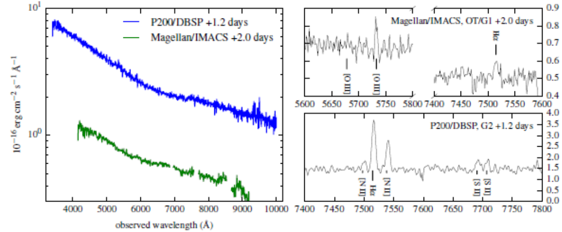

Figure 5. Optical spectra of iPTF13bxl and the nearby galaxy SDSS J142914.57+154619.3 ("G2"). Spectra in the left panel have been smoothed with a Savitzky-Golay filter. Our initial P200 spectrum of the afterglow (left panel, blue) exhibits a largely featureless blue continuum. A higher S/N spectrum taken the following night with IMACS (left panel, green) revealed faint emission features corresponding to [O III] and Hα at $z = 0.145$ (top right panel). The bottom right panel shows a spectrum of the nearby galaxy G2, which has the same redshift as iPTF13bxl.
(A color version of this figure is available in the online journal.)

clear whether this energy difference is due primarily to the release of less relativistic ejecta by the burst overall, a wider jet, or a partially off-axis view of a structured jet. Late-time radio follow-up should help distinguish these models: an intrinsically low-energy GRB should produce a much earlier jet break than a widely-beamed burst, while a structured jet will actually produce an *increase* in flux at late times as the jet core spreads and its radiation enters our sightline.

Events with similar energetics have been found by *Swift*, e.g., GRB 050826 at $z = 0.30$ and GRB 120422A at $z = 0.28$ (Mirabal et al. 2007; Zhang et al. 2012). However, given their low intrinsic luminosities and higher redshift, the afterglows were too faint to identify late-time breaks and establish their shock energies E_K, making them difficult to physically interpret. GRB 130702A's proximity avoids both these problems. Our observations suggest—and further observations should confirm—that its γ-ray and afterglow energetics are intermediate between these two previously quite-disparate classes of GRBs, helping to fill in the "gap" between the well-studied cosmological population and the class of less-luminous local GRBs and relativistic Type Ic supernovae (e.g., Soderberg et al. 2004, 2010).

similarly coarse position reconstruction. Later this decade, a network of advanced gravitational wave (GW) detectors including the Laser Interferometer GW Observatory (LIGO) and Virgo is expected to detect ~ 0.4–400 binary neutron star mergers per year (Abadie et al. 2010), but with positions uncertain to tens to hundreds of deg^2 (Fairhurst 2011; Nissanke et al. 2011; Aasi et al. 2013).

Optical counterparts to GW sources will rarely (due to jet collimation) include bright, on-axis short-hard burst afterglows. Fainter r-process–fueled kilonovae (Li & Paczyński 1998) or yet fainter off-axis afterglows (Rhoads 1997) are expected to accompany binary neutron star mergers. Both of these signatures are predicted to be several magnitudes fainter than iPTF13bxl. Optical searches will be inundated with astrophysical false positives (Nissanke et al. 2013). This problem will only be exacerbated for future surveys covering larger areas (e.g., Zwicky Transient Facility; Kulkarni 2012) and/or with larger apertures (e.g., Large Synoptic Survey Telescope; Tyson 2002). However, a breathtakingly complete astrophysical picture could reward us: masses and spins measured in GWs; host galaxy and disruption ejecta in optical; cir-