

# PROBLEMS OF QUANTUM FIELD THEORIES IN CURVED SPACETIMES

A MASTER THESIS

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# Preface

Preface

# 1 Introduction to QFT in curved spacetimes

$$dl^2 = g_{\mu\nu} dx^\mu dx^\nu \quad (1.1)$$

$$S = \int \left[ \frac{1}{2\kappa} (R - 2\Lambda) + \mathcal{L}_M \right] \sqrt{-g} d^4x \quad (1.2)$$

$$\kappa \equiv \frac{8\pi G}{c^4}$$

Variation of  $S$  with respect to the inverse metric ( $g^{\mu\nu}$ ) gives

$$\delta S = \int \left[ \frac{\sqrt{-g}}{2\kappa} \frac{\delta R}{\delta g^{\mu\nu}} + \frac{R}{2\kappa} \frac{1}{\sqrt{-g}} \frac{\delta \sqrt{-g}}{\delta g^{\mu\nu}} - \frac{\Lambda}{\kappa} \frac{1}{\sqrt{-g}} \frac{\delta \sqrt{-g}}{\delta g^{\mu\nu}} + \frac{\delta \mathcal{L}_M}{\delta g^{\mu\nu}} + \frac{\mathcal{L}_M}{\sqrt{-g}} \frac{\delta \sqrt{-g}}{\delta g^{\mu\nu}} \right] \delta g^{\mu\nu} \sqrt{-g} d^4x \quad (1.3)$$

$\delta S = 0$  and

$$\frac{\delta R}{\delta g^{\mu\nu}} = R_{\mu\nu} \quad \frac{1}{\sqrt{-g}} \frac{\delta \sqrt{-g}}{\delta g^{\mu\nu}} = -\frac{1}{2} g_{\mu\nu} \quad (1.4)$$

$$R_{\mu\nu} - \frac{1}{2} g_{\mu\nu} R + \Lambda g_{\mu\nu} = -2 \frac{8\pi G}{c^4} \left( \frac{\delta \mathcal{L}_M}{\delta g^{\mu\nu}} - \frac{1}{2} \mathcal{L}_M g_{\mu\nu} \right) \quad (1.5)$$

$$T_{\mu\nu} \equiv \mathcal{L}_M g_{\mu\nu} - \frac{\delta \mathcal{L}_M}{\delta g^{\mu\nu}} = \frac{-2}{\sqrt{-g}} \frac{\delta (\mathcal{L}_M \sqrt{-g})}{\delta g^{\mu\nu}} \quad (1.6)$$

[1]

$$\nabla_\mu T^{\mu\nu} = 0 \quad (1.7)$$

(and its symmetric)

Now, in order to obtain the equations of motion for the matter fields, consider the lagrangian

$$\mathcal{L}_M = \mathcal{L}_M[\phi^\alpha(x), \nabla_\mu \phi^\alpha(x)] \quad (1.8)$$

variations of  $S$  with respect of  $\phi^\alpha$  result in

$$\delta S = \int \left[ \frac{\partial \mathcal{L}_M}{\partial \phi^\alpha} \delta \phi^\alpha + \frac{\partial \mathcal{L}_M}{\partial (\nabla_\mu \phi^\alpha)} \nabla_\mu (\delta \phi^\alpha) \right] \sqrt{-g} d^4x \quad (1.9)$$

and thus, after applying the generalized Gauss Theorem on a curved background, and considering that field variations vanish at the boundaries, one obtains

$$\frac{\partial \mathcal{L}_M}{\partial \phi^\alpha} - \nabla_\mu \left[ \frac{\partial \mathcal{L}_M}{\partial (\nabla_\mu \phi^\alpha)} \right] = 0 \quad (1.10)$$

$$\Pi_\alpha \equiv \frac{\partial \mathcal{L}_M}{\partial (\nabla_0 \phi^\alpha)} \quad (1.11)$$

