

Commissioning of COSMIC
Savin Shynu Varghese
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1. turboSETI and seticore benchmarking on Setigen data:

The goal of this test is to understand the runtimes of the [turboSETI](#) and [seticore](#) softwares to conduct doppler drift search using the data from COSMIC pipeline. Since we are not ready with the finalized dynamic spectra (filterbank/h5 file) from the pipeline, here I am generating a simulated data similar to the expected data from VLASS survey using the [setigen](#) software.

Generation of simulated data:

Expected VLASS data product:

Observing length = 5s

Time resolution = 100 ms

Number of time samples = 50

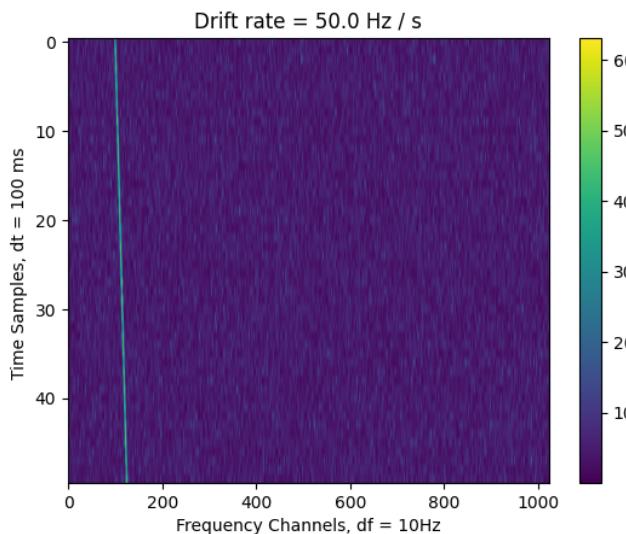
Frequency = 2-4 GHz (Here using some random value of 3.09 GHz)

Frequency resolution = 10 Hz

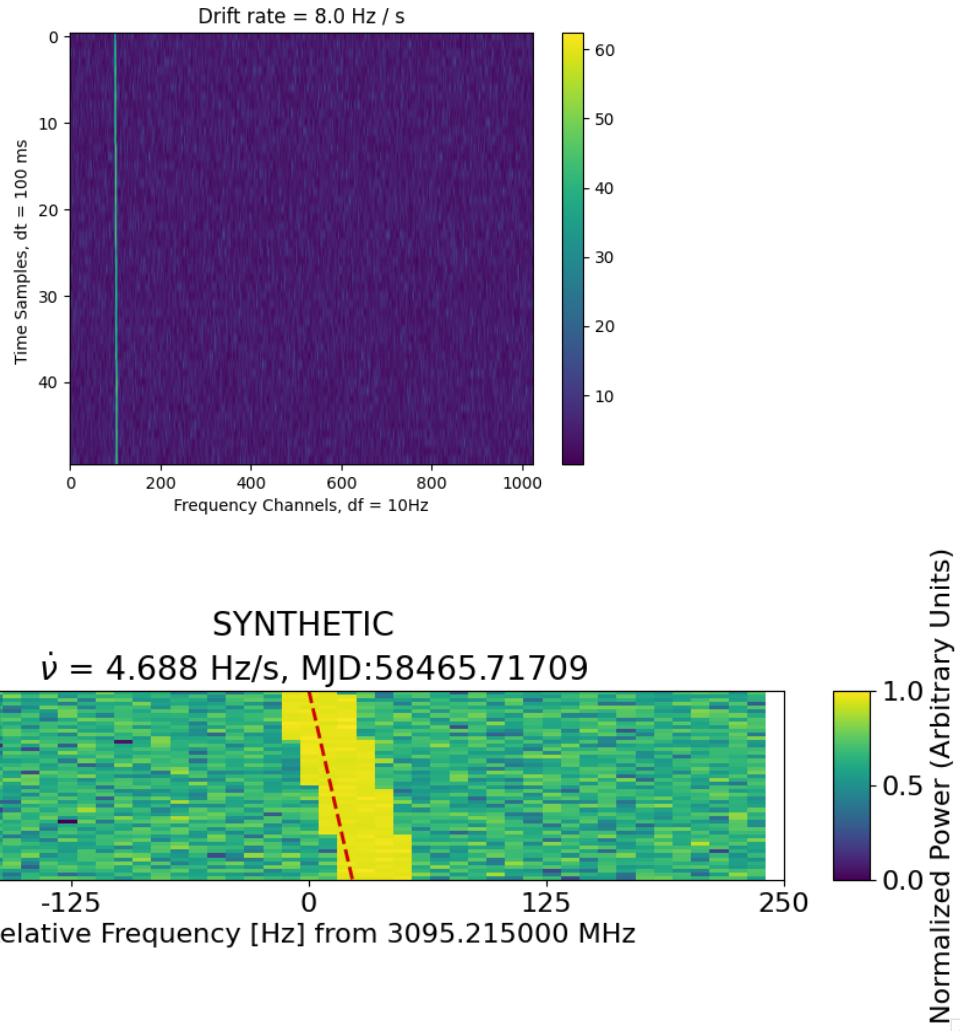
Number of fine frequency channels = 3,200,000 (if 32 MHz of data is processed in each GPU compute node)

Dynamic spectra/spectrogram = 50 (t_samples) * 3200000 (frequency channels) array

In the setigen software, first a frame is created with the above listed parameters. Then a Chi square noise (high resolution BL data follows this) is added to the frame. Then signals with preferred SNR, drift rate and width can be injected into the frame. An example of an injected signal with drift rate 50 Hz/s. Once the signals are injected, the frames are saved into h5 files and runtime tests are carried out in turboSETI.

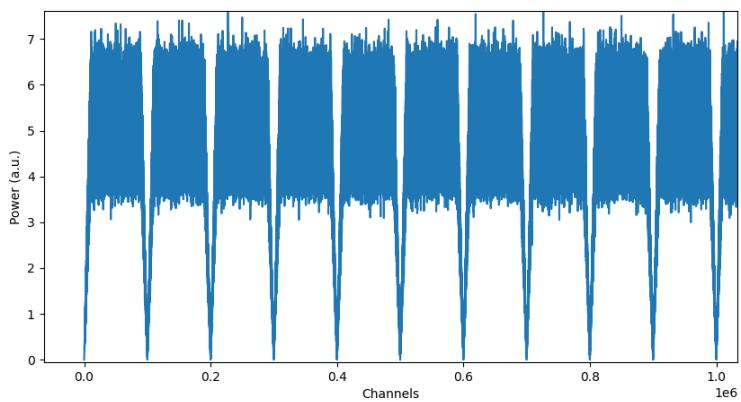


One issue: For the initial testing, I used a dynamic spectra (50 tsamples * 1024 fchannels) with signal injected at 8 Hz/s drift rate. Then I used turboSETI to confirm we are actually detecting the injected signal. Even though a strong signal was detected at a particular frequency, the drift rates of detection were different from injection. Attaching the plots of the injected signal and the one from the turboSETI identified as the hit.

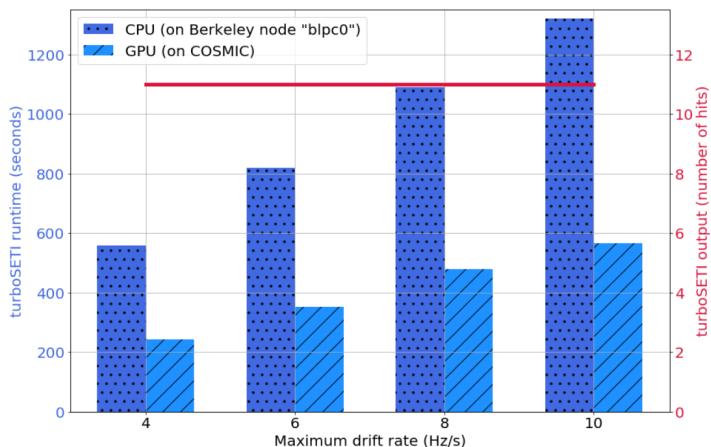


Testing from a VLASS perspective:

There would be 32 coarse channels each with 100,000 fine channels . A signal with SNR between 10-50 was injected into the middle of each coarse channel with drift rates going from 5-55 Hz/s. A bandpass response was added to each coarse channel to simulate the actual data. This created a 1.1 Gb h5 file.



From a recent turboSETI test conducted using the Meerkat data, runtime on both CPU and GPU increased as a function of drift rate. Other parameters such as SNR, frequency channels do not seem to be affecting the runtime much.



Runtimes are calculated for 2 processes: de-doppler drift search and filtering and plotting events (Both Python versions)

Turboseti search parameters:

SNR = 6, n_coarse_channel = 32,

Run times for the de-doppler drift search:

Max Drift rates (Hz/s)	GPU runtime (+/-1.2 s)	CPU runtime (+/-1.2 s)
10	3.6	6
20	3.6	6
30	3.6	6
40	3.6	6
50	3.6	6
60	4.2	6

The GPU and CPU run times seem to be independent of the max drift rates used for the search. The drift rate looks like an independent parameter when the observation times are smaller.

Finding events and plotting time:

Max Drift rates (Hz/s)	CPU runtime (s)	No of hits
10	0.46	1
20	3.49	9
30	6.71	18
40	10.22	28
50	10.42	29
60	10.51	29

I think these numbers may not be accurate. In order to confirm, I compared a test observational Meerkat file with a generated file using setigen (using observed parameters). Both files are identical based on the header information. The turboseti takes longer time for the observed data and less time (3 times lesser) for simulated data.

Maybe the simulated data is a simplified version of observed data? (Not many RFIs, etc..)

Runtimes of turboSETI and seticore:

Benchmarking of Setigen generated dynamic spectra (.h5 file) with seticore and turboseti: Comparing only the GPU run times. seticore only has a GPU version. During these tests, no process was running on any of the GPUs. seticore shows segmentation fault error messages sometimes. The search was done on two different files

First, the filterbank file were generated directly from setigen:

```
(turboseti) $varghes@cosmic-gpu-0:~/benchmark_test/test_seticore$ watutil -i data_setigen_ds32.h5
--- File Info ---
DIMENSION_LABELS : ['time' 'feed_id' 'frequency']
az_start : 0.0
data_type : 1
fch1 : 1999.5 MHz
foff : 1.017252604166666e-05 MHz
ibeam : -1
machine_id : 20
nbeams : 1
nbits : 32
nchans : 3145728
nfpc : 98304
nifs : 1
rawdatafile : Synthetic
source_name : Synthetic
src_dej : -28:22:59.16
src_raj : 17:47:15
telescope_id : 6
tsamp : 0.098304
tstart (ISOT) : 2018-12-13T17:12:37.000
tstart (MJD) : 58465.71094907406
za_start : 0.0

Num ints in file : 50
File shape : (50, 1, 3145728)
--- Selection Info ---
Data selection shape : (50, 1, 3145728)
Minimum freq (MHz) : 1999.5
Maximum freq (MHz) : 2031.4999898274739
(turboseti) $varghes@cosmic-gpu-0:~/benchmark_test/test_seticore$
```

At the same time, If I generate raw voltages for a single antenna with the same expected configurations from a single COSMIC node: 32 coarse channels (1 MHz bandwidth and 1 us resolution). Then using rawspec to achieve the time and frequency resolution (0.1 seconds and 10 Hz resolution) similar to VLASS. The results on this data product are different.

```
(turboseti) svarghes@cosmic-gpu-0:~/benchmark_test/test_seticore$ watutil -i synth_guppi_sing_ant_sig.rawspec.0000.h5
--- File Info ---
DIMENSION_LABELS : ['time' 'feed_id' 'frequency']
    az_start :          0.0
    data_type :          1
    fch1 :           1999.5 MHz
    foft : 1.017252604166666e-05 MHz
    ibeam :          -1
    machine_id :          20
    nbeams :          1
    nbits :          32
    nchans :        3145728
    nfpc :          98304
    nifs :          1
    rawdatafile : synth_guppi_sing_ant_sig.0000.raw
    source_name :      SYNTHETIC
    src_dej :        0:00:00
    src_raj :        0:00:00
    telescope_id :          -1
    tsamp :          0.098304
    tstart (ISOT) : 2000-01-02T00:00:000
    tstart (MJD) :        51545.0
    za_start :          0.0

Num ints in file :          50
File shape :        (50, 1, 3145728)
--- Selection Info ---
Data selection shape :        (50, 1, 3145728)
Minimum freq (MHz) :        1999.5
Maximum freq (MHz) :        2031.4999898274739
```

Benchmarking of seticore and turboseti on h5 file generated directly by setigen:

Max Drift rates (Hz/s)	Seticore (s)	Turboseti (s)
10	20	6
20	18	6
30	14	6
40	27	6.6
50	22	6
60	17	6

Benchmarking of seticore and turboseti on h5 file generated by channelization of raw voltages:

Max Drift rates (Hz/s)	Seticore (s)	Turboseti (s)
10	8	10.8
20	7	11.4
30	7	10.8
40	7	10.8
50	7	10.8
60	7	10.8

At some drift rates, few of the turboSETI hits are missing in seticore and vice versa. Other than that hits from turboseti and seticore agree with each other. Overall the seticore seems to be slightly faster over the turboSETI GPU version.

2. Fringes between antennas:

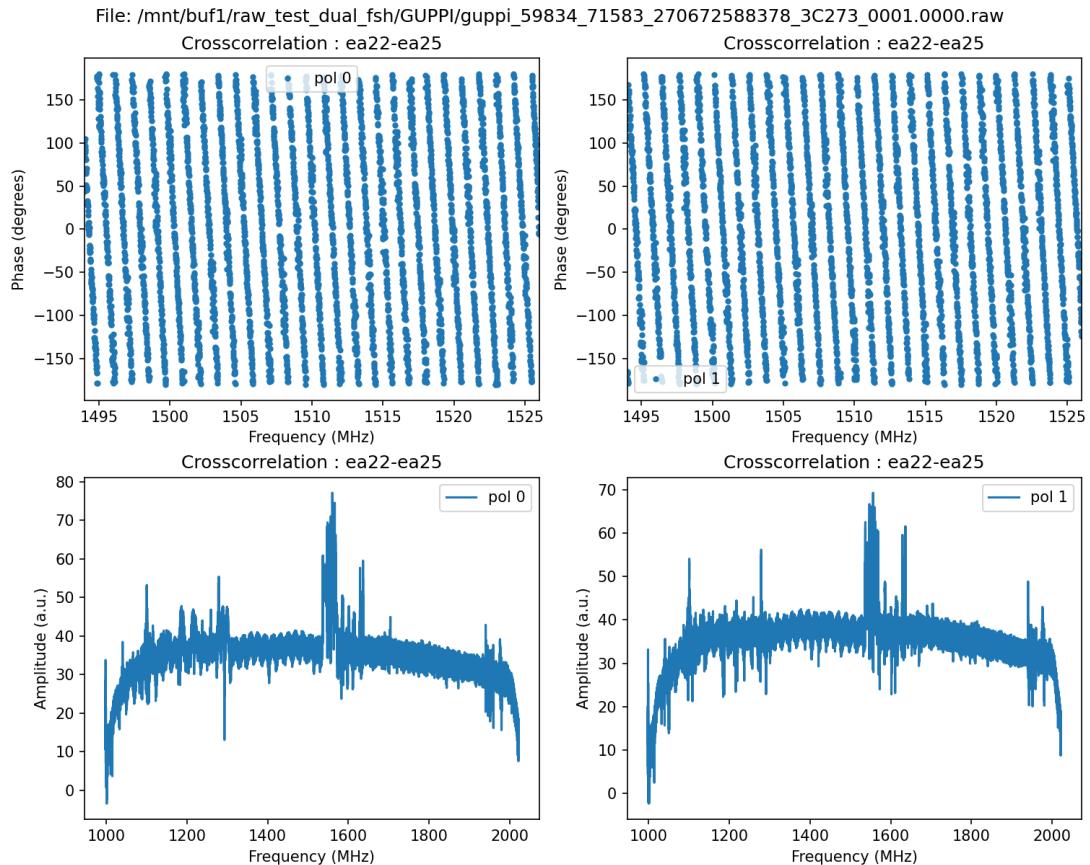
After setting up the basic hashpipe pipeline to collect data from FPGAs to GPU machines, we started conducting observations with a small number of antennas. At this stage we mainly utilized the raw mode to collect and save voltages as Guppi raw files. The first step was to get coherence from a pair of antennas. For that we cross correlate the voltage data for each baseline and average them for a short duration of time (~ 1 sec). If there is a fixed time delay between a pair of antennas, that will show up as fringes in the phase vs frequency plot for a baseline. The slope of the phase and how fast the phases are changing across the frequency depends on the delay values. If the delay values are higher, the fringes will change faster across frequency and slope would be higher. At the same time, if the delay values are smaller, they will appear as a slow slope across frequency.

