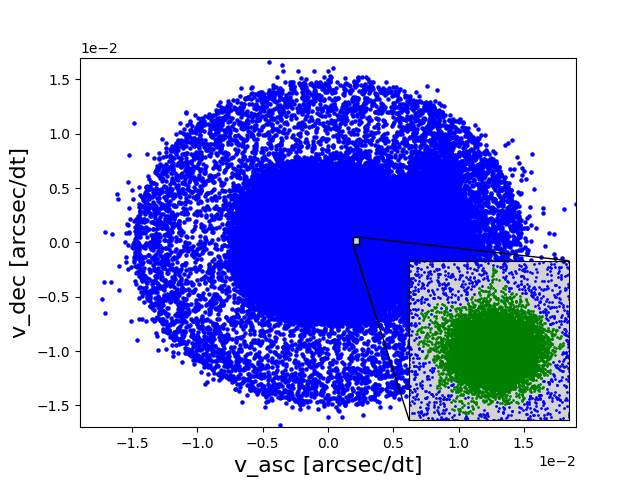
**Testing the limits of the ELT**



## Abstract[](#abstract)

Todo?: write

## Acknowledgement[](#acknowledgement)

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## Introduction[](#introduction)

The Extremely Large Telescope (ELT) is currently under construction. This first next generation 40m class telescope will have the resolution and sensitivity needed to study the motions of individual stars in the galactic centre (GC). These position and motions can be used to detect star clusters using some type of clustering algorithm and consequently estimate their Initial Mass Function (IMF). Whether the IMF of star clusters is universal is the subject of discussion [[2](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id45)] at this time. The study of Young Massive Clusters (YMCs) [[29](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id56)] near the GC will hopefully give more insight into this hotly debated topic. YMCs are tightly grouped clusters of stars, generally containing >10000 members. All original members are still present, the IMF is well sampled in all mass regimes. These facts combine to give a good picture of the end product of a star formation event.

Given detected star clusters the question remains, how reliable these results are. If the true classification of the studied set of stars is known, the performance of the clustering algorithm and hence the reliability of the results can be calculated. However, this is generally not the case.

In this master thesis an N-body simulation containing cluster stars and field stars under the influence of the milky way potential is performed. Snapshots are taken at different timesteps and fed to ScopeSim [[22](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id50)] to create mock observations. With the help of Photutils [[8](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id54)] positions are extracted and the DBSCAN algorithm used to detect cluster and field stars. Finally, the reliability of the results can be determined by comparing its results with the initially simulated stars.

## Galactic Potential[](#galactic-potential)

Modelling the gravitational potential of the Milky Way has been an ever evolving topic in recent decades. [[39](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id58)] provides a brief history. Here the potential is used to approximate the force on the simulated stars and to initialize field stars.

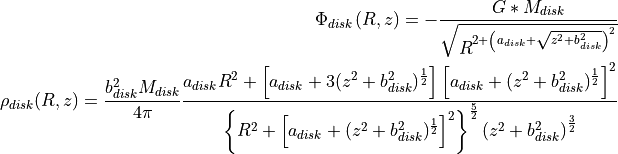
The model of choice consists of four parts: black hole, bulge, disc and the dark matter halo.

The **black hole** is represented by a Keplerian potential:

\Phi_{bh}\left ( r \right ) = -\frac{G*M_{bh}}{r}

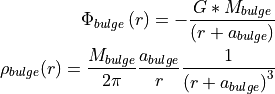
with r being the spherical radius.

The **disk** can be modelled via a Miyamoto Nagai potential [[49](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id3)]



with R being the cylindrical radius and with z being the distance to the reference plane.

For the **bulge** the Hernquist potential [[44](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id6)] is used



a being the scale-length of the spheroid potential

and the NFW potential [[26](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id59)] is used for the **dark matter halo** potential

\Phi_{halo}\left ( r \right ) = \frac{-4\pi G\rho _{s}r_{s}^{3}\ln\left ( 1+\frac{r}{r_{s}}\right )}{r}

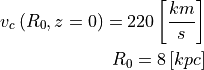
where \rho _{s} is the characteristic density and r_{s} the scale length.

Parameters for bulge and disk taken from [[30](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id4)] and the radius r_{s} for the halo from [[7](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id5)]:

Table 1 MW model parameters[](#id8)

| Parameter | Value | Unit |
| --- | --- | --- |
| M_{bh} | 4*10^{6} | M_\odot |
| M_{disk} | 10^{11} | M_\odot |
| a_{disk} | 6.5 | kpc |
| b_{disk} | 0.26 | kpc |
| M_{bulge} | 3.4*10^{10} | M_\odot |
| a_{bulge} | 0.70 | kpc |
| r_{s} | 16 | kpc |

\rho _{s} can be determined by imposing



### Circular Velocity[](#circular-velocity)

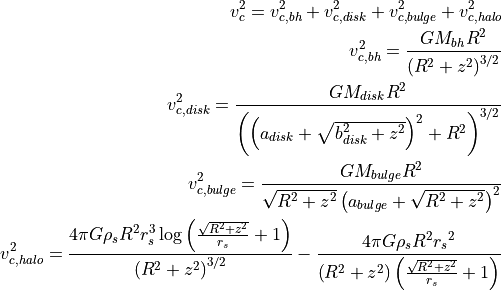
The circular velocity v_{c} is defined via

(1)[](#equation-circular-velocity)v_{c}\left ( R,z \right ) = \sqrt{R\frac{\partial \Phi \left (R,z \right )}{\partial R} }

with the total potential being in the presented case

(2)[](#equation-total-potential)\Phi \left ( R,z \right ) = \Phi_{bh} \left ( R,z \right )+\Phi_{disk} \left ( R,z \right )+\Phi_{bulge} \left ( R,z \right )+\Phi_{halo} \left ( R,z \right )

therefore



Inserting the parameters given in [Table 1](#tbl-parameters) and setting G\approx 4.302*10^{-6}\left [ \frac{kpc}{M_\odot}\frac{km^{2}}{s^{2}} \right ] results in \rho_{s}\approx 4.5*10^{6} \left [ \frac{M_\odot}{kpc^{3}} \right ]

the circular velocity can be used for fitting the model parameters [[1](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id60)]. [Fig. 2](#fig-potential-circular-velocity) displays the velocity for the chosen model and its components.

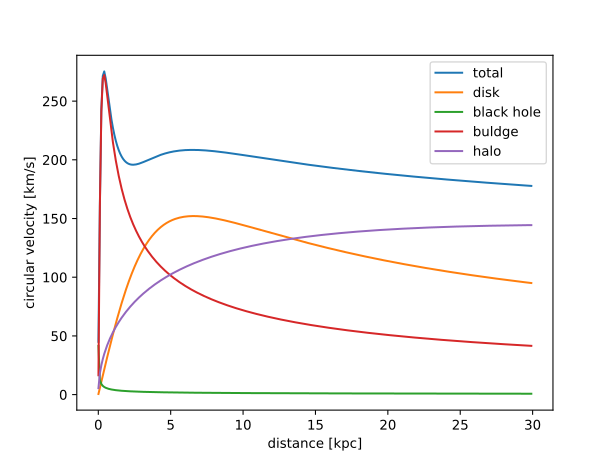


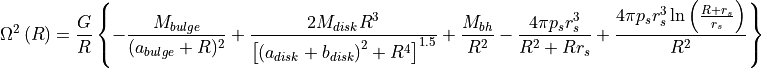
Fig. 2 Galactic rotation curve for the chosen MW model[](#id9)

### Angular Velocity[](#angular-velocity)

The angular velocity is required for initializing disc star velocities and can be obtained from the potential:

\Omega ^{2}\left ( R \right ) = \frac{1}{R}\frac{\partial \Phi \left ( R,0 \right )}{\partial r}

For the total potential [(2)](#equation-total-potential) this derivation yields:



### Mass Distribution[](#mass-distribution)

The mass inside a volume is calculated by numerical integration of the density. GSL implementation of Monte Carlo Integration is used. For further details refer to the GSL [documentation](https://www.gnu.org/software/gsl/doc/html/montecarlo.html).

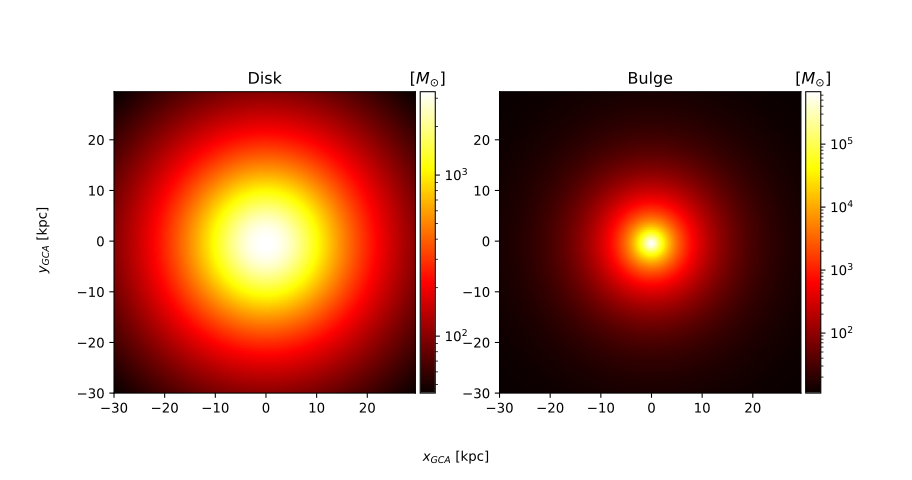


Fig. 3 Mass distribution at z_{GCA} = 1.0 [kpc][](#id10)

### Surface Mass Density[](#surface-mass-density)

The surface mass density (SMD) is used to estimate the vertical velocity dispersion [(29)](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#equation-vertical-velocity-dispersion) and is defined by

(3)[](#equation-surface-mass-density)\Sigma \left ( R \right )=2\int_{0}^{\infty}\rho \left ( R,z \right )\mathrm{d}z

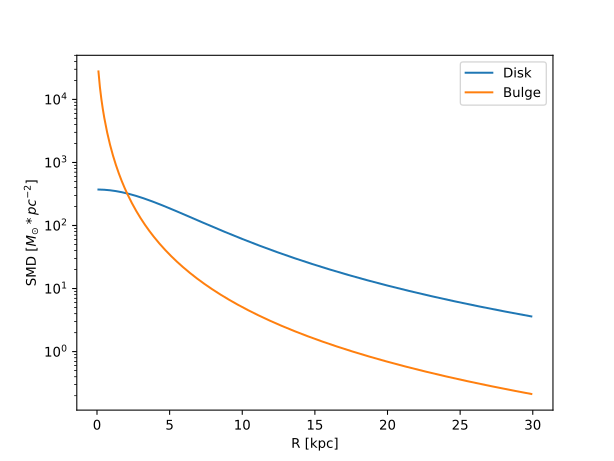


Fig. 4 The surface mass density of the disc and bulge[](#id11)

## Initialization[](#initialization)

### Sampling distribution functions[](#sampling-distribution-functions)

Space, velocity and mass distributions are sampled during various steps of the initialization using the following two methods. For further details and alternative methods see [[11](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id43)].

#### Rejection sampling[](#rejection-sampling)

A proposed distribution function g(x) is sampled and each resulting sample accepted with a probability of \frac{f(x)}{Cg(x)}. C is a constant ensuring f(x)\leqslant Cg(x) \forall x and f(x) the target distribution.

If g(x) is the pdf of a uniform distribution with x \in \left [ a,b \right ]

g(x)=\frac{1}{b-a}

and m the maximum of f(x) within \left [ a,b \right ] then C=\frac{m}{b-a} and the acceptance criterion becomes

u \leq f(x)

here *u* is a uniform distribution U(0,Cg(x)) = U(0,m). While Cg(x) has to be larger than f(x) the area between them should be as small as possible to ensure a decent acceptance fraction.

