

DETERMINATION OF STELLAR PARAMETERS FOR FGK-DWARF STARS: THE NIR
APPROACH

by

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A thesis submitted in conformity with the requirements
for the degree of Doctor of Philosophy
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Dedication

To Linnea, Henriette, and Rico

For always supporting me

Acknowledgements

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Abstract

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Chapter 5

SWEET-Cat

Part of the work during the thesis has been dedicated to regularly update SWEET-Cat¹, a catalogue with all discovered planet hosts, and the stellar parameters.

In this chapter a detailed description of SWEET-Cat will be presented. Moreover an analysis of 50 planet hosts was performed during the thesis with updated planetary parameters (mass and radius).

5.1 What is SWEET-Cat?

As mentioned above, SWEET-Cat is a catalogue of planet host stars. However, the strength of SWEET-Cat is the homogeneously analysed stars utilising the method described in Section 3.3 with FASMA or a similar tool before the creation of FASMA.

In the era with a large number of discovered exoplanets (more than 3500 confirmed exoplanet at the moment of writing), the time for in-depth statistical studies has arrived. However, when conducting these studies it is crucial to have consistent measurements of e.g. stellar atmospheric parameters. This can be obtained by using a single analysis to obtain these parameters, as it is known that different methods will lead to different results (see e.g. [Hinkel et al., 2016](#), for a recent review).

To obtain stellar atmospheric parameters from one method is an on-going goal with SWEET-Cat, where high quality spectra are obtained for stars hosting planets. These are used to determine the stellar parameters in a homogeneous way. All stars in SWEET-Cat analysed with the method from our group are marked with a flag showing whether it is analysed homogeneously or not. The columns provided in SWEET-Cat are summarised in Table 5.1. It is important to note that SWEET-Cat does not include any planetary parameters.

SWEET-Cat is updated on a weekly basis if new planet hosts are discovered, and whenever planet hosts have been analysed with the method from our group, as described in this thesis.

5.2 Data for 50 planet hosts

In this section the data for a large update to SWEET-Cat will be described. The majority of the data comes as a result from proposals submitted for observational time, while some of the data was found in the archive. In the next section the analysis of the 50 spectra will be presented along with the results.

¹ <https://www.astro.up.pt/resources/sweet-cat/>

Table 5.1: Columns in SWEET-Cat

Column	Unit	Description
Name		Popular stellar name
HD number		HD number
RA	deg	Right ascension
Dec	deg	Declination
Vmag	mag	V magnitude
$\sigma(\text{Vmag})$	mag	Error on V magnitude
π	mas	Parallax
$\sigma(\pi)$	mas	Error on parallax
Source of π		Source of parallax measurement
T_{eff}	K	Effective temperature
$\sigma(T_{\text{eff}})$	K	Error on effective temperature
$\log g$	dex	Surface gravity
$\sigma(\log g)$	dex	Error on surface gravity
$\log g_{\text{LC}}$	dex	Surface gravity corrected from light curves
$\sigma(\log g_{\text{LC}})$	dex	Error on surface gravity corrected from light curves
ξ_{micro}	km/s	Micro turbulence
$\sigma(\xi_{\text{micro}})$	km/s	Error on micro turbulence
[Fe /H]	dex	Metallicity
$\sigma([\text{Fe} /\text{H}])$	dex	Error on metallicity
Mass	M_{\odot}	Stellar mass
$\sigma(\text{Mass})$	M_{\odot}	Error on stellar mass
Reference		Reference for parameters
Homogeneity flag	0/1	0 for not homogeneous analysis, 1 otherwise
Last updated	date	Last updated
Comments		Any special remarks/comments (e.g. M star)

5.2.1 Proposals for observation time

5.2.2 Data collected from proposals

5.2.3 Data collected from archive

5.3 Analysis

The method of determining atmospheric parameters from the curve of growth analysis has been applied several times in the optical (see e.g. [Mortier et al., 2013b](#); [Santos et al., 2013](#); [Sousa et al., 2011](#); [Tsantaki et al., 2013](#)). When studying stars with planets and any correlations between stellar and planetary parameters it is important to have a homogeneous characterisation of the stars. An effort to create such a sample was initiated by [Santos et al. \(2013\)](#) with the SWEET-Cat database. The motivation to homogenise the stellar hosts is mainly to compare the hosts and make statistical studies on one consistent scale. When doing these statistical studies, the results might otherwise suffer from offsets between different methods.

