

Astrophysical S-factor calculation for selected reactions



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Abstract

Review on relevant elements of nuclear astrophysics.

Definition, motivation and applications of the astrophysical S-factor.

Survey on models that predict S-factors. They are mainly classified as empirical, potential, macroscopic and matrix models

Use of theoretical models to calculate the S-factors of selected reactions.

Comment on the applicability of the models for each reaction. Comparison between experimental data and theoretical predictions of the S-factors.

Acknowledgments

Mom for everything.

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Chapter 1

Nuclear astrophysics fundamentals

In order to apply nuclear astrophysics to the study of star element production, there are essentials to be considered. In this chapter, a general review of the main points relevant for nuclear physics are presented that emphasizes selected topics that are relevant in the field of nuclear astrophysics.

1.1 General aspects about the nucleus

Introduce a brief description of the nuclei. The structure of the nuclei The nuclear models

The nucleus is a bound state of a system of protons and neutrons. These substructures are bound states of a set of fundamental particles, which are called quarks.

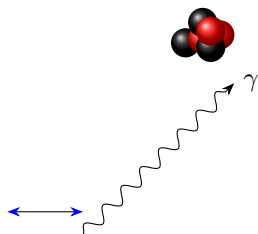
A general treatment given in [1] and basic concepts in [2].

A more particular on nuclear astrophysics [3].

Rest energy.

$$E_0 = Zm_p c^2 + (A - Z)m_n c^2 - B(A, Z), \quad (1.1)$$

where Z , A are the atomic and mass number respectively.



Each nucleus is compound of a finite number of proton and neutrons. In addition,

Elementary nuclear structure - Atomic nuclei are bound states of nucleons. They are particles. Characteristics - They obey the Pauli-exclusion principle like electrons - They have magnetic moment. - They are bonded in the nuclei by the strong force which compensates the repulsive electromagnetic force - Each nucleon is composed of elementary particles. Classification - proton: positive charge - neutron: no charge - Mass number. Atomic number.

Leptons Characteristics - They do not interact via strong force - Considerably lighter in mass than nucleons. Classification - Electron - Neutrino - Antiparticles are also considered.

1.2 Elements of quantum mechanics for nuclear physics

Quantum mechanics governs the physics at the nuclear scale. In particular, some elements of quantum mechanics theory that are essential for the description of nuclear reactions. For instance, quantum mechanical treatment of scattering theory is critical for determining the behavior of the rates of the reaction. In addition, tunneling phenomena, which can only be explained from the non-classical behavior of nuclear physics, permits the existence of nuclear fusion. Finally, transition phenomena will be considered in this section from the point of view of electromagnetic transitions.

Quantum mechanics text [4].

1.2.1 Scattering theory

Explain scattering theory relevant for the understanding of processes.

The sequence of the calculations on scattering is the following - Consider free path approximation -> Bessel Function solution - Expand free wave in terms of Bessel and Harmonic oscillator solutions - Obtain the coefficients of the expansion - Evaluate asymptotic behaviour at $r \rightarrow \infty$

When considering a potential, it appears that the oscillatory asymptotic wave function is shifted $e^{2i\delta_l}$. This factor depends on the angular momentum l . This directly relates with differential cross sections.

A reference text on collision theory [5].

$$\sigma = \frac{\text{scattered per time}}{(\text{incident per time per area})(\text{total nuclei})}. \quad (1.2)$$

Cross sections are probabilities The probabilities are ruled by the laws of scattering theory, which can be understood as a subpart of quantum mechanics.

There is of interest the differential cross section. Particularly, there is a connection with quantum mechanics scattering amplitudes with the formula

$$\frac{d\sigma}{d\Omega} = |f(\theta, \phi)|^2 \quad (1.3)$$

Cross section -> Reciprocal factors -> Scattering potentials -> Coloumb Barrier (With WKB approximation) -> Breit - Weigner formula (This is for resonant behavior)-> S factors, ...

There are several contributions to the total rate. There are four in principal Broad resonances, thin resonances, continuum, nonresonant. For particles of matter, the statistics at the star fusing scale is Maxwell - Boltzmann and for photons the Bose - Einstein statistics.

Scattering theory develops from the method of partial waves. It roughly consists on solving for the Schödinger equation in free space and, with the potential given, expand in terms of an orthonormal basis of functions, like spherical Bessel and Neumann functions and spherical harmonics. Then, conditions are adjusted via boundary condition satisfying.

$$\psi = \sum_l \sum_m c_{lm} Y_m^l j_l(kr) \quad (1.4)$$

Free space solution with $j_l(kr)$ the spherical Bessel function of l parameter. There are also the $n_l(kr)$ Neumann function and the Hankel functions $h^{(1)}(kr)$ and $h^{(2)}(kr)$ functions.

A result that relates cross sections and the probability current density is the theorem.

$$\sigma = \frac{4\pi}{k} \text{Im}(f(\theta = 0)) \quad (1.5)$$

Born approximation that is used for estimating the scattering amplitude.

$$f(\theta) = -\frac{2m}{\hbar^2} \int_0^\infty \mathcal{F}\{V(r)\} dk. \quad (1.6)$$

1.2.2 Tunneling phenomena

Quantum tunneling phenomena is critical in understanding the phenomena of alpha particles decaying in the nuclei barriers.

WKB approximation for the classical forbidden regions.

$$T = e^{2i\delta_l}. \quad (1.7)$$

Ansatz:

$$\Psi(r) = \Psi_0 \exp(\Phi(r)). \quad (1.8)$$

The Schrödinger equation transforms as

$$(\Phi'' + \Phi'^2) \Phi = \left(\frac{2m}{\hbar^2} (U(r) - E) \right) \Phi. \quad (1.9)$$

Then, if it is considered that $\Phi'' \ll \Phi'^2$, then Φ can be directly integrated and is expressed as:

$$\Phi(E) \approx -\frac{1}{\hbar} \int_{r_1}^{r_2} \sqrt{2m(U(r) - E)} dr, \quad (1.10)$$

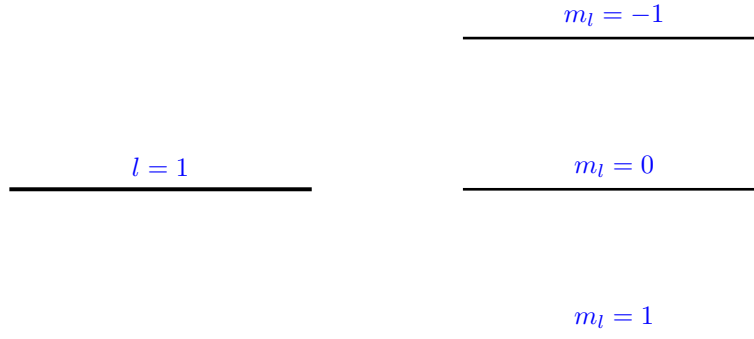
where r_1 and r_2 correspond to the lower and higher classical turning points respectively. In this opportunity, the negative choice for the square root was taken in order to account for the fact that the probability of transmission decreases within the non-classical region.

1.2.3 Perturbation theory

Generalities of perturbation theory. Cross sections. Fermi Golden rules.

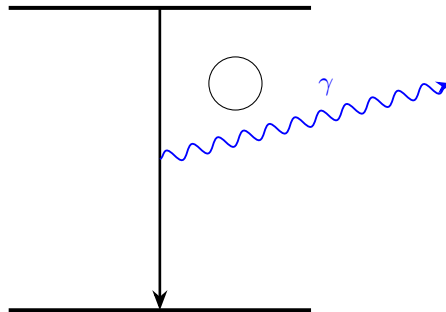
$$\sigma \propto |\langle k | \delta H | m \rangle|^2 \quad (1.11)$$

Splitting



1.2.4 Transitions of quantum states

The changes of state that nuclei exhibit obey certain rules. In particular, this selection rules consider angular momentum conservation.



Selection rules. Electromagnetic interaction and resonances [6].

1.3 Nuclear structure

Nuclear structure is presented in this section.

1.3.1 Nuclear models

Stability depends closely on the binding energy. In particular, since the rest energy of the nucleus decreases with $B(A, Z)$, the more is the binding energy, the more stable is nucleus.

Liquid drop model

A first approach when modeling the binding energy is a phenomenological model which is inspired in the physical description of a liquid drop. In particular, the predictions of the binding energies are determined by the fitting of the empirical formula:

$$B(Z, A) = a_v A - a_s A^{2/3} - a_3 \frac{Z(Z+1)}{A^{1/3}} + a_4 (Z-A)^2 + a_5 \delta. \quad (1.12)$$

However, this model does not fit accurately the binding energy of a selected nuclei which are more stable than expected from the model. In particular, it is said that the pair of number of protons and nucleons (Z, A) of these nuclei are magic numbers.