The FPGAs conduct the channelization of the voltage data to 1 MHz frequency coarse channels. If the delay is $\ll (1/1\text{MHz} = 1 \text{ us})$, then we would be able to see fringes with the coarse channelization. However, if the delay $> 1\text{us}$, then a step of upchannelization is required to see the fringes. The commissioning started in the VLA C configuration and a total delay greater than the 1 us is expected for several baselines. Here an upchannelization factor of 128 was used to get around 8 kHz frequency channels. In order to observe coherence, a target observation of a bright 3C source across 1-10 GHz was conducted during the test times which will record 5 sec observations as voltage files. The observations were done in 4 ways to verify the system:

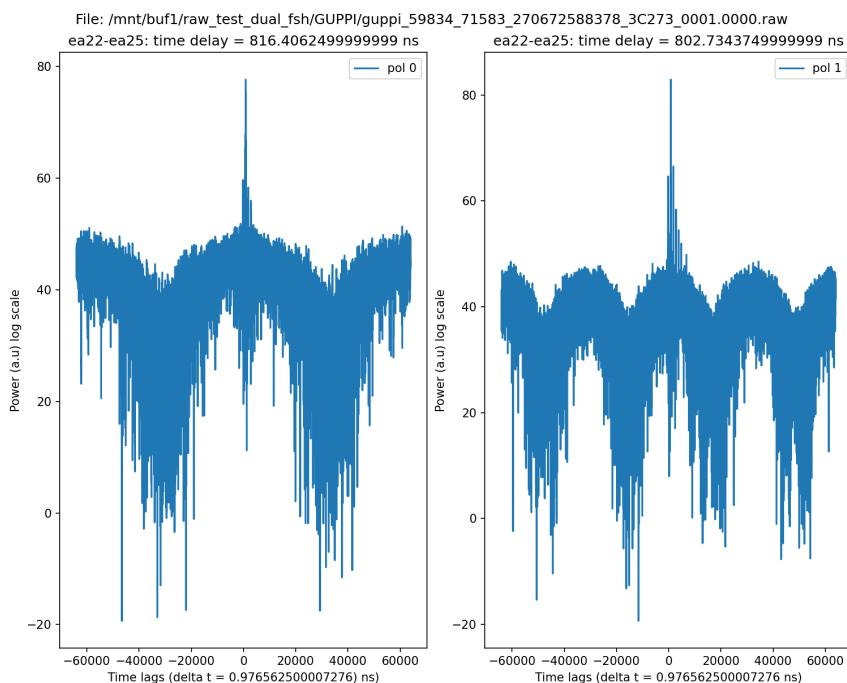
1. antenna fshift on | fpga unshift on
2. antenna fshift on | fpga unshift off
3. antenna fshift off | fpga unshift on
4. Antenna fshift off | fpga unshift off

Technically we should see strong fringes in test1 and lesser strong fringes in test4. Also, we should see decoherence in tests 2 and 4. These tests were carried out this way to make sure that the unshifting of frequencies in the FPGAs were working properly.

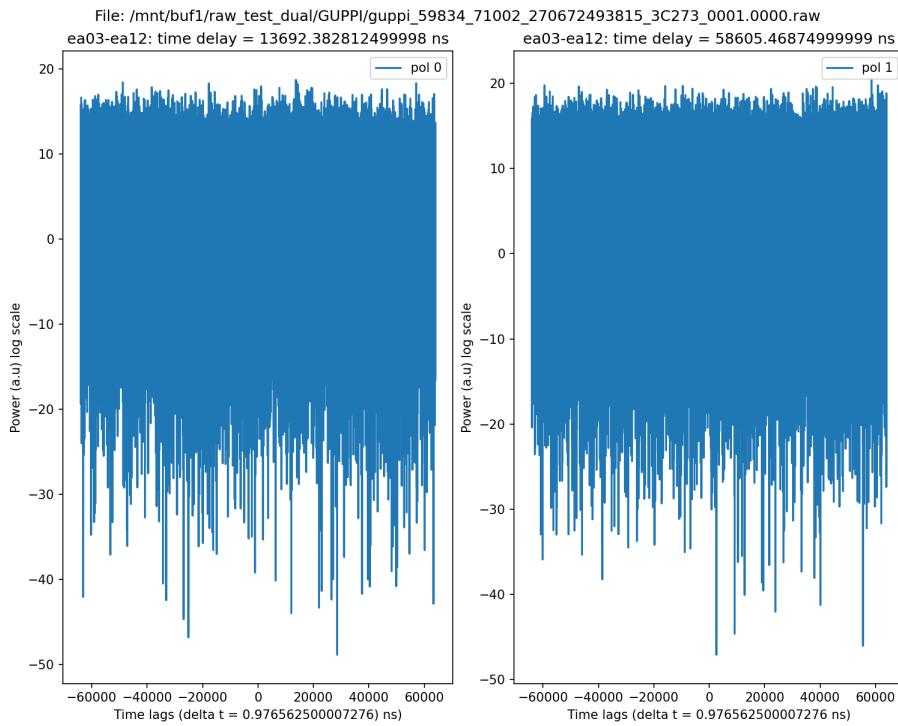
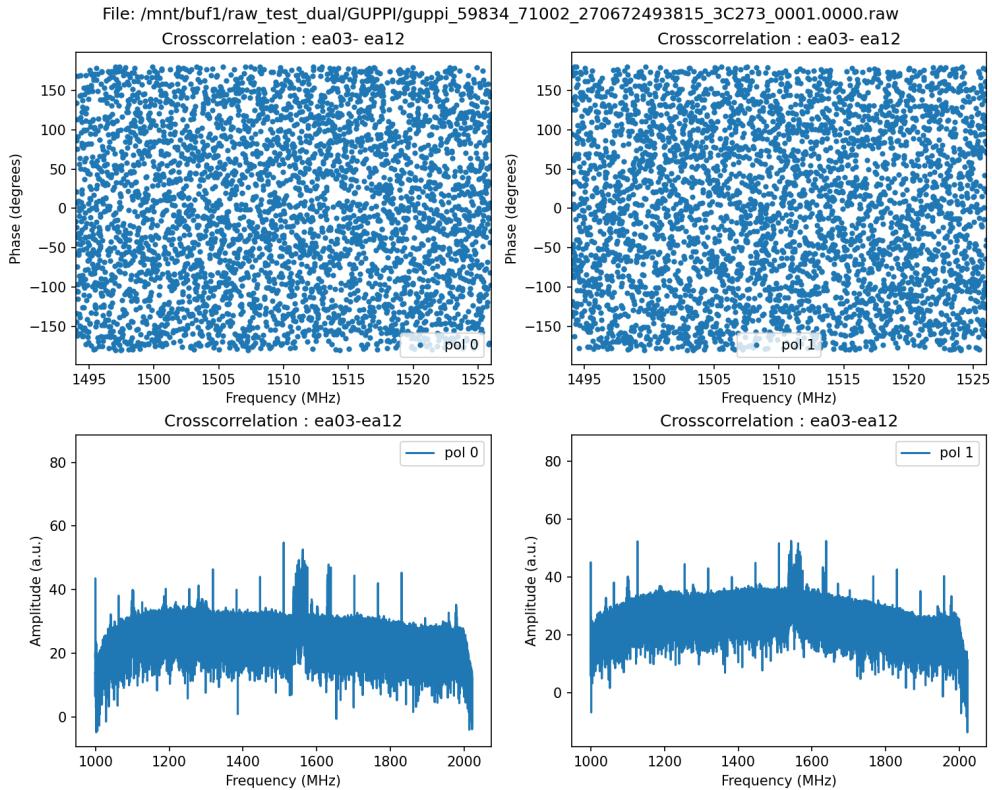
The images below show the fringes in phase vs frequency for a baseline ea22-ea25 in test1 using 3C273. The corresponding amplitude responses are also shown in the bottom panel for each polarization.



An inverse FFT of the upchannelized data would give the delay peak corresponding to each baseline. The figure below shows the power as a function of different delay lags and peak delay value is noted in the title of the plot for each polarization. In this case each polarization is measuring around ~800 ns delay.



Similar results from test 2 shows decoherence on baselines and inverse FFT of the data would look like noise without any delay peaks



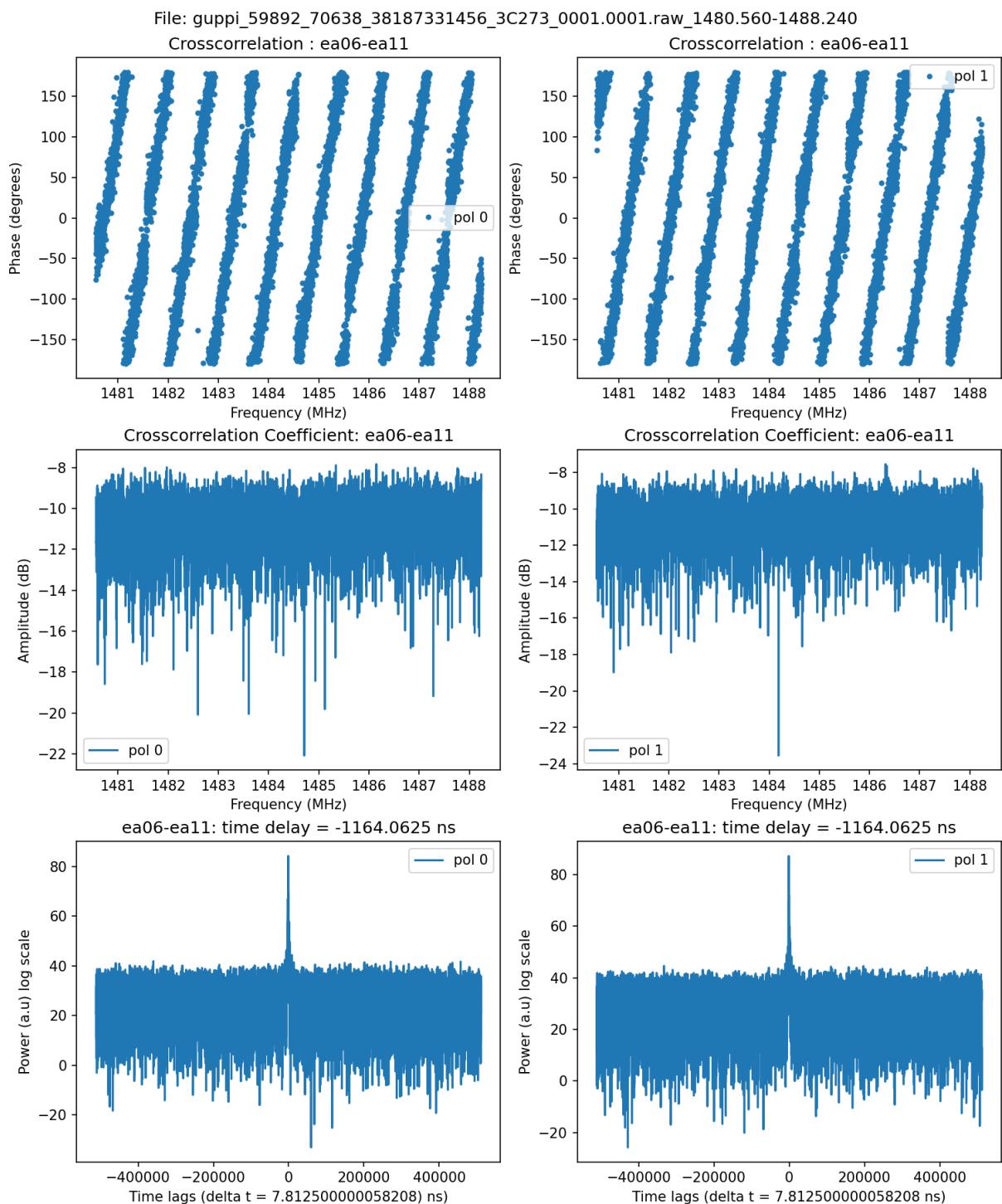
On a regular operation process, we follow test1 where the antenna fshift is on from the NRAO side and a corresponding unshifting is also on from the cosmic side.

Another example of coherence from a regular operation mode is shown below from observations of 3C273 at L band. Top panel shows the phase vs frequency in which fringes can be seen clearly. The middle panel shows the correlation coefficient and the bottom panel shows the corresponding time delay of the baseline.

The codes used for the upchannelization, cross correlation, delay calculation and plotting of the fringe patterns and amplitudes are available at

https://github.com/savindhynu/VLA_COSMIC/blob/master/obs_rawdata/upchan_coherence_opt2.py

https://github.com/COSMIC-SETI/COSMIC-VLA-CalibrationEngine/blob/rfi_mitigation_and_arrayconfig_update/upchan_coherence.py

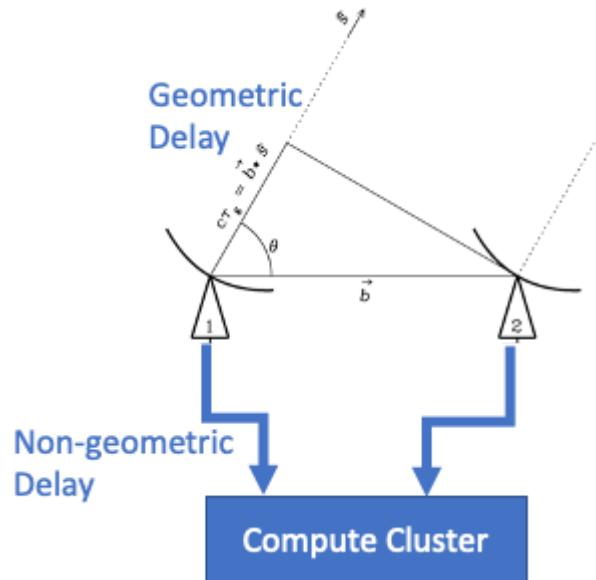


3. Delay Calibration

After observing coherence between all the baselines in an observation, the next important step is to conduct the delay calibration. COSMIC receives a copy of the digitized voltages from the antennas directly and the signal pathway is different from the WIDAR. The delay calibration consists of 2 steps

- Estimation of fixed or non-geometric delays mostly associated with the fibers, electronics etc
- Estimation of geometric delays which can be easily calculated using models

