


Figure 1: *Left panel:*  $EE$  (red) and lensing  $BB$  (green) angular power spectra and their measurement uncertainties predicted for PICO (gray), as well as the  $BB$  power spectrum produced by IGW with different values of  $r$ . Also shown are measurements of lensing from current experiments (orange) and *Planck* measurements of the  $E$  mode (dark blue) [? ? ]. The  $BB$  spectra of Galactic emission on the cleanest 60% of the sky at 75 and 155 GHz (purple) dominate the cosmological signals except at  $\ell = 1000$  and over a narrow frequency band. *Right panel:* Predicted uncertainties for a detection of primordial gravitational waves with  $r = 0.0005$  for PICO (gray), together with the signal (blue), the instrumental noise (orange), and the lensing residual after internal delensing (red).

The strength of the signal, often quantified by the tensor-to-scalar ratio  $r$ , is a direct measure of the expansion rate of the universe during inflation. Together with the Friedmann equation, this reveals one of the most important characteristics of inflation, its energy scale. PICO’s goal is to detect primordial gravitational waves if inflation occurred at an energy scale of at least  $4 \times$

$10^{15}$  GeV, or equivalently a tensor-to-scalar ratio of  $r = 3 \times 10^{-4}$ . A detection would have profound implications for fundamental physics because it would provide evidence for a new energy scale tantalizingly close to the energy scale associated with grand unified theories, and would allow us to probe physics at energies far beyond the reach of terrestrial colliders.

Even in the absence of a detection PICO’s measurements would contain invaluable information about the early universe. There are two classes of slow-roll inflation that naturally explain the observed value of the spectral index of primordial fluctuations  $n_s$ . The first class is characterized by potentials of the form  $V(\phi) \propto \phi^p$ . This class includes many of the simplest models of inflation, some of which have already been strongly disfavored by existing observations; see the right panel of Figure 1. If the constraints on the spectral index tighten by about a factor 2 with the central value unchanged, and the upper limits on  $r$  improve by an order of magnitude, this class would be ruled out. Select models in this class are shown as blue lines in Figure 2

The second class is characterized by potentials that approach a constant as a function of field value, either like a power law or exponentially. Two representative examples in this class are shown as the green and gray bands in Figure 2. This class also include  $R^2$  inflation, which predicts a tensor-to-scalar ratio of  $r \sim 0.003$ . All models in this class with a characteristic scale in the potential that is larger than the Planck scale predict a tensor-to-scalar ratio of  $r \gtrsim 0.001$ . Different values of characteristic scales are indicated by the darker lines in Figure 2. Many microphysical models in this class possess a characteristic scale that is super-Planckian, but there are models such as the Goncharov-Linde model with a somewhat smaller characteristic scale that predict a tensor-to-scalar ratio of  $r \sim 4 \times 10^{-4}$  [17].

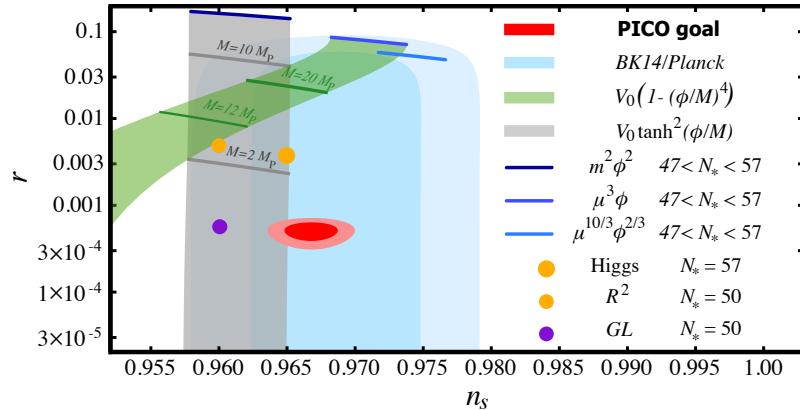


Figure 2: Current 1 and  $2\sigma$  limits on  $r$  and  $n_s$  (blue) and forecasted constraints for a fiducial model with  $r = 0.0005$  for PICO. Also shown are predictions for the selected models of inflation discussed in the text.

In the absence of a detection, PICO would limit the amount of gravitational waves to  $r < 10^{-4}$  at 95% CL and would exclude all these models.

Let us now take a closer look at the signal. As shown in Figure 1, it has two contributions, one on degree angular scales or multipoles of  $\ell \sim 80$ , typically referred to as the recombination peak, and another contribution for multipoles of  $\ell \lesssim 10$  from the epoch of reionization.

No sub-orbital experiment has yet measured modes at  $\ell < 40$ . The temporal stability, absence of atmospheric noise, and full sky coverage offered by a satellite like PICO make it the most suitable instrument to reach these lowest multipoles.

The contribution from reionization is expected to be strongest relative to the contributions from instrumental noise and ‘lensing’  $B$ -modes created from  $E$ -modes by the deflection of photons by large scale structure on their way to us from the last scattering surface.

When the tensor to scalar ratio  $r \simeq 0.01$ , the  $BB$  lensing power spectrum and the primordial  $BB$  power spectrum are comparable around the recombination peak. For lower levels of  $r$ , the lensing  $B$ -mode dominates, but the  $B$ -mode maps can be ‘delensed’ [18, 19]. The effect of lensing on  $E$  and  $B$  maps can be determined and undone if these maps are measured with few arcmin resolution and sufficient depth. Forecasts for PICO show that at least 73% of the lensing  $B$ -mode power can be removed for the baseline configuration, after accounting for foreground subtraction. 80% will be removed if the foregrounds do not degrade the inherent SNR significantly, rising to 85% for the CBE configuration. Delensing will improve PICO’s determination of  $r$  by a factor 5 – 6. We emphasize that PICO will be relying on its own data to conduct the delensing and foreground cleaning, thus avoiding reduced efficacy arising from the need to cross-calibrate experiments, identify common observing areas on the sky, not having frequency band coverage at the appropriate resolution to remove foregrounds, or from other systematic uncertainties.

Models of the early universe differ in their predictions for the scalar spectral index  $n_s$  and its scale dependence, often referred to as the running of the spectral index  $n_{\text{run}}$ . With its high resolution and low noise levels, PICO will improve the constraints on  $n_s$  and  $n_{\text{run}}$  by a factor of about two. In addition, PICO will probe the statistical properties of the primordial fluctuations over a wide range of scales with exquisite precision and improve constraints on departures from Gaussianity by a factor 2 – 3. By cross-correlating the lensing map with large-scale structure data from LSST it may even be possible to reach a theoretically important threshold (see, e.g. [20] and references therein) and constrain local non-Gaussianity to better than  $\sigma(f_{NL}) = 1$ . This is discussed in more detail in 2.2.2.

### Fundamental Particles: Light relics, Dark Matter, and Neutrinos

- **Light Relics** In the inflationary paradigm, the universe was reheated to temperatures of at least 10 MeV and perhaps as high as  $10^{12}$  GeV. At these high temperatures, even very weakly interacting or very massive particles, such as those arising in extensions of the Standard Model of particle physics, can be produced in large abundances [21, 22]. As the universe expands and cools, the particles fall out of equilibrium, leaving observable signatures in the CMB power spectra. Through these effects the CMB is a sensitive probe of neutrino and of other particles’ properties.

One particularly compelling target is the effective number of light relic particle species  $N_{\text{eff}}$ . The canonical value with three neutrino families is  $N_{\text{eff}} = 3.046$ . Additional light particles contribute a universal change to  $N_{\text{eff}}$  that is a function only of the decoupling temperature and the effective degrees of freedom of the particle,  $g$ . Furthermore, the range of  $\Delta N_{\text{eff}}$  is quite restricted even for widely varying decoupling temperatures  $T_F$  with the range  $0.027 g \leq \Delta N_{\text{eff}} \leq 0.07 g$  corresponding to decoupling at higher temperatures during post-inflation reheating ( $0.027g$ ) to lower temperatures shortly prior to the QCD phase transition ( $0.07g$ ).

Performance forecasts for  $N_{\text{eff}}$  are shown in Figure 3. For an experiment like PICO, which has sufficient resolution to reach cosmic-variance-limited measurement of  $EE$  up to  $\ell = 2300$ , the two additional most important parameters for improving constraints are the fraction of sky observed  $f_{\text{sky}}$  and the noise. Achieving both larger  $f_{\text{sky}}$  and lower noise are strengths of PICO compared to other platforms. The PICO requirement is to constrain  $\Delta N_{\text{eff}} < 0.06$  at 95%. The corresponding improvement in reach in  $T_F$  is shown in the right panel of Figure 3. The large improvement over Planck ( $\Delta N_{\text{eff}} < 0.28$ , 95%) corresponds to a factor of 400 improvement in the limit on the decoupling temperature for any particle with spin.

Many light relics of the early universe are not stable. They decay, leaving faint evidence of their



Figure 3: *Left:*  $N_{\text{eff}}$  uncertainty as a function of noise and sky fraction. The resolution assumed is 5'. Vertical lines denote the expected performance of the baseline mission. *Right:* Reach in the freeze-out temperature for various species, given a measurement of  $\Delta N_{\text{eff}}$ . We see an exclusion of  $\Delta N_{\text{eff}} < 0.06$  is a nearly two order of magnitude improvement over Planck and SO. The vertical lines are normalized to the  $T_f$  for a single vector particle.

past existence on other tracers. The relics with sufficiently long lifetime to survive few minutes, past the epoch of light element synthesis, leave a signature on the helium fraction  $Y_p$ . If they decay by the time of recombination, their existence through this period is best measured through the ratio of  $N_{\text{eff}}$  to  $Y_p$ . At both CBE and Baseline sensitivity, PICO can simultaneously measure  $N_{\text{eff}}$  and  $Y_p$  with  $\sigma(N_{\text{eff}}) = 0.08$  and  $\sigma(Y_p) = 0.005$ . Alternatively, PICO can measure  $Y_p$  at fixed  $N_{\text{eff}}$  with  $\sigma(Y_p) = 0.002$  to independently determine the primordial helium abundance with the same precision as astrophysical measurements. The combination of these measurements is a sensitive test of physics between big bang nucleosynthesis and recombination.

• **Dark Matter** Cosmological measurements have already confirmed the existence of one relic that lies beyond the Standard Model: dark matter. For a conventional WIMP candidate, the CMB places very stringent constraints on its properties through the signature of its annihilation [23–25]. Most of this information is in the  $EE$  power spectrum at  $50 < \ell < 300$  and is well-measured by *Planck* and will approach the cosmic variance limit with existing ground based surveys [26, 27]. An entirely complementary way to probe dark matter is to search for evidence of its interactions with other species in cosmological data. Since a lower mass translates to a higher number density of scattering centers, CMB is particularly sensitive to the low-mass regime and is sensitive to large, nuclear-scale cross sections.

