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2 μ/π separation using

3 Convolutional Neural Networks

4 for the MicroBooNE

5 Charged Current Inclusive Cross Section

6 Measurement

7

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μ BooNE



26

Abstract

27

The purpose of this thesis was to use Convolutional Neural Networks
28 (CNN) to separate μ' s and π' s for use in increasing the acceptance rate
29 of μ' s below the implemented 75cm track length cut in the Charged
30 Current Inclusive (CC-Inclusive) event selection for the CC-Inclusive
31 Cross-Section Measurement. In doing this, we increase acceptance
32 rate for CC-Inclusive events below a specific momentum range.

33

Dedication

34

I dedicate this dissertation to the two important women in my life; My
35 wife and my mom. Both have been there cheering me on giving me strength
36 and love as I worked towards the hardest accomplishment I've ever done.

37

Jessica Nicole Esquivel

38

Acknowledgements

39 Of the many people who deserve thanks, some are particularly prominent, such as
40 my supervisor...

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316	ple normalized to $5 * 10^{19}$ PoT. The last column gives the fraction of	
317	this signal or background type to the total selected events per CNN	
318	probability.	76

*"If they don't give you a seat at the table,
bring a folding chair."*

— Shirley Chisholm

³²¹ **Chapter 1**

³²² **Introduction**

³²³ This thesis will be a description of work done to further increase efficiency and purity
³²⁴ of the charged current inclusive cross section measurement using the MicroBooNE
³²⁵ detector. It will also describe the MicroBooNE detector, what neutrinos are, the
³²⁶ charged current inclusive cross section measurement and its importance as well as
³²⁷ convolutional neural networks and how they can be used in μ/π separation. Chapter
³²⁸ 2 will talk about the background of neutrinos and the people and detectors that
³²⁹ discovered neutrinos as well as an in depth history of neutrino oscillation and the
³³⁰ discovery that neutrinos have mass.

³³¹ Chapter 3 will discuss the MicroBooNE experiment, specifically, how Liquid
³³² Argon Time Projection Chambers work, the Light Collection System and the Electronic
³³³ and Readout Trigger systems. This chapter will also describe the Booster Neutrino
³³⁴ Beam sationed at Fermilab.

³³⁵ Chapter 4 will discuss the work that was done to detect the first neutrinos seen in
³³⁶ the MicroBooNE detector and the software reconstruction efforts required to create an
³³⁷ automated neutrino ID filter that was used to find the first neutrinos and then was
³³⁸ later expanded on to create the charged current inclusive filter that will be discussed
³³⁹ in chapter 5

³⁴⁰ Chapter 6 will give a brief description of what Convolutional Neural Networks are
³⁴¹ and how it will be used for μ/π separation in this selection. Chapter 7 will discuss
³⁴² the hardware frameworks and training methods used to train multiple Convolutional
³⁴³ Neural Networks for use in the charged current inclusive cross section measurement.
³⁴⁴ Chapters 8 and ?? will discuss the results of using Convolutional Neural Networks on
³⁴⁵ monte-carlo and data to sift out charged current inclusive neutrino events.

³⁴⁶ Chapter 2

³⁴⁷ Neutrinos

³⁴⁸ 2.1 What are Neutrinos

³⁴⁹ Neutrinos are one of the fundamental particles which make up the universe. They are
³⁵⁰ also one of the least understood. Neutrinos are not affected by the electromagnetic
³⁵¹ forces because they do not have electric charge. Neutrinos are affected by a "weak"
³⁵² sub-atomic force of much shorter range than electromagnetism, and are therefore able
³⁵³ to pass through great distances in matter without being affected by it. Until the late
³⁵⁴ 90's, neutrinos were thought to have no mass. Due to their mass, neutrinos are also
³⁵⁵ affected by gravity. Neutrinos are created by radioactive decay or nuclear reactions
³⁵⁶ such as the ones that happen in the sun, in nuclear reactors or when cosmic rays hit
³⁵⁷ atoms. There are three types of neutrinos, ν_e , ν_μ and ν_τ which correspond to their
³⁵⁸ charged lepton pairs.

³⁵⁹ As previously stated, neutrinos are very weakly interacting; in fact, neutrinos can
³⁶⁰ pass unscathed through a wall of lead several hundred light-years thick. Because
³⁶¹ neutrinos interact so rarely, studying neutrinos requires a massive detector and a
³⁶² powerful neutrino source. With that being said, we can only infer their existence when
³⁶³ they interact in a detector. In a collision, distinct charged particles are produced with
³⁶⁴ each type of neutrino. An electron neutrino will create an electron, a muon neutrino
³⁶⁵ will create a muon, and a tau neutrino will create a tau. The track the particle leaves
³⁶⁶ in the detector is how one figures out what type of neutrino interaction was "seen".
³⁶⁷ Liquid Argon Time Projection Chambers are the newest type of detectors being used to
³⁶⁸ study neutrinos due to their excellent imaging and particle identification capabilities.

³⁶⁹ 2.2 History of Neutrinos

³⁷⁰ The neutrino was first postulated by Wolfgang Pauli in 1931 to explain how beta
³⁷¹ decay could resolve the conservation of energy, momentum and angular momentum
³⁷² problem. Pauli suggested that this missing energy might be carried off, unseen, by a
³⁷³ neutral particle (he called neutron) which was escaping detection. James Chadwick
³⁷⁴ discovered a much heavier nuclear particle in 1932 that he also named neutron, leaving
³⁷⁵ two particles with the same name. Enrico Fermi was the first person to coin the
³⁷⁶ term neutrino (which means little neutral one in latin) in 1933 to fix this confusion.
³⁷⁷ Fermi's paper, which was published in 1934, unified Pauli's neutrino with Paul Dirac's
³⁷⁸ positron and Werner Heisenberg's neutron-proton model and his theory accurately
³⁷⁹ explained many experimentally observed results. Wang Ganchang first proposed the
³⁸⁰ use of beta capture to experimentally detect neutrinos and in 1959 Clyde Cowan and
³⁸¹ Frederick Reines published their work stating that they had detected the neutrino.
³⁸² The experiment called for antineutrinos created in a nuclear reactor by beta decay that
³⁸³ reacted with protons producing neutrons and positrons: $\nu_e + p^+ \rightarrow n^0 + e^+$. Once
³⁸⁴ this happens, the positron finds an electron and they annihilate each other and the
³⁸⁵ resulting gamma rays are detectable. The neutron is detected by neutron capture and
³⁸⁶ the releasing of another gamma ray. In 1962 Leon M. Lenderman, Melvin Schwartz
³⁸⁷ and Jack Steinberger were the first to detect interactions of the muon neutrino. The
³⁸⁸ first detection of the tau neutrino was announced in the summer of 2000 by the
³⁸⁹ DONUT collaboration at Fermilab. In the late 1960s, many experiments found that the
³⁹⁰ number of electron neutrinos arriving from the sun was around 1/3 to 1/2 the number
³⁹¹ predicted by the Standard Solar Model. This became known as the solar neutrino
³⁹² problem and remained unresolved for around thirty years. This problem was resolved
³⁹³ by the discovery of neutrino oscillation and mass. [1]

³⁹⁴ 2.3 Neutrino Oscillations

³⁹⁵ Neutrino oscillation was first predicted by Bruno Pontecorvo. It describes the phe-
³⁹⁶ nomenon of a neutrino created with a specific lepton flavor (electron, muon or tau)
³⁹⁷ that is later measured to have a different flavor. Neutrino oscillation is important
³⁹⁸ theoretically and experimentally due to the fact that this observation implies that the
³⁹⁹ neutrino has a non-zero mass, which is not part of the original Standard Model of
⁴⁰⁰ particle physics. [2]

401 2.3.1 Solar Oscillations and the Solar Neutrino Problem

402 The solar neutrino flux derived from Bahcall's Standard Solar Model is shown in figure
 403 2.1. Nuclear fusion and decay processes produce an abundant amount of neutrinos.
 404 The standard solar model predicts that these reactions produce several groups of
 405 neutrinos, each with differing fluxes and energy spectra. The figure also shows the
 406 ranges of detection of existing solar neutrino experiments in different shades of blue
 407 to illustrate that they sample different portions of the solar neutrino energy spectrum.
 408 Three of these experiments, plus a new one, are discussed below.

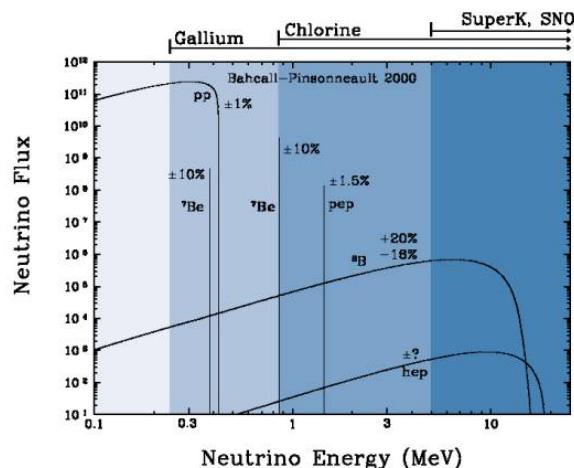


Figure 2.1: The Standard Solar Model

409 Since neutrinos rarely interact with matter, they pass through the sun and the earth
 410 undetected. About 65-billion neutrinos from the sun stream through every square
 411 centimeter on the Earth every second, yet we are oblivious to their passage in our
 412 every-day lives. [3]

413 The first experiment to detect the effects of neutrino oscillation was the Ray Davis's
 414 Homestake Experiment. The detector was stationed in the Homestake Gold Mine in
 415 Lead, South Dakota. It was 1,478 meters underground and was $380\ m^3$. The detector
 416 was filled with perchloroethylene. Perchloroethylene was chosen because of its high
 417 concentrations of chlorine. When an ν_e interacted with chlorine-37 atom, the atom
 418 would transform to argon-37 which was then extracted and counted. The neutrino
 419 capture reaction is shown in equation 2.1. Davis observed a deficit of about 1/3
 420 the flux of solar neutrinos that was predicted by Bahcall's Standard Solar Model.

- ⁴²¹ The unexplained difference between the measured solar neutrino flux and model
⁴²² predictions lead to the Solar Neutrino Problem. [4]



⁴²³ While it is now known that the Homestake Experiment detected neutrinos, some
⁴²⁴ physicist were weary of the results. Conclusive evidence of the Solar Neutrino Problem
⁴²⁵ was provided by the Kamiokande-II experiment, a water Cherenkov detector with
⁴²⁶ a low enough energy threshold to detect neutrinos through neutrino-electron elastic
⁴²⁷ scattering. In the elastic scattering interaction the electrons coming out of the point of
⁴²⁸ reaction strongly point in the direction that the neutrino was traveling, away from the
⁴²⁹ sun. While the neutrinos observed in Kamiokande-II were clearly from the sun, there
⁴³⁰ was still a discrepancy between Kamiokande-II and Homestake; The Kamiokande-
⁴³¹ II experiment measured about 1/2 the predicted flux, rather than the 1/3 that the
⁴³² Homestake Experiment saw.

⁴³³ The solution to the solar neutrino problem was finally experimentally determined
⁴³⁴ by the Sudbury Neutrino Observatory(SNO). The Ray Davis's Homestake Experiment
⁴³⁵ was only sensitive to electron neutrinos, and the Kamiokande-II Experiment was
⁴³⁶ dominated by the electron neutrino signal. The SNO experiment had the capability to
⁴³⁷ see all three neutrino flavors. Because of this, it was possible to measure the electron
⁴³⁸ neutrinos and total neutrino flux. The experiment demonstrated that the deficit was
⁴³⁹ due to the MSW effect, the conversion of electron neutrinos from their pure flavor
⁴⁴⁰ state into the second neutrino mass eigenstate as they passed through a resonance
⁴⁴¹ due to the changing density of the sun. The resonance is energy dependent, and is
⁴⁴² visible near 2MeV. The water cherenkov detectors only detect neutrinos above about
⁴⁴³ 5MeV, while the radiochemical experiments were sensitive to lower energy (0.8MeV
⁴⁴⁴ for chlorine, 0.2MeV for gallium), and this turned out to be the source of the difference
⁴⁴⁵ in the observed neutrino rates at the two types of experiments. Figure 2.2 shows
⁴⁴⁶ Homestake, Kamiokande-II and SNO experiments.

⁴⁴⁷ MSW Effect

⁴⁴⁸ The Mikheyev-Smirnov-Wolfenstein effect is a process which acts to modify neu-
⁴⁴⁹ trino oscillations in matter. The presence of electrons in matter changes the energy

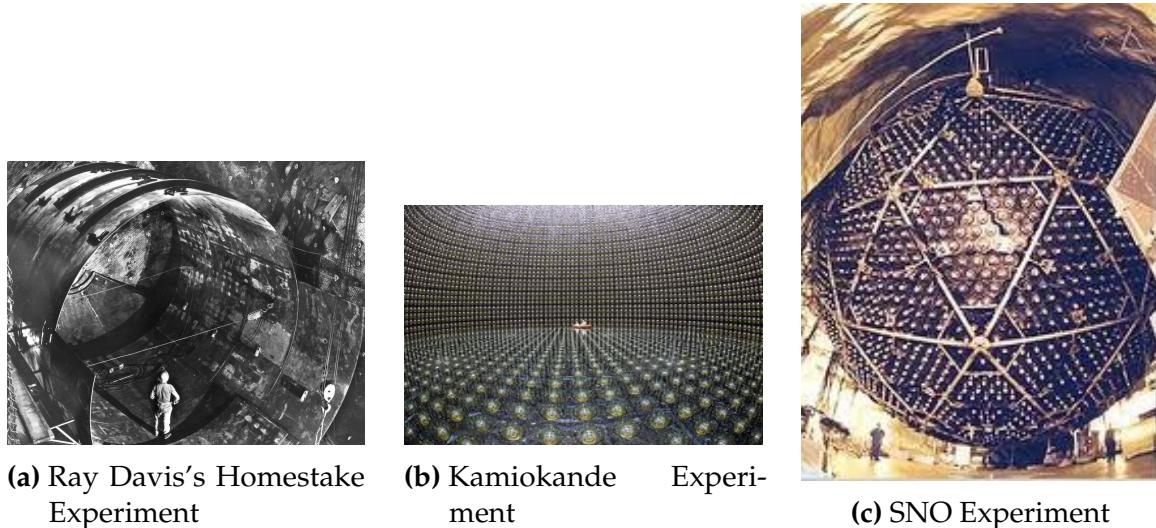


Figure 2.2: Solar Neutrino Experiments

450 levels of the mass eigenstates of neutrinos due to charged current coherent forward
 451 scattering of the electron neutrinos. This coherent forward scattering is similar to
 452 the electromagnetic process with respect to the refractive index of light in a medium.
 453 Because of this MSW Effect, neutrinos in vacuum have a different effective mass than
 454 neutrinos in matter and because neutrino oscillations depend on the squared mass
 455 difference of the neutrinos, the neutrino oscillations are different in matter than in
 456 vacuum. This effect is important at the sun where electron neutrinos are produced.
 457 The neutrinos of high energy leaving the sun are in a vacuum propagation eigenstate
 458 ν_2 that has a very small overlap with the electron neutrino $\nu_e = \nu_1 \cos(\theta) + \nu_2 \sin(\theta)$
 459 seen by the charged current reactions in Kamiokande-II and SNO. The discrepancy of
 460 the deficit between SNO, Kamiokande-II and Homestake is due to the energy of the
 461 solar neutrinos. The MSW effect "turns on" at about 2MeV and at lower energies, this
 462 MSW effect is negligible. [5]

463 2.3.2 Atmospheric Oscillations and the Atmospheric Neutrino 464 Anomaly

465 Atmospheric neutrinos are neutrinos that stem from the decay hadrons coming from
 466 primary cosmic rays. The dominant part of the decay chain is shown in equations 2.2
 467 and 2.3

$$\pi^+ \rightarrow \mu^+ \nu_\mu \mu^+ \rightarrow e^+ \nu_e \bar{\nu}_\mu \quad (2.2)$$

468

$$\pi^- \rightarrow \mu^- \bar{\nu}_\mu \mu^- \rightarrow e^- \bar{\nu}_e \nu_\mu \quad (2.3)$$



Figure 2.3: Cosmic Ray Shower

469 Figure 2.3 shows the cosmic ray shower. In general, these neutrinos have energies
 470 from 1GeV to 100s of GeV and the ratio of ν_μ s to ν_e s equals to 2 (see equation 2.4)

$$R = \frac{(\nu_\mu + \bar{\nu}_\mu)}{(\nu_e + \bar{\nu}_e)} \quad (2.4)$$

471 There have been two types of detectors used to study atmospheric neutrinos: Water
 472 Cherenkov detectors and tracking calorimeters. Super-Kamiokande is the detector we
 473 will focus on. These atmospheric detector experiments measure the ratio of ν_μ to ν_e .
 474 They also measure the zenith angle distribution of the neutrinos. These experiments
 475 report a double ratio (shown in equation 2.5). This double ratio is the ratio measured
 476 in the detector to the ratio that's expected which is 2. If the double ratio equals to 1, the
 477 data agrees with the prediction. Various measurements from multiple experiments
 478 are shown in figure 2.4. Except for Frejus, all R measurements are less than 1. This
 479 discrepancy between the predicted R and the measured R became known as the
 480 Atmospheric Neutrino Anomaly.

$$R = \frac{(N_\mu / N_e)_{DATA}}{(N_\mu / N_e)_{SIM}} \quad (2.5)$$

481 Kamiokande-II has the the capability of measuring the direction of the incoming
 482 neutrinos. The expectation of atmospheric neutrino detection is that the flux be

Experiment	Type of experiment	R
Super-Kamiokande	Water Cerenkov	0.675 ± 0.085
Soudan2	Iron Tracking Calorimeter	0.69 ± 0.13
IMB	Water Cerenkov	0.54 ± 0.12
Kamiokande	Water Cerenkov	0.60 ± 0.07
Frejus	Iron Tracking Calorimeter	1.0 ± 0.15

Figure 2.4: Measurements of the double ratio for various atmospheric neutrino experiments

483 isotropic due to the fact that atmospheric neutrinos can reach the detector from all
 484 directions. Kamiokande-II noticed that muon-like data did not agree well with this
 485 expectation. At low energies approximately half of the ν_μ are missing over the full
 486 range of zenith angles. At high energies the number of ν_μ coming down from above
 487 the detector seems to agree with expectation, but half of the same ν_μ coming up from
 488 below the detector are missing. This anomaly can be easily explained by neutrino
 489 flavor oscillations. Due to the fact that the neutrino travels less distance coming
 490 straight down into the detector (about 15km) than coming up from the bottom of the
 491 detector(13000km) changes the probability of oscillation. The probability of oscillation
 492 for the muon neutrinos coming down into the detector is roughly zero, whereas
 493 for neutrinos coming up, the oscillation probability is $\sin^2(2\theta)$. Both the solar and
 494 atmospheric neutrino problems can be explained by neutrino oscillation so its fitting
 495 to derive this phenomenon mathematically. In the next two sections, two flavor and
 496 three flavor neutrino oscillation derivations will be explained.

497 2.3.3 Two Flavor Neutrino Oscillation Formulation

498 The flavor eigenstates can oscillate between each other because they are composed
 499 of an add-mixture of mass eigenstates(ν_1, ν_2). Figure 2.5 shows the mass and flavor
 500 eigenstates rotated by an angle θ which is the mixing angle.

501 In matrix form the wavefunctions are:

$$\begin{pmatrix} \nu_\mu \\ \nu_e \end{pmatrix} = \begin{pmatrix} \cos\theta & \sin\theta \\ -\sin\theta & \cos\theta \end{pmatrix} * \begin{pmatrix} \nu_1 \\ \nu_2 \end{pmatrix} \quad (2.6)$$

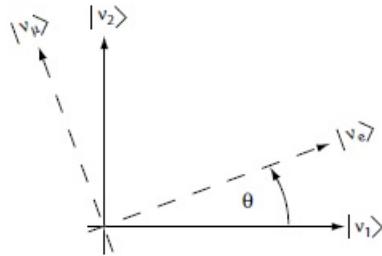


Figure 2.5: The flavor eigenstates are rotated by an angle θ with respect to the mass eigenstates

502 Applying the time evolution operator to ν_μ :

$$|\nu_\mu(t)\rangle = -\sin\theta|\nu_1\rangle e^{-i\frac{E_1 t}{\hbar}} + \cos\theta|\nu_2\rangle e^{-i\frac{E_2 t}{\hbar}} \quad (2.7)$$

503 where $E_1 = \sqrt{p^2 c^2 + m_1^2 c^4}$ and $E_2 = \sqrt{p^2 c^2 + m_2^2 c^4}$ and $p_1 = p_2$. For the time

504 being, let us assume $\hbar = c = 1$. With this assumption: $E_1 = \sqrt{p^2 + m_1^2}$ and $E_2 =$
505 $\sqrt{p^2 + m_2^2}$. The next modifications is to assume neutrinos are relativistic:

$$\gamma = \frac{E}{m_o c^2} = \frac{\sqrt{p^2 c^2 + m_o^2 c^4}}{m_o c^2} \gg 1 \quad (2.8)$$

506 because of this,

$$p \gg m_o \quad (2.9)$$

507

$$E = \sqrt{p^2 + m_o^2} = p\sqrt{1 + m_o^2/p^2} \simeq p + \frac{1}{2}\frac{m_o^2}{p} \quad (2.10)$$

508 where the binomial expansion is used. Now E_1 and E_2 can be written as:

$$E_1 \simeq p + \frac{1}{2}\frac{m_1^2}{p} \text{ and } E_2 \simeq p + \frac{1}{2}\frac{m_2^2}{p} \quad (2.11)$$

509 Now applying all these assumptions back into equation 2.7 gives us:

$$|\nu_\mu(t)\rangle = -\sin\theta|\nu_1\rangle e^{-i\left(p + \frac{1}{2}\frac{m_1^2}{p}\right)t} + \cos\theta|\nu_2\rangle e^{-i\left(p + \frac{1}{2}\frac{m_2^2}{p}\right)t} \quad (2.12)$$

$$|\nu_\mu(t)\rangle = e^{-i\left(p+\frac{1}{2}\frac{m_1^2-m_2^2}{p}\right)t} (-\sin\theta|\nu_1\rangle + \cos\theta|\nu_2\rangle) \quad (2.13)$$

₅₁₀ Substituting $\Delta m^2 = m_1^2 - m_2^2$ and $t = \frac{x}{c} = x$ and $e^{-iz} = e^{-i\left(p+\frac{1}{2}\frac{m_1^2}{p}\right)t}$ gives us:

$$|\nu_\mu(t)\rangle = e^{-iz} \left(-\sin\theta|\nu_1\rangle + \cos\theta|\nu_2\rangle e^{+ix\left(\frac{1}{2}\frac{\Delta m^2}{p}\right)} \right) \quad (2.14)$$

₅₁₁ Finding the Probability for a $\nu_\mu \rightarrow \nu_e$:

$$P(\nu_\mu \rightarrow \nu_e) = |\langle \nu_e | \nu_\mu(t) \rangle|^2 \quad (2.15)$$

₅₁₂ Remembering that $\langle \nu_i | \nu_j \rangle = \delta_{ij}$

$$\langle \nu_e | \nu_\mu(t) \rangle = e^{-iz} \left(-\sin\theta\cos\theta + \sin\theta\cos\theta e^{\frac{i\Delta m^2 x}{p}} \right) \quad (2.16)$$

₅₁₃ Taking the absolute value squared gives us:

$$P(\nu_\mu \rightarrow \nu_e) = |\langle \nu_e | \nu_\mu(t) \rangle|^2 = e^{+iz} e^{-iz} \sin^2\theta \cos^2\theta \left(-1 + e^{\frac{i\Delta m^2 x}{p}} \right) \left(-1 + e^{\frac{-i\Delta m^2 x}{p}} \right) \quad (2.17)$$

₅₁₄ Since the neutrino is relativistic we can set $p = E_\nu$ and change $x = L$. Also ₅₁₅ recognizing the trigonometric relation $(1 - \cos 2\theta)/2 = \sin^2\theta$ the above equation ₅₁₆ becomes:

$$P(\nu_\mu \rightarrow \nu_e) = \sin^2 2\theta \sin^2 \left(\frac{\Delta m^2 L}{4E_\nu} \right) \quad (2.18)$$

₅₁₇ All that's left to do now is re-introduce \hbar and c doing this we get:

$$P_{\nu_\mu \rightarrow \nu_e}(L, E) = \sin^2 2\theta \sin^2 \left(1.27 \Delta m^2 \frac{L}{E_\nu} \right) \quad (2.19)$$

₅₁₈ This equations has three important variables.

- The angle θ : This angle, as mentioned before, is called the mixing angle. It defines the difference between the flavor and the mass eigenstates. When $\theta = 0$ the mass and flavor eigenstates are identical and now oscillations occur.
- The mass squared difference, Δm^2 : Again $\Delta m^2 = m_1^2 - m_2^2$. The reason this is an important variable is because it implies that for neutrinos to oscillate, neutrinos must have mass. Furthermore, the mass squared difference also tells us that the neutrino mass eigenstates must be different.
- L/E: This is the variable that is of most interest to experimental physicists due to the fact that it is the variable that we set. L is the distance between the source and detector and E is the energy of the neutrino. For a given Δm^2 , the probability of oscillation changes with respect to L/E.

