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1

2                    $\mu/\pi$  separation using

3                   Convolutional Neural Networks

4                   for the MicroBooNE

5                   Charged Current Inclusive Cross Section

6                   Measurement

7

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$\mu$ BooNE







26

## Abstract

27

The purpose of this thesis was to use Convolutional Neural Networks  
28 (CNN) to separate  $\mu'$ s and  $\pi'$ s for use in increasing the acceptance rate  
29 of  $\mu'$ 's below the implemented 75cm track length cut in the Charged  
30 Current Inclusive (CC-Inclusive) event selection for the CC-Inclusive  
31 Cross-Section Measurement. In doing this, we increase acceptance  
32 rate for CC-Inclusive events below a specific momentum range.



33

## Dedication

34

I dedicate this dissertation to the two important women in my life; My  
35 wife and my mom. Both have been there cheering me on giving me strength  
36 and love as I worked towards the hardest accomplishment I've ever done.

37

Jessica Nicole Esquivel



38

## Acknowledgements

39        Of the many people who deserve thanks, some are particularly prominent, such as  
40        my supervisor...



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*"If they don't give you a seat at the table,  
bring a folding chair."*

— Shirley Chisholm



<sup>328</sup> **Chapter 1**

<sup>329</sup> **Introduction**

<sup>330</sup> This thesis will be a description of work done to further increase efficiency and purity  
<sup>331</sup> of the charged current inclusive cross section measurement using the MicroBooNE  
<sup>332</sup> detector. It will also describe the MicroBooNE detector, what neutrinos are, the  
<sup>333</sup> charged current inclusive cross section measurement and its importance as well as  
<sup>334</sup> convolutional neural networks and how they can be used in  $\mu/\pi$  separation. Chapter  
<sup>335</sup> 2 will talk about the background of neutrinos and the people and detectors that  
<sup>336</sup> discovered neutrinos as well as an in depth history of neutrino oscillation and the  
<sup>337</sup> discovery that neutrinos have mass.

<sup>338</sup> Chapter 3 will discuss the MicroBooNE experiment, specifically, how Liquid  
<sup>339</sup> Argon Time Projection Chambers work, the Light Collection System and the Electronic  
<sup>340</sup> and Readout Trigger systems. This chapter will also describe the Booster Neutrino  
<sup>341</sup> Beam sationed at Fermilab.

<sup>342</sup> Chapter 4 will discuss the work that was done to detect the first neutrinos seen in  
<sup>343</sup> the MicroBooNE detector and the software reconstruction efforts required to create an  
<sup>344</sup> automated neutrino ID filter that was used to find the first neutrinos and then was  
<sup>345</sup> later expanded on to create the charged current inclusive filter that will be discussed  
<sup>346</sup> in chapter 5

<sup>347</sup> Chapter 6 will give a brief description of what Convolutional Neural Networks are  
<sup>348</sup> and how it will be used for  $\mu/\pi$  separation in this selection. Chapter 7 will discuss  
<sup>349</sup> the hardware frameworks and training methods used to train multiple Convolutional  
<sup>350</sup> Neural Networks for use in the charged current inclusive cross section measurement.  
<sup>351</sup> Chapters 8 and ?? will discuss the results of using Convolutional Neural Networks on  
<sup>352</sup> monte-carlo and data to sift out charged current inclusive neutrino events.



# <sup>353</sup> Chapter 2

## <sup>354</sup> Neutrinos

### <sup>355</sup> 2.1 What are Neutrinos

<sup>356</sup> Neutrinos are one of the fundamental particles which make up the universe. They are  
<sup>357</sup> also one of the least understood. Neutrinos are not affected by the electromagnetic  
<sup>358</sup> forces because they do not have electric charge. Neutrinos are affected by a "weak"  
<sup>359</sup> sub-atomic force of much shorter range than electromagnetism, and are therefore able  
<sup>360</sup> to pass through great distances in matter without being affected by it. Until the late  
<sup>361</sup> 90's, neutrinos were thought to have no mass. Due to their mass, neutrinos are also  
<sup>362</sup> affected by gravity. Neutrinos are created by radioactive decay or nuclear reactions  
<sup>363</sup> such as the ones that happen in the sun, in nuclear reactors or when cosmic rays hit  
<sup>364</sup> atoms. There are three types of neutrinos,  $\nu_e$ ,  $\nu_\mu$  and  $\nu_\tau$  which correspond to their  
<sup>365</sup> charged lepton pairs.

<sup>366</sup> As previously stated, neutrinos are very weakly interacting; in fact, neutrinos can  
<sup>367</sup> pass unscathed through a wall of lead several hundred light-years thick. Because  
<sup>368</sup> neutrinos interact so rarely, studying neutrinos requires a massive detector and a  
<sup>369</sup> powerful neutrino source. With that being said, we can only infer their existence when  
<sup>370</sup> they interact in a detector. In a collision, distinct charged particles are produced with  
<sup>371</sup> each type of neutrino. An electron neutrino will create an electron, a muon neutrino  
<sup>372</sup> will create a muon, and a tau neutrino will create a tau. The track the particle leaves  
<sup>373</sup> in the detector is how one figures out what type of neutrino interaction was "seen".  
<sup>374</sup> Liquid Argon Time Projection Chambers are the newest type of detectors being used to  
<sup>375</sup> study neutrinos due to their excellent imaging and particle identification capabilities.

## <sup>376</sup> 2.2 History of Neutrinos

<sup>377</sup> The neutrino was first postulated by Wolfgang Pauli in 1931 to explain how beta  
<sup>378</sup> decay could resolve the conservation of energy, momentum and angular momentum  
<sup>379</sup> problem. Pauli suggested that this missing energy might be carried off, unseen, by a  
<sup>380</sup> neutral particle (he called neutron) which was escaping detection. James Chadwick  
<sup>381</sup> discovered a much heavier nuclear particle in 1932 that he also named neutron, leaving  
<sup>382</sup> two particles with the same name. Enrico Fermi was the first person to coin the  
<sup>383</sup> term neutrino (which means little neutral one in latin) in 1933 to fix this confusion.  
<sup>384</sup> Fermi's paper, which was published in 1934, unified Pauli's neutrino with Paul Dirac's  
<sup>385</sup> positron and Werner Heisenberg's neutron-proton model and his theory accurately  
<sup>386</sup> explained many experimentally observed results. Wang Ganchang first proposed the  
<sup>387</sup> use of beta capture to experimentally detect neutrinos and in 1959 Clyde Cowan and  
<sup>388</sup> Frederick Reines published their work stating that they had detected the neutrino.  
<sup>389</sup> The experiment called for antineutrinos created in a nuclear reactor by beta decay that  
<sup>390</sup> reacted with protons producing neutrons and positrons:  $\nu_e + p^+ \rightarrow n^0 + e^+$ . Once  
<sup>391</sup> this happens, the positron finds an electron and they annihilate each other and the  
<sup>392</sup> resulting gamma rays are detectable. The neutron is detected by neutron capture and  
<sup>393</sup> the releasing of another gamma ray. In 1962 Leon M. Lenderman, Melvin Schwartz  
<sup>394</sup> and Jack Steinberger were the first to detect interactions of the muon neutrino. The  
<sup>395</sup> first detection of the tau neutrino was announced in the summer of 2000 by the  
<sup>396</sup> DONUT collaboration at Fermilab. In the late 1960s, many experiments found that the  
<sup>397</sup> number of electron neutrinos arriving from the sun was around 1/3 to 1/2 the number  
<sup>398</sup> predicted by the Standard Solar Model. This became known as the solar neutrino  
<sup>399</sup> problem and remained unresolved for around thirty years. This problem was resolved  
<sup>400</sup> by the discovery of neutrino oscillation and mass. [1]

## <sup>401</sup> 2.3 Neutrino Oscillations

<sup>402</sup> Neutrino oscillation was first predicted by Bruno Pontecorvo. It describes the phe-  
<sup>403</sup> nomenon of a neutrino created with a specific lepton flavor (electron, muon or tau)  
<sup>404</sup> that is later measured to have a different flavor. Neutrino oscillation is important  
<sup>405</sup> theoretically and experimentally due to the fact that this observation implies that the  
<sup>406</sup> neutrino has a non-zero mass, which is not part of the original Standard Model of  
<sup>407</sup> particle physics. [2]

### 408 2.3.1 Solar Oscillations and the Solar Neutrino Problem

409 The solar neutrino flux derived from Bahcall's Standard Solar Model is shown in figure  
 410 2.1. Nuclear fusion and decay processes produce an abundant amount of neutrinos.  
 411 The standard solar model predicts that these reactions produce several groups of  
 412 neutrinos, each with differing fluxes and energy spectra. The figure also shows the  
 413 ranges of detection of existing solar neutrino experiments in different shades of blue  
 414 to illustrate that they sample different portions of the solar neutrino energy spectrum.  
 415 Three of these experiments, plus a new one, are discussed below.

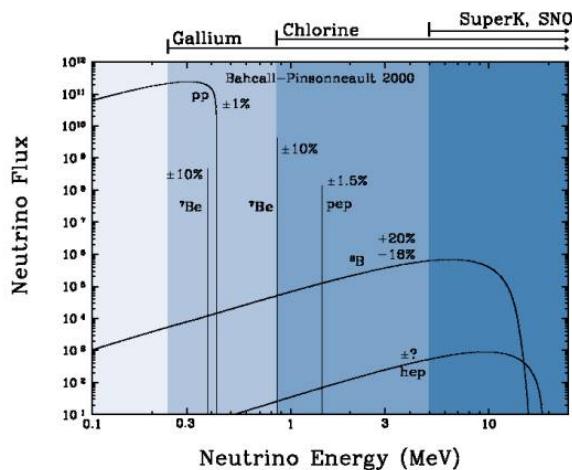


Figure 2.1: The Standard Solar Model

416 Since neutrinos rarely interact with matter, they pass through the sun and the earth  
 417 undetected. About 65-billion neutrinos from the sun stream through every square  
 418 centimeter on the Earth every second, yet we are oblivious to their passage in our  
 419 every-day lives. [3]

420 The first experiment to detect the effects of neutrino oscillation was the Ray Davis's  
 421 Homestake Experiment. The detector was stationed in the Homestake Gold Mine in  
 422 Lead, South Dakota. It was 1,478 meters underground and was  $380\text{ m}^3$ . The detector  
 423 was filled with perchloroethylene. Perchloroethylene was chosen because of its high  
 424 concentrations of chlorine. When an  $\nu_e$  interacted with chlorine-37 atom, the atom  
 425 would transform to argon-37 which was then extracted and counted. The neutrino  
 426 capture reaction is shown in equation 2.1. Davis observed a deficit of about 1/3  
 427 the flux of solar neutrinos that was predicted by Bahcall's Standard Solar Model.

- <sup>428</sup> The unexplained difference between the measured solar neutrino flux and model  
<sup>429</sup> predictions lead to the Solar Neutrino Problem. [4]

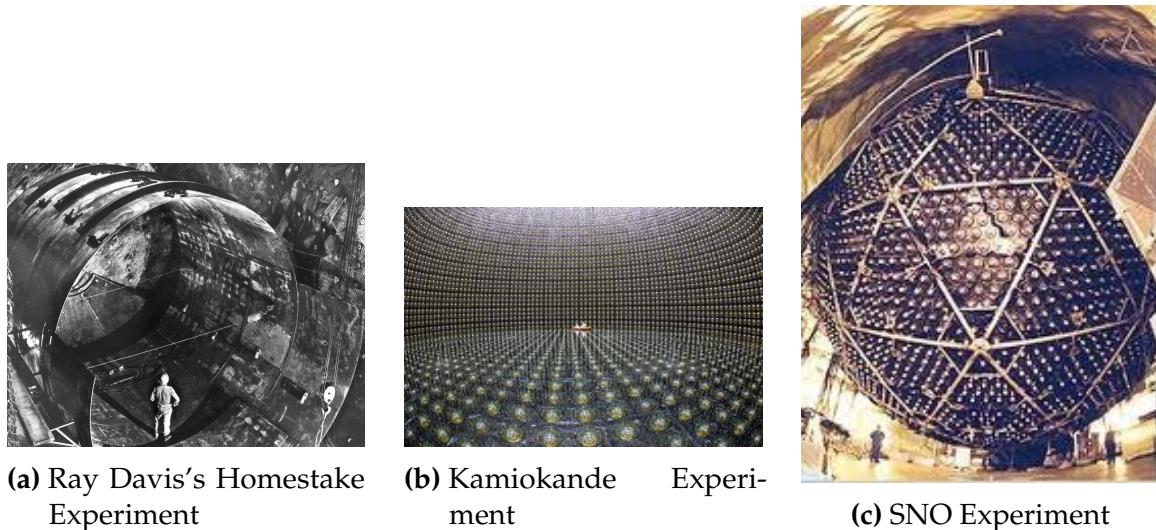


<sup>430</sup> While it is now known that the Homestake Experiment detected neutrinos, some  
<sup>431</sup> physicist were weary of the results. Conclusive evidence of the Solar Neutrino Problem  
<sup>432</sup> was provided by the Kamiokande-II experiment, a water Cherenkov detector with  
<sup>433</sup> a low enough energy threshold to detect neutrinos through neutrino-electron elastic  
<sup>434</sup> scattering. In the elastic scattering interaction the electrons coming out of the point of  
<sup>435</sup> reaction strongly point in the direction that the neutrino was traveling, away from the  
<sup>436</sup> sun. While the neutrinos observed in Kamiokande-II were clearly from the sun, there  
<sup>437</sup> was still a discrepancy between Kamiokande-II and Homestake; The Kamiokande-  
<sup>438</sup> II experiment measured about 1/2 the predicted flux, rather than the 1/3 that the  
<sup>439</sup> Homestake Experiment saw.

<sup>440</sup> The solution to the solar neutrino problem was finally experimentally determined  
<sup>441</sup> by the Sudbury Neutrino Observatory(SNO). The Ray Davis's Homestake Experiment  
<sup>442</sup> was only sensitive to electron neutrinos, and the Kamiokande-II Experiment was  
<sup>443</sup> dominated by the electron neutrino signal. The SNO experiment had the capability to  
<sup>444</sup> see all three neutrino flavors. Because of this, it was possible to measure the electron  
<sup>445</sup> neutrinos and total neutrino flux. The experiment demonstrated that the deficit was  
<sup>446</sup> due to the MSW effect, the conversion of electron neutrinos from their pure flavor  
<sup>447</sup> state into the second neutrino mass eigenstate as they passed through a resonance  
<sup>448</sup> due to the changing density of the sun. The resonance is energy dependent, and is  
<sup>449</sup> visible near 2MeV. The water cherenkov detectors only detect neutrinos above about  
<sup>450</sup> 5MeV, while the radiochemical experiments were sensitive to lower energy (0.8MeV  
<sup>451</sup> for chlorine, 0.2MeV for gallium), and this turned out to be the source of the difference  
<sup>452</sup> in the observed neutrino rates at the two types of experiments. Figure 2.2 shows  
<sup>453</sup> Homestake, Kamiokande-II and SNO experiments.

#### <sup>454</sup> MSW Effect

<sup>455</sup> The Mikheyev-Smirnov-Wolfenstein effect is a process which acts to modify neu-  
<sup>456</sup> trino oscillations in matter. The presence of electrons in matter changes the energy



**Figure 2.2:** Solar Neutrino Experiments

457 levels of the mass eigenstates of neutrinos due to charged current coherent forward  
 458 scattering of the electron neutrinos. This coherent forward scattering is similar to  
 459 the electromagnetic process with respect to the refractive index of light in a medium.  
 460 Because of this MSW Effect, neutrinos in vacuum have a different effective mass than  
 461 neutrinos in matter and because neutrino oscillations depend on the squared mass  
 462 difference of the neutrinos, the neutrino oscillations are different in matter than in  
 463 vacuum. This effect is important at the sun where electron neutrinos are produced.  
 464 The neutrinos of high energy leaving the sun are in a vacuum propagation eigenstate  
 465  $\nu_2$  that has a very small overlap with the electron neutrino  $\nu_e = \nu_1 \cos(\theta) + \nu_2 \sin(\theta)$   
 466 seen by the charged current reactions in Kamiokande-II and SNO. The discrepancy of  
 467 the deficit between SNO, Kamiokande-II and Homestake is due to the energy of the  
 468 solar neutrinos. The MSW effect "turns on" at about 2MeV and at lower energies, this  
 469 MSW effect is negligible. [5]

### 470 2.3.2 Atmospheric Oscillations and the Atmospheric Neutrino 471 Anomaly

472 Atmospheric neutrinos are neutrinos that stem from the decay hadrons coming from  
 473 primary cosmic rays. The dominant part of the decay chain is shown in equations 2.2  
 474 and 2.3

$$\pi^+ \rightarrow \mu^+ \nu_\mu \mu^+ \rightarrow e^+ \nu_e \bar{\nu}_\mu \quad (2.2)$$

475

$$\pi^- \rightarrow \mu^- \bar{\nu}_\mu \mu^- \rightarrow e^- \bar{\nu}_e \nu_\mu \quad (2.3)$$



**Figure 2.3:** Cosmic Ray Shower

476 Figure 2.3 shows the cosmic ray shower. In general, these neutrinos have energies  
 477 from 1GeV to 100s of GeV and the ratio of  $\nu_\mu$ s to  $\nu_e$ s equals to 2 (see equation 2.4)

$$R = \frac{(\nu_\mu + \bar{\nu}_\mu)}{(\nu_e + \bar{\nu}_e)} \quad (2.4)$$

478 There have been two types of detectors used to study atmospheric neutrinos: Water  
 479 Cherenkov detectors and tracking calorimeters. Super-Kamiokande is the detector we  
 480 will focus on. These atmospheric detector experiments measure the ratio of  $\nu_\mu$  to  $\nu_e$ .  
 481 They also measure the zenith angle distribution of the neutrinos. These experiments  
 482 report a double ratio (shown in equation 2.5). This double ratio is the ratio measured  
 483 in the detector to the ratio that's expected which is 2. If the double ratio equals to 1, the  
 484 data agrees with the prediction. Various measurements from multiple experiments  
 485 are shown in figure 2.4. Except for Frejus, all R measurements are less than 1. This  
 486 discrepancy between the predicted R and the measured R became known as the  
 487 Atmospheric Neutrino Anomaly.

$$R = \frac{(N_\mu / N_e)_{DATA}}{(N_\mu / N_e)_{SIM}} \quad (2.5)$$

488 Kamiokande-II has the the capability of measuring the direction of the incoming  
 489 neutrinos. The expectation of atmospheric neutrino detection is that the flux be

Experiment	Type of experiment	R
Super-Kamiokande	Water Cerenkov	$0.675 \pm 0.085$
Soudan2	Iron Tracking Calorimeter	$0.69 \pm 0.13$
IMB	Water Cerenkov	$0.54 \pm 0.12$
Kamiokande	Water Cerenkov	$0.60 \pm 0.07$
Frejus	Iron Tracking Calorimeter	$1.0 \pm 0.15$

**Figure 2.4:** Measurements of the double ratio for various atmospheric neutrino experiments

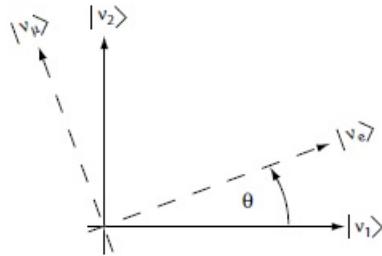
isotropic due to the fact that atmospheric neutrinos can reach the detector from all directions. Kamiokande-II noticed that muon-like data did not agree well with this expectation. At low energies approximately half of the  $\nu_\mu$  are missing over the full range of zenith angles. At high energies the number of  $\nu_\mu$  coming down from above the detector seems to agree with expectation, but half of the same  $\nu_\mu$  coming up from below the detector are missing. This anomaly can be easily explained by neutrino flavor oscillations. Due to the fact that the neutrino travels less distance coming straight down into the detector (about 15km) than coming up from the bottom of the detector(13000km) changes the probability of oscillation. The probability of oscillation for the muon neutrinos coming down into the detector is roughly zero, whereas for neutrinos coming up, the oscillation probability is  $\sin^2(2\theta)$ . Both the solar and atmospheric neutrino problems can be explained by neutrino oscillation so its fitting to derive this phenomenon mathematically. In the next two sections, two flavor and three flavor neutrino oscillation derivations will be explained.

### 2.3.3 Two Flavor Neutrino Oscillation Formulation

The flavor eigenstates can oscillate between each other because they are composed of an add-mixture of mass eigenstates( $\nu_1, \nu_2$ ). Figure 2.5 shows the mass and flavor eigenstates rotated by an angle  $\theta$  which is the mixing angle.