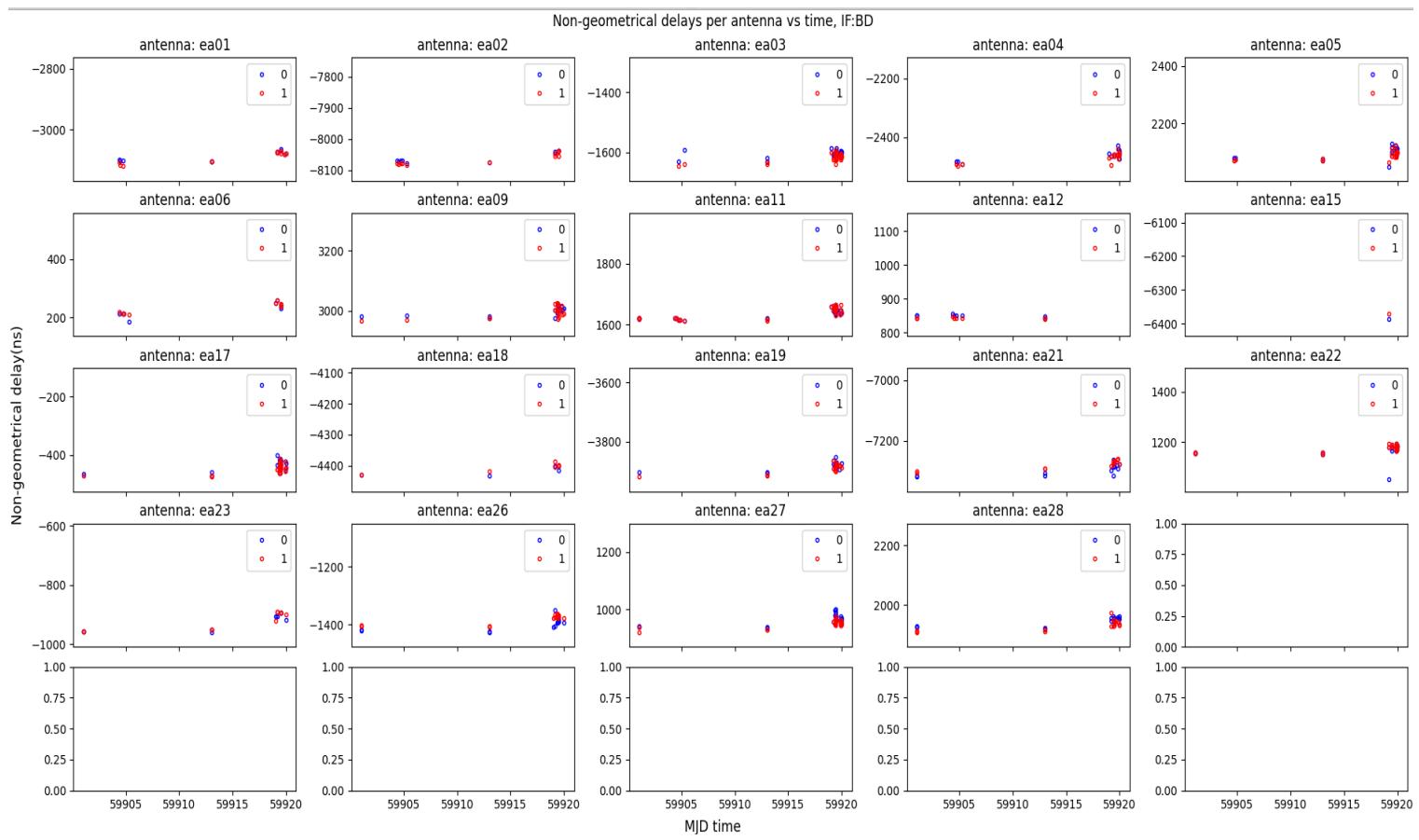
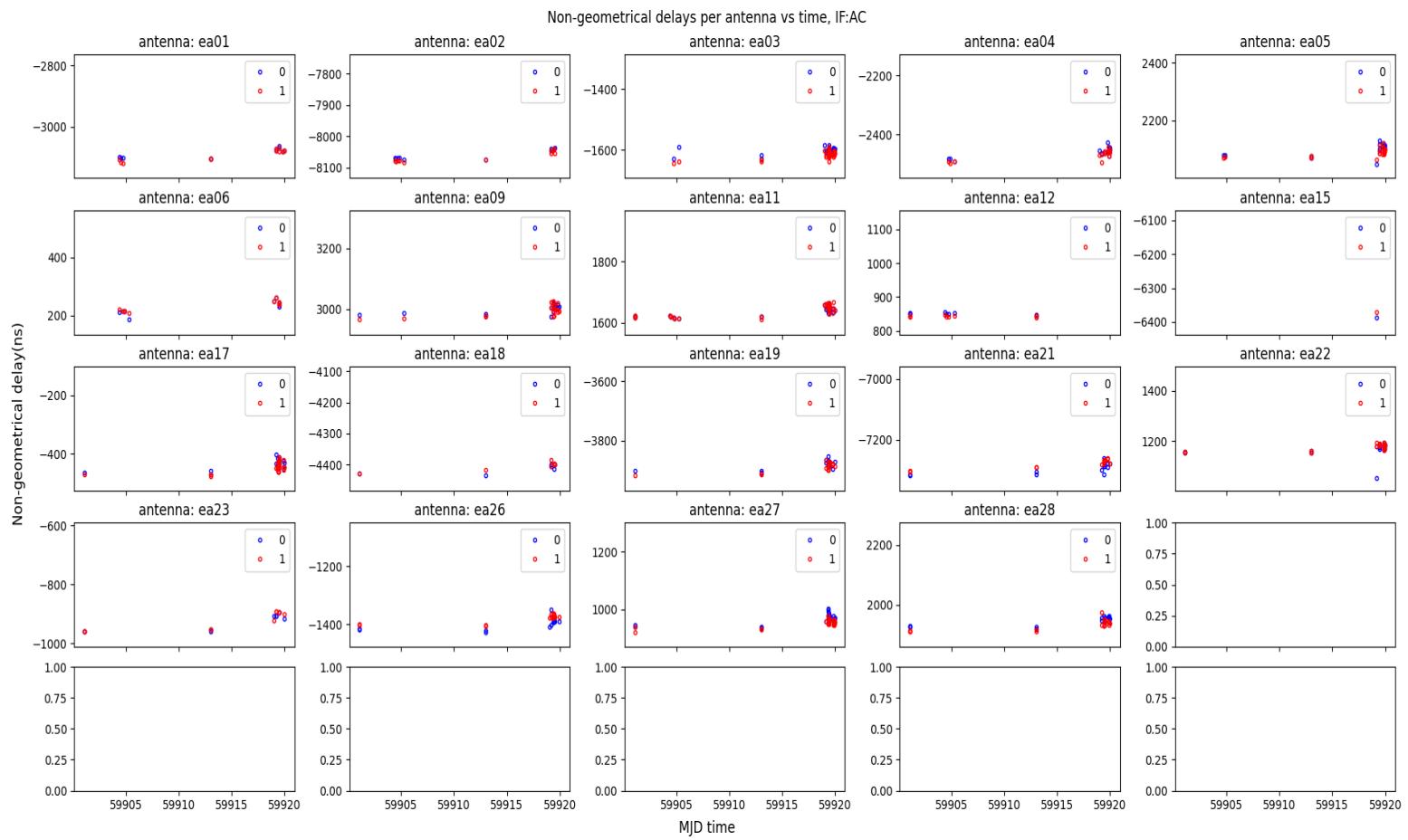
Here the non-geometric delays should be constant as a function of time and we expect up to 10-15 ns difference for different receivers. At the same time the geometric delays will change as a function of time, baseline and source direction.



The total delays are calculated using the inverse FFT of the cross correlated data for each baselines. The geometric delays are calculated using a well tested model used in CASA. The geometric delays are subtracted from the total delays to measure the fixed/non-geometrical delays for each baseline or each antenna wrt to reference antenna. These fixed delays are measured/updated during the start of each VLA reconfiguration and updated in the FPGAs. A sum of the fixed delay values and the geometrical delay based on the VLA model is compensated for each antenna in the FPGAs. This will ensure delay tracking of sources in the observation.

We conducted observation of bright 3C sources and other calibrators over a few weeks to measure the fixed delay values and measure their consistency. If the fixed delay values are not consistent over time, the delay tracking will fall apart. The plots below show the consistency of measured fixed delay values from different observations of source across multiple bands in the AC and BD tunings. An average of these delays measured across multiple days are written into a csv file which is later uploaded to the FPGAs for delay tracking. This process is repeated in a partially automated manner during each configuration change. The codes for fixed delay calculations are embedded in

https://github.com/COSMIC-SETI/COSMIC-VLA-CalibrationEngine/blob/rfi_mitigation_and_arrayconfig_update/upchan_coherence.py



4. Understanding Correlated Data Products

Once the delay tracking is enabled, the next important step of the commissioning was to move towards the correlated mode which cross correlates the voltages and outputs them in the UVH5 format.

Important tests with the correlated data:

1. Delay tracking as a function of time
2. Checking fringe rotation in the data
3. Phase tracking as a function of time/removing fringe rotation

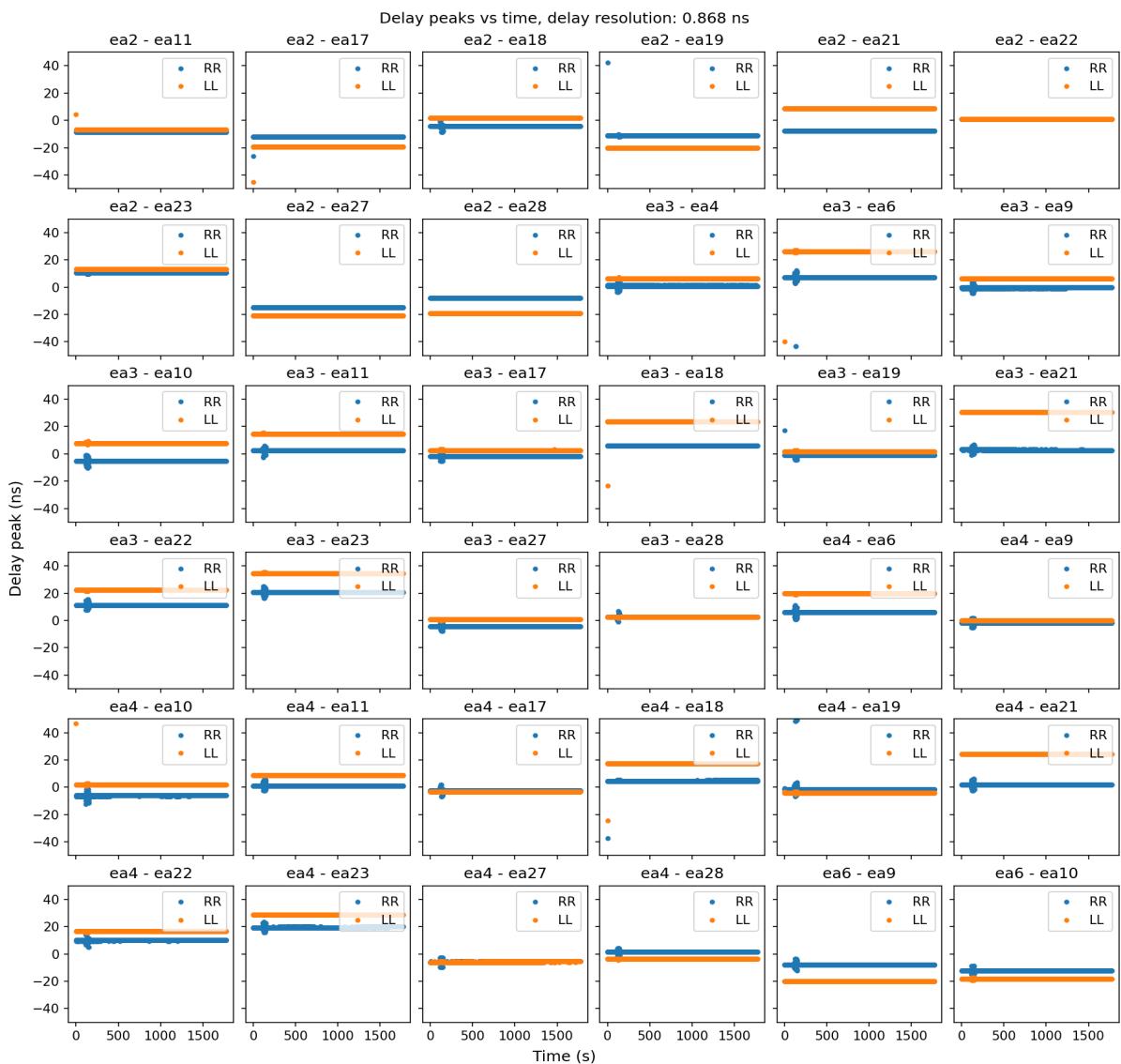
In order to conduct testing of the correlated data, observations of bright 3C radio sources were carried out in both tunings for shorter (2 min) and longer (30 min) durations. For the diagnose the UVH5 files, I wrote a python script which is available in the link below

https://github.com/COSMIC-SETI/COSMIC-VLA-CalibrationEngine/blob/rfi_mitigation_and_arrayconfig_update/calibrate_uvh5.py and
https://github.com/savinshynu/VLA_COSMIC/blob/master/calib/calibrate_uvh5.py.

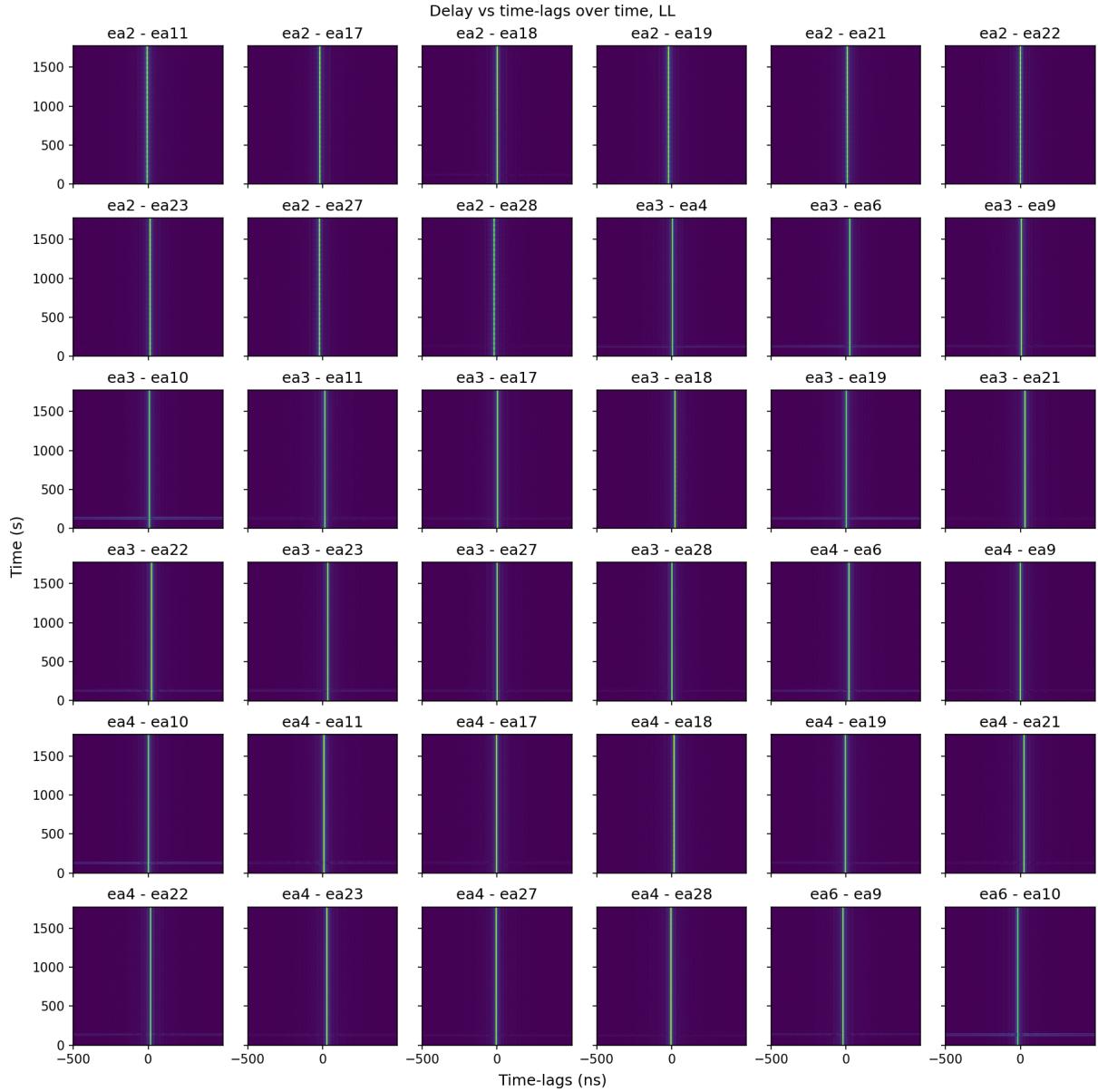
The calibrate_uvh5 class has several methods which can collect the data from uvh5 files and makes a different number of diagnostic plots.