$$S[\phi] = \int \frac{1}{2} \left[ \nabla_\nu \phi \nabla^\nu \phi - \mu^2 \phi^2 - \xi R \phi^2 \right] \sqrt{-g} d^4x \quad (1.12)$$

$$[\nabla_\nu \nabla^\nu - \mu^2 - \xi R] \phi = 0 \quad (1.13)$$

$$T_{\mu\nu} = \nabla_\mu \phi \nabla_\nu \phi - \frac{1}{2} g_{\mu\nu} [\nabla^\sigma \phi \nabla_\sigma \phi - \mu^2 \phi^2] + \xi \left[ R_{\mu\nu} - \frac{1}{2} g_{\mu\nu} + g_{\mu\nu} \nabla^\sigma \nabla_\sigma - \nabla_\mu \nabla_\nu \right] \phi^2 \quad (1.14)$$

note that for minimally coupled field ( $\xi = 0$ ) the energy-momentum tensor is equivalent to the Noether energy-momentum tensor.

## 2 Scalar field in an expanding universe

FLRW metric

$$dl^2 = c^2 dt^2 - a^2(t) \left[ \frac{dr^2}{1 - \kappa r^2} + r^2 d\Omega^2 \right] \quad (2.1)$$

Weyl tensor =0 therefore the metric is conformally flat, i.e. independently of the curvature  $\kappa$  there must exist a coordinate system where

$$dl^2 = a(t) \eta_{\mu\nu} dx^\mu dx^\nu = a(t) [c^2 dt^2 - d\mathbf{x}^2] \quad (2.2)$$

the standard action describing the dynamics of a (non-minimally coupled to gravity) real scalar field is

$$s = \int \frac{1}{2} \left[ \nabla_\nu \phi \nabla^\nu \phi - \mu^2 \phi^2 - \xi R \phi^2 \right] \sqrt{-g} d^4x \quad (2.3)$$

$$\sqrt{-g} = a^4 \quad \chi = a\phi$$

$$s = \int \frac{1}{2} \left[ \partial_\nu \chi \partial^\nu \chi - \left( \mu^2 a^2 + \xi R a^2 - c^2 \frac{a''}{a} \right) \chi^2 - \partial_t \left( c^2 \chi^2 \frac{a'}{a} \right) \right] d^4x \quad (2.4)$$

dropping the time derivative

$$s = \int \frac{1}{2} \left[ \partial_\nu \chi \partial^\nu \chi - \left( \mu^2 a^2 + \xi R a^2 - c^2 \frac{a''}{a} \right) \chi^2 \right] d^4x \quad (2.5)$$

by Euler-Lagrange

$$[\partial_\nu \partial^\nu + \mu_{\text{eff}}^2(t)] \chi = 0 \quad (2.6)$$

where

$$\mu_{\text{eff}}^2(t) = (\mu^2 + \xi R) a^2 - c^2 \frac{a''}{a} \quad (2.7)$$

solutions of previous equation have the form

$$\chi = a v(t) e^{\pm i \mathbf{k} \mathbf{x} \hbar^{-1}} \quad (2.8)$$

meaning that, the dispersion relation is

$$v'' \hbar^2 + \omega^2(t) v = 0 \quad (2.9)$$

where  $\omega(t)$  is defined as

$$\omega^2(t) = \mathbf{k}^2 + \hbar^2 \mu_{\text{eff}}^2(t) = \mathbf{k}^2 + (m^2 c^2 + \xi \hbar^2 R) a(t) - \hbar^2 c^2 \frac{a''}{a} \quad (2.10)$$

now, proof that  $\text{Im}(vv'^*)$  is constant through time

$$\frac{\partial}{\partial t} \text{Im}(vv'^*) = \frac{\partial}{\partial t} \left( \frac{vv'^* - v^*v'}{2i} \right) = \frac{vv''^* - v^*v''}{2i} = 0 \quad (2.11)$$

last step is result from dispersion relation. Since dispersion relation is scalable by a time independent function,  $\text{Im}(v'v^*)$  can be determined to be a chosen value, a particular useful choice is to consider it momentum independent.  $\text{Im}(v'v^*) = W[v, v^*]$  therefore, if its not equal to 0, they are linearly independent solutions to dispersion relation.