The skills acquired during the NIR studies as mentioned above were directly translated into deriving parameters for a sample of 50 known planet host stars that were not previously analysed by our group ([Andreasen et al., 2017b](#)). The spectra of these stars were required at UVES, FIES, HARPS, and ESPaDOnS with the mean S/N higher than 200.

A Hertzsprung-Russell diagram of the sample can be seen in Figure 5.1. The sample covers a large

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about the data
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range of T_{eff} , FGK, while there are both dwarf, sub-giant, and some giant stars. The colours indicate the $\log g$. In order to determine the luminosity of each star the simple relation

$$L = 4\pi R^2 \sigma T_{\text{eff}}^4$$

is used, where L is the luminosity, R is the stellar radius, and σ is the Stefan-Boltzmann constant. In solar units this relation is simply:

$$\frac{L}{L_{\odot}} = \left(\frac{R}{R_{\odot}} \right)^2 \left(\frac{T_{\text{eff}}}{T_{\text{eff},\odot}} \right)^4$$

In order to determine the the stellar radius, the empirical relation from [Torres et al. \(2010\)](#) was used.

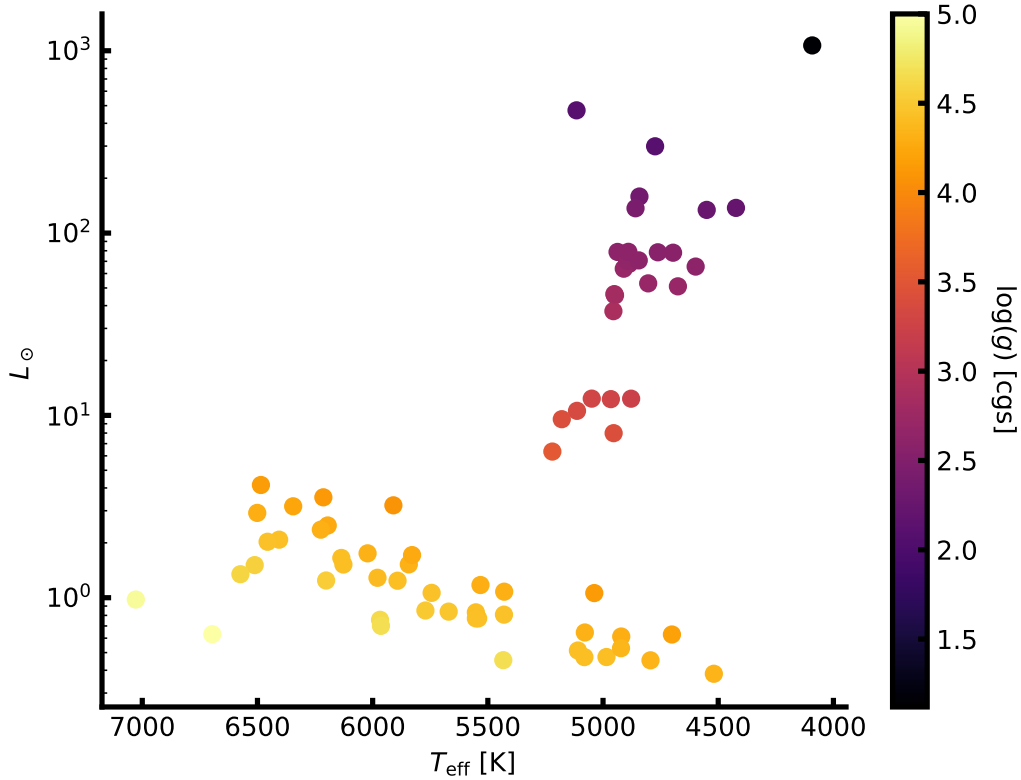


Figure 5.1: A Hertzsprung-Russell diagram of the sample of 50 planet host stars added to SWEET-Cat. The parameters were derived using optical high resolution and high S/N spectra in tandem with FASMA and an optical line list. The colour scale shows the derived $\log g$ for each star.