#### Inverse transform sampling[](#inverse-transform-sampling)

If the cumulative probability function (cpf) of the probability distribution f(y) is continuous and its inverse obtainable, this method can be utilized and is more efficient than rejection sampling. The cpf describes the probability that the random variable is smaller or equal to x

F(x)=P(X \leq x) = \int_{-\infty }^{x}f(y) dy

The inverse F^{-1}(y) may be calculated via F(F^{-1}(y)) = y. With y drawn from a uniform distribution the results of F^{-1}(y) have the desired probability distribution f(y).

### Cluster[](#cluster)

A cluster can be characterised via a number of parameters, namely its density profile, metallicity, binary fraction and initial mass function (IMF). The following sections cover methods implemented by the author. Alternatively, cluster stars can be generated using the open source code McLuster [[47](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id35)] which supports more density profiles and all other parameters mentioned above. The generated file can then be imported by setting the appropriate parameters.

### Initial Mass Function (IMF)[](#initial-mass-function-imf)

The IMF describes the initial mass distribution of stars, that is, how many stars have which mass in some set of stars at the start. It is subject of debate whether the IMF of star clusters is universal [[2](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id45)] and the present day mass function (PDMF) usually differs from the IMF. Figure 13 in [[42](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id44)] depicts the Salpeter IMF [[55](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id49)] and the PDMF of Arches. Hence, the IMF is deemed an adequate estimate for the PDMF of a generic YMSC. Sampling routines for the following two IMFs have been implemented by the author.

#### Salpeter IMF[](#salpeter-imf)

(4)[](#equation-salpeter)p(m)=\frac{dN}{dm}=A*m^{-\alpha }

This comparatively simple and widely used ([[32](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id46)], [[54](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id47)], [[43](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id48)]) IMF can be sampled using inverse transform sampling. To do so one must integrate the IMF over the mass and normalize it, yielding the normalized cumulative distribution function (NCDF).

(5)[](#equation-salpeter-ncdf-1)P(m)=\int_{m_{min}}^{m}A*m^{-\alpha } dm = \frac{A}{-\alpha +1}\left ( m^{-\alpha +1} -m_{min}^{-\alpha +1}\right )

A is defined by normalization:

P({m_{max}})\equiv 1\rightarrow A=\frac{-\alpha +1}{m_{max}^{-\alpha +1} -m_{min}^{-\alpha +1} }

With this [(5)](#equation-salpeter-ncdf-1) becomes:

P(m)=\frac{m^{-\alpha +1} -m_{min}^{-\alpha +1}}{m_{max}^{-\alpha +1} -m_{min}^{-\alpha +1}}

Since Salpeter is a power-law distribution function, the inverse of the NCPF can be calculated analytically. After some simplifications the final result reads

m_{rand} = m_{min}*\left [ 1+x*\left ( \left ( \frac{m_{max}}{m_{min}} -1\right )^{-\alpha +1} \right ) \right ]^{\frac{1}{-\alpha +1}}

where *x* is a random number in range [0,1]

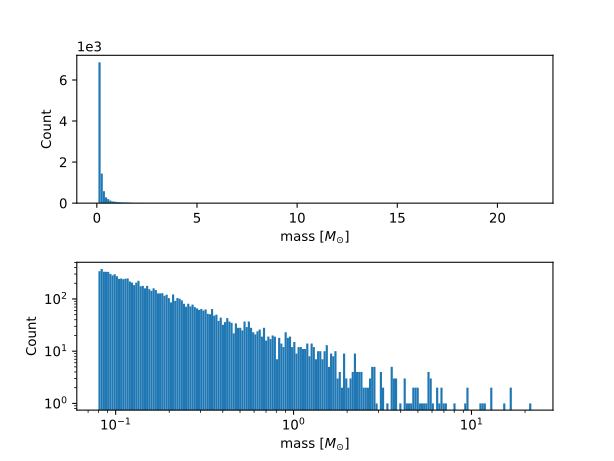
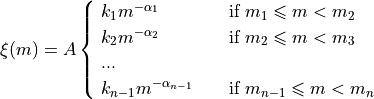


Fig. 5 10^3 stars sampled from the Salpeter IMF[](#id41)

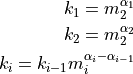
#### Broken Power Law (BPL)[](#broken-power-law-bpl)

The BPL IMF has the benefit of being highly adaptable. It can be utilized to approximate any other IMF. It has been used to describe both globular cluster IMFs [[3](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id11)] as well as young star cluster IMFs ([[28](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id12)], [[19](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id13)]).

The following is a generalization of the equation given by [[18](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id10)] for n-1 intervals.



where A is a normalization constant and k_{i} is defined as



As in the case of Salpeter, random samples are drawn with inverse transform sampling.

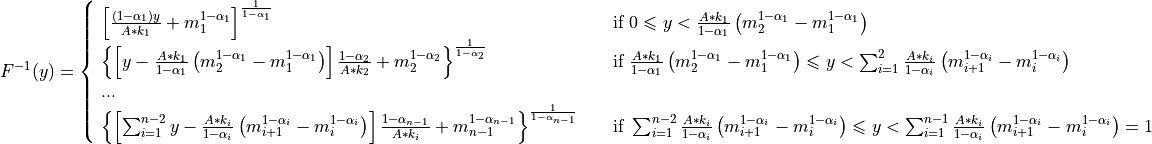
The NCDF can be obtained by integrating \xi(m) over the mass interval.

F(m)=\int_{m_{1}}^{m}\xi(m)\,dm

Where A is defined by the normalization constraint:

A * \sum_{i=1}^{n-1} \left ( k_{i}\int_{m_{i}}^{m_{i+1}}m^{-\alpha_{i}}\,dm \right ) = 1

Inverting the NCDF leads to



where *y* is a random number in range [0,1]

### Density Profile[](#density-profile)

Different density profiles (King [[46](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id40)], Plummer [[53](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id41)] and Elson, Fall & Freeman [[41](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id42)]) have been used ([[17](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id36)], [[35](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id37)], [[16](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id38)]) to simulate YMCs. The Plummer profile has been implemented here as described in [[38](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id39)].

### Field Stars (FS)[](#field-stars-fs)

Any star which does not belong to the simulated cluster is a FS and sampled from the disc and bulge potential described in [Galactic Potential](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#galactic-potential-label).

The cone of vision (COV) is defined by the angle of view \alpha, the view distance h (height of the cone), the view point vP (location of the observer) and the focus F (a point along the line of sight).

The COV is constructed by transforming a right circular cone, where the vertex is at the origin and the circular base normal to the z axis.

Per transformation the tip of the cone is displaced from the origin to the view point vP and its axis is rotated to align with the line of sight l. Consequently, the transformation consists of both translation and rotation illustrated in the following figure.

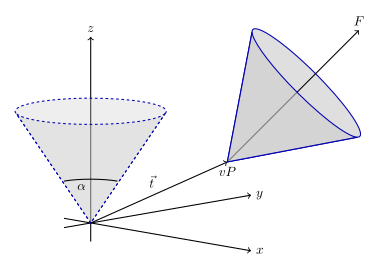


Fig. 6 Transformation of the cone of vision[](#id42)

Rotation and translation are both isometric transformations meaning area and volume of the cone are preserved [[15](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id9)] (p.175).

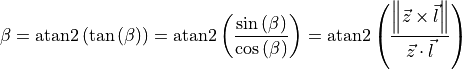
A unit quaternion \textbf{q} is used in order to construct the rotation matrix. With rotation axis \vec{b} and angle \beta the quaternion is given by

\textbf{q} = \left ( \textup{cos}\left (\frac{\beta}{2}\right ), \vec{b}\textup{ sin}\left ( \frac{\beta}{2} \right )\right )

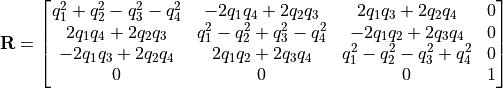
The rotation axis \vec{b} is the normalized cross product of the original (\vec{z}) and target (l) cone axis

\vec{b}=\frac{\vec{z}\times\vec{l}}{\left \| \vec{z}\times\vec{l} \right \|}

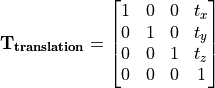
The angle \beta between the vectors of interest can be calculated as follows



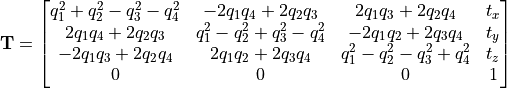
Next, quaternion is converted to the rotation matrix [[21](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id7)]. Using the homogeneous notation [[34](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id8)] (p. 57) the matrix becomes:



The translation matrix for the translation vector \vec{t} reads [[34](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id8)] (p. 66):



The transformation matrix \mathbf{T} is the product of \mathbf{R} and \mathbf{T_{translation}}

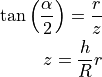
(6)[](#equation-transformation-matrix)

The total mass M of disc and bulge stars is obtained by integrating the respective density over the COV.

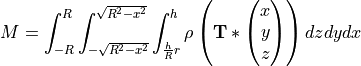
For a right circular cone

\tan\left (\frac{\alpha }{2} \right ) = \frac{R}{h}

and on its surface



where R is the base radius of the cone and r = \sqrt{x^2+y^2}. Hence \frac{h}{R}r\leq z\leq h. Both x and y are bound by the base radius. Choosing \left | x \right | \leq R implies \left | y \right | \leq \sqrt{R^2-x^2}.



#### Bulge Stellar Mass Function[](#bulge-stellar-mass-function)

By rejection sampling the following function, given by [[40](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id2)], the mass of bulge stars is determined.

For m<0.7 the log-normal distribution equation [(7)](#equation-lognormal) is used. Parameters are A=3.6*10^{-4}, m_{c}=0.22 and \sigma=0.33. For m>0.7 a Salpeter slope [(4)](#equation-salpeter) with parameters A=7.1*10^{-5} and x=1.3 is chosen.

Samples are drawn until the sum of all samples M_s = \sum_{i=1}^n m_i is larger than the total mass M. If \sum_{i=1}^{n-1} m_i +\frac{m_n}{2} > M the last sample is removed.

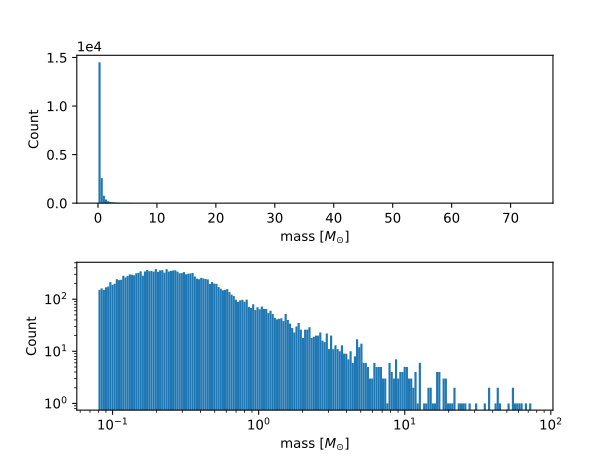


Fig. 7 10^3 stars sampled from the bulge MF[](#id43)

#### Disk Stellar Mass Function[](#disk-stellar-mass-function)

Stars belonging to the disk are given a mass by rejection sampling the PDMF as given by [[40](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id2)].

