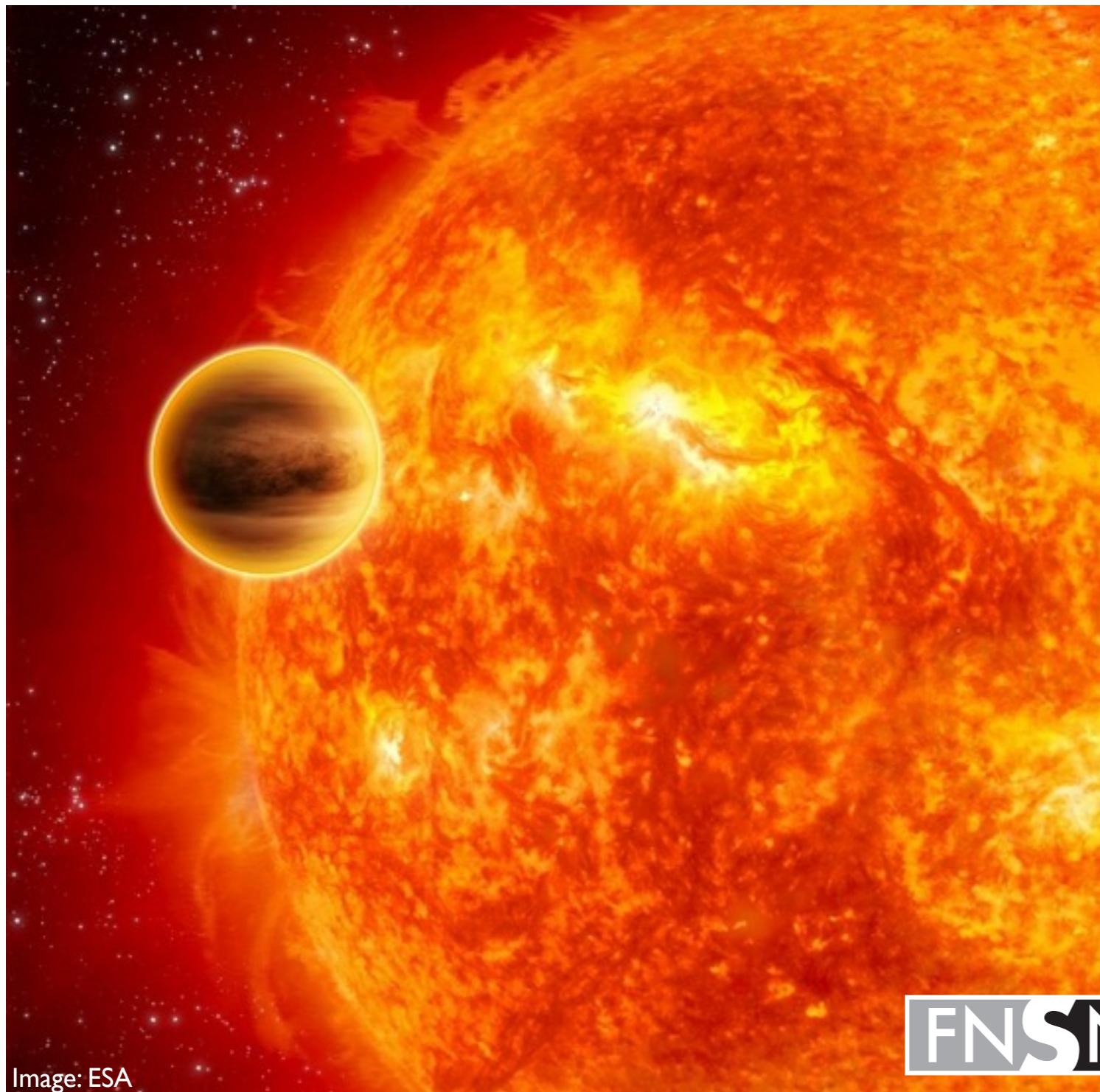


Probing the surfaces of Sun-like stars using transiting planets and 3D MHD simulations

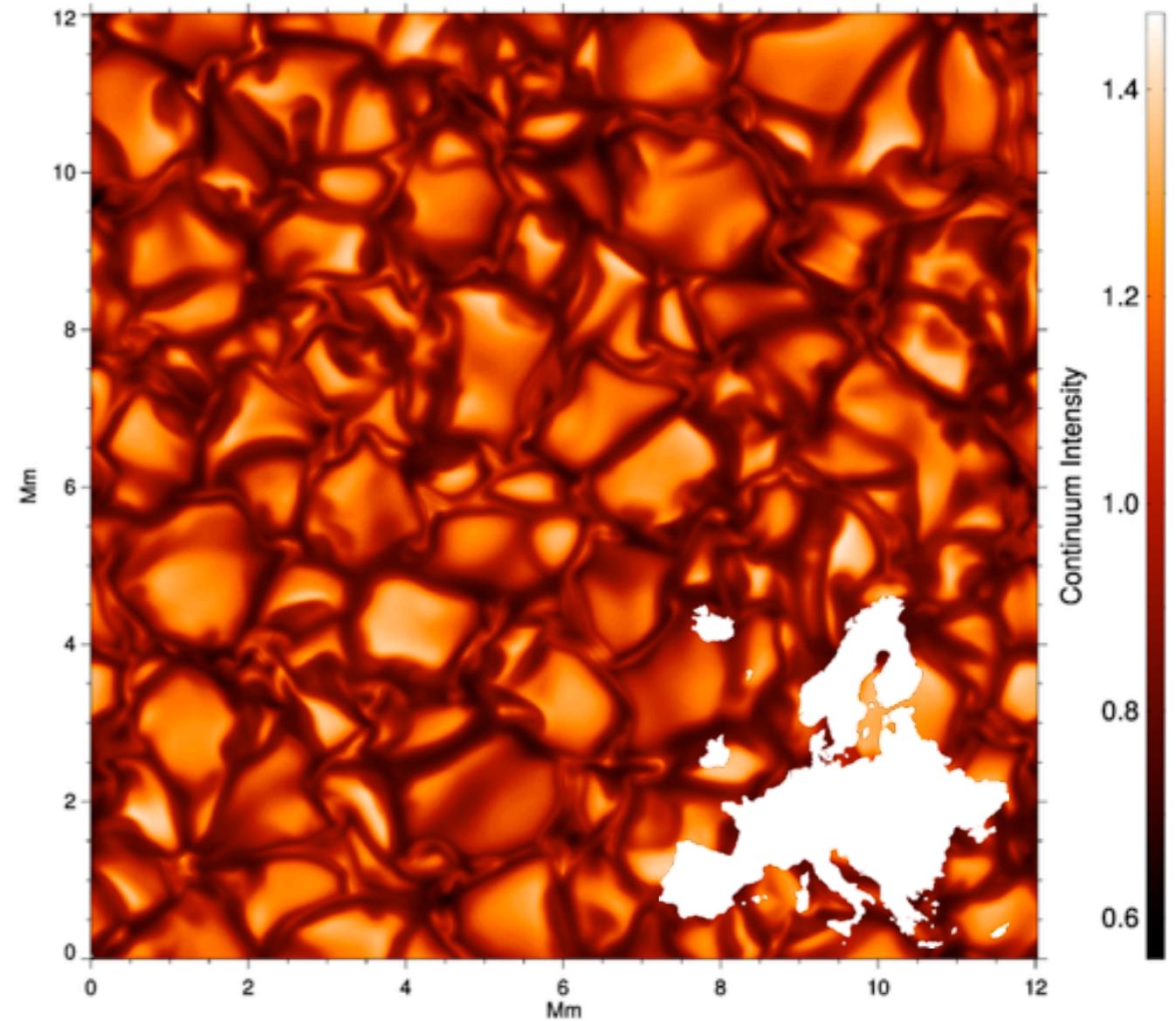
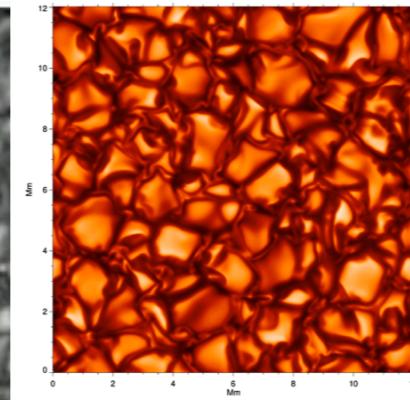
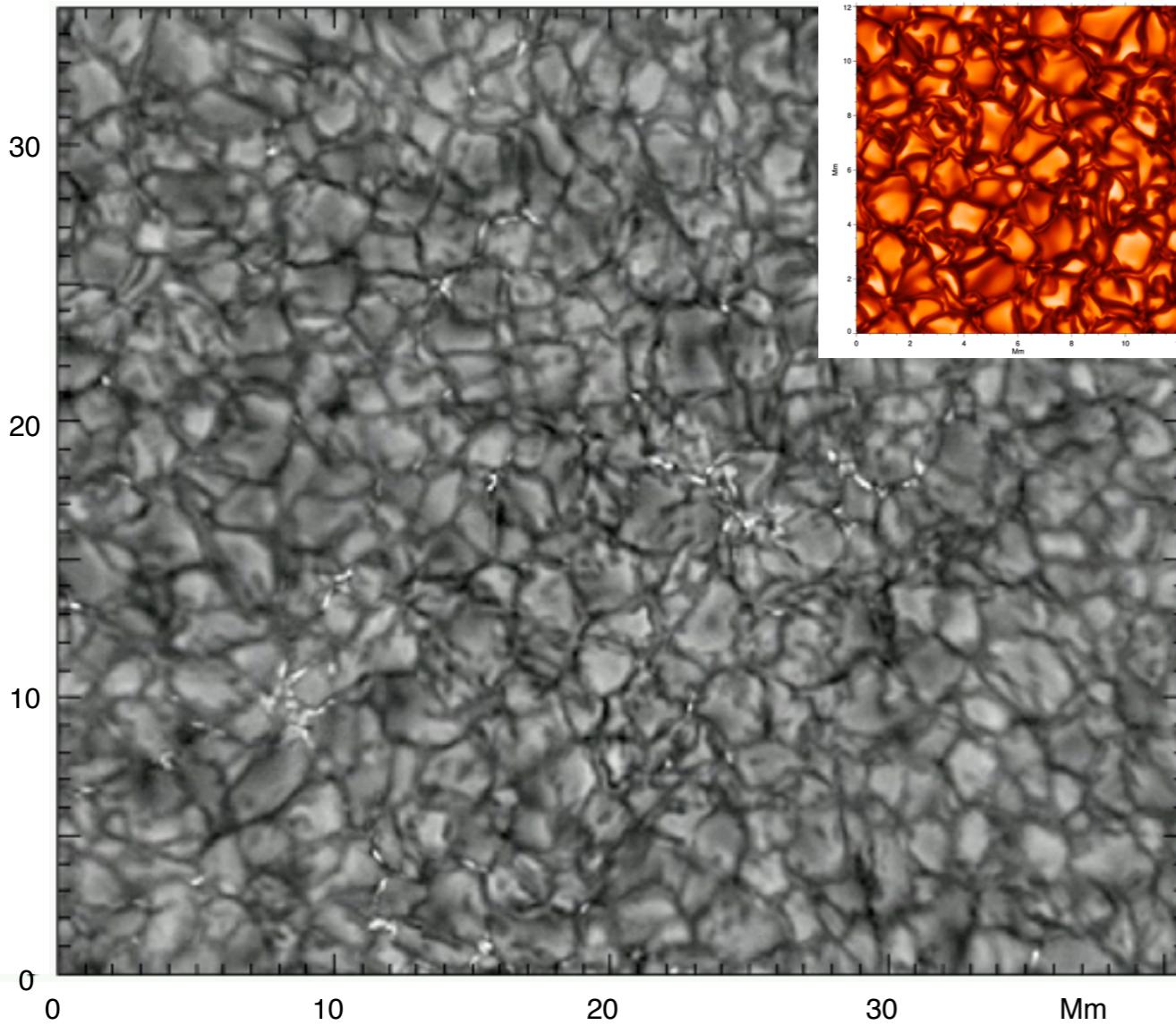