The parameters were derived using FASMA with the optical line list compiled by [Sousa et al. \(2008\)](#) and [Tsantaki et al. \(2013\)](#) for stars where T_{eff} was below 5200 K. All the new derived parameters were added to SWEET-Cat, available for the community.

With these updated parameters the completeness of SWEET-Cat for stars brighter than V magnitude 10 is 85% (77% for stars brighter than 12). For fainter stars it is time expensive to acquire spectra of the

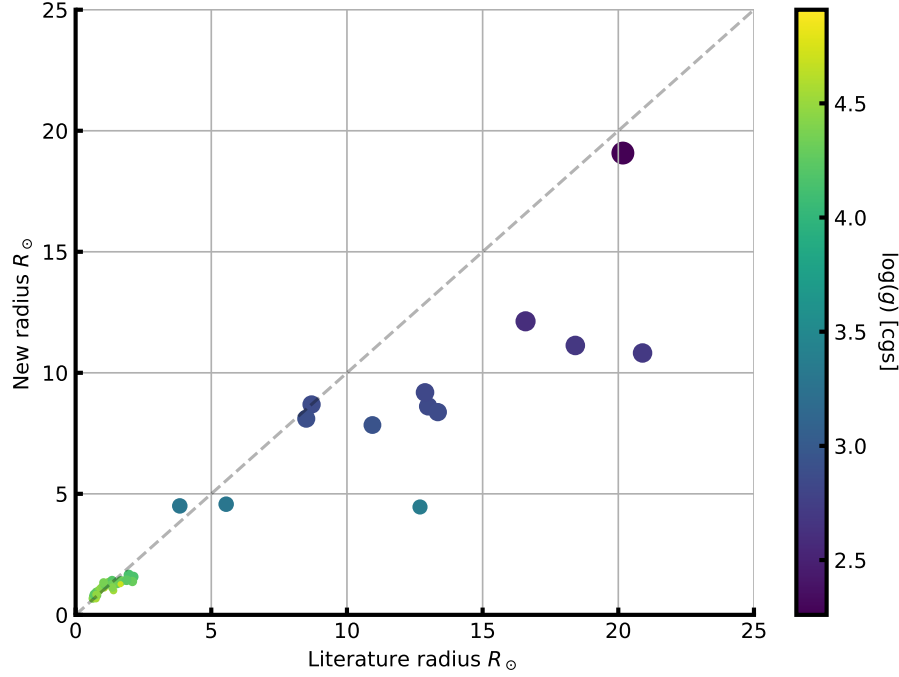


Figure 5.2: Stellar radius on both axes calculated based on [Torres et al. \(2010\)](#). The x-axis shows the stellar radius based on the atmospheric parameters from the literature, while the y-axis indicates the new homogeneous parameters presented here. The colour and size indicate the surface gravity. This clearly shows that the disagreement is biggest for more evolved stars.

quality needed for this method. Moreover, many of the fainter planet host stars have been observed with the *Kepler* space mission, where most stars are faint.

5.3.1 Changes to planetary parameters

As a results to the analysis above, it is expected that some planetary parameters will change compared with the previous literature values.

Therefore the radius and mass of all the 50 new stars updated in SWEET-Cat were computed using the empirical formula presented in [Torres et al. \(2010\)](#). Some of the stars have radii derived from different methods, usually from isochrones. These radii generally show a good correlation with radii derived from [Torres et al. \(2010\)](#) if the literature parameters of T_{eff} , $\log g$, and $[\text{Fe}/\text{H}]$ are used. However, when comparing with the new radius derived using the parameters presented here, the results can differ by up to 65%. This is shown in Figure 5.2 how the radius calculated from [Torres et al. \(2010\)](#) differs between the literature atmospheric parameters and the new homogeneous atmospheric parameters presented here. Note that stellar radii are provided by many of the authors from different discovery papers, but here the atmospheric parameters via the derivation of the stellar radius, as described above, are compared, rather than comparing the stellar radii from different methods.