In matrix form the wavefunctions are:

$$\begin{pmatrix} \nu_\mu \\ \nu_e \end{pmatrix} = \begin{pmatrix} \cos\theta & \sin\theta \\ -\sin\theta & \cos\theta \end{pmatrix} * \begin{pmatrix} \nu_1 \\ \nu_2 \end{pmatrix} \quad (2.6)$$



**Figure 2.5:** The flavor eigenstates are rotated by an angle  $\theta$  with respect to the mass eigenstates

509 Applying the time evolution operator to  $\nu_\mu$ :

$$|\nu_\mu(t)\rangle = -\sin\theta|\nu_1\rangle e^{-i\frac{E_1 t}{\hbar}} + \cos\theta|\nu_2\rangle e^{-i\frac{E_2 t}{\hbar}} \quad (2.7)$$

510 where  $E_1 = \sqrt{p^2 c^2 + m_1^2 c^4}$  and  $E_2 = \sqrt{p^2 c^2 + m_2^2 c^4}$  and  $p_1 = p_2$ . For the time  
 511 being, let us assume  $\hbar = c = 1$ . With this assumption:  $E_1 = \sqrt{p^2 + m_1^2}$  and  $E_2 =$   
 512  $\sqrt{p^2 + m_2^2}$ . The next modifications is to assume neutrinos are relativistic:

$$\gamma = \frac{E}{m_o c^2} = \frac{\sqrt{p^2 c^2 + m_o^2 c^4}}{m_o c^2} \gg 1 \quad (2.8)$$

513 because of this,

$$p \gg m_o \quad (2.9)$$

514

$$E = \sqrt{p^2 + m_o^2} = p\sqrt{1 + m_o^2/p^2} \simeq p + \frac{1}{2}\frac{m_o^2}{p} \quad (2.10)$$

515 where the binomial expansion is used. Now  $E_1$  and  $E_2$  can be written as:

$$E_1 \simeq p + \frac{1}{2}\frac{m_1^2}{p} \text{ and } E_2 \simeq p + \frac{1}{2}\frac{m_2^2}{p} \quad (2.11)$$

516 Now applying all these assumptions back into equation 2.7 gives us:

$$|\nu_\mu(t)\rangle = -\sin\theta|\nu_1\rangle e^{-i\left(p + \frac{1}{2}\frac{m_1^2}{p}\right)t} + \cos\theta|\nu_2\rangle e^{-i\left(p + \frac{1}{2}\frac{m_2^2}{p}\right)t} \quad (2.12)$$

$$|\nu_\mu(t)\rangle = e^{-i\left(p+\frac{1}{2}\frac{m_1^2-m_2^2}{p}\right)t} (-\sin\theta|\nu_1\rangle + \cos\theta|\nu_2\rangle) \quad (2.13)$$

<sub>517</sub> Substituting  $\Delta m^2 = m_1^2 - m_2^2$  and  $t = \frac{x}{c} = x$  and  $e^{-iz} = e^{-i\left(p+\frac{1}{2}\frac{m_1^2}{p}\right)t}$  gives us:

$$|\nu_\mu(t)\rangle = e^{-iz} \left( -\sin\theta|\nu_1\rangle + \cos\theta|\nu_2\rangle e^{+ix\left(\frac{1}{2}\frac{\Delta m^2}{p}\right)} \right) \quad (2.14)$$

<sub>518</sub> Finding the Probability for a  $\nu_\mu \rightarrow \nu_e$ :

$$P(\nu_\mu \rightarrow \nu_e) = |\langle \nu_e | \nu_\mu(t) \rangle|^2 \quad (2.15)$$

<sub>519</sub> Remembering that  $\langle \nu_i | \nu_j \rangle = \delta_{ij}$

$$\langle \nu_e | \nu_\mu(t) \rangle = e^{-iz} \left( -\sin\theta\cos\theta + \sin\theta\cos\theta e^{\frac{i\Delta m^2 x}{p}} \right) \quad (2.16)$$

<sub>520</sub> Taking the absolute value squared gives us:

$$P(\nu_\mu \rightarrow \nu_e) = |\langle \nu_e | \nu_\mu(t) \rangle|^2 = e^{+iz} e^{-iz} \sin^2\theta \cos^2\theta \left( -1 + e^{\frac{i\Delta m^2 x}{p}} \right) \left( -1 + e^{\frac{-i\Delta m^2 x}{p}} \right) \quad (2.17)$$

<sub>521</sub> Since the neutrino is relativistic we can set  $p = E_\nu$  and change  $x = L$ . Also <sub>522</sub> recognizing the trigonometric relation  $(1 - \cos 2\theta)/2 = \sin^2\theta$  the above equation <sub>523</sub> becomes:

$$P(\nu_\mu \rightarrow \nu_e) = \sin^2 2\theta \sin^2 \left( \frac{\Delta m^2 L}{4E_\nu} \right) \quad (2.18)$$

<sub>524</sub> All that's left to do now is re-introduce  $\hbar$  and  $c$  doing this we get:

$$P_{\nu_\mu \rightarrow \nu_e}(L, E) = \sin^2 2\theta \sin^2 \left( 1.27 \Delta m^2 \frac{L}{E_\nu} \right) \quad (2.19)$$

<sub>525</sub> This equations has three important variables.

- The angle  $\theta$ : This angle, as mentioned before, is called the mixing angle. It defines the difference between the flavor and the mass eigenstates. When  $\theta = 0$  the mass and flavor eigenstates are identical and now oscillations occur.
- The mass squared difference,  $\Delta m^2$ : Again  $\Delta m^2 = m_1^2 - m_2^2$ . The reason this is an important variable is because it implies that for neutrinos to oscillate, neutrinos must have mass. Furthermore, the mass squared difference also tells us that the neutrino mass eigenstates must be different.
- L/E: This is the variable that is of most interest to experimental physicists due to the fact that it is the variable that we set. L is the distance between the source and detector and E is the energy of the neutrino. For a given  $\Delta m^2$ , the probability of oscillation changes with respect to L/E.

### 2.3.4 Three Flavor Neutrino Oscillation Formulation

Seeing the quantum mechanics involved in deriving the probability of a two flavor neutrino oscillation, it is now possible to formulate the three flavor neutrino oscillation. The three flavor neutrino oscillation formulation begins similarly to the two flavor, but there is the Pontecorvo-Maki-Nakagawa-Sakata matrix (PMNS) instead of the 2X2 matrix in the previous section. The PMNS matrix is show below:

$$\begin{pmatrix} c_{12}c_{13} & s_{12}c_{13} & s_{13}e^{-i\delta} \\ -s_{12}c_{23} - c_{12}s_{23}s_{13}e^{i\delta} & c_{12}c_{23} - s_{12}s_{23}s_{13}e^{i\delta} & s_{23}c_{13} \\ s_{12}s_{23} - c_{12}c_{23}s_{13}e^{i\delta} & -c_{12}s_{23} - s_{12}c_{23}s_{13}e^{i\delta} & c_{23}c_{13} \end{pmatrix} * \begin{pmatrix} e^{i\alpha_1/2} & 0 & 0 \\ 0 & e^{i\alpha_2/2} & 0 \\ 0 & 0 & 1 \end{pmatrix} \quad (2.20)$$

where  $c_{ij} = \cos\theta_{ij}$  and  $s_{ij} = \sin\theta_{ij}$

Following the same steps as before we get:

$$P_{\alpha \rightarrow \beta} = \delta_{\alpha\beta} - 4\sum Re(U_{\alpha i}^* U_{\beta i} U_{\alpha j} U_{\beta j}^*) \sin^2\left(\frac{\Delta m_{ij}^2 L}{4E}\right) 2\sum Im(U_{\alpha i}^* U_{\beta i} U_{\alpha j} U_{\beta j}^*) \sin\left(\frac{\Delta m_{ij}^2 L}{2E}\right) \quad (2.21)$$

The main things to notice here are  $\delta_{ij}$  which is the CP violating term and has not been measured yet, and  $\theta_{13}$  which has just been measured. CP violation is a violation

<sup>547</sup> of the postulated CP-symmetry. CP-symmetry states that the laws of physics should  
<sup>548</sup> be the same if a particle were to be exchanged with its antiparticle and then if the left  
<sup>549</sup> hand side of a decay were switched with the right hand side.

<sup>550</sup> **2.3.5 Reactor Oscillation**

<sup>551</sup> Many experiments have searched for oscillation of electron anti-neutrinos produced at  
<sup>552</sup> nuclear reactors. Such oscillations give the value of the parameter  $\theta_{13}$ . The KamLAND  
<sup>553</sup> experiment, started in 2002, has made a high precision observation of reactor neutrino  
<sup>554</sup> oscillation. Neutrinos produced in nuclear reactors have energies similar to solar  
<sup>555</sup> neutrinos, a few MeV. The baselines of these experiments have ranged from tens  
<sup>556</sup> of meters to over 100 km. On 8 March 2012, the Daya Bay team announced a  $5.2\sigma$   
<sup>557</sup> discovery that  $\theta_{13} \neq 0$ .



558 **Chapter 3**

559 **The MicroBooNE Experiment**

560 The purpose of this chapter is to discuss and understand the details of the MicroBooNE  
561 detector. A thorough understanding of MicroBooNE and the technology behind liquid  
562 argon time projection chambers is important for understanding results as well as  
563 understanding how images were made for use in deep learning efforts that will be  
564 outlined in later chapters.

565 **3.1 Liquid argon time projection chambers**

566 Liquid Argon Time Projection Chambers (LArTPCs) are an exciting detector technol-  
567 ogy that provide excellent imaging and particle identification, and are now being  
568 used to study neutrinos. The Time Projection Chamber (TPC) was first invented by  
569 Nygren in 1974 [?] and the proposal for a LArTPC for neutrino physics was made  
570 by Rubbia [?] in 1977 with the ICARUS collaboration implementing this concept [?].  
571 A LArTPC is a three-dimensional imaging detector that uses planes of wires at the  
572 edge of an active volume to read out an interaction. When a neutrino interacts with an  
573 argon atom, the charged particles that are produced ionize the LAr as they travel away  
574 from the interaction. By placing a uniform electric field throughout the LAr volume,  
575 the ionization is made to drift towards a set of anode planes, which consist of wires  
576 spaced very closely together collecting the ionized charge, which is subsequently read  
577 out by electronics connected to the anode wires. The collected ionization creates a  
578 spatial image of what happened in the detector on each anode plane. The position  
579 resolution of the interaction along the beam direction (perpendicular to drift direction)  
580 relies on the wire pitch, while the resolution in drift direction is dependent on the

581 timing resolution of the electronics used and the longitudinal diffusion in the volume.  
 582 The drift time of the ionization relative to the time of the original signal allows the  
 583 signal to be projected back along the drift coordinate, hence the name LArTPC. Having  
 584 very small distances between each wire within an anode plane allows for very fine  
 585 granularity and detail to be captured, and having multiple wire planes at different  
 586 angles provides independent two-dimensional views that can be combined into a  
 587 three dimensional picture of the interaction. Once the charge signal is created on the  
 588 anode planes, software analysis packages identify particles in the detector by using  
 589 deposited energy on the wires along their track length. The 30 year development of the  
 590 ICARUS detector has led to LArTPCs being used as cosmic ray [?], solar neutrino [?]  
 591 and accelerator neutrino [?] detectors. The ArgoNeuT experiment at Fermilab was  
 592 the first United States based liquid argon neutrino program that has since produced  
 593 short-baseline  $\nu - Ar$  cross-section measurements in the NUMI beamline [?]. The  
 594 MicroBooNE experiment is the second experiment in the US based LArTPC neutrino  
 595 program and will be discussed thoroughly in the next sections. The next phases of  
 596 the liquid argon neutrino program are under way and are the Fermilab Short Base-  
 597 line Neutrino (SBN) program [?] and the Deep Underground Neutrino Experiment  
 598 (DUNE) [6]. The SBN program will include three LArTPC detectors, including the  
 599 MicroBooNE detector, on the Booster Neutrino Beam (BNB) to do multiple-baseline  
 600 oscillation measurements. The detector closest to the beam will be the 40 ton Short  
 601 Baseline Neutrino Detector (SBND) [?] at 150 m and the detector furthest is the 600 ton  
 602 ICARUS T600 [?] detector positioned at 600 m. The DUNE collaboration will deliver  
 603 a 30 GeV neutrino beam 1300 km from Fermilab to a 34 kiloton LArTPC detector  
 604 at Homestake, SD. DUNE will study the leptonic CP phase,  $\delta_{cp}$ , as well as measure  
 605 neutrino and antineutrino oscillations.

## 606 3.2 The MicroBooNE Time Projection Chamber

607 MicroBooNE (Micro Booster Neutrino Experiment) is a 89 ton active volume (180 ton  
 608 total mass) LArTPC which is then inserted into a cylindrical cryostat on axis of the  
 609 Booster Neutrino Beam (BNB) stationed at Fermilab in Batavia, Illinois. Understanding  
 610 LArTPC technology and detector physics is necessary to build a LArTPC the size of  
 611 DUNE, and MicroBooNE has made many advances in developing this technology [7]  
 612 [8].

MicroBooNE's Time Projection Chamber (TPC) is 10.3 m long (beamline direction), 2.3 m high and 2.5 m wide (which corresponds to the drift distance). The TPC is shown in figure ?? . MicroBooNE is the largest LArTPC currently running in the world [9]. This LArTPC has 3 wire planes: 1 plane that collects the ionization in the wires and is  $0^\circ$  to the vertical with 3456 wires spaced 3 mm apart, and 2 planes where the ionization drifts passed and induces a signal at  $\pm 60^\circ$  to the vertical each with 2400 wires also spaced 3 mm apart. Each plane has a spacing also of 3 mm from eachother. The first two planes are the induction planes and the last is the collection. The 270 V/cm electric field of the TPC is created using 64 stainless steel tubes shaped into rectangles around the TPC and held in place by G10 to form a field cage. The cathode is charged at a high voltage of -70 kV and this voltage is stepped down across the field cage tubes using a voltage divider chain with an equivalent resistance of  $240\text{ M}\Omega$  between the tubes. The field cage tubes are separated by 4 cm from center to center. The electron drift distance is 2.5 m in the x direction with a drift time of 2.3 ms. Maintaining high charge yield is done by continuously recirculating and purifying the argon. The purity is monitored using MicroBooNE's light collection system. Another use of the light collection system is initial timing and drift coordinate of the interaction.

MicroBooNE's light collection system is a crucial part for 3D reconstruction of all particle interactions in the LArTPC. The initial interaction time,  $t_0$ , and initial drift coordinate,  $x_0$ , are not known from the TPC alone. For beam events, the accelerator clock is used to determine  $t_0$  of the interaction and the  $x_0$  can be inferred using drift time. Non-beam events, however, do not have this capability, which is why scintillation light from an interaction is used. The  $\nu - Ar$  interaction produces scintillation light which is collected by photomultiplier tubes (PMTs) which allows the exact time,  $t_0$  of the neutrino interaction to be determined. The scintillation light created propagates within nanoseconds to the light collection system compared to the milliseconds it takes the ionized electrons from the interaction to reach the anode wire planes. Therefore we can precisely know where along the drift direction the particle interaction first took place. The scintillation light is also localized, so combining the PMT information with the wire plane information allows for cosmic background rejection happening outside the beam timing window.

The light collection system is made up of 32 Hamamatsu R5912-02mod cryogenic PMTs with a diameter of 8-inches. The PMTs are located behind the 3 wire anode planes and provides 0.85% photocathode coverage. Each PMT has an acrylic plate mounted in front of it that is coated with a wave-length shifting material called TPB.

648 The acrylic plates take in the scintillation light, at 128 nm, and re-emits it visible  
649 wavelengths visible to the PMTs, with a peak at 425 nm.

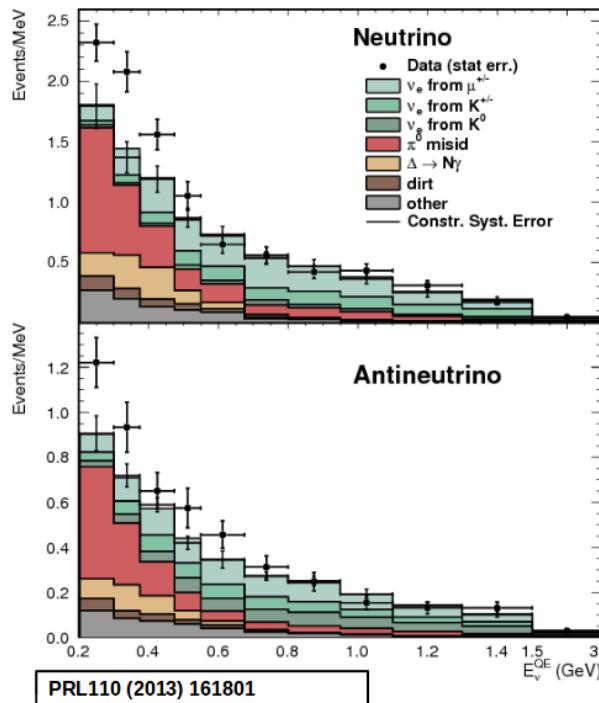
650 Both the light collection system and the TPC create analog signal that is read out and  
651 digitized by the electronics system. The process requires amplification and shaping of  
652 the signal which then goes to the data acquisition (DAQ) software for writing of the  
653 digitized data to disk. The anode plane wires are connected to detector specific circuit  
654 boards (ASICS) that are submerged and operate inside the liquid argon volume. These  
655 ASICS send amplified signal to 11 feed-throughs where further amplification of the  
656 signal happens outside the cryostat. The signal is received by custom LArTPC readout  
657 modules distributed over nine readout crates which do the digitization. The TPC wires  
658 are digitized at 16 MHz then downsampled to 2 MHz. The TPC system reads out 4  
659 frames of wire signal data per event, 1 frame before a trigger and 2 frames after the  
660 triggered frame. The four frames allows for identification of a neutrino interaction as  
661 well as cosmic background rejection. The process of digitization is similar for the light  
662 collection system. Each PMT signal undergoes a shaping with a 60 ns peaking time  
663 for digitization of multiple samples. The digitization occurs at 64 MHz but are not  
664 read out continuously during the TPC readout time. Only shaped PMT signal samples  
665 above a small threshold are read out and saved. Both the TPC and PMT readouts are  
666 initiated via triggers on a separate trigger board located in a warm electronics crate.  
667 The timing trigger is created by a timing signal from the BNB accelerator which is  
668 shaped and sent to the trigger board. The PMT trigger is generated when the PMT  
669 signal multiplicity is greater than 1 and the summed PMT pulse-height is more than 2  
670 photo-electrons summed up over all PMT channels. When the trigger board gets both  
671 a timing trigger and a PMT trigger in coincidence, at BNB trigger is then generated by  
672 the board. This signal is then passed to all readout crates initiating the readout of data.  
673 The data is then sent to the DAQ software which then saves the data to disk into one  
674 event memory.

### 675 3.3 MicroBooNE's Physics Goals

#### 676 3.3.1 The low-energy excess

677 The primary goal of the MicroBooNE experiment is to study and investigate the low-  
678 energy excess seen in MiniBooNE 3.1. MicroBooNE has the capability of confirming or

denying this excess as electrons or photons due to the detector being in the same beam, having a similar baseline, and lastly the detector being able to clearly distinguish between electrons and photons. LArTPCs use the topology of events as well as energy loss near the vertex to differentiate between single  $e^-$  tracks and photon-induced induced pair production  $\gamma \rightarrow e^+ e^-$ , which wasn't possible in MiniBooNE, a Cherenkov detector. This technique has been shown in the ArogoNeuT detector [?] and a side by side comparison of both event types in a LArTPC can be seen in figure ?? . An excess in electrons would point towards new oscillation physics beyond the standard model, while photons would be within the standard model. MicroBooNE will observe a  $4-5\sigma$  signal.



**Figure 3.1:** Low Energy excess seen in MiniBooNE

### 3.3.2 Cross sections

MicroBooNE's neutrino cross-section program will be the first  $\nu - Ar$  cross-section in the 1 GeV energy range and one of only a few cross-section measurements of  $\nu - Ar$  in the world. MicroBooNE is also the first liquid argon detector to collect the highest statistics sample of neutrino interactions. Investigating final-state-interactions in the 1GeV energy range provides information about short range nuclear correlations that affect the interpretations of neutrino oscillation experiment data.

One of the cross-section measurements MicroBooNE can make is an inclusive charged-current cross-section measurement (referred to as CC-inclusive). CC-inclusive events consist of a neutrino exchanging a  $W^\pm$  boson with an argon atom, producing a charged lepton and any number of other final state particles. In MicroBooNE's case, a CC-inclusive event will mostly have a defining muon track coming out of the vertex due to our neutrinos being predominately  $\nu_\mu$ s. A cross-section measurement is the energy dependent probability of  $\nu - Ar$  interaction in the detector. Cross-sections however are independent of the intensity or focus of the particle beam so they can be compared among different experiments. A background for a CC-inclusive cross-section measurement are the neutral-current events that contain a pion. It is possible to have a neutral current interaction with a  $\pi + p$  event signature that looks like a charged current  $\mu + p$  event. Reconstruction tools implemented to date don't efficiently separate muons from pions. A common way to separate these two particles species is to implement a track length cut. On average, muons tend to have longer track lengths in LArTPCs so by requiring that the hypothesized lepton be above a threshold track length, it is possible to increase signal to background.