The most general solution to the main equation is

$$\chi = \int \frac{d^3 \mathbf{k}}{(2\pi \hbar)^3} \left[ a_{\mathbf{k}} v_{\mathbf{k}}(t) e^{i \mathbf{k} \mathbf{x} \hbar^{-1}} + a_{\mathbf{k}}^* v_{\mathbf{k}}^*(t) e^{-i \mathbf{k} \mathbf{x} \hbar^{-1}} \right] \quad (2.12)$$

## 2 Scalar field in an expanding universe

The field  $\chi$  and its conjugate momentum  $\Pi = \partial_{ct}\chi$  are promoted to operators on the quantum Hilbert space, with the standar canonical conmutation relations

$$[\hat{\chi}(t, \mathbf{x}), \hat{\Pi}(t, \mathbf{y})] = i\hbar \delta^3(\mathbf{x} - \mathbf{y}) \quad (2.13)$$

$$[\hat{\chi}(t, \mathbf{x}), \hat{\chi}(t, \mathbf{y})] = [\hat{\Pi}(t, \mathbf{x}), \hat{\Pi}(t, \mathbf{y})] = 0 \quad (2.14)$$

where the operational nature of the fields arrise from the promotion of the mode amplitudes, i.e.

$$a_{\mathbf{k}} \longrightarrow \hat{a}_{\mathbf{k}} \quad a_{\mathbf{k}}^* \longrightarrow \hat{a}_{\mathbf{k}}^\dagger \quad (2.15)$$

this operators fulfill the following conmutation relations

$$[\hat{a}_{\mathbf{k}}, \hat{a}_{\mathbf{q}}^\dagger] = \frac{(2\pi\hbar)^3 \hbar c}{2\text{Im}(v'v^*)} \delta^3(\mathbf{k} - \mathbf{q}), \quad [\hat{a}_{\mathbf{k}}, \hat{a}_{\mathbf{q}}] = [\hat{a}_{\mathbf{k}}^\dagger, \hat{a}_{\mathbf{q}}^\dagger] = 0 \quad (2.16)$$

(note that  $\hat{a}_{\mathbf{k}} \neq \hat{a}_{-\mathbf{k}}$ )

To prove this, consider that

$$\begin{aligned} [\hat{\chi}(\mathbf{x}), \hat{\Pi}(\mathbf{y})] &= \frac{1}{c} \int \frac{d^3\mathbf{k} d^3\mathbf{q}}{(2\pi\hbar)^6} \left\{ [\hat{a}_{\mathbf{k}}, \hat{a}_{\mathbf{q}}] v_{\mathbf{k}} v'_{\mathbf{q}} e^{i(\mathbf{k}\mathbf{x} + \mathbf{q}\mathbf{y})\hbar^{-1}} + [\hat{a}_{\mathbf{k}}^\dagger, \hat{a}_{\mathbf{q}}^\dagger] v_{\mathbf{k}}^* v'_{\mathbf{q}} e^{i(\mathbf{k}\mathbf{x} - \mathbf{q}\mathbf{y})\hbar^{-1}} + \right. \\ &\quad \left. + [\hat{a}_{\mathbf{k}}, \hat{a}_{\mathbf{q}}^\dagger] v_{\mathbf{k}} v_{\mathbf{q}}^* e^{i(\mathbf{k}\mathbf{x} - \mathbf{q}\mathbf{y})\hbar^{-1}} - [\hat{a}_{\mathbf{q}}, \hat{a}_{\mathbf{k}}^\dagger] v_{\mathbf{k}}^* v'_{\mathbf{q}} e^{-i(\mathbf{k}\mathbf{x} - \mathbf{q}\mathbf{y})\hbar^{-1}} \right\} \end{aligned} \quad (2.17)$$