For m<1 the PDMF reads

(7)[](#equation-lognormal)\xi\left(\mathrm{log}(m)\right) = \frac{\mathrm{dN}}{\mathrm{dlog}(m))} = A*\mathrm{exp}[ \frac{-( \mathrm{log}(m) -\mathrm{log}( m_{c} ))^{2} }{2 \sigma^{2}}]

or equivalently (this version is sampled)

\xi\left(m\right) = \frac{\mathrm{dN}}{\mathrm{dm}} = \frac{A}{m\mathrm{ln}(10)}*\mathrm{exp}[ \frac{-( \mathrm{log}(m) -\mathrm{log}( m_{c} ))^{2} }{2 \sigma^{2}}]

For m>1 the PDMF has the form

\xi\left(\mathrm{log}(m)\right) = \frac{\mathrm{dN}}{\mathrm{dlog}(m))} = A m^{-x}

or depending on m rather than \mathrm{log}(m)

\xi\left(m\right) = \frac{\mathrm{dN}}{\mathrm{dm}} = \frac{A}{m\mathrm{ln}(10)} m^{-x}

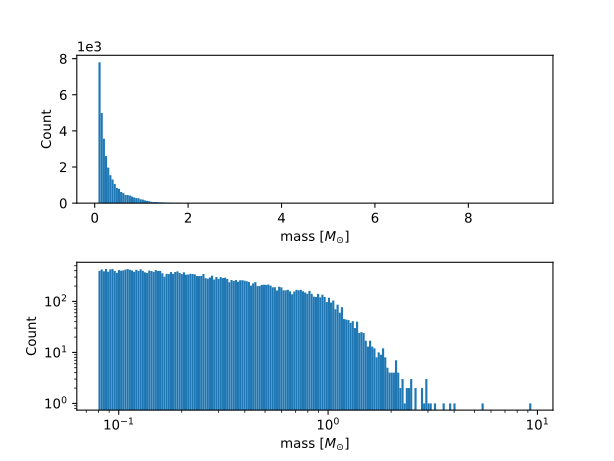


Fig. 8 Stars with total mass of 10^3 M_{\odot} sampled from disk PDMF[](#id44)

### Positions[](#positions)

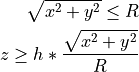
The positions of the field stars within the cone of vision are generated in two steps of rejection sampling followed by the transformation [(6)](#equation-transformation-matrix).

In the first step trial positions are drawn from a uniform distribution within a cuboid containing the cone. The boundaries of the cuboid are given by



where R=h*\textup{tan}\left ( \frac{\alpha}{2} \right ) is the base radius of the cone.

These trial positions are rejected in case they lie outside the boundaries of the cone. The conditions for acceptance are:



This method ensures that the positions are indeed homogeneously distributed, which is essential for the second step.

The second step consists of rejection sampling the density distribution. The test variable is drawn from a uniform distribution ranging from the smallest to the largest possible density within the cone volume. If this test variable is smaller than the density at the trial position generated in step two, the trial position is accepted and rejected otherwise.

Finally, the accepted position is transformed via the transformation matrix [(6)](#equation-transformation-matrix).

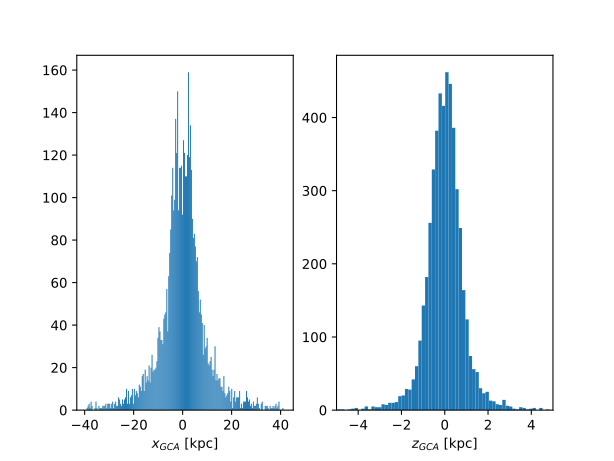
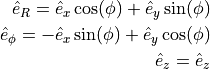


Fig. 9 Spatial distribution of a sample of disc and bulge stars[](#id45)

### Velocities[](#velocities)

#### Particle Kinematics in Cylindrical Coordinates[](#X5d12be1b06fafd82fde171acaee125cde3ca69b)

In terms of Cartesian unit vectors, cylindrical unit vectors are given by



Since \hat{e}_R depends on \phi, position vectors have the form

\vec{q} = R\hat{e}_R+z\hat{e}_z

derivation by time gives the velocity

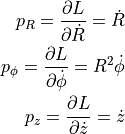
(8)[](#equation-cylindrical-velocity)\vec{v} = \dot{R}\hat{e}_R+R\dot{\phi}\hat{e}_\phi+\dot{z}\hat{e}_z

##### Lagrangian with axisymmetric potential[](#lagrangian-with-axisymmetric-potential)

Using [(8)](#equation-cylindrical-velocity) the Lagrangian per unit mass reads

L =\frac{1}{2}\left( \dot{R}^2+R^2\dot{\phi}^2+\dot{z}^2 \right)+\Phi\left(R,z\right)

and using the Euler–Lagrange equation gives the conjugate momenta

(9)[](#equation-conjugate-momenta)

##### Hamiltonian with axisymmetric potential[](#hamiltonian-with-axisymmetric-potential)

Using the momenta in cylindrical coordinates [(9)](#equation-conjugate-momenta) the Hamiltonian per unit mass with an axisymmetric potential reads ([[6](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id17)] p. 278)

(10)[](#equation-hamilton-1)H = \frac{1}{2}\left( p_{R}^2+\frac{p_\phi^2}{R^2}+p_z^2 \right)+\Phi\left(R,z\right)

using Hamilton’s equations gives

(11)[](#equation-hamilton-2-1)\dot{p}_{R} = -\frac{\partial H}{\partial R} = \frac{p_\phi^2}{R^3}-\frac{\partial \Phi }{\partial R}

(12)[](#equation-hamilton-2-2)\dot{p}_{\phi} = -\frac{\partial H}{\partial \phi} = -\frac{\partial \Phi }{\partial \phi} = 0

(13)[](#equation-hamilton-2-3)\dot{p}_{z} = -\frac{\partial H}{\partial z} = -\frac{\partial \Phi }{\partial z}

Since \vec{L} = \vec{p} \times \vec{\dot{p}} and thus

(14)[](#equation-angular-momentum)L_z = R^2\dot{\phi}

, the second equation above signifies that, in the case of an axisymmetric potential, the z component of the angular momentum is conserved.

#### The phase space distribution function (DF)[](#the-phase-space-distribution-function-df)

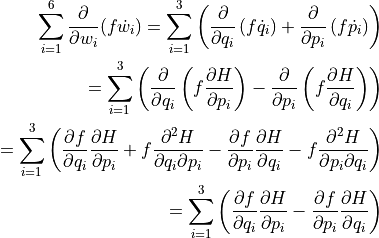
A galaxy can be viewed as a number of stars in phase space. The phase space is a six dimensional space spanned by position \vec{q} and momentum \vec{p}. One point in phase space can therefore be described by a six dimensional vector \vec{w}=(\vec{q},\vec{p}), the state of a whole galaxy as a DF f(\vec{q},\vec{p},t). f(\vec{q},\vec{p},t)d^3\vec{q}d^3\vec{p} is the probability that some specific star is inside the 6D cuboid defined by d^3\vec{q}d^3\vec{p} at time t.

The conservation of probability in phase space is, similarly to the continuity equation in fluid dynamics, given by

\frac{\partial f }{\partial t} + \sum_{i=1}^{6} \frac{\partial}{\partial w_i}(f\dot{w}_i) = 0

with \dot{\vec{w}} = (\dot{\vec{q}},\dot{\vec{p}}).

The right summand can be changed as follows, using Hamilton equations



And therefore

(15)[](#Xa9f6ecfd4d99701c78875269f7fa73c50b8605b)\frac{\partial f }{\partial t} + \sum_{i=1}^{3} \left ( \frac{\partial f}{\partial q_i} \frac{\partial H}{\partial p_i} - \frac{\partial f}{\partial p_i} \frac{\partial H}{\partial q_i} \right) = 0

which is the collisionless Boltzmann Equation (CBE).

The zeroth moment of the DF is the number density:

\nu(\vec{x}) \equiv \int f(\vec{x},\vec{v})d^3\vec{v}

Mean velocities are given by the first moment:

\overline{v_i}(\vec{x}) \equiv \frac{1}{\nu(\vec{x})}\int v_i f(\vec{x},\vec{v})d^3\vec{v}

Todo(?): Continue here (showing = 0 for 2/3 directions)

#### Jeans equations[](#jeans-equations)

Using Einstein notation for i=1,2,3 the CBE [(15)](#Xa9f6ecfd4d99701c78875269f7fa73c50b8605b) is given by ([[6](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id17)] p. 277):

\frac{\partial f}{\partial t} + \frac{\partial f}{\partial q_i}\frac{\partial H}{\partial p_i} - \frac{\partial f}{\partial p_i}\frac{\partial H}{\partial q_i} = 0

Since the galactic potential [(2)](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#equation-total-potential) is axisymmetric, it is convenient to express this equation in cylindrical coordinates.

(16)[](#equation-jeans-equations-1)\frac{\partial f}{\partial t} + p_R\frac{\partial f}{\partial R} + \frac{p_\phi}{R^2}\frac{\partial f}{\partial \phi} + p_z\frac{\partial f}{\partial z} - \left(\frac{\partial \Phi}{\partial R}-\frac{p_\phi^2}{R^3}\right)\frac{\partial f}{\partial p_R} - \frac{\partial \Phi}{\partial \phi}\frac{\partial f}{\partial p_\phi} - \frac{\partial \Phi}{\partial z}\frac{\partial f}{\partial p_z} = 0

It is assumed that the galaxy is statistically in a steady state ([[4](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id18)]) i.e. \frac{\partial f}{\partial t}=0. Due to this assumption and taking [(12)](#equation-hamilton-2-2) into account [(16)](#equation-jeans-equations-1) simplifies to

(17)[](#equation-jeans-equations-2)p_R\frac{\partial f}{\partial R} + \frac{p_\phi}{R^2}\frac{\partial f}{\partial \phi} + p_z\frac{\partial f}{\partial z} - \left(\frac{\partial \Phi}{\partial R}-\frac{p_\phi^2}{R^3}\right)\frac{\partial f}{\partial p_R} - \frac{\partial \Phi}{\partial z}\frac{\partial f}{\partial p_z} = 0

Multiplying equation [(17)](#equation-jeans-equations-2) by p_R and integrating over all momenta leads to (todo: derive?)

(18)[](#equation-jeans-equations-3)\frac{\partial \nu \overline{v_R^2}}{\partial R}+\frac{\partial \nu \overline{v_Rv_z}}{\partial z} + \nu \left ( \frac{\overline{v_R^2}-\overline{v_\phi^2}}{R} + \frac{\partial\Phi}{\partial R}\right ) = 0

#### The Epicyclic Approximation[](#the-epicyclic-approximation)

Individual stars in the disk are on nearly circular orbits. Such orbits can be approximated by circular orbits with additional retrograde elliptical orbits around the guiding center.

The derivation of this approximation starts with Hamilton’s equations for an axisymmetric potential.

Rearranging [(10)](#equation-hamilton-1) and using [(14)](#equation-angular-momentum) gives

H = \frac{1}{2}\left( p_{R}^2+p_z^2 \right)+\frac{L_z^2}{2R^2}+\Phi\left(R,z\right)

With the effective potential given by

\Phi_{\textup{eff}}\left(R,z\right) = \frac{L_z^2}{2R^2}+\Phi\left(R,z\right)

leads to

H_{\textup{eff}} = \frac{1}{2}\left( p_{R}^2+p_z^2 \right)+\Phi_{\textup{eff}}\left(R,z\right)

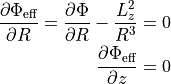
Here \frac{1}{2m}\left( p_{R}^2+p_z^2 \right) is the kinetic energy in the \left( R,z \right) plane or meridional plane. The angular momentum term in the effective potential is not a real potential energy even though sometimes called centrifugal potential. It really is the angular kinetic energy. The given definition of \Phi_{\textup{eff}} is only valid because L_z is conserved.

Using \Phi_{\textup{eff}} [(11)](#equation-hamilton-2-1) and [(13)](#equation-hamilton-2-3) can be written as

(19)[](#equation-hamilton-eff-1)\dot{p}_{R} = -\frac{\partial \Phi_{\textup{eff}} }{\partial R}

(20)[](#equation-hamilton-eff-2)\dot{p}_{z} = -\frac{\partial \Phi_{\textup{eff}} }{\partial z}

These equations describe harmonic oscillations in the effective potential. The minimum of effective potential is the minimum of the real potential energy, together with a contribution from the angular kinetic energy.



The first condition states that the attractive force (-\frac{\partial \Phi_{\textup{eff}} }{\partial R}) has to balance the “centrifugal force”. This is the condition for circular orbits with angular momentum L_z. The second condition is clearly satisfied in the equatorial plane (z=0). The coordinates of this guiding centre are defined as (R_g,\phi_g,z_g).