Interactions between dark matter and protons in the early universe creates a drag force between the two cosmological fluids, damping acoustic oscillations and suppressing power in density perturbations on small scales. As a result, the CMB temperature, polarization, and lensing power spectra are suppressed at high multipoles relative to a universe without such drag forces. This effect has been used to search for evidence of dark matter-proton scattering over a range of masses and couplings, and to provide consistency tests of dark matter in the context of the anomalous 21-cm signal reported by the EDGES collaboration [28–36]

In Figure 4, we present current and projected upper limits on the dark matter-proton interaction cross section as a function of dark matter mass, for a spin-independent velocity-independent scattering (chosen as our fiducial model). Regions above the curves are excluded at the 95% con-

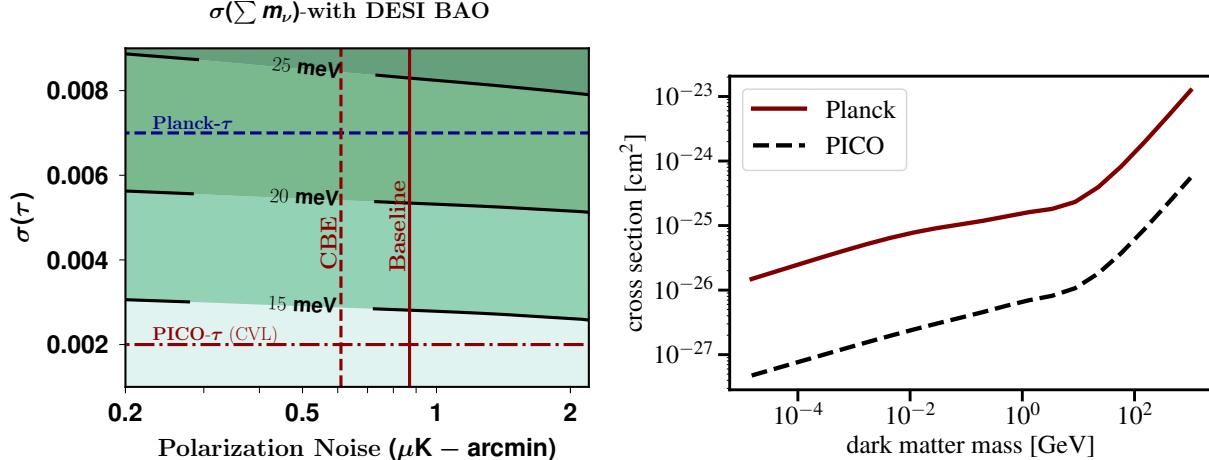


Figure 4: *Left:* Forecasts for the sum of neutrino masses uncertainty, including DESI BAO, as a function of noise and the uncertainty in the measurement of  $\tau$ , for 0.7 sky fraction. The upper blue dashed line is the current *Planck* limit; the lower grey dashed line is the limit from cosmic variance limited measurement of  $\tau$ . *Right:* Upper limits on DM-proton interaction cross section as a function of DM mass, for a spin-independent velocity-independent scattering. Areas above the curves are excluded at 95% confidence-level. Shown are the current limits from *Planck*([31]) and a forecast for PICO.

fidence level. We compare current limits obtained from *Planck* (from [31]) with projections for PICO sensitivity. We note that PICO can deliver a substantial improvement over the current limits, across the entire dark matter mass range considered. Most of the constraining power in case of PICO (and ground-based next-generation measurements with similar white-noise levels) comes from the measurement of  $C_\ell^{\phi\phi}$ .

• **Neutrino Mass** The origin and structure of the neutrino masses is one of the great outstanding questions about the nature of the Standard Model particles. Measurements of neutrinos in the lab have revealed much about the mass differences and mixing angles. Cosmology offers a measurement of the sum of the neutrino masses  $\sum m_\nu$  through the gravitational influence of the non-relativistic cosmic neutrinos. The measurement of  $N_{\text{eff}} = 2.99 \pm 0.17$  [37] already confirms the existence of these neutrinos at  $> 10\sigma$  and their mass implies that they will contribute to the matter density at low redshifts. The best current mass constraint arises from a combination of *Planck* and BOSS barion acoustic oscillations (BAO) giving  $\sum m_\nu < 0.12$  eV (95%) [37].

Cosmological measurements are primarily sensitive to the suppression of power on small scales after the neutrinos become non-relativistic, which can be measured via CMB lensing or weak lensing in a galaxy survey. However, these measurements are limited by our knowledge of the amplitude of the primordial fluctuation power spectrum,  $A_s$ . In practice, CMB observations most directly constrain  $A_s e^{-2\tau}$  and thus do not provide a high precision measurement of either  $A_s$  or  $\tau$  separately.

Although many surveys hope to detect  $\sum m_\nu$ , any detection of the minimum value expected from particle physics  $\sum m_\nu = 58$  meV at more than  $2\sigma$  will require a better measurement of  $\tau$ . The best constraints on  $\tau$  come from  $E$  modes with  $\ell < 20$  which require measurements over the largest angular scales. To date, the only proven method for such a measurement is from space. The current limit of  $\sigma(\tau) = 0.007$  is from *Planck* [38]. Forecasts for a CMB measurement of  $\sum m_\nu$  using the lensing  $B$  mode [39] are shown in Figure 3. With the current uncertainty in  $\tau$  one

is limited to  $\sigma(\sum m_\nu) \gtrsim 25$  meV (including DESI BAO); no other survey or cosmological probe would improve this constraint. But PICO will reach the cosmic variance limit of  $\tau \sim 0.002$  and will therefore reach  $\sigma(\sum m_\nu) < 15$  meV when combined with DESI’s measurements of baryon acoustic oscillations [40]. Robustly detecting neutrino mass at  $> 3\sigma$  in any cosmological setting is only possible with an improved measurement of  $\tau$  like the one achievable with PICO. The measurement would give  $\sum m_\nu > 0$  at greater than  $4\sigma$  or would exclude the inverted hierarchy ( $\sum m_\nu > 100$  meV) at 95% confidence, depending on the central value of the measurement. Lab-based measurement could determine the hierarchy before PICO but only cosmology can measure  $\sum m_\nu$ .

### Fundamental Fields: Primordial Magnetic Fields and Cosmic Birefringence

- **Primordial Magnetic Fields** One of the long standing puzzles in astrophysics is the origin of 1–10  $\mu$ G strength galactic magnetic fields [41]. Producing such fields through a dynamo mechanism would require a primordial seed field [42]. Moreover,  $\mu$ G strength fields have been observed in proto-galaxies that are too young to have gone through the number of revolutions necessary for the dynamo to work. A primordial magnetic field (PMF), present at the time of galaxy formation, could provide the seed or even eliminate the need for the dynamo altogether. Specifically, a  $\sim 0.1$  nG field in the intergalactic plasma would be adiabatically compressed in the collapse to form a  $\sim 1$   $\mu$ G galactic field [43]. PMFs could have been generated in the aftermath of phase transitions in the early universe [44], during inflation [45, 46], or at the end of inflation [47]. A detection of PMF would be a major discovery, signalling physics beyond standard models of particle physics and cosmology, and constraints on PMF offer a valuable tool for discriminating among different theories of the early universe [48–50]. While the PMF would be sustained by the primordial plasma well beyond recombination, with signatures at low redshifts, only seeing them in CMB would conclusively prove their primordial, as opposed to an astrophysical, origin.

The signature of PMF is detectable through Faraday rotation [51], which converts  $E$  modes into  $B$  modes, and through generating signatures in the  $BB$  power spectrum at high  $\ell$  [52]. The current CMB bounds on PMF strength are  $B_{1\text{Mpc}} < 1.2$  nG at 95% CL for the scale-invariant PMF spectrum [53]. PICO’s sensitivity and resolution would allow to probe PMFs as weak as 0.1 nG ( $1\sigma$ ), a limit that already includes the effects of imperfect lensing subtraction, galactic foregrounds [54–56], and other systematic effects. It would, nevertheless, be an important improvement that will conclusively rule out the purely primordial (no dynamo) origin of the largest galactic magnetic fields.

- **Cosmic Birefringence** The simplest model for late-time acceleration of the universe is with a slowly-evolving scalar field – the quintessence [57]. Such a field generically couples to electromagnetism through a Chern Simons-like term, and causes linear polarization of photons propagating cosmological distances to rotate. This is known as cosmic birefringence [57]. The birefringence converts primordial  $E$  mode into  $B$  mode. It thus produces parity-violating  $TB$  and  $EB$  cross-correlations whose magnitude depends on the statistical properties of the rotation field in the sky [58, 59]. There are no theoretical predictions for the level of birefringence, but if observed, it would be evidence for physics beyond the standard model and a potential probe of dark-energy microphysics [59–61]. Using the sensitivity of only the 155 GHz, PICO will improve current constraints on cosmic birefringence (from POLARBEAR [62]) by a factor of 300. The constraints will be even stronger when including all frequency bands.

## 2.2.2 Cosmic Structure Formation and Evolution

**The Formation of the First Luminous Sources** The reionization of the Universe imprints multiple signals in the temperature and polarization of the CMB. In polarization, the most important signal is an enhancement in the  $EE$  power spectrum at large angular scales  $\ell \lesssim 20$ ; see Figure 1. This signal gives a direct measurement of the optical depth to the reionization epoch  $\tau$ , and thus to the mean redshift of reionization  $z_{re}$ , with very little degeneracy with other cosmological parameters; see Figure 5. The mean redshift of reionization  $z_{re}$  (when 50% of the cosmic volume was reionized) depends sensitively on the nature of the ionizing sources. It is currently unknown whether star-forming galaxies or more exotic sources such as supermassive black holes drove the reionization process. What was the mean free path of ionizing photons during this epoch? What was the efficiency with which such photons were produced by ionizing sources? What were the masses and environments of the dark matter halos that hosted the sources? These properties all affect  $z_{re}$ . Furthermore, the detailed shape of the low- $\ell$   $E$ -mode power spectrum is sensitive to the reionization history itself (i.e.,  $d\tau/dz$ ), and will provide information beyond that captured in  $\tau$  alone. For example, it has been argued that *Planck* data show evidence for an extended tail of reionization out to  $z \approx 15-20$  [63]. A cosmic-variance-limited measurement of the large-scale  $E$  modes, as obtained by PICO, will settle this question.

Large-scale  $EE$  power spectrum measurements are a unique and crucial observable for many aspects of cosmology, particularly the growth of structure. If measurements of  $\tau$  are not improved beyond the current uncertainties from *Planck*, inference of several new signals of cosmological physics will be severely hindered. A canonical example is the inference of the sum of the neutrino mass (see page 9), but any cosmological inference related to the growth of structure will be affected to some extent, including constraints on dark energy and modified gravity from weak lensing, cluster counts, and similar structure probes. PICO is the ideal experiment to resolve this issue. Its noise level and frequency coverage permit a cosmic-variance-limited constraint on  $\tau$ , i.e.,  $\sigma(\tau) \approx 0.002$ , which we have verified with explicit forecasts including separation of foregrounds.

In temperature, an important imprint of reionization is that sourced at small angular scales by the “patchy” kinematic Sunyaev-Zel’dovich (kSZ) effect, due to the peculiar velocities of free electron bubbles around ionizing sources. Measurements of the small- scale kSZ power spectrum, with instruments that have higher resolution than PICO, can give constraints on the duration of reionization  $\Delta z_{re}$  [? ]. Fig. 5 presents forecasts for reionization constraints in the  $z_{re} - \Delta z_{re}$  parameter space obtained from PICO’s measurement of  $\tau$  in combination with ground-based Stage-III CMB experiments measurements of the kSZ power spectrum. The PICO measurement of  $\tau$  is essential for breaking degeneracies and allowing simultaneous, precise constraints to be placed on both the mean redshift and duration of reionization. The Figure also shows curves of constant efficiency of production of ionizing photons in the sources, and of intergalactic medium opacity. These are two parameters that quantify models of reionization. The curves shown are illustrative; families of models, that would be represented by parallel ‘source efficiency’ and ‘IGM Opacity’ lines, are allowed by current data. PICO’s data will give simultaneous constraints on these physical parameters, yielding important information on the nature of the first luminous sources. For example, galaxies or quasars predict significantly different values for these parameters.