if the operators  $\hat{a}$  and  $\hat{a}^\dagger$  are to be understood as creation and annihilation operators, they must fulfill

$$[\hat{a}_{\mathbf{k}}, \hat{a}_{\mathbf{q}}^\dagger] = \alpha \delta^3(\mathbf{k} - \mathbf{q}), \quad [\hat{a}_{\mathbf{k}}, \hat{a}_{\mathbf{q}}] = [\hat{a}_{\mathbf{k}}^\dagger, \hat{a}_{\mathbf{q}}^\dagger] = 0 \quad (2.18)$$

where  $\alpha \in \mathbb{C}$ , and thus

$$[\hat{\chi}(\mathbf{x}), \hat{\Pi}(\mathbf{y})] = \frac{\alpha}{c} \int \frac{d^3\mathbf{k}}{(2\pi\hbar)^6} 2i\text{Im}(v_{\mathbf{k}} v_{\mathbf{k}}^*) e^{i(\mathbf{k}\mathbf{x} - \mathbf{q}\mathbf{y})\hbar^{-1}} \quad (2.19)$$

considering  $\text{Im}(v'v^*)$  momentum independent, and remembering the canonical conmutation relations, one finds that

$$\alpha \text{Im}(vv^*) = \frac{1}{2} \hbar c (2\pi\hbar)^3 \quad (2.20)$$

The hamiltonian

$$\hat{\mathcal{H}}(t) = \int \frac{c}{2} \left[ \hat{\Pi}^2 + (\nabla \hat{\chi})^2 + \mu_{\text{eff}}^2(t) \hat{\chi}^2 \right] d^3\mathbf{x} \quad (2.21)$$

$$\begin{aligned} \hat{\Pi}^2 &= \frac{1}{c^2} \int \frac{d^3\mathbf{k} d^3\mathbf{q}}{(2\pi\hbar)^6} \left[ \hat{a}_{\mathbf{k}} \hat{a}_{\mathbf{q}} v'_{\mathbf{k}} v'_{\mathbf{q}} e^{i(\mathbf{k} + \mathbf{q})\mathbf{x}\hbar^{-1}} + \hat{a}_{\mathbf{k}} \hat{a}_{\mathbf{q}}^\dagger v'_{\mathbf{k}} v_{\mathbf{q}}^* e^{i(\mathbf{k} - \mathbf{q})\mathbf{x}\hbar^{-1}} + \right. \\ &\quad \left. + \hat{a}_{\mathbf{k}}^\dagger \hat{a}_{\mathbf{q}} v_{\mathbf{k}}^* v'_{\mathbf{q}} e^{-i(\mathbf{k} - \mathbf{q})\mathbf{x}\hbar^{-1}} + \hat{a}_{\mathbf{k}}^\dagger \hat{a}_{\mathbf{q}}^\dagger v_{\mathbf{k}}^* v_{\mathbf{q}}^* e^{-i(\mathbf{k} + \mathbf{q})\mathbf{x}\hbar^{-1}} \right] \end{aligned} \quad (2.22)$$

$$\begin{aligned} (\nabla \hat{\chi})^2 &= -\frac{1}{\hbar^2} \int \frac{d^3\mathbf{k} d^3\mathbf{q}}{(2\pi\hbar)^6} \mathbf{k}\mathbf{q} \left[ \hat{a}_{\mathbf{k}} \hat{a}_{\mathbf{q}} v_{\mathbf{k}} v_{\mathbf{q}} e^{i(\mathbf{k} + \mathbf{q})\mathbf{x}\hbar^{-1}} - \hat{a}_{\mathbf{k}} \hat{a}_{\mathbf{q}}^\dagger v_{\mathbf{k}} v_{\mathbf{q}}^* e^{i(\mathbf{k} - \mathbf{q})\mathbf{x}\hbar^{-1}} - \right. \\ &\quad \left. - \hat{a}_{\mathbf{k}}^\dagger \hat{a}_{\mathbf{q}} v_{\mathbf{k}}^* v_{\mathbf{q}} e^{-i(\mathbf{k} - \mathbf{q})\mathbf{x}\hbar^{-1}} + \hat{a}_{\mathbf{k}}^\dagger \hat{a}_{\mathbf{q}}^\dagger v_{\mathbf{k}}^* v_{\mathbf{q}}^* e^{-i(\mathbf{k} + \mathbf{q})\mathbf{x}\hbar^{-1}} \right] \end{aligned} \quad (2.23)$$