Dr Heather Cegla
CHEOPS Fellow

Image: ESA



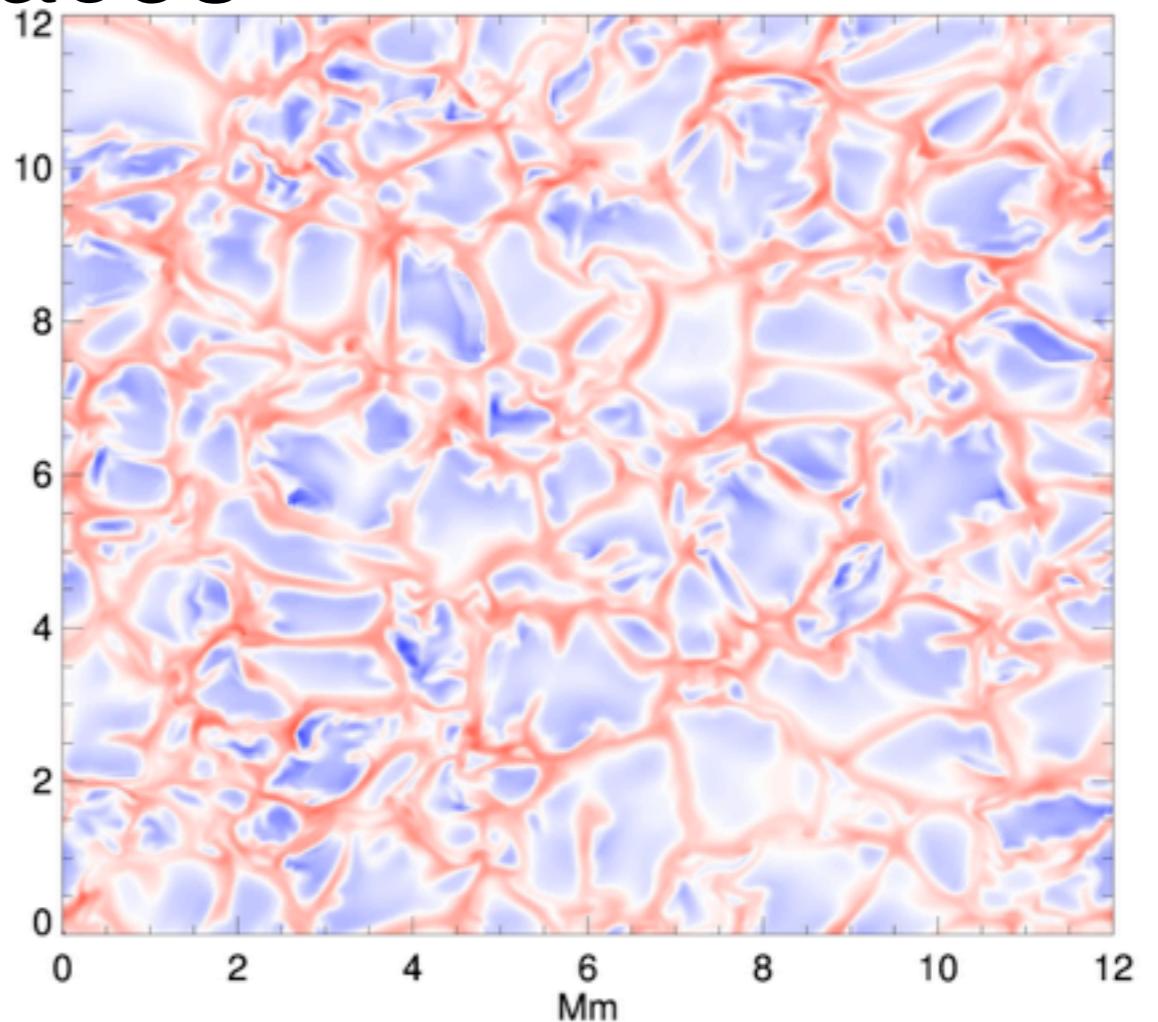
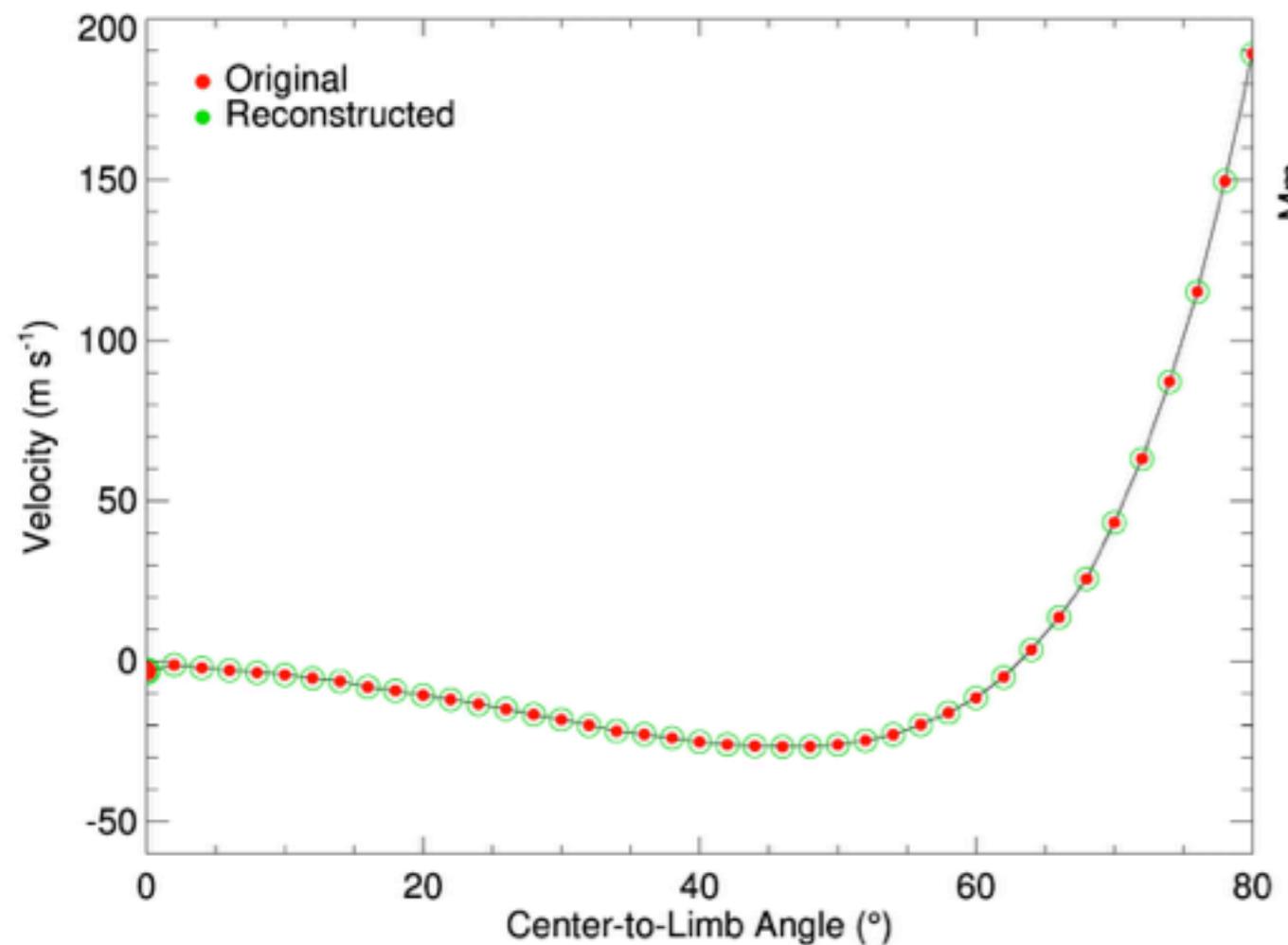
The solar surface





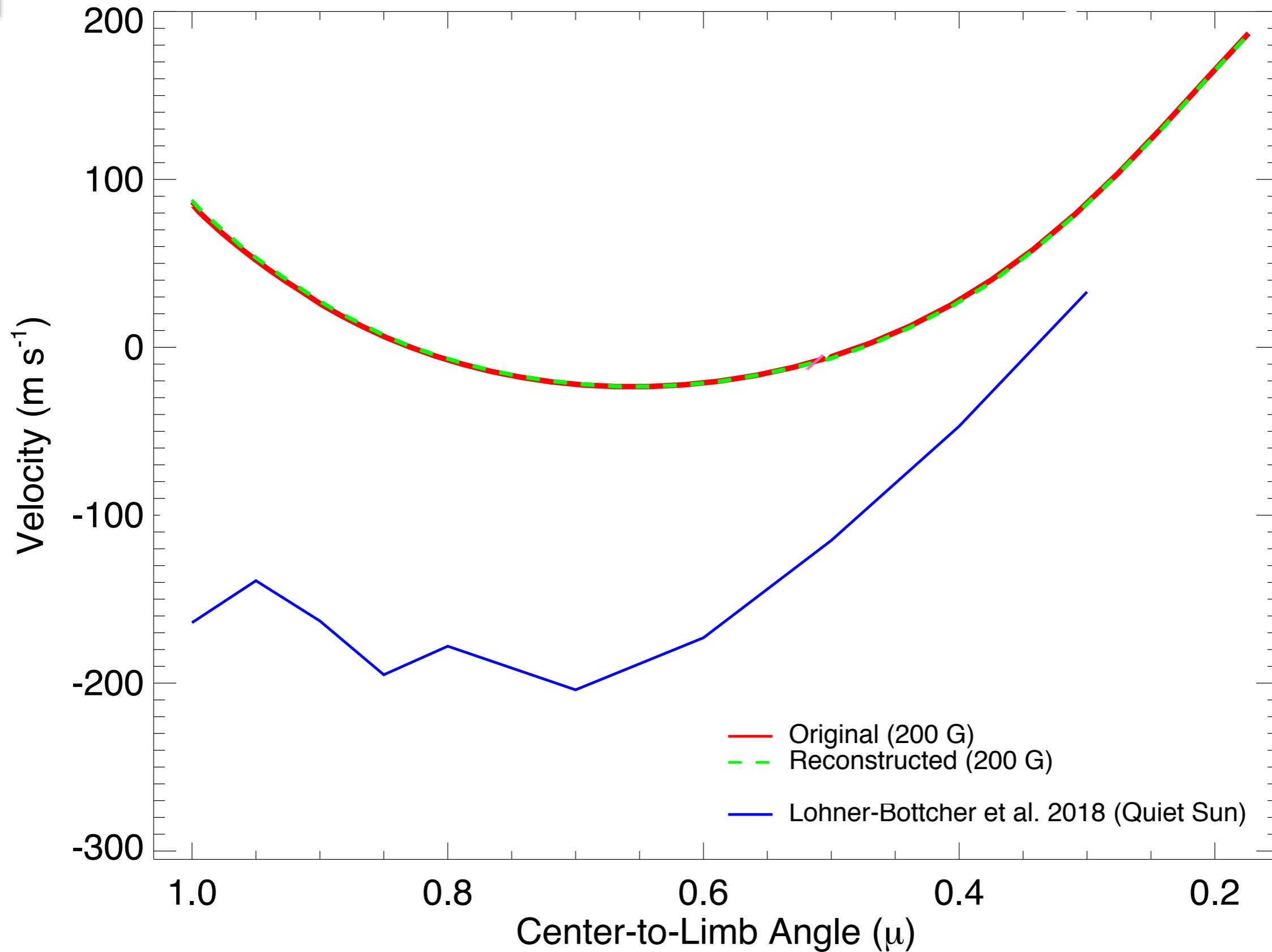
Transiting Stellar Surfaces

Line-of-sight Velocity (0°)