### 3.3.3 Liquid argon detector development

The last physics goal for the MicroBooNE collaboration is to provide important information regarding LArTPC technology. Being the first in large scare LArTPCs in the US, MicroBooNE will be able to provide improvements to High Voltage (HV) distribution, Noise Characterization [?], and Michel Electron Reconstruction [8].

## 3.4 The Booster Neutrino Beam

The MicroBooNE detector is stationed at Fermi National Accelerator Laboratory (FNAL) where it receives neutrinos from both the Booster Neutrino Beam (BNB) and Neutrinos from the Main Injector (NuMI) beams. MicroBooNE is on-axis for the BNB and off-axis by 135 mrad for NuMI. For the purpose of this analysis, only data from the BNB was used. This section will discuss how neutrinos are created using the BNB. How these neutrinos are produced as well as their flux through the MicroBooNE detector is necessary for any analysis because of the systematic uncertainties the beam

<sup>725</sup> introduces to a measurement. An aerial view of fermilab as well as the BNB is shown  
<sup>726</sup> in figure 3.2

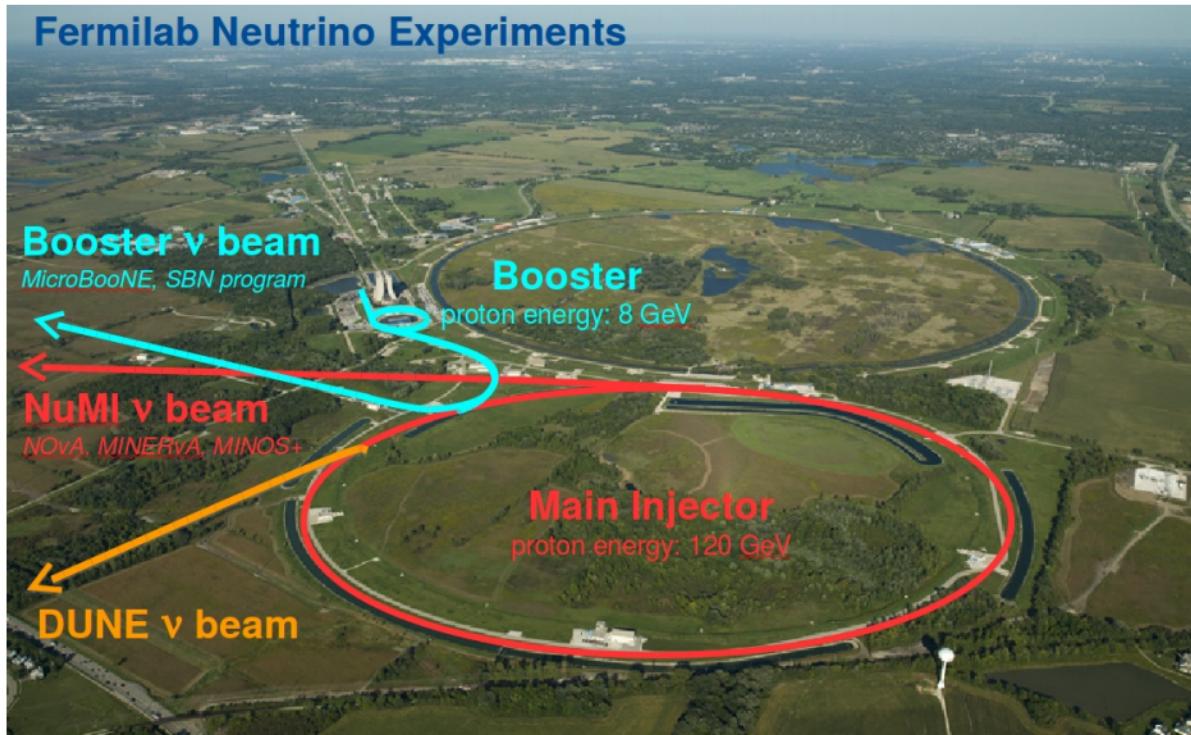


Figure 3.2: Arial view of the Main Injector and the Booster Neutrino Beam at Fermilab

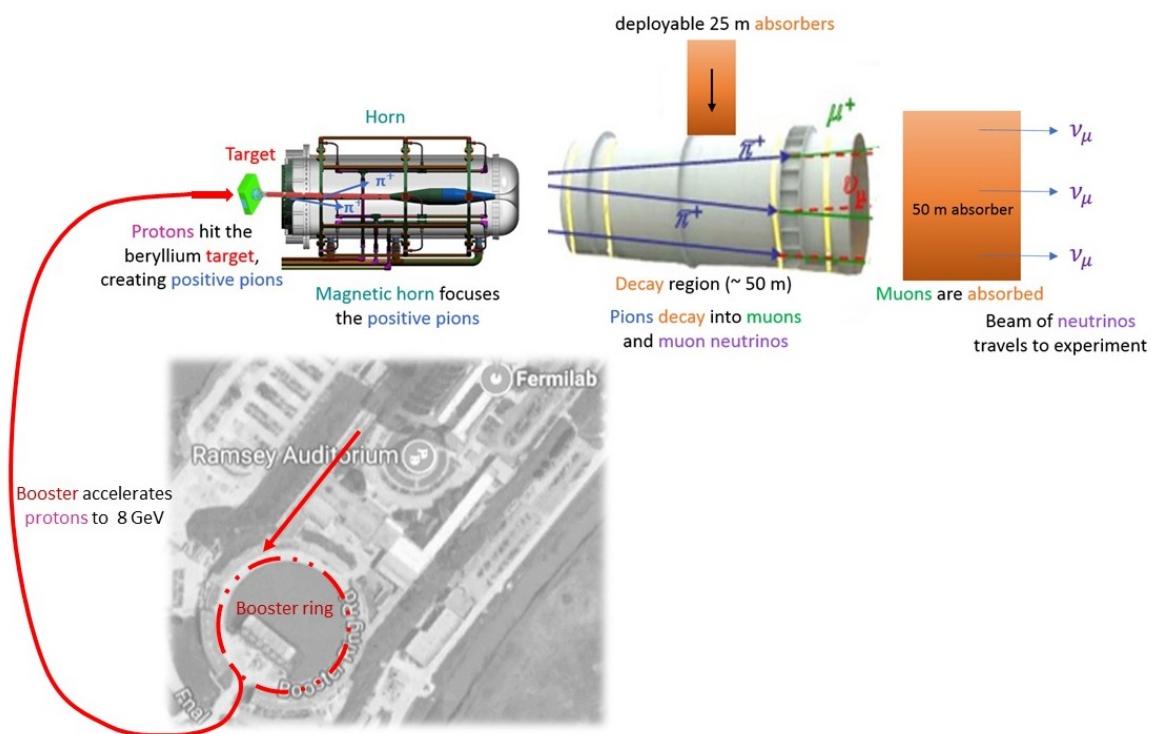
### <sup>727</sup> 3.4.1 Creating the Booster Neutrino Beam

<sup>728</sup> The BNB is a very pure  $\nu_\mu$  beam, with only 0.6% contamination from  $\nu_e$ s. The energy  
<sup>729</sup> also peaks around 700 MeV which is desired based on the probability of oscillation  
<sup>730</sup> equation which depends on the the value of  $L/E$ , where  $L$  is the distance of the  
<sup>731</sup> detector from the neutrino beam and  $E$  is the energy of the neutrino beam.  $L/E$  was  
<sup>732</sup> chosen to increase the probability of seeing neutrino oscillations in the MiniBooNE  
<sup>733</sup> Low Energy Excess (LEE) range based on the probability of oscillation equation, which  
<sup>734</sup> is  $P_{\nu_\mu \rightarrow \nu_e}(L, E) = \sin^2 2\theta \sin^2 \left( 1.27 \Delta m^2 \frac{L}{E} \right)$ . The BNB collides 8.9 GeV/c momentum  
<sup>735</sup> protons from the FNAL booster synchrotron into a beryllium target which produces a  
<sup>736</sup> high flux of neutrinos. The protons originate from  $H^2$  gas molecules that are turned  
<sup>737</sup> into  $H^-$  ions by a Cockcroft-Walton generator shown in figure ???. The  $H^-$  initially are  
<sup>738</sup> accelerated to 1MeV kinetic energy and are then passed to a linear accelerator using  
<sup>739</sup> alternating electromagnetic fields to increase their energy to 400MeV. The ions are  
<sup>740</sup> stripped of electrons by passing them through a carbon foil. The protons are bunched

741 into beam spills which contain  $4 * 10^{12}$  protons in a  $1.6 \mu\text{s}$  time window per spill. It's  
742 at this point that the protons are directed towards the beryllium target. The amount  
743 of protons directed towards the target (POT) is measured by two toroids upstream of  
744 the target with an error of 2%. Beam intensity, timing, width, position, and direction  
745 are monitored by beam position monitors, multi-wire chamber and resistive monitors.  
746 The beryllium target is 71.1 cm long, 1.7 proton interaction lengths, and is 0.51 cm in  
747 radius. The target is located inside a larger focusing electromagnet called the horn.  
748 The horn is an aluminum alloy pulsed toroidal electromagnet. The pulsed current  
749 peaks at 170 kA with a time-width of  $143 \mu\text{s}$  which coincides with the protons arriving  
750 on the target. The current flows from the inner conductor to the outer conductor  
751 with a maximum magnetic field of 1.5 Tesla. The magnetic field focuses the charged  
752 secondary particles produced by the p-Be interactions. The direction of current can be  
753 switched to change the polarity of the secondary particles being focused creating a  
754 beam of either primarily neutrinos, with positively charged secondary particles, or  
755 antineutrinos.

756 Further down the beamline is a concrete collimator which absorbs particles not  
757 necessary to the neutrino flux. The collimator is 214 cm long and 30 cm in radius.  
758 After the collimator comes a 45 meter long, 1 meter radius, air-filled cylindrical decay  
759 region which then ends in a beam-stop made of steel and concrete. The beam-stop  
760 contains an array of gas proportional counters to detect muons. The BNB is shown in  
761 figure 3.3.

762 **3.5 Event Reconstruction**



**Figure 3.3:** Aerial view of the Main Injector and the Booster Neutrino Beam at Fermilab

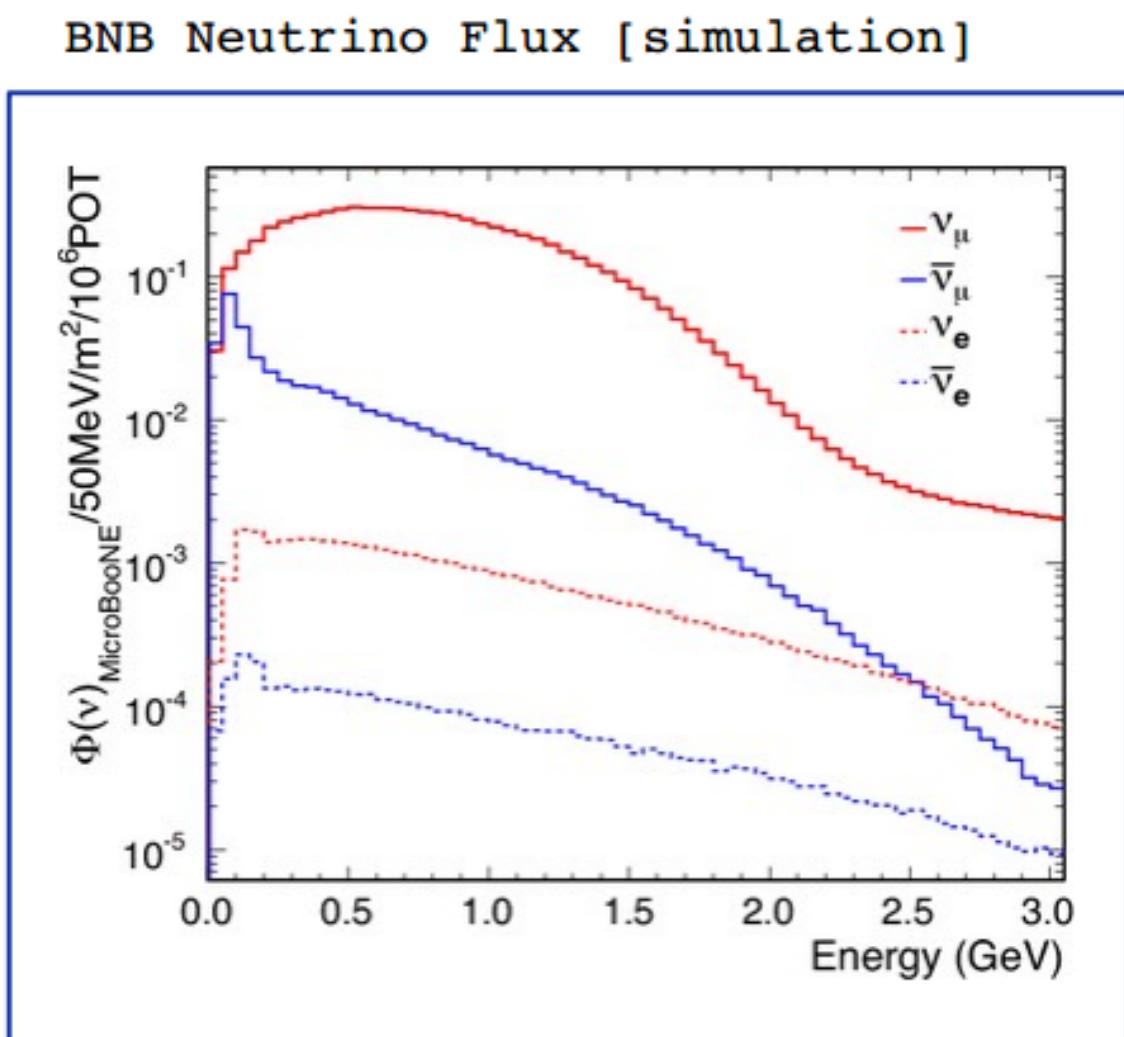


Figure 3.4: Energy spectrum of the Booster Neutrino Beam at Fermi National Laboratories

# 763 Chapter 4

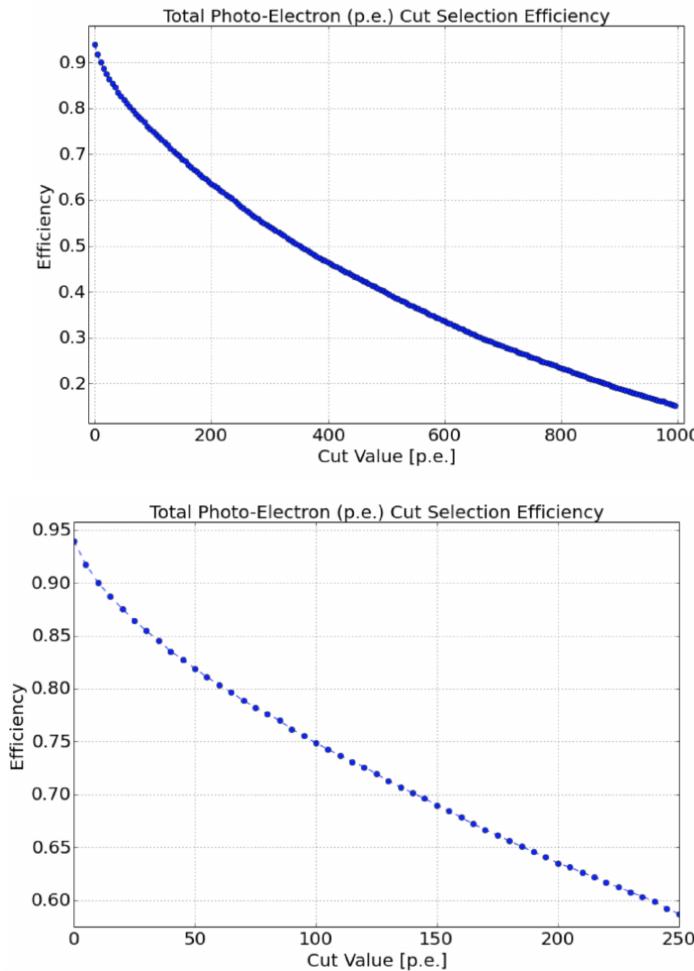
## 764 Neutrino Identification: Finding 765 MicroBooNE's first Neutrinos

766 The goal of the Neutrino Identification analysis was to positively identify BNB neutrino  
767 interactions in the MicroBooNE detector collected during the first days of running.  
768 Neutrino event candidates were identified in part by using a cut on detected flash of  
769 scintillation light during the  $1.6 \mu\text{s}$  beam-spill length of the BNB as well as identifying  
770 reconstructed object from the TPC that are neutrino like. After this selection, 2D  
771 and 3D event displays were used for verification of the selection performance. This  
772 selection was targeted to reduce the ratio of neutrino events to cosmic-only events from  
773 the initial 1 neutrino to 675 cosmics to a ratio of 1 to 0.5 or better which is equivalent to  
774 a background reduction by a factor of 1000 or more. These selected events were used  
775 for MicroBooNE's public displays of neutrino interactions. A clearly visible neutrino  
776 interaction with an identifiable vertex and at least 2 tracks originating from the vertex  
777 was what the analysis focused on. This analysis wasn't optimized for high purity  
778 or efficiency, but rather for very distinguishable neutrino interactions that could be  
779 identified by the public.

### 780 4.1 Flash Finding

781 Flash finding is the first step used in finding neutrino interactions. This section will  
782 detail how optical information is reconstructed as well as analysis scripts and event  
783 filters were used.

784 **4.1.1 Flash Reconstruction**

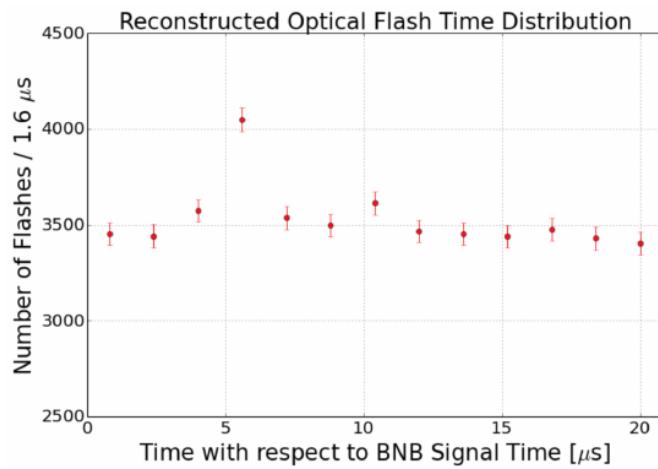


**Figure 4.1:** Efficiency for selecting beam events as a function of minimum total PE cut for all PE cuts as well as zoomed into interesting region.

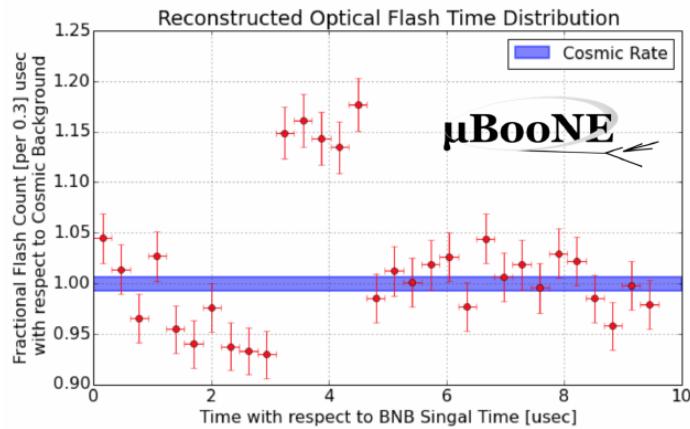
785 A flash is described as a collection of light seen at the same time within the detector.  
786 They are then reconstructed by identifying signal from the PMTs above a specific  
787 photoelectron (PE) threshold. These signals are called optical hits. Optical hits from  
788 all the PMTs are then accumulated into  $1\text{ }\mu\text{s}$  bins of time. If a specific bin is above a  
789 set PE threshold, then the optical hits that overlap in time are the labeled as the hits  
790 from the flash. All flash reconstructed properties like average time and x/y positions  
791 are then found via the flash labeled optical hits. The total size of the flash is found by  
792 summing up the total number of photoelectrons from all PMTs. Neutrino interactions  
793 and cosmic muons will have a larger flash size compared to noise and other low-energy  
794 backgrounds, therefore a total PE cut is used to reject these backgrounds. A total PE

<sup>795</sup> cut of 50 PE was deemed sufficient for this analysis. Figure 4.1 show the total PE  
<sup>796</sup> versus the selection efficency of selecting neutrino beam events.

### <sup>797</sup> 4.1.2 Beam Timing



(a) Predicted distribution of flash times with respect to trigger time for 1 day of data taking at nominal rate and intensity



(b) Measured distribution of flash times with a 50 PE threshold cut, with respect to trigger time. Shown as a ratio to the expected cosmic rate from off-beam data. A clear excess from neutrinos is visible between 3- 5  $\mu$ s after the trigger time.