- The more stable the nuclei, the greater the binding energy. Binding energy is the energy difference of rest energy of the nucleus constituents and the actual rest mass of the overall nucleus.

$$\delta = \begin{cases} 1, & Z \text{ and } A \text{ even} \\ -1, & Z \text{ and } A \text{ odd} \\ 0, & \text{otherwise} \end{cases} \quad (1.13)$$

Nuclear shell model

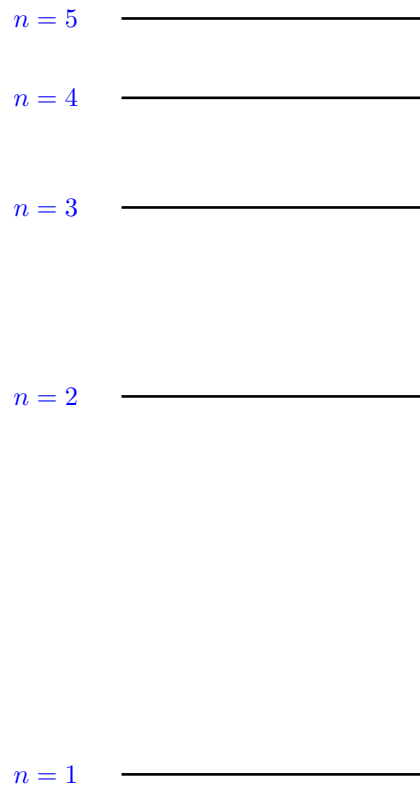
An alternative approach to the Liquid drop model is a model based on quantum mechanics. This model is closely related with the quantum mechanical modeling of the hydrogen atom. In particular, a great part of the nuclear shell model consists of solving Schrodinger equation which is expressed as:

$$-\frac{\hbar^2}{2m}\partial_{tt}\psi + V\psi = E\psi. \quad (1.14)$$

In this case, the potential V represents the effective potential of the core of the nucleus and is valid at certain radii. For example, the harmonic potential $V = m\omega^2 r^2$ models effectively the system of nucleons at most regimes.

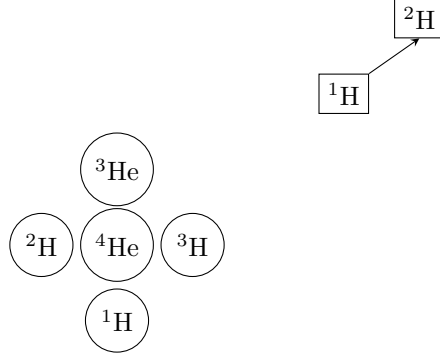
Spin orbit coupling rules the energy levels of the nucleons.

- Nuclear shell model is based on quantum mechanics.
- Nuclear shell model can predict the existence of the magic numbers.



1.4 Nuclear reactions

Nuclear reactions are reactions between nuclei. [7].



1.4.1 Classification of reactions

The relevant reactions for the document are:

$$0 + 1 \rightarrow 2 + 3 \quad (1.15)$$

$$\gamma + 3 \rightarrow 0 + 1 \quad (1.16)$$

Expand on how the fusion reactions occurs, their conditions and types as well as the differentiation factors with respect to fission processes.

Capture reactions

Also, radiative capture processes are considered at astrophysical energies. Various models could be used for determining potentials. There are clusters and potential models.

$$X + c \rightarrow \gamma + X'. \quad (1.17)$$

When c is a proton, the reaction is called radiative proton capture. [8]

Exchange reactions

$$X + x_1 \rightarrow x_2 + X' \quad (1.18)$$

A development on the (p, n) reaction kind is given by [9] and for the (n, p) in [10].

Fusion reactions

$$X + X \rightarrow a + X'. \quad (1.19)$$

1.4.2 Cross sections

Cross section of a process. Partial and total cross sections depending on the reaction type.

1.4.3 Reaction rates

The ratio that determines the quantity of a nucleus.

$$r = A \int_0^\infty E \sigma(E) e^{-\frac{E}{kT}} dE. \quad (1.20)$$

$$P(v)dv \propto v^2 \exp\left(\frac{mv^2}{2kT}\right)dv. \quad (1.21)$$

On specifying the rates, it is needed to be considered.

Reaction rates are considered from calculation of transmitted, incident and interacting - Reaction rate depends on the density of particles, the velocity and the cross section - It is averaged by a convenient distribution. (Maxwell-Boltzman for fermions, Bose Einstein (Planck) for photons) - With the rates, abundance evolution is determined

There are broad classification of reactions.

1.5 The astrophysical S-factor

1.5.1 Motivation and definition

There are S factors associated with the energy and they are the path to determine the nature of the potential, which is not necessarily a Yukawa potential.

$$S(E) = E \exp(2\pi\eta)\sigma(E). \quad (1.22)$$

It is directly proportional to the microscopic cross section of the process $\sigma(E)$ and it represents the rate of a certain nuclear reaction. In particular, the astrophysical S-factor scales the cross section with a factor proportional to the Coloumb repulsion.

The ν factor is defined as

$$\eta = \frac{Z_1 Z_2 e^2}{4\pi\epsilon_0 \hbar v} \quad (1.23)$$

where v corresponds to the relative speed between the reactants.

There is a strong relation of the astrophysical S factor with the center of mass energy.

Usually this relation depends on the square of the energy of the center of mass E_{cm} .

There are topics that are relevant in the frontiers of nuclear astrophysics [11].

1.5.2 General applications

Cross section extrapolation to low energies

Reaction rates and element abundance estimations.

Chapter 2

Reactions of astrophysical interest

Nuclear reactions are part of the core of nuclear astrophysics. They are accountable for the mechanism of production of elements in several astrophysical environments where thermonuclear reactions occur. In particular, there are specially relevant reactions for the explanation of the abundance of elements. For example, some of these reactions are part of critical fusion chains of stars, have challenging extrapolation techniques to low energies, or exhibit exotic underlying physical phenomena. In this section, a relevant selection of reactions types of will be reviewed.

Several reaction kinds and models are given in [12].

2.1 Big bang nucleosynthesis

Primordial reactions that were expected to happen before big bang occurred. For example:



2.2 Stellar fusion

The classical treatment of fusion in stars is given by the classical [13].

Introduce the needed conditions for the creation of stars in the primordial nebula. Describe how the hydrogen is burnt in the star's core and the associated products.



Consider the mechanisms for exchanging heat via radiation and convection.

Explain the required conditions that permits nuclear fusion in the star's nucleus.

2.2.1 Light heavy nuclei reactions

This reaction is governed by the weak force, due to conversion of a proton into a neutron, and by the strong and electromagnetic force.



When it comes to deuterium fusion, this channel is the most abundant.

The reaction rates of nuclear fusion processes depend on the interaction cross section and the abundance of the reactant nuclei.

$$d + d \rightarrow n + {}^3\text{He}. \quad (2.4)$$

$$d + d \rightarrow p + t. \quad (2.5)$$

The interactions of the previous processes are the electromagnetic and strong force.

$${}^3\text{He} + p \rightarrow {}^4\text{Li} + \gamma \rightarrow {}^4\text{He} + e^+ + \nu_e + \gamma. \quad (2.6)$$

This process is unlikely because of the instability of ${}^4\text{Li}$, which decay to ${}^3\text{He}$.

$${}^3\text{He} + {}^3\text{He} \rightarrow {}^4\text{He} + 2p. \quad (2.7)$$

This process is known as the pp1 chain.

Exemplify how the reaction mechanisms permits the stability of a star as well as produces He.

2.2.2 CNO cycle

The carbon burning produces the neutrons that are going to be used in heavy nuclei production. The neon burning starts by fusion Neon nuclei with ${}^4\text{He}$ nuclei to produce heavier elements. The oxygen burning starts at very high temperatures and pressures than the neon and carbon burning due to the high stability of oxygen nuclei. In fact, oxygen nucleus is double magic.

When the star has capacity for producing elements like carbon, oxygen and nitrogen, CNO cycle process can be produced. In particular, reactions take the form

$${}^{12}_6\text{C} \rightarrow {}^{13}_7\text{N} \rightarrow {}^{13}_6\text{C} \rightarrow {}^{14}_7\text{N} \rightarrow {}^{15}_8\text{O} \rightarrow {}^{15}_7\text{N} \rightarrow {}^{12}_6\text{C}. \quad (2.8)$$

There are other CNO cycles that involve oxygen and fluorine reactions. For instance, the second CNO cycle consists of the following reactions

$${}^{15}_7\text{N} \rightarrow {}^{16}_8\text{O} \rightarrow {}^{17}_9\text{F} \rightarrow {}^{17}_8\text{O} \rightarrow {}^{14}_7\text{N} \rightarrow {}^{15}_8\text{O} \rightarrow {}^{15}_7\text{N}. \quad (2.9)$$

There is even a CNO III cycle, which is far more uncommon. This branch is less dominant than the first branch.