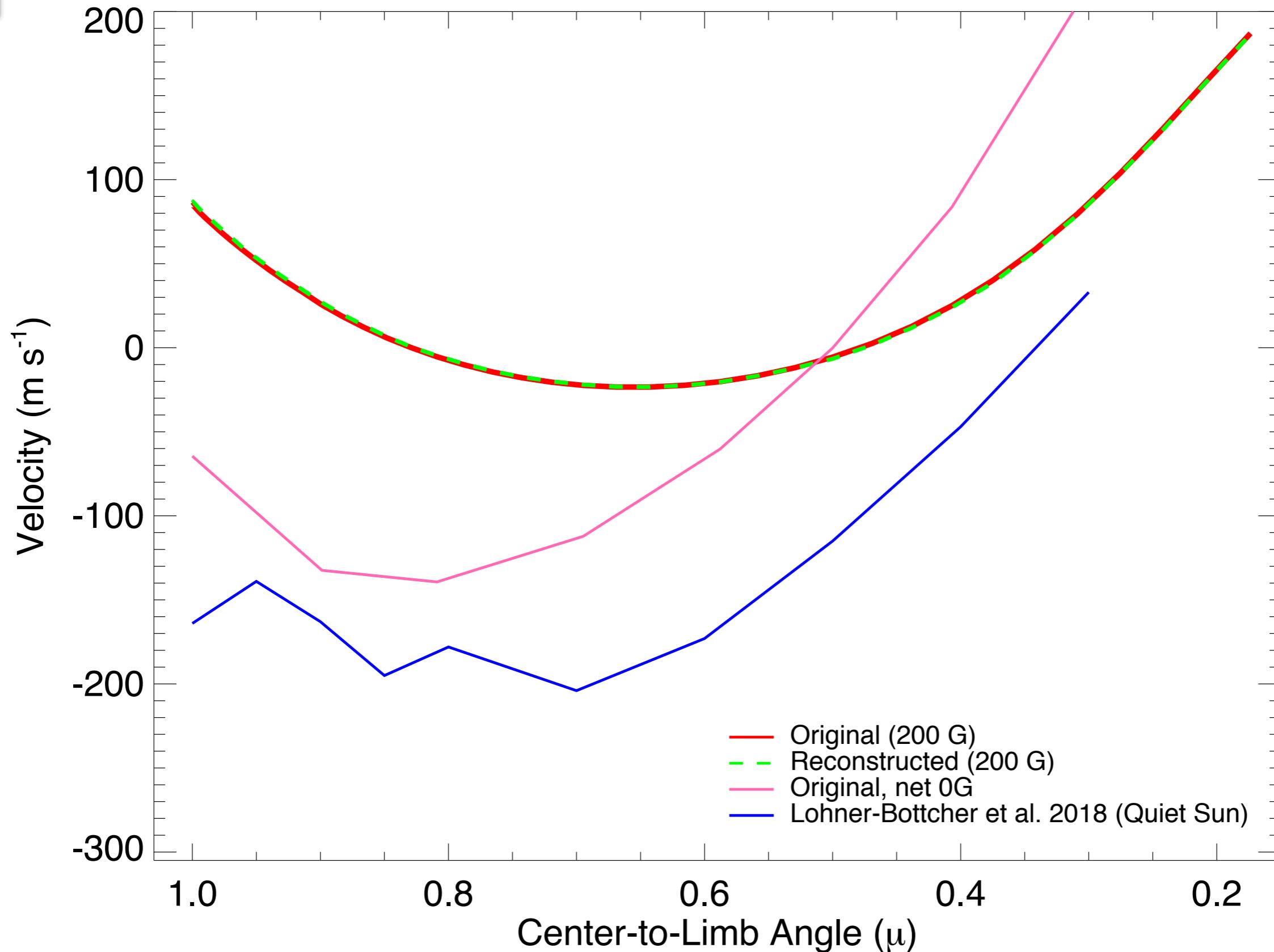


Centre-to-limb Variations on the Sun



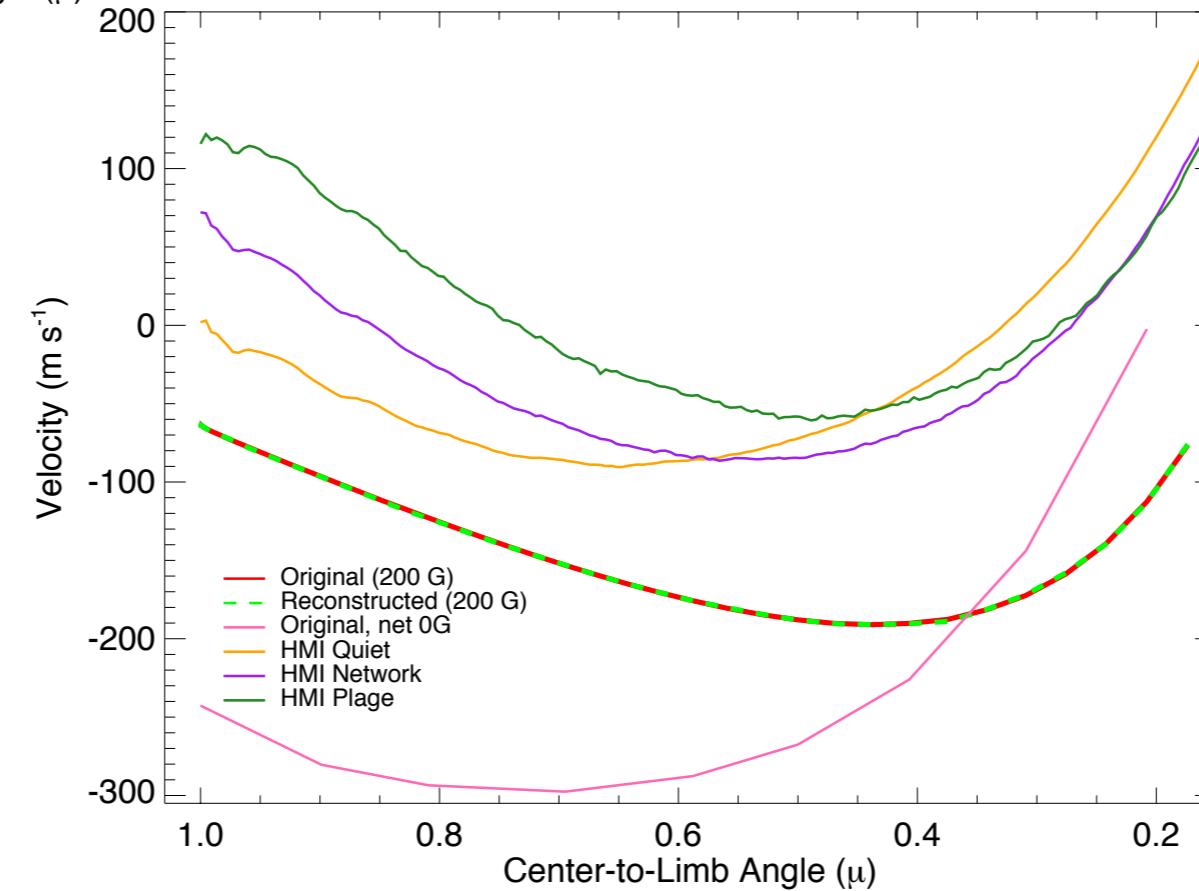
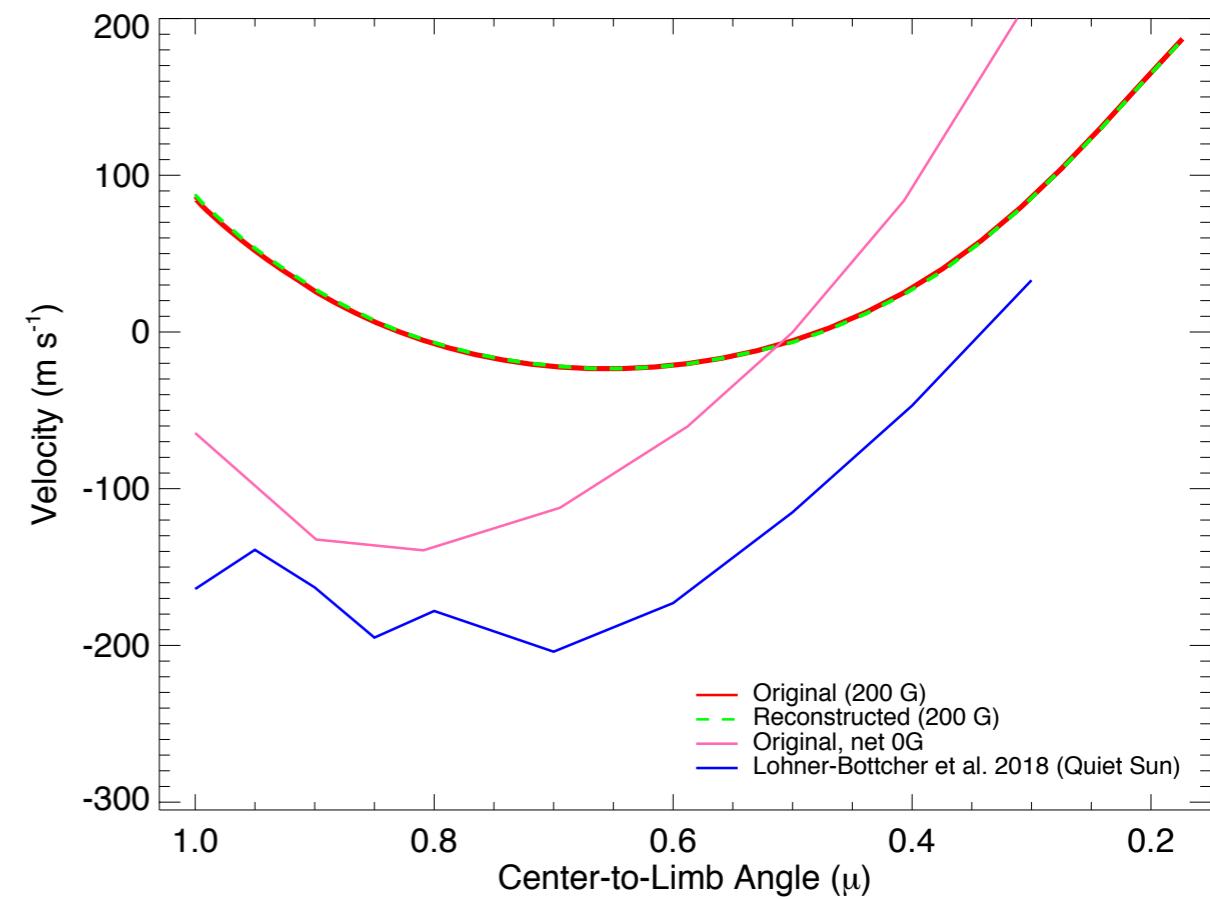
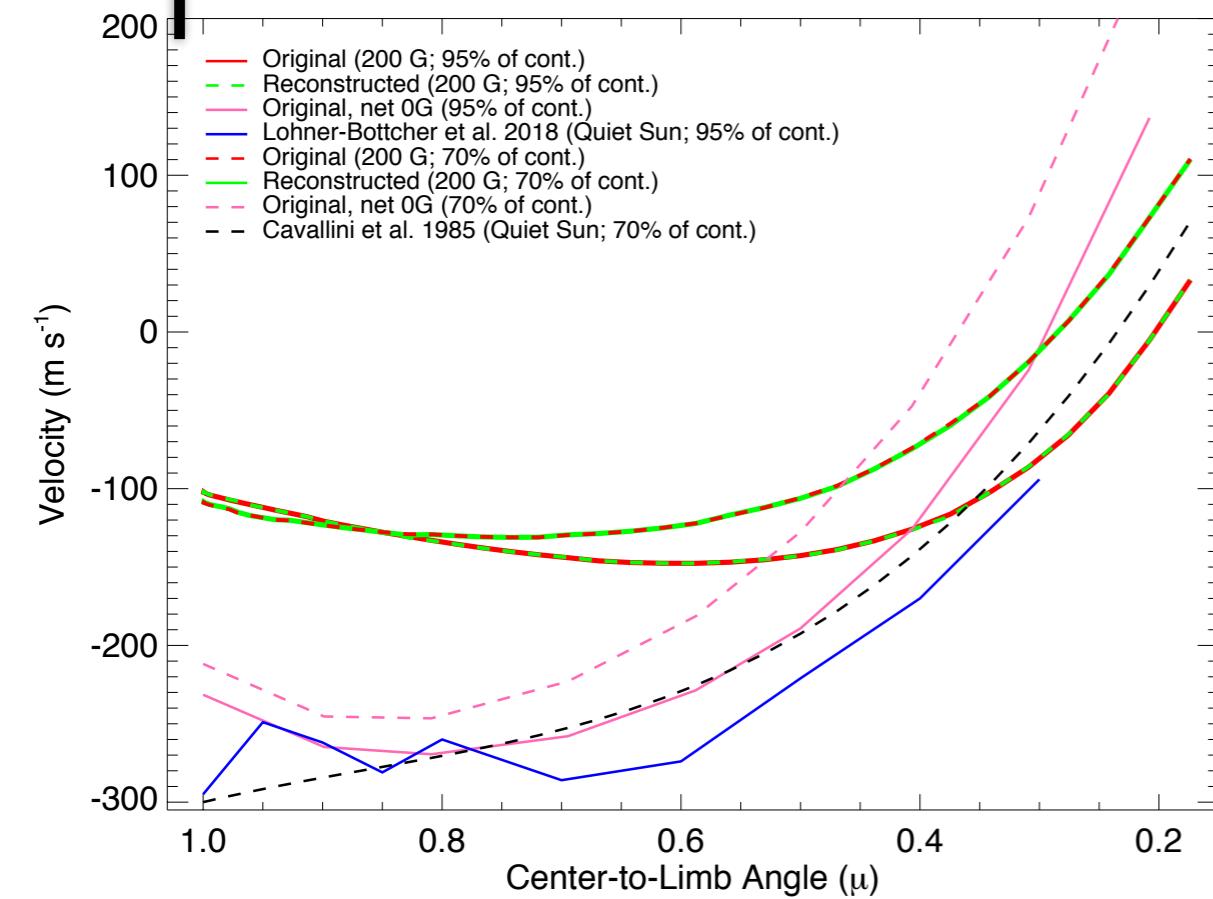