Testing delay tracking:

The plot below shows the residual peak delay measured as a function of time for different baselines showing consistency of the delay tracking. The observations used 3C 48 and showing RR and LL data products. Once the delay tracking is applied the residual delay values are less than 50 ns which are later corrected in FPGAs in subsequent calibration observations.

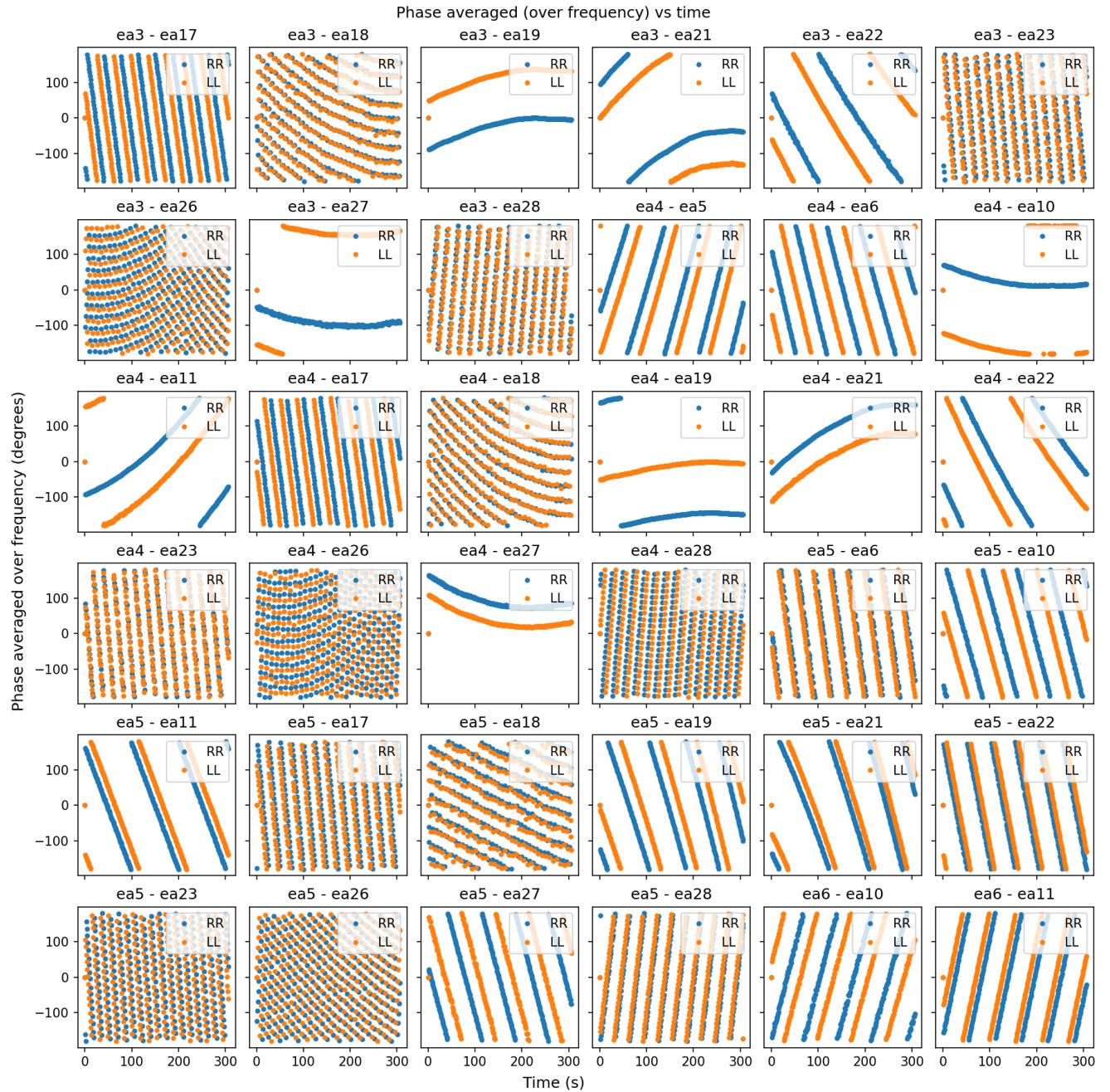


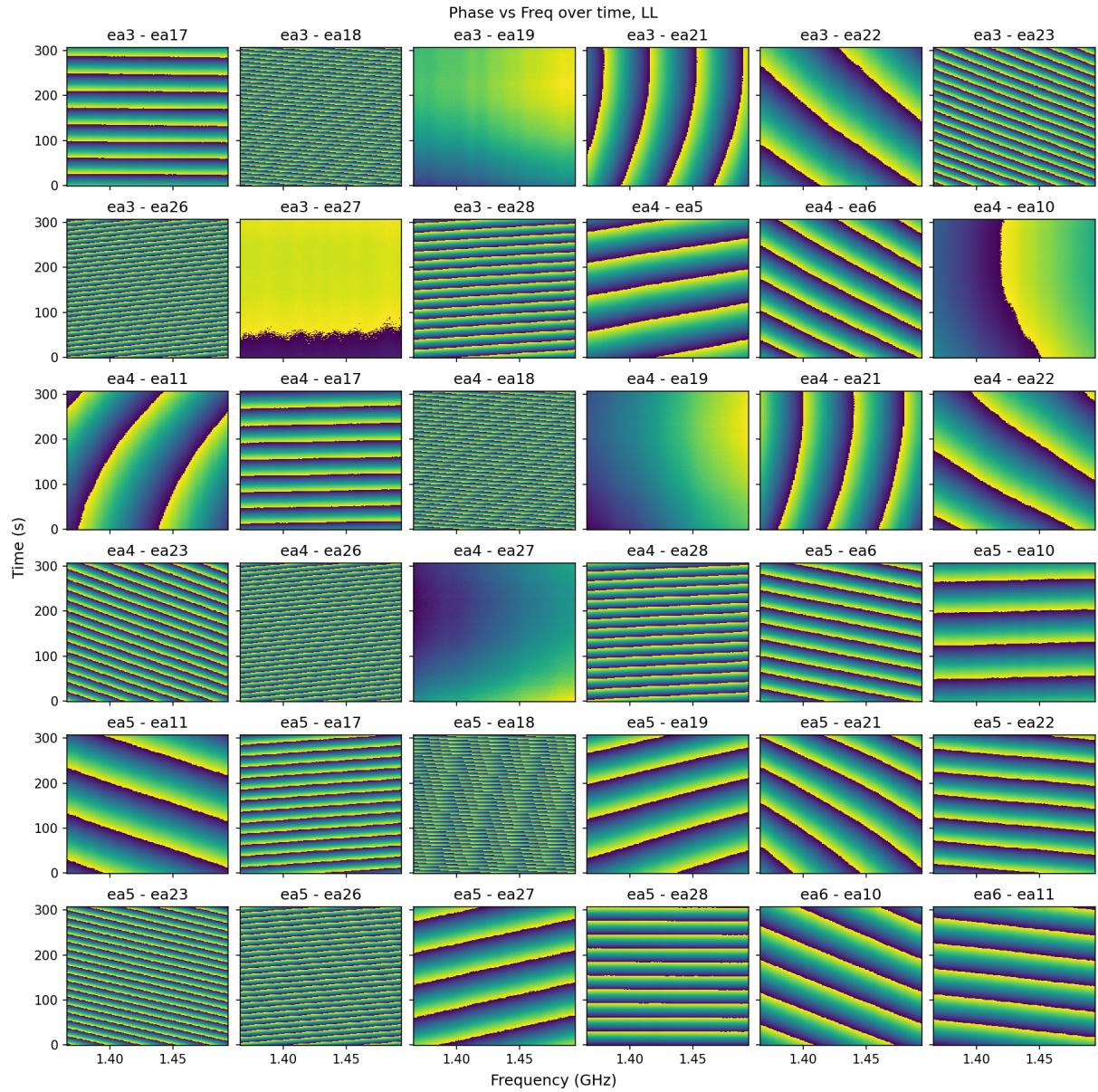
The plot below shows the delay waterfall (delay measure for each time integration) plot of the same baselines showing bright and constant delay across time.



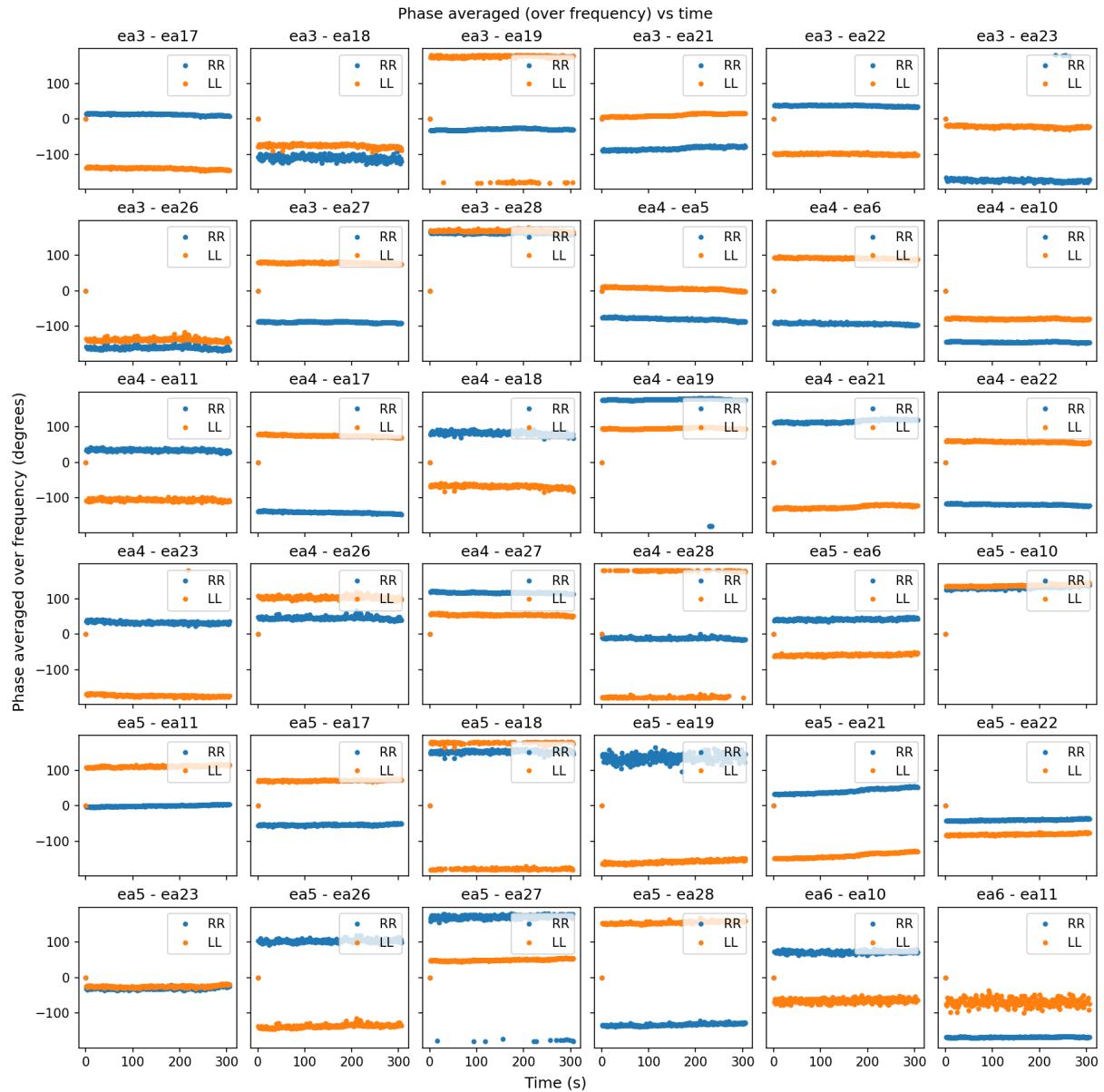
Even though the delays are calculated at the sky or RF frequencies, the antenna electronics operate at the IF frequencies and they are applied at IF frequencies. This would introduce a frequency independent phase shift over time. This causes a shift of phase fringes over time commonly known as fringe rotation or fringe oscillation. This needs to be corrected in the voltage domain before the correlation happens.

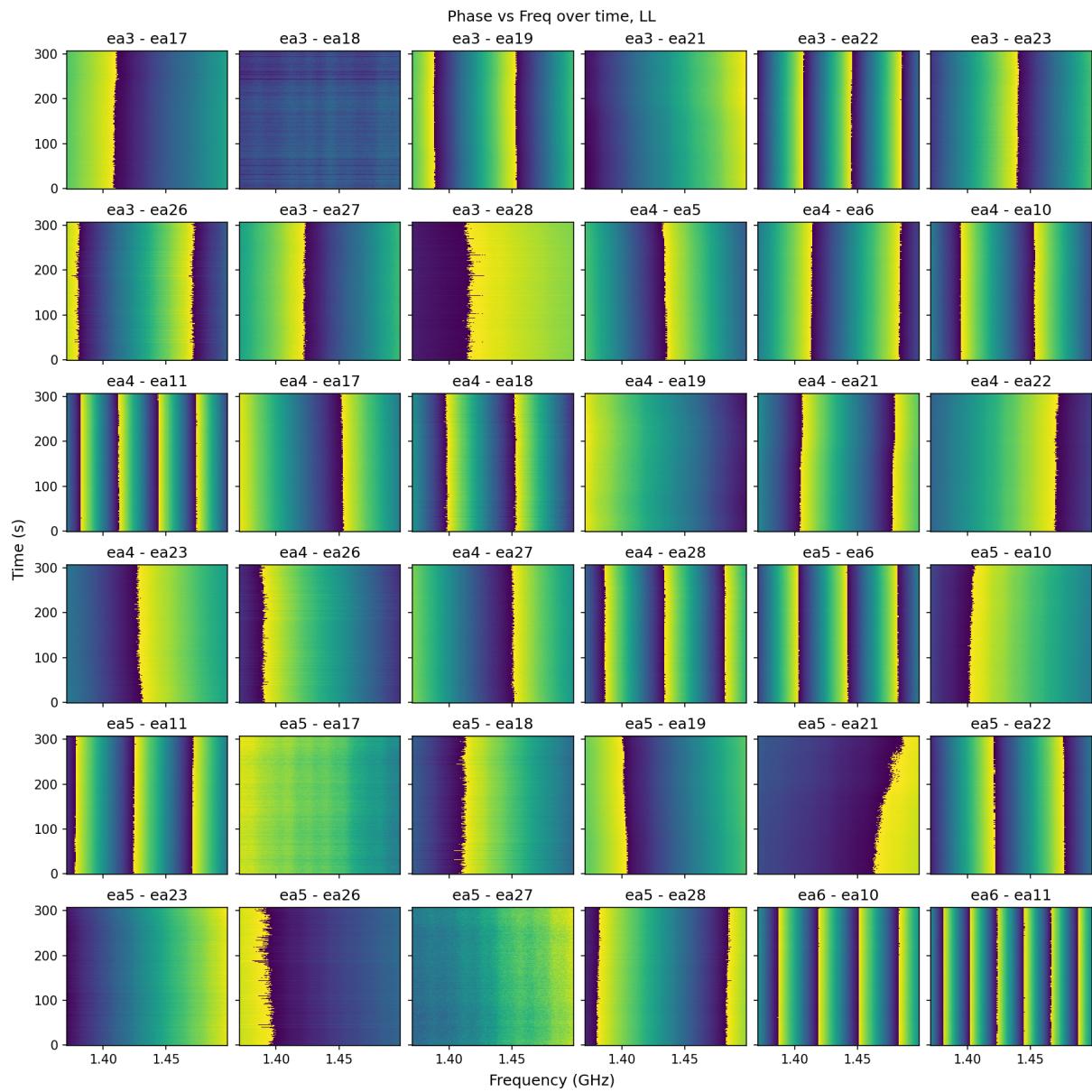
The plot below shows the phase averaged across frequency as function of time and the phase waterfall plots before stopping fringe rotation or the without phase tracking. The phase changes as a function of time and this is problematic as we need phase stability to conduct calibration and imaging.