In preparation for a Taylor series expansion about the guiding centre, x is defined as

x \equiv R - R_g

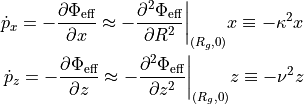
If R = R_g then x=0 and therefore the guiding centre is at (x,z) = (0,0).

\Phi_{\textup{eff}} = \Phi_{\textup{eff}}(R_g,0) + \frac{\partial \Phi_{\textup{eff}} }{\partial R}\bigg|_{(R_g,0)}x + \frac{\partial \Phi_{\textup{eff}} }{\partial z}\bigg|_{(R_g,0)}z + \frac{1}{2}\frac{\partial^2 \Phi_{\textup{eff}} }{\partial R^2}\bigg|_{(R_g,0)}x^2 + \frac{1}{2}\frac{\partial^2 \Phi_{\textup{eff}} }{\partial z^2}\bigg|_{(R_g,0)}z^2 + \frac{1}{2}\frac{\partial^2 \Phi_{\textup{eff}} }{\partial x\partial z}\bigg|_{(R_g,0)}xz + \mathcal{O}(xz^2)

The first order terms are zero (since \Phi_{\textup{eff}} is minimized at the guiding centre) and so is the xz, the later due to symmetric about z=0.

In the epicyclic approximation \mathcal{O}(xz^2) and higher order terms are neglected.

With this approximation [(19)](#equation-hamilton-eff-1) and [(20)](#equation-hamilton-eff-2) become

(21)[](#equation-hamilton-eff-approx)

where the epicyclic frequency \kappa is the frequency of small radial and the vertical frequency \nu the frequency of small vertical oscillations.

With potential \Phi\left(R,z\right) [(21)](#equation-hamilton-eff-approx) can be written as

(22)[](#equation-kappa-squared)\kappa^2 = \frac{\partial^2\Phi}{\partial R^2}\bigg|_{(R_g,0)} + \frac{3L_z^3}{R_g^4}

The combination of [(9)](#equation-conjugate-momenta) with [(11)](#equation-hamilton-2-1), given \dot{p}_{R}=0 and by definition \Omega = \dot{\phi}, results in the circular angular frequency

(23)[](#equation-circular-angular-frequency)\Omega^2 = \frac{1}{R} \frac{\partial \Phi }{\partial R}\bigg|_{(R_g,0)} = \frac{L_z^2}{R^4}

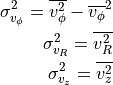
The derivative of [(23)](#equation-circular-angular-frequency) is

(24)[](#equation-circular-angular-frequency-2)\frac{\partial^2\Phi}{\partial R^2} = \Omega^2 + R \frac{d\Omega^2}{dR}

Inserting [(23)](#equation-circular-angular-frequency) and [(24)](#equation-circular-angular-frequency-2) into [(22)](#equation-kappa-squared) yields

(25)[](#equation-epicyclic-frequency)\kappa^2(R_g) = \left ( R\frac{d\Omega^2}{dR} + 4\Omega^2 \right )\bigg|_{R=R_g}

The variance is generally given as the mean of squares minus the square of means. In the epicyclic approximation the mean velocity in R and z is zero and therefor

(26)[](#equation-epicyclic-variance)

#### Disk[](#disk)

The velocity distribution of stars in the Milky Way disk is approximated with the help of Jeans equations as well as relations and constraints based on observations.

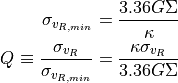
For a flat rotation curve the radial velocity dispersion exponentially decreases with increasing radius [[59](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id14)] (p. 114)

(27)[](#equation-radial-velocity-dispersion-1)\sigma_{v_{R}} \propto e^{-\frac{R}{h}}

where h in the case of the Miyamoto Nagai potential is the radial scale length a.

The relation [(27)](#equation-radial-velocity-dispersion-1) still requires a constant factor k, which can be determined by means of the Toomre parameter Q at some distance R_{ref}.

Q is the ratio between the actual and minimum velocity dispersion \sigma_{v_{R,min}} [[57](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id16)] (p. 1234)



where \kappa denotes the epicyclic frequency [(25)](#equation-epicyclic-frequency).

In the solar neighborhood Q_{\ast} = 2.7 \pm 0.4 and \sigma_{v_{R}} = (38 \pm 2) \left [ \frac{km}{s} \right ] [[6](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id17)] (p. 497)

The constant k can therefore be approximated via

(28)[](#Xd371b04e39731d8551a17d1e3233c50cc38d825)k \cong Q \sigma_{v_{R,min}}e^{\frac{R}{h}}

Under the approximation of isothermal sheets (introduced in [[59](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id14)]), the vertical velocity dispersion only depends on the surface density [[58](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id15)]

(29)[](#equation-vertical-velocity-dispersion)\sigma_{v_{z}} = \pi G \Sigma \left ( R \right )z_{0}

with z_{0} being the vertical scale length b when using the Miyamoto Nagai potential.

Combining the equation for the circular velocity [(1)](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#equation-circular-velocity) and Jeans equation [(18)](#equation-jeans-equations-3) and multiplication by \frac{R}{\nu} results in

(30)[](#equation-disk-velocity-step-1)\frac{R}{\nu} \frac{\partial \nu \overline{v_R^2}}{\partial R} + \frac{R}{\nu} \frac{\partial \nu \overline{v_Rv_z}}{\partial z} + \overline{v_R^2}-\overline{v_\phi^2} + v_c^2 = 0

For a razor thin disc the spatial density \nu can be replaced with the surface density \Sigma [[45](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id19)], which does not depend on z. With this and by using relations [(26)](#equation-epicyclic-variance), [(30)](#equation-disk-velocity-step-1) simplifies to

(31)[](#equation-disk-velocity-step-2)\frac{R}{\Sigma} \frac{\partial \nu \sigma_{v_{R}}^2}{\partial R} + R \frac{\partial \overline{v_Rv_z}}{\partial z} + \sigma_{v_{R}}^2 -\sigma_{v_{\phi}}^2 - \overline{v_\phi}^2 + v_c^2 = 0

In the epicyclic approximation, azimuthal and radial dispersion have the following relation ([[6](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id17)] p. 170)

(32)[](#equation-azimuthal-velocity-dispersion)\frac{\sigma_{v_{\phi}}^2}{\sigma_{v_{R}}^2} = \frac{\kappa^2}{4\Omega^2}

With this relation [(31)](#equation-disk-velocity-step-2) can be written as

\frac{R}{\Sigma} \frac{\partial \left ( \Sigma \sigma_{v_{R}}^2 \right ) }{\partial R} + R \frac{\partial \overline{v_Rv_z}}{\partial z} + \sigma_{v_{R}}^2 - \sigma_{v_{R}}^2 \frac{\kappa^2}{4\Omega^2} - \overline{v_\phi}^2 + v_c^2 = 0

rearrangements lead to

(33)[](#equation-disk-velocity-step-3)\overline{v_\phi}^2 - v_c^2 = \sigma_{v_{R}}^2 \left ( 1 - \frac{\kappa^2}{4\Omega^2} + \frac{R}{\Sigma \sigma_{v_{R}}^2 } \frac{\partial \left ( \Sigma \sigma_{v_{R}}^2 \right ) }{\partial R} + \frac{R}{\sigma_{v_{R}}^2} \frac{\partial \overline{v_Rv_z}}{\partial z} \right )

Due to the exponential dependency of the surface mass density ([[6](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id17)] p. 325) and the radial dispersion \sigma_{v_{R}} [(27)](#equation-radial-velocity-dispersion-1) on R, the derivation after R summand can be solved as follows

\frac{R}{\Sigma \sigma_{v_{R}}^2 } \frac{\partial \left ( \Sigma \sigma_{v_{R}}^2 \right ) }{\partial R} = \frac{1}{\Sigma \sigma_{v_{R}}^2 } \frac{\partial R }{\partial \ln(R)}\frac{\partial \left ( \Sigma \sigma_{v_{R}}^2 \right ) }{\partial R} = \frac{1}{\Sigma \sigma_{v_{R}}^2 } \frac{\partial \left ( \Sigma \sigma_{v_{R}}^2 \right ) }{\partial \ln(R)} = \frac{1}{\Sigma \sigma_{v_{R}}^2 } \left ( - \frac{R\Sigma\sigma_{v_{R}}^2}{a} - \frac{R\sigma_{v_{R}}^2\Sigma}{a} \right ) = - 2 \frac{R}{a}

With this and under the assumption that the retrograde elliptical orbit is aligned with the coordinate axes [(33)](#equation-disk-velocity-step-3) becomes

(34)[](#equation-disk-velocity-step-4)\overline{v_\phi}^2 - v_c^2 = \sigma_{v_{R}}^2 \left (1 - \frac{\kappa^2}{4\Omega^2} -\frac{R}{a } \right )

The velocity components can be sampled from Gaussian distributions. Dispersions are obtained by first evaluating the epicyclic frequency [(25)](#equation-epicyclic-frequency) and the surface mass density [(3)](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#equation-surface-mass-density), followed by the radial velocity dispersion [(27)](#equation-radial-velocity-dispersion-1) with constant factor [(28)](#Xd371b04e39731d8551a17d1e3233c50cc38d825), the vertical [(29)](#equation-vertical-velocity-dispersion) and the azimuthal velocity dispersion [(32)](#equation-azimuthal-velocity-dispersion). As noted before, the mean velocity in R and z is zero. For \phi the mean velocity is calculated from [(34)](#equation-disk-velocity-step-4), where the circular velocity v_c is given by [(1)](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#equation-circular-velocity).

#### Bulge[](#bulge)

For spherically symmetric distribution functions, such as, with the present approximation, the bulge, the spherically symmetric Jeans equation can be used [[37](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id20)]

(35)[](#equation-bulge-velocity-1)\frac{\partial (\rho \sigma_r^2)}{\partial r} + 2\frac{\beta\sigma_r^2}{r} + \rho\frac{\partial \Phi}{\partial r} = 0

where \beta is an indicator for anisotropy in radial and tangential direction.

\beta = 1 - \frac{\sigma^2_\phi}{\sigma^2_r}

If the distribution is isotropic, as assumed here, \beta=0 and [(35)](#equation-bulge-velocity-1) simplifies to

\frac{\partial (\rho \sigma_r^2)}{\partial r} + \rho\frac{\partial \Phi}{\partial r} = 0

and the radial dispersion can be calculated by integration

\sigma_r^2 = \frac{1}{\rho}\int_{r}^{\infty}\rho \frac{\partial \Phi}{\partial r}\textup{dr}

However, \Phi is the composite potential consisting of bulge, disc and halo and therefore not spherically symmetric. As suggested in [[45](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id19)], contribution by the disc to the bulge dispersion may be accounted for by spherically averaging the disc potential. This is achieved by expressing the cylinder in spherical coordinates and integrating over the relevant angle

\sigma_{r,disk}^2 = \frac{1}{\rho}\int_{r}^{\infty} \frac{\rho }{2\pi} \int_0^{2\pi} \frac{\partial \Phi\left ( r\sin(\theta ),r\cos(\theta ) \right )}{\partial r} \textup{d}\theta \textup{dr} \\

As neither of these integrals have an analytical solution, this calculation is computationally intensive. Therefore a lookup table has been implemented in which discrete r and corresponding bulge dispersion are stored.

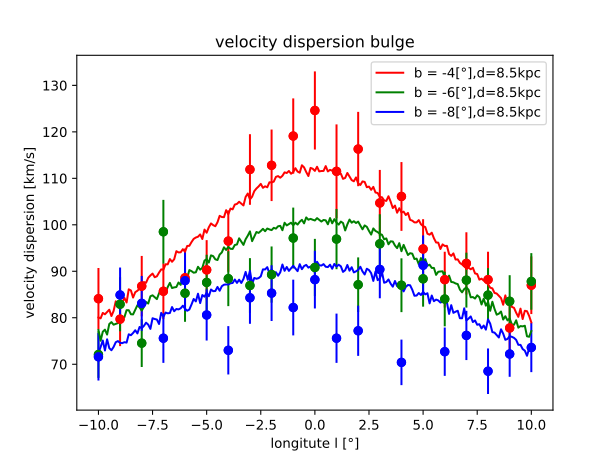


Fig. 10 Velocity dispersion of bulge stars[](#id46)

In the figure above results are compared with observations from [[20](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id21)]. These results were generated using a view distance of 7.5kpc. The other parameters are given in [Table 1](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#tbl-parameters). The radial velocity dispersion heavily depends on the scale parameter of the bulge potential. With the current choice, the dispersion at small r is lower than realistic. This issue could be solved by adding an additional potential with smaller scale radius. A multi component bulge model has been suggested for instance by [[31](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id22)].