Centre-to-limb Variations on the Sun



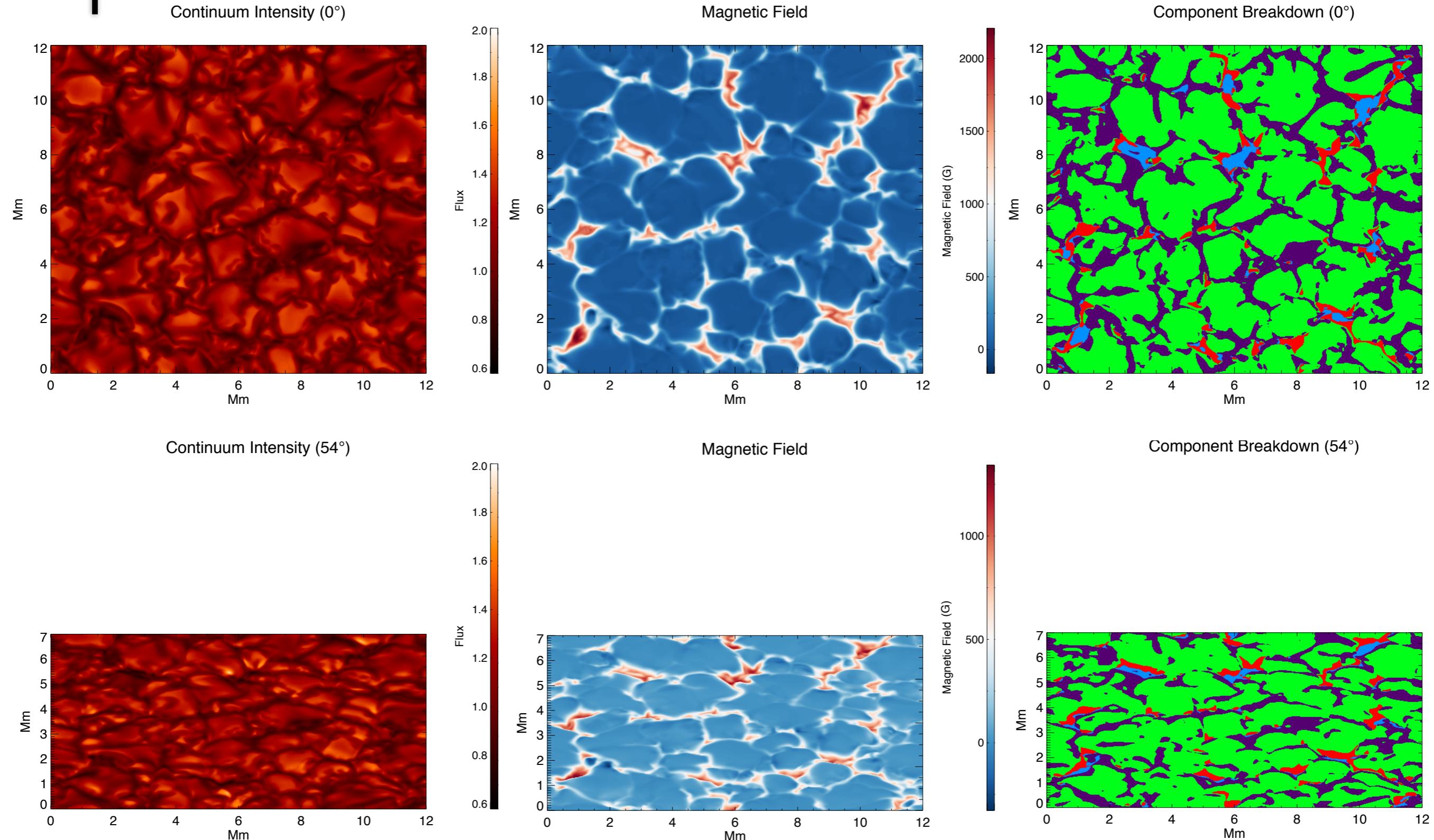


Centre-to-limb Variations on the Sun

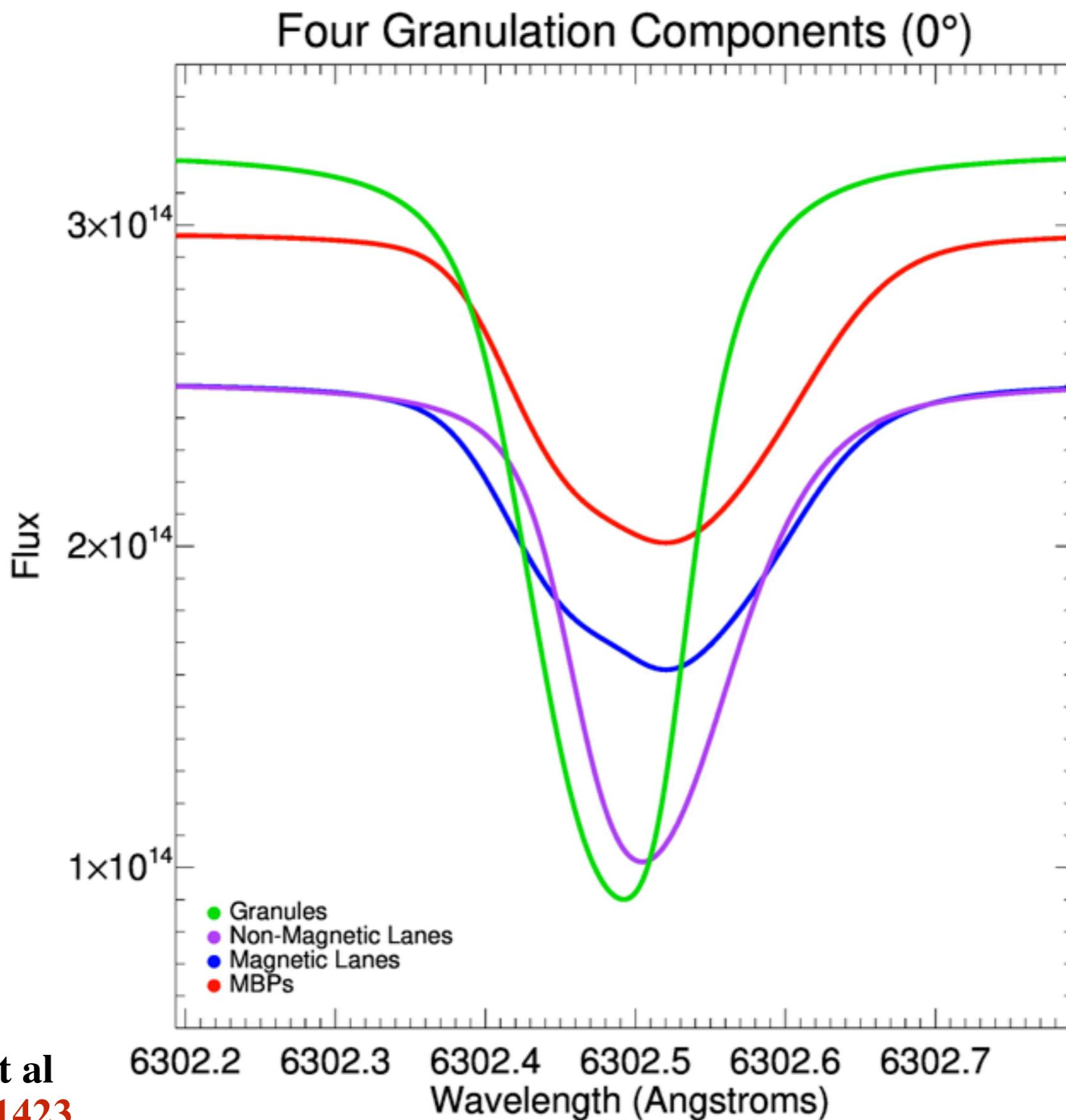


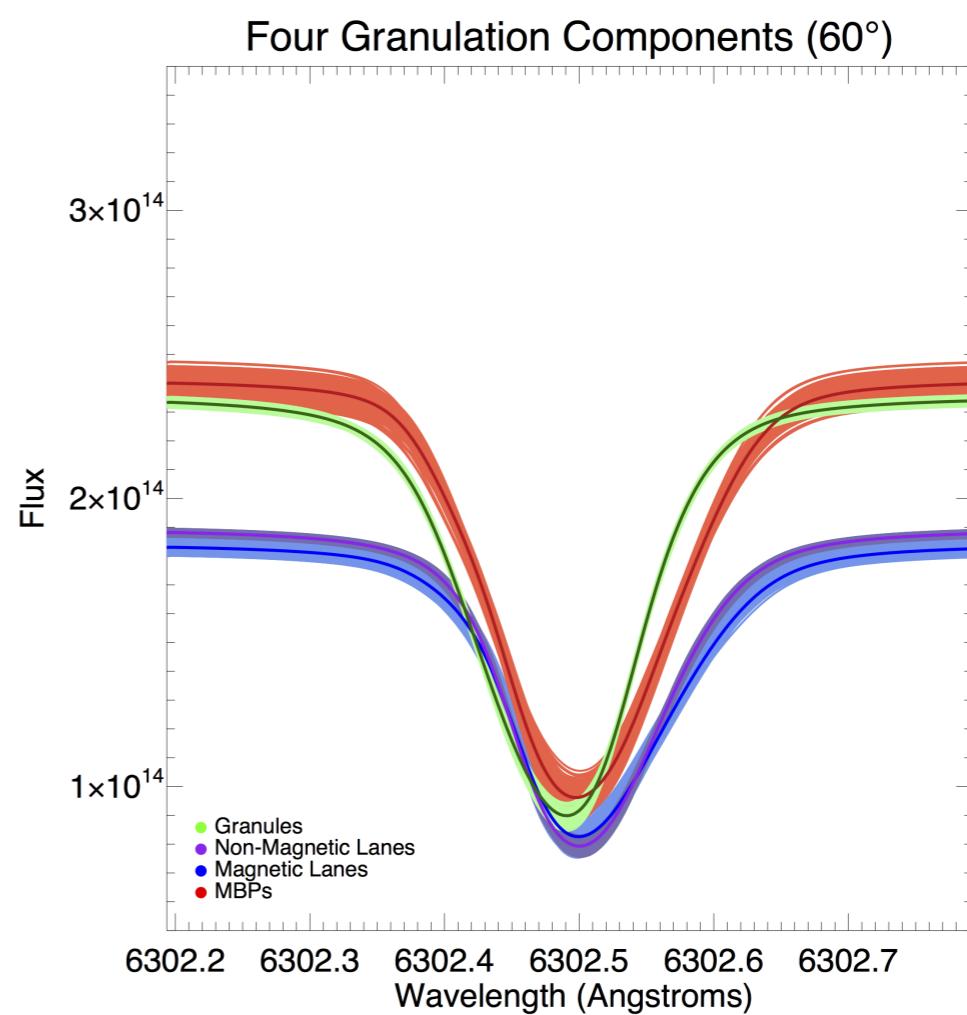
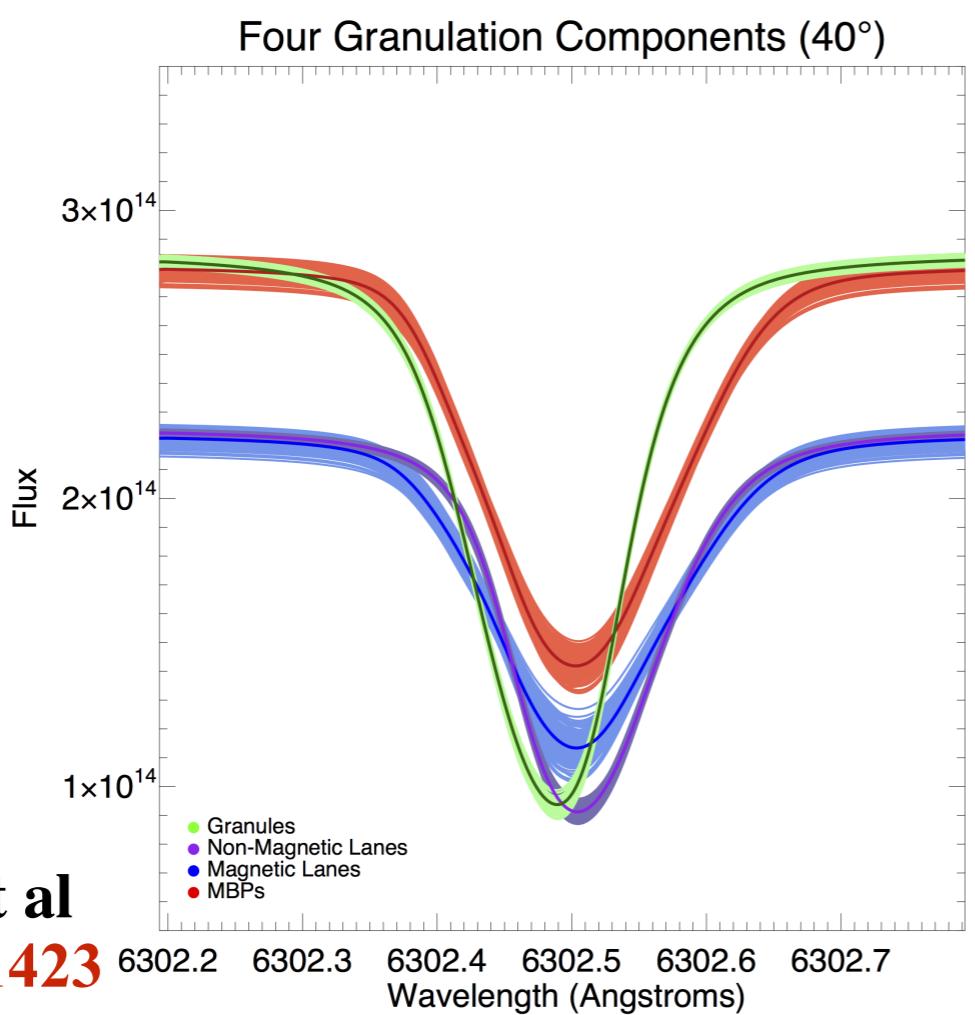
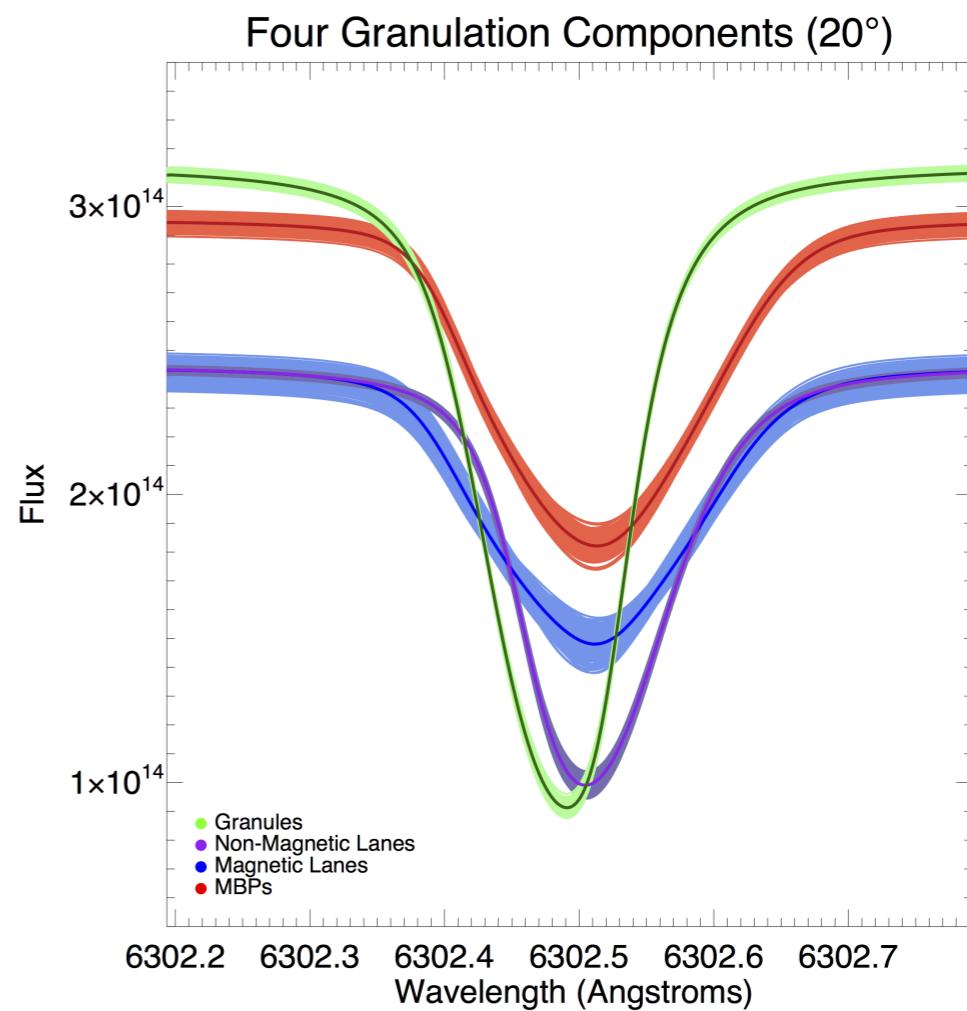
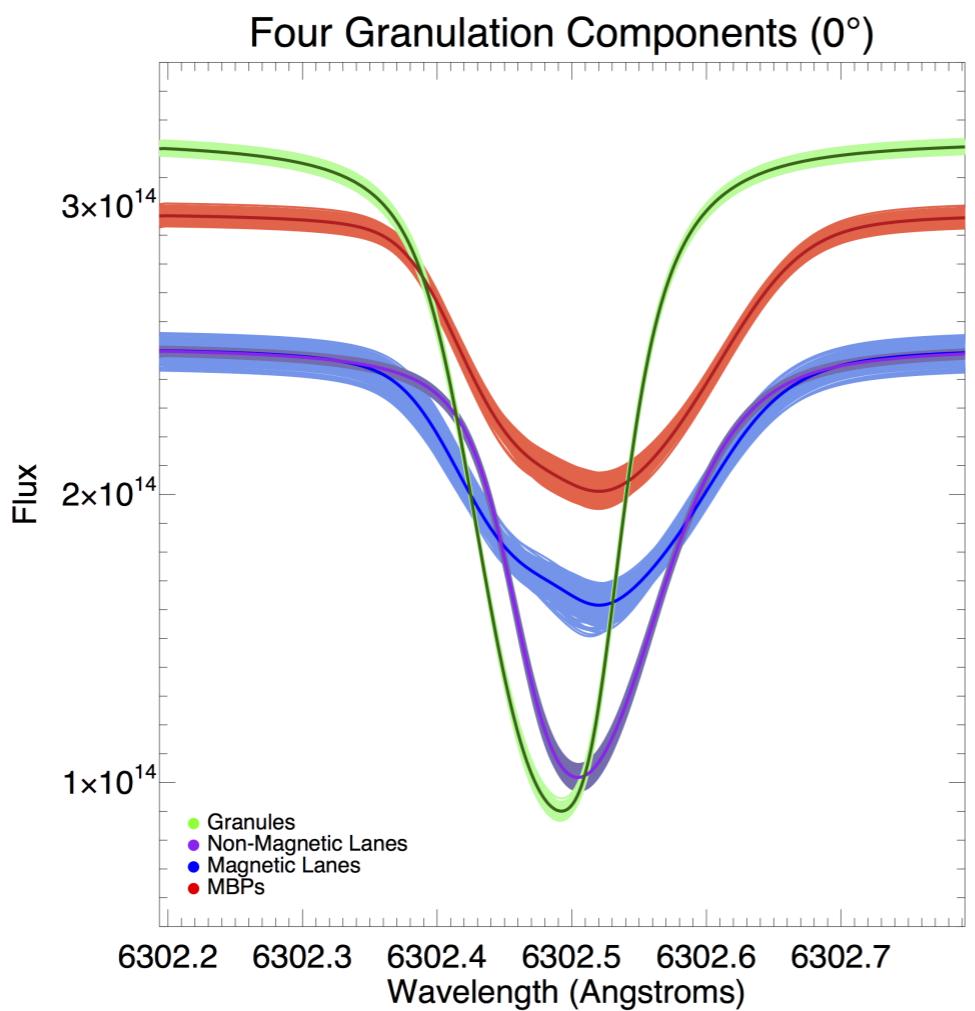