2.3.4 Three Flavor Neutrino Oscillation Formulation

Seeing the quantum mechanics involved in deriving the probability of a two flavor neutrino oscillation, it is now possible to formulate the three flavor neutrino oscillation. The three flavor neutrino oscillation formulation begins similarly to the two flavor, but there is the Pontecorvo-Maki-Nakagawa-Sakata matrix (PMNS) instead of the 2X2 matrix in the previous section. The PMNS matrix is show below:

$$\begin{pmatrix} c_{12}c_{13} & s_{12}c_{13} & s_{13}e^{-i\delta} \\ -s_{12}c_{23} - c_{12}s_{23}s_{13}e^{i\delta} & c_{12}c_{23} - s_{12}s_{23}s_{13}e^{i\delta} & s_{23}c_{13} \\ s_{12}s_{23} - c_{12}c_{23}s_{13}e^{i\delta} & -c_{12}s_{23} - s_{12}c_{23}s_{13}e^{i\delta} & c_{23}c_{13} \end{pmatrix} * \begin{pmatrix} e^{i\alpha_1/2} & 0 & 0 \\ 0 & e^{i\alpha_2/2} & 0 \\ 0 & 0 & 1 \end{pmatrix} \quad (2.20)$$

where $c_{ij} = \cos\theta_{ij}$ and $s_{ij} = \sin\theta_{ij}$

Following the same steps as before we get:

$$P_{\alpha \rightarrow \beta} = \delta_{\alpha\beta} - 4\sum Re(U_{\alpha i}^* U_{\beta i} U_{\alpha j} U_{\beta j}^*) \sin^2\left(\frac{\Delta m_{ij}^2 L}{4E}\right) 2\sum Im(U_{\alpha i}^* U_{\beta i} U_{\alpha j} U_{\beta j}^*) \sin\left(\frac{\Delta m_{ij}^2 L}{2E}\right) \quad (2.21)$$

The main things to notice here are δ_{ij} which is the CP violating term and has not been measured yet, and θ_{13} which has just been measured. CP violation is a violation

⁵⁴⁰ of the postulated CP-symmetry. CP-symmetry states that the laws of physics should
⁵⁴¹ be the same if a particle were to be exchanged with its antiparticle and then if the left
⁵⁴² hand side of a decay were switched with the right hand side.

⁵⁴³ **2.3.5 Reactor Oscillation**

⁵⁴⁴ Many experiments have searched for oscillation of electron anti-neutrinos produced at
⁵⁴⁵ nuclear reactors. Such oscillations give the value of the parameter θ_{13} . The KamLAND
⁵⁴⁶ experiment, started in 2002, has made a high precision observation of reactor neutrino
⁵⁴⁷ oscillation. Neutrinos produced in nuclear reactors have energies similar to solar
⁵⁴⁸ neutrinos, a few MeV. The baselines of these experiments have ranged from tens
⁵⁴⁹ of meters to over 100 km. On 8 March 2012, the Daya Bay team announced a 5.2σ
⁵⁵⁰ discovery that $\theta_{13} \neq 0$.

551 **Chapter 3**

552 **The MicroBooNE Experiment**

553 The purpose of this chapter is to discuss and understand the details of the MicroBooNE
554 detector. A thorough understanding of MicroBooNE and the technology behind liquid
555 argon time projection chambers is important for understanding results as well as
556 understanding how images were made for use in deep learning efforts that will be
557 outlined in later chapters.

558 **3.1 Liquid argon time projection chambers**

559 Liquid Argon Time Projection Chambers (LArTPCs) are an exciting detector technol-
560 ogy that provide excellent imaging and particle identification, and are now being
561 used to study neutrinos. The Time Projection Chamber (TPC) was first invented by
562 Nygren in 1974 [?] and the proposal for a LArTPC for neutrino physics was made
563 by Rubbia [?] in 1977 with the ICARUS collaboration implementing this concept [?].
564 A LArTPC is a three-dimensional imaging detector that uses planes of wires at the
565 edge of an active volume to read out an interaction. When a neutrino interacts with an
566 argon atom, the charged particles that are produced ionize the LAr as they travel away
567 from the interaction. By placing a uniform electric field throughout the LAr volume,
568 the ionization is made to drift towards a set of anode planes, which consist of wires
569 spaced very closely together collecting the ionized charge, which is subsequently read
570 out by electronics connected to the anode wires. The collected ionization creates a
571 spatial image of what happened in the detector on each anode plane. The position
572 resolution of the interaction along the beam direction (perpendicular to drift direction)
573 relies on the wire pitch, while the resolution in drift direction is dependent on the

574 timing resolution of the electronics used and the longitudinal diffusion in the volume.
 575 The drift time of the ionization relative to the time of the original signal allows the
 576 signal to be projected back along the drift coordinate, hence the name LArTPC. Having
 577 very small distances between each wire within an anode plane allows for very fine
 578 granularity and detail to be captured, and having multiple wire planes at different
 579 angles provides independent two-dimensional views that can be combined into a
 580 three dimensional picture of the interaction. Once the charge signal is created on the
 581 anode planes, software analysis packages identify particles in the detector by using
 582 deposited energy on the wires along their track length. The 30 year development of the
 583 ICARUS detector has led to LArTPCs being used as cosmic ray [?], solar neutrino [?]
 584 and accelerator neutrino [?] detectors. The ArgoNeuT experiment at Fermilab was
 585 the first United States based liquid argon neutrino program that has since produced
 586 short-baseline $\nu - Ar$ cross-section measurements in the NUMI beamline [?]. The
 587 MicroBooNE experiment is the second experiment in the US based LArTPC neutrino
 588 program and will be discussed thoroughly in the next sections. The next phases of
 589 the liquid argon neutrino program are under way and are the Fermilab Short Base-
 590 line Neutrino (SBN) program [?] and the Deep Underground Neutrino Experiment
 591 (DUNE) [6]. The SBN program will include three LArTPC detectors, including the
 592 MicroBooNE detector, on the Booster Neutrino Beam (BNB) to do multiple-baseline
 593 oscillation measurements. The detector closest to the beam will be the 40 ton Short
 594 Baseline Neutrino Detector (SBND) [?] at 150 m and the detector furthest is the 600 ton
 595 ICARUS T600 [?] detector positioned at 600 m. The DUNE collaboration will deliver
 596 a 30 GeV neutrino beam 1300 km from Fermilab to a 34 kiloton LArTPC detector
 597 at Homestake, SD. DUNE will study the leptonic CP phase, δ_{cp} , as well as measure
 598 neutrino and antineutrino oscillations.

599 3.2 The MicroBooNE Time Projection Chamber

600 MicroBooNE (Micro Booster Neutrino Experiment) is a 89 ton active volume (180 ton
 601 total mass) LArTPC which is then inserted into a cylindrical cryostat on axis of the
 602 Booster Neutrino Beam (BNB) stationed at Fermilab in Batavia, Illinois. Understanding
 603 LArTPC technology and detector physics is necessary to build a LArTPC the size of
 604 DUNE, and MicroBooNE has made many advances in developing this technology [7]
 605 [8].

MicroBooNE's Time Projection Chamber (TPC) is 10.3 m long (beamline direction), 2.3 m high and 2.5 m wide (which corresponds to the drift distance). The TPC is shown in figure ?? . MicroBooNE is the largest LArTPC currently running in the world [9]. This LArTPC has 3 wire planes: 1 plane that collects the ionization in the wires and is 0° to the vertical with 3456 wires spaced 3 mm apart, and 2 planes where the ionization drifts passed and induces a signal at $\pm 60^\circ$ to the vertical each with 2400 wires also spaced 3 mm apart. Each plane has a spacing also of 3 mm from eachother. The first two planes are the induction planes and the last is the collection. The 270 V/cm electric field of the TPC is created using 64 stainless steel tubes shaped into rectangles around the TPC and held in place by G10 to form a field cage. The cathode is charged at a high voltage of -70 kV and this voltage is stepped down across the field cage tubes using a voltage divider chain with an equivalent resistance of $240\text{ M}\Omega$ between the tubes. The field cage tubes are separated by 4 cm from center to center. The electron drift distance is 2.5 m in the x direction with a drift time of 2.3 ms. Maintaining high charge yield is done by continuously recirculating and purifying the argon. The purity is monitored using MicroBooNE's light collection system. Another use of the light collection system is initial timing and drift coordinate of the interaction.

MicroBooNE's light collection system is a crucial part for 3D reconstruction of all particle interactions in the LArTPC. The initial interaction time, t_0 , and initial drift coordinate, x_0 , are not known from the TPC alone. For beam events, the accelerator clock is used to determine t_0 of the interaction and the x_0 can be inferred using drift time. Non-beam events, however, do not have this capability, which is why scintillation light from an interaction is used. The $\nu - Ar$ interaction produces scintillation light which is collected by photomultiplier tubes (PMTs) which allows the exact time, t_0 of the neutrino interaction to be determined. The scintillation light created propagates within nanoseconds to the light collection system compared to the milliseconds it takes the ionized electrons from the interaction to reach the anode wire planes. Therefore we can precisely know where along the drift direction the particle interaction first took place. The scintillation light is also localized, so combining the PMT information with the wire plane information allows for cosmic background rejection happening outside the beam timing window.

The light collection system is made up of 32 Hamamatsu R5912-02mod cryogenic PMTs with a diameter of 8-inches. The PMTs are located behind the 3 wire anode planes and provides 0.85% photocathode coverage. Each PMT has an acrylic plate mounted in front of it that is coated with a wave-length shifting material called TPB.

641 The acrylic plates take in the scintillation light, at 128 nm, and re-emits it visible
642 wavelengths visible to the PMTs, with a peak at 425 nm.

643 Both the light collection system and the TPC create analog signal that is read out and
644 digitized by the electronics system. The process requires amplification and shaping of
645 the signal which then goes to the data acquisition (DAQ) software for writing of the
646 digitized data to disk. The anode plane wires are connected to detector specific circuit
647 boards (ASICS) that are submerged and operate inside the liquid argon volume. These
648 ASICS send amplified signal to 11 feed-throughs where further amplification of the
649 signal happens outside the cryostat. The signal is received by custom LArTPC readout
650 modules distributed over nine readout crates which do the digitization. The TPC wires
651 are digitized at 16 MHz then downsampled to 2 MHz. The TPC system reads out 4
652 frames of wire signal data per event, 1 frame before a trigger and 2 frames after the
653 triggered frame. The four frames allows for identification of a neutrino interaction as
654 well as cosmic background rejection. The process of digitization is similar for the light
655 collection system. Each PMT signal undergoes a shaping with a 60 ns peaking time
656 for digitization of multiple samples. The digitization occurs at 64 MHz but are not
657 read out continuously during the TPC readout time. Only shaped PMT signal samples
658 above a small threshold are read out and saved. Both the TPC and PMT readouts are
659 initiated via triggers on a separate trigger board located in a warm electronics crate.
660 The timing trigger is created by a timing signal from the BNB accelerator which is
661 shaped and sent to the trigger board. The PMT trigger is generated when the PMT
662 signal multiplicity is greater than 1 and the summed PMT pulse-height is more than 2
663 photo-electrons summed up over all PMT channels. When the trigger board gets both
664 a timing trigger and a PMT trigger in coincidence, at BNB trigger is then generated by
665 the board. This signal is then passed to all readout crates initiating the readout of data.
666 The data is then sent to the DAQ software which then saves the data to disk into one
667 event memory.

668 3.3 MicroBooNE's Physics Goals

669 3.3.1 The low-energy excess

670 The primary goal of the MicroBooNE experiment is to study and investigate the low-
671 energy excess seen in MiniBooNE 3.1. MicroBooNE has the capability of confirming or

denying this excess as electrons or photons due to the detector being in the same beam, having a similar baseline, and lastly the detector being able to clearly distinguish between electrons and photons. LArTPCs use the topology of events as well as energy loss near the vertex to differentiate between single e^- tracks and photon-induced induced pair production $\gamma \rightarrow e^+ e^-$, which wasn't possible in MiniBooNE, a Cherenkov detector. This technique has been shown in the ArogoNeuT detector [?] and a side by side comparison of both event types in a LArTPC can be seen in figure ?? . An excess in electrons would point towards new oscillation physics beyond the standard model, while photons would be within the standard model. MicroBooNE will observe a $4-5\sigma$ signal.

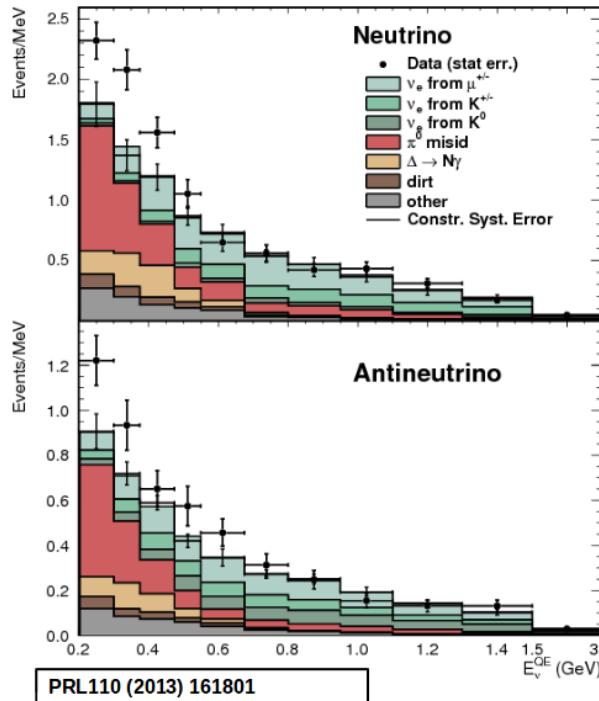


Figure 3.1: Low Energy excess seen in MiniBooNE

3.3.2 Cross sections

MicroBooNE's neutrino cross-section program will be the first $\nu - Ar$ cross-section in the 1 GeV energy range and one of only a few cross-section measurements of $\nu - Ar$ in the world. MicroBooNE is also the first liquid argon detector to collect the highest statistics sample of neutrino interactions. Investigating final-state-interactions in the 1GeV energy range provides information about short range nuclear correlations that affect the interpretations of neutrino oscillation experiment data.

689 One of the cross-section measurements MicroBooNE can make is an inclusive
 690 charged-current cross-section measurement (referred to as CC-inclusive). CC-inclusive
 691 events consist of a neutrino exchanging a W^\pm boson with an argon atom, producing a
 692 charged lepton and any number of other final state particles. In MicroBooNE's case, a
 693 CC-inclusive event will mostly have a defining muon track coming out of the vertex
 694 due to our neutrinos being predominately ν_μ s. A cross-section measurement is the
 695 energy dependent probability of $\nu - Ar$ interaction in the detector. Cross-sections
 696 however are independent of the intensity or focus of the particle beam so they can
 697 be compared among different experiments. A background for a CC-inclusive cross-
 698 section measurement are the neutral-current events that contain a pion. It is possible
 699 to have a neutral current interaction with a $\pi + p$ event signature that looks like a
 700 charged current $\mu + p$ event. Reconstruction tools implemented to date don't efficiently
 701 separate muons from pions. A common way to separate these two particles species is
 702 to implement a track length cut. On average, muons tend to have longer track lengths
 703 in LArTPCs so by requiring that the hypothesized lepton be above a threshold track
 704 length, it is possible to increase signal to background.

705 3.3.3 Liquid argon detector development

706 The last physics goal for the MicroBooNE collaboration is to provide important infor-
 707 mation regarding LArTPC technology. Being the first in large scare LArTPCs in the US,
 708 MicroBooNE will be albe to provide improvements to High Voltage (HV) distribution,
 709 Noise Characterization [?], and Michel Electron Reconstruction [8].

710 3.4 The Booster Neutrino Beam

711 The MicroBooNE detector is stationed at Fermi National Accelerator Laboratory
 712 (FNAL) where it receives neutrinos from both the Booster Neutrino Beam (BNB)
 713 and Neutrinos from the Main Injector (NuMI) beams. MicroBooNE is on-axis for the
 714 BNB and off-axis by 135 mrad for NuMI. For the purpose of this analysis, only data
 715 from the BNB was used. This section will discuss how neutrinos are created using the
 716 BNB. How these neutrinos are produced as well as their flux through the MicroBooNE
 717 detector is necessary for any analysis because of the systematic uncertainties the beam

⁷¹⁸ introduces to a measurement. An aerial view of fermilab as well as the BNB is shown
⁷¹⁹ in figure 3.2

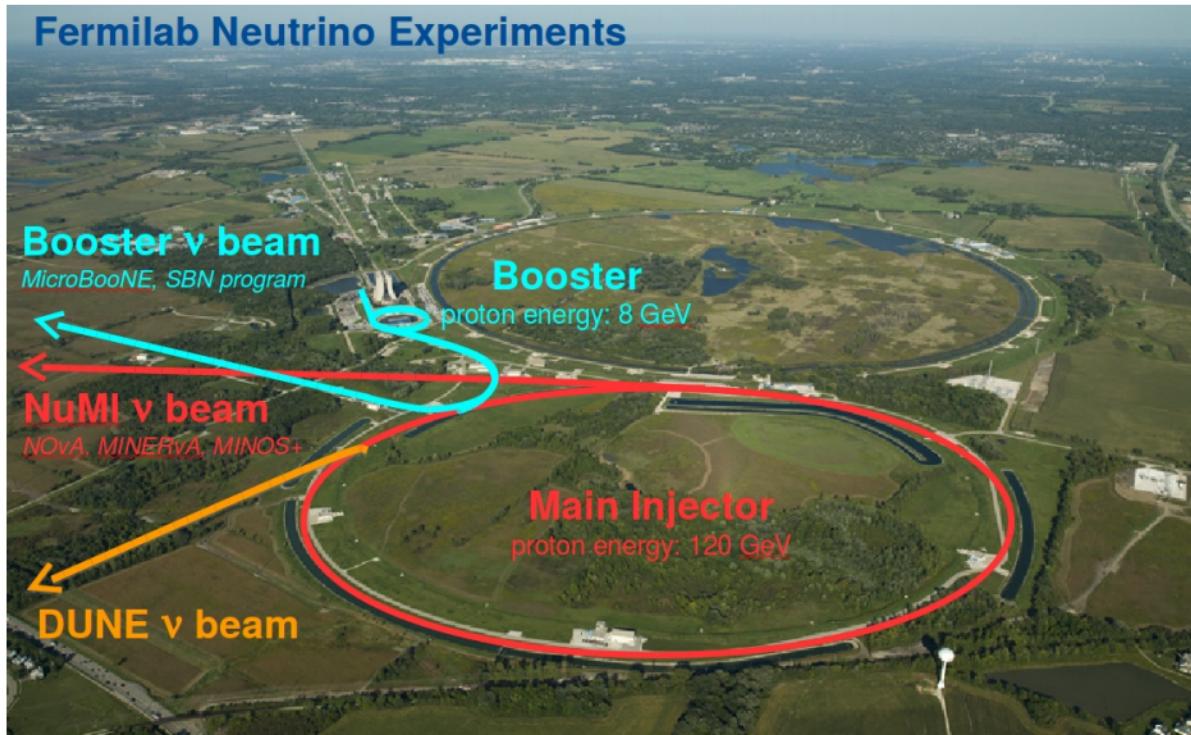


Figure 3.2: Arial view of the Main Injector and the Booster Neutrino Beam at Fermilab

⁷²⁰ 3.4.1 Creating the Booster Neutrino Beam

⁷²¹ The BNB is a very pure ν_μ beam, with only 0.6% contamination from ν_e s. The energy
⁷²² also peaks around 700 MeV which is desired based on the probability of oscillation
⁷²³ equation which depends on the the value of L/E , where L is the distance of the
⁷²⁴ detector from the neutrino beam and E is the energy of the neutrino beam. L/E was
⁷²⁵ chosen to increase the probability of seeing neutrino oscillations in the MiniBooNE
⁷²⁶ Low Energy Excess (LEE) range based on the probability of oscillation equation, which
⁷²⁷ is $P_{\nu_\mu \rightarrow \nu_e}(L, E) = \sin^2 2\theta \sin^2 \left(1.27 \Delta m^2 \frac{L}{E} \right)$. The BNB collides 8.9 GeV/c momentum
⁷²⁸ protons from the FNAL booster synchrotron into a beryllium target which produces a
⁷²⁹ high flux of neutrinos. The protons originate from H^2 gas molecules that are turned
⁷³⁰ into H^- ions by a Cockcroft-Walton generator shown in figure ???. The H^- initially are
⁷³¹ accelerated to 1MeV kinetic energy and are then passed to a linear accelerator using
⁷³² alternating electromagnetic fields to increase their energy to 400MeV. The ions are
⁷³³ stripped of electrons by passing them through a carbon foil. The protons are bunched

⁷³⁴ into beam spills which contain $4 * 10^{12}$ protons in a $1.6 \mu\text{s}$ time window per spill. It's
⁷³⁵ at this point that the protons are directed towards the beryllium target. The amount
⁷³⁶ of protons directed towards the target (POT) is measured by two toroids upstream of
⁷³⁷ the target with an error of 2%. Beam intensity, timing, width, position, and direction
⁷³⁸ are monitored by beam position monitors, multi-wire chamber and resistive monitors.
⁷³⁹ The beryllium target is 71.1 cm long, 1.7 proton interaction lengths, and is 0.51 cm in
⁷⁴⁰ radius. The target is located inside a larger focusing electromagnet called the horn.
⁷⁴¹ The horn is an aluminum alloy pulsed toroidal electromagnet. The pulsed current
⁷⁴² peaks at 170 kA with a time-width of $143 \mu\text{s}$ which coincides with the protons arriving
⁷⁴³ on the target. The current flows from the inner conductor to the outer conductor
⁷⁴⁴ with a maximum magnetic field of 1.5 Tesla. The magnetic field focuses the charged
⁷⁴⁵ secondary particles produced by the p-Be interactions. The direction of current can be
⁷⁴⁶ switched to change the polarity of the secondary particles being focused creating a
⁷⁴⁷ beam of either primarily neutrinos, with positively charged secondary particles, or
⁷⁴⁸ antineutrinos.

⁷⁴⁹ Further down the beamline is a concrete collimator which absorbs particles not
⁷⁵⁰ necessary to the neutrino flux. The collimator is 214 cm long and 30 cm in radius.
⁷⁵¹ After the collimator comes a 45 meter long, 1 meter radius, air-filled cylindrical decay
⁷⁵² region which then ends in a beam-stop made of steel and concrete. The beam-stop
⁷⁵³ contains an array of gas proportional counters to detect muons. The BNB is shown in
⁷⁵⁴ figure 3.3.