## Integration[](#integration)

Numerical integration is needed during initialization and simulation. Various options for time integration have been implemented. The GNU Scientific Library ([[14](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id24)], GSL) is used for integrations performed during initialization. The relevant functions are all based on QAG or QAGI, which have been ported from the Fortran library QUADPACK [[52](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id25)] to C in GLS. The decision trees given on page 79-80 in [[52](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id25)] help with the decision on when and how to use the respective methods.

### Quadrature, Adaptive, General-purpose (QAG)[](#quadrature-adaptive-general-purpose-qag)

This algorithm makes use of adaptive Gauss-Kronrod quadrature to estimate the definite integral of a given function.

Quadrature sums are defined as

Q_n[a,b] \equiv \sum_i^nw_if(x_i) \cong \int_a^bw(x)f(x)\textup{d}x

where w_i are weights, x_i nodes, w(x) a weight function. The highest possible degree of precision is 2n-1. With this maximum precision Q_n is exact for polynomials of a degree smaller or equal to 2n-1.

Using classical Gaussian quadrature formulae, error estimation, by increasing n to n+1, requires n+1 evaluations of f(x) in addition to n evaluations from calculation of the original sum, since the respective nodes have no common points. By doing so, the degree of precision is only increased from 2n-1 to 2n+1. Therefore, the error estimation obtained by subtracting the two sums could be unreliable.

Adding n+1 points to the Gauss-Legendre formula - here w(x)=1 and the nodes are zeros of the Legendre polynomial - Kronrod introduced the option of increasing the precision to 3n+1, again requiring n+1 additional evaluations of f(x) ([[27](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id26)]).

Q_n^K[a,b] \equiv \sum_i^nw_if(x_i) + \sum_j^{n+1}w_j^*f(x_j^*) \cong \int_{-1}^1f(x)\textup{d}x

QAG makes use of this option, bisecting the interval with the largest local absolute error estimate in each step. This division is repeated until either the absolute or relative global error estimate is smaller than required by the caller.

### Quadrature, Adaptive, General-purpose, Infinite interval (QAGI)[](#Xb2a0c1f53142e4fe7ef31cb84369df8c38e8fef)

In case of a semi-infinite interval (a,\pm \infty), the integration variable is transformed

x = a\pm\frac{1-t}{t}

leading to

\int_a^{\pm \infty}f(x)\textup{d}x = \pm \int_0^1 f \left(a\pm\frac{1-t}{t}\right)t^{-2}\textup{d}t

For an infinite interval

\int_{-\infty}^{\infty}f(x)\textup{d}x = \int_0^\infty f(x)+f(-x)\textup{d}x = \int_0^1 \left ( f \left(\frac{1-t}{t}\right) + f \left(\frac{t-1}{t}\right) \right ) t^{-2}\textup{d}t

After the transformation QAGS with the 15-point Kronrod rule is used. QAGS, in addition to the adaptive bisection (see QAG), makes use of the Wynn \epsilon-algorithm to accelerate the convergence.

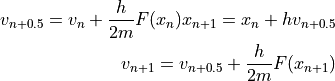
### Velocity Verlet Algorithm[](#velocity-verlet-algorithm)

For cluster members the acceleration is a combination of the force resulting from the presence of all other cluster stars (see Barens Huts Algorithm) and from the milky way potential. The acceleration of field stars solely comes from the milky way potential. In each time step both velocity and acceleration of each star is evaluated.

Since the velocity, v(t) changes over time, its value at the midpoint between the current (t_n) and the next timestep t_{n+1} is intuitively a better approximation than v(t_n) or v(t_{n+1}). The same holds true for the acceleration. This leads to [[13](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id27)]

(36)[](#equation-leapfrog-algorithm)x_{n+1} = x_{n} + hv_{n+0.5}\\ v_{n+1.5} = v_{n+0.5} + \frac{h}{m}F(x_{n+1})

which is the Leapfrog algorithm. If one needs x and v at the same time, [(36)](#equation-leapfrog-algorithm) can be split into two half steps.



does not have to be calculated twice, because F(x_{n+1}) can be used as F(x_{n}) in the next timestep.



### Barnes-Hut Algorithm (BH)[](#barnes-hut-algorithm-bh)

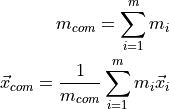
When determining the gravitational force acting on a star belonging to a cluster, the positions of all other stars in that cluster have to be taken into account.

m_i\vec{x}_i = G\sum_{j=1,j\not\equiv i}^{N}\frac{m_im_j\left ( \vec{x}_j-\vec{x}_i \right )}{\left | \vec{x}_j-\vec{x}_i \right |^3}

Calculating this force for all stars requires O(n^2) operations. However, the simulated clusters consist of \sim 10^4 - 10^5 stars. While the resulting number of calculations is possible, it is not feasible for a typical desktop computer. Therefore, the Barnes-Hut algorithm (BH) has been implemented which is of order O(n\log(n)).

The gist of the BH is to approximate a set of stars by their total mass and center of mass (com) if the distance between them and the star, for which the force is to be calculated, is large enough.

The total mass and com of a set of m stars is



All cluster stars are stored in an octree. An octree is a data structure where each node in the tree has up to eight child nodes. These nodes split the space represented by their parent node into eight cubes. External nodes are nodes without any children. Each external node contains at most one star. Internal nodes have at least one child. They represent stars stored in their child nodes by storing their total mass and com. The root node contains the whole space occupied by the cluster. Each node stores the following information: total mass, amount and center of mass of stars contained within the cube, two points defining the volume of the cube, one point at the center of the cube, since cpu time is more valuable than ram, as well as links (pointers) to each child node and to the parent node. If a child pointer is null, it does not exist yet.

Stars are added recursively starting at the root node. If the current node is already an internal node, the star is passed to the appropriate child. Mass and com of the internal node are updated. The appropriate child is determined by comparing the position of the star with the center of the node. If the considered node is an external node but already contains a star, both the newly added and already present star are passed down to the appropriate child or children. Consequently, the current node becomes an internal node. Since both stars can lie in the same octant, this can lead to additional recursions until the stars are assigned to different child nodes. If the current node is external and does not yet contain a star, the star is added to the node and the recursion ends.

When calculating the gravitational force on a star, the octree is travelled through recursively beginning with the root node. In case the distance d between the star and a node is sufficiently large, the stars within that node are approximated by the mass and com of that node, otherwise all child nodes within the current node are considered. Whenever the distance criterion is met, the acceleration vector stemming from the force is calculated, added to the overall acceleration of the star and the recursion for the considered branch stops.

Whether or not d is big enough, is determined by the quotient \theta.

\theta = \frac{s}{d} < \theta_{max}

with s the side length of the cube and \theta_{max} a set threshold value. In the special case of \theta_{max}=0, BH becomes a direct-sum algorithm. \theta_{max}=0.5 is a commonly chosen value.

todo?: explain smoothing

## Coordinate Systems[](#coordinate-systems)

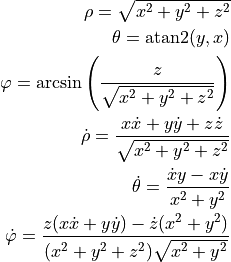
In the simulation, locations and velocities of stars are stored in galactocentric cartesian coordinates (GCA). However, the observer/scopesim expect heliocentric equatorial polar coordinates. Moreover, observational data is usually given in heliocentric galactic polar coordinates. This data is used for initial cluster positions in the simulation and to compare results. The implementation of transformations is therefore inevitable. The code has been adapted from GalPot ([[48](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id23)]).

### Galactocentric Cartesian (GCA)[](#galactocentric-cartesian-gca)

GCA is a right-handed coordinate system with the galactic center in its origin. The projection of \hat{e}_x onto the galactic equator (or midplane) points to the initial location of the sun and \hat{e}_z towards the galactic north pole. Therefore, the direction of galactic rotation at the location of the sun is the negative z axis.

### Galactocentric Polar (GCP)[](#galactocentric-polar-gcp)

GCP is a spherical coordinate system and similar to GCA. Position and velocity transformation between the two systems goes as follows.

(37)[](#equation-gca-to-gcp)

### Local Standard of Rest (LSR)[](#local-standard-of-rest-lsr)

Like GCA, LSR is a right-handed coordinate system. The origin of positions is the location of the sun and the origin of velocity is the velocity of a star on a circular orbit with mean velocity of stars in the solar neighborhood. \hat{e}_x points towards the galactic center, \hat{e}_y towards the direction of galactic rotation and \hat{e}_z approximately towards the galactic north pole.

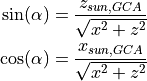
The location of the sun is given by ([[25](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id28)], [[5](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id29)])

\vec{x}_{sun,GCA} \cong (8.20,0,0.014)[kpc]

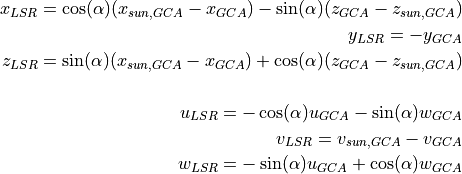
and the mean velocity ([[25](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id28)])

\vec{v}_{mean,GCA} \cong (0,-232.8,0)[km/s]

Since the sun is generally not in the galactic midplane, there is an angle between the planes spanned by (x,y)_{GCA} and (x,y)_{LSR}. This angle can be expressed as



The transformation of position and velocity vectors from GCA to LSR has to contain a rotation by -\alpha about the y axis.



### Heliocentric Cartesian (HCA)[](#heliocentric-cartesian-hca)

The only distinction between HCA and LSR is the origin of the velocity. In HCA the origin is the velocity of the sun. The difference between the average velocity in the solar neighborhood and the sun itself, also called the peculiar motion of the sun, is given by ([[56](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id30)])

\vec{v}_{sun,LSR} \cong (11.1, 12.24, 7.25)[km/s]

Transformation from LSR to HCA can be written as

\vec{x}_{HCA} = \vec{x}_{LSR} \\ \vec{v}_{HCA} = \vec{v}_{LSR} - \vec{v}_{sun,LSR}

### Heliocentric Galactic Polar (HGP)[](#heliocentric-galactic-polar-hgp)

HGP is a spherical coordinate system with origins for position and velocity identical to those defined in HCA. Coordinates given in this system are often called galactic coordinates. The galactic longitude l is the angular distance in the galactic midplane with l=0^{\circ} towards the galactic center. The galactic latitude b denotes the angle below and above the galactic midplane ranging from -90^{\circ} to 90^{\circ}. and r the radial distance.

The transformation from HCA to HGP is identical to the transformation from GCA to GCP [(37)](#equation-gca-to-gcp) with l \equiv \varphi, b \equiv \theta and \rho \equiv r

### Heliocentric Equatorial Polar (HEQ)[](#heliocentric-equatorial-polar-heq)

The following positions and the transformation between HGP and HEQ are described in [[9](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id32)] (p 1044-1047), the transformation between different epochs in [[33](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id31)] (p 95-105) . The latter is dated. The international celestial reference system (ICRS) should be preferred if highest possible accuracy is of concern.

HEQ, like HGP, is a spherical coordinate system having the same origins for position and velocity. However, angles are given in, and normal to, the celestial equator which is not parallel to the galactic midplane. The right ascension (a) is the angular distance in the equator with a=0^{\circ} towards the northward equinox. The equinox is the intersection of the ecliptic - the plane in which the earth orbits the sun - and the celestial equator. The declination (d) is the angular distance above or below in the equator.

Since the ecliptic and the equator are in motion, a reference frame is needed. A reference frame consists of quantities defining the coordinate system at a specific time as well as methods to calculate those quantities for any other date. A commonly used reference frame is defined for the J2000.0 epoch (\epsilon_0).