In the sections below there follow a discussion of the systems (seven stars, eight exoplanets) where the radius or mass of the stars changes more than 25% and how this influences the planetary parameters. The changes in radius for a star is primarily due to changes in $\log g$, which can be used as an indicator of

the evolutionary stage of a star.

The planetary radius, mass, and semi-major axis were re-derived when possible following the three simple scaling relations based on Newton’s law of gravity (Newton, 1687) for deriving mass and distance and simple geometry for radius (see e.g. Torres et al., 2008)

$$M_{\text{pl,new}} = \left(\frac{M_{*,\text{lit}}}{M_{*,\text{new}}} \right)^{-2/3} M_{\text{pl,lit}} \quad (5.1)$$

$$R_{\text{pl,new}} = \left(\frac{R_{*,\text{lit}}}{R_{*,\text{new}}} \right) R_{\text{pl,lit}} \quad (5.2)$$

$$a_{\text{pl,new}} = \left(\frac{M_{*,\text{lit}}}{M_{*,\text{new}}} \right)^{1/3} a_{\text{pl,lit}}, \quad (5.3)$$

where the subscript “lit” denotes the value from the literature used in the comparison, the subscript “new” indicates the new computed values, the subscript “pl” is short for planet, and the subscript “*” is short for star; M , R , and a are mass, radius, and semi-major axis, respectively. Note that for the literature values, the values reported directly from the literature were used and not the derived radius and mass from Torres et al. (2010). To identify outliers, the radii and masses were compared when derived from Torres et al. (2010) since this is a measure of how the atmospheric parameters have changed.

5.3.1.1 HAT-P-46

HAT-P-46 has two known exoplanets according to ?. The outer planet HAT-P-46 c is not transiting, hence we do not have a radius for this planet. The results we present in this paper for this star come from UVES/VLT data with a S/N of 208. ? derives the following spectroscopic parameters: $T_{\text{eff}} = (6120 \pm 100) \text{ K}$, $\log g = 4.25 \pm 0.11$, and $[\text{Fe}/\text{H}] = 0.30 \pm 0.10$. We note that for this star the asteroseismic correction we apply (see Section ??) results in a corrected $\log g$ below 4.2dex, so we used the spectroscopic $\log g$ for this star.

If we derive the mass and radius of HAT-P-46 b with our new parameters, we obtain $R_{\text{pl}} = 0.93R_J$, while ? derived $R_{\text{pl}} = 1.28R_J$. We see no change in mass (? found $M_{\text{pl}} = 0.49M_J$); however, there is a decrease in the radius, and we end up with a more dense planet, $\rho_{\text{pl}} = 0.76 \text{ g/cm}^3$ from $\rho_{\text{pl}} = (0.28 \pm 0.10) \text{ g/cm}^3$.

As the secondary companion does not transit we only have a limit on the minimum mass for this planet. Here we get $M \sin i_{\text{pl}} = 1.97M_J$ and ? presented $M \sin i_{\text{pl}} = 2.00M_J$, so a very small change, as expected.

5.3.1.2 HD 120084

The exoplanet orbiting this star with a period of 2082 days and a quite eccentric orbit at 0.66 was discovered by Sato et al. (2013). The atmospheric parameters were derived by Takeda et al. (2008) using a similar method to that described in this paper. The quality of the spectra they analysed, however, were not as high as those used here. Using the HIDES spectrograph at the 188cm reflector at NAOJ, Takeda et al. (2008) reported an average S/N for their sample of 100-300 objects at a resolving power of 67 000. We used data from ESPaDOnS with a resolving power of 81 000, and with a S/N for this star of 850. With our new parameters we obtain a slightly lower stellar mass for the star at $1.93M_{\odot}$ compared to $2.39M_{\odot}$ obtained by Takeda et al. (2008), hence the minimum planetary mass is also slightly lower,

from $m_{\text{pl}} \sin i = 4.5M_J$ to $m_{\text{pl}} \sin i = 3.9M_J$. We see a 28% decrease in the stellar radius, from $9.12R_\odot$ to $7.81R_\odot$. Since there are no observations of the planet transiting, the planetary radius has not been computed.