$$\begin{aligned} \hat{\chi}^2 &= \int \frac{d^3\mathbf{k} d^3\mathbf{q}}{(2\pi\hbar)^6} \left[ \hat{a}_{\mathbf{k}} \hat{a}_{\mathbf{q}} v_{\mathbf{k}} v_{\mathbf{q}} e^{i(\mathbf{k} + \mathbf{q})\mathbf{x}\hbar^{-1}} + \hat{a}_{\mathbf{k}} \hat{a}_{\mathbf{q}}^\dagger v_{\mathbf{k}} v_{\mathbf{q}}^* e^{i(\mathbf{k} - \mathbf{q})\mathbf{x}\hbar^{-1}} + \right. \\ &\quad \left. + \hat{a}_{\mathbf{k}}^\dagger \hat{a}_{\mathbf{q}} v_{\mathbf{k}}^* v_{\mathbf{q}} e^{-i(\mathbf{k} - \mathbf{q})\mathbf{x}\hbar^{-1}} + \hat{a}_{\mathbf{k}}^\dagger \hat{a}_{\mathbf{q}}^\dagger v_{\mathbf{k}}^* v_{\mathbf{q}}^* e^{-i(\mathbf{k} + \mathbf{q})\mathbf{x}\hbar^{-1}} \right] \end{aligned} \quad (2.24)$$

$$\begin{aligned}
 \hat{\mathcal{H}} = \frac{c}{2} \int \frac{d^3\mathbf{k}d^3\mathbf{q}}{(2\pi\hbar)^3} \left\{ \hat{a}_{\mathbf{k}}\hat{a}_{\mathbf{q}} \left[ \frac{1}{c^2}v'_{\mathbf{k}}v'_{\mathbf{q}} - \left( \frac{1}{\hbar^2}\mathbf{k}\mathbf{q} - \mu_{\text{eff}}^2 \right) v_{\mathbf{k}}v_{\mathbf{q}} \right] \delta^3(\mathbf{k} + \mathbf{q}) + \right. \\
 + \hat{a}_{\mathbf{k}}\hat{a}_{\mathbf{q}}^\dagger \left[ \frac{1}{c^2}v'_{\mathbf{k}}v'_{\mathbf{q}}^* + \left( \frac{1}{\hbar^2}\mathbf{k}\mathbf{q} + \mu_{\text{eff}}^2 \right) v_{\mathbf{k}}v_{\mathbf{q}}^* \right] \delta^3(\mathbf{k} - \mathbf{q}) + \\
 + \hat{a}_{\mathbf{k}}^\dagger\hat{a}_{\mathbf{q}} \left[ \frac{1}{c^2}v_{\mathbf{k}}^*v'_{\mathbf{q}} + \left( \frac{1}{\hbar^2}\mathbf{k}\mathbf{q} + \mu_{\text{eff}}^2 \right) v_{\mathbf{k}}^*v_{\mathbf{q}} \right] \delta^3(\mathbf{k} - \mathbf{q}) + \\
 \left. + \hat{a}_{\mathbf{k}}^\dagger\hat{a}_{\mathbf{q}}^\dagger \left[ \frac{1}{c^2}v_{\mathbf{k}}^*v_{\mathbf{q}}^* - \left( \frac{1}{\hbar^2}\mathbf{k}\mathbf{q} - \mu_{\text{eff}}^2 \right) v_{\mathbf{k}}^*v_{\mathbf{q}}^* \right] \delta^3(\mathbf{k} + \mathbf{q}) \right\} \quad (2.25)
 \end{aligned}$$