⁷⁵⁵ 3.5 Event Reconstruction

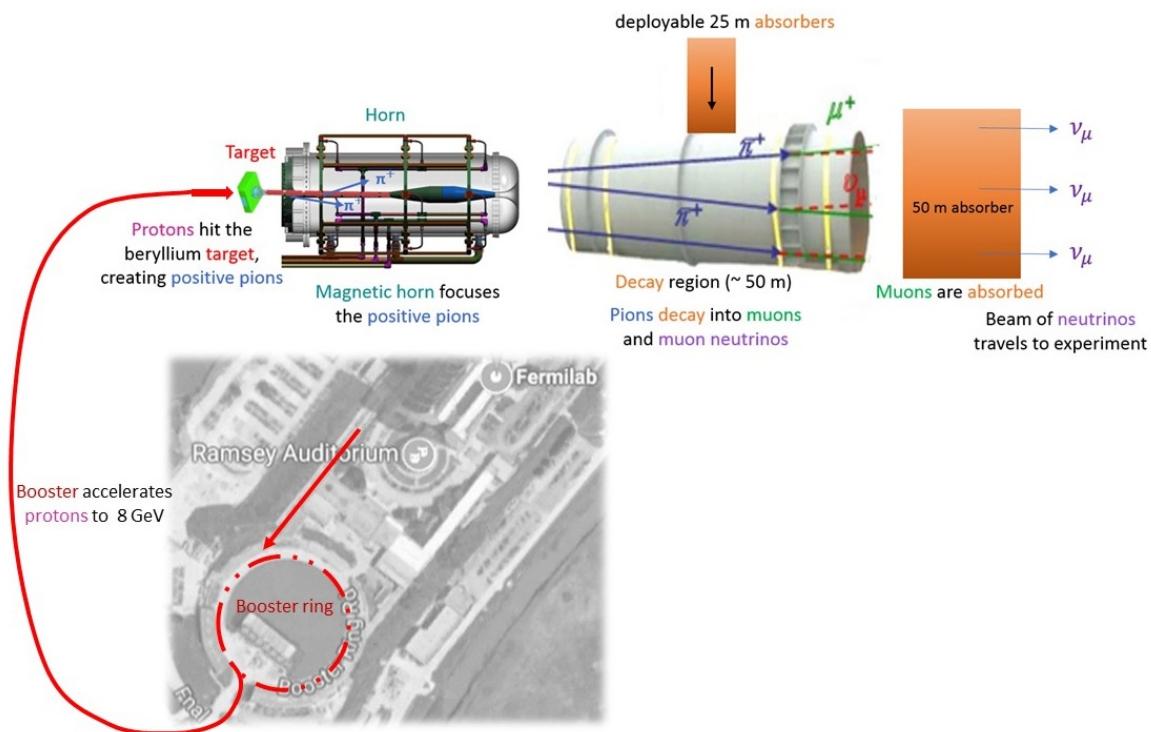


Figure 3.3: Aerial view of the Main Injector and the Booster Neutrino Beam at Fermilab

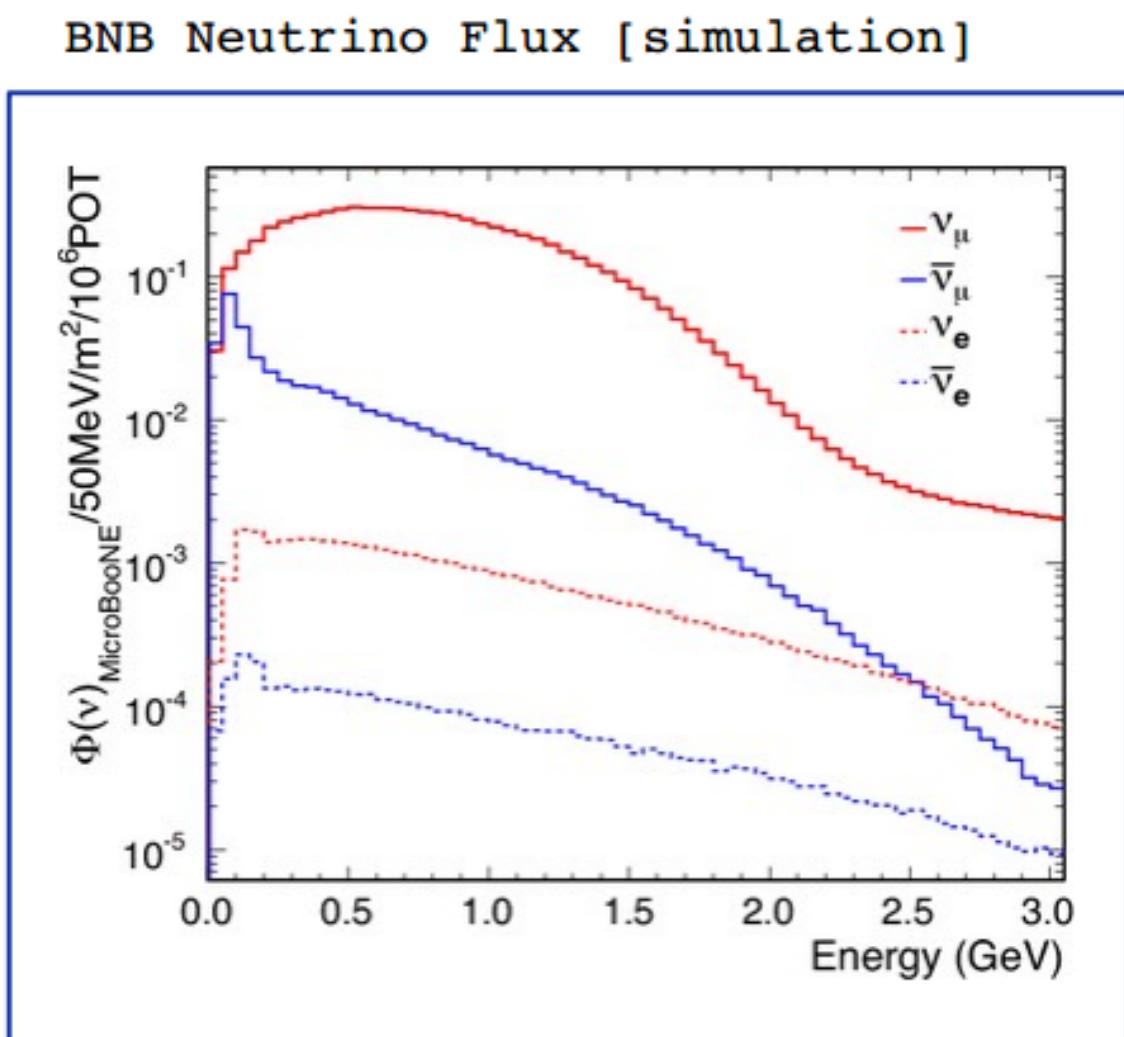


Figure 3.4: Energy spectrum of the Booster Neutrino Beam at Fermi National Laboratories

756 **Chapter 4**

757 **Neutrino Identification: Finding
758 MicroBooNE's first Neutrinos**

759 The goal of the Neutrino Identification analysis was to positively identify BNB neutrino
760 interactions in the MicroBooNE detector collected during the first days of running.
761 Neutrino event candidates were identified in part by using a cut on detected flash of
762 scintillation light during the $1.6 \mu\text{s}$ beam-spill length of the BNB as well as identifying
763 reconstructed object from the TPC that are neutrino like. After this selection, 2D
764 and 3D event displays were used for verification of the selection performance. This
765 selection was targeted to reduce the ratio of neutrino events to cosmic-only events from
766 the initial 1 neutrino to 675 cosmics to a ratio of 1 to 0.5 or better which is equivalent to
767 a background reduction by a factor of 1000 or more. These selected events were used
768 for MicroBooNE's public displays of neutrino interactions. A clearly visible neutrino
769 interaction with an identifiable vertex and at least 2 tracks originating from the vertex
770 was what the analysis focused on. This analysis wasn't optimized for high purity
771 or efficiency, but rather for very distinguishable neutrino interactions that could be
772 identified by the public.

773 **4.1 Flash Finding**

774 Flash finding is the first step used in finding neutrino interactions. This section will
775 detail how optical information is reconstructed as well as analysis scripts and event
776 filters were used.

777 **4.1.1 Flash Reconstruction**

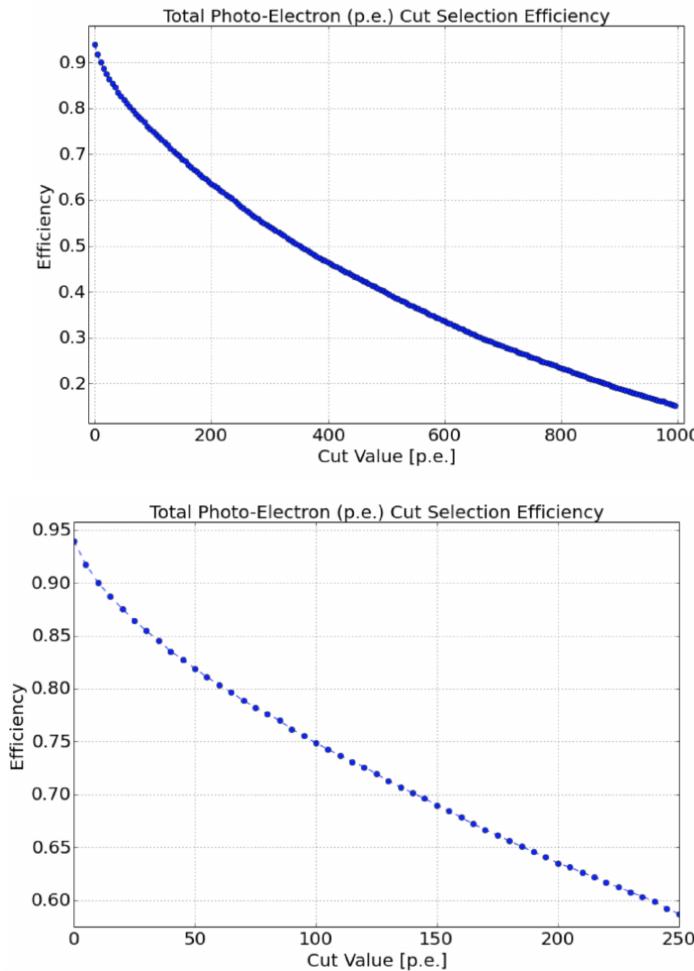
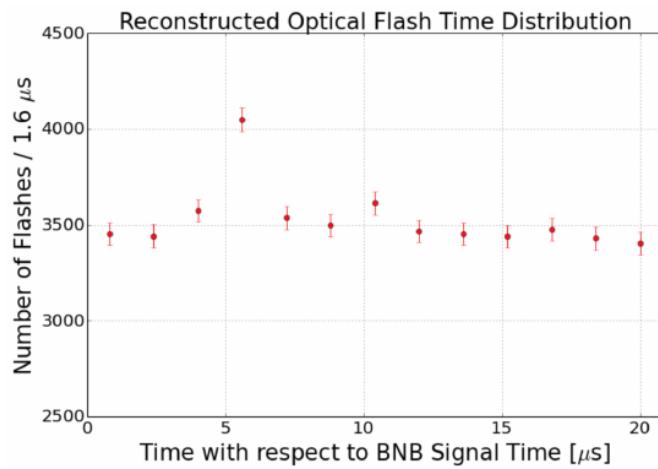


Figure 4.1: Efficiency for selecting beam events as a function of minimum total PE cut for all PE cuts as well as zoomed into interesting region.

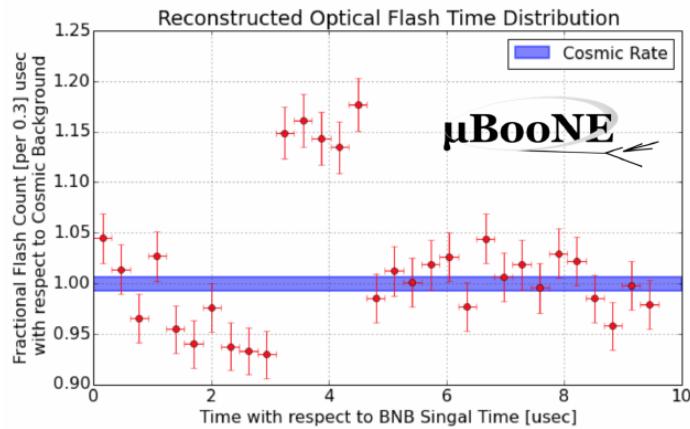
778 A flash is described as a collection of light seen at the same time within the detector.
779 They are then reconstructed by identifying signal from the PMTs above a specific
780 photoelectron (PE) threshold. These signals are called optical hits. Optical hits from
781 all the PMTs are then accumulated into $1\text{ }\mu\text{s}$ bins of time. If a specific bin is above a
782 set PE threshold, then the optical hits that overlap in time are the labeled as the hits
783 from the flash. All flash reconstructed properties like average time and x/y positions
784 are then found via the flash labeled optical hits. The total size of the flash is found by
785 summing up the total number of photoelectrons from all PMTs. Neutrino interactions
786 and cosmic muons will have a larger flash size compared to noise and other low-energy
787 backgrounds, therefore a total PE cut is used to reject these backgrounds. A total PE

⁷⁸⁸ cut of 50 PE was deemed sufficient for this analysis. Figure 4.1 show the total PE
⁷⁸⁹ versus the selection efficency of selecting neutrino beam events.

⁷⁹⁰ **4.1.2 Beam Timing**



(a) Predicted distribution of flash times with respect to trigger time for 1 day of data taking at nominal rate and intensity



(b) Measured distribution of flash times with a 50 PE threshold cut, with respect to trigger time. Shown as a ratio to the expected cosmic rate from off-beam data. A clear excess from neutrinos is visible between 3- 5 μ s after the trigger time.

⁷⁹¹ It is necessary to get the specific time from flashes if one uses flashes to filter out
⁷⁹² neutrino interactions coincident with the neutrino beam spill period and background.
⁷⁹³ Before a filter can be applied, an understanding of the timing of the trigger and PMT

794 readout with respect to the arrival of neutrinos from the BNB. To do this, a $1.6\ \mu\text{s}$
795 window near the expected beamtime was created and verified by finding that the
796 number of flashes was significantly above the cosmic-ray background flashes. Beam
797 data during the first week of running, October 16th 2016 through October 22nd 2016
798 and were used for a timing measurement. The total POT uses corresponds to roughly
799 24 hours of data taking at nominal intensity ($4\times 10^{12}\ \text{ppp}$) and a 5 Hz repetition rate.
800 Figure 4.2a shows size of the expected neutrino signal in time using Monte Carlo
801 predictions and figure 4.2b shows the neutrino signal in data. The intensity in data is
802 lower, however there can still be seen a significant excess above data.

803 4.1.3 Event Rates

804 Applying a 50 PE threshold cut inside a $1.6\ \mu\text{s}$ window reduces the cosmic-ray passing
805 rate to 0.8%. With a 5 Hz beam rate, this corresponds to 135 cosmics passing per
806 hour. The neutrino passing rate for this filter is about 22 events per hour. To further
807 increase the neutrino to cosmic ratio, TPC topology cuts were implemented and will
808 be discussed in the following section.

809 4.2 TPC Topology Selection

810 In order to further reduce the background of cosmic events, two independent selection
811 streams using TPC wire data reconstruction was implemented. The first using 2D
812 reconstructed clusters, and the second using 3D reconstructed tracks. Both streams
813 look for neutrino interactions in the active TPC volume which are identifiable by two
814 or more tracks originating from the same vertex.

815 Both 2D and 3D channels were optimized using monte carlo simulation which
816 used a 128 kV cathode voltage. Passing rates were calculated using a 0.008 efficiency
817 factor for cosmic events passing to simulate the flash finding described in section 4.1.
818 This efficiency factor was an overestimation and was just used to get a general feel of
819 what signal and background rates we would actually see in data.

820 4.2.1 Cosmic Tagging

821 The first step in TPC selection was based on the geometry of cosmic tracks in an event.
822 The cosmic ray muon geometry tagger runs on 3D tracks and assigns a score to each
823 reconstructed track on the likeliness of the track originating from a cosmic. The cosmic
824 scores are detailed below:

- 825 • 1: The track is tagged as entering or exiting the TPC
- 826 • 0.95: The track is a delta ray associated with a tagged track
- 827 • 0.5: The track is either entering or exiting, but not both
- 828 • 0.4: The track is entering or exiting through the Z boundary
- 829 • 0: The track isn't tagged

830 Clusters are assigned either a 0 or 1, 1 being a cosmic. In simulation, 90% of cosmics
831 are tagged as cosmics. These tracks are no longer considered when looking for a
832 neutrino topology. Requiring that the tracks be contained in turn affects the neutrino
833 efficiency by 20%. The algorithm checks that each track is contained within a boundary
834 region of 10 cm from all sides of the TPC. This boundary region was optimized via
835 handscanning of experimental data.

836 As can be expected, cosmic tagging is more efficient in the 3D channel (tracks) than
837 the 2D channel (clusters) because the reconstructed tracks can use the full 3D position
838 information of the entering and exiting points while the 2D channel mainly use the
839 reconstructed x position of the cluster which is associated to timing.

840 Cosmic tagging uses timing information to reject tracks and clusters that are outside
841 of drift window. The drift window for 128 kV is 1.6 μ s while for 70 kV, the actual
842 voltage MicroBooNE is running at, is 2.3 μ s. Due to this variation between simulation
843 and data, we expect to see $2.3/1.6 = 1.44$ times more cosmic induced tracks or clusters
844 in the drift window.

845 4.2.2 2D Cluster Selection

846 This selection was spearheaded by myself and Katherine Woodruff. After looking at
847 experimental cosmics data, 2D clustering performs well, while 3D track reconstruction
848 is affected by more variations in simulation, for example noise filters. This was the

849 motivation for having a selection only on 2D clusters in the collection (Y) plane. As
850 stated previously, the goal of this analysis was to find identifiable neutrino interactions
851 for use in public event displays, in future analyses, the 3D track reconstruction has
852 been modified to further increase the tracking efficiency and has more information
853 that just the clusters. For this analysis, however, 2D cluster information was sufficient
854 enough for neutrino selection.

855 **Primary Cuts**

856 The first cuts were used to select which clusters to consider. First the clusters must
857 have at least ten hits on the collection plane and have a cosmic tagging score < 0.4.
858 Only events that have at least two clusters that satisfy these primary cuts continue on.

859 After the initial cosmic tagging is applied, the following cuts are used to further
860 separate identifiable neutrinos from background cosmics.

861 The next cut was to remove long, vertical clusters. This was applied after seeing
862 that most cosmic induced clusters passing were long with high angles, while neutrino
863 induced clusters were mainly forward going. We required a good cluster to either
864 have a projected start angle less than 30 degrees from the z axis or be less than 200
865 wires long. The length cut was added to make sure we don't cut any short high angle
866 clusters that can correspond with a proton, or other highly ionizing particle associated
867 with a long muon cluster. The 200 wire cut roughly equates to 0.6 m in the z direction,
868 with a 3 mm wire pitch. Also, the projected angle is defined by $\tan \alpha = \Delta T / \Delta W$ where
869 T is the time ticks and W is the wires.

870 The last cut requires the clusters to be either 30 time ticks or 30 wires. This cut was
871 applied to reduce small delta rays associated with a cosmic without removing proton
872 clusters associated with a long muon cluster, which saves ideal neutrino events that
873 have both a long minimum ionizing muon like cluster and a short highly ionizing
874 proton like cluster.

875 **Secondary Cuts**

876 The secondary cuts look to match long, low-angle clusters with short, high-charge
877 clusters. Only clusters that have passed previous cuts are used. First clusters with
878 length greater than 100 wires are chosen, which is approximately 0.3 m in the z

Cluster set	No Cuts	Primary Cuts	Secondary Cuts
Neutrinos only	570	303	32
Cosmics only (no flash)	308,016	291,879	602
Cosmics only (w/ flash)	2464	2335	5
Neutrinos/Cosmics	0.23	0.13	6.4

Table 4.1: Passing rates for 2D cluster cuts for neutrino on MC set and a cosmic only MC set. First column shows event rates with no cuts applied to both sets. Columns two and three show event rates after primary and secondary cuts are applied. Line three shows the second line scaled with the flash finding factor of 0.008. All events are normalized to per day assuming we are running at 5 Hz.

879 direction. Then we search for any cluster that is within approximately 3 cm (10 wires
 880 and 30 time ticks) away from the low-z end of the long cluster. This cluster must also
 881 be shorter than the first. In our reconstruction, the start and end point of a cluster can
 882 be swapped so both ends of the short cluster are compared to the long cluster.

883 Now that there is a vertex match, cuts based on charge and projected opening angle
 884 are implemented. We require the short cluster to have a higher start charge than the
 885 long cluster or the long cluster be longer than 500 wires. Start charge is defined as
 886 the charge on the first wire in ADC counts. The projected opening angle must also
 887 be between 11 and 90 degrees. This last cut is intended to remove clusters that are
 888 entirely overlapping or are part of the same long track. The resulting neutrino/cosmic
 889 event rate per day is shown in table 4.1. Figures 4.3 and 4.4 shows the percentages of
 890 clusters that pass each primary and secondary cuts.

891 4.2.3 3D Tracks and vertices Selection

892 The neutrino selection for the 3D channel was based on a reconstructed vertex and
 893 two tracks. All vertices and tracks were looped over that had a cosmic tag score < 0.4
 894 and the distances below were calculated:

- 895 • d : distance between the start points of the two tracks.
- 896 • d_1 : distance between vertex and start of track 1.
- 897 • d_2 : distance between vertex and start of track 2.

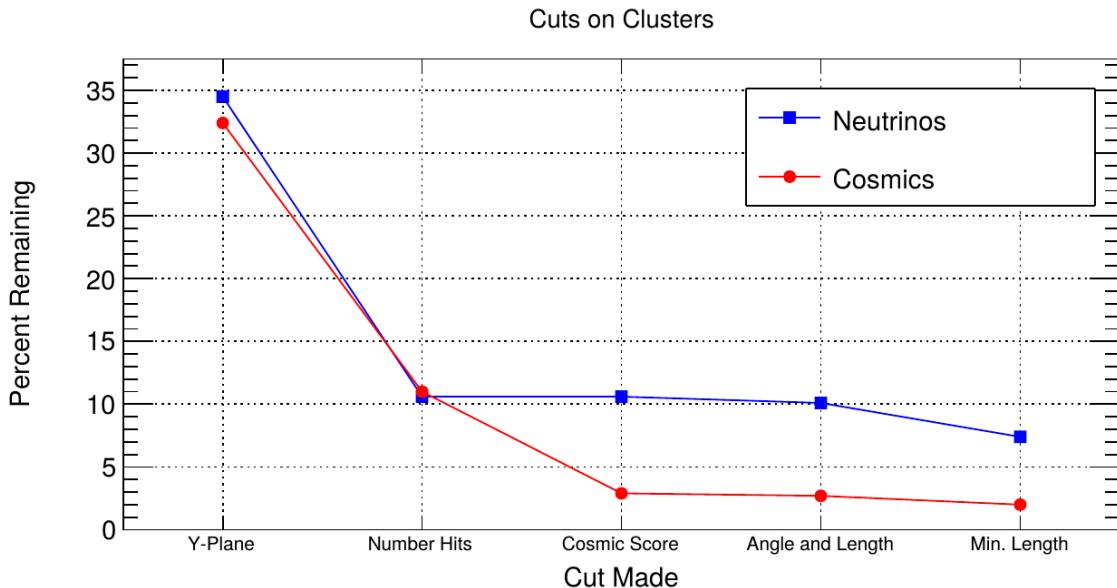


Figure 4.3: Percent of good clusters remaining for neutrinos and cosmics after the primary cuts were applied. This is relative to total number of initial clusters.

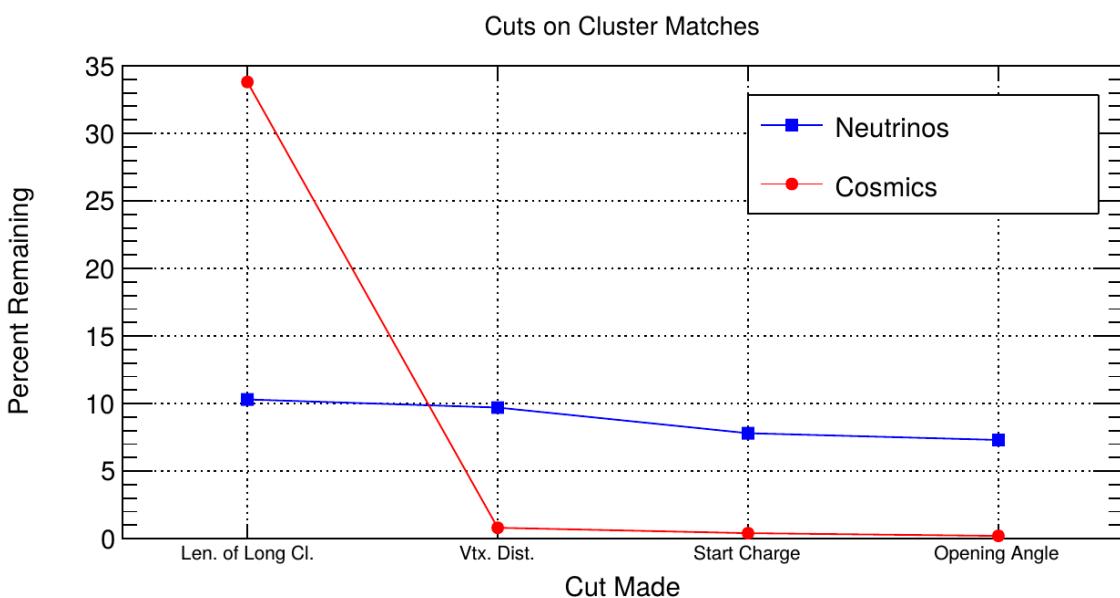


Figure 4.4: Percent matched cluster pairs remaining for neutrinos and cosmics after secondary cuts applied. This is relative to the number of events that contain clusters which pass the primary cuts.

898 The maximum distance of all three is then selected as the important characteristic per
899 trio. The best trio is the one that has the smallest maximum distance. The $\min(\max_d)$
900 for all trios in an event were plotted for BNB neutrino events and for cosmics to
901 find the best cut value for each tracking algorithm. The distribution of $\min(\max_{d,i})$
902 is smaller for neutrinos than for cosmics. The cut values for different tracking and
903 clustering algorithms are shown below. These cut values were chosen to minimize the
904 cosmic background to 20%.

- 905 • trackkalmanhit with cccluster $\min(\max_{d,i}) < 3$ cm.
906 • trackkalmanhit with pandoraNu $\min(\max_{d,i}) < 4.5$ cm.
907 • pandoraNu with cccluster $\min(\max_{d,i}) < 5$ cm.

908 4.2.4 TPC Updates

909 After doing a visual hand-scanning of the first beam data processed with the filters
910 detailed above, the events passing had a larger contamination of background than
911 expected. This was mainly in part due to the reconstruction performing better on
912 simulation than on data. Due to this, additional cuts on both streams needed to be
913 implemented in order to increase signal/background ratio. These cuts were added on
914 top of the filters described above and further reduce the event count.