2.2.3 Medium heavy nuclei reactions

These reactions concern nuclei with $A < 28$. Examples of this kind of reaction include the ${}^{12}\text{C} + {}^{12}\text{C}$, ${}^{12}\text{C} + {}^{16}\text{O}$ and the ${}^{16}\text{O} + {}^{16}\text{O}$ fusion reactions. Most of these reactions result in different channels.

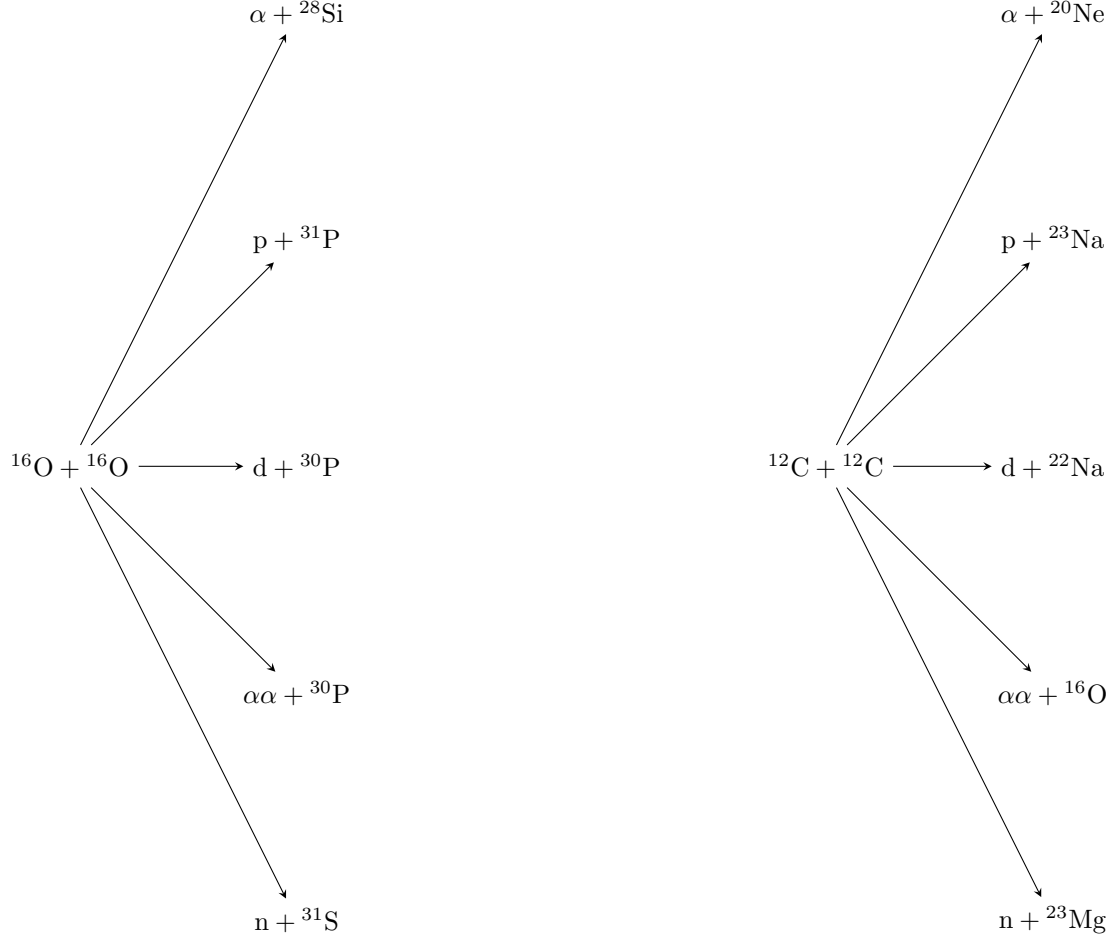
With respect to the carbon fusion reaction, the principal channel is given by:

$${}^{12}\text{C} + {}^{12}\text{C} \rightarrow {}^4\text{He} + {}^{20}\text{Ne}. \quad (2.10)$$

On the other hand, the main channel of the oxygen fusion reaction occurs in the following manner:

$${}^{16}\text{O} + {}^{16}\text{O} \rightarrow {}^4\text{He} + {}^{28}\text{Si}. \quad (2.11)$$

Additionally, there are inelastic channels associated with ${}^{16}\text{O}$ which could increase the probability of occurrence for the overall reactions.



There are more cycles of capture reaction. An example regarding the MgAl cycle is given in [14].

Explain the different chains going on during the carbon-oxygen combustion to heavier elements like iron or silicon.

The Trojan horse method[15]

nucleosynthesis in massive stars [16]

$^{12}\text{C} + ^{12}\text{C}$ Wave packet dynamics[17] nonresonant Bayesian [18] Modified sfactor [19] Trojan horse [20] [21] full-microscopical [22] multichannel folding [23] screening effects [24]

correlation between carbon fusion reactions [25] constraint $^{12}\text{C} + ^{12}\text{C}$ based on $^{12}\text{C} + ^{13}\text{C}$ reaction [26]

$^{16}\text{O} + ^{16}\text{O}$ cross sections [27] [28]

oxygen isotopes fusion [29] coupled channels [30]

$^{12}\text{C} + ^{16}\text{O}$ semi-microscopic cluster [31] low-energy resonances [32] $^{12}\text{C} + ^{16}\text{O}$ based on yields of $^{12}\text{C} + ^{16}\text{O}$ and $^{12}\text{C} + ^{18}\text{O}$ [33]

selected oxygen-oxygen and oxygen-carbon reactions [34] [35] $^{12}\text{C} + ^{12}\text{C}$ and $^{16}\text{O} + ^{16}\text{O}$ [36]

^{12}C and ^{24}Mg resonances [37] Fusion hindrance $^{12}\text{C} + ^{24}\text{Mg}$ [38]

heavy-ion reactions [39]

2.3 Supernovae and other explosive environments

Additional fusion reactions with nuclei with $A > 28$ are present at the last stages of the life of a star. There are two main environments: The stellar fusion and explosive events fusion.

2.3.1 Heavy nuclei reactions

There is a silicon burning process with the following description



Silicon burning is another process.



This reaction is essential since the ${}^{56}\text{Fe}$ nucleus has the maximum binding energy. Then, fusion reactions of heavier nuclei tend to be less spontaneous. Then, this chain of reactions continues until the process in there ${}^{56}\text{Ni}$ is reached.



2.3.2 The s and p processes

There exists the s and the r processes. Each permits to form heavier nuclear elements products. Now, there are aspects of the particularities of the production of the nuclei that are currently unknown

On the other hand, protons are captured at the ending of the life of a gigantic star. This process, which occurs at high densities and temperatures, is referred as proton capture. In particular, under astrophysical scenarios like supernovas, the rapid proton capture, in addition to the neutron capture, produces a considerable amount of the heavy nuclei. This process is referred as rapid proton capture, or rp process.

Further processes are not viable because they are not exothermic and potential heavier nuclei than ${}^{56}\text{Ni}$ are photodisintegrated. At this scale of energy, nuclear reactions start to behave similarly than chemical equilibrium reactions.

The conditions are associated with the mass. Particularly, depending on the mass regime, star evolution can differentiate. For instance, there are brown dwarfs, sun like stars.

There is an evolution process that is ruled by different instances. Each instance can be regarded as a region in the Hertzsprung-Russell diagram. This diagram consists of a two dimensional scatter plot of the intensity with respect to the spectral type of a star. Particularly, there are regions associated to the different instances of the life of a star: Main sequence (Hydrogen burning stars), red stars, neutron stars, supernova of kinds Ia Ib and other instances.

Electron degeneracy prevents the white dwarf to collapse and the neutron degeneracy pressure prevents the neutron star to collapse. These limits account for relativistic Fermi gas and neutron degeneracy with a model from general relativity.

Indicate that there exists different channels of production depending on the existence of neutron capture or fission products. Additionally, explain the needed conditions in order to those processes to occur.

If mass increases further, there is no place to any degeneracy compensation so the star would collapse to a black hole.

2.3.3 Active research topics

Enumerate the broad ranges of phenomena that are left to be explained.