$$\begin{aligned}
 \hat{\mathcal{H}} = \frac{c}{2} \int \frac{d^3\mathbf{k}}{(2\pi\hbar)^3} \left\{ \hat{a}_{\mathbf{k}}\hat{a}_{-\mathbf{k}} \left[ \frac{1}{c^2}v'_{\mathbf{k}}v'_{\mathbf{k}} + \frac{1}{\hbar^2}\omega_{\mathbf{k}}^2(t)v_{\mathbf{k}}v_{\mathbf{k}} \right] + \right. \\
 + \hat{a}_{\mathbf{k}}\hat{a}_{\mathbf{k}}^\dagger \left[ \frac{1}{c^2}v'_{\mathbf{k}}v'_{\mathbf{k}}^* + \frac{1}{\hbar^2}\omega_{\mathbf{k}}^2(t)v_{\mathbf{k}}v_{\mathbf{k}}^* \right] + \\
 + \hat{a}_{\mathbf{k}}^\dagger\hat{a}_{\mathbf{k}} \left[ \frac{1}{c^2}v_{\mathbf{k}}^*v'_{\mathbf{k}} + \frac{1}{\hbar^2}\omega_{\mathbf{k}}^2(t)v_{\mathbf{k}}^*v_{\mathbf{k}} \right] + \\
 \left. + \hat{a}_{\mathbf{k}}^\dagger\hat{a}_{-\mathbf{k}}^\dagger \left[ \frac{1}{c^2}v_{\mathbf{k}}^*v_{-\mathbf{k}}^* + \frac{1}{\hbar^2}\omega_{\mathbf{k}}^2(t)v_{\mathbf{k}}^*v_{-\mathbf{k}}^* \right] \right\} \quad (2.26)
 \end{aligned}$$

$$\hat{\mathcal{H}} = \frac{c}{2} \int \frac{d^3\mathbf{k}}{(2\pi\hbar)^3} \left[ \hat{a}_{\mathbf{k}}\hat{a}_{-\mathbf{k}}F_{\mathbf{k}} + \hat{a}_{\mathbf{k}}^\dagger\hat{a}_{-\mathbf{k}}^\dagger F_{\mathbf{k}}^* + \left( 2\hat{a}_{\mathbf{k}}^\dagger\hat{a}_{\mathbf{k}} + \frac{(2\pi\hbar)^3\hbar c}{2\text{Im}(v'v^*)}\delta^3(\mathbf{0}) \right) E_{\mathbf{k}} \right] \quad (2.27)$$

where

$$F_{\mathbf{k}}(t) = \left( \frac{1}{\hbar c} \right)^2 \left[ \hbar^2 v_{\mathbf{k}}'^2 + \omega_{\mathbf{k}}^2(t) c^2 v_{\mathbf{k}}^2 \right] \quad (2.28)$$

$$E_{\mathbf{k}}(t) = \left( \frac{1}{\hbar c} \right)^2 \left[ \hbar^2 |v_{\mathbf{k}}'|^2 + \omega_{\mathbf{k}}^2(t) c^2 |v_{\mathbf{k}}|^2 \right] \quad (2.29)$$

Now, the expectation value of the hamiltonian at time  $t_0$  in the state  $|(v)0\rangle$

$$\langle (v)0 | \hat{\mathcal{H}}(t_0) | (v)0 \rangle = \rho(t_0)\delta^3(\mathbf{0}) = \frac{\hbar c^2 \delta^3(\mathbf{0})}{4\text{Im}(v'v^*)} \int d^3\mathbf{k} E_{\mathbf{k}} \quad (2.30)$$