In addition to these signals, reionization also leaves specific non-Gaussian signatures in the CMB. In particular, patchy reionization induces non-trivial 4-point functions in both temperature [?] and polarization [? ]. The temperature 4-point function can be used to separate reion-

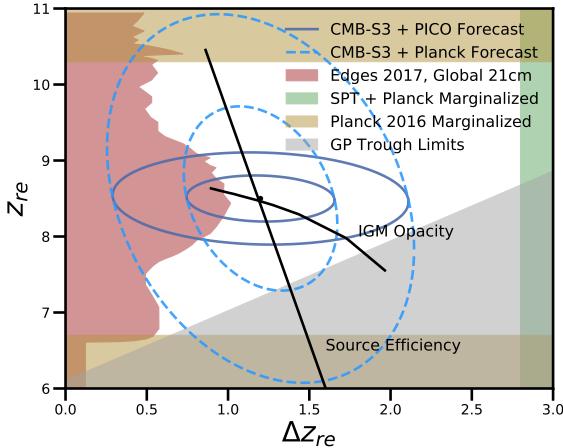


Figure 5: Contours of 1 and  $2\sigma$  constraints on the mean redshift and duration of reionization using PICO and CMB-S3 data (solid dark blue), and comparison with *Planck* and CMB-S3 (dash light blue). Source Efficiency and IGM Opacity (dark lines) are two physical parameters controlling the reionization process in current models. The PICO measurements together with higher resolution data of the kSZ effect will significantly the range of models allowed. We also include other constraints from *Planck*, EDGES, the Gunn-Peterson (GP) trough , and *Planck*+ the South Pole Telescope (SPT) [37? ? ? ].

ization and late-time kSZ contributions. Combinations of temperature and polarization data can be used to build quadratic estimators for reconstruction of the patchy  $\tau$  field, analogous to CMB lensing reconstruction. These estimators generally require high angular resolution, but also rely on foreground-cleaned CMB maps. Thus, while PICO alone may not enable high SNR reconstructions, its high-frequency bands — which have better than 2 arcmin resolution and cover frequencies that are not suitable for observations from the ground — will enable these estimators to be robustly applied to ground-based CMB data sets, a strong example of ground-space complementarity.

**Structure Formation via Gravitational Lensing** Matter between us and the last-scattering surface deflects the path of photons through gravitational lensing, imprinting the 3-dimensional matter distribution across the volume of the universe onto the CMB maps. The specific quantity being mapped by the data is the projected gravitational potential  $\phi$  that is lensing the photons. From the lensing map, which receives contributions from all redshifts between us and the CMB with the peak of the distribution at  $z \simeq 2$ , we infer the angular power spectrum  $C_L^{\phi\phi}$ . Both the temperature and polarization maps of the CMB, and by extension the angular power spectra, are affected by lensing.

*Planck*'s  $\phi$  map had SNR of  $\sim 1$  per  $L$  mode over a narrow range of scales,  $30 < L < 50$ . PICO's map would represent true mapping, with  $\text{SNR} \gg 1$  per each mode down to scales of approximately ten arcminutes ( $L \sim 1000$ ); see Figure 6. On smaller scales, the map will still contain statistical information. While *Planck* had an SNR of 40 integrated across the  $C_L^{\phi\phi}$  power spectrum [64], the PICO combination of resolution, sensitivity, and sky coverage enables a measurement with SNR of 638 and 737 for the baseline and CBE configurations, respectively. When accounting for possible foreground contamination, its broad frequency coverage leads to a reduction of SNR of less than 20%; see Figure 6.

The value of the reconstructed lensing map is immense, as has already been demonstrated with the much lower SNR map from *Planck*. The unprecedented constraints on neutrino mass, discussed in page 9, are a direct result of this deep map. Tomographic cross-correlations of the lensing map with samples of galaxies and quasars will yield constraints on structure formation. The measurements will constrain the properties of quasars and other high-redshift astrophysics, e.g., a precise determination of the quasar bias (and hence host halo mass) as a function of their properties, such as (non-)obscuration. The map will be cross-correlated with other large scale tracers to probe fundamental physics. For instance, one can use correlations between large scale structure tracers with different clustering bias factors to effectively cancel cosmic variance [65, 66]

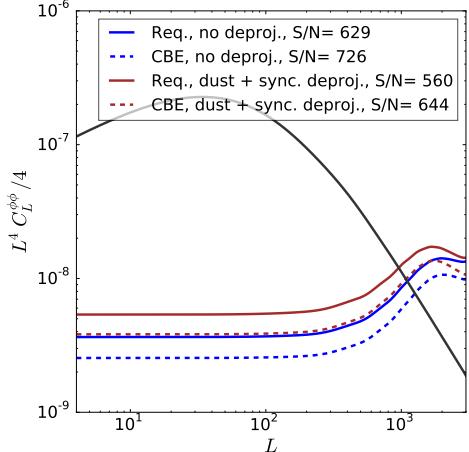


Figure 6: The theoretically predicted lensing power spectrum  $C_L^{\phi\phi}$  (black) and forecasted PICO noise levels, with (red), and without (blue) deprojection, that is removal, of foregrounds. PICO will make a map of  $\phi$  at angular scales where the noise is below the signal.

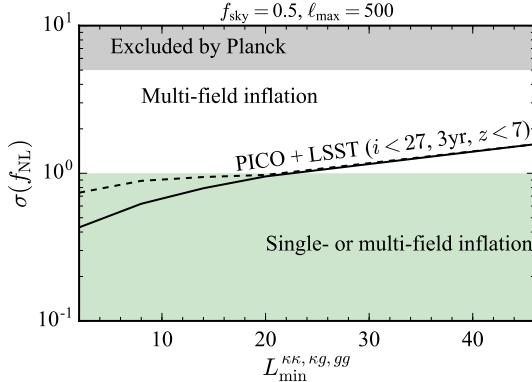


Figure 7: Forecasted sensitivity to the parameter describing primordial non-Gaussianity of the local type for the PICO CMB lensing map together with three years of the LSST survey, as a function of the minimal multipole used in the analysis. A value of  $\sigma(f_{NL}) \simeq 1$  is a well-motivated theoretical target.

and constrain physics that affects the biasing of objects on large scales, such as primordial local non-Gaussianity [67]. In Fig. 7 we show the expected constraints for the CMB lensing field as reconstructed with PICO, in cross correlation with three years of the LSST survey. It can be seen that depending on the minimal multipole that can be used in the cross correlation, which is uncertain in both LSST and the PICO lensing map, the well-motivated theory target of  $\sigma(f_{NL}) \simeq 1$  [20] can be within reach. Values of  $f_{NL}$  at or above this level are a generic prediction of multi-field inflationary models.

Using the same cross-correlation techniques, it is also possible to constrain the evolution of the amplitude of structure as a function of redshift. Figure 8 shows constraints on the amplitude of linear structure in several redshift bins. This is a model-independent representation of the structure growth constraints; these measurements will yield constraints on dark energy or modified gravity, in the context of specific models. The measurements can also be used for a neutrino mass constraint that is complementary to and competitive with that inferred from the CMB lensing auto-power spectrum described earlier.

Lensing will also be used to weigh dark matter halos hosting galaxies, groups, and clusters of galaxies. Calibrating the masses of galaxy clusters is the most uncertain and crucial step in the cluster cosmology program, in which CMB lensing has already begun to play an important role. In this approach, known as CMB halo lensing, we focus on the small-scale effects of gravitational lensing around these objects [68–70]. The technique holds great potential for measuring halo masses out to high redshifts where gravitational lensing of galaxies (i.e., gravitational shear) no longer works because of the lack of background sources.

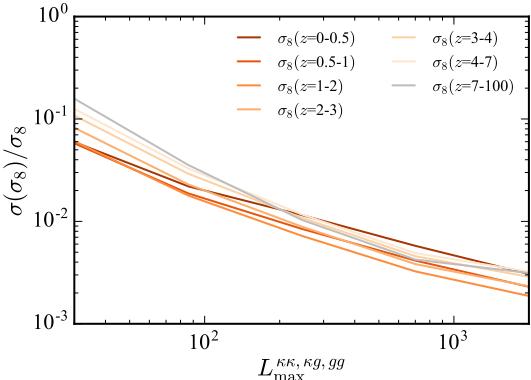


Figure 8: Forecasted sensitivity to the parameter describing the amplitude of structure in various redshift bins, as a function of the maximal multipole used in the analysis. Percent-level constraints on these parameters allow for stringent tests of physics beyond  $\Lambda$ CDM that modify the rate of growth of structure.

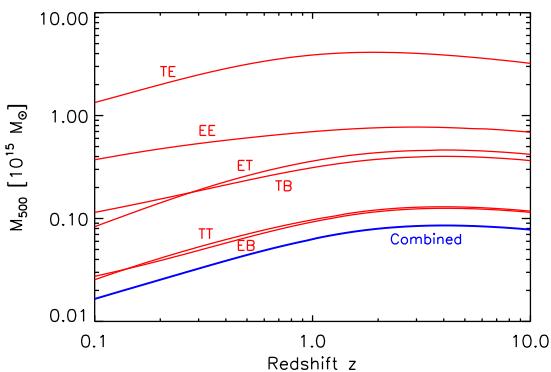


Figure 9: PICO sensitivity for CMB halo lensing. Curves for different CMB signal correlations give the one-sigma sensitivity of an optimal mass filter [71]. The curves are flat at high redshift, demonstrating the essential property that CMB halo lensing can be applied over a very wide redshift range, and well beyond the  $z = 2$  limit of the figure. For PICO, the  $EB$  and  $TT$  estimators are roughly equivalent, offering important cross-validation of measurements because the systematics are very different for temperature and polarization.

This is illustrated in Fig. 9, which shows the mass sensitivity of PICO using a spatial filter optimized for extracting the mass of halos [71]. The curves give the one-sigma noise in a mass measurement through the filter as a function of redshift. Their flattening at high redshift reflects the fact that CMB lensing is sensitive over a broad range of redshifts, extending well beyond the limit of  $z = 2$  of the figure. We see that PICO can measure the mass of individual low-mass clusters ( $\sim 10^{14} M_\odot$ ) over a wide redshift range, and by stacking we can determine the mean mass of much smaller halos, including those hosting individual galaxies.

Halo lensing will enable calibration of the galaxy cluster mass scale, which is critical for our cosmological analysis of PICO cluster counts, as mentioned above. It also gives a unique tool for measuring the relation between galaxies and their dark matter halos during the key epochs of cosmic star formation at  $z \geq 2$ , not reachable by other means. This will provide valuable insight into the role of environment on galaxy formation during the rise to and fall from the peak of cosmic star formation at  $z \sim 2$ . From a complementarity perspective, the high-resolution, high-frequency PICO channels will play an essential role in cleaning foregrounds for high-resolution ground-based halo lensing measurements at lower frequencies, particularly those derived from the temperature-based estimator, which is most contaminated by foregrounds.

**Galaxy Formation via the Sunyaev-Zel'dovich (SZ) Effects** Not all CMB photons propagate through the universe freely; about 6% are Thomson-scattered by free electrons in the intergalactic medium (IGM) and intracluster medium (ICM). These scattering events leave a measurable imprint on CMB temperature fluctuations, which thereby contain a wealth of information about the growth of structures and the thermodynamic history of baryons. A fraction of these photons are responsible for the thermal and kinetic Sunyaev–Zel'dovich effects (tSZ and kSZ) [? ? ]. The amplitudes of the

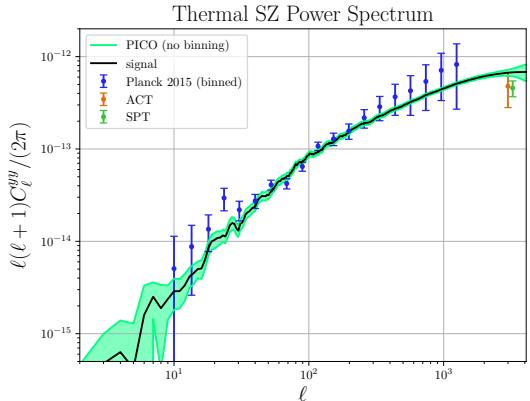


Figure 10: Constraints on the tSZ power spectrum from PICO and current data. A simulated tSZ power spectrum (black) is constrained at each multipole by PICO’s data ( $1\sigma$ , green). Binning in  $\ell$ , would increase the SNR beyond the calculated value of 1270, which is already nearly 100 times larger than *Planck* (blue). We include current measurements by SPT and ACT, two ground-based programs [? ? ]

tSZ and kSZ signals are proportional to the integrated electron pressure and momentum along the line of sight, respectively. They thus contain information about the thermodynamic properties of the IGM and ICM, which are highly sensitive to astrophysical ‘feedback’. Feedback is the process of energy injection into the IGM and ICM from accreting supermassive black holes, supernovae, stellar winds, and other sources. The tSZ effect will be used to measure ensemble statistics of galaxy clusters, which contain cosmological information, as well as to provide uniform cluster samples for galaxy formation studies in dense environments.