The plots below show the averaged phases and phase waterfall as a function of time with phase tracking / correcting the fringe rotation. We can clearly see consistent phases over time in the averaged plot and the waterfall plots.





Once the delay tracking and phase consistency is achieved, the correlated data is good to derive the calibration and can be used for imaging purposes.

5. Gain Calibration

There are different types of calibration when it comes to correlated data. The data product at this point is mostly delay calibrated. Still there could be delays on the order of ~ a couple of ns which need to be corrected. The first priority of COSMIC is detection of an ETI signal and the only calibration that we focus at the moment is the complex gain or ‘phase’ calibration. For the same reason, we are not conducting any bandpass or the amplitude calibration in real time. However the correlated data of amplitude and bandpass calibrators are saved in case we need to calibrate the bandpass response and amplitudes post ETI detection.

The gain calibration is usually antenna based and mainly accounts for the time variable factors associated with the instrument (power level setting within the receivers, changes in antenna gains, etc) and the atmospheric effects (troposphere at high frequencies, ionosphere at low frequencies, water content, solar flares, etc).

Here we are using a modified version of the Sdmpy package which was initially written to calibrate the VLA Science Data Model (SDM) datasets. This codebase was modified to ingest COSMIC data and derive calibrations. The derive_gains method in the calibrate_uvh5 class of

https://github.com/COSMIC-SETI/COSMIC-VLA-CalibrationEngine/blob/rfi_mitigation_and_arrayconfig_update/calibrate_uvh5.py is used to derive the antenna gains. The cross correlation / visibility matrix is modeled as the outer product of the gains from the antennas. First the data matrix is converted from baseline axis to antenna axis. Then an eigen value decomposition is used to calculate the eigen vectors and eigen values which is later used to get the gain values per antenna per time per frequency per polarization.

Deeping dive into the mathematics:

The data matrix is converted to a reduced matrix using the low rank approximation technique. The largest eigen value and the corresponding eigen vector from the eigen value decomposition of the data matrix is used to create the reduced data matrix. The autocorrelation values in the data matrix before decomposition are zero filled for this process. However they are replaced with the values from the reduced matrix.

Every complex matrix can be decomposed into its eigen values and eigen vectors using singular value decomposition and an approximate reduced matrix can be generated using highest eigen values and the corresponding vector as shown below.

$$V_{obs} = U * \lambda U = U_1 * \lambda_1 U_1 \quad (\text{reduced matrix})$$

$$V_{ij} = \lambda_{ij} u_i * u_j \quad (1) \quad (\text{single baseline visibility})$$

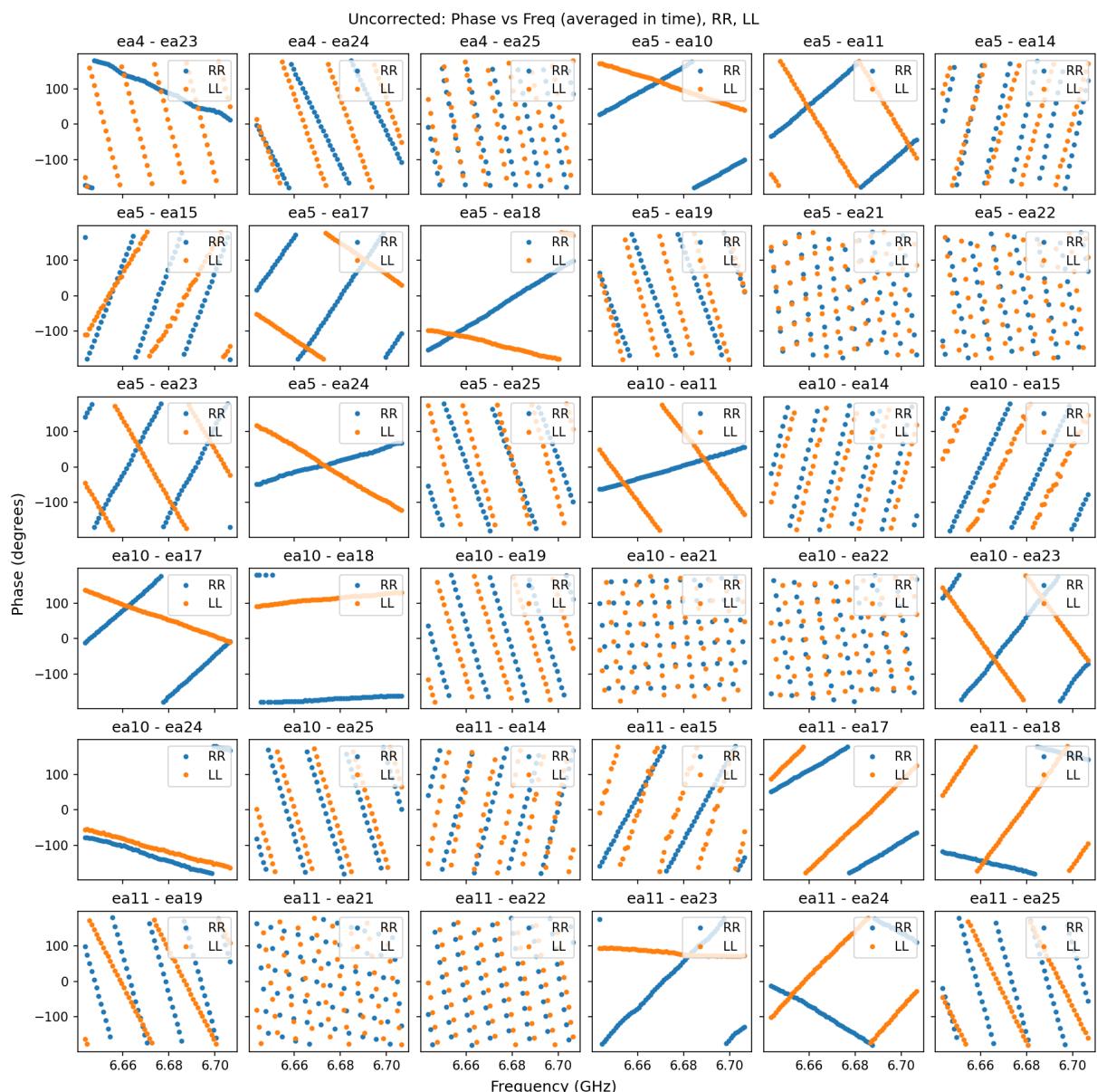
If we have a point source, their visibility for a baseline can be written as

$$V_{ij} = So \ G_i \ G_j^* = \sqrt{So} \ G_i \ \sqrt{So} \ G_j^* \quad (2)$$

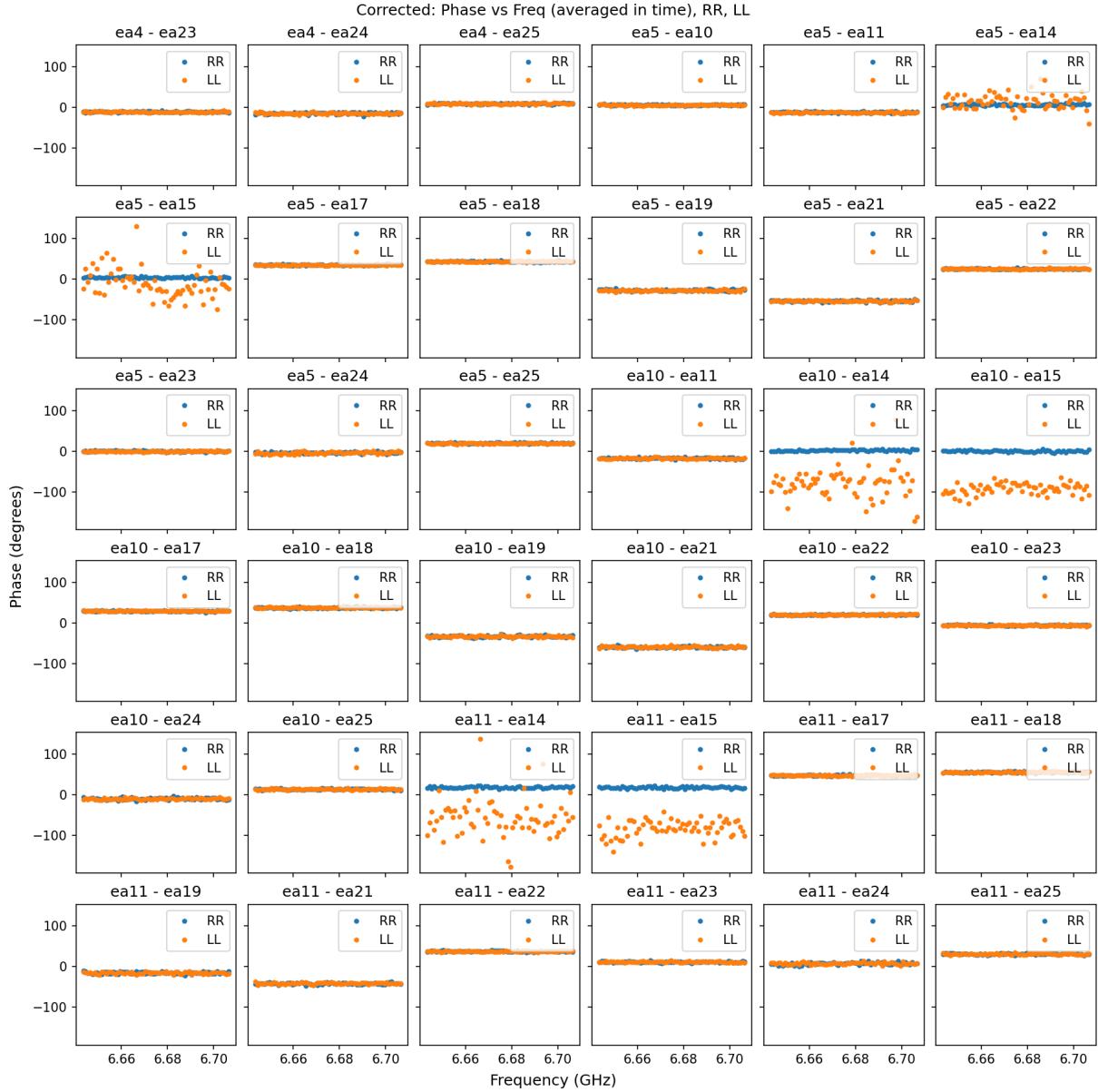
Where V is the visibility, So is the flux density and G are the gains of each antennas. Or in other words, a single baseline response to a point source ideally has zero phase and constant flux density. Any variation from it can be modeled as a gain variable for an antenna multiplied by the constant flux density parameter. If we compare equations 1 & 2, gains of each antenna are equal to the reduced eigen vectors.

I guess there is a simple way to explain this whole thing: More questions ask Paul Demorest.

Does this work ? Yes, it works. The plot below shows the phase vs frequency for different baselines before and after gain calibration.



The plots phase vs frequency plots after calibration has flat and zero phases.



6. Offline calibration and First images from COSMIC

Before online calibration, the gain calibration algorithm was tested using a back to back observation of 2 close calibrators. For this purpose, we conducted 3 minute observations 3C286 and J1310+3220 bright sources.

Tests with gain calibration code:

1. Derive the gain calibration from 3C 286 and apply it back to the same source.
2. Apply the derived gain calibration from 3C 286 to the J1310+3220
3. Look at the phase solutions in each case

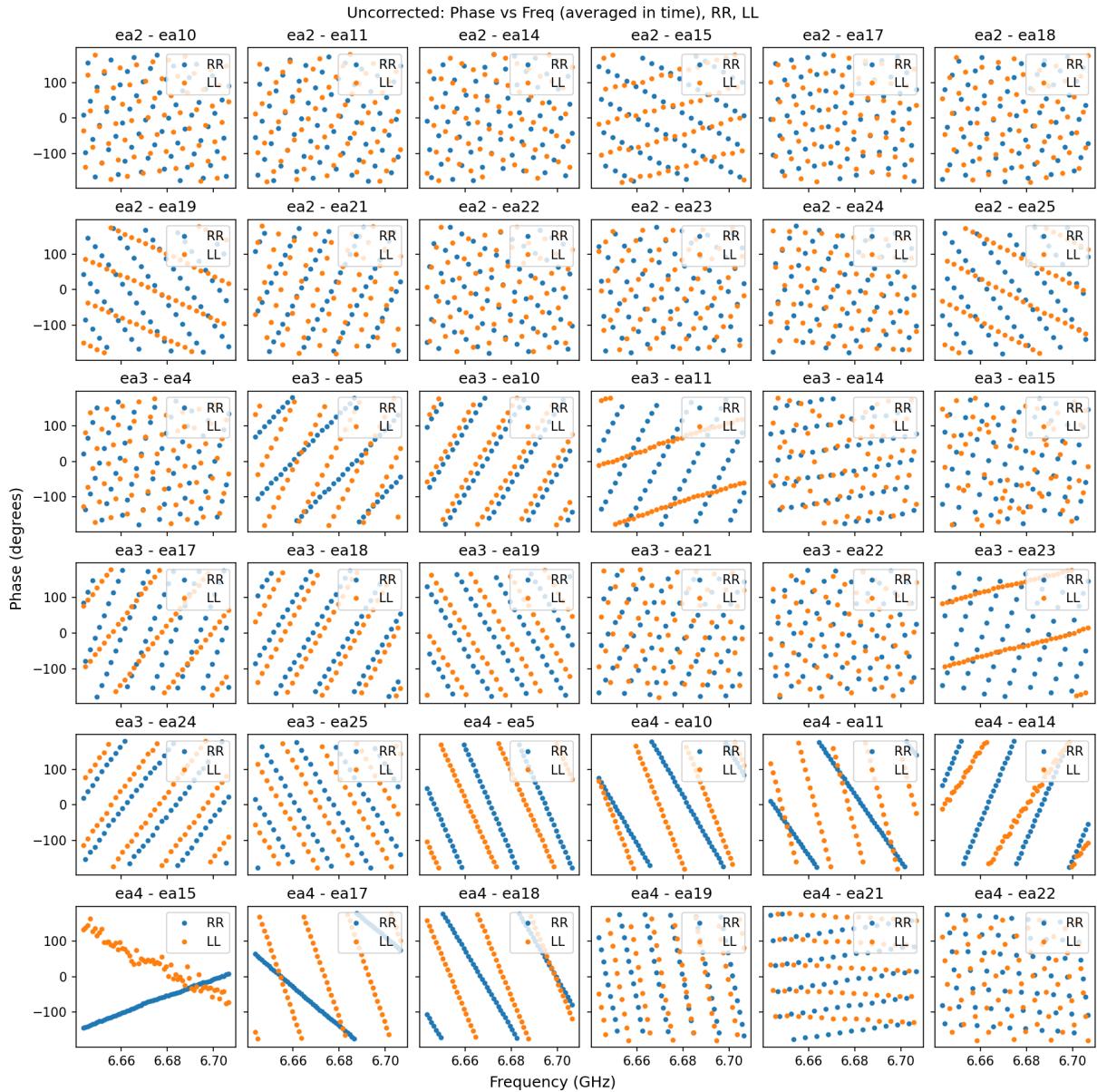
Tests with CAS

4. Conduct calibration and imaging of each datasets independently in CASA

Test1:

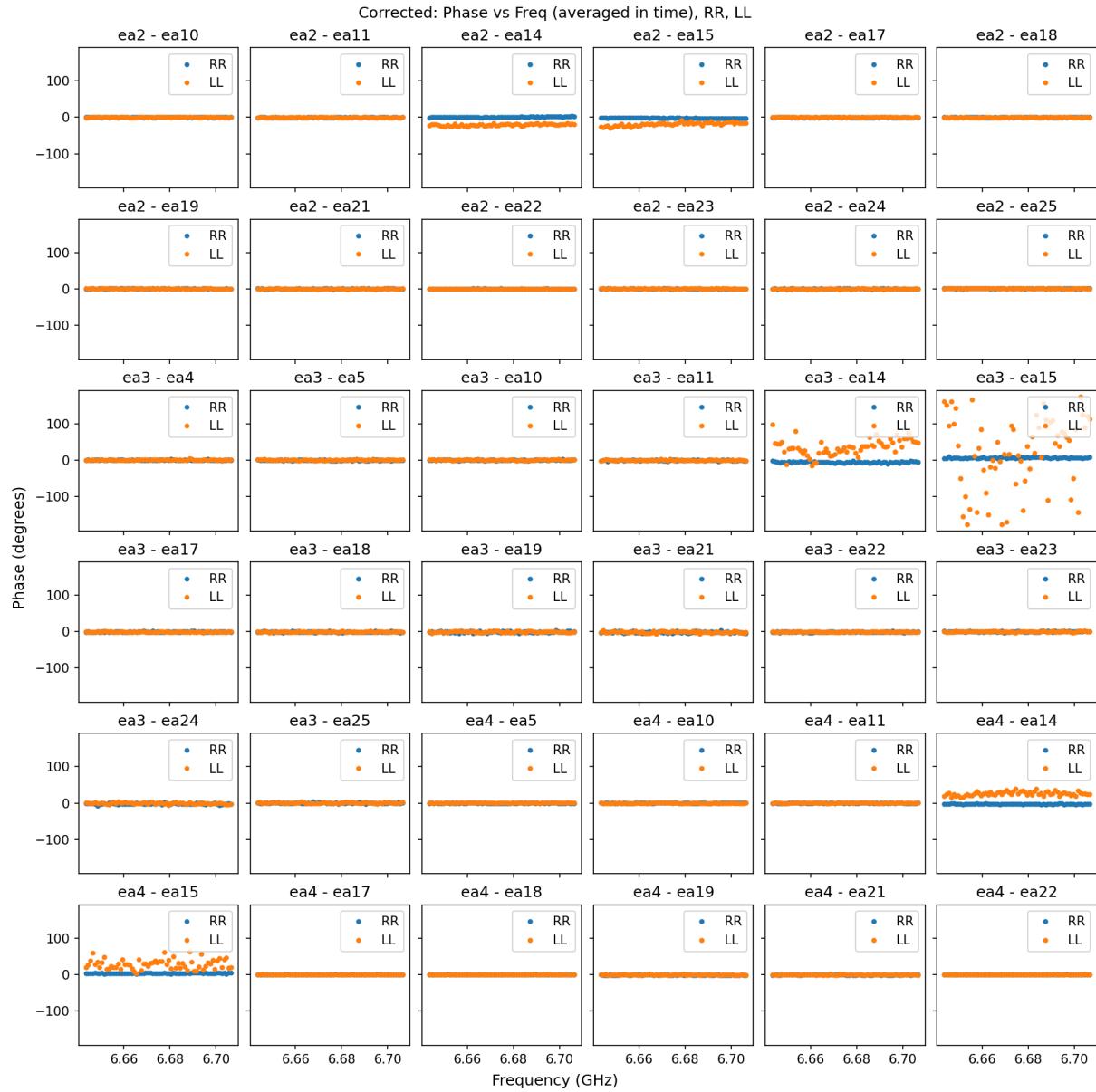
Plot below shows the phase vs freq for each baseline averaged over time before calibration on 3C 286 showing good fringes.

Before Calibration:



Now the gain calibration derived from 3C286 is applied on the same dataset. The phases vs frequency plot after applying the correction is shown below for all the baselines. The phases are flat and zero for both RR and LL in most of the baselines. The phases deviates from baselines with bad fringes which could be due to RFI or data recording issue in certain antennas.

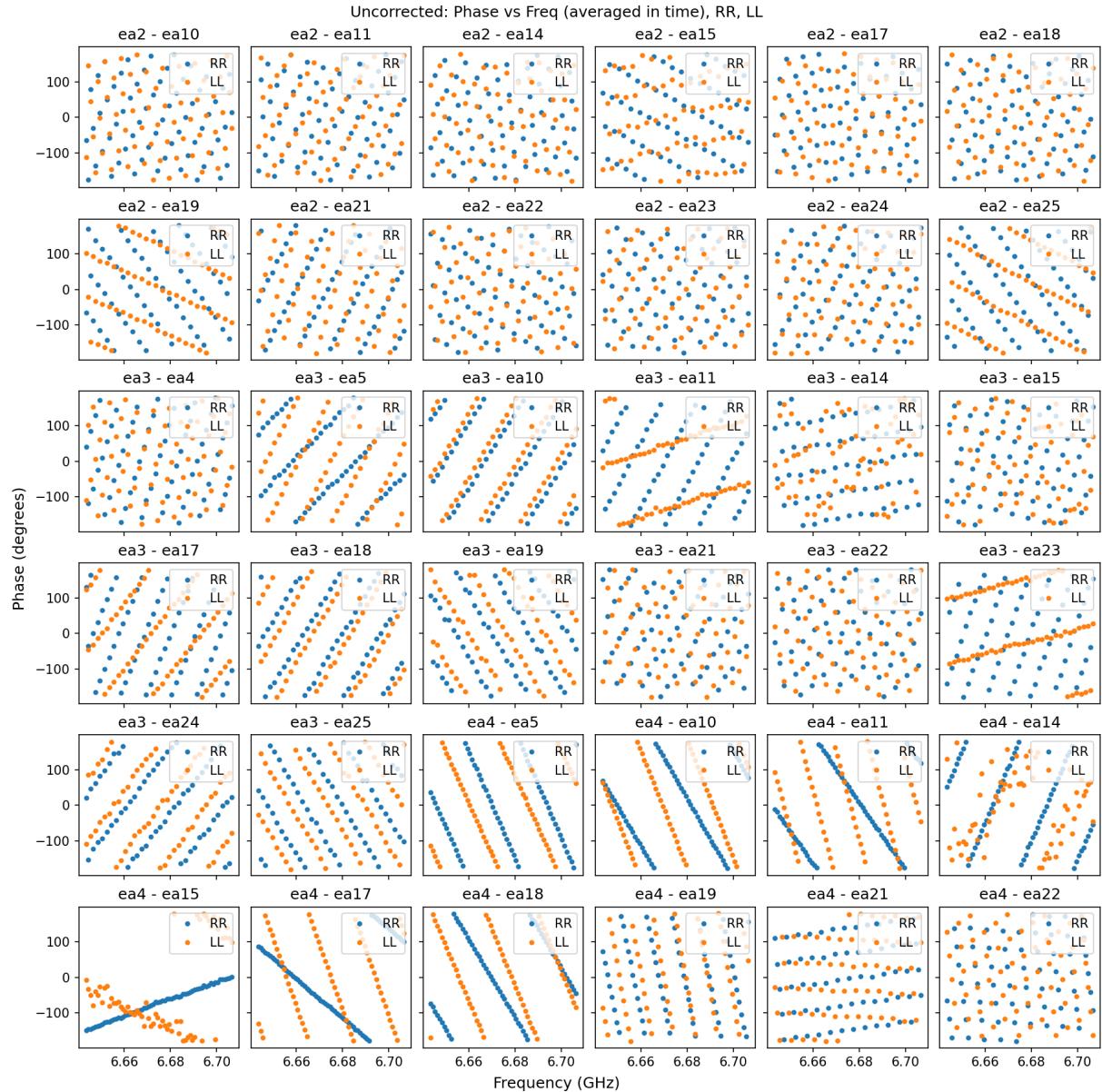
After Calibration:



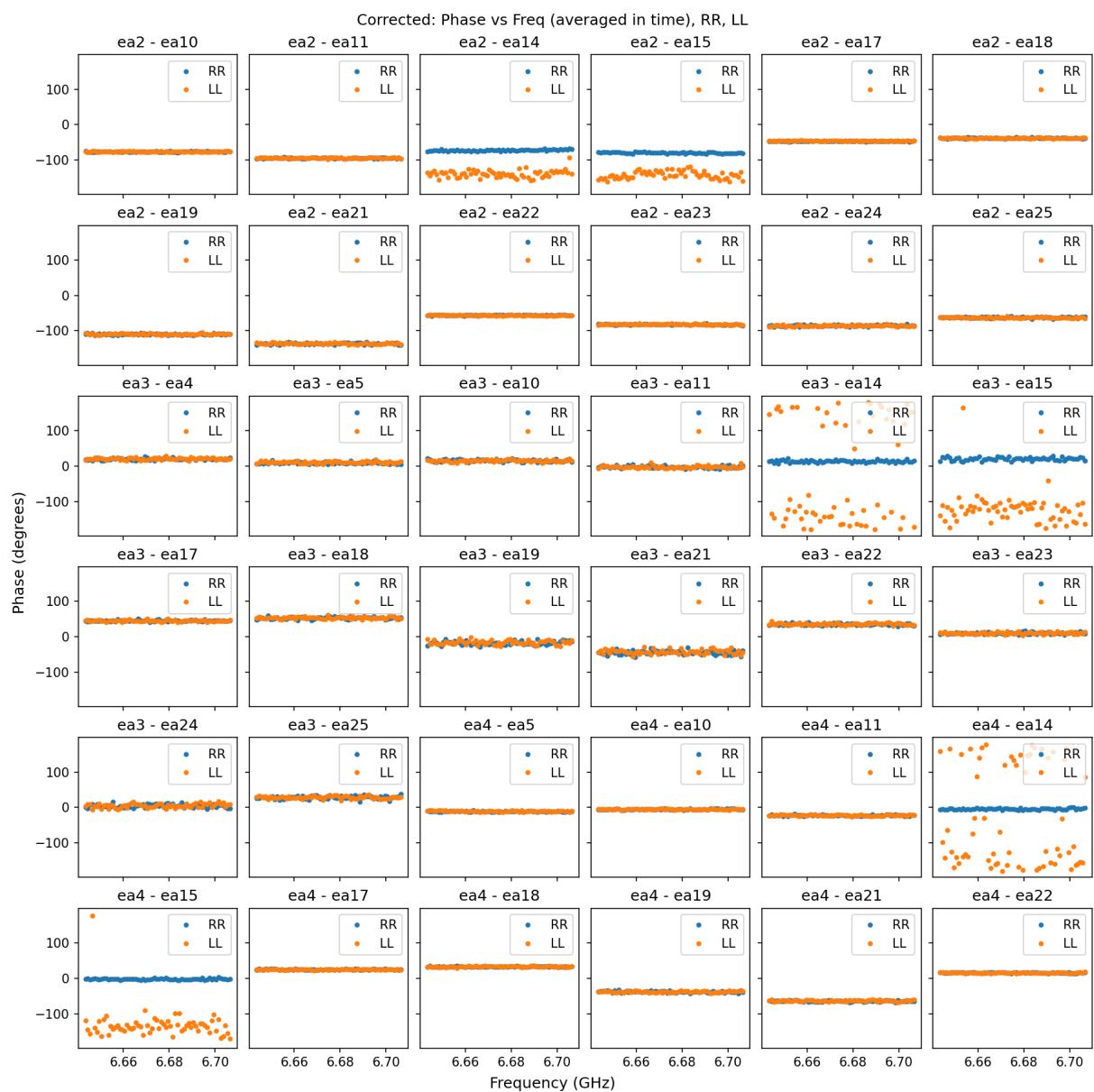
Test 2:

Plots below shows the phases vs freq for J1310+3220 before calibration and after applying the derived calibration from 3C 286

Before calibration:

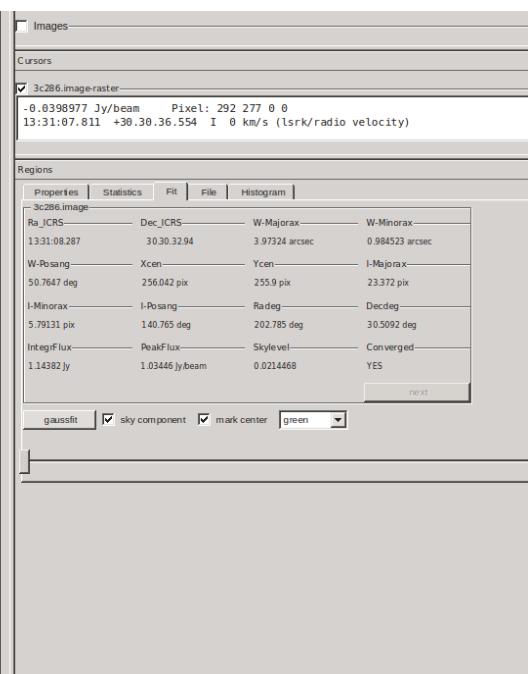
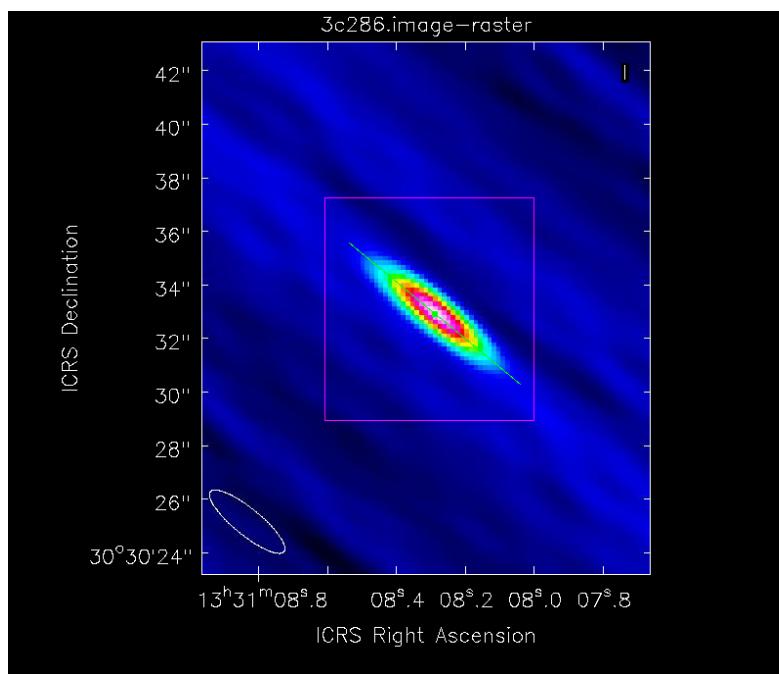
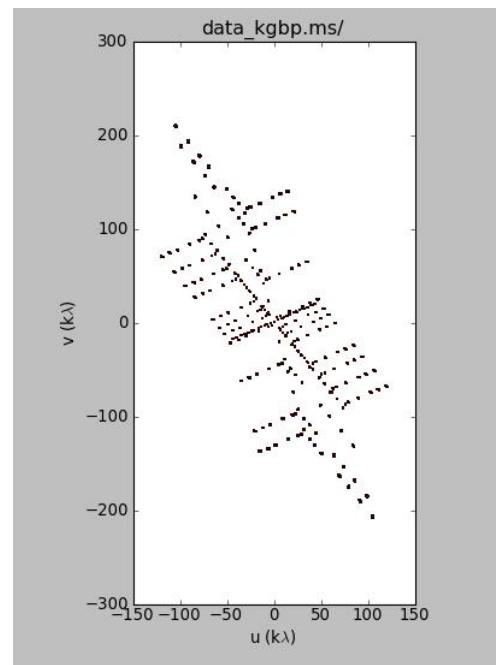
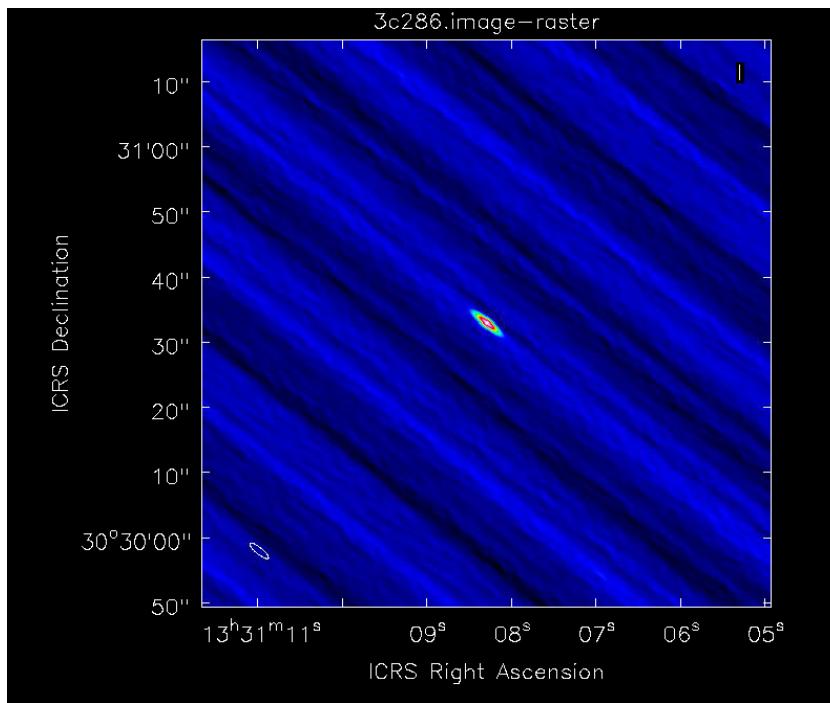


After calibration: The phases are flat and zero across many baselines. However, in some baselines phases deviates from zero. This means the gains solutions from one source to another source are transferable but not perfect. But still good enough to conduct the imaging.



Test 4: The same datasets of 3C286 and J1310+3220 were converted to measurement sets and calibrated in CASA to compare the results with the cosmic gain calibration code. Here, instead of a single step of gain calibration, a combination of delay, bandpass and gain calibration was conducted in CASA. The solutions derived from 3C 286 were applied to the same dataset and the J1310+3220 dataset. Then datasets were imaged using the CASA clean task.

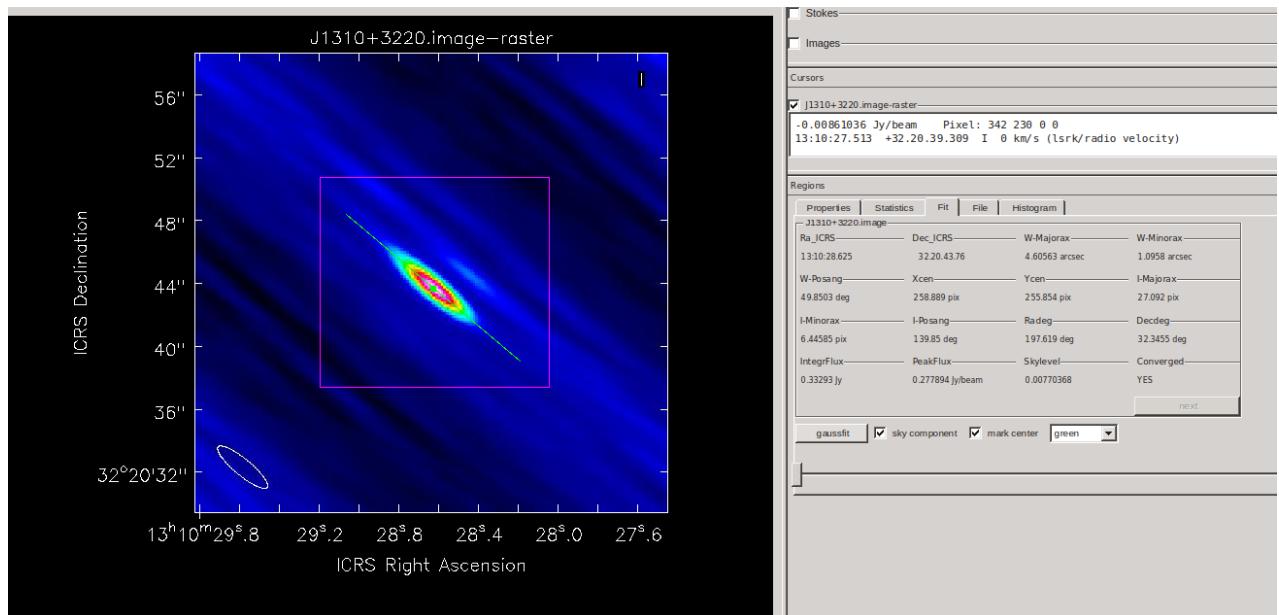
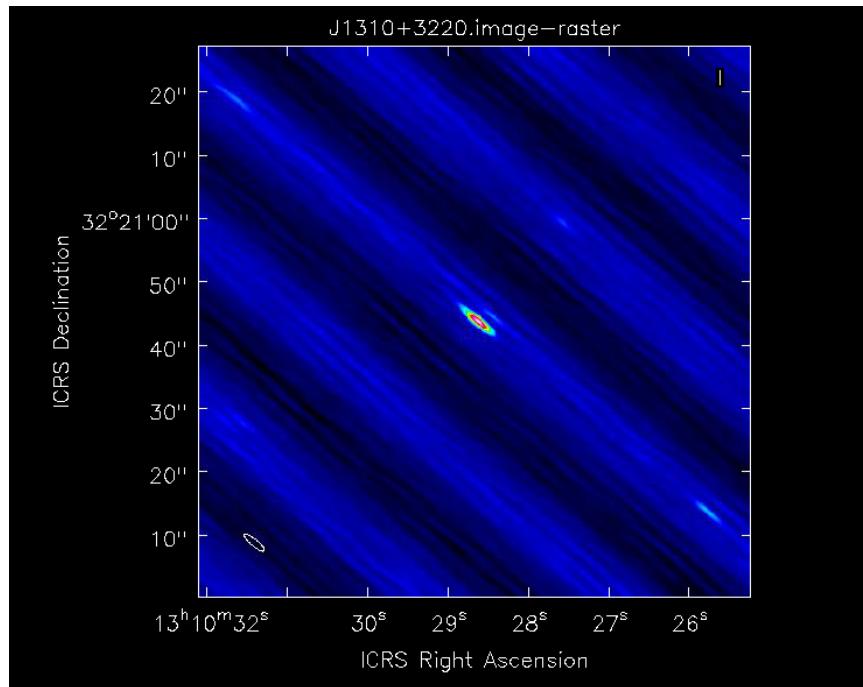
Images of 3C 286:



The data was collected at 6.5GHz for ~140 seconds. One bad antenna and a couple of bad baselines with delay tracking issues were flagged. This resulted in a total of 16 antennas after flagging with 32 MHz of bandwidth. The UV coverage is extended in one direction due to limited time, bandwidth and number of antennas. The clean task was used with a deconvolution of 500 iterations, a cell size of 0.17 arcsec (5 cells across the synthesized beam of 0.86 arcsec) and an image size of 512 pixels.

3C 286 looks like a pointed source, but slightly extended in one direction basically due to the poor UV coverage. The ripples seen across the image are also due to the bad UV coverage. The signal to noise ratio of the source is really good as evident from the image. After doing a 2D gaussian fit on the source, fitted coordinates (13h31m08.287, +30d30m32.94) match pretty well with the NED values (13h31m08.2879s, +30d30m32.958s)

Images of J1310+3220:

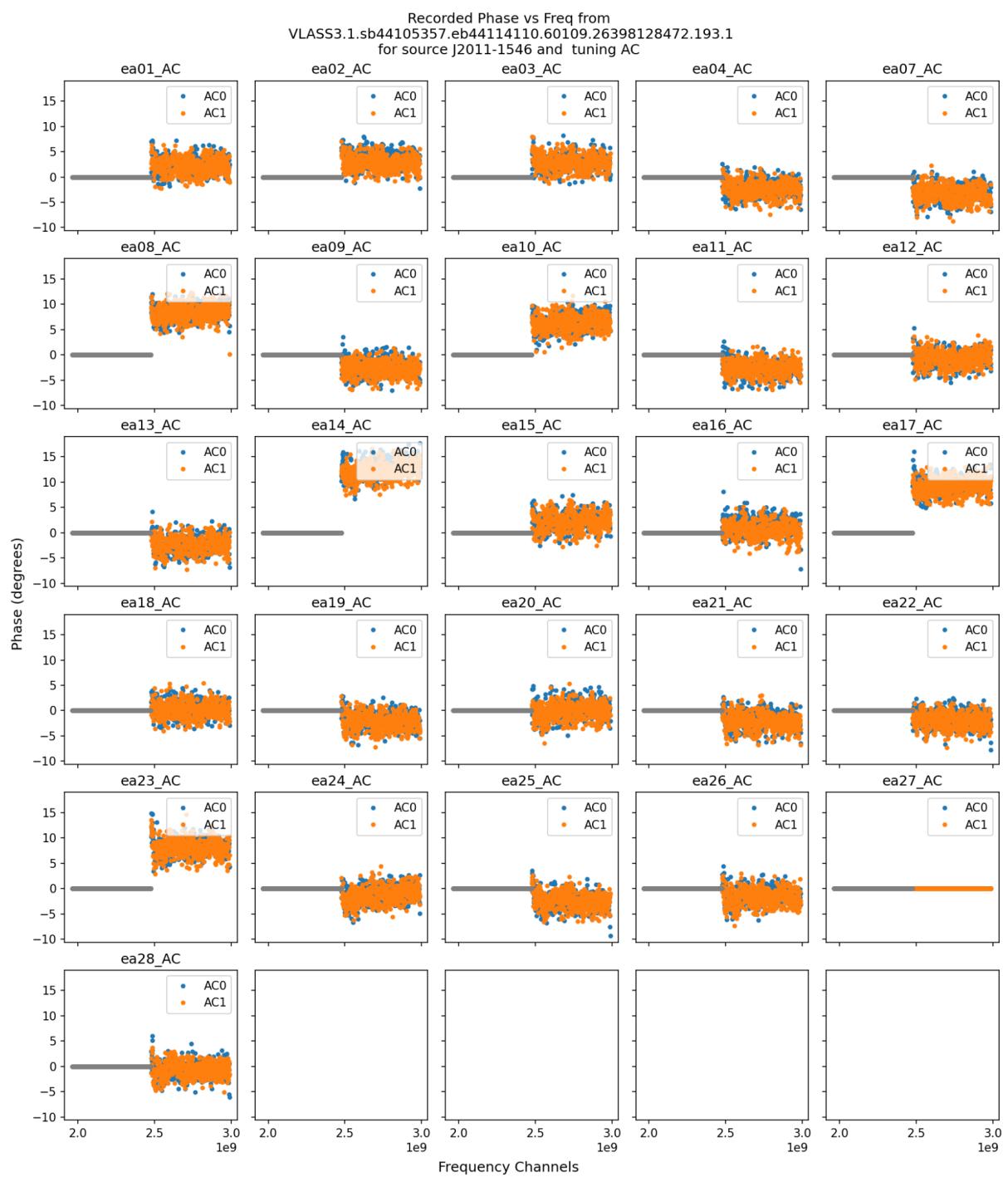


The calibrations derived from 3C286 are applied to the J1310+3220 dataset and it looks good with point source structure, but extended in one direction due to bad UV coverage. The 2D Gaussian fit gives an Ra, Dec [13:10:28.625, 32:20:43.76] and the actual values of radio position from NED are [13:10:28.6638, 32:20:43.783]. These values agree pretty well with each other (only 0.3" difference). This shows that the COSMIC data can be calibrated well with the cosmic gain calibration script and CASA. The astrometry of the calibrators from the back to back observations are pretty good and within 0.3 arcsecs.

7. Online or Real Time Calibration

Once the offline calibration was successful, then back to back sources were observed and the calibrations were applied in real time. This is how the calibration works in real time. The complex gains collected from each compute node is sent to the head node, where it is collated across frequencies in both tunings. A previous phase applied to the FPGAs is subtracted from the gain solutions. Then IFFT of the gains will give the residual delay per antenna. Then a phase factor corresponding to the residual delay is subtracted from the gains to get the phase solutions. The residual delays and phases are applied to FPGAs in real time. The residual delay values and phases will get updated in the subsequent calibration runs

This part was mostly done by Talon, Jack and Paul. I tested the same algorithms on a couple of datasets and they all looked to work fine. However, a spread of phases was noted in some subsequent calibration runs due to unknown reasons. The phases are reset once the dataset Id changes during observations. This removes the bad phases propagated from bad observations with technical issues or the ones affected by RFIs. The calibration slack channel produces a phase vs freq for the gains received per antenna from a calibration dataset which gives an instantaneous view of the quality of calibration. The plot below shows such a plot from the slack channel showing the phases of the gains per antennas in the AC tuning. The following data correspond to a RFI free VLASS observation and it shows close to zero flat phases across frequency which is what we want. Similar results are observed for back to back observations of nearby calibrators as well.