To minimise the energy density of de vacuum state is to fin the set of functions  $v_{\mathbf{k}}$  that minimise  $E_{\mathbf{k}}$ . Suppose that  $v_{\mathbf{k}}$  can be written as

$$v_{\mathbf{k}} = r_{\mathbf{k}} e^{i\alpha_{\mathbf{k}}} \quad (2.31)$$

since  $\text{Im}(vv'^*)$  was constant through time

$$\text{Im}(v_{\mathbf{k}}v_{\mathbf{k}}'^*) = -r_{\mathbf{k}}^2 \alpha_{\mathbf{k}}' \quad (2.32)$$

this means

$$E_{\mathbf{k}} = \left( \frac{1}{\hbar c} \right)^2 \left\{ \hbar^2 \left[ r_{\mathbf{k}}'^2 + \text{Im}^2(v_{\mathbf{k}}v_{\mathbf{k}}'^*) \frac{1}{r_{\mathbf{k}}^2} \right] + \omega_{\mathbf{k}}^2 c^2 r_{\mathbf{k}}^2 \right\} \quad (2.33)$$

the minimum of this function must fulfil  $r_{\mathbf{k}}'(t_0) = 0$ . Now, if  $\omega_{\mathbf{k}}^2(t_0)$  and  $\text{Im}(v_{\mathbf{k}}v_{\mathbf{k}}'^*)$  have the same sign, the minimum of  $E_{\mathbf{k}}$  happens when  $r_{\mathbf{k}}(t_0) = \left[ \frac{\hbar \text{Im}(v_{\mathbf{k}}v_{\mathbf{k}}'^*)}{\omega_{\mathbf{k}}(t_0) c} \right]^{1/2}$ .

If there is a minimum, then

$$v_{\mathbf{k}}(t_0) = \left[ \frac{\hbar \text{Im}(v_{\mathbf{k}}v_{\mathbf{k}}'^*)}{\omega_{\mathbf{k}}(t_0) c} \right]^{1/2} e^{i\alpha_{\mathbf{k}}(t_0)} \quad v_{\mathbf{k}}'(t_0) = -c \frac{\omega_{\mathbf{k}}(t_0)}{i\hbar} v_{\mathbf{k}}(t_0) \quad (2.34)$$

under these functions,

$$E_{\mathbf{k}}(t_0) = 2 \frac{\text{Im}(v_{\mathbf{k}} v_{\mathbf{k}}'^*)}{\hbar c} \omega_{\mathbf{k}}(t_0) \quad F_{\mathbf{k}}(t_0) = 0 \quad (2.35)$$

meaning

$$\hat{\mathcal{H}}(t_0) = \text{Im}(v v'^*) \frac{1}{\hbar} \int \frac{d^3 \mathbf{k}}{(2\pi\hbar)^3} \left( 2\hat{a}_{\mathbf{k}}^\dagger \hat{a}_{\mathbf{k}} + \frac{(2\pi\hbar)^3 \hbar c}{2\text{Im}(v' v^*)} \delta^3(\mathbf{0}) \right) \omega_{\mathbf{k}}(t_0) \quad (2.36)$$

which is equivalent to the standard Hamiltonian for a scalar field without the presence of gravity.

### Bogolyubov Transformation

$$u_{\mathbf{k}}(t) = \alpha_{\mathbf{k}} v_{\mathbf{k}}(t) + \beta_{\mathbf{k}} v_{\mathbf{k}}^*(t) \quad (2.37)$$

$\alpha_{\mathbf{k}}, \beta_{\mathbf{k}} \in \mathbb{C}$  (time independent)

$$\text{Im}(u'_{\mathbf{k}} u_{\mathbf{k}}^*) = \text{Im}(v'_{\mathbf{k}} v_{\mathbf{k}}^*) \left( |\alpha_{\mathbf{k}}|^2 - |\beta_{\mathbf{k}}|^2 \right) \quad (2.38)$$