<sup>798</sup> It is necessary to get the specific time from flashes if one uses flashes to filter out  
<sup>799</sup> neutrino interactions coincident with the neutrino beam spill period and background.  
<sup>800</sup> Before a filter can be applied, an understanding of the timing of the trigger and PMT

801 readout with respect to the arrival of neutrinos from the BNB. To do this, a  $1.6 \mu\text{s}$   
 802 window near the expected beamtime was created and verified by finding that the  
 803 number of flashes was significantly above the cosmic-ray background flashes. Beam  
 804 data during the first week of running, October 16th 2016 through October 22nd 2016  
 805 and were used for a timing measurement. The total POT uses corresponds to roughly  
 806 24 hours of data taking at nominal intensity ( $4 \times 10^{12} \text{ ppp}$ ) and a 5 Hz repetition rate.  
 807 Figure 4.2a shows size of the expected neutrino signal in time using Monte Carlo  
 808 predictions and figure 4.2b shows the neutrino signal in data. The intensity in data is  
 809 lower, however there can still be seen a significant excess above data.

810 **4.1.3 Event Rates**

811 Applying a 50 PE threshold cut inside a  $1.6 \mu\text{s}$  window reduces the cosmic-ray passing  
 812 rate to 0.8%. With a 5 Hz beam rate, this corresponds to 135 cosmics passing per  
 813 hour. The neutrino passing rate for this filter is about 22 events per hour. To further  
 814 increase the neutrino to cosmic ratio, TPC topology cuts were implemented and will  
 815 be discussed in the following section.

816 **4.2 TPC Topology Selection**

817 In order to further reduce the background of cosmic events, two independent selection  
 818 streams using TPC wire data reconstruction was implemented. The first using 2D  
 819 reconstructed clusters, and the second using 3D reconstructed tracks. Both streams  
 820 look for neutrino interactions in the active TPC volume which are identifiable by two  
 821 or more tracks originating from the same vertex.

822 Both 2D and 3D channels were optimized using monte carlo simulation which  
 823 used a 128 kV cathode voltage. Passing rates were calculated using a 0.008 efficiency  
 824 factor for cosmic events passing to simulate the flash finding described in section 4.1.  
 825 This efficiency factor was an overestimation and was just used to get a general feel of  
 826 what signal and background rates we would actually see in data.

### <sup>827</sup> 4.2.1 Cosmic Tagging

<sup>828</sup> The first step in TPC selection was based on the geometry of cosmic tracks in an event.  
<sup>829</sup> The cosmic ray muon geometry tagger runs on 3D tracks and assigns a score to each  
<sup>830</sup> reconstructed track on the likeliness of the track originating from a cosmic. The cosmic  
<sup>831</sup> scores are detailed below:

- <sup>832</sup> • 1: The track is tagged as entering or exiting the TPC
- <sup>833</sup> • 0.95: The track is a delta ray associated with a tagged track
- <sup>834</sup> • 0.5: The track is either entering or exiting, but not both
- <sup>835</sup> • 0.4: The track is entering or exiting through the Z boundary
- <sup>836</sup> • 0: The track isn't tagged

<sup>837</sup> Clusters are assigned either a 0 or 1, 1 being a cosmic. In simulation, 90% of cosmics  
<sup>838</sup> are tagged as cosmics. These tracks are no longer considered when looking for a  
<sup>839</sup> neutrino topology. Requiring that the tracks be contained in turn affects the neutrino  
<sup>840</sup> efficiency by 20%. The algorithm checks that each track is contained within a boundary  
<sup>841</sup> region of 10 cm from all sides of the TPC. This boundary region was optimized via  
<sup>842</sup> handscanning of experimental data.

<sup>843</sup> As can be expected, cosmic tagging is more efficient in the 3D channel (tracks) than  
<sup>844</sup> the 2D channel (clusters) because the reconstructed tracks can use the full 3D position  
<sup>845</sup> information of the entering and exiting points while the 2D channel mainly use the  
<sup>846</sup> reconstructed x position of the cluster which is associated to timing.

<sup>847</sup> Cosmic tagging uses timing information to reject tracks and clusters that are outside  
<sup>848</sup> of drift window. The drift window for 128 kV is  $1.6 \mu\text{s}$  while for 70 kV, the actual  
<sup>849</sup> voltage MicroBooNE is running at, is  $2.3 \mu\text{s}$ . Due to this variation between simulation  
<sup>850</sup> and data, we expect to see  $2.3/1.6 = 1.44$  times more cosmic induced tracks or clusters  
<sup>851</sup> in the drift window.

### <sup>852</sup> 4.2.2 2D Cluster Selection

<sup>853</sup> This selection was spearheaded by myself and Katherine Woodruff. After looking at  
<sup>854</sup> experimental cosmics data, 2D clustering performs well, while 3D track reconstruction  
<sup>855</sup> is affected by more variations in simulation, for example noise filters. This was the

856 motivation for having a selection only on 2D clusters in the collection (Y) plane. As  
 857 stated previously, the goal of this analysis was to find identifiable neutrino interactions  
 858 for use in public event displays, in future analyses, the 3D track reconstruction has  
 859 been modified to further increase the tracking efficiency and has more information  
 860 than just the clusters. For this analysis, however, 2D cluster information was sufficient  
 861 enough for neutrino selection.

862 **Primary Cuts**

863 The first cuts were used to select which clusters to consider. First the clusters must  
 864 have at least ten hits on the collection plane and have a cosmic tagging score < 0.4.  
 865 Only events that have at least two clusters that satisfy these primary cuts continue on.

866 After the initial cosmic tagging is applied, the following cuts are used to further  
 867 separate identifiable neutrinos from background cosmics.

868 The next cut was to remove long, vertical clusters. This was applied after seeing  
 869 that most cosmic induced clusters passing were long with high angles, while neutrino  
 870 induced clusters were mainly forward going. We required a good cluster to either  
 871 have a projected start angle less than 30 degrees from the z axis or be less than 200  
 872 wires long. The length cut was added to make sure we don't cut any short high angle  
 873 clusters that can correspond with a proton, or other highly ionizing particle associated  
 874 with a long muon cluster. The 200 wire cut roughly equates to 0.6 m in the z direction,  
 875 with a 3 mm wire pitch. Also, the projected angle is defined by  $\tan \alpha = \Delta T / \Delta W$  where  
 876 T is the time ticks and W is the wires.

877 The last cut requires the clusters to be either 30 time ticks or 30 wires. This cut was  
 878 applied to reduce small delta rays associated with a cosmic without removing proton  
 879 clusters associated with a long muon cluster, which saves ideal neutrino events that  
 880 have both a long minimum ionizing muon like cluster and a short highly ionizing  
 881 proton like cluster.

882 **Secondary Cuts**

883 The secondary cuts look to match long, low-angle clusters with short, high-charge  
 884 clusters. Only clusters that have passed previous cuts are used. First clusters with  
 885 length greater than 100 wires are chosen, which is approximately 0.3 m in the z

Cluster set	No Cuts	Primary Cuts	Secondary Cuts
Neutrinos only	570	303	32
Cosmics only ( no flash)	308,016	291,879	602
Cosmics only (w/ flash)	2464	2335	5
Neutrinos/Cosmics	0.23	0.13	6.4

**Table 4.1:** Passing rates for 2D cluster cuts for neutrino on MC set and a cosmic only MC set. First column shows event rates with no cuts applied to both sets. Columns two and three show event rates after primary and secondary cuts are applied. Line three shows the second line scaled with the flash finding factor of 0.008. All events are normalized to per day assuming we are running at 5 Hz.

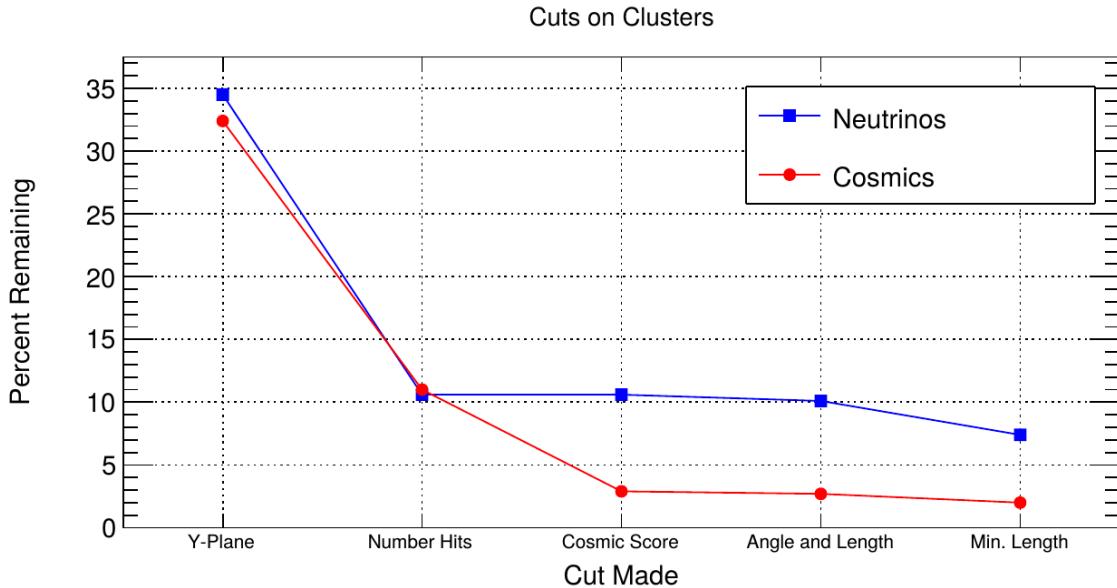
886 direction. Then we search for any cluster that is within approximately 3 cm ( 10 wires  
 887 and 30 time ticks) away from the low-z end of the long cluster. This cluster must also  
 888 be shorter than the first. In our reconstruction, the start and end point of a cluster can  
 889 be swapped so both ends of the short cluster are compared to the long cluster.

890 Now that there is a vertex match, cuts based on charge and projected opening angle  
 891 are implemented. We require the short cluster to have a higher start charge than the  
 892 long cluster or the long cluster be longer than 500 wires. Start charge is defined as  
 893 the charge on the first wire in ADC counts. The projected opening angle must also  
 894 be between 11 and 90 degrees. This last cut is intended to remove clusters that are  
 895 entirely overlapping or are part of the same long track. The resulting neutrino/cosmic  
 896 event rate per day is shown in table 4.1. Figures 4.3 and 4.4 shows the percentages of  
 897 clusters that pass each primary and secondary cuts.

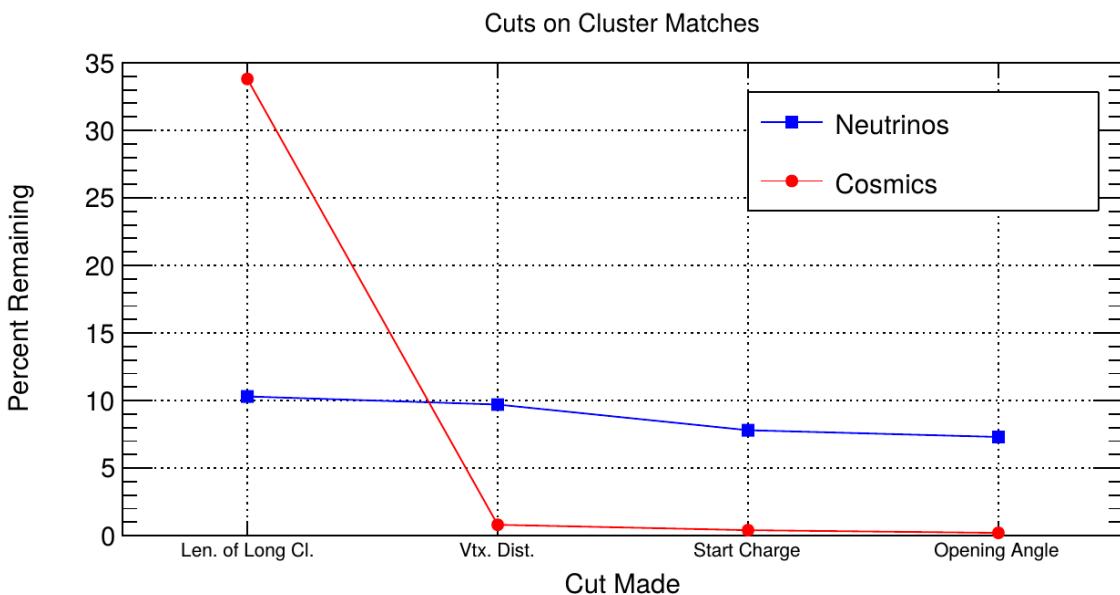
### 898 4.2.3 3D Tracks and vertices Selection

899 The neutrino selection for the 3D channel was based on a reconstructed vertex and  
 900 two tracks. All vertices and tracks were looped over that had a cosmic tag score < 0.4  
 901 and the distances below were calculated:

- 902 •  $d$ : distance between the start points of the two tracks.
- 903 •  $d_1$ : distance between vertex and start of track 1.
- 904 •  $d_2$ : distance between vertex and start of track 2.



**Figure 4.3:** Percent of good clusters remaining for neutrinos and cosmics after the primary cuts were applied. This is relative to total number of initial clusters.



**Figure 4.4:** Percent matched cluster pairs remaining for neutrinos and cosmics after secondary cuts applied. This is relative to the number of events that contain clusters which pass the primary cuts.

905 The maximum distance of all three is then selected as the important characteristic per  
906 trio. The best trio is the one that has the smallest maximum distance. The  $\min(\max_d)$   
907 for all trios in an event were plotted for BNB neutrino events and for cosmics to  
908 find the best cut value for each tracking algorithm. The distribution of  $\min(\max_{d,i})$   
909 is smaller for neutrinos than for cosmics. The cut values for different tracking and  
910 clustering algorithms are shown below. These cut values were chosen to minimize the  
911 cosmic background to 20%.

- 912 • trackkalmanhit with cccluster  $\min(\max_{d,i}) < 3$  cm.  
913 • trackkalmanhit with pandoraNu  $\min(\max_{d,i}) < 4.5$  cm.  
914 • pandoraNu with cccluster  $\min(\max_{d,i}) < 5$  cm.

#### 915 4.2.4 TPC Updates

916 After doing a visual hand-scanning of the first beam data processed with the filters  
917 detailed above, the events passing had a larger contamination of background than  
918 expected. This was mainly in part due to the reconstruction performing better on  
919 simulation than on data. Due to this, additional cuts on both streams needed to be  
920 implemented in order to increase signal/background ratio. These cuts were added on  
921 top of the filters described above and further reduce the event count.

#### 922 2D Filter Updates

923 The main background observed in the 2D filter were Michel events, where the muon  
924 and electron formed two connected clusters. These events were rejected by comparing  
925 the start and end charge deposition of the long cluster (i.e muon particle). The start  
926 charge deposition must be less than the end charge deposition. This cut is implemented  
927 because muons have a higher ionizaiton loss at the end.

#### 928 3D Filter Updates

929 It was seen that cosmic tracks can often originate or end at the same point, therefore  
930 faking a signal. Cosmic tracks, however, are mostly vertical. By requiring the angle  
931 of the longer track have a cosine greater than 0.85 with respect to the z-axis as well

<sup>932</sup> as requiring the longer track to have a length greater than 10 cm, we can reduce this  
<sup>933</sup> background.

<sup>934</sup> **4.3 Conclusion**

<sup>935</sup> After proccesing these filters in parallel, it was shown that the 3D filter had a higher  
<sup>936</sup> purity than the 2D filter because of the higher cosmic rejection being used due to 3D  
<sup>937</sup> reconstruction. The 2D filter is blind to track entering/exiting from the top or bottom  
<sup>938</sup> of the TPC. Although the 3D filter had a higher purity, the 2D filter was still able to  
<sup>939</sup> find identifiable events in data that were used as public event displays. A sample of  
<sup>940</sup> event displays are shown in figures ?? and ??.

<sub>941</sub> **Chapter 5**

<sub>942</sub> **CC-Inclusive Cross Section Selection  
943 Filter**

<sub>944</sub> The CC-Inclusive cross-section selection I and selection I modified filters used in this  
<sub>945</sub> analysis will be described in the following sections below. These filters are an expan-  
<sub>946</sub> sion of the Neutrino ID filter. The work done in this thesis was to further improve these  
<sub>947</sub> selections by increasing both efficiency and purity as well as increasing acceptance  
<sub>948</sub> without further affecting the kinematic distributions of the selected neutrino events.

<sub>949</sub> MicroBooNE requires fully automated event reconstruction and selection algo-  
<sub>950</sub> rithms for use in the many physics measurements being worked on to date due to  
<sub>951</sub> the large data rate MicroBooNE receives. Being able to automatically pluck out the  
<sub>952</sub> neutrino interaction among a sea of cosmics proved to be challenging but was accom-  
<sub>953</sub> plished. MicroBooNE has developed two complementary and preliminary selection  
<sub>954</sub> algorithms to select charged-current  $\nu_\mu - Ar$  interactions. Both are fully automated  
<sub>955</sub> and cut based. The results of this thesis will focus on selection I and selection I modi-  
<sub>956</sub> fied and will focus on further improving these algorithms using Convolutional Neural  
<sub>957</sub> Network (CNN) implementations. These selections identify the muon from a neutrino  
<sub>958</sub> interaction without biasing towards track multiplicity. To combat cosmic and neutral  
<sub>959</sub> current background, the analysis is strongly biased towards forward-going long tracks  
<sub>960</sub> which are contained. This limits phase space and reduces acceptance.

## 961 5.1 Data and MC Processing Chain

962 The data used for this analysis were based on hardware and software triggers. Events  
 963 used came from the *BNB\_INCLUSIVE* and *EXT\_BNB\_INCLUSIVE* streams and were  
 964 used for signal and background. The *BNB\_INCLUSIVE* stream is chosen by requiring  
 965 that the hardware trigger bit is fired and that the event passed an optical software  
 966 trigger within a BNB spill timing window. The *EXT\_BNB\_INCLUSIVE* stream requires  
 967 the EXT hardware trigger to fire as well as pass the same optical software trigger  
 968 within a BNB spill size timing window similar to the *BNB\_INCLUSIVE*.

969 The two MC samples used in this analysis and for determining selection efficiencies  
 970 and purities were GENIE BNB neutrino interactions with CORSIKA cosmic ray overlay  
 971 within the readout window and inTime CORSIKA cosmic rays. The MC samples  
 972 generated used *uboonecode v04\_36\_00* and are based on the following packages:

- 973     • larsoft v04\_36\_00
- 974     • GEANT v04\_09\_06\_p04d
- 975     • GENIE v02\_08\_06d
- 976     • GENIE xsec v02\_08\_06a
- 977     • pandora v02\_03\_0a
- 978     • CORSIKA v07\_4003

979 Both data and MC samples were processed using the same reconstruction release,  
 980 *uboonecode v05\_08\_00* and the fcl files used for reconstruction are listed below:

- 981     • MC fcl files
  - 982         – reco\_uboone\_mcc7\_driver\_stage1.fcl
  - 983         – reco\_uboone\_mcc7\_driver\_stage2.fcl
- 984     • Data fcl files
  - 985         • reco\_uboone\_data\_Feb2016\_driver\_stage1.fcl
  - 986         • reco\_uboone\_data\_Feb2016\_driver\_stage2.fcl

987 On top of the hardware and software triggers, the data also had to pass more  
 988 criteria to be identified as part of the good run list. The criteria is detailed below.

- **Detector conditions:** the detector has to be in a good operating condition. The detector conditions are read from the slow monitoring database and are required to be within the alarm thresholds. The variables of interest for events passing the good run list criteria include DAQ, PMT, HV, Drift HV, wire bias, electron lifetime and detector power. These conditions need to be met on a run-by-run basis in order to pass the selection.
- **Data quality:** normal and stable behavior for basic reconstruction quantities. These reconstruction variables include average number of tracks, hits, and flashes in each event, the average length of tracks, the average amplitude and area of hits, the average PE and the average spread of each one of these quantities.
- **Beam Conditions:** the BNB must be on and stable and the POT per spill needs to be above the intensity threshold. Beam quality conditions include checking the fraction of proton beam interacting within the target, the horn current, and the intensity of protons per spill. The final sample is  $5 * 10^{19}$  and a per-spill intensity of  $4 * 10^{12}$
- **Run processed:** the full run must be processed completely without missing subruns or crashes in the data processing.

The selection begins with a cut that requires an optical flash greater than 50 photo electrons (PE) in the  $1.6 \mu\text{s}$  beam window. Next, two or more 3D reconstructed tracks must be within 5 cm from a 3D reconstructed vertex. The most forward going track vertex-track association is then selected for further cuts. The vertex from the chosen association must be in the fiducial volume, and the longest track from this association must be matched to a flash 80 cm in z. Lastly the longest track must be contained and longer than 75 cm.