- **Galaxy Clusters** Galaxy clusters found via the tSZ effect provide a well-defined sample with a simple-to-model selection function. Sample of clusters such as these are straightforward to use for cosmological inference and studies of galaxy evolution in dense environments. The tSZ-selected sample from PICO will provide all clusters with masses above **double check**  $\sim 3 \times 10^{14} M_{\odot}$  (defined with respect to the radius within which the average density reaches 200 times the critical) out to high redshifts **which high redshift?**, as long as the clusters have started to virialize. We forecast that PICO will find  $\sim 150,000$  galaxy clusters, assuming the cosmological parameters from *Planck* and applying a galaxy mask, using only 70% of the sky. With redshifts provided by optical surveys and infrared follow-up observations, the PICO tSZ-selected cluster sample will be an excellent cosmological probe, with mass calibration provided by CMB halo lensing described above and optical weak lensing for clusters with  $z < 1.5$ . **why is it an excellent cosmological probe? didn’t we want to add forecasts on parameters such as neutrino mass, and dark energy? wasn’t there a point about a sample that is not restricted to a particular location in the sky?**

- **Compton-y map and tSZ auto-power spectrum** In addition to finding individual clusters, multifrequency CMB data also allow the reconstruction of full-sky maps of the tSZ signal. These are called ‘Compton-y maps’. With its extremely low noise and broad frequency coverage, which is essential for separating out other signals, PICO will yield a definitive Compton-y map over the full sky, with high SNR down to angular scales of a few arcminutes. We quantify this expectation by reconstructing the Compton-y field using the needlet internal linear combination (NILC) algorithm [? ] applied to sky simulations generated with the *Planck* sky model, with maps at all PICO frequencies (with appropriate noise added). The error bars on the reconstructed tSZ power spectrum are shown in Fig. 10, in comparison to current measurements. The total SNR is 1270 for the PICO CBE configuration, with the PICO baseline configuration only  $\approx 10\%$  lower. This is nearly two orders of magnitude higher SNR compared to *Planck*, which has already provided data with much higher SNR compared to ground-based experiments.

Extremely strong constraints on models of astrophysical feedback will be obtained from the

analysis of the PICO  $y$ -map, both from its auto-power spectrum and from cross-correlations with galaxy, group, cluster, and quasar samples. Like the CMB lensing map described above, the legacy value of the PICO  $y$ -map will be immense. As an example, we forecast the detection of cross-correlations between the PICO  $y$ -map and galaxy weak lensing maps constructed from LSST and WFIRST data. Considering the LSST “gold” sample with a source density of 26 galaxies/arcmin<sup>2</sup> covering 40% of the sky, we forecast a detection of the tSZ – weak lensing cross-correlation with S/N = 3000. At this immense significance, the signal can be broken down into dozens of tomographic redshift bins, yielding a precise breakdown of the evolution of thermal pressure over cosmic time. For PICO and WFIRST (assuming 45 galaxies/arcmin<sup>2</sup> covering 5.3% of the sky), we forecast S/N = 1100 for the tSZ – weak lensing cross-correlation. The WFIRST galaxy sample extends to higher redshift, and thus this high-S/N measurement will allow the evolution of the thermal gas pressure to be probed to  $z \approx 2$  and beyond, the peak of the cosmic star formation history. These transformative measurements will revolutionize our understanding of galaxy formation and evolution by distinguishing between models of feedback energy injection at high significance. Additional cross-correlations of the PICO  $y$ -map with quasar samples, filament catalogs, and other large-scale structure tracers will further demonstrate its immense legacy value, providing valuable information on baryonic physics that is complementary to inferences from the lensing cross-correlations described earlier.

### 2.2.3 Galactic Structure and Star Formation

Observations of Galactic polarization are a highlight and a lasting legacy of the *Planck* space mission. Spectacular images combining the intensity of dust with the texture derived from polarization data have received world-wide attention and have become part of the general scientific culture [72]. Beyond their popular impact, the *Planck* polarization maps represented an immense step forward for Galactic astrophysics [73]. We expect an even greater leap forward from PICO based on the higher angular resolution dust polarization images obtained with the balloon experiment BLAST-Pol. PICO will provide all-sky maps of dust polarization at higher resolution than BLASTPol and with significantly higher sensitivity than *Planck*. Such a data set can only be obtained from a space mission. *Planck* made hundreds of thousands of measurements of magnetic field orientation across the sky; with PICO we expect  $\sim 1.5 \times 10^8$  independent measurements in the 799 GHz band alone. The data will complement a rich array of polarization observations including stellar polarization surveys to be combined with Gaia astrometry and synchrotron observations measuring Faraday rotation at radio wavelengths with the Square Kilometer Array and its precursors. In this section we focus on two key crucial Galactic science measurements that only PICO can obtain.

(1) *Testing Composition Models of Interstellar Dust:* The analysis of the PICO data will involve the spectral characterization of Galactic polarization. This aspect of the data analysis will continue to update and test models of dust emission and of grain alignment, which are of interest for the interpretation of dust polarization data at large.

(2) *Determining how magnetic fields affect the processes of molecular cloud and star formation:* By virtue of the strong dynamical coupling of dust and gas and the systematic alignment of dust grains with magnetic fields, dust polarization probes magnetic fields in the cold and warm neutral phases of the ISM and in molecular gas. If the magnetic fields are sufficiently strong, they can prevent the gravitational collapse of gas across magnetic field lines and can slow down or limit the processes of star and planet formation. In the diffuse ISM the neutral phase of the interstellar medium contains the bulk of the gas mass and of its turbulent kinetic energy.

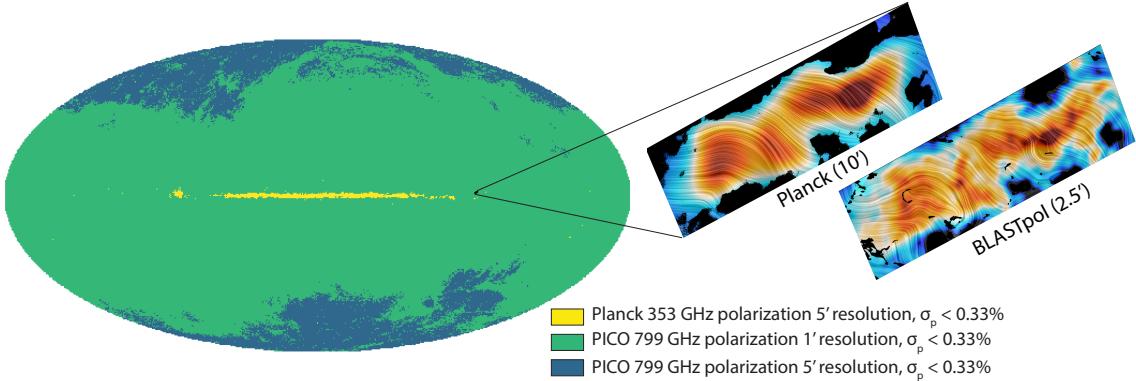


Figure 11: At 799 GHz, the PICO Baseline mission will map nearly the entire sky at 1' resolution. As an example of the current state-of-the-art, Planck (10') and BLASTpol (2.5') maps of the Vela C region are shown [74]. These observations will enable PICO to characterize magnetized turbulence from the diffuse ISM down to dense star forming cores.

## Dust Physics

Strong extinction features at 9.7 and 18  $\mu\text{m}$  indicate much interstellar dust is in the form of amorphous silicates while features at 2175 Å, 3.3  $\mu\text{m}$ , and 3.4  $\mu\text{m}$  attest to abundant hydrocarbons. It is unknown, however, whether the silicate and carbonaceous materials coexist on the same grains or whether they are segregated into distinct grain populations. If there are indeed multiple grain species, this will induce additional challenges for modeling the emission from interstellar dust in both total intensity and polarization at levels relevant for B-mode science [75].

Spectropolarimetry of dust extinction features reveals robust polarization in the 9.7  $\mu\text{m}$  silicate feature [e.g., 76], indicating that the silicate grains are aligned with the interstellar magnetic field. In contrast, searches for polarization in the 3.4  $\mu\text{m}$  carbonaceous feature have yielded only upper limits, even along sightlines where silicate polarization is observed [77, 78]. These data suggest that most of the silicate and carbonaceous materials do not exist on the same grains. However, these studies are limited to only a few highly-extincted sightlines that may not typify the diffuse ISM.

At odds with the spectropolarimetric evidence from dust extinction, current measurements of the polarization fraction of the far-infrared dust emission with *Planck* [79] and BLASTPol [80] betray little to no frequency dependence, as would be expected if two components with distinct polarization properties were contributing to the total emission. However, current uncertainties are relatively large and the data with  $\nu > 353 \text{ GHz}$  are from high density sightlines that may not be representative of the diffuse ISM. With great polarization sensitivity even in diffuse regions, PICO will provide a definitive test of the two component paradigm.

To assess PICO's ability to discriminate quantitatively, we employ the analytic two component dust model of [81] which provided a better fit to IRAS and *Planck* data than one component models.

Applying the noise estimates from PICO, 1000 simulations were run for different combinations of polarization fractions of the two components in this model. Only frequency channels 107 GHz and above were used, and the simulated data were binned to the 7.9' beam of PICO's 107 GHz channel. Based on the variance of the simulation results, PICO can determine the intrinsic polarization fractions of the two components to a precision of 1-2%. PICO will therefore be able to validate or reject state-of-the-art dust models [e.g. 82, Hensley & Draine, in prep] and test for

the presence of additional grain species with distinct polarization signatures, such as magnetic nanoparticles [83].

### Are Magnetic Fields Responsible For Low Star Formation Efficiency?

Stars form out of dense, gravitationally unstable regions within molecular gas clouds. The efficiency of this conversion from molecular gas to stars is very low, due to regulation from supersonic turbulent gas motions, magnetic fields, and feedback from young stars [84]. Magnetic fields may play an important role in slowing the process of star formation by inhibiting movement of gas in the direction perpendicular to the field lines. Observations to date suggest that the outer envelopes of clouds can be supported against gravity by magnetic fields, but in dense cores gravity tends to dominate, and so these dense structures can collapse to form stars [85].

On larger scales, the formation of gravitationally unstable clouds is regulated by the flow of diffuse material into the molecular phase, a process that is mediated by magnetized turbulence in the low-density ISM. Structure formation in the diffuse ISM is poorly understood, but as a precursor to star formation it is crucial to understand what drives molecular cloud formation. Recent observations suggest that the structure of the diffuse medium is highly anisotropic, and strongly coupled to the local magnetic field [86–89].

However, the degree to which magnetic fields affect the formation of molecular clouds as well as stars within these clouds is poorly constrained, in large part due to the difficulty of making detailed maps of magnetic fields in the interstellar medium.