Parameterising the Granulation



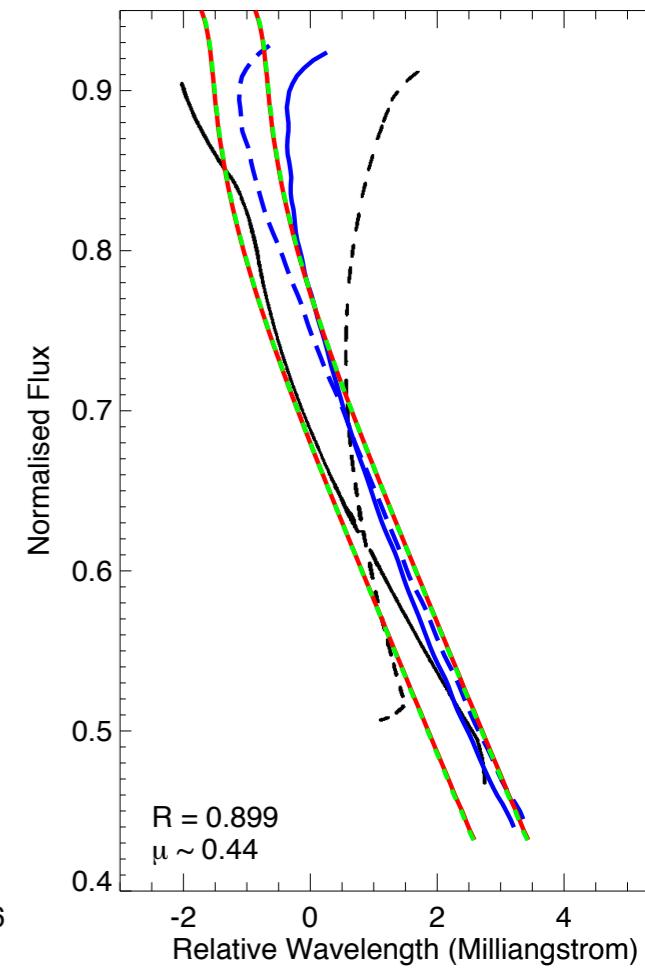
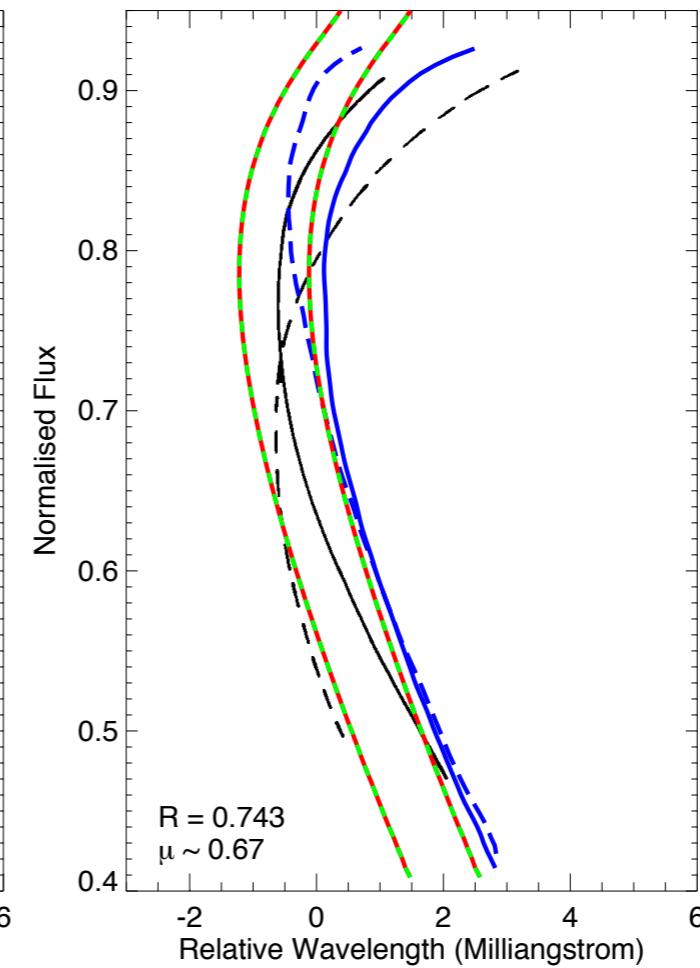
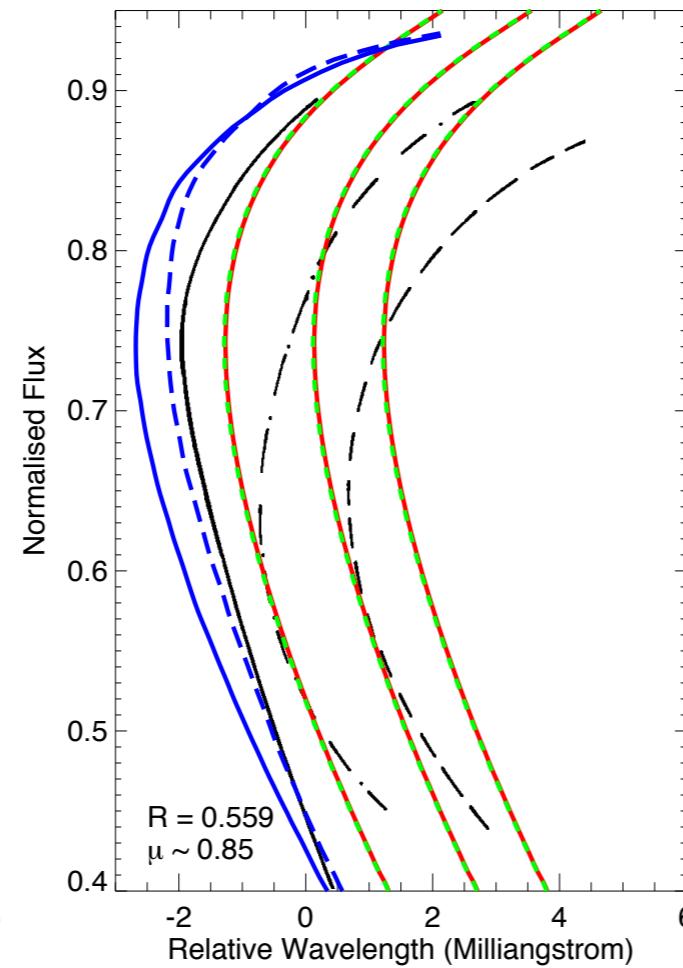
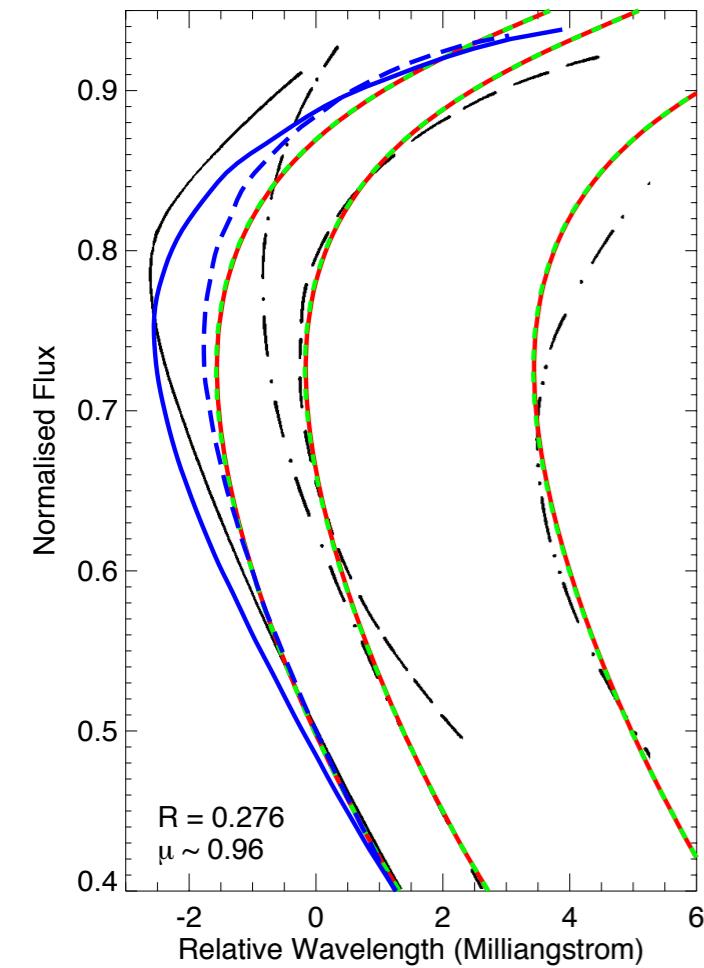
Centre-to-limb Variations on the Sun





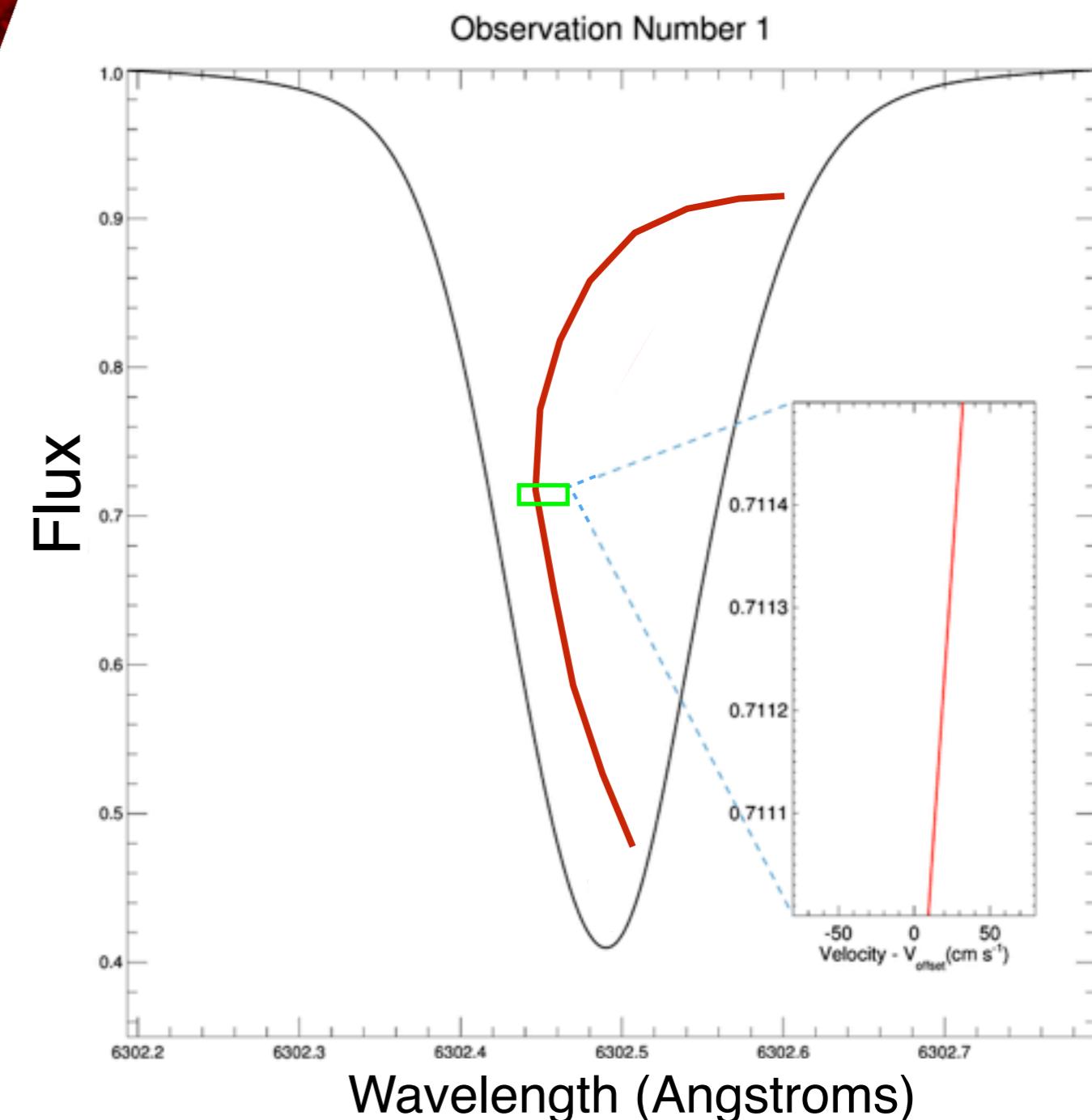
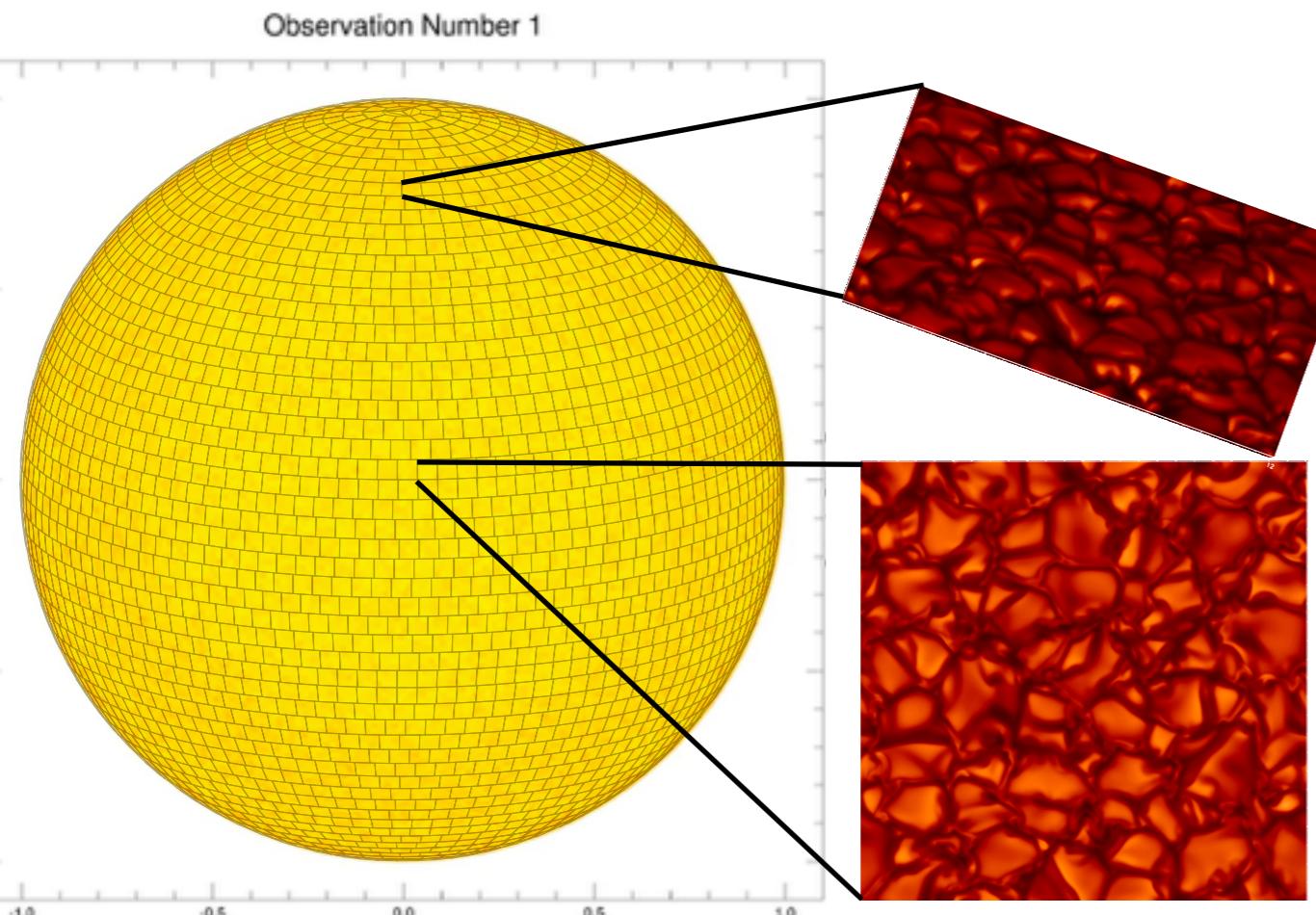