915 2D Filter Updates

916 The main background observed in the 2D filter were Michel events, where the muon
917 and electron formed two connected clusters. These events were rejected by comparing
918 the start and end charge deposition of the long cluster (i.e muon particle). The start
919 charge deposition must be less than the end charge deposition. This cut is implemented
920 because muons have a higher ionizaiton loss at the end.

921 3D Filter Updates

922 It was seen that cosmic tracks can often originate or end at the same point, therefore
923 faking a signal. Cosmic tracks, however, are mostly vertical. By requiring the angle
924 of the longer track have a cosine greater than 0.85 with respect to the z-axis as well

925 as requiring the longer track to have a length greater than 10 cm, we can reduce this
926 background.

927 **4.3 Conclusion**

928 After proccesing these filters in parallel, it was shown that the 3D filter had a higher
929 purity than the 2D filter because of the higher cosmic rejection being used due to 3D
930 reconstruction. The 2D filter is blind to track entering/exiting from the top or bottom
931 of the TPC. Although the 3D filter had a higher purity, the 2D filter was still able to
932 find identifiable events in data that were used as public event displays. A sample of
933 event displays are shown in figures ?? and ??.

⁹³⁴ **Chapter 5**

⁹³⁵ **CC-Inclusive Cross Section Selection
Filter**

⁹³⁷ The CC-Inclusive cross-section selection I and selection I modified filters used in this
⁹³⁸ analysis will be described in the following sections below. These filters are an expan-
⁹³⁹ sion of the Neutrino ID filter. The work done in this thesis was to further improve these
⁹⁴⁰ selections by increasing both efficiency and purity as well as increasing acceptance
⁹⁴¹ without further affecting the kinematic distributions of the selected neutrino events.

⁹⁴² MicroBooNE requires fully automated event reconstruction and selection algo-
⁹⁴³ rithms for use in the many physics measurements being worked on to date due to
⁹⁴⁴ the large data rate MicroBooNE receives. Being able to automatically pluck out the
⁹⁴⁵ neutrino interaction among a sea of cosmics proved to be challenging but was accom-
⁹⁴⁶ plished. MicroBooNE has developed two complementary and preliminary selection
⁹⁴⁷ algorithms to select charged-current $\nu_\mu - Ar$ interactions. Both are fully automated
⁹⁴⁸ and cut based. The results of this thesis will focus on selection I and selection I modi-
⁹⁴⁹ fied and will focus on further improving these algorithms using Convolutional Neural
⁹⁵⁰ Network (CNN) implementations. These selections identify the muon from a neutrino
⁹⁵¹ interaction without biasing towards track multiplicity. To combat cosmic and neutral
⁹⁵² current background, the analysis is strongly biased towards forward-going long tracks
⁹⁵³ which are contained. This limits phase space and reduces acceptance.

954 **5.1 Data and MC Processing Chain**

955 The data used for this analysis were based on hardware and software triggers. Events
956 used came from the *BNB_INCLUSIVE* and *EXT_BNB_INCLUSIVE* streams and were
957 used for signal and background. The *BNB_INCLUSIVE* stream is chosen by requiring
958 that the hardware trigger bit is fired and that the event passed an optical software
959 trigger within a BNB spill timing window. The *EXT_BNB_INCLUSIVE* stream requires
960 the EXT hardware trigger to fire as well as pass the same optical software trigger
961 within a BNB spill size timing window similar to the *BNB_INCLUSIVE*.

962 The two MC samples used in this analysis and for determining selection efficiencies
963 and purities were GENIE BNB neutrino interactions with CORSIKA cosmic ray overlay
964 within the readout window and inTime CORSIKA cosmic rays. The MC samples
965 generated used *uboonecode v04_36_00* and are based on the following packages:

- 966 • larsoft v04_36_00
- 967 • GEANT v04_09_06_p04d
- 968 • GENIE v02_08_06d
- 969 • GENIE xsec v02_08_06a
- 970 • pandora v02_03_0a
- 971 • CORSIKA v07_4003

972 Both data and MC samples were processed using the same reconstruction release,
973 *uboonecode v05_08_00* and the fcl files used for reconstruction are listed below:

- 974 • MC fcl files
 - 975 – reco_uboone_mcc7_driver_stage1.fcl
 - 976 – reco_uboone_mcc7_driver_stage2.fcl
- 977 • Data fcl files
 - 978 • reco_uboone_data_Feb2016_driver_stage1.fcl
 - 979 • reco_uboone_data_Feb2016_driver_stage2.fcl

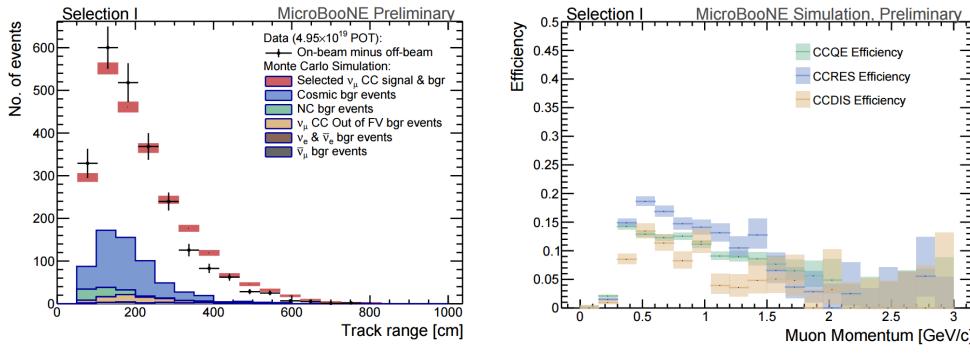
980 On top of the hardware and software triggers, the data also had to pass more
981 criteria to be identified as part of the good run list. The criteria is detailed below.

- **Detector conditions:** the detector has to be in a good operating condition. The detector conditions are read from the slow monitoring database and are required to be within the alarm thresholds. The variables of interest for events passing the good run list criteria include DAQ, PMT, HV, Drift HV, wire bias, electron lifetime and detector power. These conditions need to be met on a run-by-run basis in order to pass the selection.
- **Data quality:** normal and stable behavior for basic reconstruction quantities. These reconstruction variables include average number of tracks, hits, and flashes in each event, the average length of tracks, the average amplitude and area of hits, the average PE and the average spread of each one of these quantities.
- **Beam Conditions:** the BNB must be on and stable and the POT per spill needs to be above the intensity threshold. Beam quality conditions include checking the fraction of proton beam interacting within the target, the horn current, and the intensity of protons per spill. The final sample is 5×10^{19} and a per-spill intensity of 4×10^{12}
- **Run processed:** the full run must be processed completely without missing subruns or crashes in the data processing.

The selection begins with a cut that requires an optical flash greater than 50 photo electrons (PE) in the $1.6 \mu\text{s}$ beam window. Next, two or more 3D reconstructed tracks must be within 5 cm from a 3D reconstructed vertex. The most forward going track vertex-track association is then selected for further cuts. The vertex from the chosen association must be in the fiducial volume, and the longest track from this association must be matched to a flash 80 cm in z. Lastly the longest track must be contained and longer than 75 cm.

The efficiency and purity are used as performance values of selection I. Efficiency is described as the number of selected true ν_μ CC events divided by the number of expected true ν_μ CC events. The purity is described as the number of selected true ν_μ CC events divided by the sum of itself and all the backgrounds. The efficiency of selection I is 12% and the purity is 39.7%. The poster related to this proceedings will focus on the last cut which requires the longest track to be longer than 75 cm. This cut has a passing rate of 30% w.r.t the previous cut and is implemented in part to separate charged-current events from neutral-current events that mimic our signal. Implementing a CNN for $\mu - \pi$ separation picks out differences in these two particles that are track range independent therefore eliminating the need for the 75 cm track

length cut and increase efficiency and passing rate at low muon momentum. Figure 5.1a shows the track distribution of selection I and the lack of data below the 75 cm track length cut. Figure 5.1b shows the efficiency of selection I as a function of muon momentum.



(a) Track range distribution of selection I

(b) Selection efficiency as a function of the true muon momentum

Figure 5.1: 5.1a Track range distribution for selection I. The track range is defined as the 3D distance between the start and end of the muon candidate track. No data is shown below 75 cm due to the track length cut described previously. 5.1b Efficiency of the selected events by process quasi-elastic (QE), resonant (RES), and deep-inelastic (DIS). Statistical uncertainty is shown in the bands and the distributions are a function of true muon momentum. The rise of the efficiency between 0 GeV and 0.5 GeV is due to the minimum track length cut and the decreasing efficiency for higher momentum tracks is caused by the containment requirement.

1019

1020 5.2 The importance of μ/π separation

1021 **Chapter 6**

1022 **Background on Convolutional Neural
1023 Networks**

1024 Convolutional neural networks (CNNs) have been one of the most influential inno-
1025 vations in the field of computer vision. Neural networks became popular in 2012
1026 when Alex Krizhevsky used them to win that year's ImageNet competition [?] by
1027 dropping the error from 26% to 15%. Since then, many companies are using deep
1028 learning including Facebook's tagging algorithms, Google for their photo search and
1029 Amazon for product recommendations. For the purpose of this thesis CNNs were
1030 used for image classification, specifically, images of varying particles created using
1031 LArTPC data.

1032 **6.1 Image Classification**

1033 Image classification is the process of inputting an image into the CNN and receiving a
1034 probability of classes that best describes what is happening in the image. As humans,
1035 image classification is something that is learned at a very young age and is easy to
1036 do without much effort. This is also apparent when hand-scanning LArTPC images.
1037 After learning what a neutrino event looks like in MicroBooNE, it is relatively easy
1038 to recognize simple neutrino events from cosmic ray background as well as highly
1039 ionizing particles like protons from minimum ionizing particles like muons. The very
1040 detailed images LArTPC detectors output are prime candidates for input images into
1041 a CNN. CNNs mimic a human's ability to classify objects by creating an architecture
1042 that can learn differences between all the images it's given as well as figure out the

1043 unique features that make up each object. CNNs are modeled after the visual cortex.
1044 Hubel and Wiesel [?] found that there are small regions of neuronal cells in the brain
1045 that respond to specific regions of the visual field. They saw that some neurons fired
1046 when exposed to vertical edges while others fired when shown horizontal or diagonal
1047 edges. They also saw that these neurons were organized in columns. The idea of
1048 specific neurons inside of the brain firing to specific characteristics is the basis behind
1049 CNNs.

1050 **6.2 CNN Structure**

1051 When used for image recognition, convolutional neural networks consist of multiple
1052 layers that extract different information on small portions of the input image. How
1053 many layers is tunable to increase the accuracy. The output of these collections are
1054 then tiled so that they overlap to gain a better representation of the original image
1055 and allow for translation. The first of these layers is always a convolution layer. To
1056 the CNN, an image is an array of pixel values. For a RGB color image with width
1057 and height equal to 32x32 the corresponding array is 32x32x3. Filters, also known as
1058 neurons, of any size set by the user is then convolved with the receptive field of the
1059 image. If the filter is 5x5, the receptive field will be a patch of 5x5 on the input image.
1060 The filter is also an array of numbers called weights. The convolution of the filter and
1061 image are matrix multiplications of the weights and the pixel values. By stepping the
1062 receptive field by 1 unit, for an input image of 32x32x3 and a filter of 5x5x3 you'd get
1063 an output array of 28x28x1. This output array is called an activation map or feature
1064 map. The use of more filters preserves the spatial dimensions better. The filters can
1065 be described as feature identifiers. Examples of features in an image consist of edges,
1066 curves, and changes in colors. The first filters in a CNN will primarily be straight
1067 line and curve feature identifiers. An example of a curve filter is shown in figure 6.2.
1068 When a curve in the same concavity is found in the input image, the corresponding
1069 pixel in the output feature map will be activated. Going back to our example of a
1070 32x32 input image and a 5x5 filter, if there were to be a curve in the top left corner
1071 of the input image, our output feature map would have a high pixel value in the top
1072 left. Therefore, feature maps tell us where a specific feature is located in the original
1073 image. Figure ?? shows a visualization of filters found in the first layers of many CNN
1074 architectures. These filters in the first layer convolve around the image and activate
1075 when the specific feature it is looking for is in the receptive field.

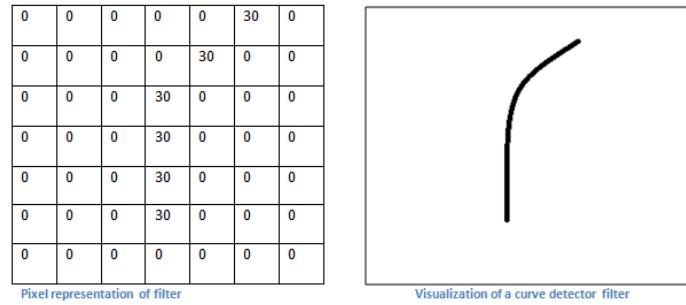


Figure 6.1: Pixel representation and visualization of a curve detector filter. As you can see, in the pixel representation, the weights of this filter are greater along a curve we are trying to find in the input image

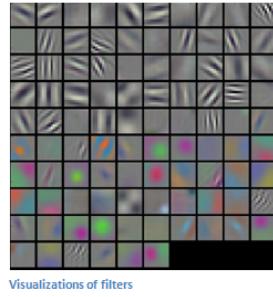


Figure 6.2: Visualization of filters found in first layer of a CNN.

1076 In figure 6.3 you can see how an edge detection filter is used to save only necessary
 1077 information for recognizing different types of clothes. You can also see by having
 1078 multiple filters you can get more detail or less detail from an image which can then
 1079 simplify or complicate the object recognition task. Being able to distinguish between a
 1080 shirt or a leg garment is as much information you want, having a filter that extracts
 1081 outline edge or shape information would be all that you need. But if instead you
 1082 wanted to distinguish between a formal cocktail dress or a summer dress, more
 1083 information would need to be saved equating to many more filters for one image.
 1084 Rather than trying to come up with how many filters and what features are important
 1085 for detection, CNNs do this automatically. CNNs take input parameters, called
 1086 hyperparameters, for example number of layers, number of filters per layers, number
 1087 of weights per filter, and uses these to create the output feature maps. The layers build
 1088 upon each-other, for example if we were creating a CNN for facial recognition the
 1089 convolutional layers will start learning feature combinations off of the previous layers.
 1090 The low level features like edges, gradients, and corners of the first layers become high
 1091 level features like eyes, noses, and hairs. This process is visualized in figure 6.4



Figure 6.3: Applying a feature mask over a set of fashion items to extract necessary information for auto-encoding. Unnecessary information for example color or brand emblems are not saved. This feature map is an edge detection mask that leaves only shape information which helps to distinguish between different types of clothes.

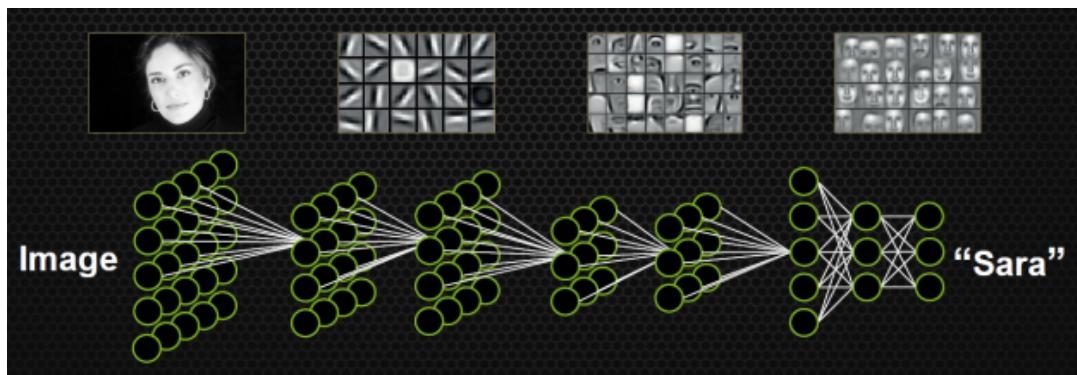


Figure 6.4: Pictorial Representation of Convolutional Neural Networks as well as a visual representation on CNN's complexity of layer feature extraction

1092 There are other layers in a CNN architecture that will not be covered in the scope
 1093 of this thesis but in a general sense, these layers are interspersed between convolution
 1094 layers to preserve dimensionality and control overfitting of the network. The last layer
 1095 is called a fully connected layer and it's job is to output an N dimensional vector where
 1096 N is the number of classes the network has been trained on. Each number in this vector
 1097 represents the probability that the input image is a certain class. Fully connected layers
 1098 use the feature maps of the high level features to compute the products between the
 1099 weights of the previous layer to get the probabilities of each class. These weights are
 1100 then adjusted through the training process using backpropagation.

1101 6.2.1 Backpropagation

1102 A CNN at it's onset has weights that are randomized. The filters themselves don't
 1103 know how to pull out identifying information per class. For a neural network to learn,
 1104 it must be trained on a training set that is labeled. Backpropagation has four seperate
 1105 steps: foward pass, loss function, backward pass and updating weights. In the forward
 1106 pass, a training image is passed through the whole network. All of our weights at this
 1107 time are randomly initialized so the output for the first image will have no preference
 1108 to a specific class. A common loss function is mean squared error (MSE):

$$E_{total} = \sum \frac{1}{2} (actualclass - predictedclass)^2 \quad (6.1)$$

1109 If we assume that the MSE is the loss of our CNN, the goal would be that our
 1110 predicted label (output of CNN) is the same as our training label. To do this, we need
 1111 to minimize the loss function. To do this, it is necessary to find out which weights most
 1112 directly affect the loss of the network i.e $\frac{dL}{dW}$ where L is our loss function and W are
 1113 the weights of a specific layer. The next step is the backward pass which determines
 1114 which weights contribute the most to the loss and finds ways to adjust these weights
 1115 so that the loss decreases. After the derivative is computed, the last step updates the
 1116 weights in the opposite direction of the gradient.

$$w = w_i - \eta \frac{dL}{dW} \quad (6.2)$$

$$w = \text{Weight} \quad (6.3)$$

$$w_i = \text{Initial Weight} \quad (6.4)$$

$$\eta = \text{Learning Rate} \quad (6.5)$$

1117 The learning rate is a parameter given to the CNN and it describes the steps the
 1118 network takes to update the weights. Higher learning rate equals large steps and a
 1119 lower training time, but a learning rate that is too large can mean the CNN never
 1120 converges.

1121 Going through backpropagation consists of one training iteration. Once the net-
 1122 work completes a specific number of iterations, another parameter given, and runs
 1123 over all training images that are split up into batches, the process is considered com-
 1124 plete. User input parameters, called hyperparameters, help the network converge to

optimal weights for each layer. Batch size, learning rate, and training iteration are just some of the user input hyperparameters that help. Lastly, to check if the network has learned, a different set of labeled images are fed to the CNN iteratively through the training process to see how well it's learning. This process is especially important to make sure the network architecture isn't being affected by overfitting (memorizing training input rather than learning).

6.3 Choosing Hyperparameters

Convolutional neural networks are a relatively new tool in computer vision. Choosing hyperparameters for your specific dataset is a non-trivial task. Hyperparameters can range from the amount of layers and filters per layer in an CNN architecture to the stride the receptive field of a filter takes, not to mention training hyperparameters such as learning rate and batch size described above. They're ways to optimize these hyperparameters via hyperparameter optimization using Bayesian Optimization [?] but as you can imagine, optimizing an CNN architecture from scratch can be very computationally intensive. For the purpose of this thesis, two well known CNN architectures were used, AlexNet [?], which won the ImageNet Large-Scale Visual Recognition Challenge (ILSVRC) in 2012 and therefore bringing awareness of CNNs, and GoogleNet [?], which won the ILSVRC in 2014, giving rise to deep networks. Both AlexNet and GoogleNet architectures were used to train on LArTPC images and their low level filter weights. Higher level filter weights were randomly initialized before training so the network can learn high level features of LArTPC image classes. The AlexNet architecture is shown in figure 6.5 and the GoogleNet architecture is shown in figure 6.6

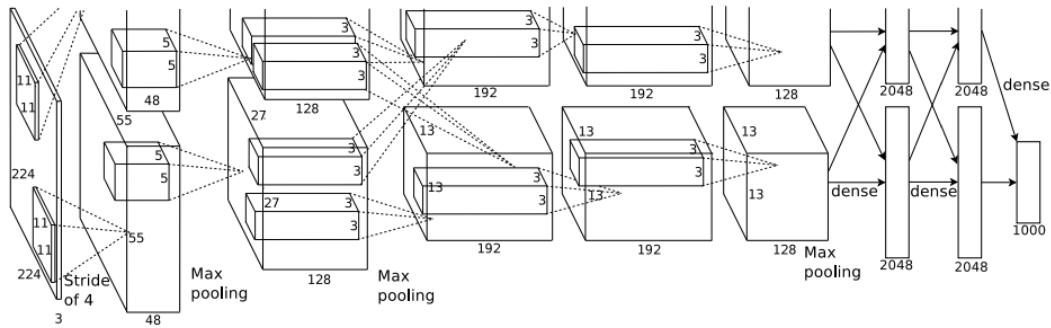


Figure 6.5: Pictoral representation of the AlexNet model. The AlexNet model consists of 5 convolution layers and 3 fully connected layers.

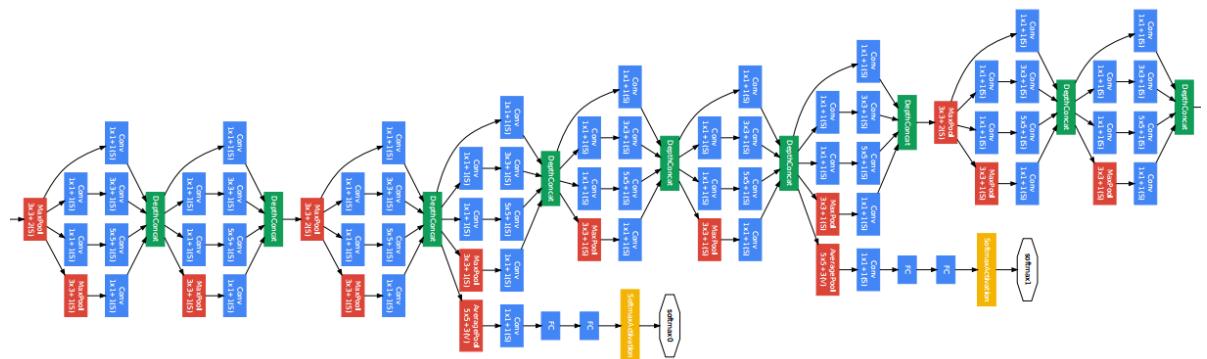


Figure 6.6: Pictoral representation of the GoogleNet model. The GoogleNet model consists of 22 layers. The model implements 9 Inception modules which performs covolution and pooling in parallel and strays away from the basis that CNN layers need to be stacked up sequentially. The GoogleNet model also doesn't use fully connected layers, instead it uses average pooling which greatly reduces the amount of parameters. GoogleNet has 12x fewer parameters than AlexNet.

₁₁₄₈ **Chapter 7**

₁₁₄₉ **Training Convolutional Neural**
₁₁₅₀ **Networks on particles WORKING**
₁₁₅₁ **TITLE**

₁₁₅₂ Three Convolutional Neural Networks CNNs were trained throughout this analysis.
₁₁₅₃ There are differences to each CNN and will be described fully in the next sections but
₁₁₅₄ the main difference are the amount of particle images used for training and validation.
₁₁₅₅ CNN1075 used 1,075 muons and 10,75 pions for training and the same amount of each
₁₁₅₆ particle for validation. CNN10000 used 10,000 muons and 10,000 pions split in half
₁₁₅₇ for testing and training. Lastly CNN100000 had muons, pions, protons, electrons,
₁₁₅₈ and gammas in its training and validation set. Each particle had 20,000 images and
₁₁₅₉ training and validation was split 90% training, 10% validation. This chapter will also
₁₁₆₀ describe the different hardware frameworks used for training beginning on a CPU
₁₁₆₁ and ending on a GPU cluster.