Introduce current approaches of some of the unsolved questions on nuclear astrophysics, particularly on the last instances of fusion in the star's lifetime.

There are various types of supernovae. However, there are mechanisms that are not considered and yet not well explained.

Current experiments in detection of heavy nuclei are essential to determine the distribution of elements, specially in events like the fusion of neutron stars.

Review on active research topics and current challenges in nuclear astrophysics in general [40] and on low-energy nuclear physics [41] are given.

Chapter 3

Astrophysical S-factor models

There are several methods for estimating the astrophysical S-factor. Each method is more convenient depending on the nature of the reaction. For example, method accuracy can vary depending on the reaction type or the existence of resonant phenomena. Additionally, the methods are based on a specific approach when modeling the S-factors. In this section, the principal methods of estimating the astrophysical S-factors will be reviewed.

3.1 Empirical formulas

For resonant reactions, the Breit-Wigner formula

$$S(E) = S_r \frac{\Gamma_r^2/4}{(E - E_r)^2 + (\Gamma_r/2)^2}. \quad (3.1)$$

And some empirical formulas can be given.

On the other hand, non-resonant reactions S-factors can be calculated with polynomial interpolations and fits.

In particular, there is a polynomial expansion that is reported in literature to be useful.

$$S(E) = S(0) \exp(g_0 + g_1 E + g_2 E^2 + g_3 E^3 + \dots). \quad (3.2)$$

Also, more generic expansions are modeled by fitting a Laurent series expansion.

$$S(E) = \frac{a_{-1}}{E} + a_0 E + a_1 E^2 + \dots. \quad (3.3)$$

For some reactions, $a_{-1} = 0$, then they are described as a Taylor series expansion.

On the other hand, empirical formulas can mix resonant and non-resonant terms. For example, consider the expression:

$$S(E) = \sum_{k=0}^N a_k E^k + \sum_{l=1}^M \frac{\Gamma_l^2/4}{(E - E_l)^2 + \Gamma_l^2/4}, \quad (3.4)$$

where N is the maximum order of the polynomial to be fitted, which accounts for the non-resonant behavior of the reaction, and M represents the number resonances present in the reaction.

Applications in direct capture [42]. Exchange reactions (p, n) kind [43].

Charged particle collisions non-resonant [44] and resonant [45].

Fusion reactions, which is a more sophisticated model [46].

3.2 Potential models

The electromagnetic potential can be modeled as an uniform charged sphere potential.

$$V_{\text{EM}} = \begin{cases} \frac{Z_1 Z_2 e^2}{4\pi\epsilon_0 R} \left(\frac{3}{2} - \frac{r^2}{2R^2} \right), & r \leq R \\ \frac{Z_1 Z_2 e^2}{4\pi\epsilon_0 r}, & r > R \end{cases} \quad (3.5)$$

In addition, in order to account for the nuclear force, there are several potentials. For instance, in a first approximation, there is the step potential.

$$V_{\text{SqW}} = V_0 \theta(r). \quad (3.6)$$

Yakovlev et. al model [47] consists of an estimation of the potential that is similar to the parabolic potentials with free parameters that attempt to account for the effects of the nuclear and electromagnetic interactions.

$$V_{\text{YAK}} = \begin{cases} E_c \left(3 - \frac{r^2}{R_c^2} \right), & r \leq R_{c1} \\ \frac{\alpha}{r}, & r > R_{c1} \end{cases} \quad (3.7)$$

Additionally, the model distinguish between classical allowed and forbidden regions. For the first region, where $E < E_c$, the astrophysical S-factor is estimated as:

$$S(E) = S(0) \sqrt{\frac{Ec}{E}} [\exp(\Psi(E)) + k\xi], \quad (3.8)$$

where ξ is a free parameter of the model that accounts for the deviation from free particle behaviour of the S-factor. On the other hand, $S(0)$ corresponds to the zero-energy extrapolation of the S-factor.

uses the WKB approximation for estimating the wave function. Then, the astrophysical S-factor is determined by a phenomenological equation.

$$\Psi(E) = 2\pi\eta(E) + \Phi(E), \quad (3.9)$$

where $\eta(E)$ is the Sommerfeld parameter, and $\Phi(E)$ is the wavefunction calculated by the WKB. Further details on the calculation of the last parameter are presented in [Tunneling phenomena].

In order to the potential to be more realistic, a Saxon-Woods potential consider a smooth behavior.

$$V_{\text{SW}}(r) = \frac{V_0}{1 + \exp\left(\frac{r-R}{a}\right)}, \quad (3.10)$$

where V_0 represents the potential depth, a the parametrization and R a computed radius which is defined as:

$$R = r_0(A_1^{1/3} + A_2^{1/3}). \quad (3.11)$$

On the other hand, it is relevant, specially at low energies, the inclusion of angular momentum effects in the potential. In order to account for that effect, there is the potential with angular moment.

Also, spin orbit coupling effect is noticeable for certain kind of reactions. The contribution to the potential associated with this effect is expressed as:

$$V_{\text{SO}} = \langle L \cdot S \rangle \frac{d}{dr} V_{\text{WS}}, \quad (3.12)$$

where the coupling $\langle L \cdot S \rangle$ is expressed as:

$$\langle L \cdot S \rangle = J(J+1) - l(l+1) - s(s+1). \quad (3.13)$$

In the radiative capture reactions, cross sections are computed based on electromagnetic field operators. In particular

$$\sigma_m = \sum_k \langle k | \mathcal{O} | m \rangle C_{klm}. \quad (3.14)$$

Double exponential potentials:

$$V_{\text{DE}}(r) = -\frac{V_0}{e^{\alpha_1 r} + e^{\alpha_2 r}}. \quad (3.15)$$

Yukawa potential which includes screening effect.

$$V_Y = -\frac{c}{4\pi\epsilon_0} e^{-\mu r}. \quad (3.16)$$

And, potentials with non-locality are also found in literature. For example, there is the São Pablo potential that accounts for the effect of non-local interactions between the interactant nuclei.

Capture reaction review light nuclei[48] [49]

7Be radiative capture [50] [51]

9Be radiative capture [52]

p2H radiative capture [53]

2Hdp reaction [54]

13C radiative capture [55] [56]

11C radiative capture[57]

16O fusion [58]

Sao pablo [59]

$$V_{\text{SPP}} = \left(1 - \frac{V^2}{c^2}\right) \exp(V_F). \quad (3.17)$$

$$V_F(r) = \int \int \rho_1(\vec{r}_1) \rho_2(\vec{r}_2) f(\vec{r}_1, \vec{r}_2) d\vec{r}_1^3 d\vec{r}_2^3. \quad (3.18)$$

Optical model [60]

$$V(r) = V_R(r) + iW_R(r). \quad (3.19)$$

Woods-Saxon [61] [62]

$$V_{\text{WS}} = -V_0 f(r), \quad (3.20)$$

$$f(r) = \frac{1}{1 + \exp\left(\frac{r-R}{a}\right)}. \quad (3.21)$$

Energy dependent model [63]

$V_0(E)$

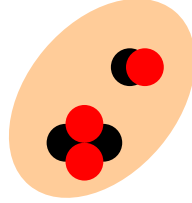
Cross section deduced potential [64] [65]

3.3 Microscopic models

Microscopic models are based on first principles and usually require assumptions about the nucleus and its structure.

$$H = \sum_k T_k + \sum_k V_k(r) + \sum_{k \neq j} V_{kj}(r). \quad (3.22)$$

On the other, clusters consider the system of interactant nuclei as a whole. In particular, there are models based on *ab initio* assumptions about a particular set of nuclei.



The *ab initio* no core shell model is presented in [66].

A more detailed survey on cluster models [67].

$$\int \int \rho(r_1) \rho(r_2) dr_1 dr_2. \quad (3.23)$$

Non-local calculations. Multi-body calculations.

Also, nuclei deformation can be considered in models. For example, there are models with no-core shell.

More sophisticated models use effective field theory for calculations. The general principle of these theories is the low-energy approximate solution of the physics described in quantum chromodynamics.

$$\mathcal{L} = N^\dagger i \hbar \gamma^\mu \partial_\mu N + \dots \quad (3.24)$$

Diagrams associated with the processes of low-energy nuclear force. Mediators are mesons like pions.