• **Formation of Stars within Magnetized Molecular Clouds** With full-sky coverage and a best resolution of  $1.1'$ , PICO will be able to map all molecular clouds with better than 1 pc resolution, out to a distance of 3.4 kpc. Extrapolating from the Bolocam Galactic Plane Survey [BGPS, 90], PICO is expected to make highly detailed magnetic field maps of over 2,000 molecular clouds with thousands to hundreds of thousands of independent measurements per cloud.

Our goal is to constrain both the strength of the magnetic field,  $B$ , within these clouds, as well as the energetic importance of the field compared to self-gravity (parameterized by the mass-to-flux ratio  $\mu$ ) and turbulence (parameterized by the Alfvén Mach number  $M_A$ ) as a function of density. To measure these quantities we will apply a series of established polarization analysis techniques: (1) characterizing the relative orientation of cloud structures and the magnetic field [91–94]; (2) making probability distributions functions of polarization measurables [74, 95]; (3) comparing between the magnetic field and velocity gradient directions [96–98]; and (4) measuring the angular dispersion of the magnetic field [99–102]. By applying all four techniques to both PICO observations and synthetic polarization maps made from “observing” numerical simulations of star formation, we will quantitatively compare theory and observations. PICO’s large number of frequency bands will be used to better model the temperature and polarization efficiency of the cloud dust [103], which can then be used to generate more realistic generation of synthetic observations from simulations for comparison with PICO observations [104]. We can then compare the observed magnetization levels derived from the PICO observations to the levels of turbulence derived from molecular gas surveys (e.g.: Ellsworth-Bowers et al. 90, Miville-Deschénes et al. 105), and the efficiency of star formation, measured from near and far-IR observations of dense cores and protostars with *Herschel*, *Spitzer*, and *WISE*.

*PICO’s ability to map thousands of clouds is not possible with any other current or proposed polarimeter. Planck*, for example, was only able to map 10 nearby clouds to a similar level of detail [94]. This large sample of clouds is crucial because dust polarization observations are sensitive to

only the magnetic field projected on the plane of the sky, and therefore polarization maps will look very different for molecular clouds observed at different viewing angles. *By observing thousands of molecular clouds PICO will determine the role of magnetic fields in star formation as a function of cloud age and mass.* ga

• **Formation of Magnetized Molecular Clouds from The Diffuse Interstellar Medium** Structure formation in the diffuse ISM is a key area of study motivating observations across the electromagnetic spectrum. PICO’s observations will complement recently completed high dynamic range neutral hydrogen (HI) surveys, such as HI4PI [106] and GALFA-HI [107], as well as planned surveys of interstellar gas, most prominently with the Square Kilometer Array (SKA) and its pathfinders. One of the open questions in diffuse structure formation is how gas flows within and between phases of the ISM. A planned all-sky absorption line survey with SKA-1 will increase the number of measurements of the ISM gas temperature by several orders of magnitude [108]. Quantitative comparisons of the ISM temperature distribution from SKA-1 and estimates of the magnetic field strength and coherence length scale from PICO will elucidate the role of the magnetic field in ISM phase transitions.

Despite its importance, a comprehensive understanding of the magnetized diffuse ISM is challenging because of its diverse composition, its sheer expanse, and the multi-scale nature of the physics that shapes it. How are matter and energy exchanged between the diffuse and dense media? This question must be addressed by measuring the properties of the magnetic field over many orders of magnitude in column density. PICO is unique in its ability to do this in the diffuse ISM. *Planck* achieved measurements of the diffuse sky at 60' resolution, resulting in  $\sim$ 30,000 independent measurements of the magnetic field direction in the diffuse ISM. With 1.1' resolution PICO will expand the number of independent polarization measurements in the diffuse ISM to  $\sim$ 86,000,000. This will allow us to robustly characterize turbulent properties like  $M_A$  across a previously unexplored regime of parameter space.

## Galactic Legacy Science

PICO will also produce legacy datasets that will revolutionize our understanding of how magnetic fields influence physical processes ranging from planet formation to galaxy evolution. For 10 nearby clouds ( $d < 500$  pc) PICO will resolve magnetic fields on the crucial 0.1 pc size scale associated with dense cores and filaments, and observe how the magnetic fields on these scales directly influence the formation structure of cores. By comparing the orientation of the core-scale magnetic field with respect to the orientation and sizes of protoplanetary disks, PICO will directly test whether there is evidence that magnetic breaking inhibits the growth of protoplanetary disks [109, 110].

On larger scales, PICO’s tens of millions of independent measurements of magnetic field orientation will allow us to directly probe magnetized turbulence and study how magnetic fields are generated through a combination of turbulence and large scale gas motions [111]. Key processes in the diffuse ISM, including heat transport [112], streaming of cosmic rays [113], and magnetic reconnection [114] are dramatically dependent on the level of magnetization.

Finally, PICO observations will create detailed magnetic field maps of approximately 70 nearby galaxies, with more than 100 measurements of magnetic field direction per galaxy. These observations will be used to study the turbulence on galactic scales, determine whether the magnetic fields of the Milky Way in the diffuse ISM are consistent with other galaxies, and directly study how interaction between large scale magnetic fields, turbulence, and feedback from previous generations

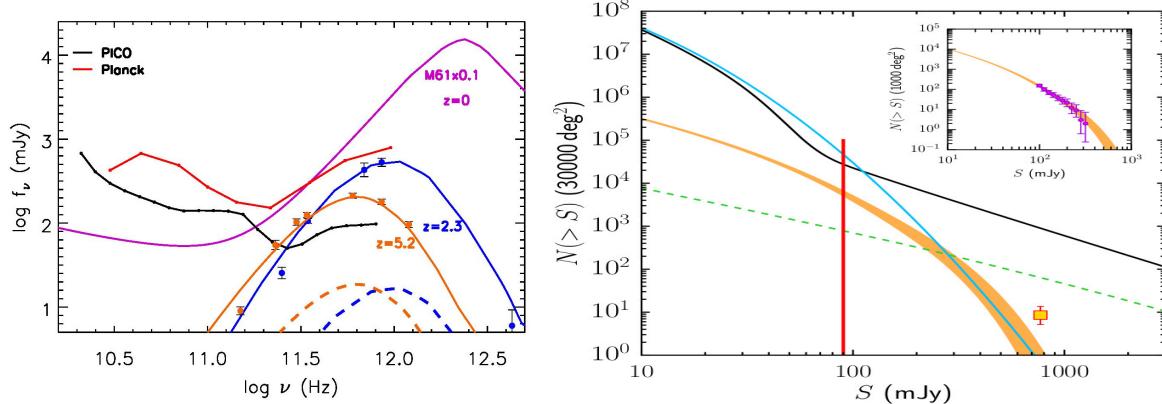


Figure 12: **Left panel.** Examples of spectral energy distributions (SEDs) of dusty star-forming galaxies detectable by PICO, compared with its point source detection limits (broken black line) and with the *Planck* 90% completeness limits (broken red line; [115]). PICO will detect nearby galaxies, like M 61, whose SED was scaled down by a factor of 10, and high- $z$  redshift strongly lensed galaxies, like SMM J2133-0102 at  $z = 2.3$  [116] and HLSJ091828.6 + 514223 at  $z = 5.2$  [117]. The dashed lines show the SEDs of the two galaxies corrected for the lensing magnification. **Right panel.** Integral counts at  $500 \mu\text{m}$  (600 GHz). The solid black line and the orange band show counts of unlensed and strongly lensed star-forming galaxies, respectively, based on fits of *Herschel* counts. Above the PICO detection limit (vertical red line) the counts of unlensed galaxies have an Euclidean slope, indicative of low  $z$ . The yellow square on the bottom right shows the *Planck* count of strongly lensed galaxies, while their *Herschel* counts [118] are shown in the inset. The solid blue line shows the counts of proto-clusters of galaxies yielded by the model of ref. [119].

of star formation affect galaxy evolution and star formation efficiency.

For each of the science cases described, PICO will provide the crucial large number statistics afforded by its all-sky coverage and will bridge the spatial scales covered by its predecessor *Planck* and high resolution ground based telescopes like ALMA.

### 2.3 Cosmological Legacy Surveys

#### 2.3.1 Early phases of galaxy evolution

PICO will have a crucial role in providing answers to major, still open issues on galaxy formation and evolution. Which are the main physical mechanisms shaping the galaxy properties [120, 121]: in situ processes, interactions, mergers, or cold flows from the intergalactic medium? How do feedback processes work? To settle these issues we need direct information on the structure and the dynamics of high- $z$  galaxies. But these are compact, with typical sizes of 1–2 kpc [122]), corresponding to angular sizes of 0.1–0.2 arcsec at  $z \simeq 2$ –3. Thus they are hardly resolved even by ALMA and by HST. If they are resolved, high enough SNR per resolution element are achieved only for the brightest galaxies, which are probably not representative of the general population.

Strong gravitational lensing provides a solution to these problems. PICO will detect galaxies whose flux densities are boosted by large factors; see the right panel of Fig. 12. Since lensing conserves the surface brightness, the effective angular size is stretched on average by a factor  $\mu^{1/2}$ , where  $\mu$  is the gravitational magnification, thus substantially increasing the resolving power. A spectacular example are ALMA observations of the strongly lensed galaxy PLCK\_G244.8+54.9 at  $z \simeq 3.0$  with  $\mu \simeq 30$  [123]. ALMA observation with a  $0.1''$  resolution reached the astounding spatial resolution of  $\simeq 60\text{pc}$ , substantially smaller than the size of Galactic giant molecular

clouds. Other high- $z$  galaxies spatially resolved thanks to gravitational lensing, with less extreme magnifications, are reported by Dye et al. [124], and others [125, 126].

Cañameras et al. [123] have also obtained CO spectroscopy, measuring the kinematics of the molecular gas with an uncertainty of 40–50 km/s. This spectral resolution makes possible a direct investigation of massive outflows driven by AGN feedback at high  $z$ . In this way Spilker et al. [127] were able to detect a fast (800 km/s) molecular outflow due to feedback in a strongly lensed galaxy at  $z = 5.3$ . The outflow carries mass at a rate close to the SFR and can thus remove a large fraction of the gas available for star-formation.

*Herschel* surveys have demonstrated that, at the PICO detection limit at  $\simeq 500 \mu\text{m}$  (600 GHz), about 25% of all detected extragalactic sources are strongly lensed; for comparison, at optical/near-IR and radio wavelengths, were intensive searches have been carried out for many years, the yield is of only 0.1%, i.e. more than two orders of magnitude lower [128]. To add to the extraordinary sub-mm bonanza, the selection of strongly lensed galaxies detected by sub-mm surveys is extremely easy because of their peculiar sub-mm colors – see the left panel of Fig. 12 – resulting in a selection efficiency close to 100% [129].

A straightforward extrapolation of the *Herschel* counts to the much larger area covered by PICO shows that its surveys will yield  $\sim 4,500$  strongly lensed galaxies with a redshift distribution peaking at  $2 \lesssim z \lesssim 3$  [118] but extending up to  $z > 5$ ; see the left panel of Fig. 12. If objects like the  $z = 5.2$  strongly lensed galaxy HLSJ091828.6 + 514223 exist at higher redshifts, they will be detectable by PICO up to  $z > 10$ .

An intensive high spectral and spatial resolution follow up campaign of such a large sample will be challenging but also extremely rewarding since it will allow a giant leap forward towards the understanding of the processes driving early galaxy evolution, in addition to opening many other exciting prospects both on the astrophysical and on the cosmological side (cf., e.g., ref. [128]). The PICO all-sky surveys will select the brightest objects in the sky, maximizing the efficiency of the effort.