5.3.1.3 HD 233604

HD 233604 b was discovered by [Nowak et al. \(2013\)](#), while the atmospheric parameters of the star were derived by [Zieliński et al. \(2012\)](#), who used the same method as described in this paper using the HRS spectrograph at HET with a resolving power of 60 000 with a typical S/N at 200-250. We obtained the spectrum for this star using the FIES spectrograph with a slightly higher resolution at 67 000, and similar but also slightly higher S/N at 320 for this star.

This planet is in a very close orbit with a semi-major axis of $\sim 15R_*$ (R_* is the stellar radius) using the parameters from [Nowak et al. \(2013\)](#). Using the updated parameters presented in this paper we see a slight increase in the stellar mass from $1.5M_\odot$ to $1.9M_\odot$, and a decrease in stellar radius from $10.5R_\odot$ to $8.6R_\odot$. This increases the semi-major axis to $\sim 21R_*$. We note that the correct stellar radii are used to describe the semi-major axis in both cases. The increase in stellar mass leads to an increase in the minimum planetary mass, from $m_{\text{pl}} \sin i = 6.58M_J$ to $m_{\text{pl}} \sin i = 7.79M_J$.

[Nowak et al. \(2013\)](#) found a high Li abundance at $A(\text{Li})_{\text{LTE}} = 1.400 \pm 0.042$ for this star and speculated that this star might have engulfed a planet. A more likely explanation is that this star has not yet reached the first dredge-up process ([Nowak et al., 2013](#)). We found a much lower value, $A(\text{Li})_{\text{LTE}} = 0.92$, and hence do not find the star to be Li rich. The Li abundance we find is in excellent agreement with [Adamów et al. \(2014\)](#). Even applying a NLTE correction, as was done in [Adamów et al. \(2014\)](#) ($A(\text{Li})_{\text{NLTE}} = 1.08$), this star is not Li rich.

5.3.1.4 HD 5583

This exoplanet was discovered by [Niedzielski et al. \(2016\)](#) with an orbital period of 139 days around a K giant. This exoplanet was discovered with the radial velocity technique, and we do not have a planetary radius. The stellar parameters were derived in a similar manner to that presented here (see [Niedzielski et al., 2016](#), and references therein); our biggest disagreement is in the surface gravity. We derive a $\log g$ that is higher by 0.34 dex, which gives a stellar radius that is smaller by 37%. The derived mass is 15% higher, which in turn increases the minimum planetary mass from $m_{\text{pl}} \sin i = 5.78M_J$ to $m_{\text{pl}} \sin i = 8.63M_J$. Even with the increase in mass, it is still within the planetary regime for most inclinations, as was noted by [Niedzielski et al. \(2016\)](#).

5.3.1.5 HD 81688

This exoplanet was discovered by [Sato et al. \(2008\)](#) with the RV method. The host star is a metal-poor K giant. The atmospheric parameters presented in [Sato et al. \(2008\)](#) are obtained via the same method as presented in this paper, and the agreement is quite good. Once again the big disagreement is in the surface gravity: ours is 0.48 dex higher. Even though the stellar parameters, and hence the planetary parameters, do change, the radius and mass we derive are not far from the values presented in the paper by [Sato et al. \(2008\)](#). This is a case where the star was marked as an outlier, due to the comparison between the radius and mass derived from [Torres et al. \(2010\)](#).

The new stellar mass is the same as before, $2.1M_{\odot}$. The stellar radius changed from $13.0R_{\odot}$ to $10.8R_{\odot}$. Since a transit of this star has not been observed and the stellar mass remains the same, we do not see any change in the planetary parameters.