Modern theories [68]

Three body [69]

Ab initio [70] [71] [72] [73] tensor force [74]

No core shell model [75]

Shell model [76] [77]

Effective field theory [78] [79]

Multicluster [80] [81]

Hauser-Feshbach [82]

Break up effects [83]

Linear chain [84]

Hartree-Fock [85]

3.4 Matrix models

The matrix models permit the modeling of resonances with an appropriate fitting of free parameters. In particular, the main formalism in this approach is the R-matrix model. [86].

$$R_{cc'} = \sum_k \frac{\gamma_{ck}\gamma_{c'k}}{E - E_k}. \quad (3.25)$$

$$S(E) = S_r \frac{\Gamma^2/4}{(E - E_r + \Delta)^2 + (\Gamma/2)^2}. \quad (3.26)$$

There is parameter $P(E)$ which is expressed as:

$$P = \frac{\Gamma}{F_l^2 + G_l^2}, \quad (3.27)$$

where the functions $F_l = F_l(ka)$ and $G_l = G_l(ka)$ are the Coulomb functions evaluated at a particular $\rho = ka$. In addition, another parameter Q is given:

$$Q = \frac{\Gamma}{F_l G_l' + G_l F_l'}. \quad (3.28)$$

On the other hand, the phase shift is given by:

$$\delta = \delta_{\text{HS}} + \delta_R. \quad (3.29)$$

The model includes free parameter a which is called the radius channel.

Resonances arrives from poles. A description on the physical consideration of poles in the scattering (S-matrix) is given in [87].

In primordial nucleosynthesis [88] big-bang ^3He and several reactions [89].

Bayesian fitting methods [90].

Lithium detueron capture [91].

^{10}B proton non-radiative [92] and similarly [93]. ^{10}B radiative and resonances in ^{11}C inverse kinematics [94].

Alpha capture reactions of carbon [95]. Hybrid potential/R-matrix [96].

^{13}C alpha capture and ^{16}O determination [97].

Proton capture carbon [98] for C^{12} and [99] for C^{13} and [100] for the same reaction .

^{15}N capture in [101] and ^{14}N in [102]. also on ^{14}N proton capture [103].

There is also the K-matrix model as an alternative of the R-matrix model [104].

Alternative formulation of R-matrix [105].

Chapter 4

S-factor calculations for selected reactions

In this chapter, the S-factor calculation for a selected list of reactions will be presented. The list of reactions distinguish between resonant and non-resonant reactions. Additionally, this list includes reactions from Big Bang nucleosynthesis, pp-chain, CNO cycle and middle heavy nuclei fusion astrophysical environments.

In order to account for the best predictions for the astrophysical S-factors of the referred reactions, a motivation on the selection of a particular model will be mentioned, as well as the calculation consideration used to perform the calculations will be detailed for each reaction.

Further details about capture reactions with single-particle states is given by [106].

Additional, about pp-chain is given by [107].

Computer libraries that were used:

Python [108], Pandas [109], NumPy [110], SciPy [111], Matplotlib [112].

4.1 Non-resonant reactions

${}^2\text{H} + \text{p} \rightarrow \gamma + {}^3\text{He}$ and ${}^2\text{H} + \text{d} \rightarrow \text{p} + {}^3\text{H}$ Yakovlev and other fitting methods. can be also used with selected microscopical and potential models.

4.1.1 Calculation considerations

Yakovlev is used for middle-heavy fusion reactions. However, this model is not expected to work for radiative capture reactions since it predicts a decrement of the S-factor for all energies. In contrast, the S-factor increases in radiative capture reactions.

In the Yakovlev model, there is a fitting of the parameters in the following way. The free parameters are δ , E_c and ξ . In particular, E_c can be determined by the following expression:

$$E_c = \frac{\alpha}{R}, \quad (4.1)$$

where $\alpha = \sqrt{Z_1 Z_2 e^2}$ and R is a parameter that is determined as:

$$R = R_0 + \Delta R_1 |2Z_1 - A_1| + \Delta R_2 |2Z_2 - A_2|. \quad (4.2)$$

In particular, the parameter R_0 depends on the reactant nuclei while the parameters ΔR_1 and ΔR_2 can assume two values per pair of reactant nuclei. For instance, the values are given in the following way:

$$\Delta R_1 = \begin{cases} \Delta R_{10}, & 2Z_1 > A_1 \\ \Delta R_{11}, & 2Z_1 < A_1 \end{cases} \quad (4.3)$$

Analogously, ΔR_2 is determined as:

$$\Delta R_2 = \begin{cases} \Delta R_{20}, & 2Z_2 > A_2 \\ \Delta R_{21}, & 2Z_2 < A_2 \end{cases} \quad (4.4)$$