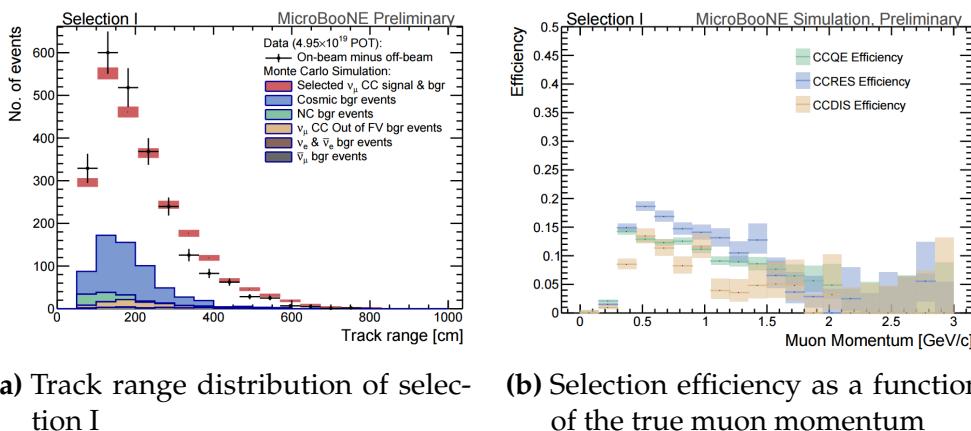
## 5.2 Normalization of data and MC

The off-beam sample is used to measure beam unrelated backgrounds. For normalization, one needs the total number of BNB spills ( $N_{BNB}$ ) and the total number of external triggers. The BNB spills used need to pass the beam quality cuts. The normalization factor is then  $N_{BNB} / N_{EXT}$  which is 1.23.

To normalize generated BNB MC events to POT, we used the following:

- 1019     •  $5 * 10^{19} POT = 41524.3$  generated events

1020 where this scaling factor only applies to mcc7 generated events. The inTime cosmic  
 1021 sample is normalized with respect to the open cosmic sample so an understanding  
 1022 of both is necessary. The POT per beam spill for mcc7 BNB samples is  $5 * 10^{12}$ . To  
 1023 calculate how many spills are necessary to produce a specific POT one would multiply  
 1024 the total POT by the average 1/POT per spill. For a total POT of  $5 * 10^{19}$  the amount of  
 1025 spills necessary is  $\frac{5*10^{19}}{5*10^{12}} = 1 * 10^7$ . This is only one in 241 events therefore each cosmic  
 1026 event needs to be scaled up by a factor of 240.8 when comparing to BNB MC. For  
 1027 inTime cosmics however, two filters are applied to reduce computing and processing  
 1028 time and only leave cosmics that will interact within the detector. The passing rate  
 1029 after these two filters is 0.02125, therefore the total inTime cosmic scaling factor to  
 1030 compare inTime cosmics to BNB is  $0.02125 * 240.8 = 5.12$ .



**Figure 5.1:** **5.1a** Track range distribution for selection I. The track range is defined as the 3D distance between the start and end of the muon candidate track. No data is shown below 75 cm due to the track length cut described previously. **5.1b** Efficiency of the selected events by process quasi-elastic (QE), resonant (RES), and deep-inelastic (DIS). Statistical uncertainty is shown in the bands and the distributions are a function of true muon momentum. The rise of the efficiency between 0 GeV and 0.5 GeV is due to the minimum track length cut and the decreasing efficiency for higher momentum tracks is caused by the containment requirement.

## 1031 5.3 Optical Software Trigger and Reconstruction

### 1032 5.3.1 Software Trigger

1033 Most of the BNB spills from the accelerator do not have a neutrino interaction in  
1034 MicroBooNE. To save computation resources and reduce data-rates, we require a  
1035 burst of light in the light collection system in coincidence with the 1.6  $\mu$ s beam spill.  
1036 Requiring light activity in coincidence with the beam spill eliminates the vast majority  
1037 of triggers with no neutrino interaction in the detector, however, it doesn't guarantee  
1038 the activity in the detector is a neutrino interaction since a cosmic ray can interact in  
1039 coincidence with the beam spill as well.

1040 To implement this, a software trigger was used on the PMT waveforms to decide  
1041 whether or not to keep that event. The software trigger is implemented after the event  
1042 builder combines data from the PMTs and triggers into a single event. The software  
1043 trigger uses the digitized output of the 32 PMT channels in the light collection system.  
1044 Only the waveform region in coincidence with the beam spill is used to search for  
1045 possible triggers. For each PMT, a waveform is found by taking the difference of ADC  
1046 values is calculated between  $t$  and  $t + s$ . This waveform is then scanned for ADC  
1047 values above a threshold  $X_0$ . Once an ADC is above this threshold, a discriminator  
1048 window is opened for a fixed number of time ticks ( $W_0$ ). If the ADC count within this  
1049 window  $W_0$  is greater than a second larger threshold  $X_3$ , a final window of width  $W_3$   
1050 is opened. The max ADC value within this final window is set as the peak amplitude  
1051 for the PMT and then summed across all 32 PMTs and set to the variable PHMAX. The  
1052 software trigger places a final cut on the PHMAX variable to decide whether or not  
1053 to keep the event. The thresholds were found by the Trigger task force using Monte  
1054 Carlo Studies and are as follows:

- 1055 •  $X_0 = 5$  ADC
- 1056 •  $X_3 = 10$  ADC
- 1057 •  $W_0 = 6$  Ticks
- 1058 •  $W_3 = 6$  Ticks
- 1059 • PHMAX cut = 130 ADC

**1060 5.3.2 Flash Reconstruction**

1061 MicroBooNE collects light from each of the 32 PMTs either in a continuous readout  
 1062 window of 23.4  $\mu$ s activated by a beam gate signal on the trigger board, or in discrimi-  
 1063 nated pulses of 1  $\mu$ s duration activated if the ADC count for any PMT goes above 80  
 1064 ADC count. These two formats are saved as output waveforms and put onto an event.  
 1065 Additionally, each PMT can provide two output streams, high-gain ( 20 ADC/PE) and  
 1066 low-gain ( 2 ADC/PE) channels. The first step in the reconstruction is to merge both  
 1067 these channels into a “saturation corrected waveform” which uses information from  
 1068 the low-gain waveform to correct for saturating high-gain pulses.

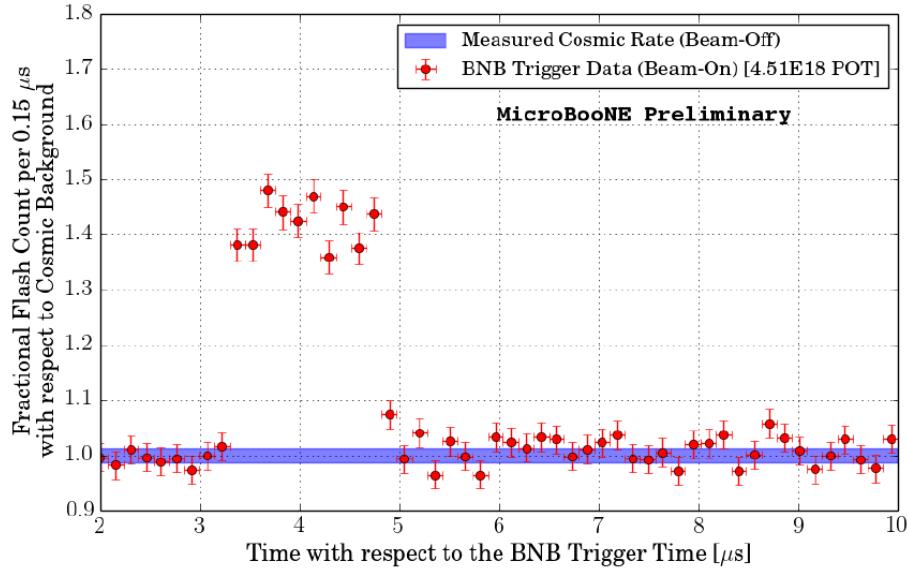
1069 The saturation corrected waveform in the continuous readout window is used to  
 1070 reconstruct optical hits. Each PMT’s waveform is scanned for hits then a threshold  
 1071 based hit reconstruction algorithm is applied which requires pulses of a minimum  
 1072 area in order to be reconstructed. Each reconstructed hit is associated to a PMT, a time  
 1073 in  $\mu$ s, and a PE count.

1074 Once hits are reconstructed for all 32 PMTs, all PMT information is then combined  
 1075 into optical flashes which represent optical information seen by the PMTs from interac-  
 1076 tions in the detector. Each flash has information on total light seen per interaction, the  
 1077 distribution of the light across all 32 PMTs, the flash time with respect to the trigger  
 1078 time of the flash, and lastly, the spacial information of the flash in Y-Z plane of the  
 1079 detector. These flashes are reconstructed by requiring that there is a 1  $\mu$ s coincidence  
 1080 between the reconstructed hits in all 32 PMTs. The total PE is summed up among  
 1081 all coincident hits across the PMTs and if the total PE is greater than 2 PE, a flash is  
 1082 reconstructed. There are also safe guards in place to take care of late scintillation light.

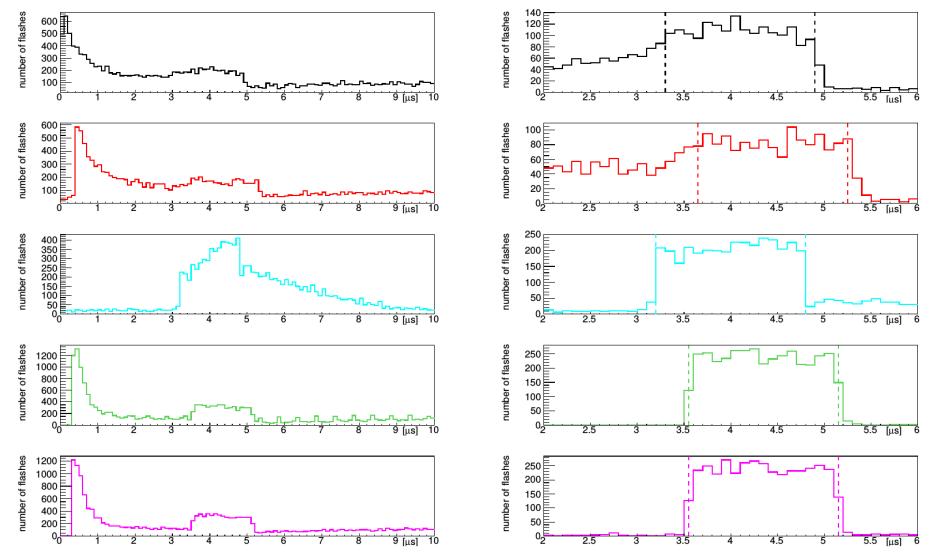
1083 Figure 5.2 shows the time distribution of reconstructed optical flashes using the  
 1084 BNB continuous stream. You can see a clear excess in coincidence with the expected  
 1085 arrival time of neutrinos. The same flash reconstruction that was used in the cc-  
 1086 inclusive filter detailed here was used to create this plot in data.

**1087 5.3.3 Beam Window**

1088 Figure ?? shows the distribution of flashes for on-beam, off-beam and various MC  
 1089 samples. The software trigger has been applied to these samples. The pile-up seen just  
 1090 after 0  $\mu$ s is a feature of the flash finding algorithm and consists of low PE flashes and



**Figure 5.2:** Time distribution of reconstructed optical flashes with a PE value of 50 or more for a sample of BNB unbiased triggered events.



**Figure 5.3:** Flash time distribution for all flashes (left plot) and flashes  $> 20\text{PE}$  (right plot). The different curves are as follows: on-beam data (black), off-beam data (red), CORSIKA inTime MC (light blue), BNB only MC (green), and BNB+Cosmic MC (purple). The dashed vertical lines mark the time window that was chosen for each sample

1091 is removed in the second column of distributions with a low 20 PE threshold cut. The  
1092 plots show that the time window for the distributions are shifted a small amount from  
1093 each-other. This is caused by different hardware configurations per sample. Using  
1094 these distributions, the windows chosen per sample are as follows:

- 1095 • On-Beam: 3.3 to 4.9  $\mu$ s
  - 1096 • Off-Beam: 3.65 to 5.25  $\mu$ s
  - 1097 • CORSIKA inTime: 3.2 to 4.8  $\mu$ s
  - 1098 • BNB only: 3.55 to 5.15  $\mu$ s
  - 1099 • BNB+Cosmic: 3.55 to 5.15  $\mu$ s
- 1100 Each window has a width of 1.6  $\mu$ s.

## 1101 5.4 TPC Reconstruction

## 1102 5.5 Event Selection

<sub>1103</sub> **Chapter 6**

<sub>1104</sub> **Background on Convolutional Neural  
1105 Networks**

<sub>1106</sub> Convolutional neural networks (CNNs) have been one of the most influential inno-  
<sub>1107</sub> vations in the field of computer vision. Neural networks became popular in 2012  
<sub>1108</sub> when Alex Krizhevsky used them to win that year's ImageNet competition [?] by  
<sub>1109</sub> dropping the error from 26% to 15%. Since then, many companies are using deep  
<sub>1110</sub> learning including Facebook's tagging algorithms, Google for their photo search and  
<sub>1111</sub> Amazon for product recommendations. For the purpose of this thesis CNNs were  
<sub>1112</sub> used for image classification, specifically, images of varying particles created using  
<sub>1113</sub> LArTPC data.

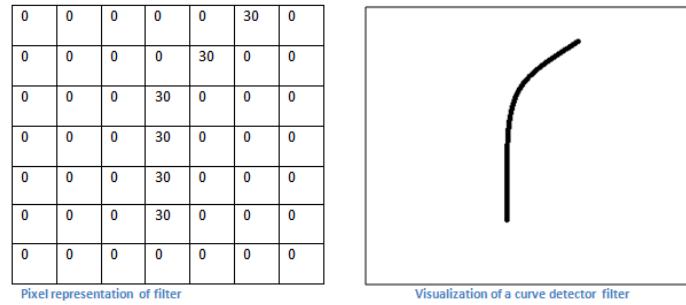
<sub>1114</sub> **6.1 Image Classification**

<sub>1115</sub> Image classification is the process of inputting an image into the CNN and receiving a  
<sub>1116</sub> probability of classes that best describes what is happening in the image. As humans,  
<sub>1117</sub> image classification is something that is learned at a very young age and is easy to  
<sub>1118</sub> do without much effort. This is also apparent when hand-scanning LArTPC images.  
<sub>1119</sub> After learning what a neutrino event looks like in MicroBooNE, it is relatively easy  
<sub>1120</sub> to recognize simple neutrino events from cosmic ray background as well as highly  
<sub>1121</sub> ionizing particles like protons from minimum ionizing particles like muons. The very  
<sub>1122</sub> detailed images LArTPC detectors output are prime candidates for input images into  
<sub>1123</sub> a CNN. CNNs mimic a human's ability to classify objects by creating an architecture  
<sub>1124</sub> that can learn differences between all the images it's given as well as figure out the

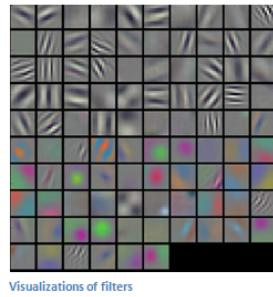
unique features that make up each object. CNNs are modeled after the visual cortex. Hubel and Wiesel [?] found that there are small regions of neuronal cells in the brain that respond to specific regions of the visual field. They saw that some neurons fired when exposed to vertical edges while others fired when shown horizontal or diagonal edges. They also saw that these neurons were organized in columns. The idea of specific neurons inside of the brain firing to specific characteristics is the basis behind CNNs.

## 6.2 CNN Structure

When used for image recognition, convolutional neural networks consist of multiple layers that extract different information on small portions of the input image. How many layers is tunable to increase the accuracy. The output of these collections are then tiled so that they overlap to gain a better representation of the original image and allow for translation. The first of these layers is always a convolution layer. To the CNN, an image is an array of pixel values. For a RGB color image with width and height equal to 32x32 the corresponding array is 32x32x3. Filters, also known as neurons, of any size set by the user is then convolved with the receptive field of the image. If the filter is 5x5, the receptive field will be a patch of 5x5 on the input image. The filter is also an array of numbers called weights. The convolution of the filter and image are matrix multiplications of the weights and the pixel values. By stepping the receptive field by 1 unit, for an input image of 32x32x3 and a filter of 5x5x3 you'd get an output array of 28x28x1. This output array is called an activation map or feature map. The use of more filters preserves the spatial dimensions better. The filters can be described as feature identifiers. Examples of features in an image consist of edges, curves, and changes in colors. The first filters in a CNN will primarily be straight line and curve feature identifiers. An example of a curve filter is shown in figure 6.2. When a curve in the same concavity is found in the input image, the corresponding pixel in the output feature map will be activated. Going back to our example of a 32x32 input image and a 5x5 filter, if there were to be a curve in the top left corner of the input image, our output feature map would have a high pixel value in the top left. Therefore, feature maps tell us where a specific feature is located in the original image. Figure ?? shows a visualization of filters found in the first layers of many CNN architectures. These filters in the first layer convolve around the image and activate when the specific feature it is looking for is in the receptive field.



**Figure 6.1:** Pixel representation and visualization of a curve detector filter. As you can see, in the pixel representation, the weights of this filter are greater along a curve we are trying to find in the input image

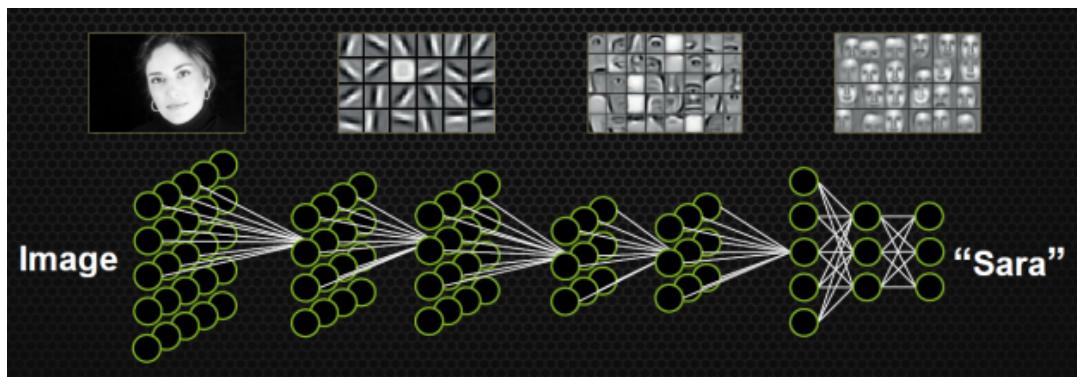


**Figure 6.2:** Visualization of filters found in first layer of a CNN.

In figure 6.3 you can see how an edge detection filter is used to save only necessary information for recognizing different types of clothes. You can also see by having multiple filters you can get more detail or less detail from an image which can then simplify or complicate the object recognition task. Being able to distinguish between a shirt or a leg garment is as much information you want, having a filter that extracts outline edge or shape information would be all that you need. But if instead you wanted to distinguish between a formal cocktail dress or a summer dress, more information would need to be saved equating to many more filters for one image. Rather than trying to come up with how many filters and what features are important for detection, CNNs do this automatically. CNNs take input parameters, called hyperparameters, for example number of layers, number of filters per layers, number of weights per filter, and uses these to create the output feature maps. The layers build upon each-other, for example if we were creating a CNN for facial recognition the convolutional layers will start learning feature combinations off of the previous layers. The low level features like edges, gradients, and corners of the first layers become high level features like eyes, noses, and hairs. This process is visualized in figure 6.4



**Figure 6.3:** Applying a feature mask over a set of fashion items to extract necessary information for auto-encoding. Unnecessary information for example color or brand emblems are not saved. This feature map is an edge detection mask that leaves only shape information which helps to distinguish between different types of clothes.



**Figure 6.4:** Pictorial Representation of Convolutional Neural Networks as well as a visual representation on CNN's complexity of layer feature extraction

1174 There are other layers in a CNN architecture that will not be covered in the scope  
 1175 of this thesis but in a general sense, these layers are interspersed between convolution  
 1176 layers to preserve dimensionality and control overfitting of the network. The last layer  
 1177 is called a fully connected layer and its job is to output an  $N$  dimensional vector where  
 1178  $N$  is the number of classes the network has been trained on. Each number in this vector  
 1179 represents the probability that the input image is a certain class. Fully connected layers  
 1180 use the feature maps of the high level features to compute the products between the  
 1181 weights of the previous layer to get the probabilities of each class. These weights are  
 1182 then adjusted through the training process using backpropagation.