₁₁₆₂ **7.1 Hardware Frameworks used for Training**

₁₁₆₃ **7.1.1 Syracuse CPU Machine setup**

₁₁₆₄ **7.1.2 Syracuse University GPU Cluster Setup**

₁₁₆₅ **7.2 Convolutional Neural Network Training**

₁₁₆₆ **7.2.1 Image Making Scheme**

₁₁₆₇ **Images used for Traing/Validation of Convolutional Neural Networks**

₁₁₆₈ **add image making for CNN1075** The μ/π image dataset used to train and validate
₁₁₆₉ the CNN10000 was created using single generated isotropic muons and pions from
₁₁₇₀ 0-2 GeV energ range. 10,000 muons and 10,000 pions were used for training and
₁₁₇₁ testing split 50%. The images were created based on wire number and time tick in the
₁₁₇₂ collection plane. Uboonecode v06_23_00 was used instead of v05_08_00 which was
₁₁₇₃ used previously. The wire signal was the raw ADC value after noise filtering. Each
₁₁₇₄ collection plane grayscale image was 3456x1280x1 where 5 time ticks were pooled into
₁₁₇₅ 1 bin which is different than the previous dataset and was implemented due to the fact
₁₁₇₆ that the time ticks of an event went from 9400 to 6400 with the change of uboonecode
₁₁₇₇ version. The grayscale color standard is 8bit therefore the ADC value of wire and time
₁₁₇₈ tick was also downsampled due to the 12bit ADC value MicroBooNE has. To do this,
₁₁₇₉ the highest ADC pixel in the image was found and then this was divided by the rest
₁₁₈₀ placing all pixel values between 0-1. From there, all pixel values are then multiplied
₁₁₈₁ by 255. All images were made using a LArSoft module. Once the images were created,
₁₁₈₂ using and image manipulation framework called OpenCV images were read into a
₁₁₈₃ numpy array and cropped to the region of interest by only keeping rows and columns
₁₁₈₄ where all ADC values are higher than 0 and then resized it to 224x224 using OpenCV's
₁₁₈₅ resize function. This downsampling of ADC values creates a problem of information
₁₁₈₆ loss for example, a proton which is highly ionizing will have the same brightness as a
₁₁₈₇ minimum ionizing muon by virtue of how the images are created. Issues that arose
₁₁₈₈ in CNN1075 that were fixed in CNN10000 include zero-padding images in X and Y
₁₁₈₉ that are smaller than 224X224 to eliminate over-zooming effect and fixing a bug that
₁₁₉₀ shifted pixels separated by a dead-wire region.

1191 Images were also made from events that passed the cc-inclusive selection 1 filter
1192 right before the 75 cm track length cut and were classified using the CNN10000. The
1193 dataset used to create these images is the same one used in [?], prodgenie_bnb_nu_cosmic_uboone_mc
1194 These images were created using information from the track candidate that passed
1195 the filter. Only wire number and time ticks associated to the track candidate were
1196 drawn on the image to mimic a single particle generated image. These images were
1197 then classified using CNN10000. Two approaches were taken in making these images.
1198 The first was using the image normalization above where the maximum pixel in each
1199 image is used as a normalization constant to get all pixels between 0-1 then multiply
1200 all pixels by 255. As described above, this is the incorrect way to normalize; it should
1201 be normalized by dataset not by event, which is the second way the images were
1202 created. The results of CNN10000 performance are shown in section 7.2.

1203 7.2.2 Training CNN1075

1204 The work shown in these next sections are based on the previous work done described
1205 in [?]. That CNN (now referred to as CNN1075) was trained using single generated
1206 isotropic muons and pions from 0-2 GeV energy range. 1,075 muons and pions were
1207 used to train the network and 1,075 μ/π were used as a validation set. The accuracy is
1208 how well CNN1075 is doing by epoch and was 74.5%. The loss is gradient descent
1209 or minimization of the error of the weights and biases used in each neuron of each
1210 layer of CNN1075 and was 58% with a trend sloping downwards on the loss curve
1211 as well as a trend sloping upward in the accuracy curve. Due to the depth of the
1212 neural network framework, it was necessary to train with a larger dataset and for
1213 more epochs, however, the downward slope of the loss curve is an indication that once
1214 trained for longer with a higher training sample, neural networks can be used for μ/π
1215 separation. Updates in the image making and downsampling algorithm were made to
1216 fix issues that arose in CNN1075.

1217 7.2.3 Training CNN10000

1218 The hyperparameters used for CNN10000 are shown. The batch size for the training
1219 and testing as well as the test iter were chosen to encompass the whole training/testing
1220 image set when doing accuracy/loss calculations. To do this, multiplying the test

1221 iter by the test batch size give you the amount of images used when calculating
1222 accuracy/loss curves. For reference, the accuracy and loss are defined as well.

```
1223     • train_batch_size: 100
1224     • test_batch_size: 100
1225     • test_iter: 100
1226     • test_interval: 100
1227     • base_lr: 0.001
1228     • lr_policy: "step"
1229     • gamma: 0.1
1230     • stepsize: 1000
1231     • display: 100
1232     • max_iter: 10000
1233     • momentum: 0.99
1234     • weight_decay: 0.0005
1235     • snapshot: 100
1236     • Accuracy: How often the CNN predicts the truth over total number of images
1237     • Loss: Error between truth and prediction. Minimize loss by gradient descent to
1238       update weights and biases of CNN
```

1239 The same architecure that was used to train CNN1075 was employed on CNN10000,
1240 Imagenet. Caffe [?] was the software package used for both CNNs. The differences
1241 include batch size and test_iter and momentum to account for the larger dataset. Both
1242 CNNs were trained on a CPU machine, Syracuse01. Further training will be done
1243 on a GPU cluster stationed at Syracuse University. Figure 7.1 shows the loss and
1244 accuracy of CNN10000. There is around a 10% increase in accuracy from CNN1075 to
1245 CNN10000, 85%, and around a 20% decrease in loss, 36%.

1246 Figure 7.2 show a breakdown of μ/π separation for CNN10000. It also shows
1247 the network is not being overtrained due to the Accuracy of both the training and
1248 testing datasets being within .01% of eachother. The CNN is doing a very good job of

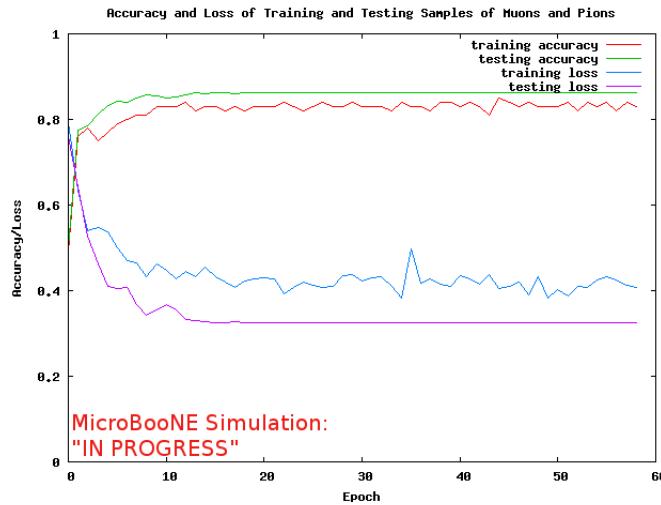
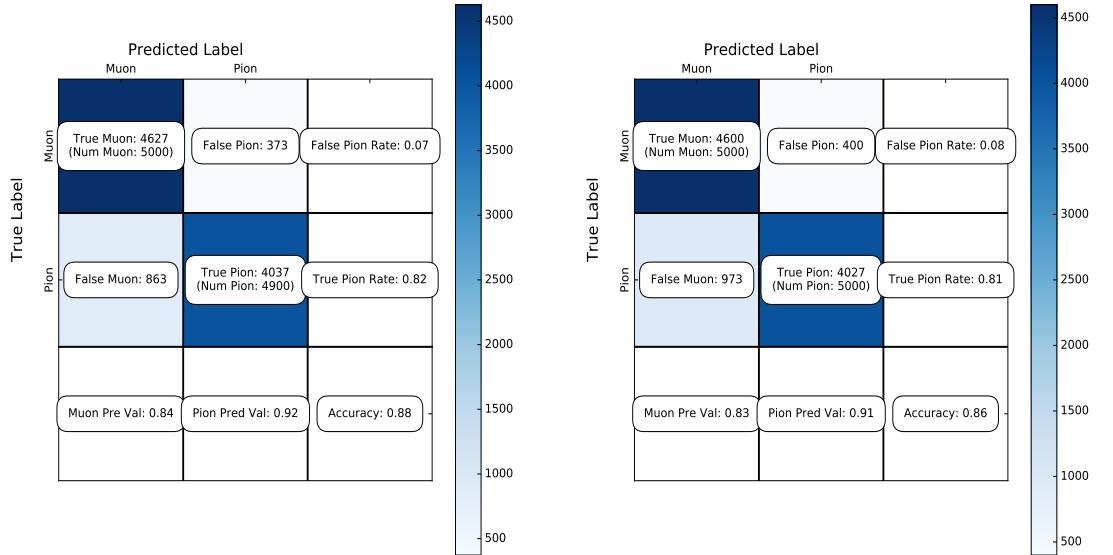


Figure 7.1: Accuracy vs. Loss of ImageNet 2-output μ/π sample consisting of 10000 images each.

1249 classifying true muons as muons, and our loss increase from CNN1075 is due to the
1250 increase in accurately classifying pions as pions.

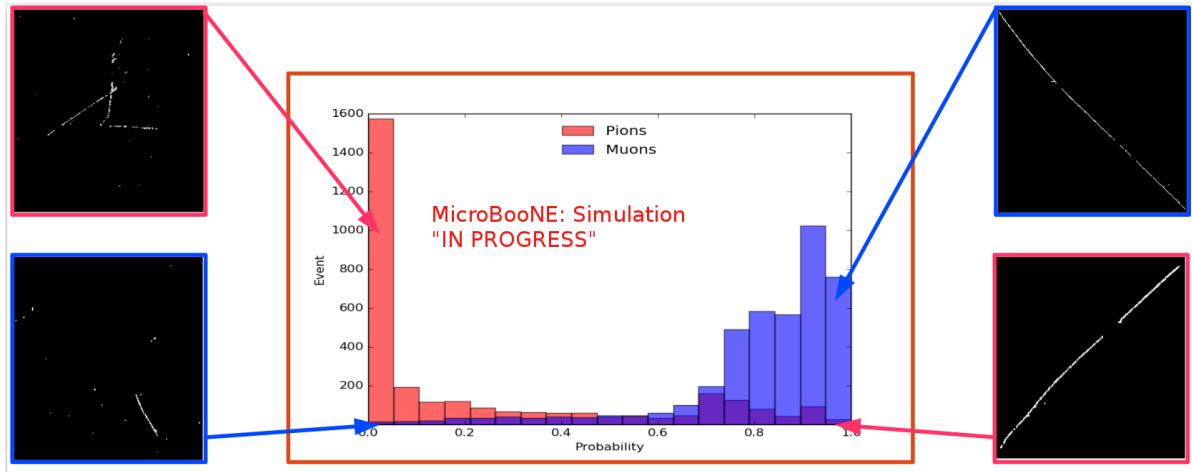
1251 7.2.4 Training CNN100000

1252 Results of training using 100,000 images, 20,000 images per $\mu/\pi/p/\gamma/e$.



(a) Confusion Matrix showing Accuracy of CNN using training data

(b) Confusion Matrix showing Accuracy of CNN using testing data



(c) Probability plot of muons and pions from testing set. Images surrounding histogram are a random event from lowest bin and highest bin for each particle.

Figure 7.2: Description of confusion matrix variables: False pion rate = $false\pi / total\pi$ True pion rate = $true\pi / total\pi$ Accuracy = $(true\pi rate + true\mu rate) / 2$ Pion prediction value = $true\pi / (true\pi + false\pi)$ Muon prediction value = $true\mu / (true\mu + false\mu)$

7.2c The probability plot includes muons and pions that are classified as primary particles.

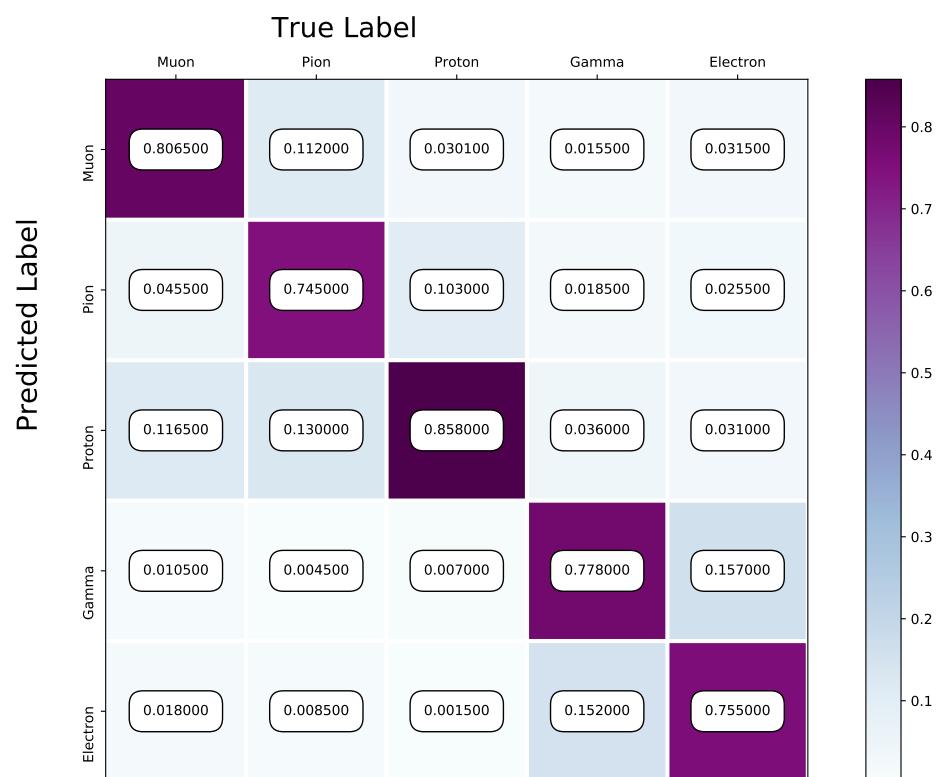


Figure 7.3: Confusion Matrix of all five particles

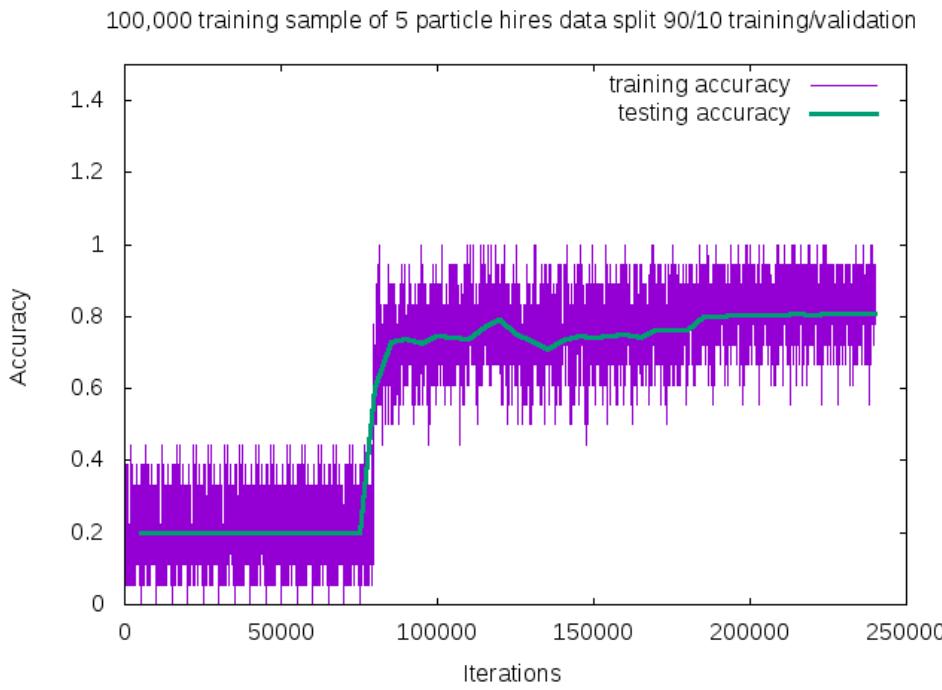


Figure 7.4: Training and testing accuracy of CNN trained on 100,000 images of $\mu/\pi/p/\gamma/e$ with 20,000 images of each particle. Each image was a size of 576x576 and the images per particle were split 90% use for training and 10% used for testing the network

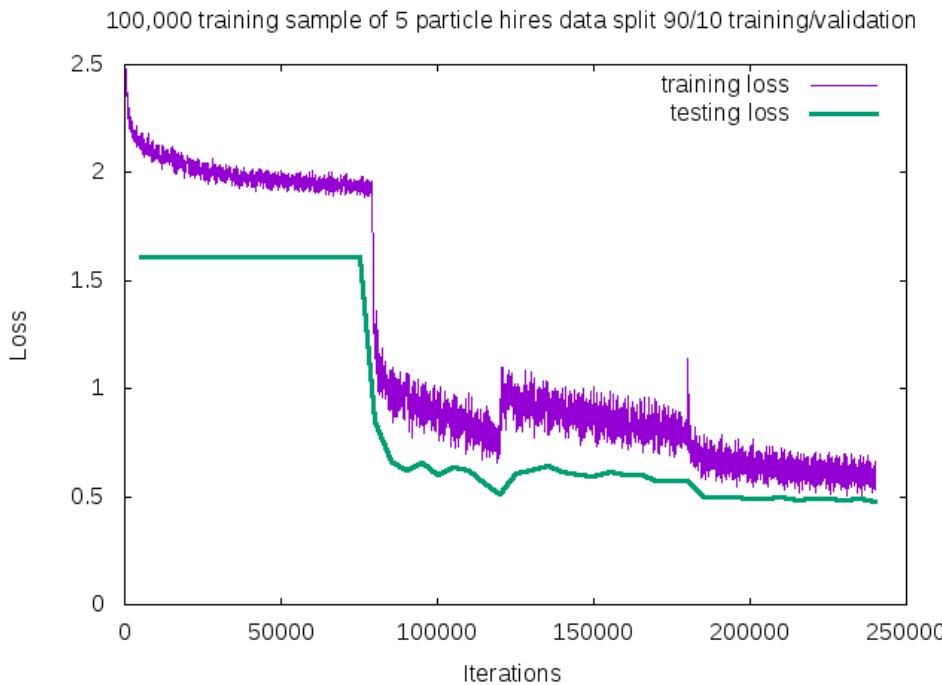


Figure 7.5: Training and testing loss of CNN trained on 100,000 images of $\mu/\pi/p/\gamma/e$