4.1.2 Results

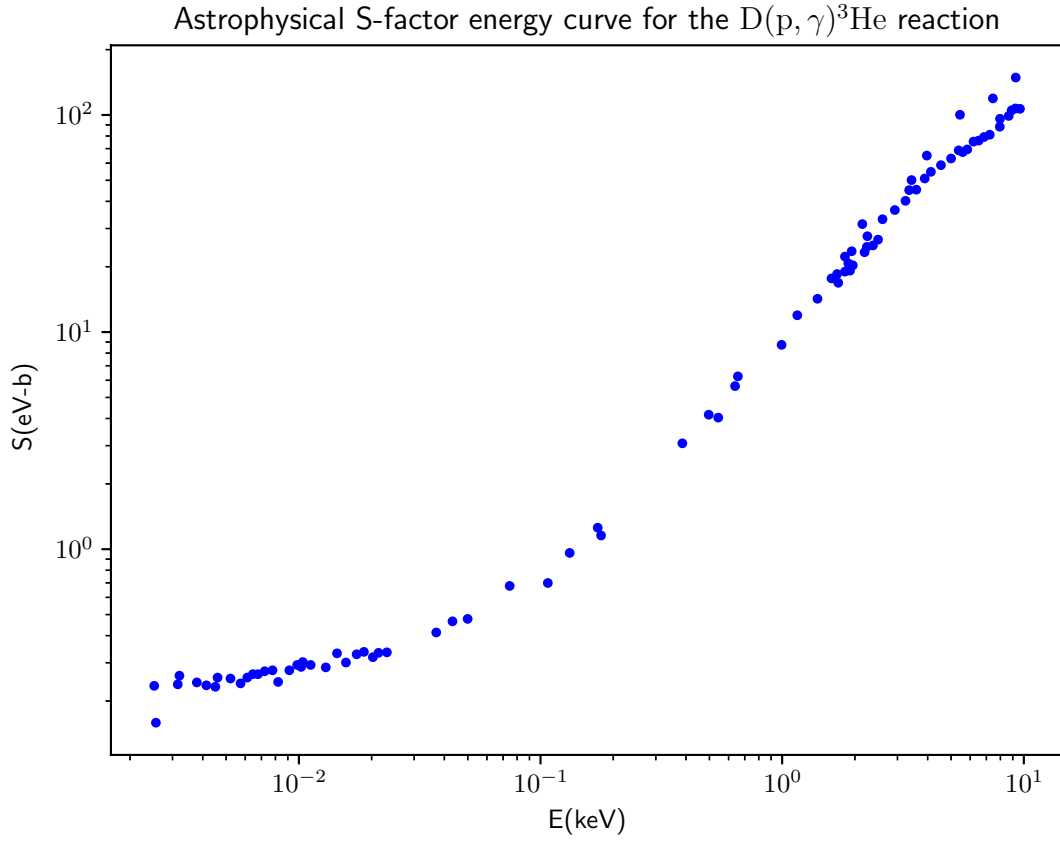


Figure 4.1: Caption

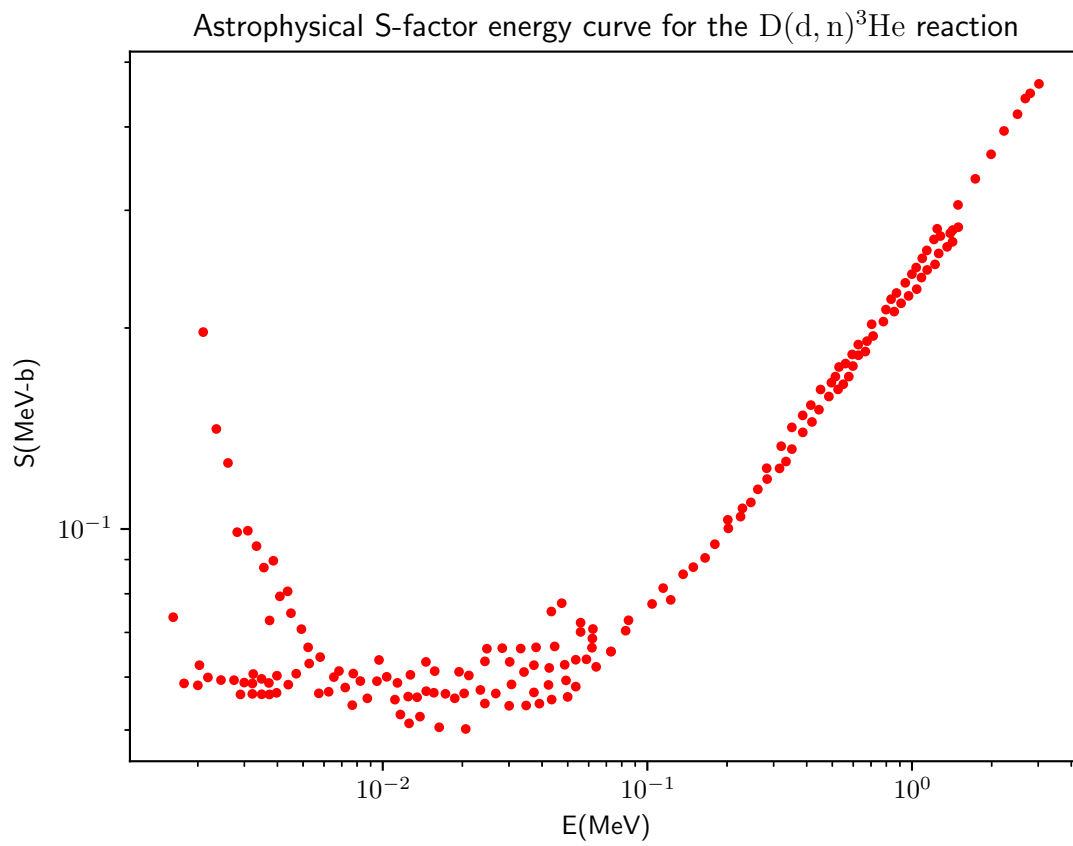


Figure 4.2: Caption

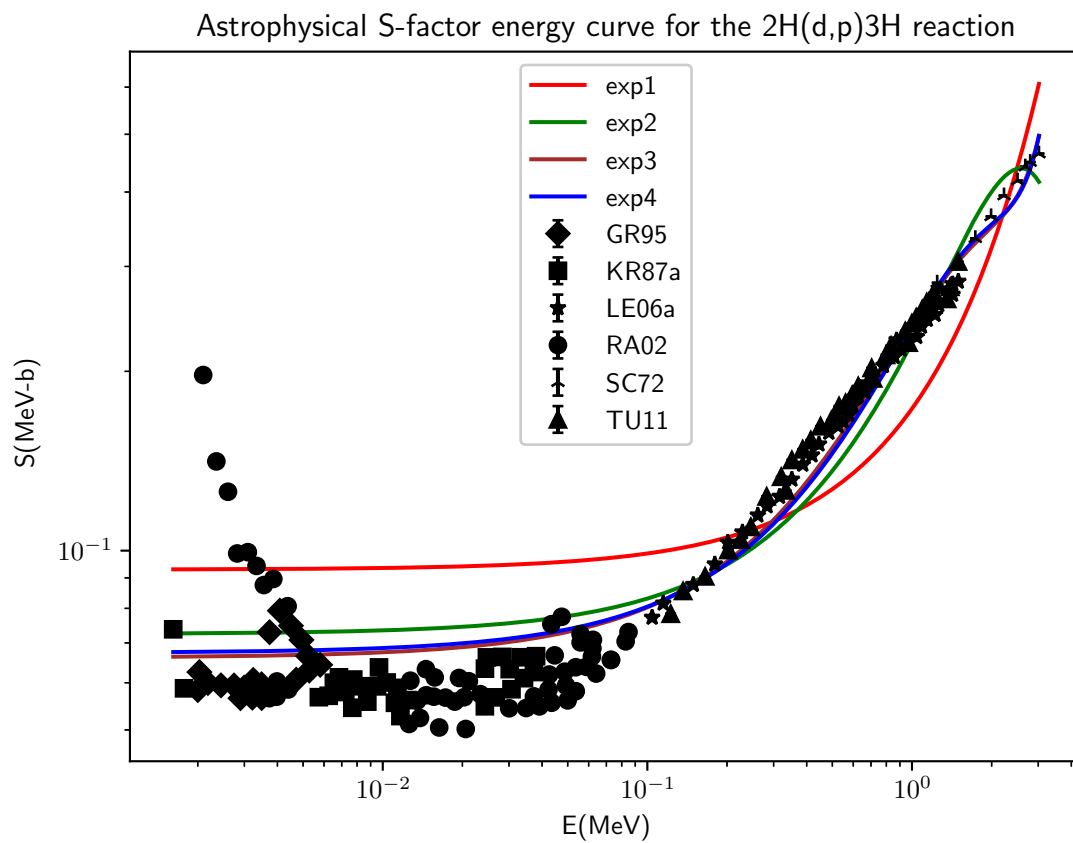


Figure 4.3: Caption

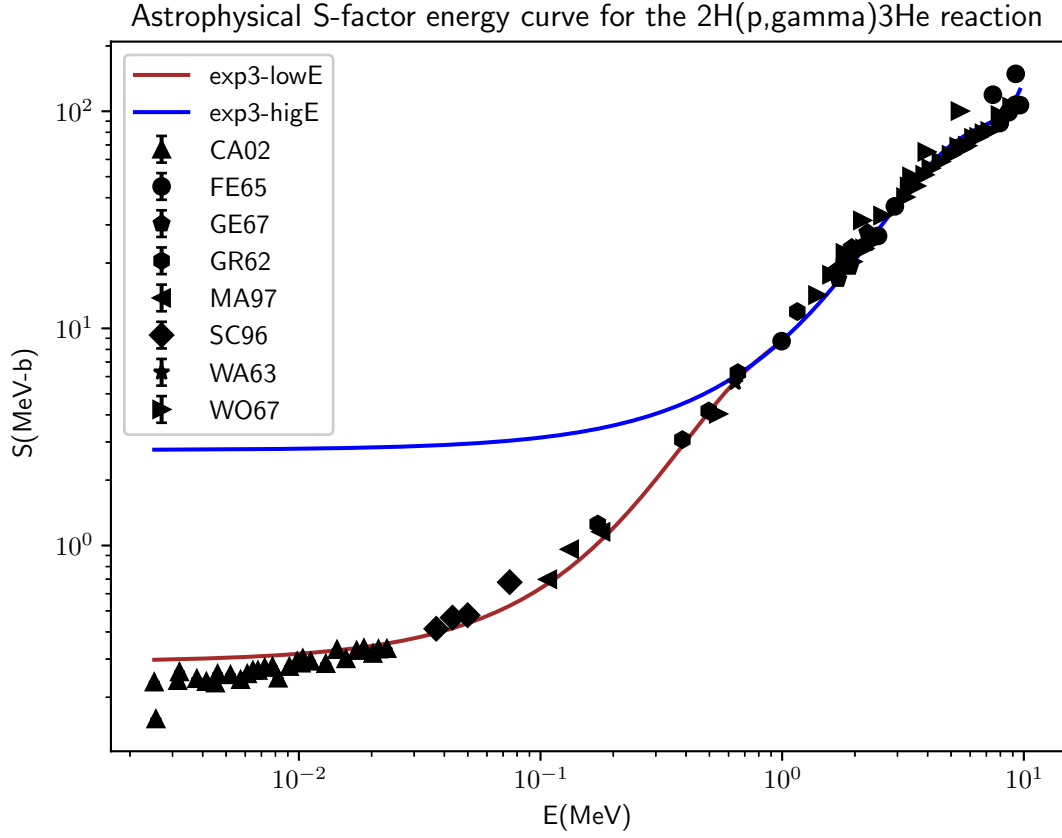


Figure 4.4: Caption

4.1.3 Analysis

4.2 Resonant reactions

Breit-Wigner. Single channel approximation.



A compilation of additional R-matrix treatment is given by [113].