---

### 1183 6.2.1 Backpropagation

1184 A CNN at it's onset has weights that are randomized. The filters themselves don't  
 1185 know how to pull out identifying information per class. For a neural network to learn,  
 1186 it must be trained on a training set that is labeled. Backpropagation has four seperate  
 1187 steps: foward pass, loss function, backward pass and updating weights. In the forward  
 1188 pass, a training image is passed through the whole network. All of our weights at this  
 1189 time are randomly initialized so the output for the first image will have no preference  
 1190 to a specific class. A common loss function is mean squared error (MSE):

$$E_{total} = \sum \frac{1}{2} (actualclass - predictedclass)^2 \quad (6.1)$$

1191 If we assume that the MSE is the loss of our CNN, the goal would be that our  
 1192 predicted label (output of CNN) is the same as our training label. To do this, we need  
 1193 to minimize the loss function. To do this, it is necessary to find out which weights most  
 1194 directly affect the loss of the network i.e  $\frac{dL}{dW}$  where L is our loss function and W are  
 1195 the weights of a specific layer. The next step is the backward pass which determines  
 1196 which weights contribute the most to the loss and finds ways to adjust these weights  
 1197 so that the loss decreases. After the derivative is computed, the last step updates the  
 1198 weights in the opposite direction of the gradient.

$$w = w_i - \eta \frac{dL}{dW} \quad (6.2)$$

$$w = \text{Weight} \quad (6.3)$$

$$w_i = \text{Initial Weight} \quad (6.4)$$

$$\eta = \text{Learning Rate} \quad (6.5)$$

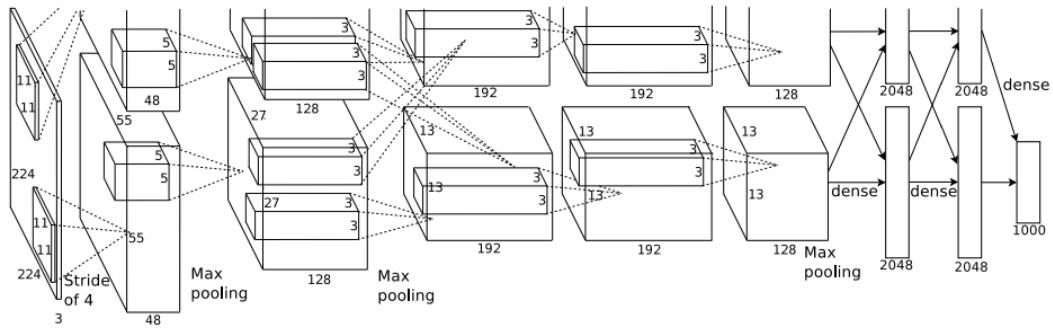
1199 The learning rate is a parameter given to the CNN and it describes the steps the  
 1200 network takes to update the weights. Higher learning rate equals large steps and a  
 1201 lower training time, but a learning rate that is too large can mean the CNN never  
 1202 converges.

1203 Going through backpropagation consists of one training iteration. Once the net-  
 1204 work completes a specific number of iterations, another parameter given, and runs  
 1205 over all training images that are split up into batches, the process is considered com-  
 1206 plete. User input parameters, called hyperparameters, help the network converge to

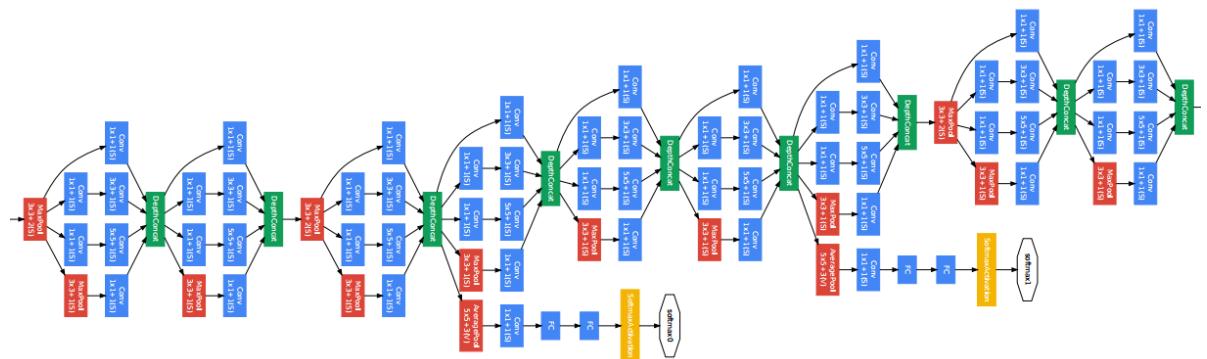
optimal weights for each layer. Batch size, learning rate, and training iteration are just some of the user input hyperparameters that help. Lastly, to check if the network has learned, a different set of labeled images are fed to the CNN iteratively through the training process to see how well it's learning. This process is especially important to make sure the network architecture isn't being affected by overfitting (memorizing training input rather than learning).

## 6.3 Choosing Hyperparameters

Convolutional neural networks are a relatively new tool in computer vision. Choosing hyperparameters for your specific dataset is a non-trivial task. Hyperparameters can range from the amount of layers and filters per layer in an CNN architecture to the stride the receptive field of a filter takes, not to mention training hyperparameters such as learning rate and batch size described above. They're ways to optimize these hyperparameters via hyperparameter optimization using Bayesian Optimization [?] but as you can imagine, optimizing an CNN architecture from scratch can be very computationally intensive. For the purpose of this thesis, two well known CNN architectures were used, AlexNet [?], which won the ImageNet Large-Scale Visual Recognition Challenge (ILSVRC) in 2012 and therefore bringing awareness of CNNs, and GoogleNet [?], which won the ILSVRC in 2014, giving rise to deep networks. Both AlexNet and GoogleNet architectures were used to train on LArTPC images and their low level filter weights. Higher level filter weights were randomly initialized before training so the network can learn high level features of LArTPC image classes. The AlexNet architecture is shown in figure 6.5 and the GoogleNet architecture is shown in figure 6.6



**Figure 6.5:** Pictoral representation of the AlexNet model. The AlexNet model consists of 5 convolution layers and 3 fully connected layers.



**Figure 6.6:** Pictoral representation of the GoogleNet model. The GoogleNet model consists of 22 layers. The model implements 9 Inception modules which performs covolution and pooling in parallel and strays away from the basis that CNN layers need to be stacked up sequentially. The GoogleNet model also doesn't use fully connected layers, instead it uses average pooling which greatly reduces the amount of parameters. GoogleNet has 12x fewer parameters than AlexNet.



1230 **Chapter 7**

1231 **Training Convolutional Neural  
1232 Networks on particles WORKING  
1233 TITLE**

1234 Three Convolutional Neural Networks CNNs were trained throughout this analysis.  
1235 There are differences to each CNN and will be described fully in the next sections but  
1236 the main difference are the amount of particle images used for training and validation.  
1237 CNN1075 used 1,075 muons and 10,75 pions for training and the same amount of each  
1238 particle for validation. CNN10000 used 10,000 muons and 10,000 pions split in half  
1239 for testing and training. Lastly CNN100000 had muons, pions, protons, electrons,  
1240 and gammas in its training and validation set. Each particle had 20,000 images and  
1241 training and validation was split 90% training, 10% validation. This chapter will also  
1242 describe the different hardware frameworks used for training beginning on a CPU  
1243 and ending on a GPU cluster.

---

<sub>1244</sub> **7.1 Hardware Frameworks used for Training**

<sub>1245</sub> **7.1.1 Syracuse CPU Machine setup**

<sub>1246</sub> **7.1.2 Syracuse University GPU Cluster Setup**

<sub>1247</sub> **7.2 Convolutional Neural Network Training**

<sub>1248</sub> **7.2.1 Image Making Scheme**

<sub>1249</sub> **Images used for Traing/Validation of Convolutional Neural Networks**

<sub>1250</sub> **add image making for CNN1075** The  $\mu/\pi$  image dataset used to train and validate  
<sub>1251</sub> the CNN10000 was created using single generated isotropic muons and pions from  
<sub>1252</sub> 0-2 GeV energ range. 10,000 muons and 10,000 pions were used for training and  
<sub>1253</sub> testing split 50%. The images were created based on wire number and time tick in the  
<sub>1254</sub> collection plane. Uboonecode v06\_23\_00 was used instead of v05\_08\_00 which was  
<sub>1255</sub> used previously. The wire signal was the raw ADC value after noise filtering. Each  
<sub>1256</sub> collection plane grayscale image was 3456x1280x1 where 5 time ticks were pooled into  
<sub>1257</sub> 1 bin which is different than the previous dataset and was implemented due to the fact  
<sub>1258</sub> that the time ticks of an event went from 9400 to 6400 with the change of uboonecode  
<sub>1259</sub> version. The grayscale color standard is 8bit therefore the ADC value of wire and time  
<sub>1260</sub> tick was also downsampled due to the 12bit ADC value MicroBooNE has. To do this,  
<sub>1261</sub> the highest ADC pixel in the image was found and then this was divided by the rest  
<sub>1262</sub> placing all pixel values between 0-1. From there, all pixel values are then multiplied  
<sub>1263</sub> by 255. All images were made using a LArSoft module. Once the images were created,  
<sub>1264</sub> using and image manipulation framework called OpenCV images were read into a  
<sub>1265</sub> numpy array and cropped to the region of interest by only keeping rows and columns  
<sub>1266</sub> where all ADC values are higher than 0 and then resized it to 224x224 using OpenCV's  
<sub>1267</sub> resize function. This downsampling of ADC values creates a problem of information  
<sub>1268</sub> loss for example, a proton which is highly ionizing will have the same brightness as a  
<sub>1269</sub> minimum ionizing muon by virtue of how the images are created. Issues that arose  
<sub>1270</sub> in CNN1075 that were fixed in CNN10000 include zero-padding images in X and Y  
<sub>1271</sub> that are smaller than 224X224 to eliminate over-zooming effect and fixing a bug that  
<sub>1272</sub> shifted pixels separated by a dead-wire region.

1273 Images were also made from events that passed the cc-inclusive selection 1 filter  
1274 right before the 75 cm track length cut and were classified using the CNN10000. The  
1275 dataset used to create these images is the same one used in [?], prodgenie\_bnb\_nu\_cosmic\_uboone\_mc  
1276 These images were created using information from the track candidate that passed  
1277 the filter. Only wire number and time ticks associated to the track candidate were  
1278 drawn on the image to mimic a single particle generated image. These images were  
1279 then classified using CNN10000. Two approaches were taken in making these images.  
1280 The first was using the image normalization above where the maximum pixel in each  
1281 image is used as a normalization constant to get all pixels between 0-1 then multiply  
1282 all pixels by 255. As described above, this is the incorrect way to normalize; it should  
1283 be normalized by dataset not by event, which is the second way the images were  
1284 created. The results of CNN10000 performance are shown in section 7.2.

### 1285 7.2.2 Training CNN1075

1286 The work shown in these next sections are based on the previous work done described  
1287 in [?]. That CNN (now referred to as CNN1075) was trained using single generated  
1288 isotropic muons and pions from 0-2 GeV energy range. 1,075 muons and pions were  
1289 used to train the network and 1,075  $\mu/\pi$  were used as a validation set. The accuracy is  
1290 how well CNN1075 is doing by epoch and was 74.5%. The loss is gradient descent  
1291 or minimization of the error of the weights and biases used in each neuron of each  
1292 layer of CNN1075 and was 58% with a trend sloping downwards on the loss curve  
1293 as well as a trend sloping upward in the accuracy curve. Due to the depth of the  
1294 neural network framework, it was necessary to train with a larger dataset and for  
1295 more epochs, however, the downward slope of the loss curve is an indication that once  
1296 trained for longer with a higher training sample, neural networks can be used for  $\mu/\pi$   
1297 separation. Updates in the image making and downsampling algorithm were made to  
1298 fix issues that arose in CNN1075.

### 1299 7.2.3 Training CNN10000

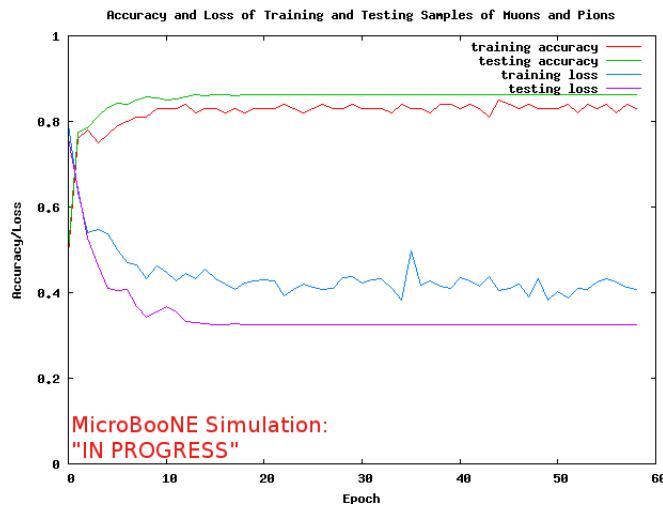
1300 The hyperparameters used for CNN10000 are shown. The batch size for the training  
1301 and testing as well as the test iter were chosen to encompass the whole training/testing  
1302 image set when doing accuracy/loss calculations. To do this, multiplying the test

1303 iter by the test batch size give you the amount of images used when calculating  
1304 accuracy/loss curves. For reference, the accuracy and loss are defined as well.

```
1305     • train_batch_size: 100
1306     • test_batch_size: 100
1307     • test_iter: 100
1308     • test_interval: 100
1309     • base_lr: 0.001
1310     • lr_policy: "step"
1311     • gamma: 0.1
1312     • stepsize: 1000
1313     • display: 100
1314     • max_iter: 10000
1315     • momentum: 0.99
1316     • weight_decay: 0.0005
1317     • snapshot: 100
1318     • Accuracy: How often the CNN predicts the truth over total number of images
1319     • Loss: Error between truth and prediction. Minimize loss by gradient descent to
1320       update weights and biases of CNN
```

1321 The same architecure that was used to train CNN1075 was employed on CNN10000,
1322 Imagenet. Caffe [?] was the software package used for both CNNs. The differences
1323 include batch size and test\_iter and momentum to account for the larger dataset. Both
1324 CNNs were trained on a CPU machine, Syracuse01. Further training will be done
1325 on a GPU cluster stationed at Syracuse University. Figure 7.1 shows the loss and
1326 accuracy of CNN10000. There is around a 10% increase in accuracy from CNN1075 to
1327 CNN10000, 85%, and around a 20% decrease in loss, 36%.

1328 Figure 7.2 show a breakdown of  $\mu/\pi$  separation for CNN10000. It also shows
1329 the network is not being overtrained due to the Accuracy of both the training and
1330 testing datasets being within .01% of eachother. The CNN is doing a very good job of

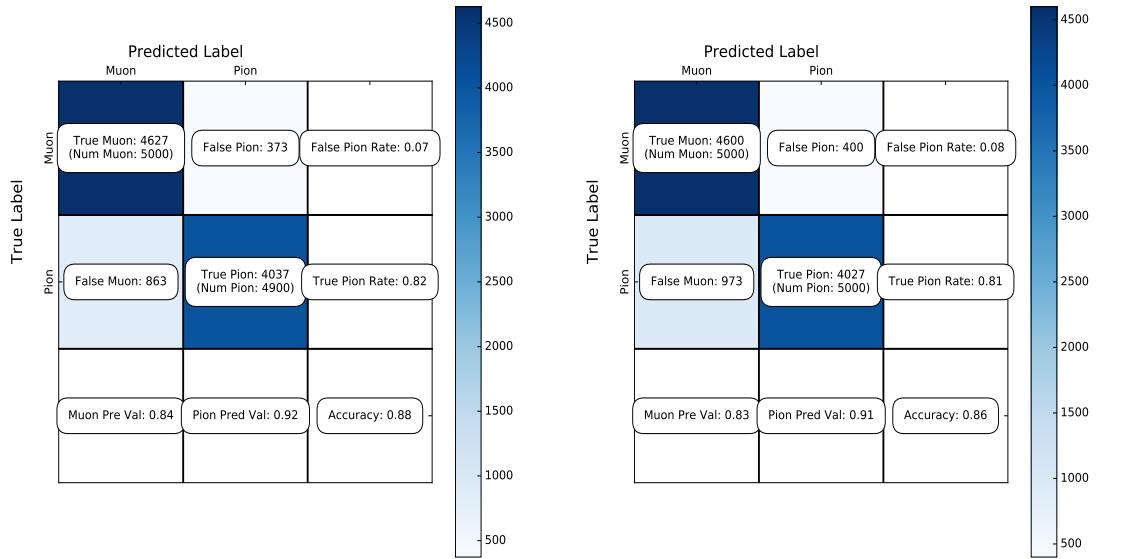


**Figure 7.1:** Accuracy vs. Loss of ImageNet 2-output  $\mu/\pi$  sample consisting of 10000 images each.

<sub>1331</sub> classifying true muons as muons, and our loss increase from CNN1075 is due to the  
<sub>1332</sub> increase in accurately classifying pions as pions.

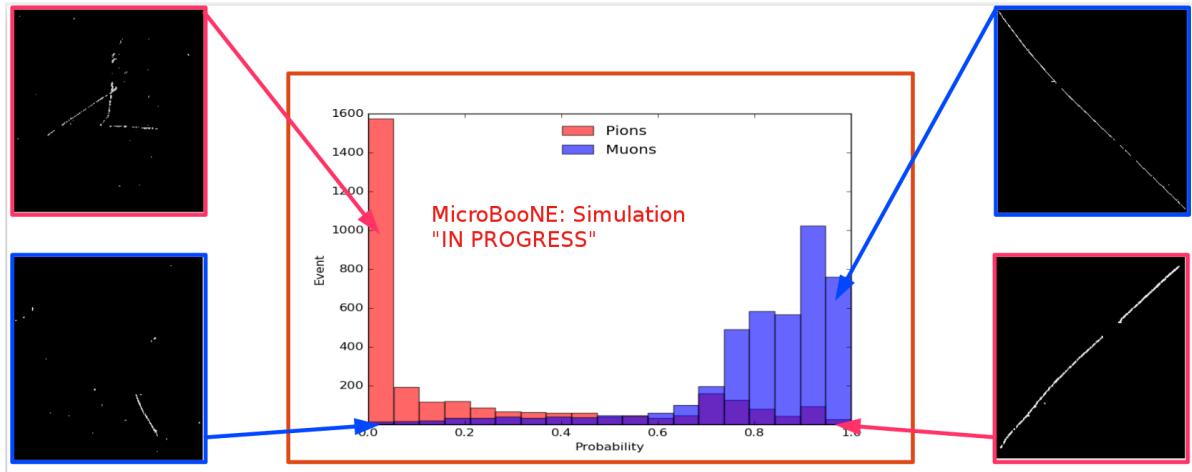
#### <sub>1333</sub> 7.2.4 Training CNN100000

<sub>1334</sub> Results of training using 100,000 images, 20,000 images per  $\mu/\pi/p/\gamma/e$ .



**(a)** Confusion Matrix showing Accuracy of CNN using training data

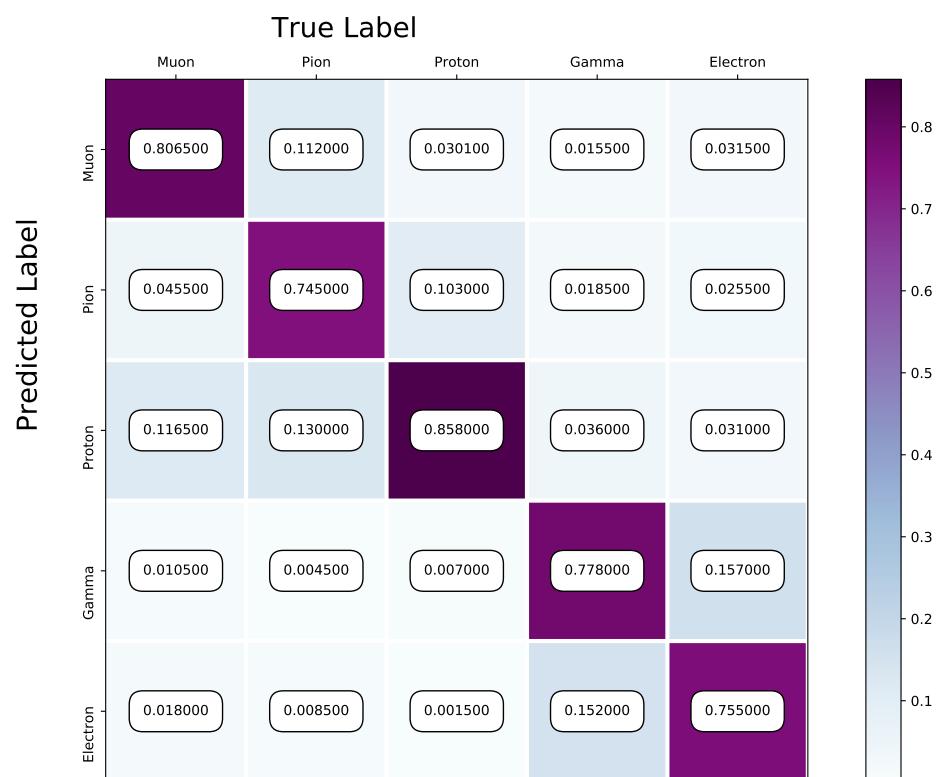
**(b)** Confusion Matrix showing Accuracy of CNN using testing data



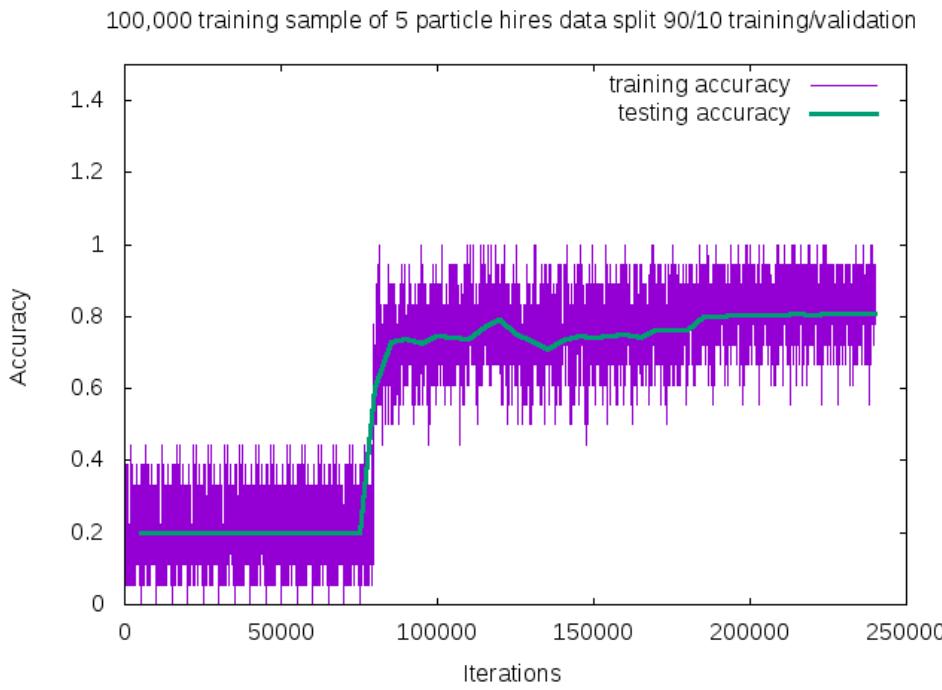
**(c)** Probability plot of muons and pions from testing set. Images surrounding histogram are a random event from lowest bin and highest bin for each particle.