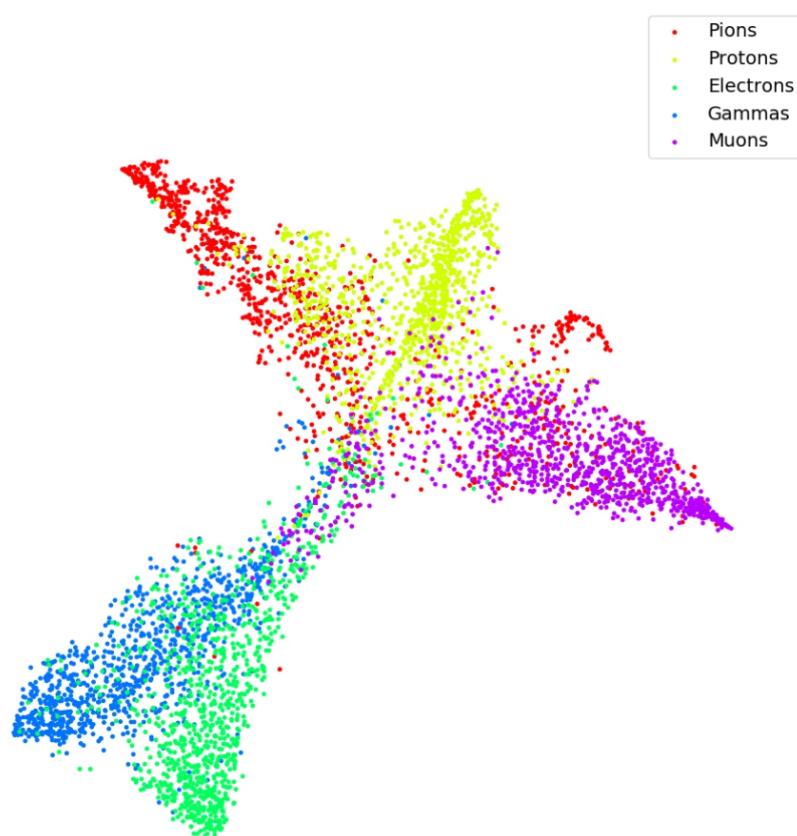


Figure 7.6: t-SNE of CNN

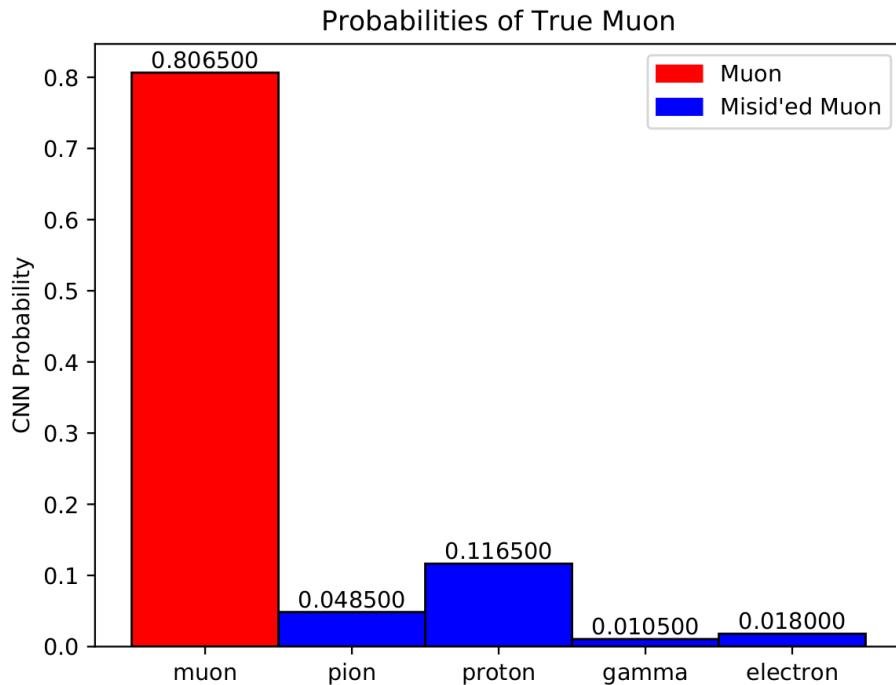


Figure 7.7: Muon Prob

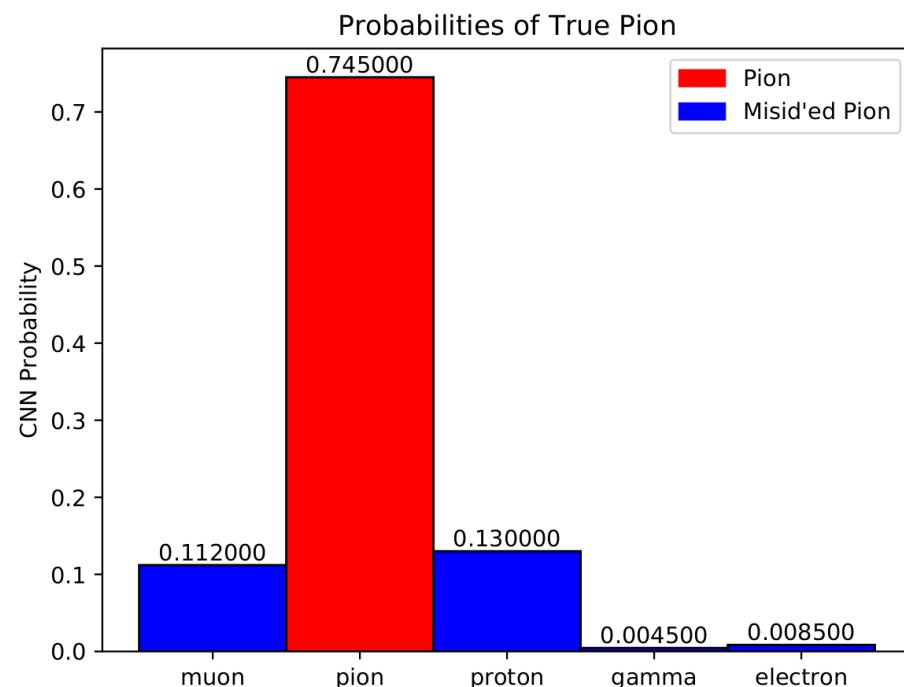


Figure 7.8: Pion Prob

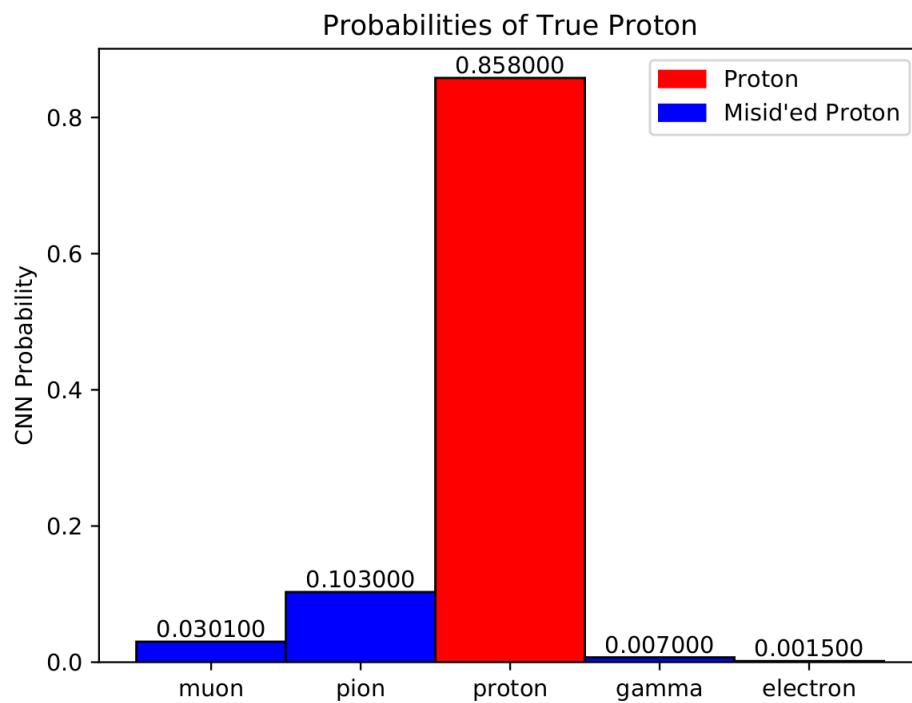


Figure 7.9: Proton Prob

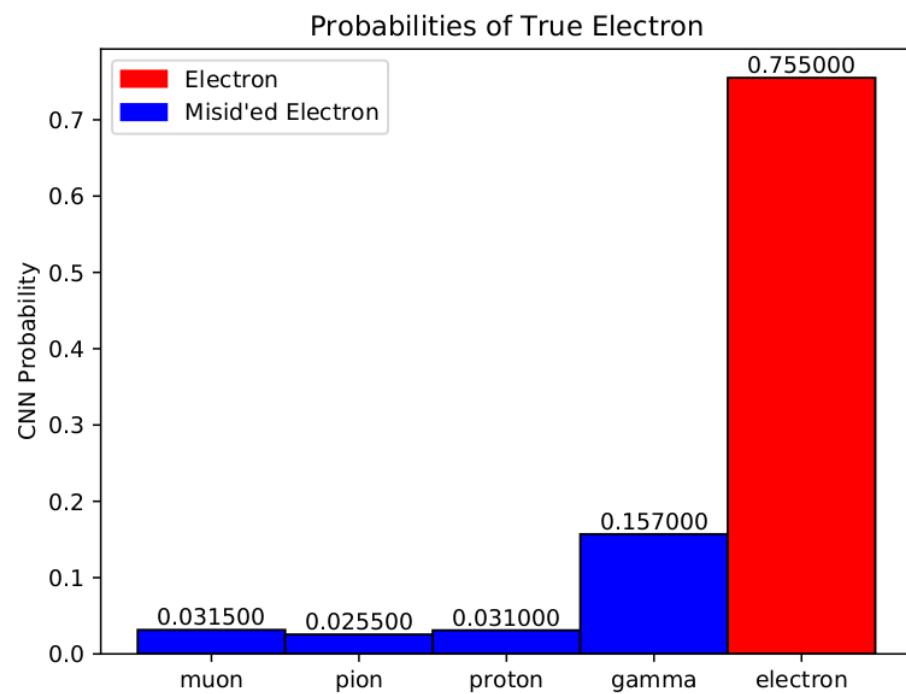


Figure 7.10: Electron Prob

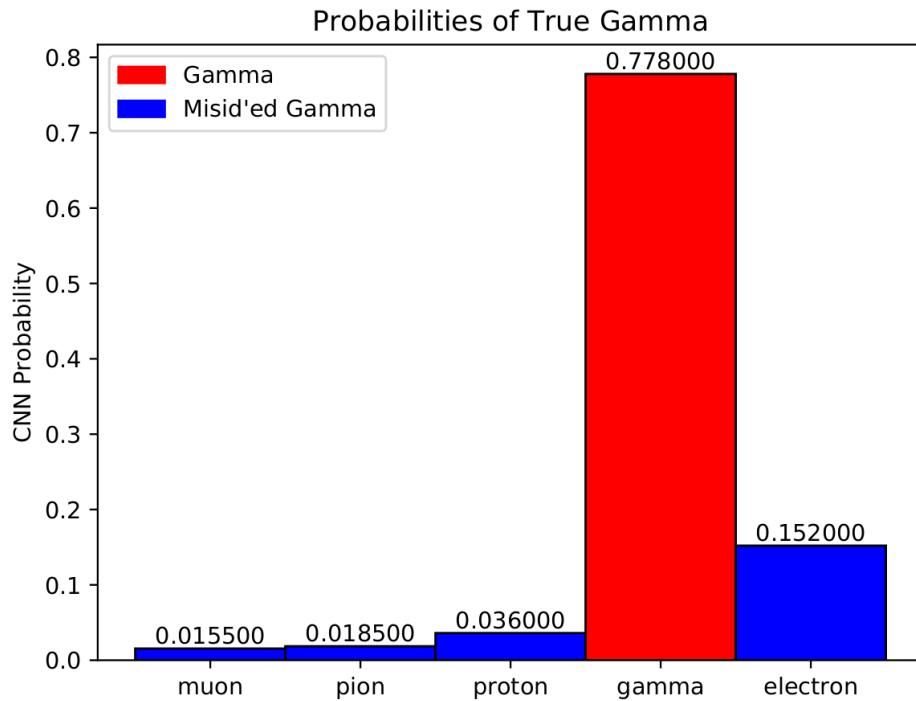


Figure 7.11: Gamma Prob

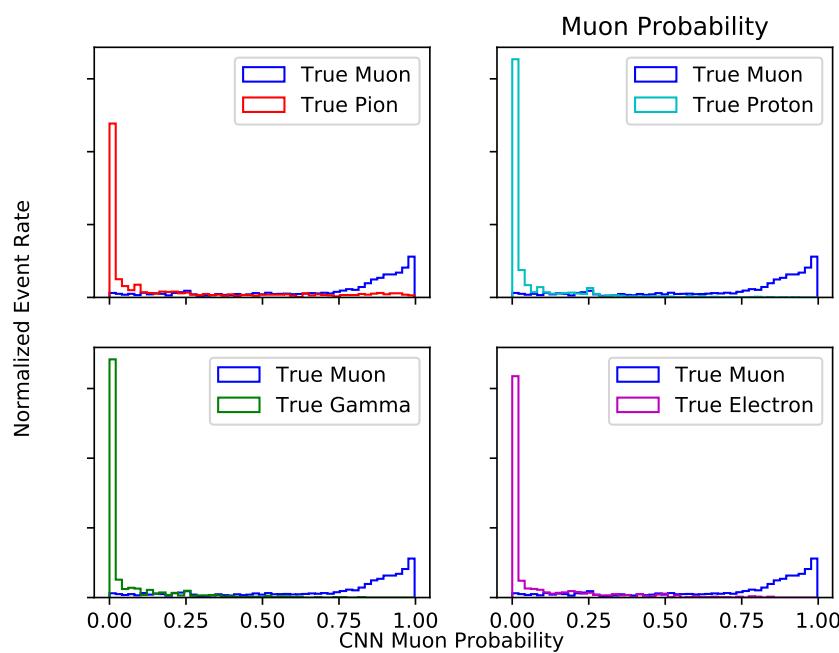


Figure 7.12: Prob

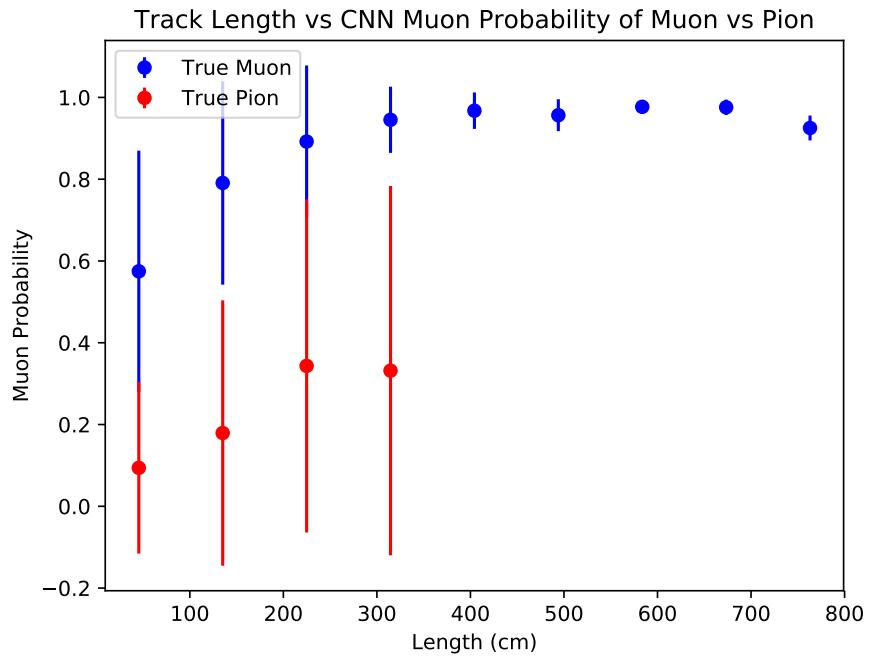


Figure 7.13: mupi

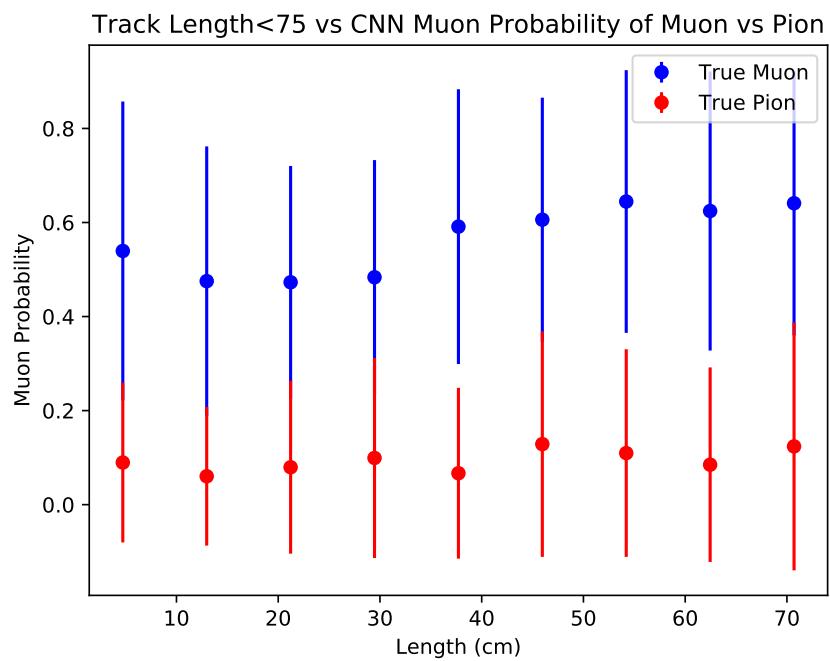
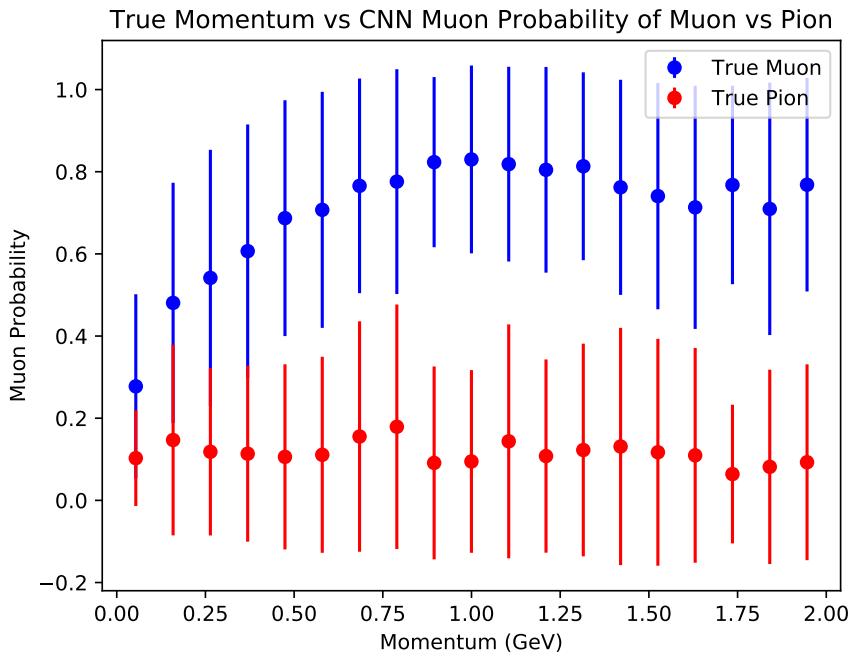
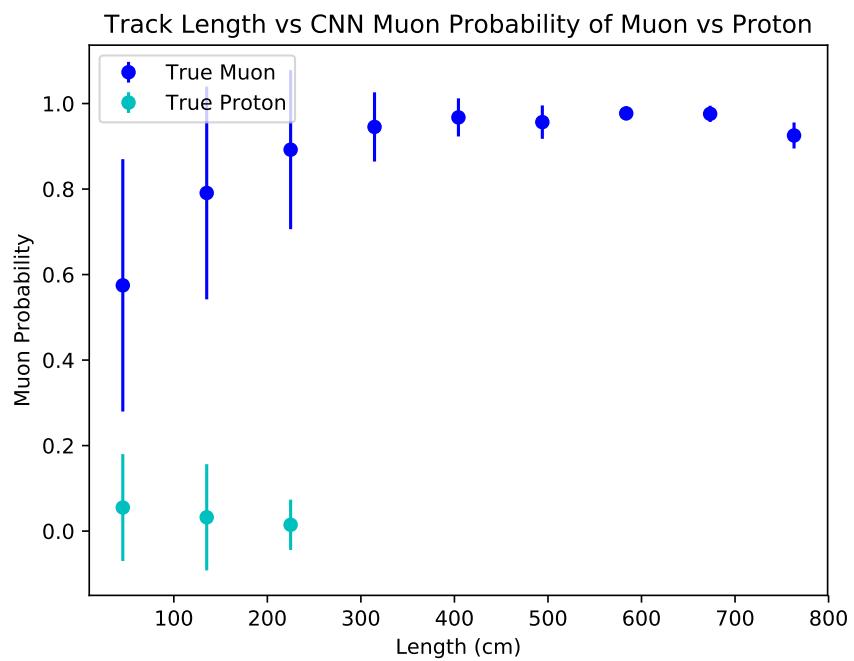
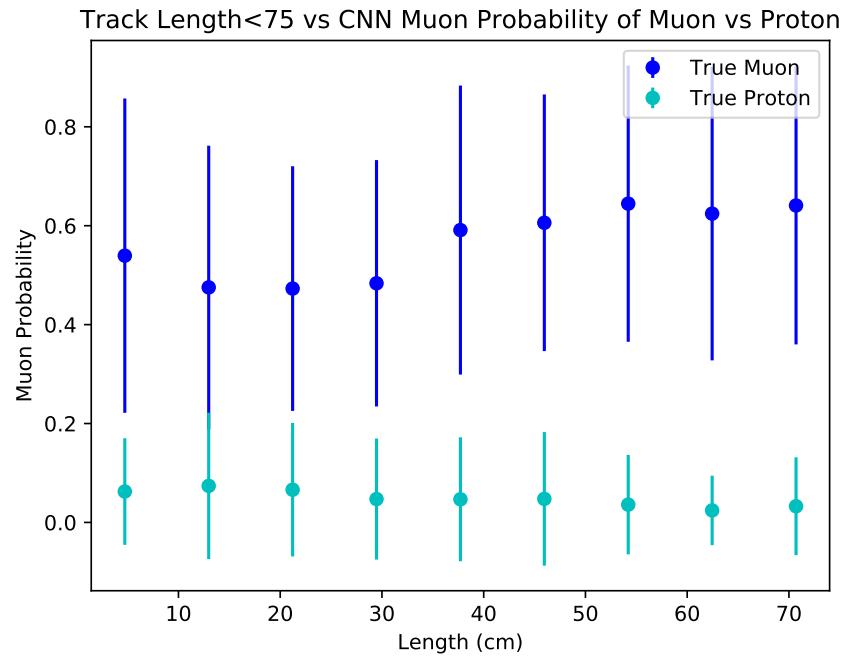
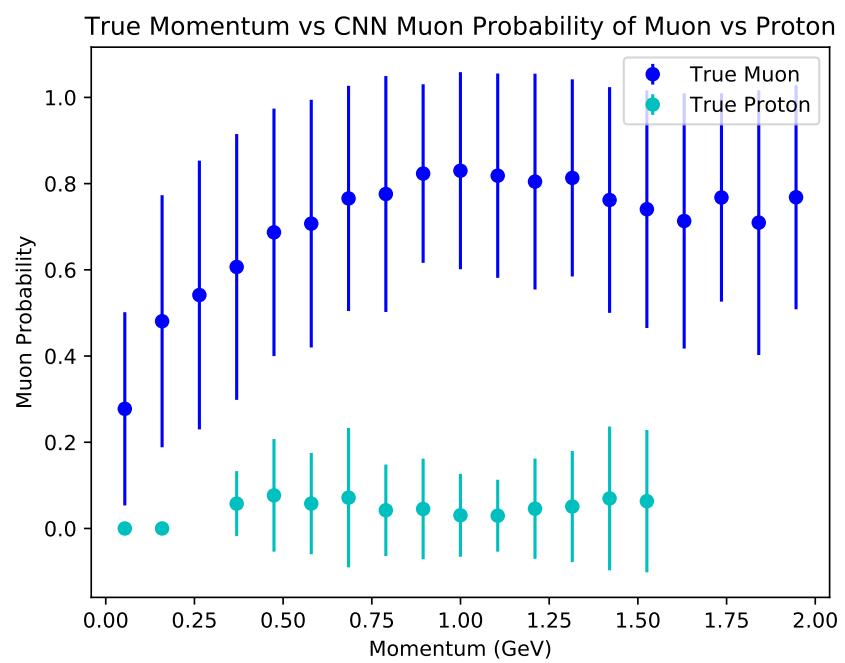