4.2.1 Calculation considerations

4.2.2 Results

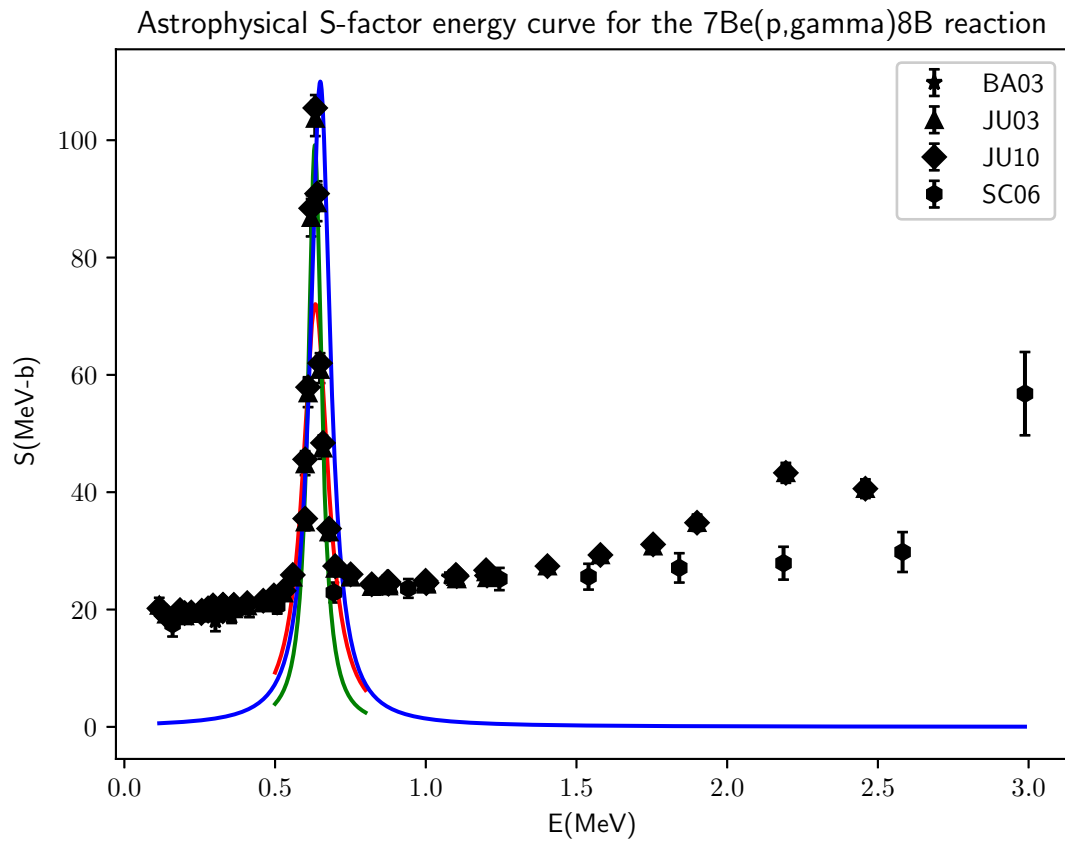


Figure 4.5: Caption

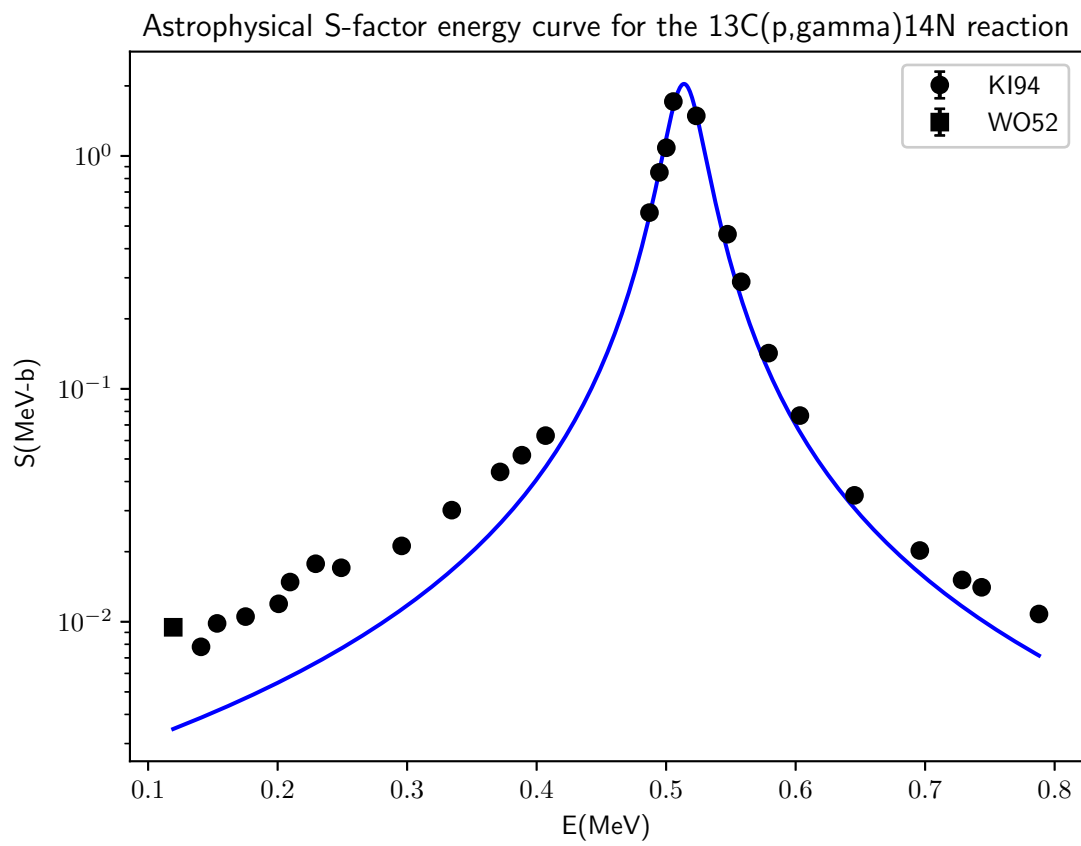


Figure 4.6: Caption

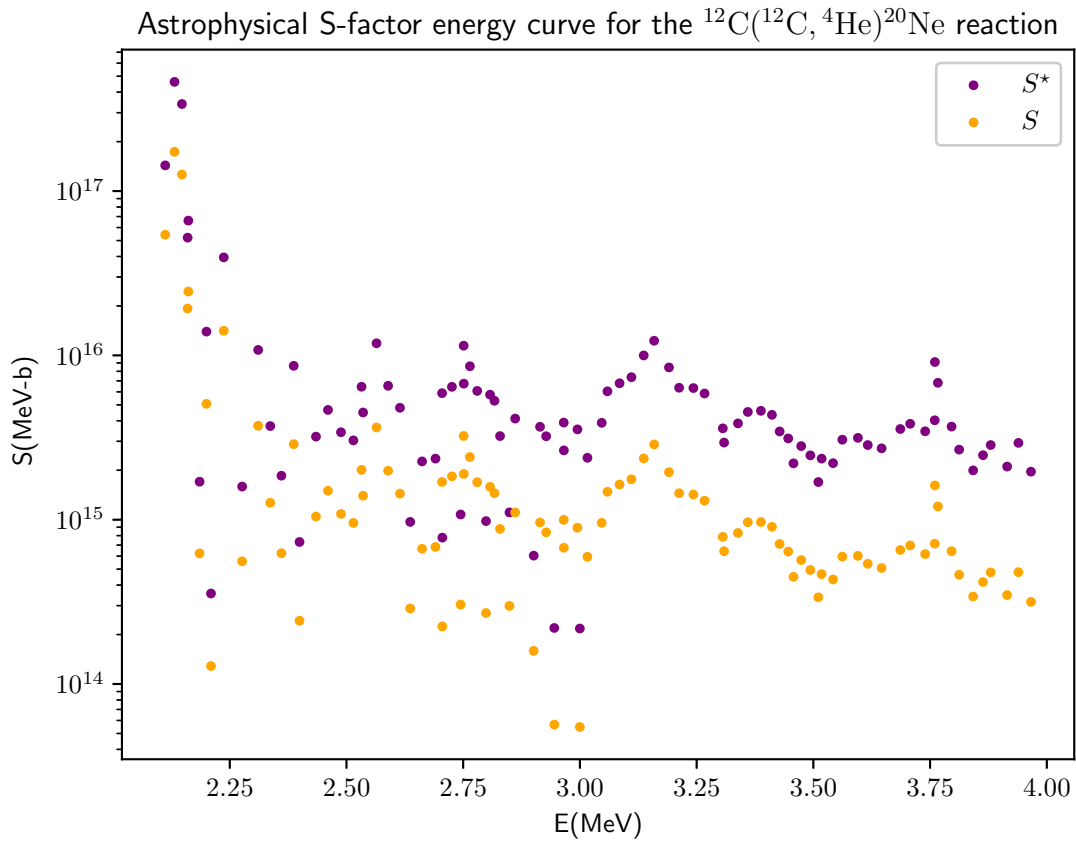


Figure 4.7: Caption

4.2.3 Analysis

Chapter 5

Conclusions

Astrophysical S-factors were calculated for the selected reactions and compared with available experimental data.

Nuclear physics of stars. Christian Iliadis.

Cauldrons in the cosmos. Claus Rolfs.

Hyde and Basdevant Nuclear physics.

Appendix A

Literature selected data

A.1 General data of nuclei

In this section general data about the structure of the nuclei will be presented.