Centre-to-limb Variations on the Sun



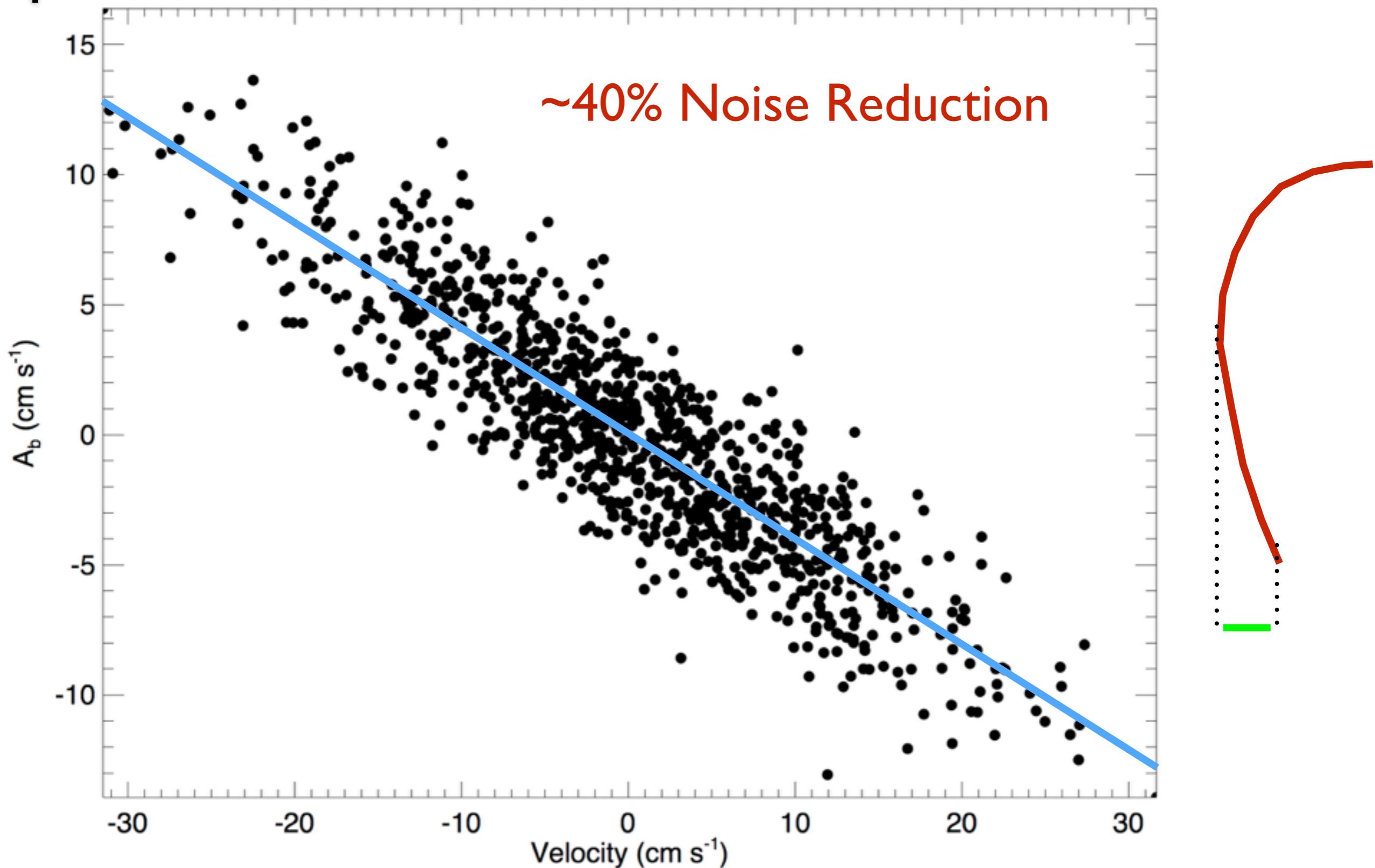


Sun-as-a-star Model Observations



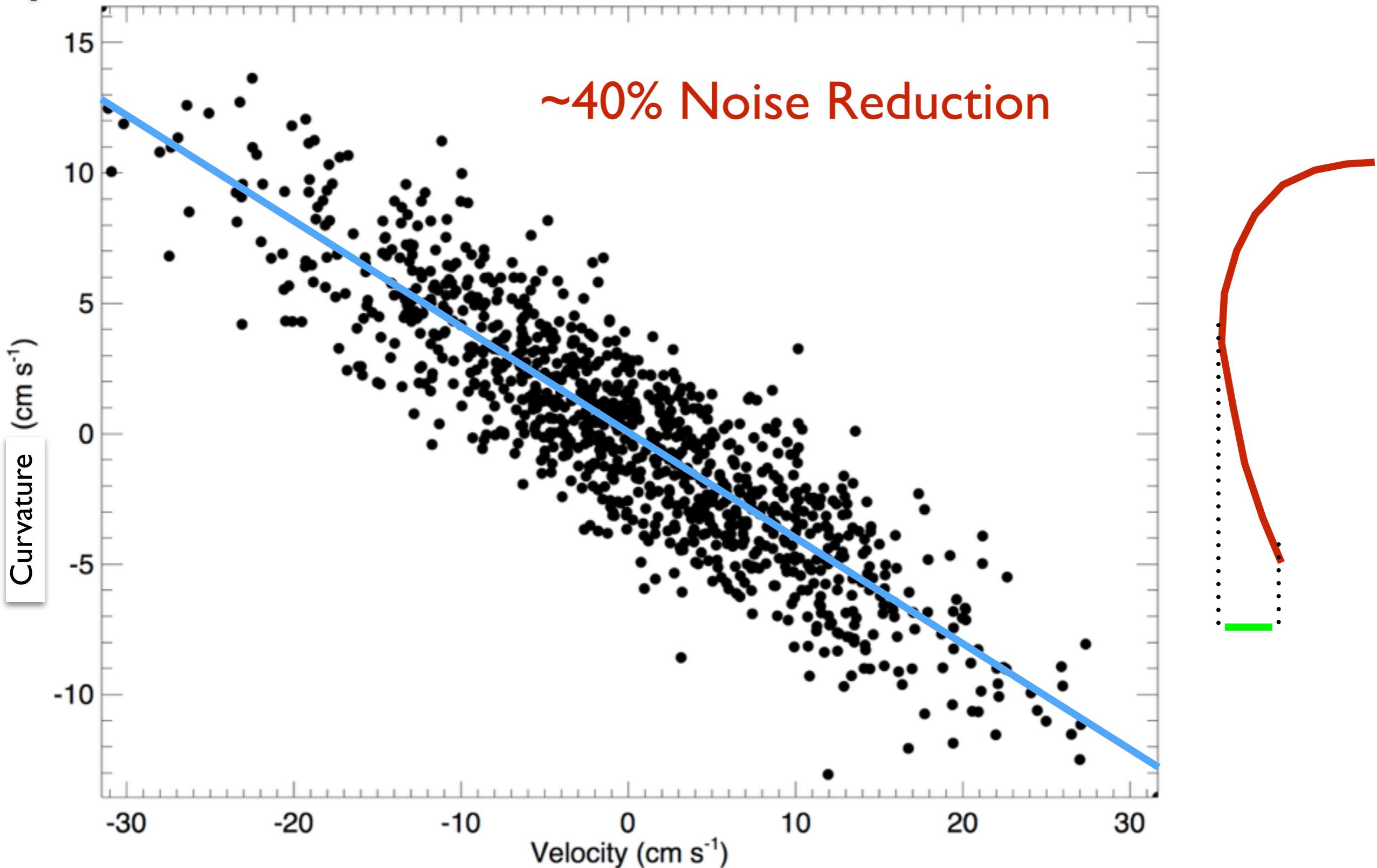


Correlations



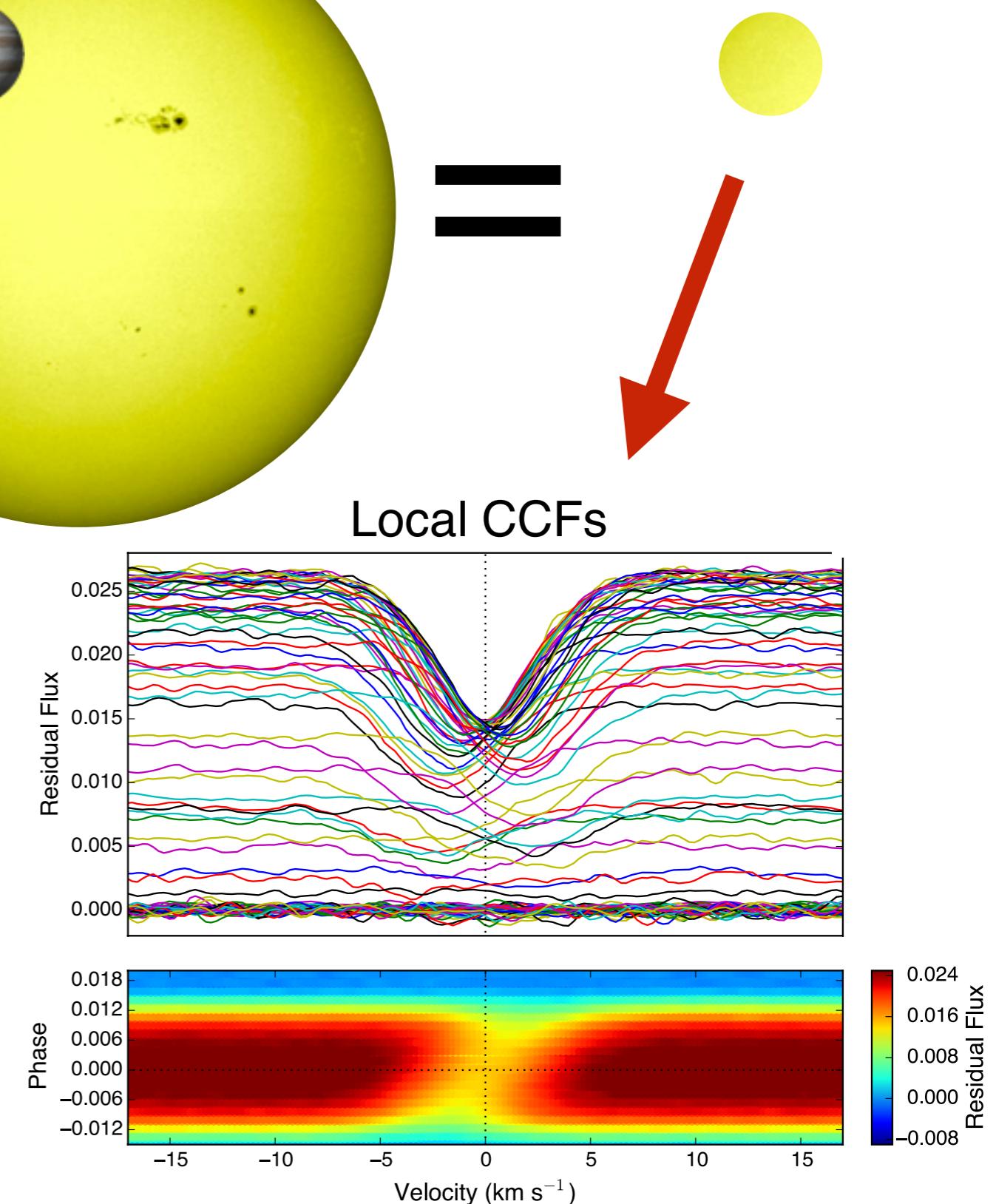
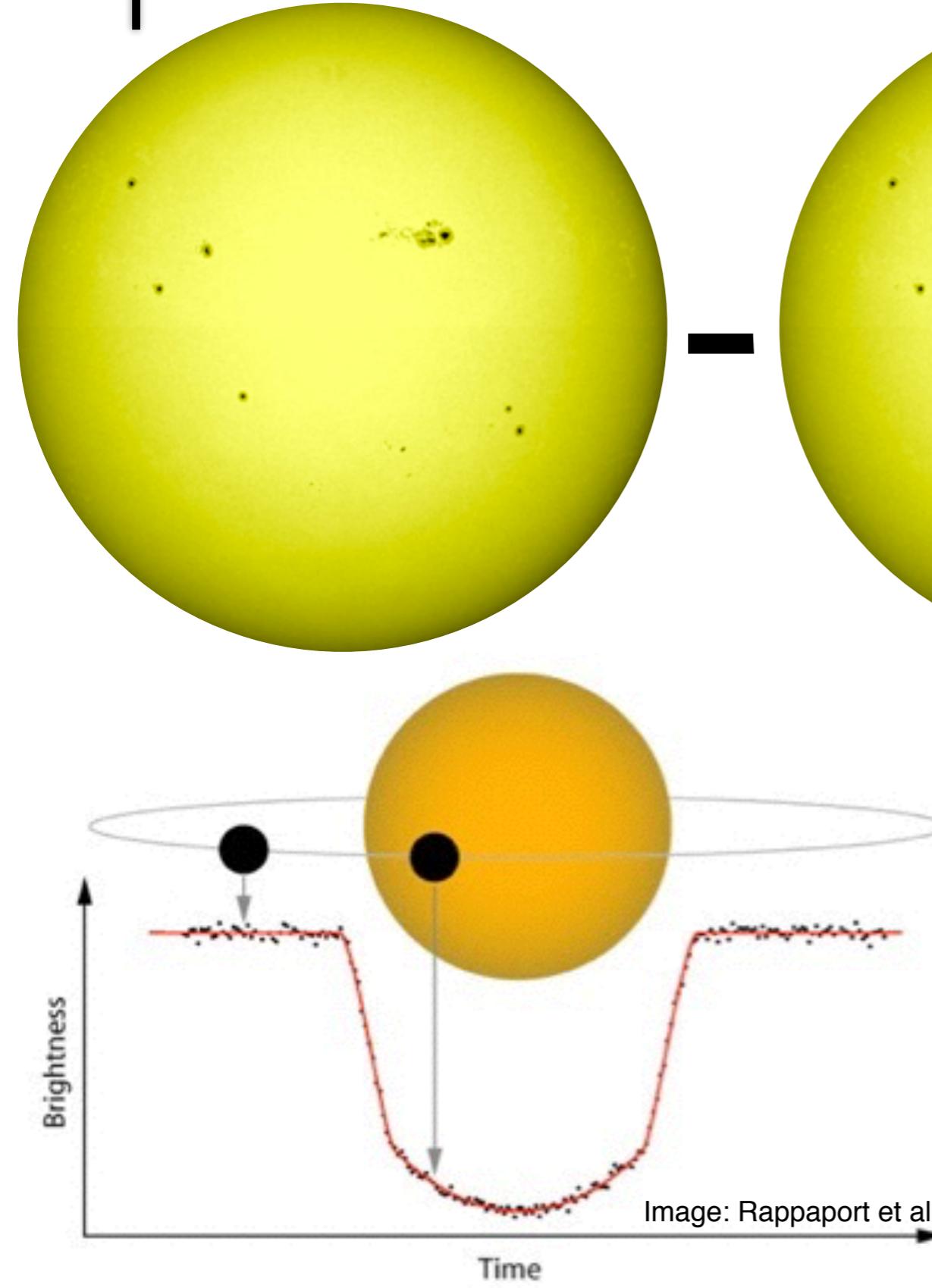