Figure 7.14: mupi

**Figure 7.15:** mupi**Figure 7.16:** mup

**Figure 7.17: mup****Figure 7.18: mup**

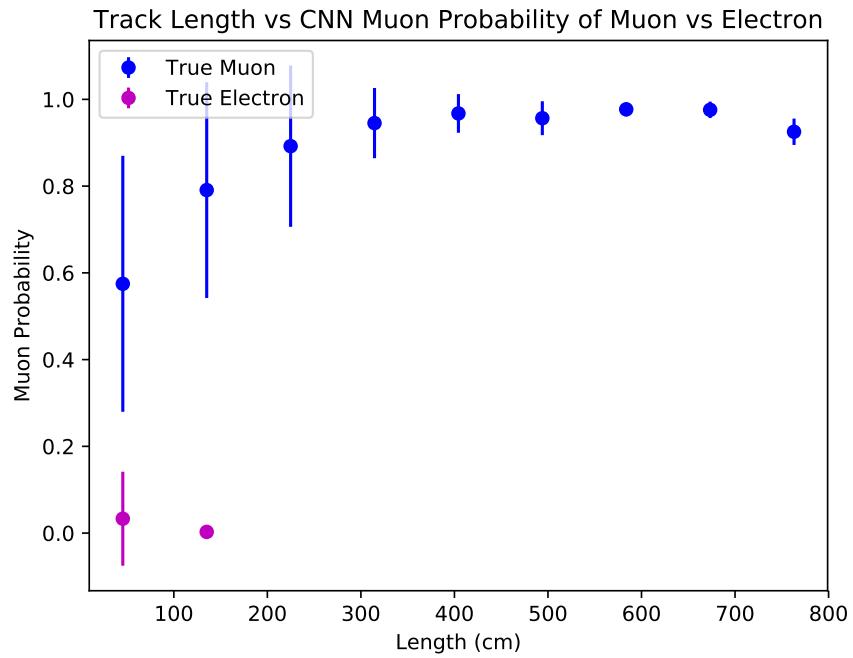


Figure 7.19: mue

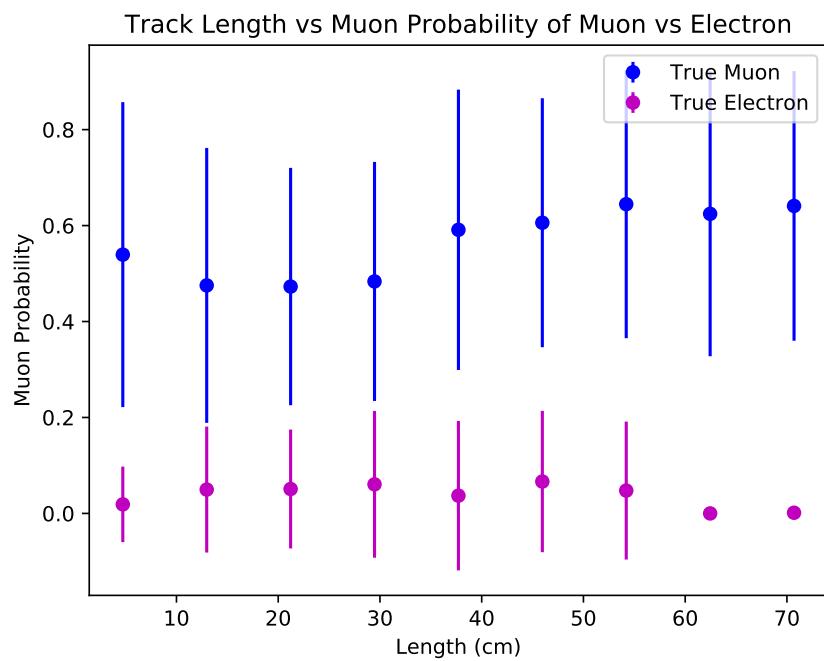


Figure 7.20: mue

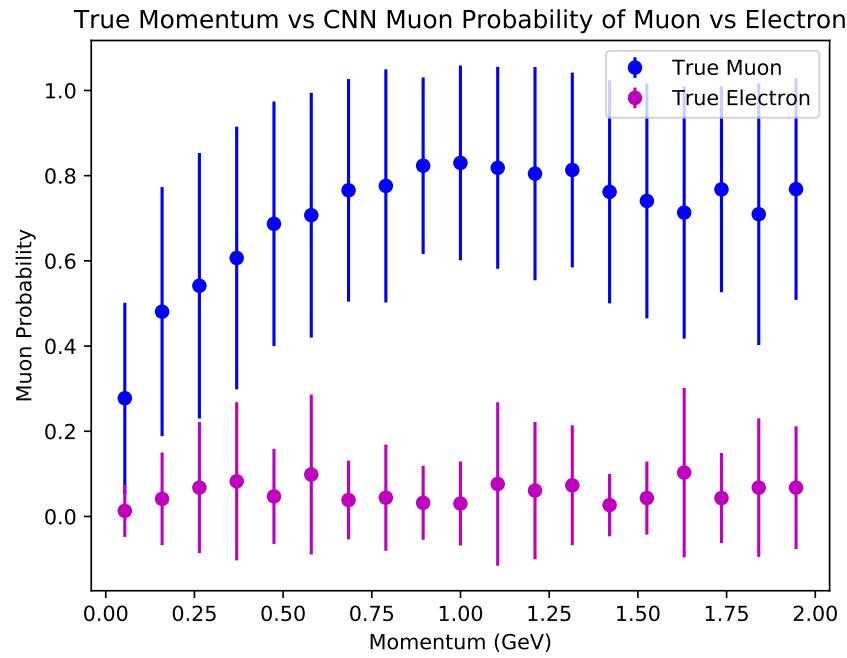


Figure 7.21: mue

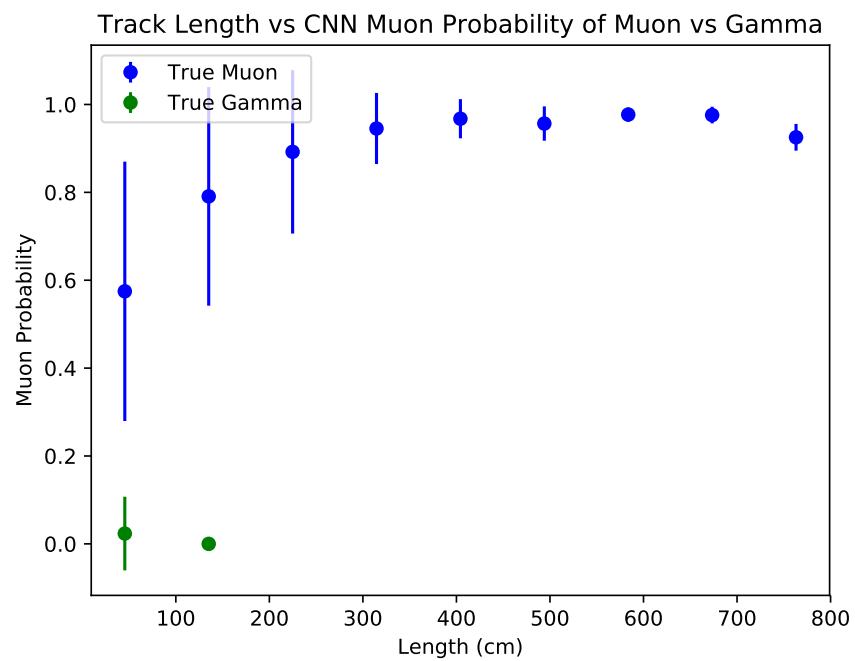


Figure 7.22: mug

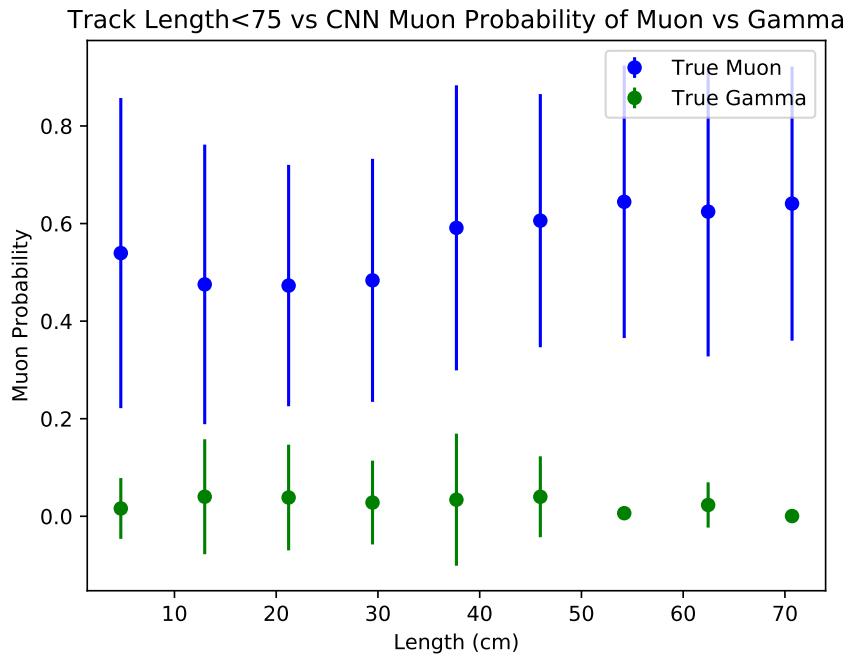


Figure 7.23: mug

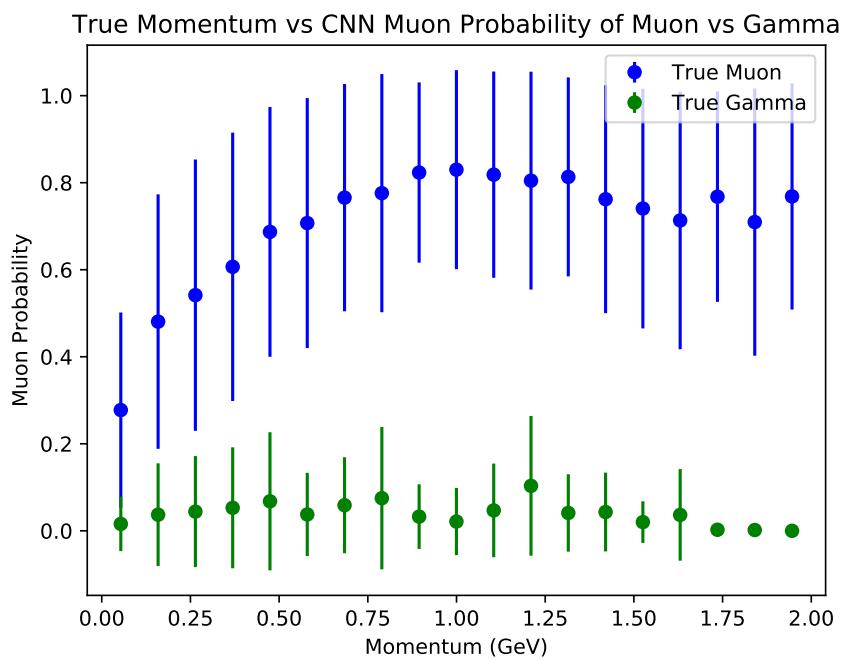


Figure 7.24: mug

1253 **Chapter 8**

1254 **Results of Convolutional Neural**
1255 **Networks on particles WORKING**
1256 **TITLE**

1257 **8.1 Classification using CNN10000**

1258 **8.1.1 Classification of MC data using Selection I Original**
1259 **CC-Inclusive Filter**

1260 The next step that was taken was to use CNN10000 to classify track candidate images
1261 that were identified by the selection I original cc-inclusive filter described in [?].
1262 Passing rates for each cut in cc-inclusive filter are show in figure 8.1. For the incorrect
1263 image making normalization dataset, out of 188,880 events, 7438 passed the cut right
1264 before 75 cm track length cut which is 3.9% of total data. Discrepancies in passing rates
1265 are due to grid submission issues, however, this dataset is used to check if changes
1266 in image making normalization affects μ/π separation probability due to CNN10000
1267 being trained with incorrectly image making normalized data. For the second dataset
1268 with correct image making normalization, out of 188,880 events, 9552 events passed the
1269 cut right before the 75 cm track length cut which is 5.1% passing rate and is comparable
1270 to figure 8.1. In time cosmics were also run over for efficiency and purity calculations.
1271 Out of 14395 in time cosmic events, 175 passed the cut right before the 75 cm track
1272 length cut which is a passing rate of 1.2% compared to 1.3% shown in figure 8.1.

Table 3: **Selection I: Original** The table shows passing rates for the above described event selection. Numbers are absolute event counts and Cosmic background is not scaled appropriately. The BNB+Cosmic sample contains all events, not just ν_μ CC inclusive. The selected events are further broken up in the following subsection. The numbers in brackets give the passing rate wrt the step before (first percentage) and wrt the generated events (second percentage). In the BNB+Cosmic MC Truth column shows how many true ν_μ CC inclusive events (in FV) are left in the sample. This number includes possible mis-identifications where a cosmic track is picked by the selection instead of the neutrino interaction in the same event. The cosmic only sample has low statistics, but please note that it is not used in any plots, it is just for illustrating the cut efficiency. The last column Signal:Cosmic only gives an estimate of the ν_μ CC events wrt the cosmic only background at each step. For this number, the cosmic background has been scaled as described in chapter 2. Note that this numbers is not a purity, since other backgrounds can't be determined at this step.

	BNB + Cosmic Selection	MC-Truth	Cosmic only	Signal: Cosmic only
Generated events	191362	45273	4804	1:22
≥ 1 flash with ≥ 50 PE	136219 (71%/71%)	44002 (97%/97%)	2979 (62%/62%)	1:14
≥ 1 vertex in FV	131170 (96%/69%)	43794 (99%/97%)	2805 (94%/58%)	1:13
≥ 1 track within 5 cm of vertex	129784 (99%/68%)	43689 (99%/97%)	2756 (98%/58%)	1:13
flash matching of longest track	44775 (34%/23%)	23647 (54%/52%)	647 (23%/13%)	1:5.7
track containment	10114 (23%/5.3%)	6882 (29%/15%)	61 (9.4%/1.3%)	1:1.9
track ≥ 75 cm	7358 (73%/3.8%)	5801 (84%/13%)	31 (51%/0.6%)	1:1.1

Figure 8.1: Snapshot of passing rates of Selection I from CC-Inclusive Filter

Figures 8.2a, 8.2b, 8.2c and 8.2d show the accuracy and μ/π separation of both the correct and incorrect normalized images. The confusion matrices are only composed of μ/π data. Other particles passed the cc-inclusive filter before the 75 cm track length cut and were all mis-id'ed as muons. Since CNN10000 has not seen any particles other than muons and pions, it makes sense that those get mis-id'ed. Figures 8.2b and 8.2d don't have μ/π separation comparable to 7.2c, but 8.2b does skew to higher probabilities compared to 8.2d. This is to be expected and further work on quantifying the performance of CNN10000 should use the incorrect image making normalization. It is also expected that the separation isn't as defined as the testing dataset for CNN10000. CNN10000 was trained and tested using single particle muons and pions and the track candidate dataset come from BNB+Cosmic events, not to mentions all track candidates have passed the cc-inclusive filter that tags "muon-like" tracks therefore the pions in this sample look much closer in muon topology than the network has seen. Also, these images were made from wire and time ticks associated to hits from the track candidate that passed the cc-inclusive filter. This is different from the training images where a bounding box was drawn over the total μ or π interaction. Spurious energy deposition from a $\pi - Ar$ interaction is most likely not included in the BNB+Cosmic images due to the tracking algorithm. To remedy this, the neural network needs to see more "muon-like" pions and muons and pions from a neutrino interaction passing the cc-inclusive filter as well as a larger particle variety including protons, photons

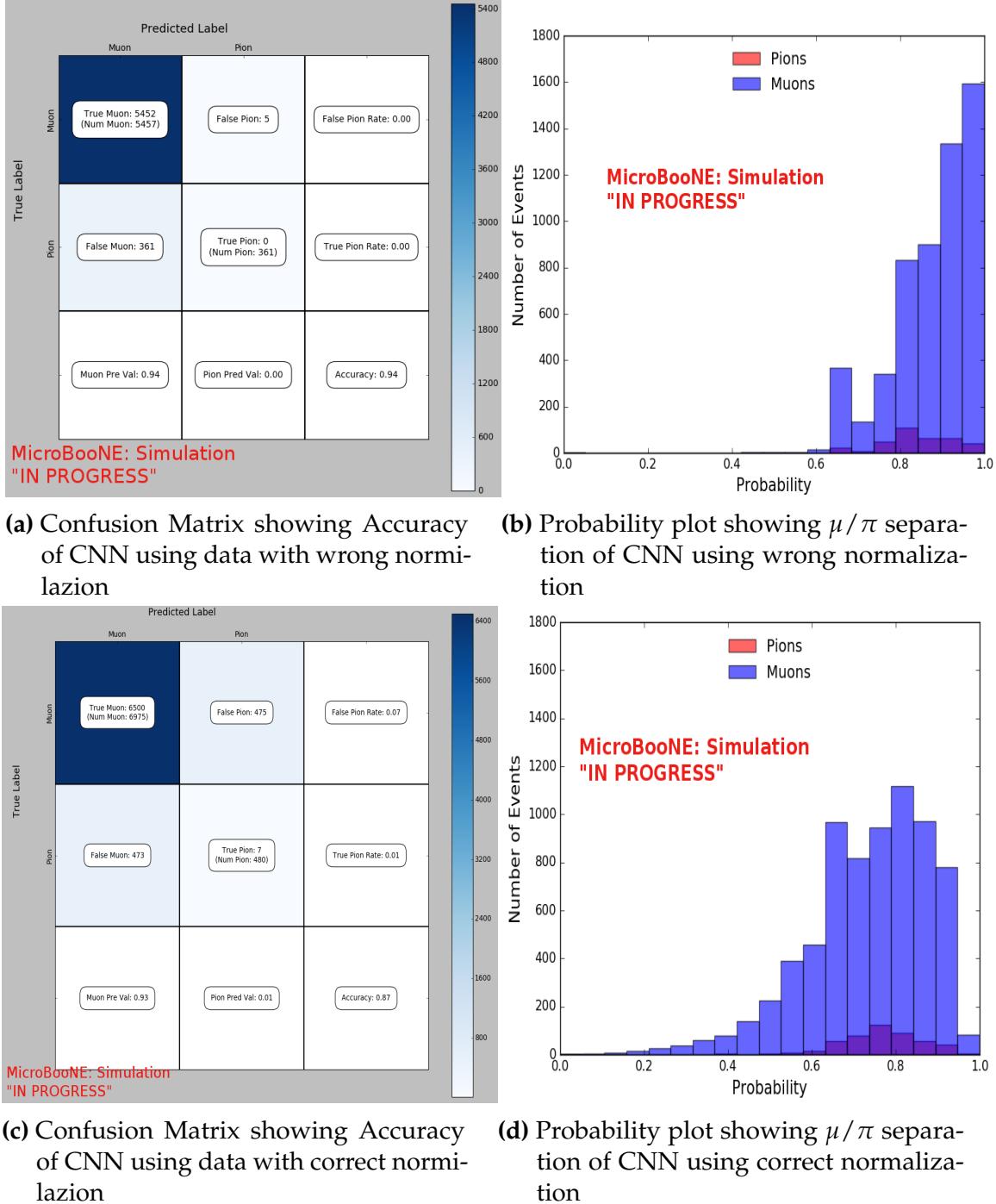


Figure 8.2: Results of CNN10000 classification of track candidate images output from cc-inclusive filter.

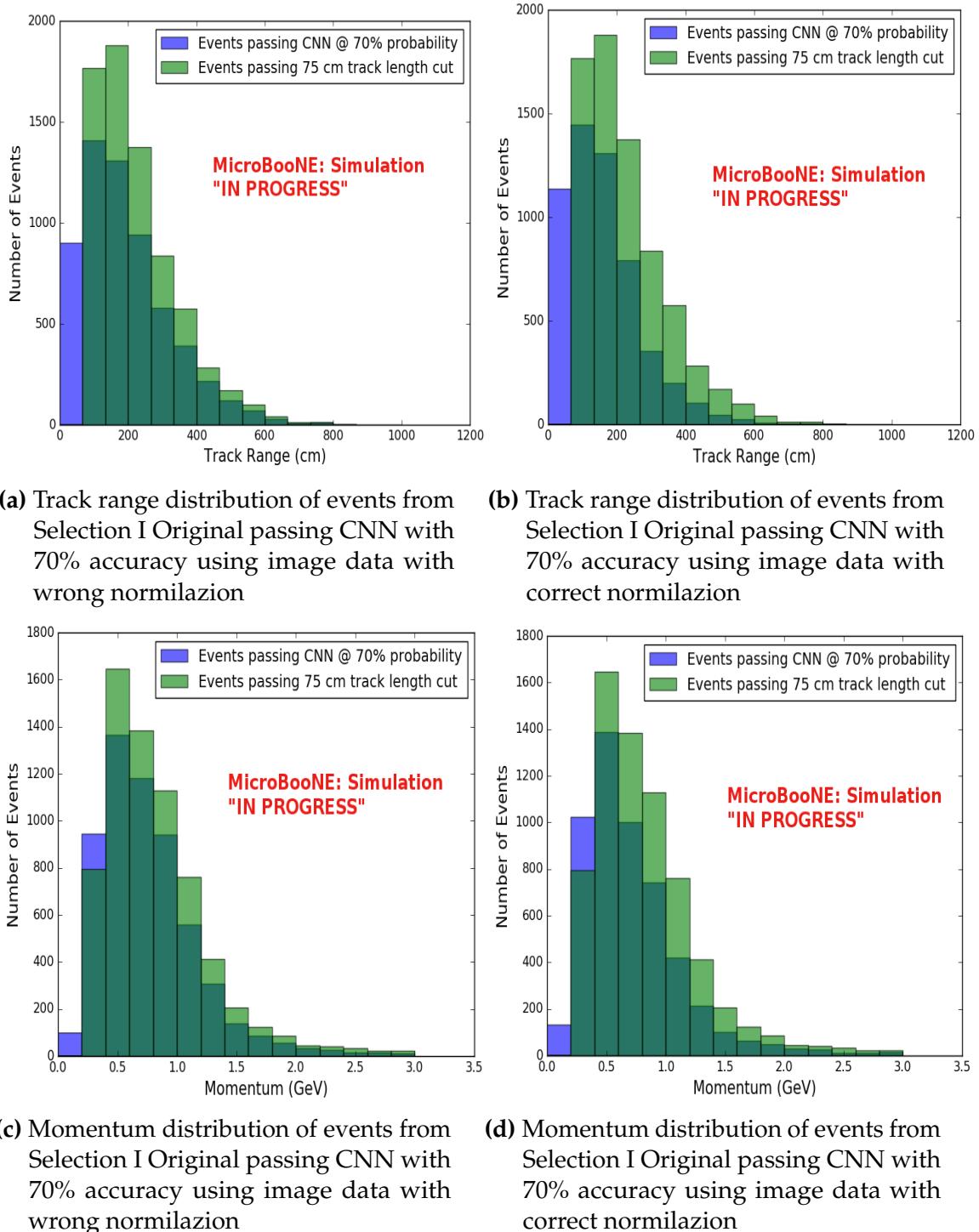


Figure 8.3: CNN10000 distributions of track candidate images output from Selection I Original cc-inclusive filter with different image data normalizations

and electrons. Although μ/π separation is lacking, CNN10000 does an excellent job of classifying muons and using higher CNN probability can increase purity. Figures 8.3a, 8.3b, 8.3c and 8.3d show the track and momentum distributions for these two datasets. In both sets you have an increase in data in the bin below 75 cm and at bins below 0.5 GeV. These distributions were made with events classified with 70% probability of being a muon regardless of true particle type.

8.1.2 Classification of MC data using Selection I Modified CC-Inclusive Filter

Table 5: Selection I: Modified The table shows passing rates for the above described event selection. Numbers are absolute event counts and Cosmic background is not scaled appropriately. The BNB+Cosmic sample contains all events, not just ν_μ CC inclusive. The selected events are further broken up in the following subsection. The numbers in brackets give the passing rate wrt the step before (first percentage) and wrt the generated events (second percentage). In the BNB+Cosmic Mc Truth column shows how many true ν_μ CC inclusive events (in FV) are left in the sample. This number includes possible mis-identifications where a cosmic track is picked by the selection instead of the neutrino interaction in the same event. The cosmic only sample has low statistics, but please note that it is not used in any plots, it is just for illustrating the cut efficiency. The last column Signal:Cosmic only gives an estimate of the ν_μ CC events wrt the cosmic only background at each step. For this number, the cosmic background has been scaled as described in chapter 2. Note that this numbers is not a purity, since other backgrounds can't be determined at this step.

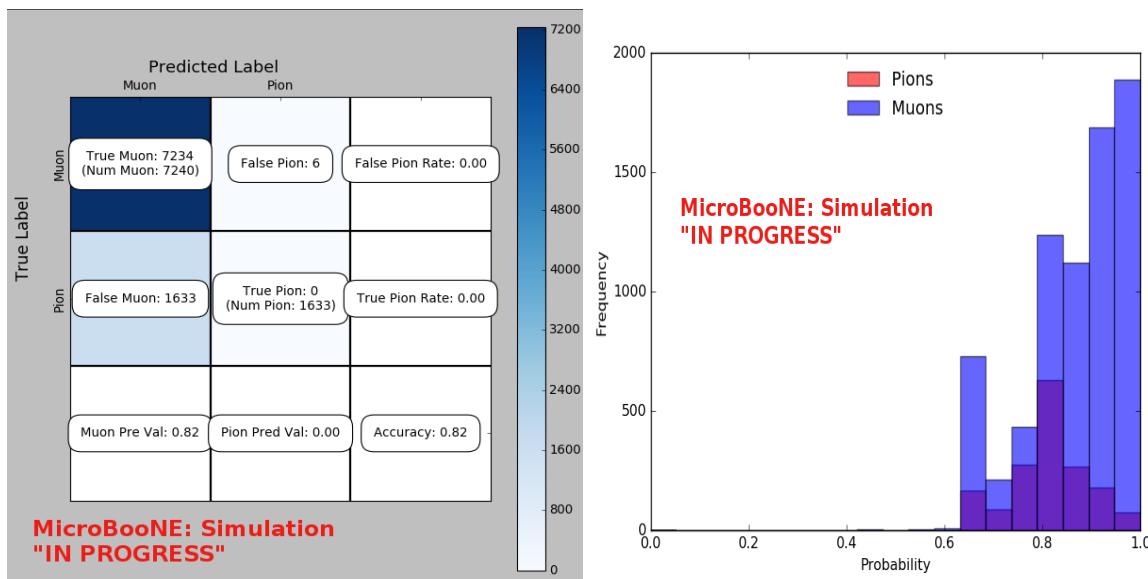
	BNB + Cosmic Selection	MC-Truth	Cosmic only	Signal: Cosmic only
Generated events	191362	45273	4804	1:22
≥ 1 flash with ≥ 50 PE	136219 (71%/71%)	44002 (97%/97%)	2979 (62%/62%)	1:14
≥ 1 track within 5 cm of vertex	135830 (99%/71%)	43974 (99%/97%)	2975 (99%/62%)	1:14
vertex candidate in FV	79112 (58%/41%)	34891 (79%/77%)	1482 (50%/31%)	1:8.9
flash matching of longest track	40267 (51%/21%)	25891 (74%/57%)	340 (23%/7.1%)	1:2.8
track containment	19391 (48%/10%)	11693 (45%/26%)	129 (38%/2.7%)	1:2.3
track ≥ 75 cm	6920 (36%/3.6%)	5780 (49%/13%)	17 (13%/0.4%)	1:0.6

Figure 8.4: Snapshot of passing rates of all cuts from Selection I Modified cc-inclusive filter

CNN10000 was also used to classify track candidate images that were identified by the selection I modified cc-inclusive filter described in [?]. Passing rates for each cut in this filter are shown in figure 8.4. As seen in section 8.1.1, wrong image normalization had a higher muon classification probability so all work done using selection I modified cc-inclusive filter was done using this normalization. Out of 188,880 events, 19,112 passed the cut right before the 75 cm track length cut which is a 10.1% passing rate and comparable to the 10% passing rate shown in figure 8.4. In time cosmics were also run over, out of 14,606 in time cosmics events, 302 passed the cut right before the 75 cm track length cut which is a 2.1% passing rate comparable to the 2.7% passing rate in the cc-inclusive tech-note. Figures 8.5a and 8.5b show the accuracy and μ/π separation. Both plots are only composed of muons and pions and like selection I original data,

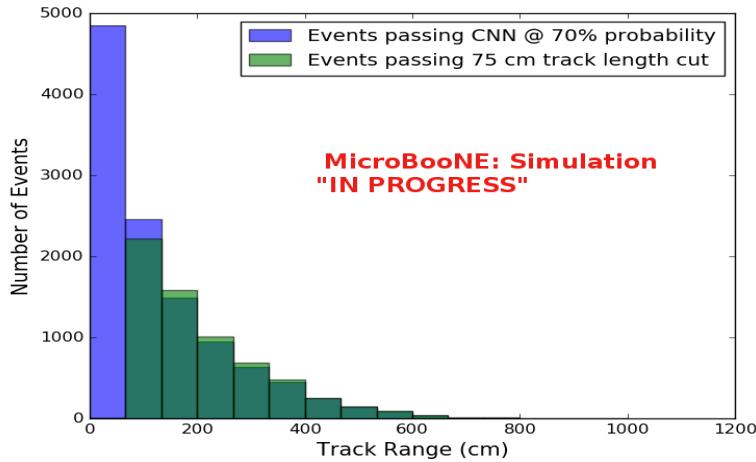
1312 all other particles were id'ed as muons. Also like selection I original data, muons are
 1313 being identified at a very high rate. Figure 8.6a shows the track range distributions
 1314 of all events from selection I modified being classified by the CNN as a muon with a
 1315 probability of 70% regardless of true particle type. We get entries for the CNN curve
 1316 in the lowest bin and none for the 75 cm curve. To see how many true CC events
 1317 were identified by CNN10000 breaking down figure 8.6a by event type was necessary.
 1318 Figures 8.6b and 8.6c show track range distributions separated by signal and various
 1319 backgrounds. Particle type was not taken into consideration in these plots so true CC
 1320 event images can be any track candidate particle passing selection I modified cut right
 1321 before track length cut including pions and protons.

1322 To gain an even deeper understanding on how CNN10000 is performing, plotting
 1323 these distributions with only muons and pions was done due to the fact that CNN10000
 1324 was trained with only those particles for μ/π separation. Figures 8.6d-8.7d show the
 1325 stacked histograms of signal and background of the track range distributions with
 1326 varying CNN probabilities starting from 70% and ending at 90% probability. With
 1327 higher probabilities we get a purer sample in the lower bin but we end up losing
 1328 events as well. Momentum distributions for all signal/background events are shown
 1329 in figure 8.8.