EXFOR [114].

Japanese database [115].

A.1.1 Selected constants

NIST selected constants and conversion factors to convenient units.

The units used for calculations are those suitable for calculations of astrophysical S-factors. In particular, the unit of energy is given in mega-electronvolts (MeV), and the distances between nuclei are given in femtometers (fm), and the cross sections in barns ($1\text{b} = 10^{-28}\text{cm}$). Therefore, various fundamental constant are converted to consistent units with respect to the mentioned convention as shown in the next table.

Some constants to be considered: c , h , \hbar , m_e , m_p , m_n .

In the units to be used for nuclear physics calculations: $4\pi\epsilon_0 = 1$. So, the fine structure constant is expressed as:

$$\alpha = \frac{e^2}{\hbar c}. \quad (\text{A.1})$$

Therefore, the elementary charge is given by:

$$e = \sqrt{\alpha\hbar c} \approx \sqrt{1.41\dots}\sqrt{\text{MeV fm}}. \quad (\text{A.2})$$

Masses can be expressed in terms of MeV/c^2 or $\text{MeV fm}^{-2} \text{ s}^2$ depending on the context.

A.1.2 Nuclei structure data

atomic, massic, mass excess, spin, charge, data. L^π .

A.2 Astrophysical S-factor data

In this section, the reference to the experimental data on astrophysical S-factors for the selected reactions will be presented. In particular, the Nacre II database was widely used for reactions with $A < 10$. On the other hand, for reactions with $A > 10$, more specific references were used to obtain the experimental data.

Additional, light heavy experimental data pd reaction [116] and 2H proton capture c[117] and He3 photodisintegration [118] with [119]

A.2.1 Nacre II database

Selected reactions reference table. [120].

A.2.2 Middle heavy nuclei data

Selected reaction reference table for the $^{12}\text{C} + ^{12}\text{C}$, $^{12}\text{C} + ^{16}\text{O}$, $^{16}\text{O} + ^{16}\text{O}$, $^{13}\text{C} + ^{13}\text{C}$ reactions.

A.2.3 Transitions data

Transition type data, energy levels.

A.3 Resonances and transitions

Resonant specific data will be presented. In particular, for those reactions with the available information, the experimental peak as well as the reaction peak will be detailed. Additionally, with special observance on the radiative capture reactions, the transitions will be specified.

Data on the resonant $^7\text{Be} + \text{p}$ and experimental methods [121].

A.3.1 Resonance data

Selection reaction reference list.

A.3.2 Transitions data

Transition type data, energy levels, energy peak and widths.

A.4 Fitting parameters

The fitting parameters in a selected list of articles will be presented. In particular, this section distinguish between specific model parameters, like empirical, potential or microscopical models, and R-matrix fitting, whose special calculation considerations will be detailed.

A.4.1 Specific model parameters

Tables of models and parameters, with uncertainties.

A.4.2 R-matrix parameters

Tables of parameters that were chosen for the selected reactions

Appendix B

Special functions

In this section will be encountered special functions to be used in scattering theory and solution of analytical equations.

B.1 Bessel functions

Differential equation, solutions, 1st and 2nd kind and some useful properties. There are standard and spherical Bessel functions.

The spherical functions are solutions for the differential equation:

$$\frac{d^2x}{d\rho^2} + 2\rho\frac{dx}{d\rho} + (\rho^2 - l(l+1))x = 0. \quad (\text{B.1})$$

In particular, the solutions $J_l(\rho)$ and $Y_l(\rho)$ are called the Bessel and Von-Neumann solutions. One difference between these solutions is that $J_l(\rho)$ is well defined while $Y_l(\rho)$ has a pole at that $\rho = 0$.

In addition, the asymptotic behavior when $\rho \rightarrow \infty$ is different for each function as shown:

$$J_l(\rho) \rightarrow \sqrt{\frac{2}{\pi\rho}} \cos\left(\rho - \frac{\pi}{4} - \frac{l\pi}{2}\right). \quad (\text{B.2})$$

and

$$Y_l(\rho) \rightarrow \sqrt{\frac{2}{\pi\rho}} \sin\left(\rho - \frac{\pi}{4} - \frac{l\pi}{2}\right). \quad (\text{B.3})$$

B.2 Coulomb functions

Differential equations, solutions and more properties.

$$\frac{d^2x}{d\rho^2} + (\rho^2 - 2\eta\rho - l(l+1))x = 0. \quad (\text{B.4})$$

There are two solutions for this equations. In the context of scattering theory, the solutions $F_{l\eta}(\rho)$ and $G_{l\eta}(\rho)$ are analogous to the $J_l(\rho)$ and $Y_l(\rho)$ solutions of the Bessel differential equation respectively.

The formulas are connected as [122].

$$G_{l\eta} = \frac{F_{l\eta} - iF_{-l\eta}}{2i}. \quad (\text{B.5})$$

In a similar way than the Bessel functions, the Coulomb functions have harmonic like asymptotic behavior. In particular, $F_{l\eta} \rightarrow \sin \theta_{l\eta}$ and $G_{l\eta} \rightarrow \cos \theta_{l\eta}$ with $\theta_{l\eta}$ defined as :

$$\theta_{l\eta}(\rho) = \rho - l\frac{\pi}{2} - \eta \ln(2\rho) + \arg \Gamma(l+1+i\eta). \quad (\text{B.6})$$

B.3 Additional selected functions

Hypergeometric equations and connection with coulomb functions with implementation.

Spherical harmonics and Legendre polynomials for the expansions used in scattering theory. Hyperspherical harmonics

B.4 Clebsch-Gordan coefficients

Motivation, definition and some computations. Generalization to further spins and couplings that are useful for determining effects like the spin-orbit coupling.

The total angular momentum and change of basis.

$$J = L + S \quad (\text{B.7})$$

Tensor products.

$$a \otimes b. \quad (\text{B.8})$$

$$a \oplus b. \quad (\text{B.9})$$

And the Wigner notation for the Clebsh-Gordan coefficients.

$$\begin{pmatrix} s & l & I_x \\ j_s & j & J \end{pmatrix}. \quad (\text{B.10})$$

B.5 Microscopic model functions

Those relevant microscopic model methods. Specially, those that account for cluster model and effective field theory mod els.

Hamiltonians and Lagrangians of related field theories.

Appendix C

Fitting

This section is about fitting considerations on the different approaches used to calculate the astrophysical S-factor.

C.1 Empirical formulas fitting

Parameter list with uncertainty for its respective formula.

C.2 Free parameters fitting on potential models

Parameter list, uncertainty and its respective formula.

[123] Sao Pablo potential.

[124] Woods-Saxon potential.

C.3 Microscopic model fitting

Parameter list, uncertainty and its respective formula.

[125]

C.4 R-matrix fitting

Parameter list, uncertainty and its respective formula. Widths and channels.

Reaction	$E_R(\text{MeV})$	$\Gamma(\text{MeV})$	References
1	0.504	0.102	REF1
4	0.607	0.304	REF2

[126] [127] [128]

Appendix D

Numerical integration and differential equation solving

In this section will be introduced the generalities of the numerical solution of integrals and differential equation that were used throughout the document. In particular, a special subsection on the numerical solution for the Schrödinger equation for scattering phenomena will be presented.

D.1 Integration of selected potentials

WKB numerical implementation based on the Gaussian quadrature algorithm.

D.2 Numerical solution of the Schrödinger equation

The numerical solution of the Schrödinger equation is found to be useful for the approaches where it is not possible to find a solution for the problem in closed form, mainly analytical solutions.

[129]

Coupled channels [130].

D.2.1 Main approach

The different methods to be solved the Schrödinger equation and overall differential equation solving strategy.

D.2.2 Implementation of the potential functions

Different potential considerations.

D.2.3 Boundary conditions implementation

Expected boundary conditions at $r \rightarrow \infty$ and their numerical implementation.

Appendix E

Computer codes implementation

In this section is going to be presented the aspects related with the `sfactors` python module.

E.1 Structure of the computer program

The full tree with the files with explanation of the detailed packages.

E.2 User manual

In this section the user manual should be presented.

E.2.1 Installation

In this section a full installation sequence will be presented

E.2.2 Plotting

Plotting of S-factors.

E.2.3 Empirical fitting

Empirical fitting based on predetermined and custom functions.

E.2.4 Specific model fitting

Model fitting.

E.2.5 Model testing

Different tests of the program code.

E.3 Documentation

- E.3.1 Databases**
- E.3.2 Non resonant**
- E.3.3 Resonant**
- E.3.4 Plots**
- E.3.5 Reconstructed Plots**
- E.3.6 Solver**

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