In order to transform between HCA and HEQ at \epsilon_0, the direction of the north Galactic pole (NGP) and the galactic center (GC) are needed in both basis.

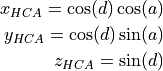
In HCA the NGP is simply \vec{x}_{NGP,HCA}=(0, 0, 1). In HGP, since the direction is normal to the fundamental plane, b=90^{\circ}_{GC,HGP}. In HEQ at \epsilon_0 the direction is

a_{NGP} = 12^h51^m26.28^s \\ d_{NGP} = 27^{\circ}7^\prime41.7^{\prime\prime}

The GC defines the x axis of HCA: \vec{x}_{GC,HCA}=(1, 0, 0). In GC the same direction is

a_{GC,HEQ} = 17^h45^m40.0409^s \\ d_{GC,HEQ} = -29^{\circ}0^\prime28.118^{\prime\prime}

To express these basis vectors in HCA basis, they can to be transformed as follows



The third basis vector is the cross product of \vec{x}_{NGP} and \vec{x}_{GC}. With these basis vectors the change of basis matrix is

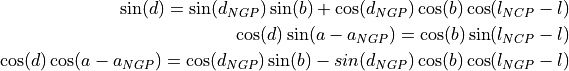
M = \left [\hat{e}_x,\hat{e}_y,\hat{e}_z\right ]

The full transformation from HCA to HEQ consists of the two steps: the multiplication with M followed by the transformation from Cartesian to spherical as given in [(37)](#equation-gca-to-gcp).

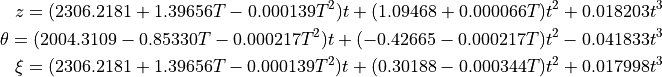
For the transformation between HGP and HEQ the direction of the north celestial pole (NCP) is required. NCP is perpendicular to the celestial equator, hence d_{NGP} = 90^{\circ}. In HGP at \epsilon_0, NCP is

l_{NCP} = 123^{\circ}55^\prime55.2^{\prime\prime}\\ b_{NCP} = 27^{\circ}7^\prime41.7^{\prime\prime}

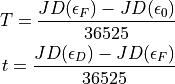
Using NGP and NCP the transformation from HGP to HEQ at \epsilon_0 is



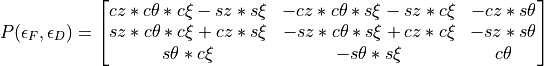
Three angles describing the precision of both planes are needed in order to transform between epochs \epsilon_F and \epsilon_D.



where t and T are time differences given in units of Julian century.



With these three rotations, a precession matrix P as well as its inverse can be formalized.



where cz = \cos(z), sz = \sin(z), c\theta = \cos(\theta), s\theta = \sin(\theta), c\xi = \cos(\xi) and s\xi = \sin(\xi).

Before multiplication with P, the target vector has to be expressed in Cartesian coordinates.

### Heliocentric Telescope Polar (HTP)[](#heliocentric-telescope-polar-htp)

This coordinate system aims to describe positions as they are observed with a telescope. As with HEQ the right ascension a describes the angular distance in and the declination d the angular distance below and above the midplane. The origin, however is defined by the line of sight vector, that is \hat{v}_{los} = [r,a,d] = [1,0,0]_{HTP} . Hence one image, taken with a telescope, is a circle in the ad-plane.

It may seem as if the transformation from HGP to HTP and vice versa simply consists of a shift by \hat{v}_{los} in HGP basis. However for any line of sight vector with b\neq0, the result would be an ellipse rather than a circle. Casually speaking, the cone of vision in any spherical coordinate system is a circle when projected onto the unit sphere. While this circle always covers the same range in b, the range in l depends on the distance (b) between the circle and the midplane. With increasing absolute value of b the circles around the unit sphere described by l decrease in size.

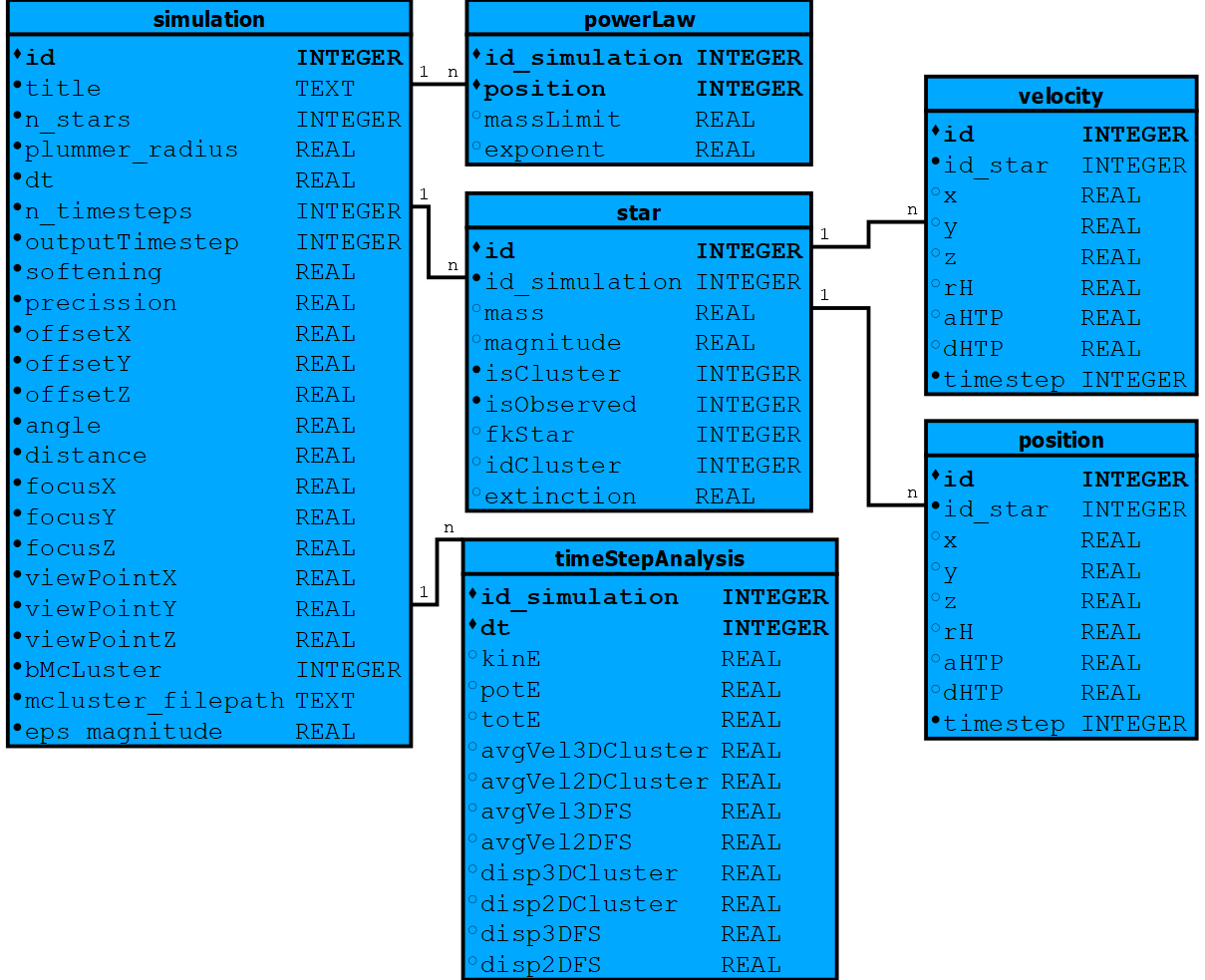
While there may be a solution for this issue, a different transformation was devised. Positions can be transformed from HCA to HTP by aligning \hat{v}_{los} with the x axis hat{x} using a rotation matrix generated as described in [(6)](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#equation-transformation-matrix) with the rotation axis defined by [1,0,0]_{HCA} and \hat{v}_{los,HCA} Subsequent transformation from Cartesian to polar coordinates as described in the section GCP above completes the transformation.

## Database[](#database)

A SQLite Database is used to store simulations including all relevant parameters, stars with their respective positions and velocities as well as data resulting from analysis. SQLite supports both C++ and python. And is therefore used in most data related parts of the project. Generated data can be loaded and plotted directly from the database. Previously performed simulations can be loaded from the database in order to carry out analysis like energy vs time or average star velocity vs time.

Due to bad performance, no foreign key constraints where set for position.id\_star and velocity.id\_star. A multi-column index is used to improve the execution time of queries containing both velocity.timestep and velocity.id\_star. NULL values are avoided where possible. However, at many stages during execution, some information is unknown, but creating entries is still required. Observed stars contain no information about their mass or GCA phase space coordinates, some cannot be mapped to simulated stars, etc.

### Entity Relationship Diagram[](#entity-relationship-diagram)



## Mock Observations[](#mock-observations)

Images of stars are generated using ScopeSim [[22](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id50)] with the MICADO\_Sci instrument, the ELT telescope and the Armazones location package. Parameters differing from default values are

1. SIM.sub\_pixel.flag - Initial tests indicated the necessity of integrating sub pixel shifts
2. DET.width and DET.height - Width and height of the detector plane in pix are set ensuring the image contains the whole COV. The pixelscale scale (arcsec/pixel) the ELT is 0.004
3. OBS.dit - The exposure time is adjusted to one hour.
4. scao\_const\_psf.convolve\_mode = “same” ??? what does this do.
5. scao\_const\_psf.rotational\_blur\_angle - to reflect the rotation of the point spread function (PSF) occurring during one hour, this angle is set to 15 degrees. However, with the approach discussed in [Masking bright stars](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#masking-bright-stars-label) this will have little to no impact on the results.

For each simulation two images, one for t=0 and one for math:t=dt, are stored as FITS files.

### Creating Source Objects[](#creating-source-objects)

ScopeSim requires one source object for each simulated star. Source objects contain spatial and spectral information. The spatial information consists of x and y in arcsec relative to the FOV axis, here generated by a chain of transformations from GCA to HTP and stored in the appropriate database table.

The spectral information is acquired using multiple functions from the ScopeSim Templates library involving the following steps:

1. The spectral type of a star is determined by looking up the closest spectral type from “A Modern Mean Dwarf Stellar Color and Effective Temperature Sequence” [[50](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id51)] given its mass.
2. The spectrum is taken from Pickles catalogue [[51](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id52)] requiring metallicity in range 0.5 - 2 solar and age of about 5 Gyr.
3. That spectrum is scaled such that it has 0 apparent Vega mag in the V filter.
4. The absolute magnitude M_V of the star in the V filter is determined from its mass using [[50](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id51)] with linear interpolation.
5. The apparent magnitude m_V is obtained from m_V = M_V + 5\log_{10}\left ( d \right )-5
6. The interstellar extinction e is taken into account using [[24](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id53)]
7. The weight w of the spectrum s is defined as

F\left (x,y \right ) = s * w

and calculated via

w = \frac{F}{F_0} = 10^{0.4\left ( m_0-(m_V+e) \right )}=10^{-0.4 (m_V+e)}

## Data Reduction[](#data-reduction)

With the help of Photutils [[8](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id54)] the 2D HTP positions and fluxes of stars are detected from the FITS files generated during observation. In a first step, all stars are stored in the database, each having exactly one location corresponding to the timestep of the FITS file.

The following method, while yielding decent results, is certainly not optimal. Due to hardware and time constraints, options such as Image Segmentation were not feasible. Moreover, the parameters of the chosen method were not optimized beyond some spot checks.

### Detecting Stars[](#detecting-stars)

#### Background[](#background)

Testing FITS files generated with no input sources still yield some amount of detected sources. With the [Parameters](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#parameters-label) used for the 25 simulations, 125 sources were detected from an empty background. The FITS files contain a raster of 64 images, sometimes overlapping and sometimes separated by one pixel due to rounding. This leads to wrong detections at the corners. However, this effect only explains a fraction of the detections.

#### Masking bright stars[](#masking-bright-stars)

PSFs of bright stars were wrongfully detected as stars. Increasing the detection threshold would have meant not detecting real faint stars in different areas. Hence the decision was made to mask square areas around bright stars leading to only one detection within those areas.