**Figure 7.2:** Description of confusion matrix variables: False pion rate =  $false\pi / total\pi$  True pion rate =  $true\pi / total\pi$  Accuracy =  $(true\pi rate + true\mu rate) / 2$  Pion prediction value =  $true\pi / (true\pi + false\pi)$  Muon prediction value =  $true\mu / (true\mu + false\mu)$

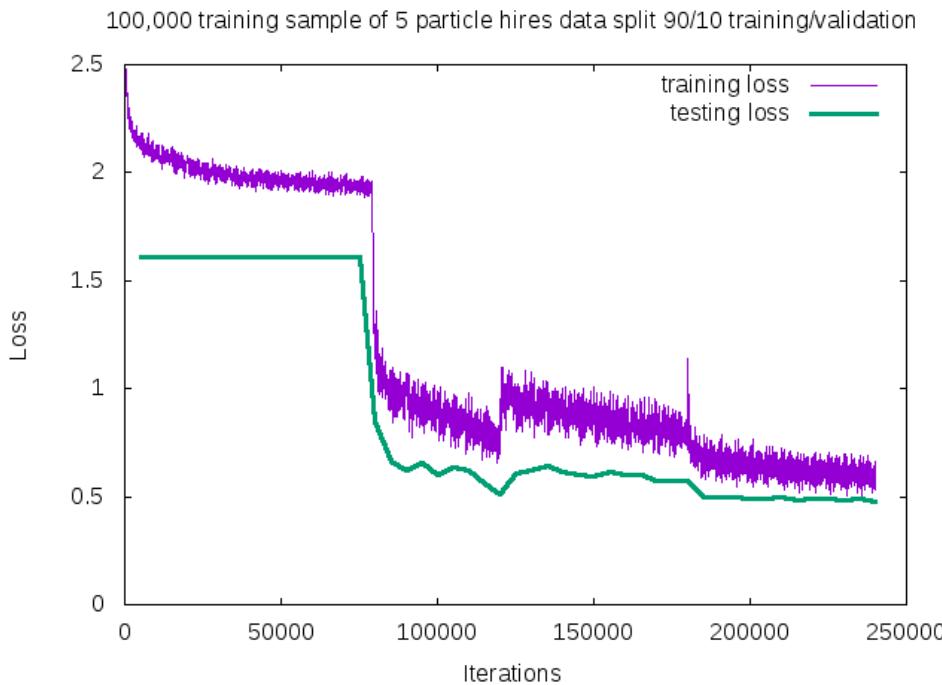
**7.2c** The probability plot includes muons and pions that are classified as primary particles.



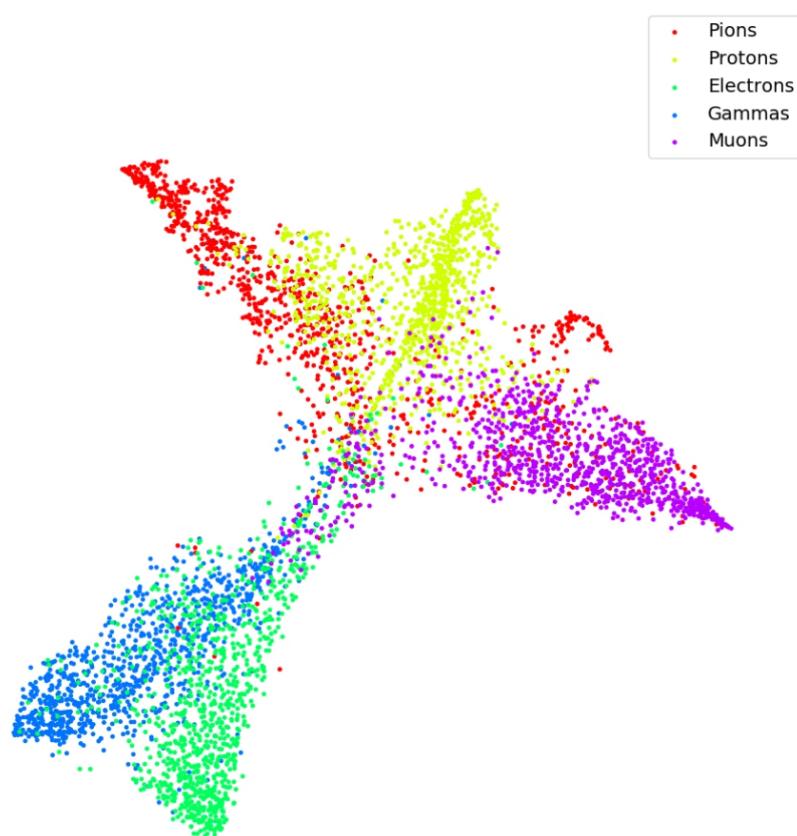
**Figure 7.3:** Confusion Matrix of all five particles



**Figure 7.4:** Training and testing accuracy of CNN trained on 100,000 images of  $\mu/\pi/p/\gamma/e$  with 20,000 images of each particle. Each image was a size of 576x576 and the images per particle were split 90% use for training and 10% used for testing the network



**Figure 7.5:** Training and testing loss of CNN trained on 100,000 images of  $\mu/\pi/p/\gamma/e$



**Figure 7.6:** t-SNE of CNN

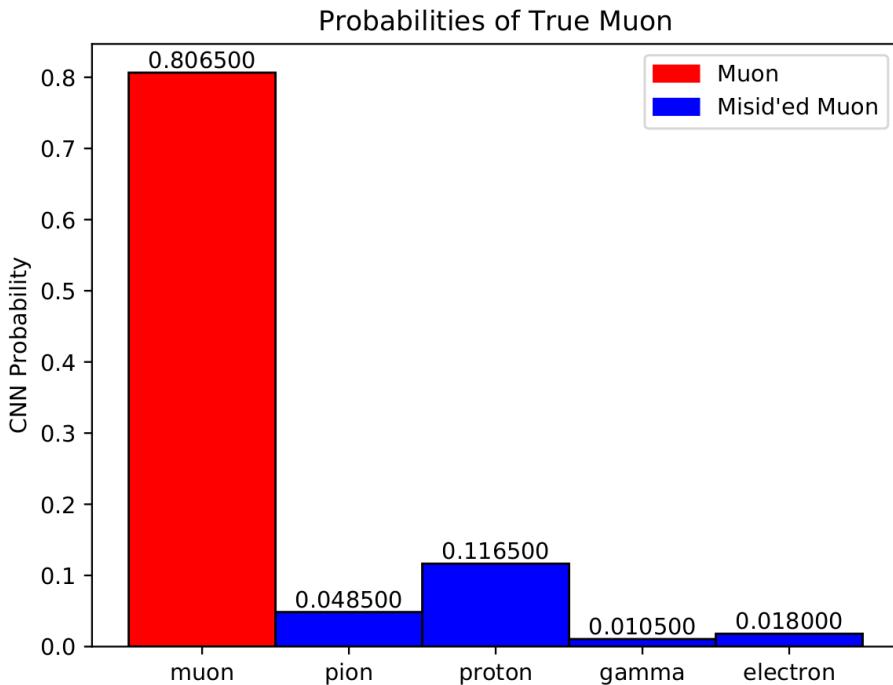


Figure 7.7: Muon Prob

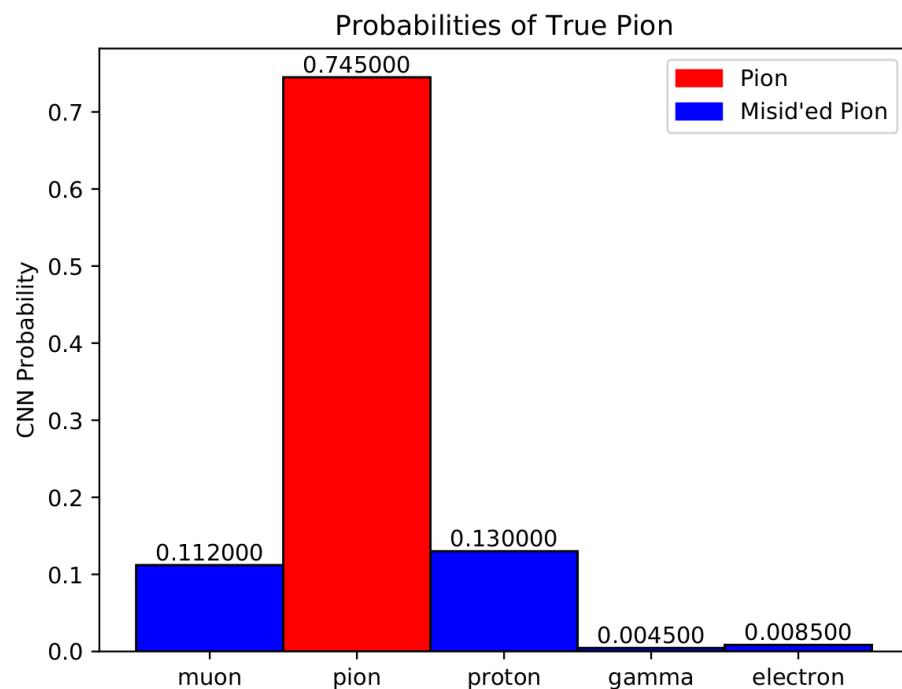
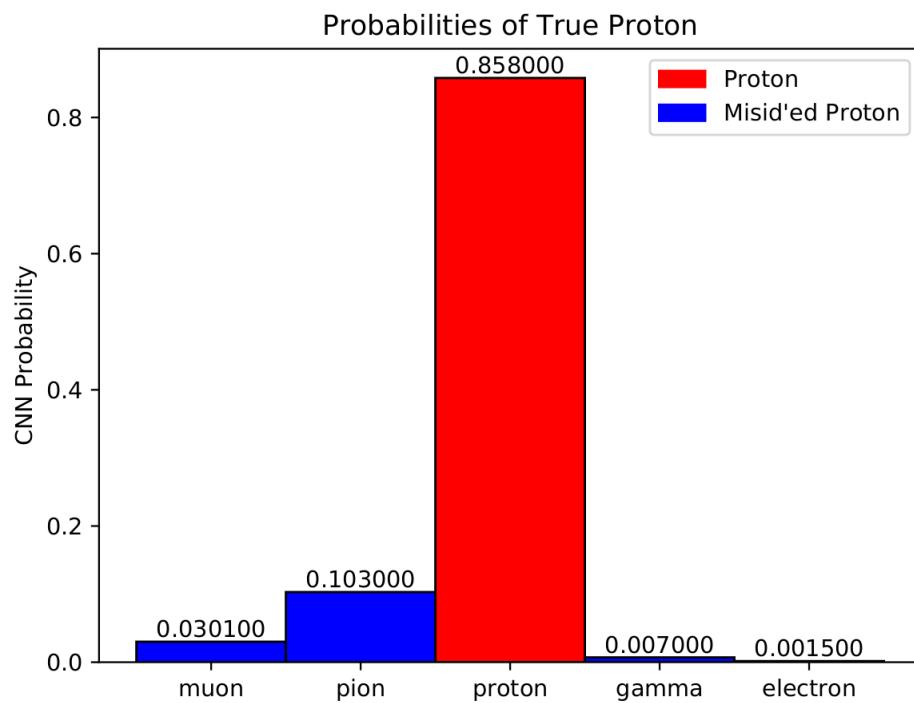
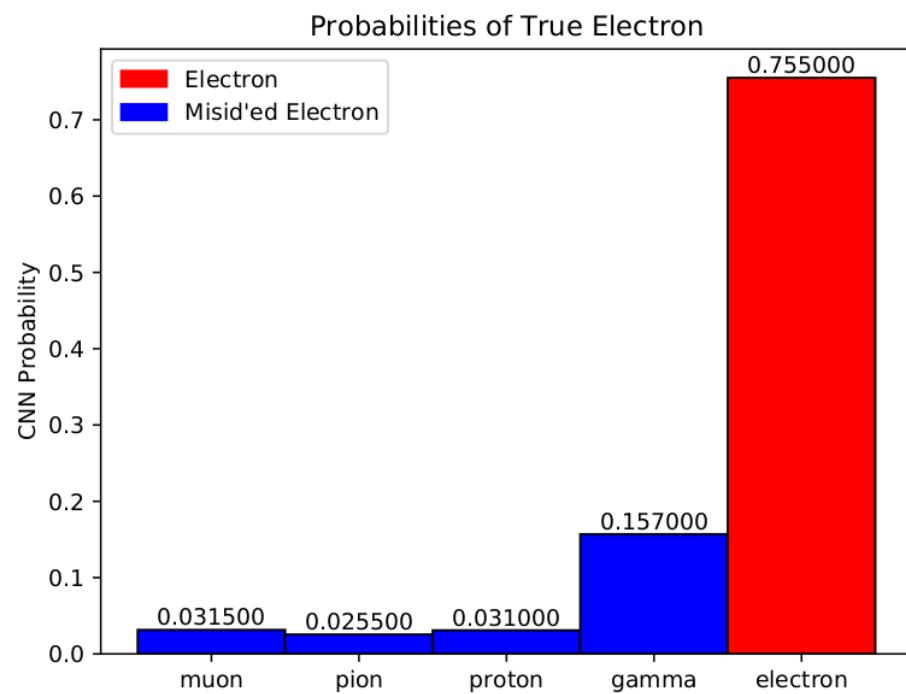
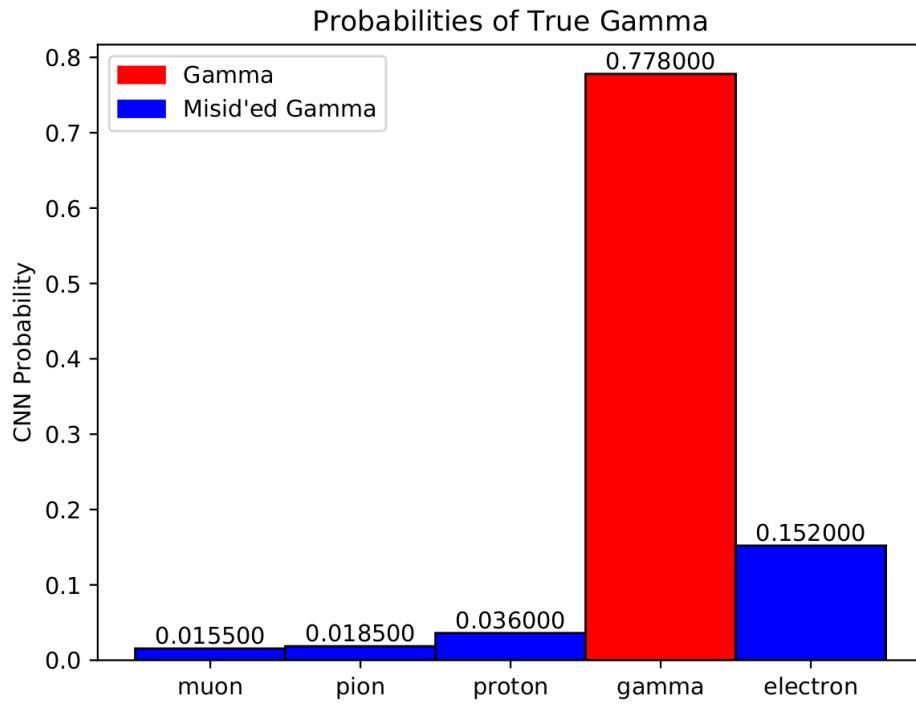
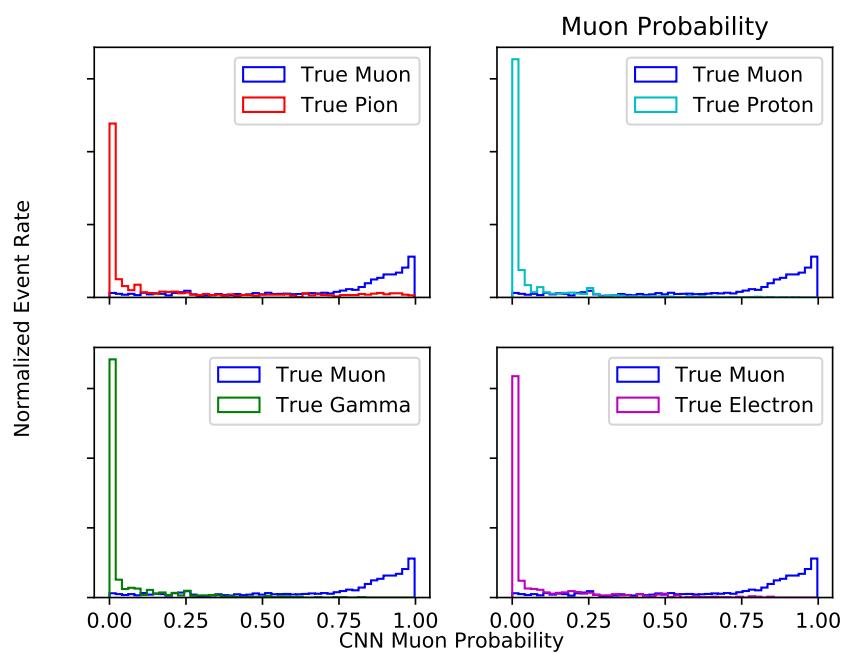
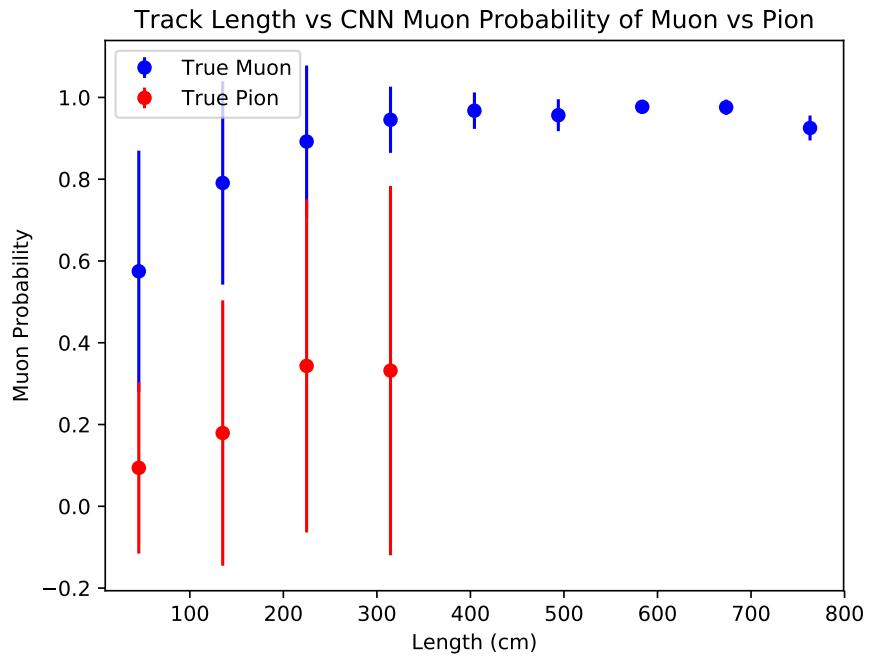