We note that this system is in an interesting configuration with a very close orbit around an evolved star. This system, among others, has been the subject of work on planet engulfment (see e.g. [Kunitomo et al., 2011](#)).

5.3.1.6 HIP 107773

The planetary companion was presented in [Jones et al. \(2015\)](#) as an exoplanet around an intermediate-mass evolved star. The stellar parameters were obtained from the analysis by [Jones et al. \(2011\)](#) using the same method as presented here, but with a different line list, which might lead to some disagreements. For this star we derive a higher $\log g$ (2.83 dex compared to 2.60 dex), thus we derive a slightly smaller star with $11.6R_{\odot}$ to $9.2R_{\odot}$ and $2.4M_{\odot}$ to $2.1M_{\odot}$ for radius and mass of the star, respectively. The other atmospheric parameters are very similar to those derived by [Jones et al. \(2011\)](#). This leads to a reduced minimum mass of the planetary companion from $m \sin i = 1.98M_J$ to $m \sin i = 1.78M_J$. The planetary radius has not been measured.

5.3.1.7 WASP-97

The exoplanet orbiting WASP-97 was discovered by [Hellier et al. \(2014\)](#). The host star parameters were derived using a similar method to that described in this paper after co-adding several spectra from the CORALIE spectrograph. They reach a S/N of 100 with a spectral resolution of 50 000. The parameters presented here come from the UVES spectrograph with a S/N of more than 200.

The parameters do not change much for this planet. The planetary mass changes from $m_{\text{pl}} \sin i = 1.32M_J$ to $m_{\text{pl}} \sin i = 1.37M_J$ and the radius from $1.13R_J$ to $1.42R_J$. This affects the density quite strongly; it changes from 1.13 g/cm^3 to 0.59 g/cm^3 . This exoplanet is then in the same category as Saturn; its density is lower than water, but it is slightly larger than Jupiter.

5.3.1.8 ω Serpentis (ome Ser)

The exoplanet orbiting this star with a period of 277 days and an eccentric orbit at 0.11 was also presented by [Sato et al. \(2013\)](#). The atmospheric parameters were derived in the same way as for HD 120084. We used data from FIES with a resolving power of 67 000, and with a S/N for this star of 1168. With our new parameters we obtain a slightly higher stellar mass for the star at $2.19M_{\odot}$ compared to the value of $2.17M_{\odot}$ obtained by [Takeda et al. \(2008\)](#). This change is not significant enough to change the minimum planetary mass at $m_{\text{pl}} \sin i = 1.7M_J$. The stellar radius decreases by more than one solar radius, from $12.3R_{\odot}$ to $11.1R_{\odot}$. However, since there are no observations of transiting exoplanets, we cannot see the change in the planetary radius.

5.3.1.9 α Ursa Major (omi UMa)

omi UMa b was discovered by [Sato et al. \(2012\)](#) using the RV method. The stellar parameters are from [Takeda et al. \(2008\)](#), as discussed above. The spectrum used for this star is from ESPaDOnS with a S/N of more than 500 compared to the value of 100-300 reached for the large sample presented in [Takeda et al. \(2008\)](#). The luminosity and mass for omi UMa were obtained from theoretical evolutionary tracks

(see [Sato et al., 2012](#), and references therein). The radius was then estimated using the Stefan-Boltzmann relationship, using the measured luminosity and T_{eff} .

The parameters presented here mainly differ in the surface gravity: ours is 0.72 dex higher at $\log g = 3.36$. This leads to a big change in stellar mass and radius from $3.1M_{\odot}$ to $1.6M_{\odot}$ and $14.1R_{\odot}$ to $4.5R_{\odot}$, respectively. [Sato et al. \(2012\)](#) have reported that omi UMa b is the first planet candidate around a star more massive than $3M_{\odot}$. With these updated results, the minimum mass of the planet is now $m \sin i = 2.7M_J$, whereas previously it was $m \sin i = 4.1M_J$ ([Sato et al., 2012](#)). The exoplanet is not reported to transit, as seen from Earth, so we do not have a radius for this exoplanet, which would have changed a great deal with these new results.