### 2.3.2 Early phases of cluster evolution

PICO will open a new window for the investigation of early phases of cluster evolution, when their member galaxies were actively star forming but the hot IGM was not necessarily in place. In this phase, traditional approaches to cluster detection (X-ray and SZ surveys, searches for galaxy red sequences) work only for the more evolved objects; indeed these methods have yielded only a handful of confirmed proto-clusters at  $z \gtrsim 1.5$  [130]<sup>2</sup>. *Planck* has demonstrated the power of low-resolution surveys for the study of large-scale structure [131] but its resolution was too poor to detect individual proto-clusters [119]. Studies of the high- $z$  2-point correlation function [92, 119] and *Herschel* images of the few sub-mm bright protoclusters detected so far, at  $z$  of up to 4 [132–134], all of which will be detected by PICO, indicate sizes of  $\simeq 1'$  for the cluster cores, nicely matching the PICO FWHM at the highest frequencies.

PICO will detect many tens of thousands of these objects – this is the blue line in the right-hand panel of Fig. 12 – as peaks in its sub-mm maps, in addition to the evolved ones, detected by the SZ effect. This will constitute a real breakthrough in the observational validation of the formation history of the most massive dark matter halos, traced by clusters, a crucial test of models for structure formation. Follow-up observations will characterize the properties of member galaxies,

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<sup>2</sup>More high- $z$  proto-clusters have been found targeting the environment of tracers of very massive halos, such as radio-galaxies, QSOs, sub-mm galaxies. These searches are however obviously biased.

Table 2: Legacy Science

| Catalog                      | Impact   | Science  |
|------------------------------|--|--|
| 1. Proto-Clusters            | Discover 3000 <sup>a</sup> mm/sub-mm proto-clusters distributed over the sky and across redshift.<br>Current knowledge: <i>Planck</i> data expected to yield a few tens.   | Probe the earliest observable dust-enshrouded galaxy clusters to determine the initial stages of their formation and evolution.  |
| 2. Strongly Lensed Galaxies  | Discover 3000 <sup>a</sup> highly magnified dusty galaxies across redshift.<br>Current knowledge: 13 sources confirmed in <i>Planck</i> data; few hundred candidates in <i>Herschel</i> , SPT and ACT data. <sup>c</sup> | Learn about dark matter sub-structure in the lensing galaxies; probe star formation history in high- $z$ dust-enshrouded galaxies, a population in which star formation history cannot be probed in any other way.                 |
| 3. High- $z$ Galaxy Clusters | Find 1000 <sup>a</sup> mm/submm emitting clusters at $1 < z < 1.5$ and $\sim 20$ at $z > 2$ .<br>Current knowledge: <i>Planck</i> and <i>Herschel</i> identified mm/sub-mm emission of $\sim 100$ known sources.         | Probe star formation history at high $z$ and in dust-enshrouded environments.  |
| 4. Polarized Point Sources   | Detect 4000 <sup>a,b</sup> radio and dusty galaxies in polarization.<br>Current knowledge: <a href="#">Shaul checking w/ Gianfranco</a>  | Determine the structure of magnetic fields in dusty galaxies, and the mechanism for relativistic jet formation in radio-loud galaxies; Determine the importance of polarized sources as a foreground for CMB polarization science. |

<sup>a</sup> Confusion (not noise) limited

<sup>b</sup> Noise and confusion limited

<sup>c</sup> Many *Planck* candidates are not expected to be real because of confusion. *Herschel*, SPT, and ACT source candidates are expected to be real. PICO source candidates will mostly be real.

probing the galaxy evolution in dense environments and shedding light on the complex physical processes driving it.

### 2.3.3 Additional products of PICO surveys

PICO will also yield a complete census of cold dust, available to sustain star formation in the nearby universe, by detecting tens of thousands galaxies mostly at  $z \lesssim 0.1$ . Its statistics will allow us to investigate the distribution of such dust as a function of galaxy properties (morphology, stellar mass, etc.).

Moreover, PICO will increase by orders of magnitude the number of blazars selected at sub-mm wavelengths and will determine the SEDs of many hundreds of them up to 800 GHz and up to  $z > 5$ . Blazar searches are the most effective way to sample the most massive BHs at high  $z$  because of the Doppler boosting of their flux densities. Its surveys of the largely unexplored mm/sub-mm spectral region will also offer the possibility to discover new transient sources [135] or events, such as blazar outbursts.

PICO will also make a giant leap forward in the determination of polarization properties of both radio sources and of dusty galaxies over a frequency range where ground based surveys are impractical or impossible. Thanks to its high sensitivity, it will detect in polarization both populations over a substantial flux density range, determining directly, for the first time, number counts in polarized flux density and allowing an accurate correction for their contamination of CMB maps.

| Characteristic                             | Ground   | Balloon   | Space                   |
|--|--|---|-------------------------|
| Sky coverage                               | Partial from single site   | Partial from single flight                                      | Full                    |
| Frequency coverage                         | 70 GHz inaccessible <sup>a</sup><br>$\nu \geq 300$ GHz unusable<br>limited atmospheric windows | 70 GHz inaccessible <sup>a</sup><br>otherwise, almost unlimited | Unrestricted            |
| Angular resolution at 150 GHz <sup>b</sup> | 1.5' with 6 m telescope  | 6' with 1.5 m telescope   | 6' with 1.5 m telescope |
| Detector Noise                             | $\geq xx$ microK rt(s);  | $\geq xx$ microK rt(s)  | $xx$ microK rt(s)       |
| Integration time                           | Unlimited  | Weeks to a Month  | Continuous, for years   |
| Accessibility, repairability               | Good   | None. Multiple flights possible.                                | None                    |

<sup>a</sup> 70 GHz is the frequency at which Galactic emissions are expected to have a minimum for sky coverage of  $\sim 50\%$  or more.

<sup>b</sup> We give representative approximate telescope aperture values. Significantly larger apertures for balloons and in space result in higher mass, volume, and cost.

Table 3: Relative characteristics of ground, balloon, and space platforms for experiments in the CMB bands.

## 2.4 Complementarity with Other Surveys and with Sub-Orbital Measurements

### 2.4.1 Complementarity with Astrophysical Surveys in the 2020s

PICO has strong complementarity with forthcoming surveys. Here we summarize areas of synergy that have been mentioned in a number of earlier sections.

There is no known way to achieve any cosmological constraint on the sum of the neutrino mass  $\sigma(\sum m_\nu) < 25$  meV without improving *Planck*'s measurement of the optical depth  $\tau$ . In particular, this applies to all methods that rely on comparing low-redshift structures with the amplitude of the CMB at high redshift, such as galaxy clustering, weak lensing, or cluster counts. PICO therefore complements all efforts that probe the late time structure of the Universe; combining PICO with these low-redshift observations extends the scientific reach of all these experiments well beyond what they could achieve on their own.

Reconstructing the CMB lensing  $\phi$  map on very large angular scales,  $L < 20$ , requires exquisite control of systematic uncertainties over a large sky fraction, with sufficient angular resolution to perform the lensing reconstruction, and with breadth in frequency band to robustly separate Galactic emissions (see Section 2.5). PICO will provide these, complementing ground-based CMB lensing reconstructions that typically observe a smaller sky fraction, with a smaller number of frequency bands, and without access to the largest angular scales. As discussed in page ??, PICO will robustly measure the lensing signal with a power spectrum SNR larger than 10 *per mode* on very large scales. Such high-significance CMB lensing measurements on the very largest scales will be useful when combined with measurements of galaxy clustering, from example from LSST, to search for local primordial non-Gaussianity via its scale-dependent effect on galaxy bias; see Section ???. While such a measurement would ultimately be limited by limitations of the LSST data on the very largest scales, space based observations of galaxy clustering with Euclid or SPHEREx may enable further improvements.

what about wfirst? jwst? DESI? CIB?

### 2.4.2 Complementarity with Sub-Orbital Measurements

Since the first CMB measurements, more than 50 years ago, important observations have been made from the ground, from balloons, and from space. Each of the CMB satellites flown to date - COBE, WMAP, and *Planck*- has relied crucially on technologies and techniques that were first proved on ground and balloon flights, making these also crucial to the success of PICO. The phenomenal success of, and the immense science outcomes produced by, past space missions is a direct consequence of their relative advantages, as listed in Table 3. In every respect, with the exception of integration time exceeding several years, space has the advantage. These advantages used to come with higher relative costs. However, with the advent of massive ground-based experiments this balance shifts; the costs for a CMB experiment planned for the next decade are squarely

within the cost window of this Probe. We can thus point to the following general guidelines for the next decade.

When the entire sky is needed, as for fluctuations on the largest angular scales, space is by far the most suitable platform, and for the search for the IGW signal it is absolutely necessary. When broad frequency coverage is needed, space will very likely be required to reach the ultimate limits set by astronomical foregrounds. As Figures 1 and 13 demonstrate Galactic emission overwhelm the IGW signal on the largest angular scales, and they are dominant even at high  $\ell$  potentially limiting the process of delensing that is necessary for reaching levels of  $r \lesssim 0.001$ . The stability offered in space can not be matched on any other platform and translates to superb control of systematic uncertainties. There is a broad consensus within the CMB community that for levels of  $r \lesssim 0.001$  the challenges in the measurement are the ability to control systematic uncertainties and to remove Galactic emissions; modern focal plane arrays, like the one employed by PICO have ample raw sensitivity. The PICO  $r$  goal of  $10^{-4}$  is beyond the reach of ground observations. However, for science requiring higher angular resolution, such as observations of galaxy clusters with  $\sim 1$  arcmin resolution at 150 GHz, the ground has a clear advantage. An appropriately large aperture on the ground will also provide high resolution information at lower frequencies, which may be important for separating Galactic emissions at high  $\ell$ . A recommended plan for the next decade is therefore to pursue a space mission, and complement it with an aggressive ground program that will overlap in  $\ell$  space, and will add science at the highest angular resolution, beyond the reach of a space mission.

Balloon observations have been exceedingly valuable in the past. They co-lead discoveries of the temperature anisotropy and polarization, provided proving grounds for the technologies enabling the success of COBE, WMAP and *Planck*, and trained the scientists that then led NASA’s space missions. There are specific areas for which balloon missions can continue to play an important role, despite their inherently limited observing time. Balloon payload can access frequency bands above 280 GHz; currently there are no plans for any ground program to conduct observations at higher frequencies. These frequency bands will provide important, and perhaps critical information about polarized emission by Galactic dust, a foreground that is currently known to limit knowledge of the CMB signals. With flights above 99% of the atmosphere, balloon-borne observations are free from the noise induced by atmospheric turbulence, making them good platforms for observations of the low  $\ell$  multipoles, and for characterizing foregrounds on these very large angular scales. And balloon-platforms continue to be an excellent arena for training the scientists of tomorrow.

## 2.5 Signal Separation

**Component separation? Signal separation? Foreground removal? need to be consistent.** Diffuse Milky Way emissions dominate the sky’s polarized intensity on the largest angular scales; see Figure 13. **also reference to Figure 1?** Even in the cleanest, smaller patches of the sky, far from the galactic plane and thus relatively low in galactic emissions, their levels are expected to dominate the CMB inflationary *BB* signal for  $r \lesssim 0.01$ , and overwhelm it for  $r \lesssim 0.001$ ; see Figure 13. Foreground separation together with control of systematic uncertainties are *the* challenges facing any next decade experiment attempting to reach these levels of constraints on  $r$ .