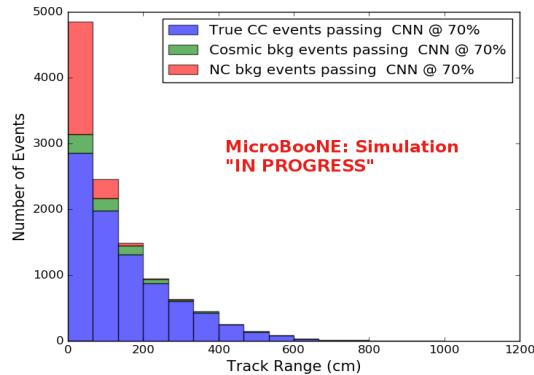


(a) Confusion Matrix for CNN10000 classified events from selection I modified (b) Probability plot for CNN10000 classified events from selection I modified

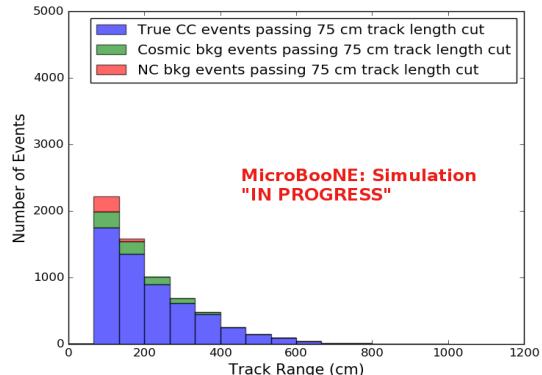
Figure 8.5: Confusion matrix and probability plot of events passing selection I modified cc-inclusive cuts right before 75cm track length cut



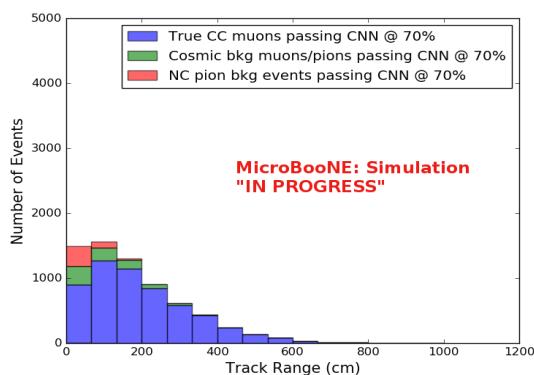
(a) Track range distribution of events from Selection I Modified passing CNN with 70% accuracy



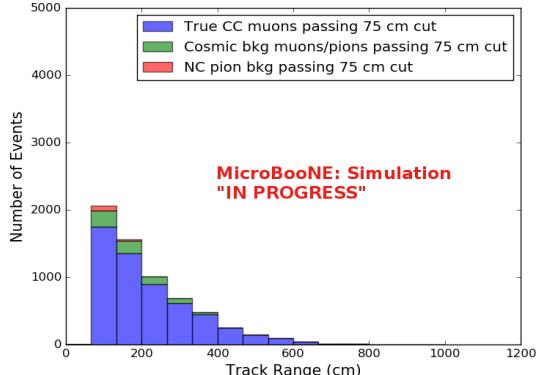
(b) Stacked signal and background track range distributions from Selection I Modified passing CNN with 70% accuracy



(c) Stacked signal and background track range distributions from Selection I Modified passing 75 cm track length cut



(d) Stacked signal muons and background muons/pions of track range distributions from Selection I Modified passing CNN with 70% accuracy



(e) Stacked signal muons and background muons/pions of track range distributions from Selection I Modified passing 75 cm track length cut

Figure 8.6: CNN10000 distributions of track candidate images output from Selection I Modified cc-inclusive filter

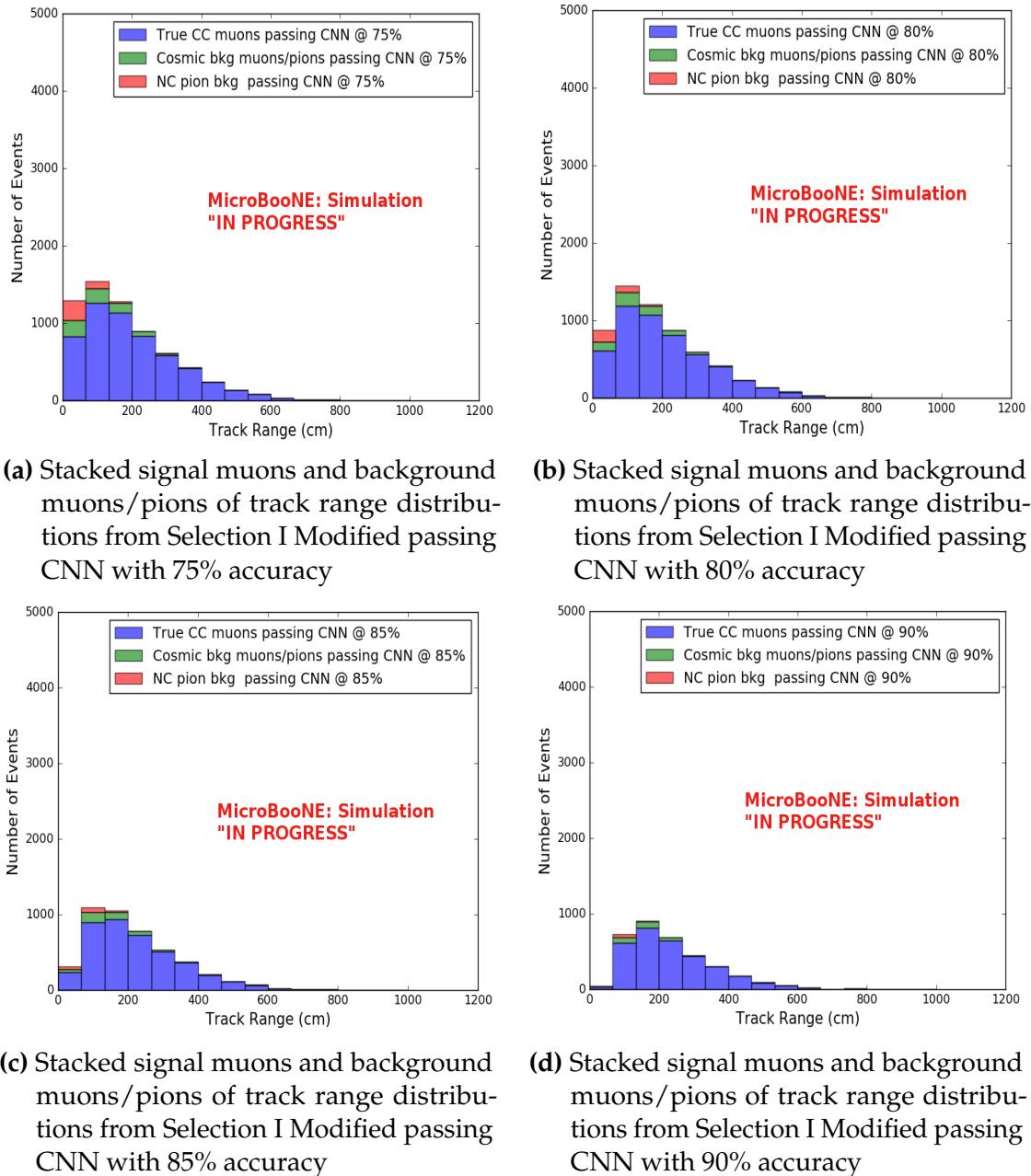
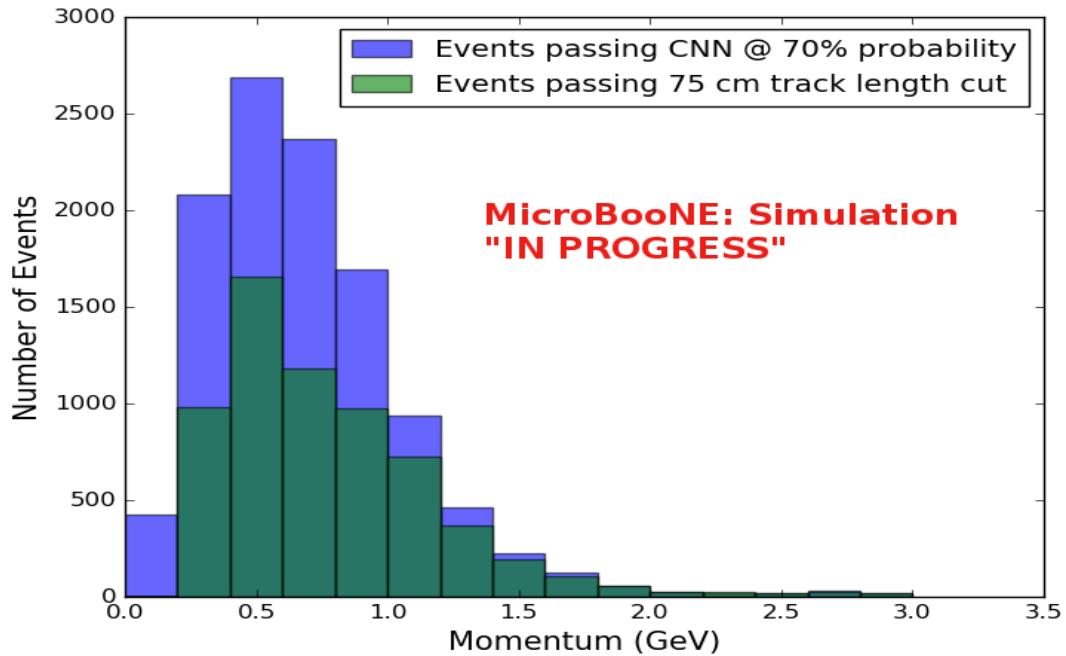
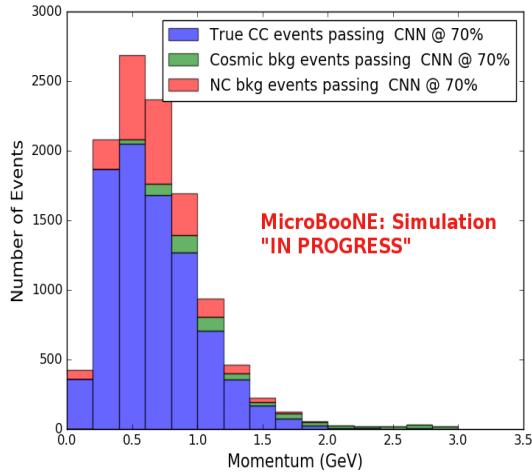


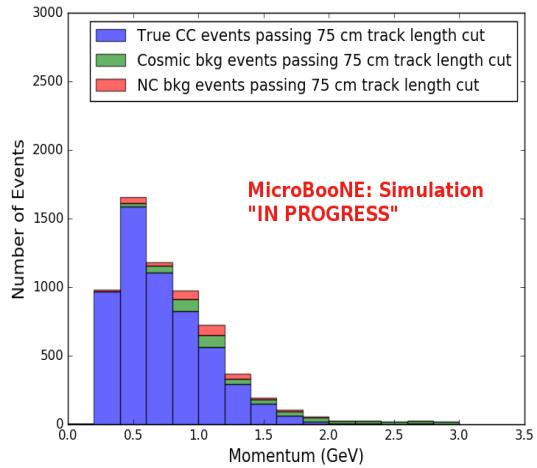
Figure 8.7: CNN10000 stacked signal/background track range distributions of track candidate images output from Selection I Modified cc-inclusive filter



(a) Momentum distribution of events from Selection I Modified passing CNN with 70% accuracy



(b) Stacked signal and background momentum distributions from Selection I Modified passing CNN with 70% accuracy



(c) Stacked signal and background momentum distributions from Selection I Modified passing 75 cm track length cut

Figure 8.8: CNN10000 momentum distributions of track candidate images output from Selection I Modified cc-inclusive filter

Another check was to see if any true CC pions were passing through the cut right before the 75 cm track length cut. Figure 8.9 shows the comparison of the stacked track range distribution with only true CC muon signal versus the stacked distribution with true CC muons and pions signal. As you can see, we gain more events when plotting CC events with a particle type of either muons or pions due to the CNN classifying all pions in this dataset as muons. This is an interesting scenario and a sample of topologies of these images are represented in figure 8.10, at least 3 tracks are coming out of the vertex for these types of events. With the 75 cm track length cut, the selection is cutting event topologies like this where the pion is the tagged track candidate. Figure 8.10a has a defined longer muon track, but because of dead wires through the track, the reconstructed range is 1. less than 75 cm and 2. shorter than the reconstructed pion whose length is also less than 75 cm. This is a very interesting event, but because of issues with the tracking algorithm, the 75 cm cut would get rid of this event. The CNN was able to recover this event only because it has classified all pions as muons. Figure 8.10b shows the second case to think about, the pion, while still less than 75 cm has a reconstructed track length longer than the muon. Again, the CNN recovered this event due to pions being classified as muons. Lastly, figure 8.10c shows a pion with a reconstructed track length greater than 75 cm and the muon. These three cases show that a broader question must be asked when training the network other than is it a muon or pion. There are different routes to recover interesting events like these. One route is to ask the network “Is it a CC event or is it an NC event?” and obtain an image dataset consisting of whole CC/NC events that will train the network to answer this question. The other route is to ask the network “Is this a $\mu/\pi/p/$ from a CC event or NC event and obtain an image dataset consisting of primary particles from a CC/NC event. Both these paths will be explored in future work.

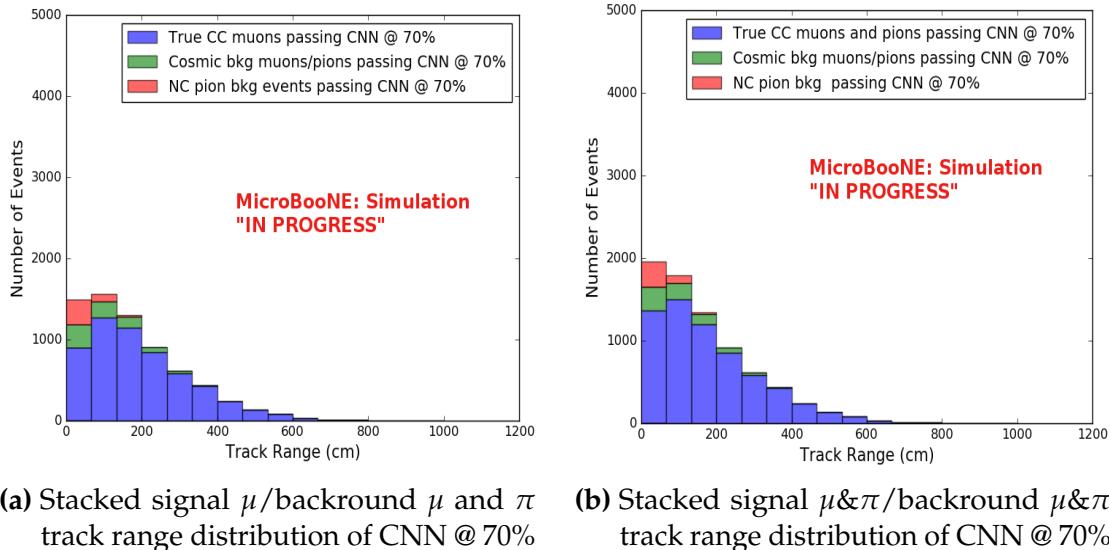


Figure 8.9: Track distribution comparisons of true CC muons plotted vs true CC muons and pions plotted

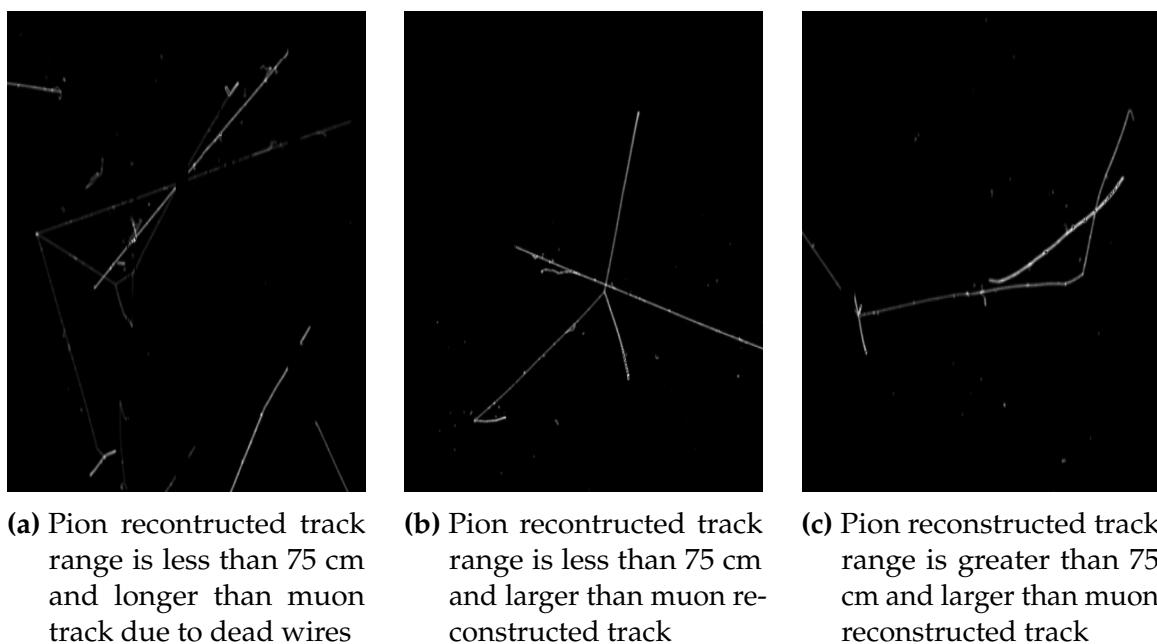


Figure 8.10: Images of true CC events where the pion was the tagged track candidate

		BNB + Cosmics		Cosmic Only	Signal: Cosmic Only
		Selection	MC Truth		
75 cm Cut passing rates	Generated Events	191362	45723	4804	1:22
	Track Containment	19391 (48%/10%)	11693 (45%/26%)	129 (38%/2.7%)	1:2.3
	track \geq 75 cm	6920 (36%/3.6%)	5780 (49%/13%)	17 (13%/0.4%)	1:0.6
CNN passing rates	Generated Events	188880	44689	14606	1:21
	Track Containment	19112 (/10%)	11554 (/26%)	302 (/2.1%)	1:1.73
	CNN cut @ 70% Probability	16502 (86%/8.7%)	10605 (92%/23%)	205 (68%/14%)	1:1.28
	CNN cut @ 83% Probability	7511 (46%/4.0%)	6142 (58%/14%)	32 (16%/0.2%)	1:0.4

Table 8.1: Comparing passing rates of CNN at different probabilities versus 75 cm track length cut: Numbers are absolute event counts and Cosmic background is not scaled appropriately. The BNB+Cosmic sample contains all events. The numbers in brackets give the passing rate wrt the step before (first percentage) and wrt the generated events (second percentage). In the BNB+Cosmic MC Truth column shows how many true ν_μ CC-inclusive events (in FV) are left in the sample. This number includes possible mis-identifications where a cosmic track is picked by the selection instead of the neutrino interaction in the same event. The CNN MC True generated events were scaled wrt the MC True generated events for the 75 cm cut passing rates due to only running over 188,880 generated events versus the 191362 generated events. The last column Signal:Cosmic only gives an estimate of the ν_μ CC events wrt the cosmic only background at each step. For this number, the cosmic background has been scaled as described in [?]. Note that these numbers are not a purity, since other backgrounds can't be determined at this step.

Signal	ν_μ CC events with true vertex in FV	#Events(Fraction)	#Events(Fraction)
		passing CNN @ 70% Probability	
Backgrounds		10605(35%)	#Events(Fraction) passing CNN @ 83% Probability
	Cosmics Only Events	13573(45%)	
	Cosmics in BNB Events	2249(7.4%)	
	NC Events	3412(11%)	
	ν_e and $\bar{\nu}_e$ Events	139(0.5%)	
	$\bar{\nu}_\mu$ Events	97(0.3%)	

Table 8.2: Signal and background event numbers at modified selection level with CNN cut estimated from a BNB+Cosmic sample and Cosmic only sample normalized to $5 * 10^{19}$ PoT. The last column gives the fraction of this signal or background type to the total selected events per CNN probability.

Table 8.1 shows the passing rates for the 75 cm track length cut and the CNN cut at 70% and 83%. The passing rates at the track containment level for the 75 cm track length cut compared to the CNN are comparable with only a 0.6% difference in the in time cosmic bin which may be due in part to the larger in time cosmic statistics used for the CNN dataset. These passing rates need to be comparable to then be able to compare the passing rates after the CNN cut to the 75 cm cut. Again, the same BNB+Cosmic sample was used for both selection I modified with 75 cm cut and selection I modified with CNN cut. As it stands, a CNN cut at 83% probability has

1363 a MC true CC event passing rate of 14% compared to the 13% passing rate of the 75
 1364 cm track length cut. The Signal:Cosmic Only background is also reduced from 1:0.6
 1365 to 1:0.4 The total passing rate is also higher than the 75 cm cut, 3.6% vs 4.0%. Table
 1366 8.2 shows the breakdown of signal and backgrounds for the CNN at the different
 1367 probabilities. We have a 61% signal passing rate with the CNN cut @ 83% versus the
 1368 53.8% signal passing rate of the 75 cm cut.

1369 Based on these numbers, the following performance values of the modified selec-
 1370 tion with 75 cm cut versus modified selection with CNN @ 83% probability cut were
 1371 calculated:

- 1372 • Efficiency: Number of selected true ν_μ CC events divided by the number of
 1373 expected true ν_μ CC events with interaction in the FV.
 - 1374 – Selection I modified: 13%
 - 1375 – Selection I modified with CNN cut @ 83% probability: 14%
- 1376 • Purity: Number of selected true ν_μ CC events divided by sum of itself and the
 1377 number of all backgrounds.
 - 1378 – Selection I modified: 53.8%
 - 1379 – Selection I modified with CNN cut @ 83% probability: 61%

1380 Lastly, figure 8.12 shows a more representative performance of the CNN. Due to
 1381 the fact that the CNN was trained on muons and pions, showing the performance
 1382 of CC muon events versus NC pion events with respect to CNN probability gives a
 1383 better picture of how the network is performing. Figure 8.12 shows that at 83% we
 1384 are below the 75 cm cut NC pion threshold and still above the CC muon threshold.
 1385 Using 83% probability not only reduced the NC pion background, it also dramatically
 1386 reduced the in time cosmics and cosmics in the BNB.

1387 8.1.3 Conclusions and Future Work

1388 It was shown that even though CNN10000 was trained with single particle generated
 1389 muons and pions, it performs fairly well at classifying track candidate images from
 1390 BNB+Cosmic events. Events have been regained below the 75 cm track length cut and
 1391 the momentum and track range distributions have similar shapes to the distributions of
 1392 Selection I original and modified. Efficiencies and purities were calculated for selection

Table 8: **Selection I: Modified** Signal and background event numbers at modified selection level estimated from a BNB+Cosmic sample and Cosmic only sample normalized to 5×10^{19} PoT. The last column gives the fraction of this signal or background type to the total 2189 selected events.

Signal	#Events	
ν_μ CC events with true vertex in FV	1168	53.8%
Backgrounds		
Cosmics only events	725	33.4%
Cosmics in BNB events	144	6.6%
NC events	75	3.5%
ν_e and $\bar{\nu}_e$ events	4	0.2%
$\bar{\nu}_\mu$ events	15	0.7%
ν_μ CC events with true vertex outside FV	40	1.8%

Figure 8.11: Snapshot of signal and background event numbers of Selection I modified from cc-inclusive note [?]

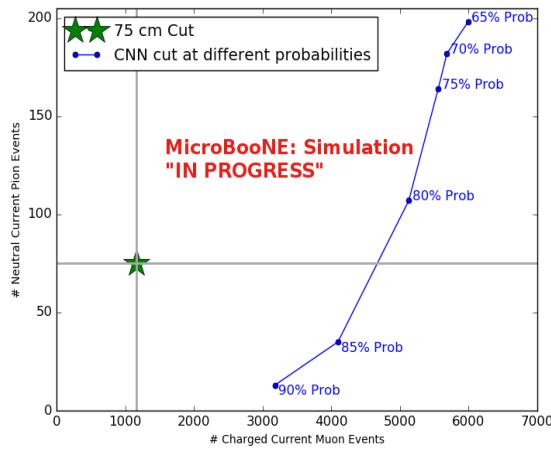


Figure 8.12: CNN performance of classified muons and pions compared to the already implemented 75 cm track length cut

¹³⁹³ I modified events before 75 cm track length cut with the CNN at 83% probability and
¹³⁹⁴ are 14% and 62% respectively. Although the CNN doesn't have separation between
¹³⁹⁵ muons and pions and although all particles passing CNN are classified as muon,
¹³⁹⁶ increasing CNN probability allows us to increase the purity as well as maintain an
¹³⁹⁷ efficiency comparable to the 75 cm track length cut all while recovering events below
¹³⁹⁸ that 75 cm cut. Out of the 6142 events that passed the CNN @ 83% 1470 events were
¹³⁹⁹ below the 75 cm cut, a recovery of 3.3% of data with a purity of 15%. Although
¹⁴⁰⁰ these numbers are low, it is an improvement from the selection I modified in both total
¹⁴⁰¹ efficiency and purity and an increase in phase space by recovering these events.

¹⁴⁰² **8.2 Classification using CNN100000**

¹⁴⁰³ All future classifications will be done using Selection I Modified CC-Inclusive Filter
¹⁴⁰⁴ because it has a higher efficiency and purity than Selection I Original CC-Inclusive
¹⁴⁰⁵ Filter. To reiterate, CNN100000 was trained using 20,000 images of each $\mu/\pi/p/\gamma/e$.
¹⁴⁰⁶ The results of using CNN100000 to classify BNB+Cosmics will be outlined below.

¹⁴⁰⁷ **8.2.1 Classification of MC data using Selection I Modified
1408 CC-Inclusive Filter**

¹⁴⁰⁹ **8.2.2 Classification of MicroBooNE data using Selection I Modified
1410 CC-Inclusive Filter**

¹⁴¹¹ **8.2.3 Comparing two CC-Inclusive Cross Section Selection Filters**

¹⁴¹² **Chapter 9**

¹⁴¹³ **Conclusion**

¹⁴¹⁴ Your Conclusions here.

¹⁴¹⁵

¹⁴¹⁶ Bibliography

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