5.4 Discovering two giant planet populations

SWEET-Cat was recently combined with planetary masses to see two distinctive populations for giant planets by [Santos et al. \(2017\)](#). This can be seen in the mass histogram in Figure 5.3 for the full sample of giant planets, with masses higher than 1 Jupiter mass and lower than 20 Jupiter masses, and for a sample constrained by: $4000 \text{ K} \leq T_{\text{eff}} \leq 6500 \text{ K}$ in order to have reliably atmospheric parameters from spectroscopic data, orbital periods above 10 days to avoid hot jupiters whose formation and migration process is debated ([Ngo et al., 2016](#), see e.g.), orbital periods below 5 years to allow for the sample to be reasonable complete. Last only stars brighter than 13 magnitude were included to ensure that the planetary masses can have been derived with reasonable confidence using the radial velocities.

By separating the distribution into two at $4M_{Jup}$, it can be shown (see [Santos et al., 2017](#), for details) that the stars hosting the more massive giant planets are in average more metal-poor compared to the stars hosting the lower mass giant planets. This suggest two different stellar populations forming giant planets.

5.5 Future work

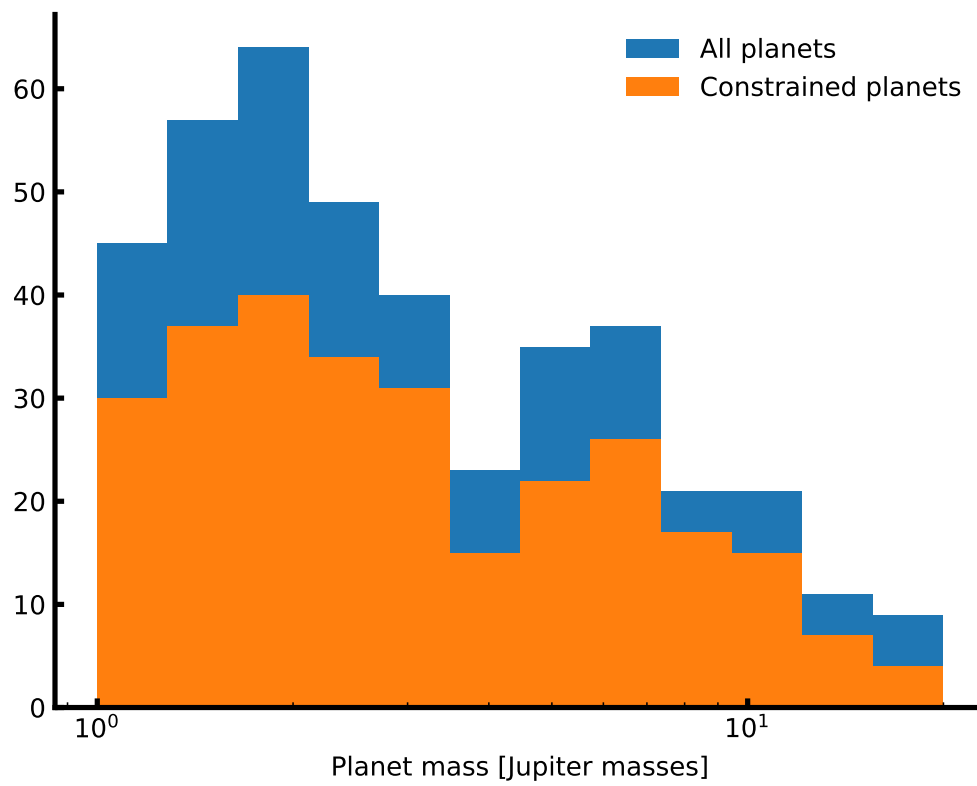


Figure 5.3: Giant planet masses for the full sample and constrained sample (see text for details). This study was performed by Santos et al. (2017) to distinct two giant planet populations.

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