8. Commissioning the BLADE beamformer code

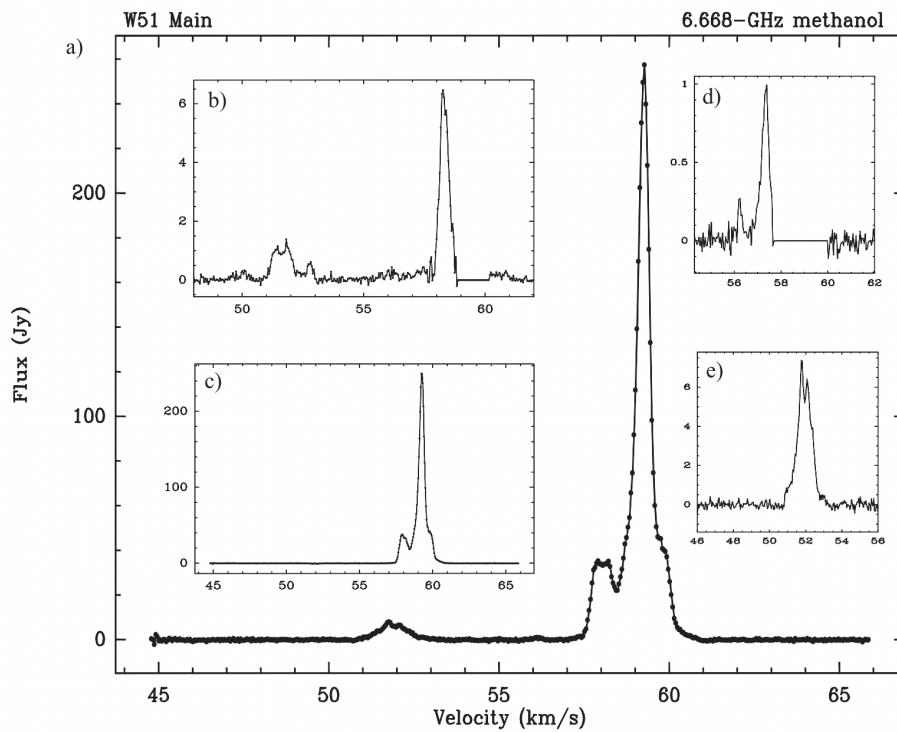
BLADE is the production ready real time beamformer code written in CUDA C++ for COSMIC. A good documentation of BLADE is given [here](#). Some of the commissioning tests of BLADE conducted by Chenoa Tremblay is listed in another document. In the initial days of testing, the BLADE code was not production ready and several fixes were needed in the software. This is the time I realized, there are 1000s of ways a beamformer code can go wrong.

In order to test the findings of the BLADE, I wrote a prototype python based beamformer which is available [here](#). The beamformer takes in the BFR5 file which contains the antenna locations, frequency information, beamforming coefficients, the coordinates to form the beams and the Guppi raw file stem. The results from the prototype code is compared with BLADE to measure the consistency between the 2 codebase.

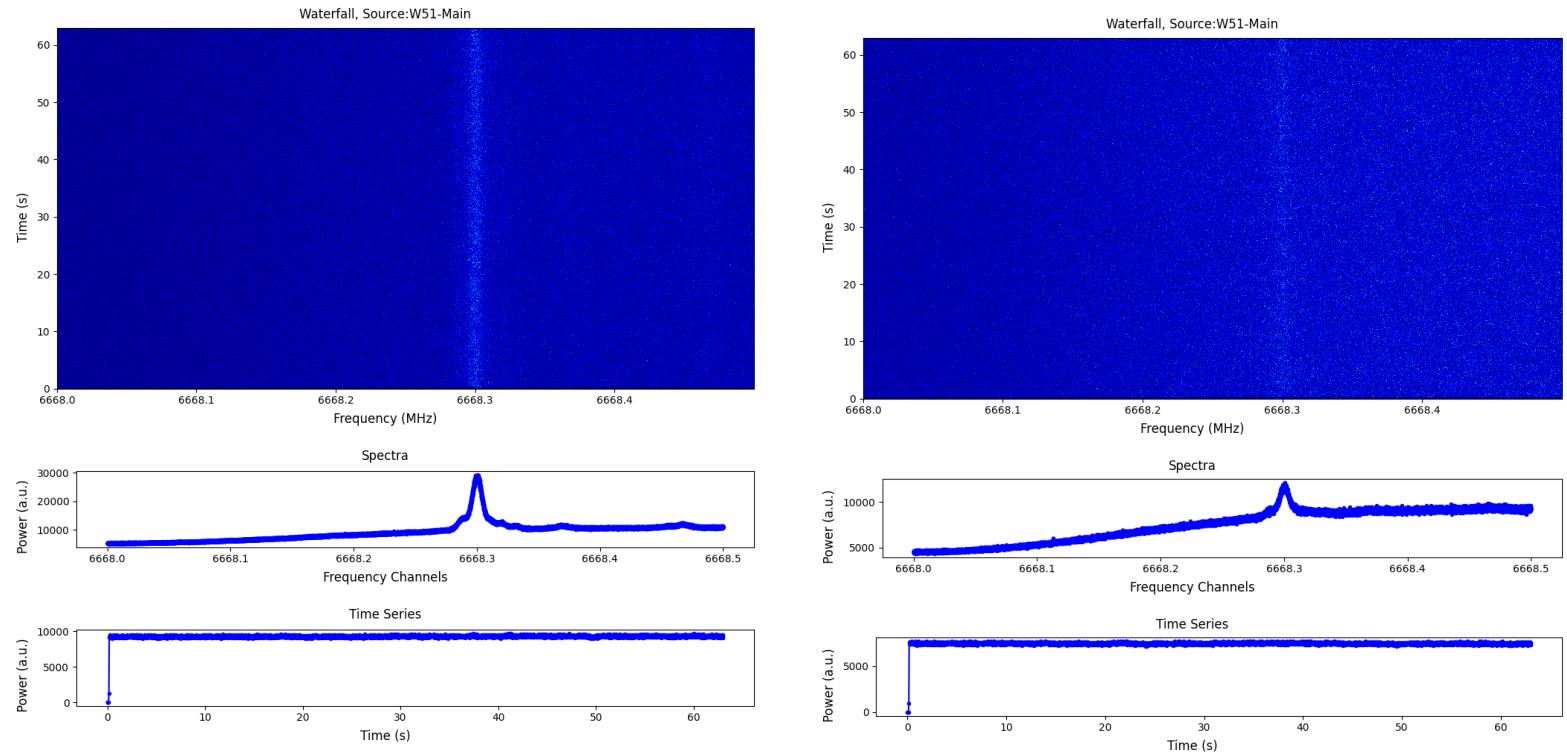
Tests with Methanol Maser:

In order to test the accuracy of the beamformer, we observed a bright W51 methanol maser at 6 GHz, which has a flux density of ~ 274 Jy and ~ 0.5 MHz line width.

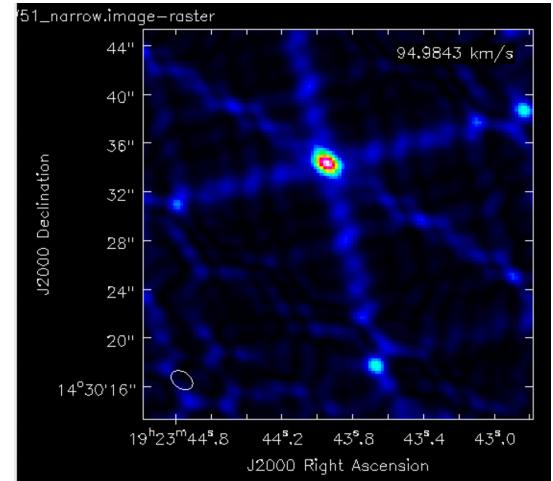
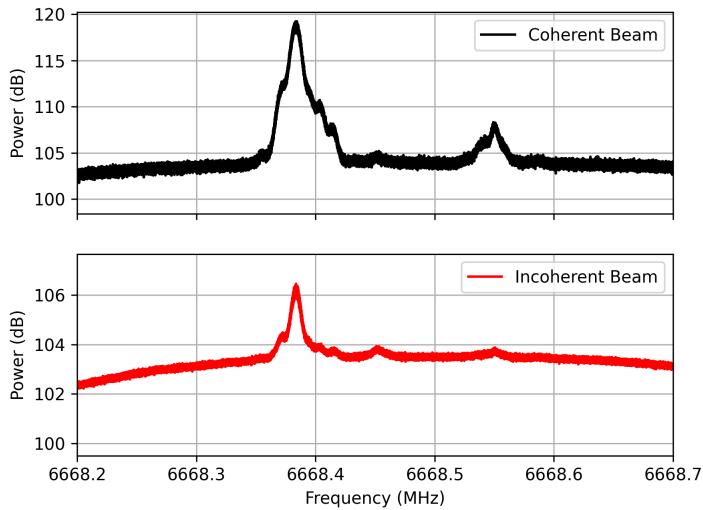
This is the profile of the maser source from S. Etoka, M.D. Gray1 & G.A. Fuller 2012 and it utilized observations from the Merlin telescope. The plot shows the spectra of resolved components with larger spectra showing the main component.



Before running the beamformer on the data, we needed to make sure that all the antennas had detected a maser signal. In order to do that, the associated GUPPI raw files were upchannelized to 30 kHz frequency channels and converted to Stokes I filterbank files for each antenna. The MASER source is bright enough to show up in a single antenna data. The output filterbank files showed that the maser was detected in all the antennas. However, some antennas with low gain detected weak pulses compared to the rest of the antennas. The plots (waterfall, integrated spectra and time series) below shows a bright signal detected in ea08 and a weak signal detected in ea10. The initial dip in the time series is associated with some packet loss in the beginning of the recording.

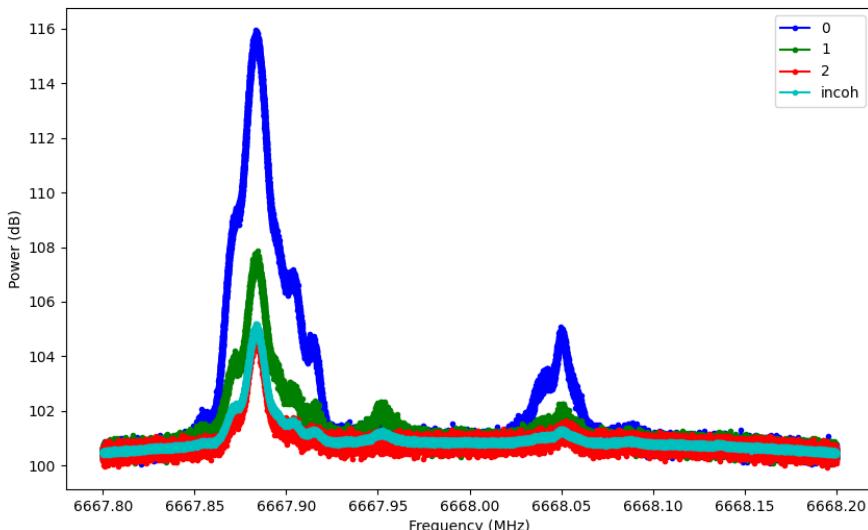


Once we verified that the maser is detected on all antennas, we moved on to the beamforming part. These are the results of BLADE on maser data taken from Cheona's document. The plot below on the left shows the power measured by the coherent beam on the maser position and the incoherent beam. The plot on the right is an image of the Maser made by Chenoa showing a really good bright point source.



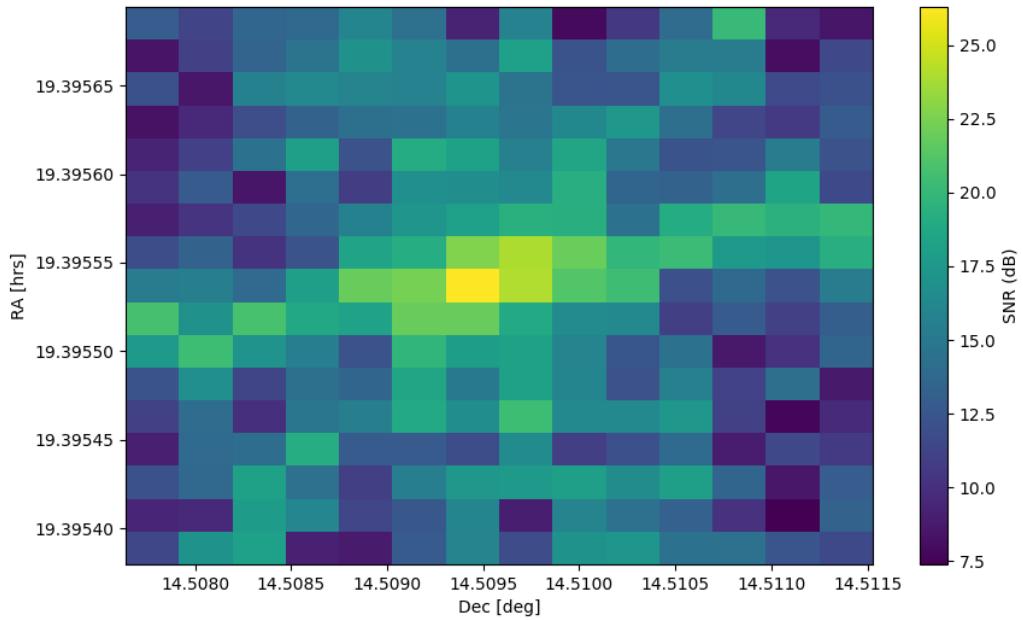
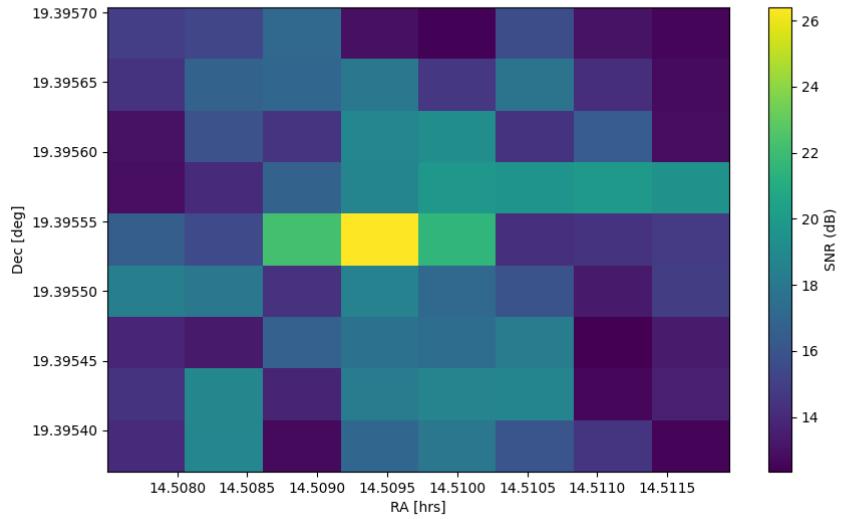
When the observation of the maser took place, the boresight of the telescope was pointed towards the maser coordinates from the literature RA, Dec: (19:23:43.873, +14:30:29.453 J2000). However due to small calibration errors the peak point of the maser from the image was obtained at RA, Dec (19:23:43.94, +14:30:34.34) which is a couple of arcsec away. At C band and VLA B configuration, this is a couple of beams away and the beamforming coordinates were updated accordingly. The plot below shows the spectra from beams pointed at different locations.

- 0 = Updated location of the maser source
- 1 = boresight where we thought maser would be
- 2 = point at edge of primary beam/FoV
- Incoh = incoherent beam



We observe ~11.5 dB difference between the coherent and incoherent beam even though ideally we expect around 13 dB ($10 \cdot \log_{10}(20)$ antennas) difference. The incoherent beam could be picking up other point sources within the FoV. The power in the beam couple of arcsec away is going down by 8 dB and power of the beam at the edge of primary FoV goes down by 12 dB. These results agrees with the observations from BLADE.

Along with the multi point beamforming, a raster scan was conducted around the maser to map out the point source structure. The plots below show the power measured in 2 raster scans with different beam resolutions.



In both the images, we can clearly see a point source and it matches with the image of the maser made from the visibilities. The dirty beam structures associated with the maser are clearly seen in the high resolution map. BLADE was having issues in retrieving the point source structure. After some bug fixes, raster scan from the

python code matched with the BLADE version. Currently the BLADE has a pointing accuracy less than 1/s

Beamformer Test with setigen data:

In order to test the efficiency of the beamformer, we conducted tests on simulated data generated by the setigen software which is discussed in the first session. Raw files similar to COSMIC were generated and narrowband drifting signals were injected into the raw files at specific frequencies. The idea was to study how the beamformed signals change as the beam is pointed at different directions in the sky. During the study we realized that the raw data generated from setigen does not incorporate the effects of sky or the primary beam pattern of the telescopes. Due to this reason, the results from beamforming tests on simulated data were inconclusive.