The side length l_i of the square depends on the flux F_i of the stars: l_i = f(F_i).

To find the appropriate function, FITS for single stars at a given distance and with varying mass were generated, the sources in each file detected, using DAOStarFinder, and their flux and maximum distance between the correct and any wrong sources calculated. A linear fit of the resulting dataset ([Appendix](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#appendix-label)) lead to the “empirical” function

(38)[](#equation-side-length)l_i = \begin{cases} & 0\text{ if } F_i < 100\\ & 0.01*F_i+28\text{ if } F_i \geqslant 100\\ \end{cases}

ToDo?: Add 3 images “without mask” “drawn mask” “after mask”

The DAOStarFinder method is called twice:

1. To find the bright stars and generate the mask. The resulting table contains one row for each source.

* This table is sorted by the flux column in descending order and iterated from top to bottom until the current entry has F_i < 100.
* Elements of the mask - a 2D Boolean array with the same size as the image - are updated.
* All elements inside the box with side length [(38)](#equation-side-length) are set to true and the current table entry stored in a new table if located outside a masked area.

1. For passing the mask parameter generated in the previous step and returning sources outside the masked areas.

Both the bright sources recorded after the first and the faint sources returned from second call are stored in the database.

## Cluster analysis[](#cluster-analysis)

The objective is to detect all cluster stars present in the observed data with a certain accuracy. One image contains a set of stars s_{i} with positions \vec{x}_{i}\left ( t \right ) in pixel coordinates and apparent magnitudes m_{i}. This information is not enough to detect cluster memberships. The number density of cluster stars in position space is similar to that of the field stars. Furthermore, cluster and field star positions overlap significantly. Moreover, depending on the distance of the cluster to the observer, the cluster size and the field of view angle, cluster stars often cover a significant area of if not the whole image. The proper motion \vec{v_{i}}(t) density of cluster stars on the other hand is highly relative to that of field stars and hardly any field stars have \vec{v_{i}}(t) close to cluster stars.

### Velocity approximation[](#velocity-approximation)

The proper motion of a star can be approximated from its position at two different timestamps via first order Taylor expansion

v_{i} = \frac{x_{i}\left ( t+dt \right )-x_{i}\left ( t \right )}{dt}

However, this is only possible if the positions at two different timesteps can be accurately associated with the same star s_{i}. The assumption that \vec{x}(t+dt) belongs to the same star as \vec{x}(t) if their Euclidian distance is smaller than the distance of \vec{x}(t) to any other \vec{x}(t+dt), only hold true for small dt. However, due to the discreteness of images, dt has to be large enough so that the change in position is detectable between timesteps. The pixel scale - the ratio of arcsec to pixel - dictates a lower bound for dt.

Since the data results from a simulation, it is trivial to verify whether or not the attribution of the two positions is indeed correct. In fact, tests using only the Euclidian distance as a metric with dt = 1 day lead to a small but significant amount of wrong assignments.

An additional condition was introduced. Stars at \vec{x}(t) and \vec{x}(t+dt) have to have a similar apparent magnitude m to be considered the same star.

\left | 1-\frac{m_{i}\left ( t \right )}{m_{j}\left ( t+dt \right )} \right | < \varepsilon_{m}

where \varepsilon_{m} is the maximum relative difference in apparent magnitude.

With this constraint, all positions where correctly assigned during further tests.

### DBSCAN[](#dbscan)

The clustering algorithm of choice is DBSCAN since it is density based and able to detect clusters of arbitrary shape. Additionally and contrary to other algorithms the amount of clusters to find is not a parameter. DBSCAN detects any clusters present in the data based on two parameters:

1. \epsilon: the maximum distance between points to be considered neighbors
2. nPoints: the minimum amount of neighbors for a point to be classified as core point.

During execution all data points are classified as one of the following:

1. core point: a point with at least nPoints within \epsilon:
2. border point: a point having at least one core point but less than nPoints within \epsilon:
3. noise/outlier: any other point

The implementation of DBSCAN can be summarized as follows: Iterate the list of points. If the current point is not already classified, check if it meets the requirements to be classified as core point. Once a core point has been found, the neighboring points of that point are tested. If they too have enough neighbors the recursion continues until all neighbors are classified as either core or border points.

In a naive implementation, the distance of each point to every other point is checked. The time complexity of such an implementation is O(n^2). Moreover, for large datasets the recursion can lead to stack overflow.

R\*-tree or similar data structures can be used to improve the performance to an average of O(n\log{n}) ([[12](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id33)]) The library mlpack ([[10](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id34)]) was used, which includes an implementation of DBSCAN supporting R\*-tree and many other trees.

Initially and in addition to the standard condition \epsilon_{x} for the spatial distance between points a supplementary condition \epsilon_{v} was introduced. The difference in velocity between two stars has to be smaller than \epsilon_{v} for them to be classified as neighbors:

\left \|\vec{v}_{1}-\vec{v}_{2} \right \|_{2}< \epsilon_{v}

Large \epsilon_{x} lead to more accurate membership detection. It turned out, the spatial distance condition does not benefit the results at all and it was dropped. For larger areas than used here, constraining the spatial distance, for instance via subdivision, should be beneficial.

### Performance[](#performance)

Observed stars are mapped to simulated stars via their proximity in order to measure the performance. If a star is not the closest observed star to any simulated star it remains not mapped. Not mapped stars exist due to Background (Todo: Link) and parts of the PSFs of bright stars being detected as stars. Observed stars can then be attributed one of the following types:

1. True Positive (TP): correctly classified as cluster star.
2. False Positive (FP): wrongly classified as cluster star.
3. True Negative (TN): correctly classified as field star.
4. False Negative (FN): wrongly classified as field star.
5. Unconfirmed Positive (UP): not mapped star classified as cluster star. Treated as FP.
6. Unconfirmed Negative (UN): not mapped star classified as field star. Not taken into account since FN is unlikely.

#### Accuracy[](#accuracy)

A = \frac{TP+TN}{TP+TN+FP+FN}

With a large amount of field stars relative to cluster stars, this metric is not ideal as it will give a good rating even if most cluster stars are FN.

#### Precision and Recall[](#precision-and-recall)

When FPs are more problematic than FNs, the precision P should be high

P = \frac{TP}{TP+FP+UP}

On the flip side, if FNs are a big concern, but FPs tolerable, the recall R is a good metric

R = \frac{TP}{TP+FN}

#### F1 Score[](#f1-score)

This metric is a balance between P and R. Contrary to A TN is not taken into account.

F_1 = 2 \frac{P*R}{P+R} = \frac{TP}{TP+0.5(FP+UP+FN)}

## Experiments[](#experiments)

### Parameter optimization[](#parameter-optimization)

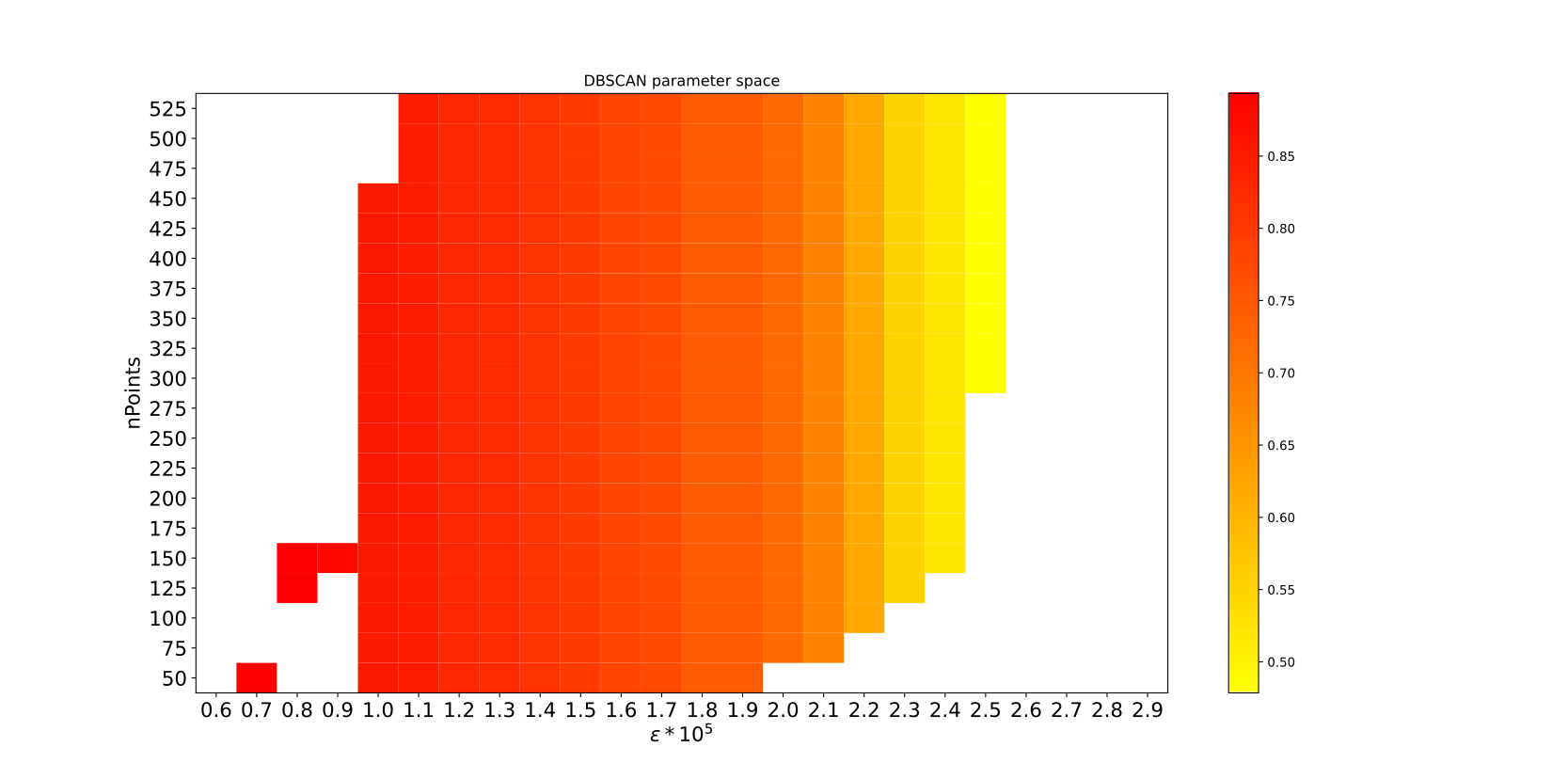


Fig. 11 Precision depending on DBSCAN parameters[](#id2)

The quality of the cluster analysis with DBSCAN depends on the choice of its parameters. [Fig. 11](#fig-dbscan) depicts the precision for a 1 [kM_\odot] cluster 300 \textup{[pc]} from the GC. Decreasing \epsilon leads to a higher precision until it is too small, at which point multiple clusters are detected. Based on this observation \epsilon was set to 1.1*10^{-5} and nPoints to 200 for all simulations.

### Setup[](#setup)

A total of 25 simulations with varying focus and cluster mass were carried out in order to study the effects of these parameters on the cluster detection performance. The focus was set to l=0^{\circ} and b \in \left \{0,5,10,25,180 \right \}[^\circ] HGP and the cluster mass \in \left \{0.64, 1.6, 4.0, 10, 25 \right \} [\textup{kM}_\odot]. Each simulation was repeated 10 times for error estimation.

#### Parameters[](#parameters)

The following parameters remained unchanged between the simulations.

**General simulation parameters**

* FOV angle: 54 \textup{[arcsec]}
* The field of view angle, chosen large enough to cover most of the clusters with the highest mass. See \alpha in [Fig. 6](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#fig-cone)
* View distance: 9 \textup{[kpc]}
* The height of the COV or line of sight distance reaching behind the GC when looking towards it.
* Cluster distance: 8 \textup{[kpc]}
* The distance between the observer and the cluster. When looking straight at the GC the distance between the cluster and the GC is 300 \textup{[pc]}
* The mean cluster velocity is set to the circular velocity at this location.
* View point: \begin{bmatrix}8300 & 0 & 27\end{bmatrix}^{T}_{GCA} \textup{[pc]}
* The position of the observer in GCA.
* Timestep: 28 \textup{[day]}
* Amount of time between the two recorded snapshots. The time per integration is 7 \textup{[day]}. Therefor snapshots are taken every 4 integrations.