Correlations

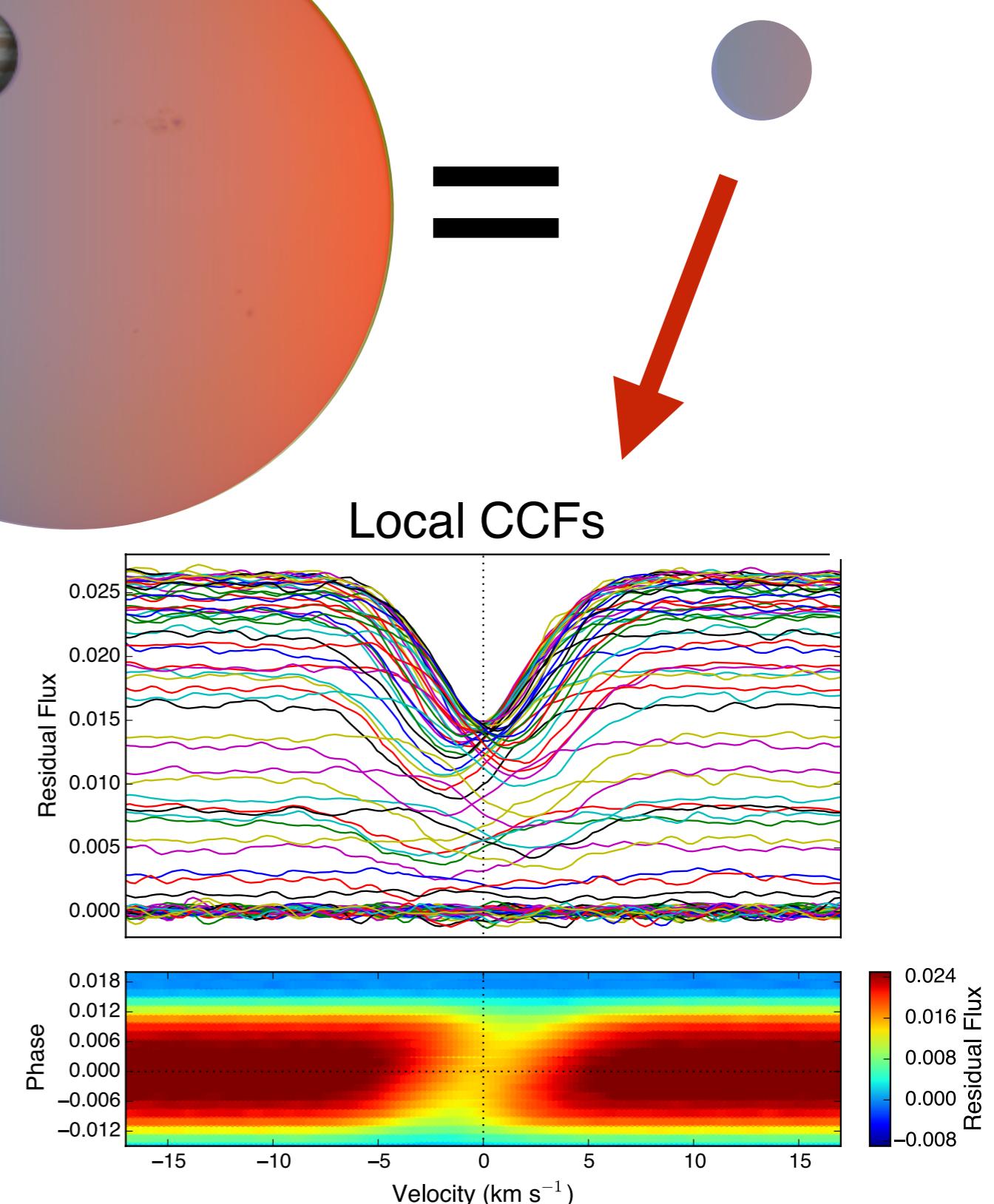
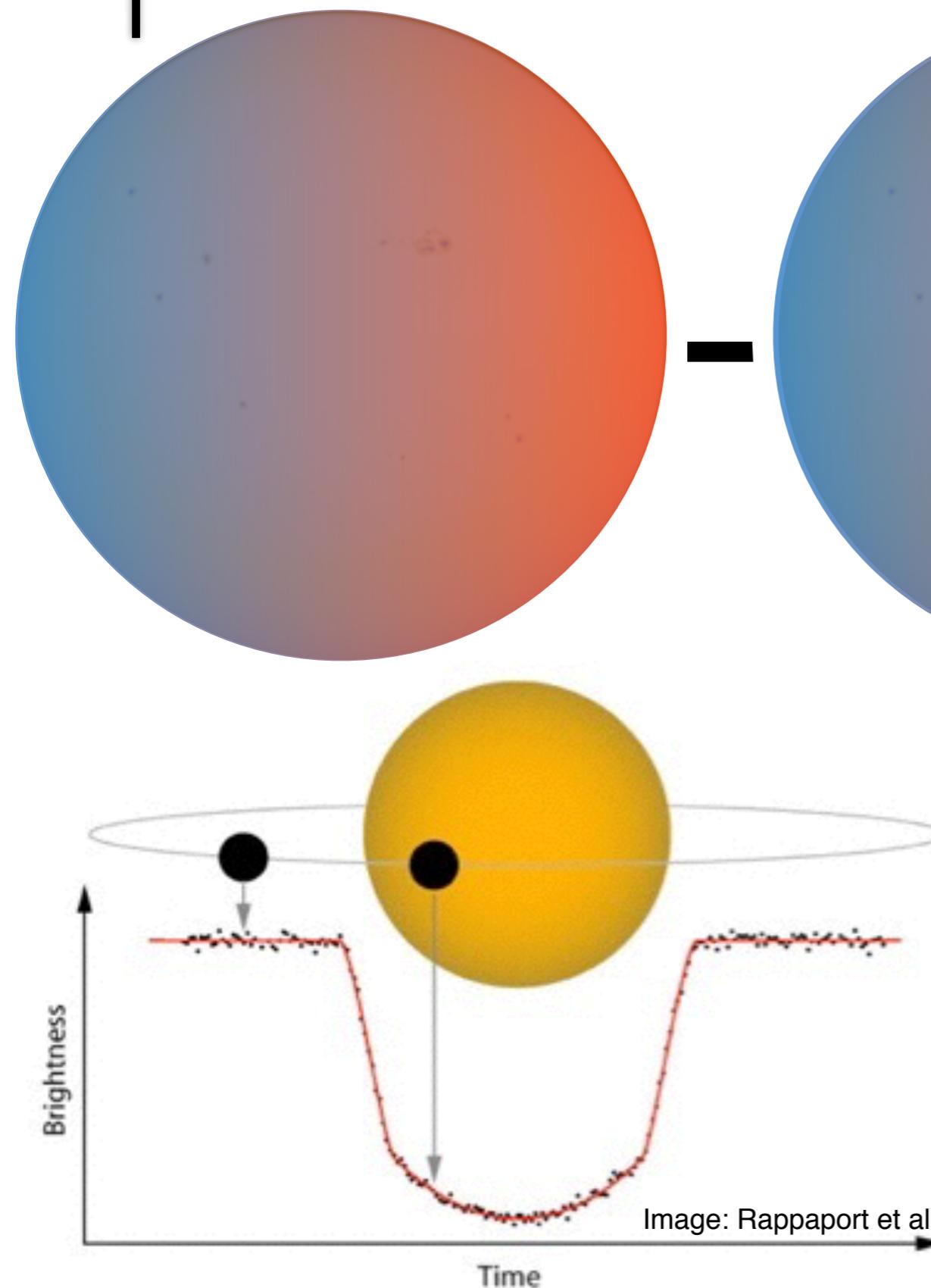




Planets as Probes

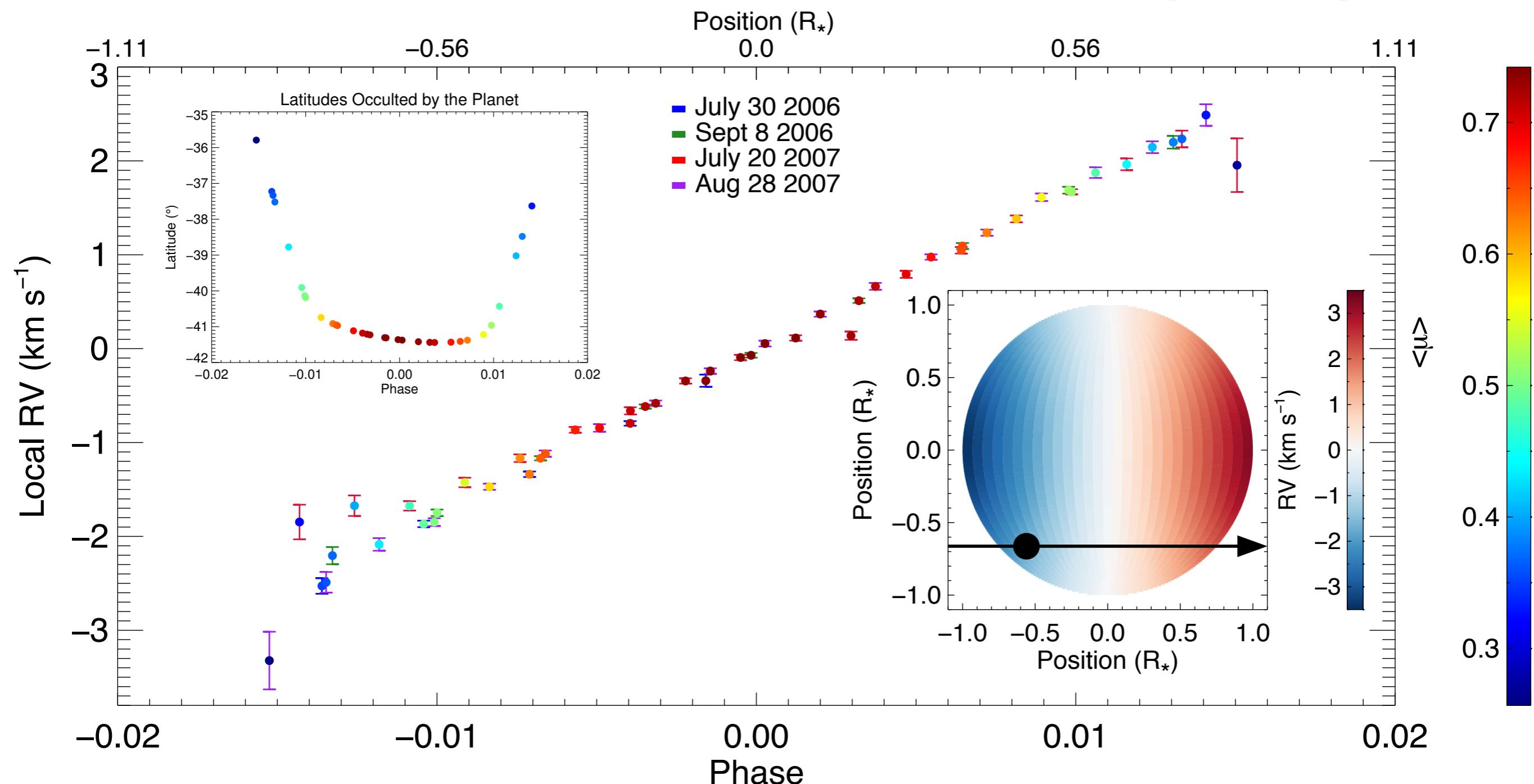
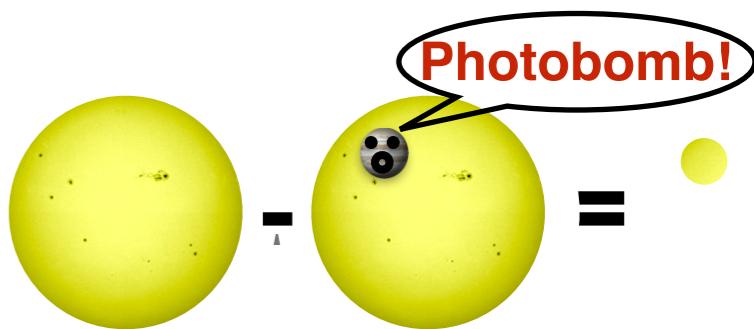