Changing the  $v$  functions would entail a change in the creation and annihilation, therefore if we could write the field as

$$\hat{\chi} = \int \frac{d^3 \mathbf{k}}{(2\pi\hbar)^3} \left[ \hat{b}_{\mathbf{k}} u_{\mathbf{k}} e^{i\mathbf{k}\mathbf{x}\hbar^{-1}} + \hat{b}_{\mathbf{k}}^\dagger u_{\mathbf{k}}^* e^{-i\mathbf{k}\mathbf{x}\hbar^{-1}} \right] \quad (2.39)$$

the field must be the same as if it was written with the  $v$  functions and  $\hat{a}$  operators, that means that

$$\hat{b}_{\mathbf{k}} u_{\mathbf{k}} + \hat{b}_{-\mathbf{k}}^\dagger u_{\mathbf{k}}^* = \hat{a}_{\mathbf{k}} v_{\mathbf{k}} + \hat{a}_{-\mathbf{k}}^\dagger v_{\mathbf{k}}^* \quad (2.40)$$

and thus, the relation between the operators would be

$$\hat{a}_{\mathbf{k}} = \alpha_{\mathbf{k}} \hat{b}_{\mathbf{k}} + \beta_{\mathbf{k}}^* \hat{b}_{-\mathbf{k}}^\dagger \quad \hat{a}_{\mathbf{k}}^\dagger = \beta_{\mathbf{k}} \hat{b}_{-\mathbf{k}} + \alpha_{\mathbf{k}}^* \hat{b}_{\mathbf{k}}^\dagger \quad (2.41)$$

now, there are ' $a$ ' particles in the ' $b$ ' vacuum

$$\langle {}_{(b)}0 | \hat{\mathcal{N}}_{\mathbf{k}}^{(a)} | {}_{(b)}0 \rangle = \langle {}_{(b)}0 | \hat{a}_{\mathbf{k}}^\dagger \hat{a}_{\mathbf{k}} | {}_{(b)}0 \rangle = |\beta_{\mathbf{k}}|^2 \frac{(2\pi\hbar)^3 \hbar c}{2\text{Im}(u' u^*)} \delta^3(\mathbf{0}) \quad (2.42)$$

therefore

$$\langle {}_{(t_0)}0 | \hat{\mathcal{H}}(t) | {}_{(t_0)}0 \rangle = \delta^3(\mathbf{0}) \int d^3 \mathbf{k} \left( \frac{|\beta_{\mathbf{k}}|^2}{|\alpha_{\mathbf{k}}|^2 - |\beta_{\mathbf{k}}|^2} + \frac{1}{2} \right) c \omega_{\mathbf{k}}(t) \quad (2.43)$$

meaning, if  $\beta_{\mathbf{k}} \neq 0$  for all  $\mathbf{k}$  then, at a time  $t > t_0$  the energy density will be different in relation to the original vacuum.

## 3 de Sitter Universe

The de Sitter Universe is a flat FLRW metric with no matter or radiation, but it does have a positive cosmological constant  $\Lambda$ . Per the Friedmann equations,

$$\left( \frac{\dot{a}}{a} \right)^2 = \frac{8\pi G + \Lambda c^2}{3} - \frac{\kappa c^2}{a^2} \quad (3.1)$$

the expansion parameter  $a(t)$  will be equal to

$$a(t) = a_1 e^{H_\Lambda t} + a_2 e^{-H_\Lambda t}, \quad H_\Lambda = \sqrt{\frac{\Lambda c^2}{3}} \quad (3.2)$$



# Notas sobre unidades

- $[s] = [\hbar]$
- $[a] = [\xi] = 1$
- $[\mu] = [L]^{-1}$
- $[R] = [L]^{-2}$
- $[\phi] = [\chi] = [\hbar]^{1/2}[L]^{-1}$
- $[\Pi] = [\hbar]^{1/2}[L]^{-2}$
- $[a_{\mathbf{k}}] = [\hbar]^{1/2}[L]^2$

# Bibliography

- [1] Mark Robert Baker, Natalia Kiriushcheva, and Sergei Kuzmin. Noether and hilbert (metric) energy-momentum tensors are not, in general, equivalent. *Nuclear Physics B*, 962:115240, 2021.
- [2] Viatcheslav Mukhanov and Sergei Winitzki. *Introduction to Quantum Effects in Gravity*. Cambridge University Press, 2007.