The challenge would be easily surmountable if the Galactic emissions were already precisely characterized, or were known to have simple, fittable spectral emission laws; But neither is true. Until recently, the spectrum of Galactic synchrotron emission, arising from free electrons spiral-

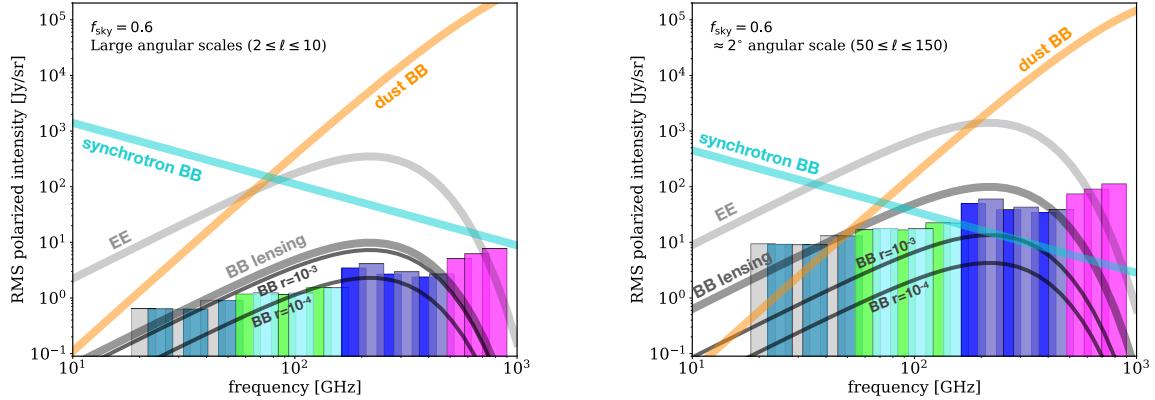


Figure 13: Polarization  $B$  modes of Galactic synchrotron and dust, compared to CMB polarization  $E$ -modes and  $B$ -modes of different origin, for two values of the tensor-to-scalar ratio  $r$ . The location and sensitivity of the PICO frequency channels is shown as vertical bands. Left: integrated r.m.s. emission on the largest angular scales ( $2 \leq \ell \leq 10$ ), corresponding to the reionization peak; Right, integrated r.m.s. emission on  $\simeq 2^\circ$  angular scales ( $50 \leq \ell \leq 150$ ), corresponding to the expected recombination peak in CMB primordial  $B$  modes.

ing around Galactic magnetic fields, was modeled as a power law  $I_{\text{sync}} \propto v^\alpha$ , with  $\alpha \simeq -1$  (in brightness units). The spectrum of Galactic dust emission, arising from emission by Galactic dust grains, was modeled as  $I_{\text{dust}} \propto v^\beta B_v(T_{\text{dust}})$ , where  $\beta \simeq 1.6$ ,  $T_{\text{dust}} \simeq 19$  K, and  $B_v(T)$  is the Planck function; this is referred to as ‘modified black body emission’. **the values are averages across the sky?** An experiment that has 6 frequency bands could determine the three emission parameters as well as the three amplitudes corresponding to that of dust, synchrotron, and the CMB. However, WMAP and Planck observations have shown that neither emission law is universal and that spectral parameters vary with the region of sky **is that true? is there evidence from Planck; add references** (thus that the values given above are valid as averages across the sky). Also, while both emission laws are well-motivated phenomenological descriptions, the fundamental physics of emissions from grains of different material, sizes and temperatures, and of electrons spiraling around magnetic fields implies that these laws are not expected to be exact, nor universal.

We know that we don’t know enough about synchrotron and dust emission. We know even less about the polarization level of ‘anomalous microwave emission’, an excess of dust emission at frequencies between 10 and 100 GHz, and of infra-red sources. Depending on reasonable levels of polarization assumed their contributions to the total polarized signal may be appreciable or negligible (for  $r \lesssim 0.001$ ) [? ].

Faced with these uncertainties, but also with the opportunity provided by a platform that can host a broad range of frequencies - ground-based experiments are limited to several atmospheric windows and to frequencies of less than 300 GHz - PICO is designed with 21 frequency bands between 20 and 800 GHz; see Figure 13 and Table **which in the instrument section**. This is the broadest frequency lever arm proposed by any imaging instrument to characterize and enable separation of Galactic emissions.

Foreground uncertainties, and the level of fidelity required in their characterization, also compel a transition in the way we assess and forecast the performance of a future experiment. The ‘gold standard’ in the community has become to construct sky maps that are constrained by available

data, but otherwise have a mixtures of foreground properties, observe these maps just like a realistic experiment will do, and then apply foreground separation techniques to separate the Galactic and CMB emissions. This approach is distinct from past techniques that used analytic calculations to estimate the efficacy of foreground separation, or others in which the simulated sky map inherently assumed specific Galactic emission models, or made other simplifying assumptions.

**now need to connect to the next paragraphs** As we show below, the PICO broad frequency coverage, coupled with high sensitivity enables studying the sky using the PICO data themselves rather than assuming what the sky is, and fitting our models to these assumptions.

### 2.5.1 PICO Component Separation Methodology

**Sky Maps** For assessing the efficacy of component separation with PICO we used 8 different full sky models. All models were consistent with and constrained by available data and uncertainties from WMAP and *Planck*. The range of models included one test case that had a simple Gaussian realizations of synchrotron and dust at the level observed in the BICEP2 field and then scaling in frequency with a single modified blackbody emission, to models in which the spectral parameters of foregrounds vary across the sky and along the line of sight, anomalous microwave emission **is up to 2%** polarized, dust polarization rotates slightly as a function of frequency because of projection effects, or the dust spectral energy distribution departs from a simple modified blackbody. All foreground maps are generated at native resolution of ?? arcmin pixels [? ]. They are generated using PySM and/or PSM codes [? ]. Distinctly different realizations of the sky are allowed by current data, as demonstrated by Figure ???. More details of the models are available at [? ].

For each of the 8 models we added CMB signals in both intensity and polarization matching a  $\Lambda$ CDM universe. The *BB*-lensing signal matched the level of 85% delensing forecasted for PICO. Each of these sky models had 100 different realization of the PICO CBE noise levels; 50 realizations had no IGW signal and 50 others had a level of  $r = 0.003$ .

**Component Separation** The sky models were analyzed with a variety of techniques which are based on two broad categories: correlation methods, which exploit the fact that foreground emission is strongly correlated from frequency to frequency, but uncorrelated with the CMB, and parametric methods, which model the sky emission using specific (parametric) emission laws, and use spectral fits in independent pixels or sky regions to infer the amplitude and spectral parameters of each of the components in the sky. Correlation methods include SEVEM [? ], and variants of the independent linear combination (ILC) algorithm, such as the needlet-space ILC (NILC) and a version generalised to multidimensional components (GNILC). Parametric methods include the Commander algorithm and **the X-forecast method?**.

### 2.5.2 Results

### 2.5.3 Discussion

## 2.6 Systematic Uncertainties

Some of the PICO science goals attempt to detect extremely faint signals. The most ambitious one is to reach the nanoK-level signals characterizing an inflationary gravity wave with  $r \lesssim 0.001$ . **be more explicit about the level of the  $\ell = 80$  peak** It has long been recognized that exquisite control of systematic uncertainties will be required for any experiment attempting to reach these levels, and it is widely accepted that the stability provided aboard a space platform makes it best suited to control systematic uncertainties compared to other platforms. This is one of the most

compelling reasons to observe the CMB from space. As WMAP<sup>?</sup> and *Planck* demonstrated, the L2 environment offers excellent stability as well as the ability to observe large fractions of the sky on many time scales without interference from the Sun, Earth, or Moon. The redundancy of observations allows the checking of consistency of results and an improved ability to calibrate and to correct systematic errors in post-processing analysis. **there are several arguments here lumped in one sentence: suggest to make more explicit to PICO, and tie to the scan strategy. Also need to fold in the issue of 1/f (and absence of HWP.)**

A rich literature investigates the types of systematic errors due to the environment, the instrumentation, observation strategies, and data analysis that confound the polarization measurement by creating a bias or an increased variance[136–138]. Every measurement to date has reached a systematic error limit, and have advanced many sophisticated techniques to mitigate systematics, finding both new technological solutions and new analysis techniques. As an example, the BICEP’s systematics limited it to  $r=0.1$ [139] while through additional effort within the program, BICEP2 achieved a systematics limit of  $r=6 \times 10^{-3}$ [140]). In the near term, the ground based and suborbital CMB community will continue to develop new techniques in handling systematics, particularly in developing the CMB-S4 project.

All prior on-orbit measurements of CMB polarization were limited by systematic errors until an in-depth study of the systematics was performed and the post-processing data analysis suppressed them[38, 141, 142]. Particularly we note Fig. 3 of the Planck legacy paper which indicates Planck’s systematic error limits on the polarization power spectral measurements. Recently studied space missions, such as EPIC-IM, LiteBird and *CORE*, have placed systematic error mitigation at the forefront of the case for their mission and have developed tools and strategies for estimating and mitigating these[143–145].

Systematics are coupled with the spacecraft scan strategy, and the details of the data analysis pipeline. Thus, end-to-end simulation of the experiment is an essential tool, including realistic instabilities and non-idealities of the spacecraft, telescope, instrument and folding in data post-processing techniques used to mitigate the effects.

## 2.6.1 List of Systematics

The systematic errors faced by PICO can be categorized into three broad categories: 1) Intensity-to-polarization leakage, 2) stability, and 3) straylight, and are listed in Table ???. These were prioritized for further study using a risk factor incorporating the working group’s assessment of how mission-limiting the effect is, how well these effects are understood by the community and whether mitigation techniques exist.

The three highest risk systematic errors were studied further and are discussed in subsections below. The PICO team used simulation and analysis tools developed for Planck[146] and *CORE*, adapting them for PICO.

## 2.6.2 Absolute polarization angle calibration

CMB polarization can be rotated due to 1. a birefringent primordial Universe, or a Faraday rotation due a primordial magnetic field [147], 2. birefringent foregrounds, or interaction with the Galactic magnetic field, 3. systematic effects in the instrument, and in particular an error on the direction of polarization measured by each detector. While the first two sources create a rotation that may depend on scale, position and/or frequency, the latter depends mainly on the detector.

A rotation  $\alpha$  of the direction of polarization mixes the  $Q$  and  $U$  Stokes parameters via  $Q \pm$

| Name                                      | Risk | Effect                             |                  |
|---|------|------------------------------------|------------------|
| <b>Leakage</b>                            |      |                                    |                  |
| Polarization Angle Calibration.....       | 5    | $E \rightarrow B$                  | See Sect. 2.6.2. |
| Bandpass Mismatch.....                    | 4    | $T \rightarrow P, E \rightarrow B$ |                  |
| Beam mismatch .....                       | 4    | $T \rightarrow P, E \rightarrow B$ | See Sect. 2.6.2  |
| Time Response Accuracy and Stability..... | 4    | $T \rightarrow P, E \rightarrow B$ |                  |
| Readout Cross-talk.....                   | 4    | spurious P                         |                  |
| Chromatic beam shape .....                | 4    | spurious P                         |                  |
| Gain mismatch .....                       | 3    | $T \rightarrow P$                  |                  |
| Cross-polarization .....                  | 3    | $E \rightarrow B$                  |                  |
| <b>Stability</b>                          |      |                                    |                  |
| Gain Stability .....                      | 5    | $T \rightarrow P, E \rightarrow B$ | See Sect. 2.6.3  |
| Pointing jitter.....                      | 3    | $T \rightarrow P, E \rightarrow B$ |                  |
| <b>Straylight</b>                         |      |                                    |                  |
| Far Sidelobes.....                        | 5    | spurious P                         | See Sect. 2.6.4. |
| <b>Other</b>                              |      |                                    |                  |
| Residual correlated cosmic ray hits ..... | 3    | increased variance                 |                  |

Table 4: Systematic errors expected in PICO’s measurement of CMB polarization. Each source of systematic errors was given a rating of the risk that a given systematic error will dominate the B-mode measurement. A risk level of 5 indicates that a systematic effect is highly significant because it is design-driving, has limited past experiments, and/or isn’t well understood. Risk level of 4 indicates a systematic that is either known to be large but is understood reasonably well or a smaller effect that requires precise modeling. Risk level of 3 indicates that we expect the effect to be small, but it isn’t necessarily well understood enough that modeling it should be done in detail in a mission Phase A. This study investigated the systematics with risk levels of 5 via simulations.