**Cluster specific parameters (McLuster)**

* 1. Profile: 0 Plummer density profile
* (R) Radius: -1
* With this setting the radius is calculated by mcluster via a mass to half-mass radius relation as described in [[23](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id55)]
* (Q) Virial ratio: 0.5
* The cluster is in virial equilibrium.
  1. IMF: 1 Kroupa ranging from 0.08 Msun to 100 Msun
  2. Output: 3 The resulting mass, position and velocity for each star is written into a file.

### Results and Interpretation[](#results-and-interpretation)

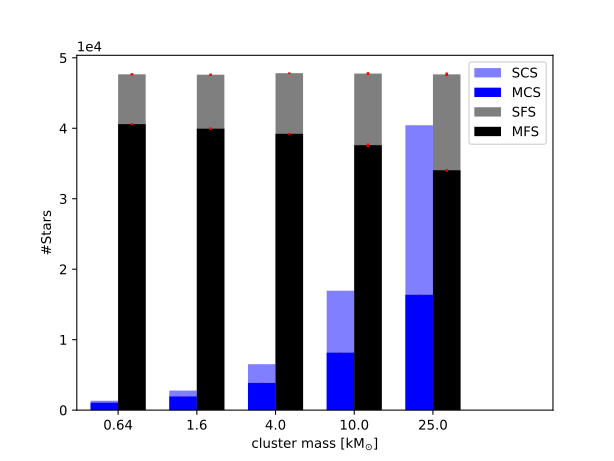


Fig. 12 Number of simulated/detected cluster and field stars at 10^\circ depending on the cluster masses[](#id3)

Depending on the total cluster mass the amount of simulated CS (SCS) ranges from 1.3e3 to 40.4e3 while the amount of mapped CS (MCS) lies between 1.0e3 and 16.4e3. The decrease in detectability of CS is due to increasing CS density and has no direct impact on the clustering performance. The difference between observed stars (OS) and mapped stars (MS) is negligible, in fact the amount of not mapped stars decreases with increasing number of SCS.

While the number of simulated FS (SFS) remains constant, the number of mapped FS (MFS) decreases with increasing number of CS because bright CS hide FS. This inverse relationship does impact the clustering performance to some degree, less MFS means less potential FPs as well as TNs, the latter are not relevant for Precision and F1 score.

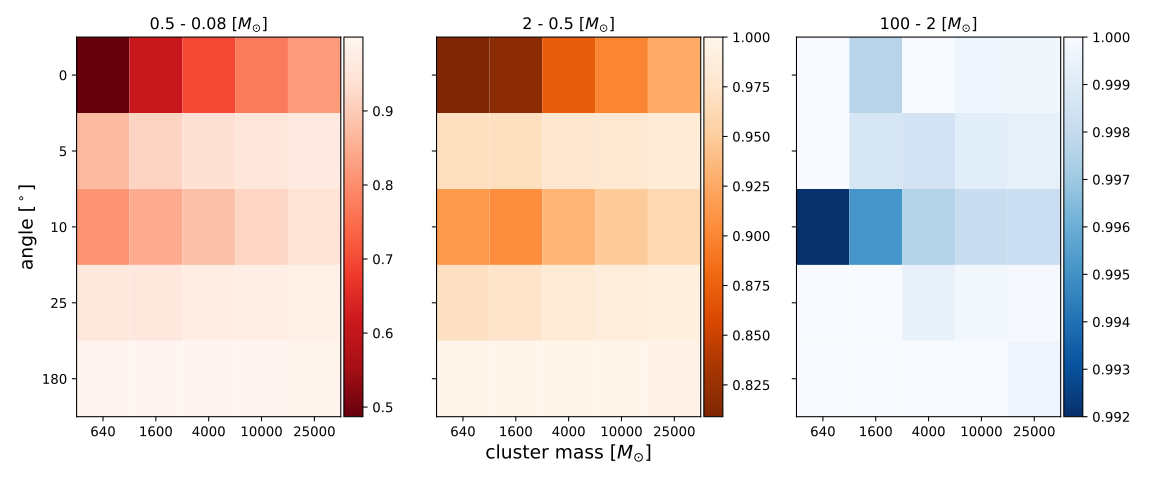


Fig. 13 Precision for different mass bins, angles and cluster masses[](#id4)

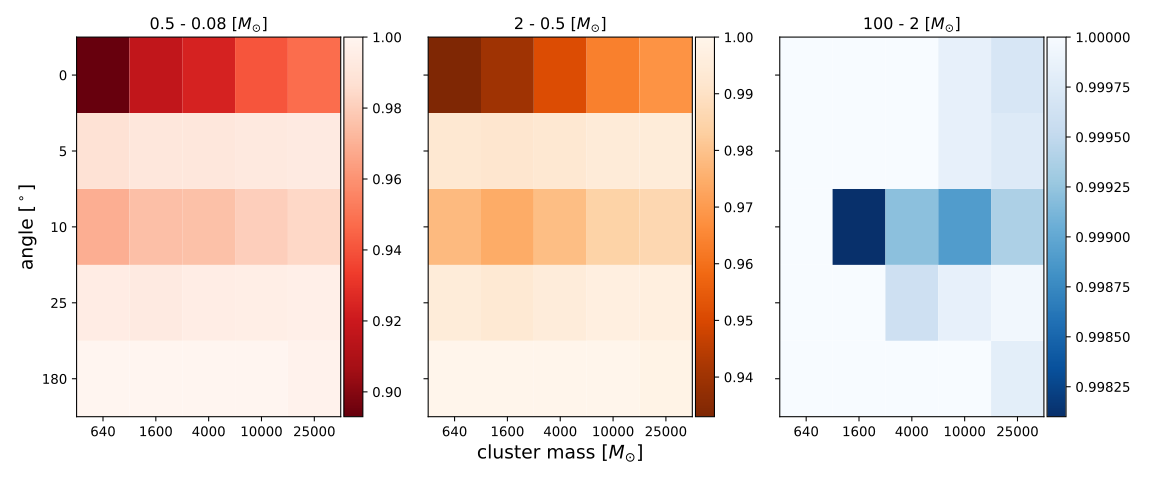


Fig. 14 Precision using accurate velocities[](#id5)

[Fig. 13](#fig-25-precision) displays the precision using the velocity of OS and [Fig. 14](#fig-25-precision-sim) for the velocity of SS. The only relevant difference between simulated and observed HTP proper motion is the accuracy of position and consequently of velocity. While both figures show the same relationships, the drop in overall performance due to inaccuracies introduced during observation and source detection are painfully apparent.

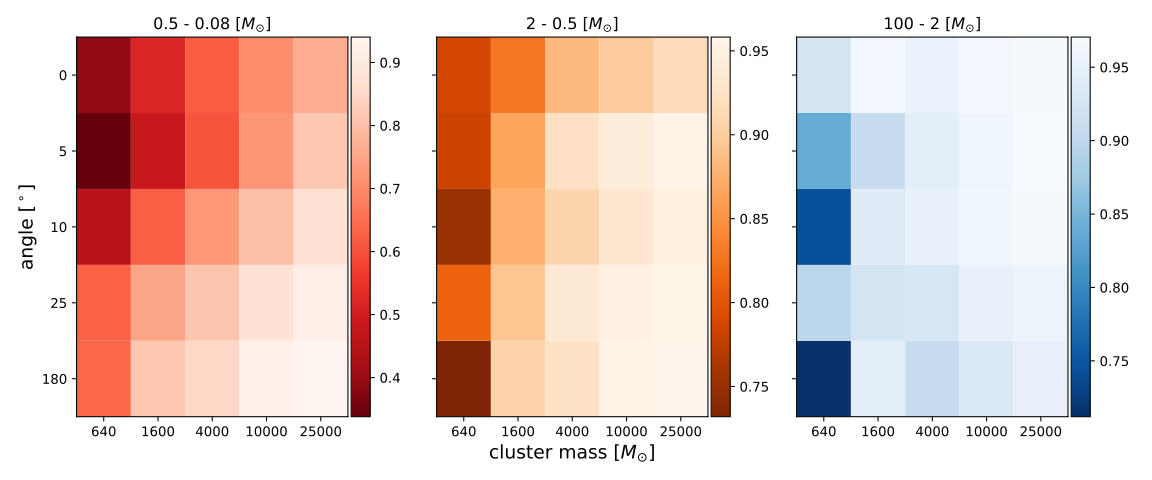


Fig. 15 F1 score for different mass bins, angles and cluster masses[](#id6)

As is discernable in [Fig. 13](#fig-25-precision) and [Fig. 14](#fig-25-precision-sim) the precision, with one exception, is correlated with the angle. Curiously for the same cluster mass the precision is lower at 10^\circ than at 5^\circ.

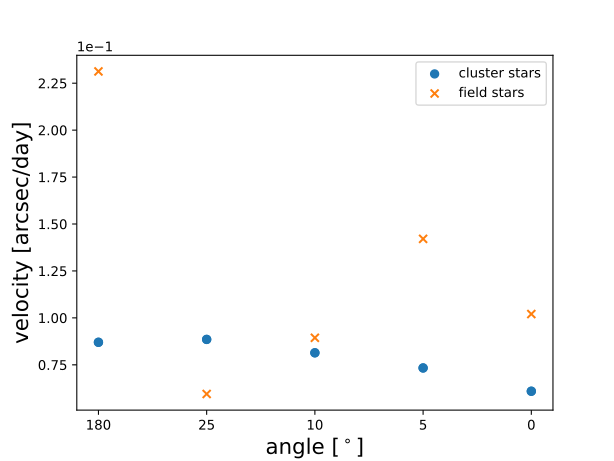


Fig. 16 Average cluster and field star velocity at different angles[](#id7)

[Fig. 16](#fig-25-avg-vel-640) provides the explanation for this outlier. At 10^\circ the average field star velocity is closer to the average cluster star velocity than at any other angle, making it harder to differentiate between cluster and noise.

The bigger the cluster mass, the higher the cluster star velocity density, which implies the second correlation - precision with cluster mass - presented in [Fig. 13](#fig-25-precision).

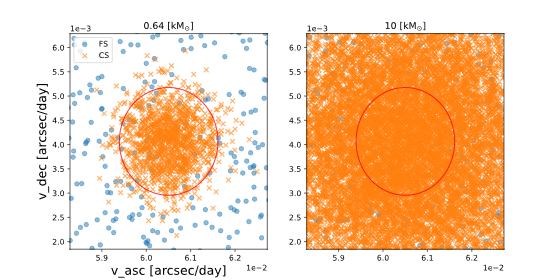


Fig. 17 2D HTP velocity of simulated clusters[](#id8)

[Fig. 17](#fig-25-vel-scatter) displays examples for the 2D HTP velocity space of two simulated clusters near the GC.

In this example the 0.64 [kM_{\odot}] cluster only has 1143 stars inside the circle while 10 [kM_{\odot}] has 3158. In both cases statistically the same amount of field stars fall within that area, leading to a higher ratio of FPs and therefore a lower precision for the lower mass cluster.

## Conclusions[](#conclusions)

The results presented, particularly [Fig. 13](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#fig-25-precision), indicate that data gathered with the ELT will indeed be accurate enough to study the IMF of YMCs near the GC. However, it remains questionable if this is still the case for stars with mass < 0.08 [\textup{M}_\odot]. Information about this mass regime, such as spectral type, MW disc and core IMFs is sparse. Moreover, the precision obtained here in the range 0.5 - 0.08 [\textup{M}_{\odot}] is already quite low. That is not to say it will not be possible. Many improvements and additions could be made to the presented work such as

1. A more accurate bulge model, for instance [[36](file:///C:\Users\kemu\Desktop\Texte\Alarich\private_beta\index.html#id57)]
2. Additional snapshots to recover temporarily hidden stars and get a more accurate velocity estimation.
3. A sophisticated method for extracting positional data like point spread function fitting.
4. Accurately determine the ideal time step size.

## Appendix[](#appendix)

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