The Rossiter-McLaughlin effect RELOADED





Planets as Probes: HD 189733

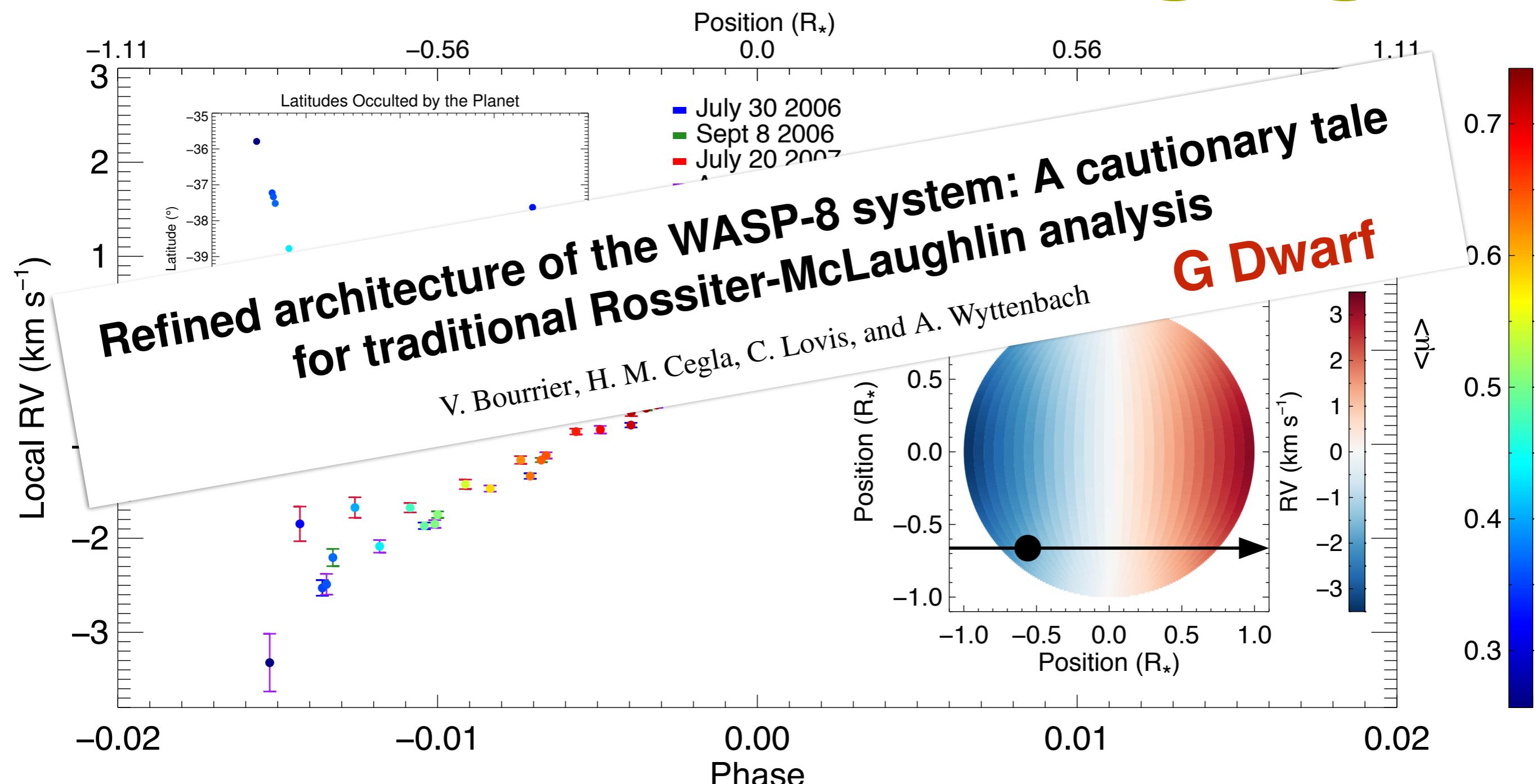
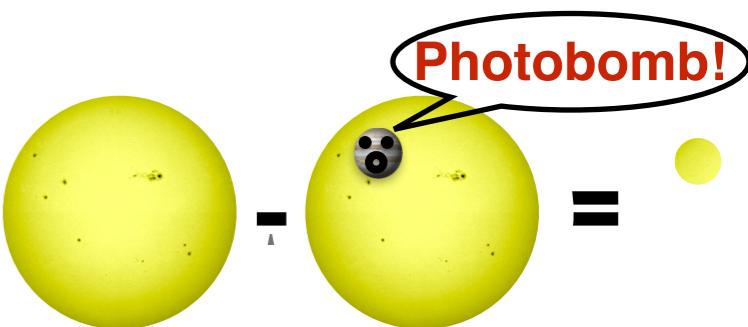


$$V_{\text{eq}} \approx 4.5^{+0.5}_{-0.4} \text{ km s}^{-1}; \alpha : 0.3-0.9; \gamma > 0.1; i_{\star} \approx 92^{+12}_{-14} \text{ }^{\circ};$$

$$\lambda \approx -0.4 \pm 0.2^{\circ}; \psi \approx 7^{+12}_{-4} \text{ }^{\circ}$$



Planets as Probes: HD 189733

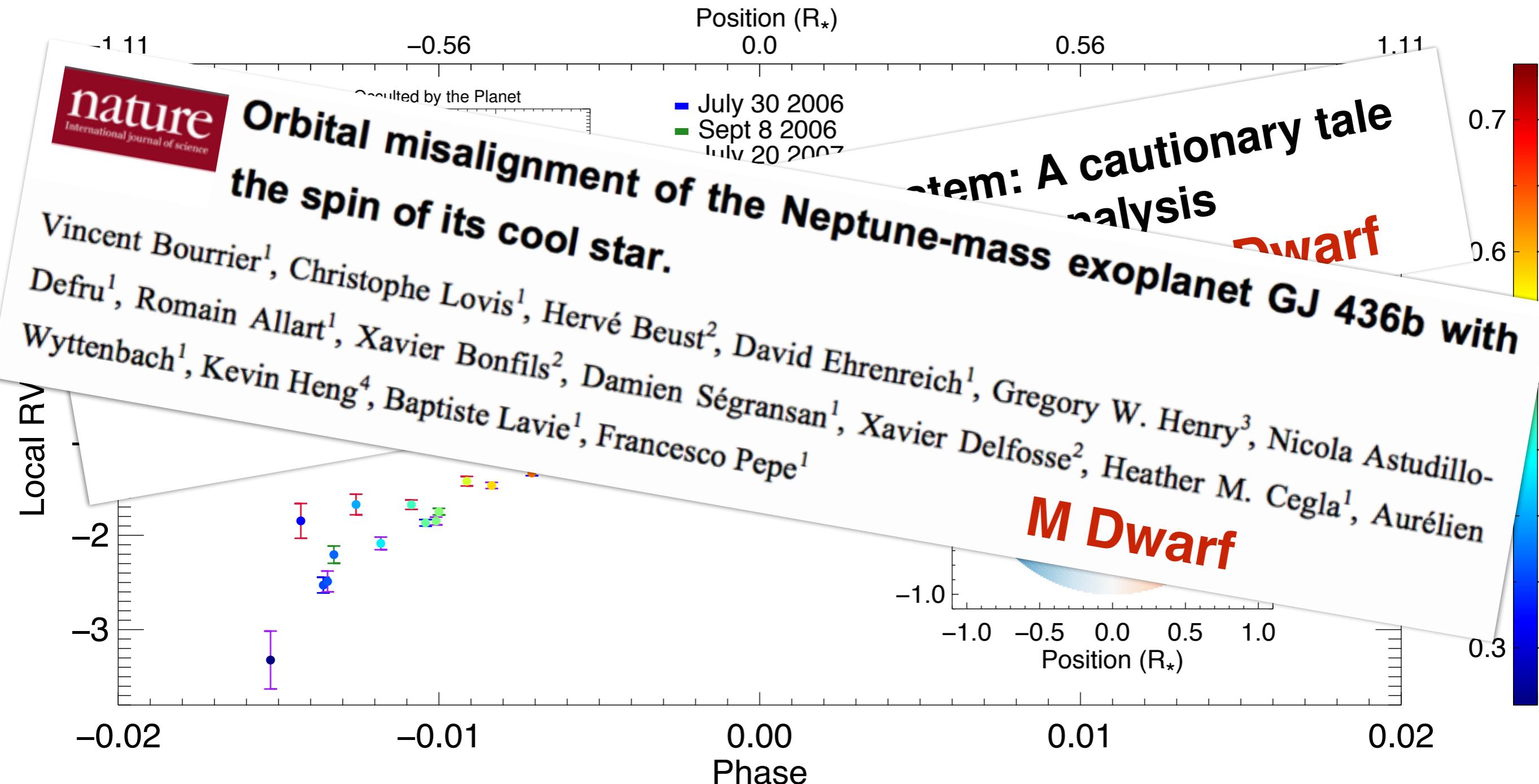


$$V_{\text{eq}} \approx 4.5^{+0.5}_{-0.4} \text{ km s}^{-1}; \alpha : 0.3-0.9; \gamma > 0.1; i_{\star} \approx 92^{+12}_{-14} \text{ }^{\circ};$$

$$\lambda \approx -0.4 \pm 0.2^{\circ}; \psi \approx 7^{+12}_{-4} \text{ }^{\circ}$$



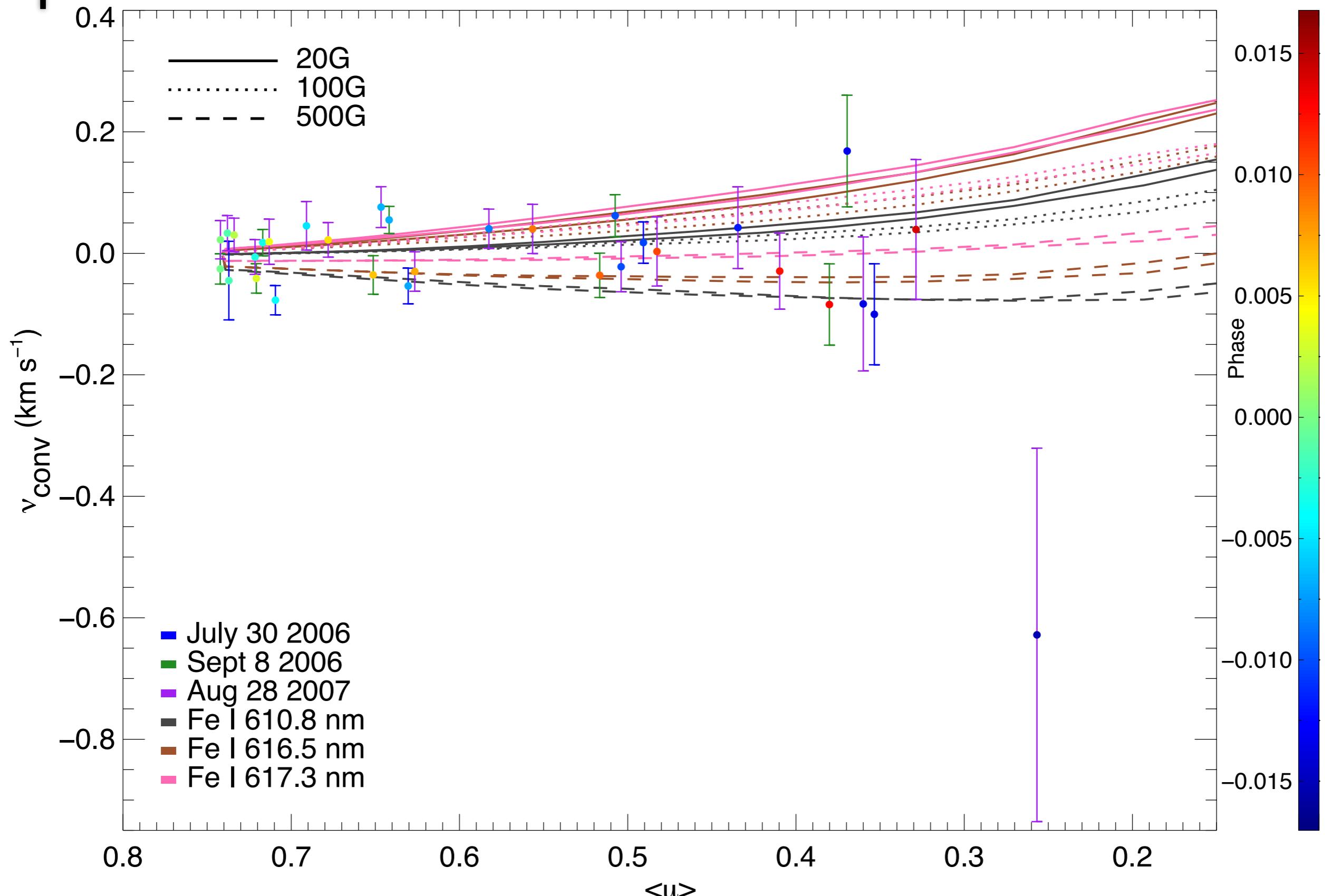
Planets as Probes: HD 189733



$$V_{\text{eq}} \approx 4.5^{+0.5}_{-0.4} \text{ km s}^{-1}; \alpha : 0.3-0.9; \gamma > 0.1; i_{\star} \approx 92^{+12}_{-14} \text{ }^{\circ};$$

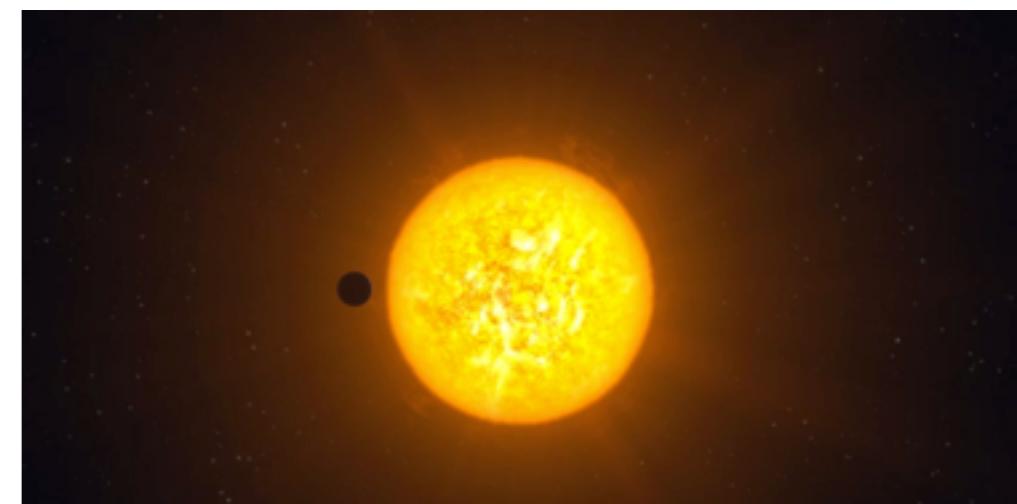
$$\lambda \approx -0.4 \pm 0.2^{\circ}; \psi \approx 7^{+12}_{-4} \text{ }^{\circ}$$

Planets as Probes: HD 189733



Summary

- Stellar surface phenomena alter line profiles & RVs
 - Impacts planet detection/confirmation/characterisation
 - Poses fundamental RV precision limit
 - MHD simulations offer pathway to characterise and disentangle
- Can use planets to spatially resolve stars
 - Probe convection, differential rotation etc.
 - Validate MHD simulations (beyond the Sun!)
 - Study evolution of star-planet systems
- Ask me about oscillations!



S What about oscillations...?