$iU \longrightarrow e^{\mp i2\alpha}(Q \pm iU)$  and thus mixes the the power spectra and their correlations as illustrated in Fig. 14.

The most recent constraints on cosmological birefringence Planck collaboration [148] were limited by uncertainties on the detector orientations. In Planck, the detectors were characterized pre-launch to  $\pm 0.9^\circ$  (rel.)  $\pm 0.3^\circ$  (abs.) [149]. For PICO, the relative rotation of the detectors will be measured to a few  $0.1'$  using the CMB, but the overall rotation is unlikely to be known pre-launch to better than Planck. Known polarized sources, such as the Crab Nebula, are not characterized well enough independently to serve as calibrators; Aumont et al. [150] show that the current uncertainty of  $0.33^\circ = 20'$  on the Crab polarization orientation, limits a  $B$  mode measurement to  $r \sim 0.01$ , far from PICO’s target.

In the absence of other systematics and foregrounds, a polarization rotation error  $\alpha$  of  $10'$  degrades the error bar of  $r$  by 30%, while  $EB$ ,  $TB$  and  $BB$  spectra can measure a rotation  $\alpha$  at  $3\sigma$  when  $\alpha \sim 0.07, 0.2$  and  $0.9'$  respectively on perfectly delensed maps, and  $0.25, 0.9$  and  $4.5'$  on raw maps.

In principle, the technique of using the  $TB$  and  $EB$  spectra can detect and measure a global polarization rotation error at levels ( $0.1'$ ) below those affecting  $r$  measurements in  $BB$  ( $> 1'$ ). However, a future mission should simulate additional aspects, such as delensing, the interaction with foregrounds, and  $1/f$  noise in simulating and assessing the impact of an angle calibration error.

### 2.6.3 Gain Stability

Photometric calibration is the process of converting the raw output of the receivers into astrophysical units via the characterization of the *gain factor*  $G(t)$  which we allow to vary with time. In

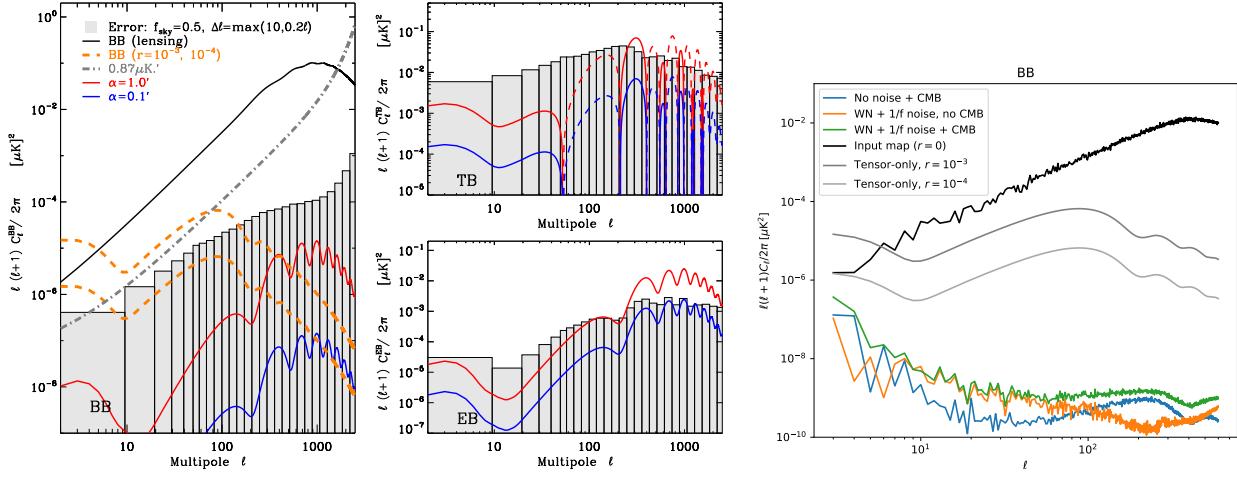


Figure 14: Effect of a rotation of the angle of polarization, assuming the Planck 2018  $\Lambda$ -CDM best fit model [37] with  $\tau = 0.054$  and expected PICO noise performance, assuming perfect delensing.

space,  $G(t)$  can be measured with the dipole. For the PICO concept study, we evaluated the impact of noise in the estimation of  $G(t)$  using the tools developed for the Planck/LFI instrument and the CORE mission proposal. The quality of the estimate depends on the noise level of the receivers, but also on the details of the scanning strategy. To analyze the impact of calibration uncertainties on PICO, we performed the following analysis: 1. We simulated the observation of the sky, assuming four receivers, the nominal scanning strategy, and  $1/f$  noise. The simulated sky contained CMB anisotropies, plus the CMB dipole. 2. We ran the calibration code to fit the dipole against the raw data simulated during step 1. 3. We again simulated the observation of the sky, this time using the values of  $G$  computed during step 2, which contain errors due to the presence of noise and the CMB signal.

The presence of large-scale Galactic emission features can bias the estimation of calibration factors. Ideally, a full data analysis pipeline would pair the calibration step with the component separation step, following a schema similar to Planck/LFI’s legacy data processing[151]: the calibration code is followed by a component separation analysis, and these two steps are iterated until the solution converges.

Results of the simulation (neglecting foregrounds) are shown as power spectrum residuals in Fig. ???. We estimate the gain fluctuations to better than  $10^{-4}$  solving for the gain every 40 hours (4 precession periods). The scanning strategy employed by PICO allows for a much better calibration than Planck, thanks to the much faster precession.

#### 2.6.4 Far Sidelobe Pickup

Measurement of each detector’s response to signals off axis, which tends to be weak ( $-80\text{dB}$  less than the peak response) but spread over a very large solid angle, is difficult to do pre-launch, and may not even be done accurately after launch. Nonetheless, this far sidelobe can couple bright Galactic signal from many tens of degrees off-axis and confuse it with polarized signal from the CMB off the Galactic plane. To evaluate this systematic error, GRASP software<sup>3</sup> was used to compute the PICO telescope’s response over the full sky. The computed full-sky beams showed

<sup>3</sup><https://www.ticra.com>

features peaking at about -70 dB of the on-axis beam. This full-sky beam was convolved with a polarized Galactic signal and a one-year PICO mission scan using the simulation pipeline and preliminarily shows that the far sidelobe pickup must be calculated accurately down to the 90 dB level in order to be removed from the measured B-mode signal to a level that does not appreciably increase the variance on the B-mode power measurement.

### 2.6.5 Key Findings

Properly modeling, engineering for, and controlling the effects of systematic errors in a next-generation CMB probe is critical. As of today, we conclude that there is a clear path to demonstrate that state-of-the-art technology and data processing can take advantage of the L2 environment and control systematic errors to a level that enables the science goals of PICO. In particular we note:

- The raw sensitivity of the instrument should include enough margin that data subsets can independently achieve the science goals. This allows testing of the results in the data analysis and additional data cuts, if needed.
- In a PICO mission, a physical optics model of the telescope should be developed, enabling full-sky beam calculations, which should be validated as much as possible on the ground. This will be needed to characterize and remove far sidelobe pickup seen during the mission.
- NASA’s support of ground-based and suborbital CMB missions will mitigate risk to a future space mission as PICO by continuing to develop analysis techniques and technology for mitigation of systematic errors.
- In a PICO mission’s phase A, a complete end-to-end system-level simulation software facility would be developed to assist the team in setting requirements and conducting trades between subsystem requirements while realistically accounting for post-processing mitigation. Any future CMB mission is likely to have similar orbit and scan characteristics to those of PICO, thus there is an opportunity for NASA and the CMB community to invest in further development of this capability now.

### 2.7 Measurement Requirements

The set of physical parameters and observables that derive from the PICO science objectives place requirements on the depth of the mission, the fraction of sky the instrument scans, the frequency range the instrument probes and the number of frequency bands, the angular resolution provided by the reflectors, and the specific pattern with which PICO will observe the sky. We discuss each of these aspects.

• **Depth** We quantify survey depth in terms of the RMS fluctuations that would give a signal-to-noise ratio of 1 on a sky pixel that is 1 arcminute on a side. Depth in any frequency band is determined by detector sensitivity, the number of detectors in the focal plane, the sky area covered, and the duration of the mission. The science objective driving the depth requirement is SO1, the search for the IGW signal which requires a depth of  $0.87 \mu\text{K} \cdot \text{arcmin}$ . This requirement is a combination of the low-level of the signal, the need to separate the various signals detected in each band, and the need to detect and subtract systematic effects to the required levels.

• **Sky Coverage** There are several science goals driving a full sky survey for PICO. The term ‘full sky’ refers to the entire area of sky available after separating other astrophysical sources of

confusion. In practice this implies an area of 50-70% of the full sky for probing non-Galactic signals, and the rest of the sky for achieving the Galactic science goals.

(1) Probing the optical depth to the epoch of reionization (STM SO5) requires full sky coverage as the signal peaks in the  $EE$  power spectrum on angular scales of 20 to 90 degrees. Measuring this optical depth to limits imposed by the statics of the small number of available  $\ell$  modes is crucial for minimizing the error on the neutrino mass measurement.

(2) If  $r \neq 0$ , the  $BB$  power spectrum due to IGW (STM SO1) also has a local maximum in the same range of angular scales (20 to 90 degrees). For  $r \gtrsim 0.001$  (CHECK) this local maximum is at a higher level than the  $BB$  lensing spectrum, making this range of scales appealing to survey, as there is no need to separate the signatures of two cosmological signals.

(3) The PICO constraint on  $N_{eff}$  (STM SO4) requires a determination of the  $EE$  power spectrum to limits imposed by the statics of available  $\ell$  modes. Full sky coverage is required to achieve this limit. (4) PICO’s survey of the Galactic plane and regions outside of it is essential to achieving its Galactic structure and star formation science goals (SO6, 7, 8).

- **Frequency Bands** The multitude of astrophysical signals that PICO will characterize determine the frequency range and number of sub-bands that PICO uses. The IGW signal peaks in the frequency range between 30 and 300 GHz. However, Galactic signals, which are themselves signals PICO strives to characterize, are a source of confusion for the IGW. The Galactic signals and the IGW are separable using their spectral signature. Simulations indicate that 21 bands, each with  $\sim 25\%$  bandwidth, that are spread across the range of 20 - 800 GHz can achieve the separation at the level of fidelity required by PICO.

Characterizing the Galactic signals, specifically the make up of Galactic dust (SO7), requires spectral characterization of galactic dust in frequencies between 100 and 800 GHz. [Aren’t there synchrotron questions that are answerable with spectral information?](#)

- **Resolution** Several science objectives require an aperture of 1.5 m and the resolution listed in Table 1. To reach  $\sigma(r) = ??$  we will need to ‘delens’ the  $E$ - and  $B$ -mode maps that PICO will generate; see Section ???. Delensing is enabled with a map that has a native resolution of 2-3 arcminutes at frequencies between 100 and 300 GHz. Similar resolution is required to achieve the constraints on the number of light relics (SO??), which will be extracted from the  $EE$  power spectrum at multipoles  $100 \lesssim \ell \lesssim 2500$ . The process of delensing may be affected by other signals, primarily the signal due to Galactic dust. It is thus required to map Galactic dust to at least the same resolution as at 300 GHz. Higher resolution is mandated by science objectives 6,7, and 8, which require resolution of 1 arcminute at 800 GHz. We have thus chosen to implement diffracted limited resolution between 20 and 800 GHz.

- **Sky Scan Pattern** [polarization systematics, 1/f noise](#)

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