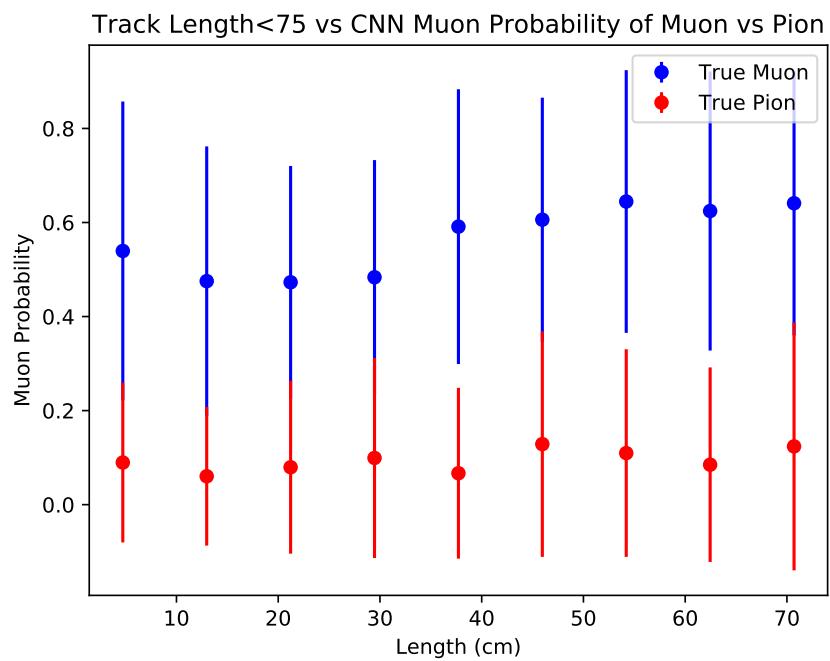
Figure 7.8: Pion Prob

**Figure 7.9:** Proton Prob**Figure 7.10:** Electron Prob

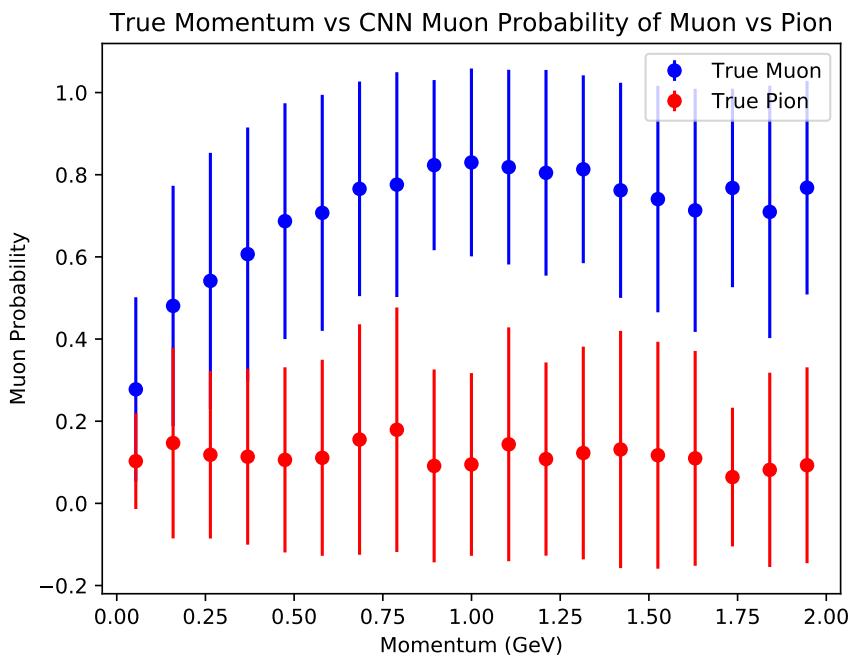
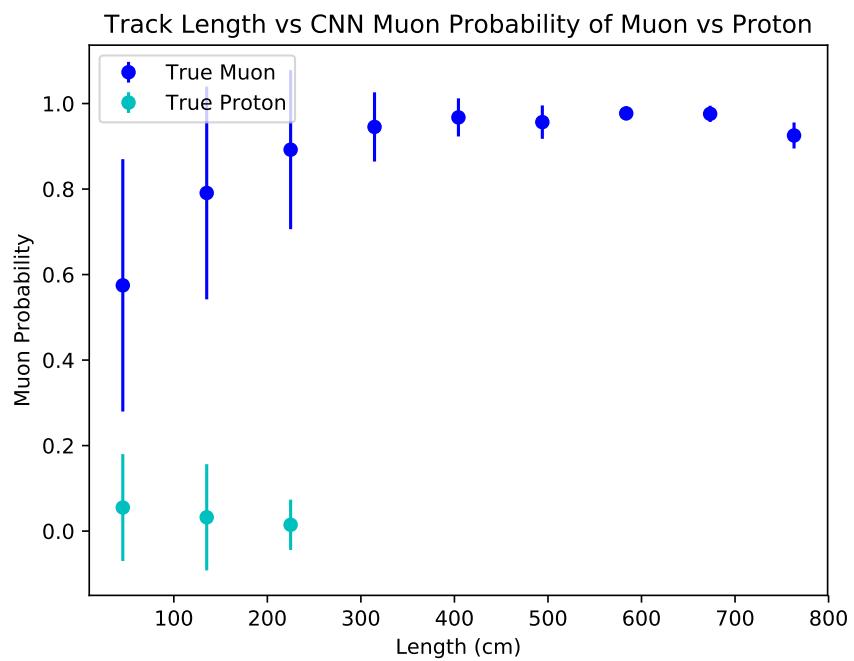
**Figure 7.11:** Gamma Prob**Figure 7.12:** Prob

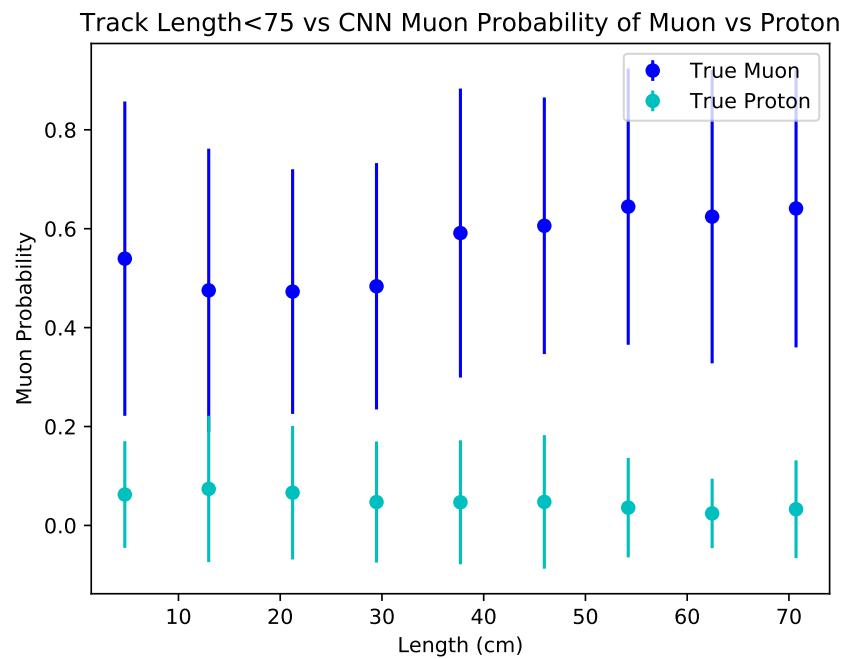


**Figure 7.13:** mupi

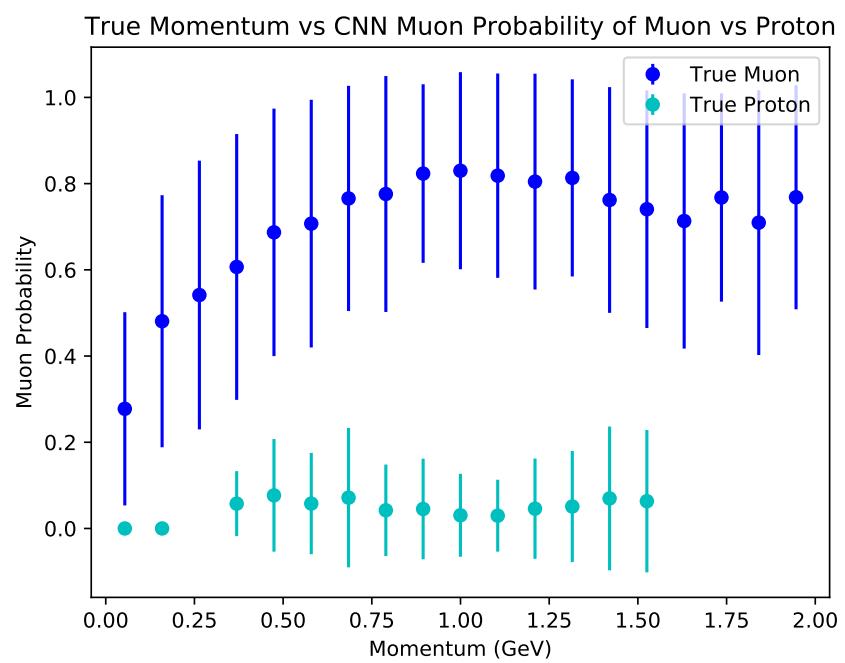


**Figure 7.14:** mupi

**Figure 7.15:** mupi**Figure 7.16:** mup



**Figure 7.17:** mup



**Figure 7.18:** mup

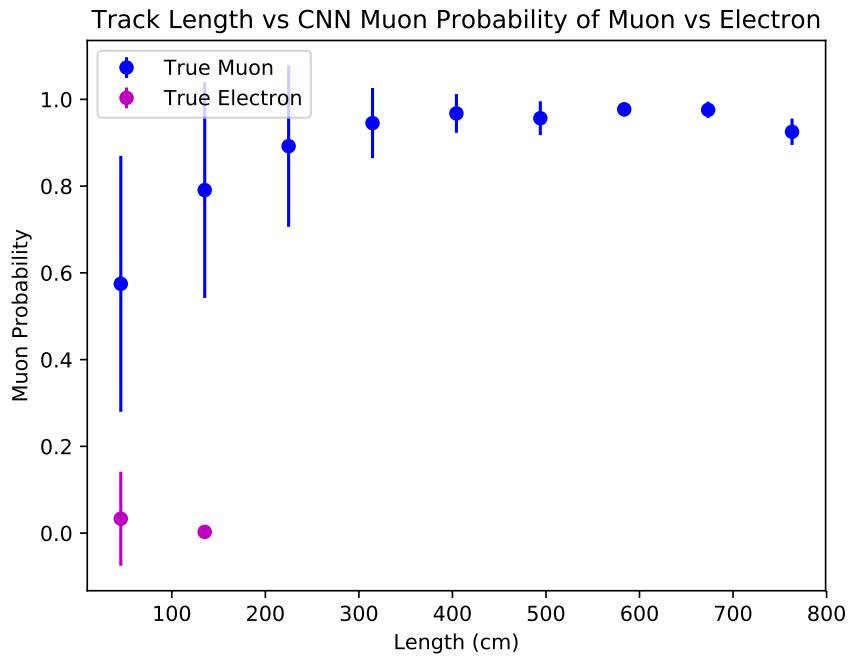


Figure 7.19: mue

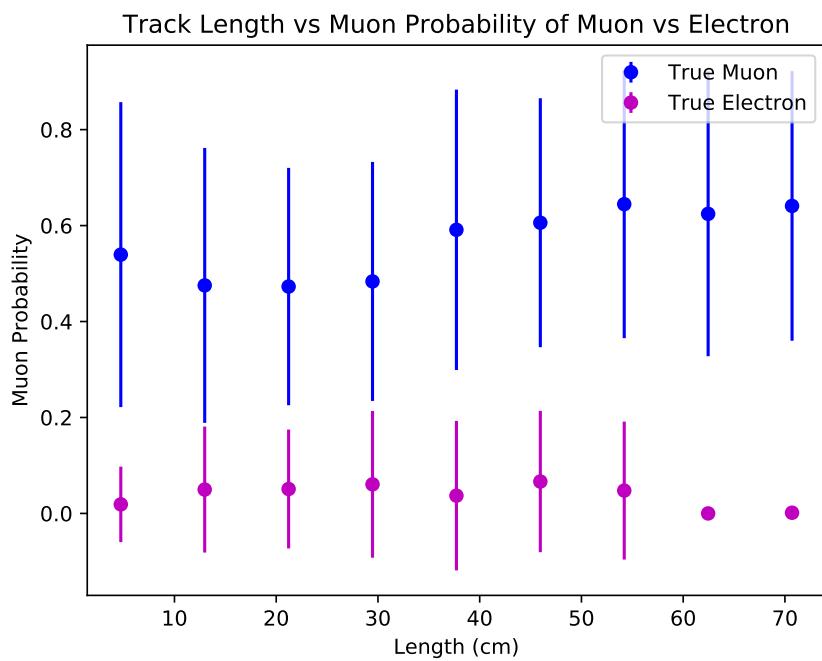


Figure 7.20: mue

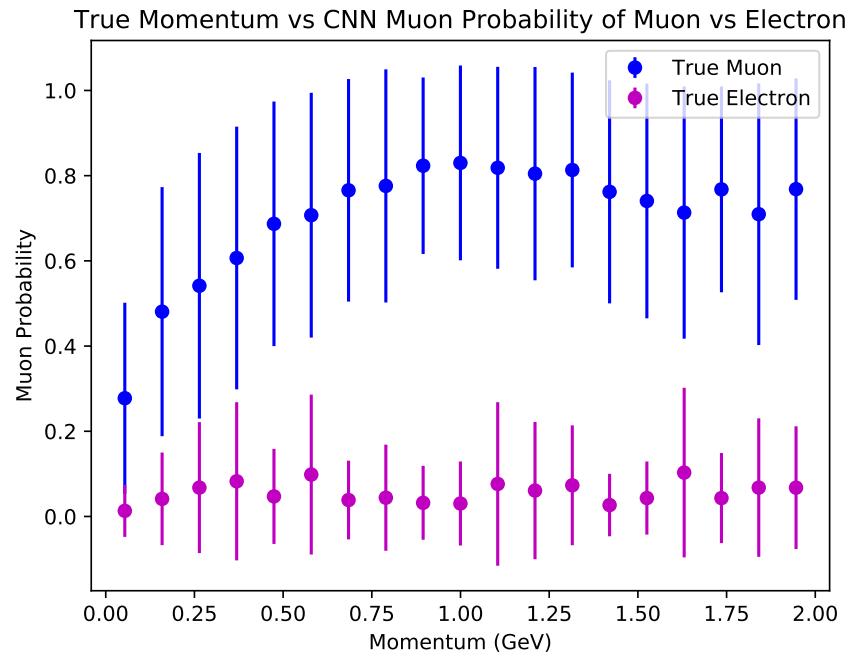


Figure 7.21: mue

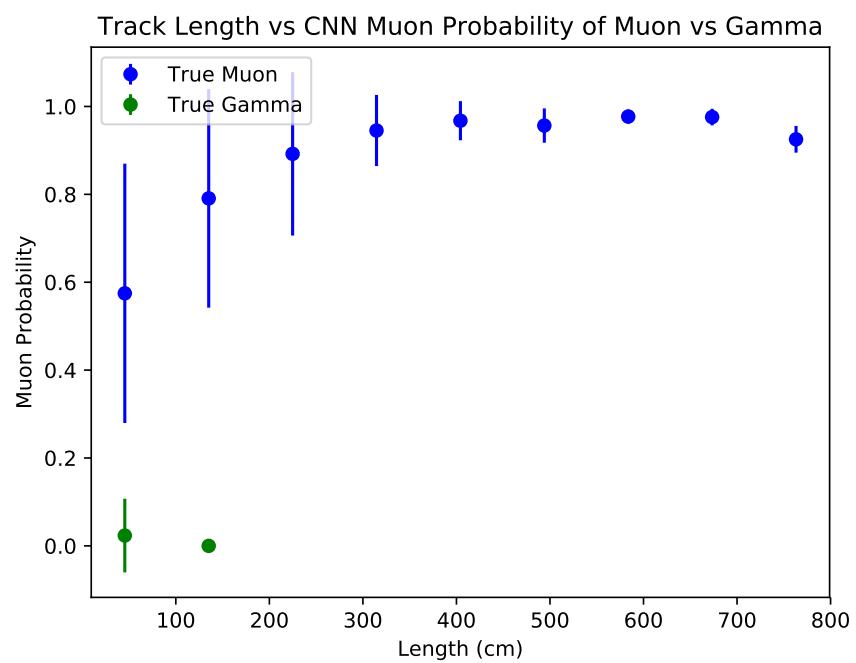


Figure 7.22: mug

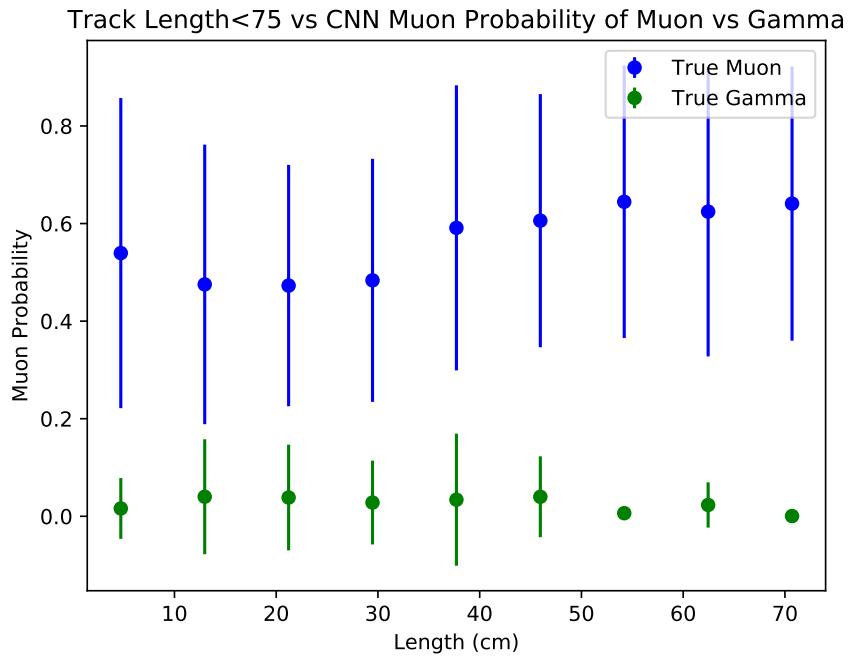


Figure 7.23: mug

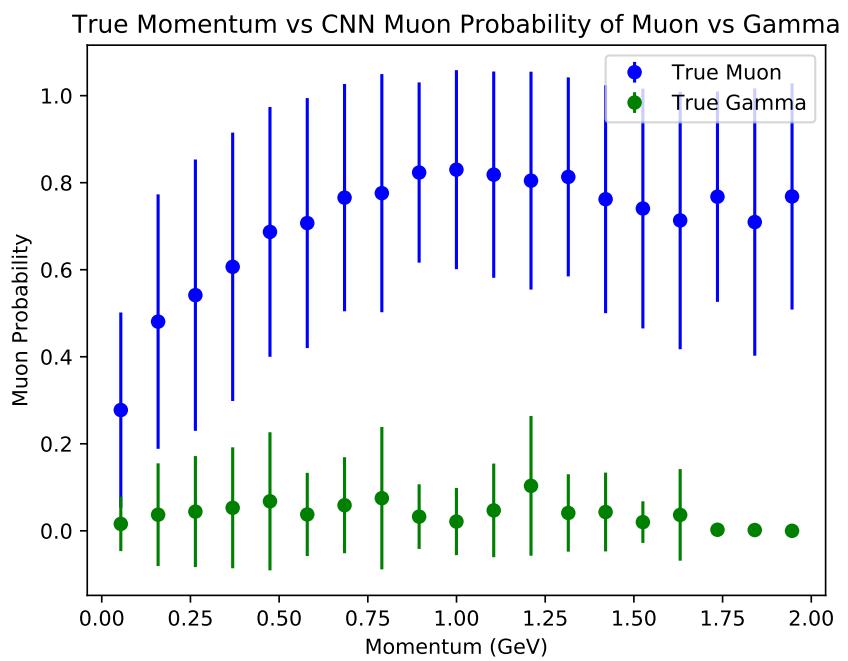


Figure 7.24: mug

# <sup>1335</sup> Chapter 8

## <sup>1336</sup> Results of Convolutional Neural Networks on particles <sup>1337</sup> WORKING <sup>1338</sup> TITLE

### <sup>1339</sup> 8.1 Classification using CNN10000

#### <sup>1340</sup> 8.1.1 Classification of MC data using Selection I Original CC-Inclusive Filter

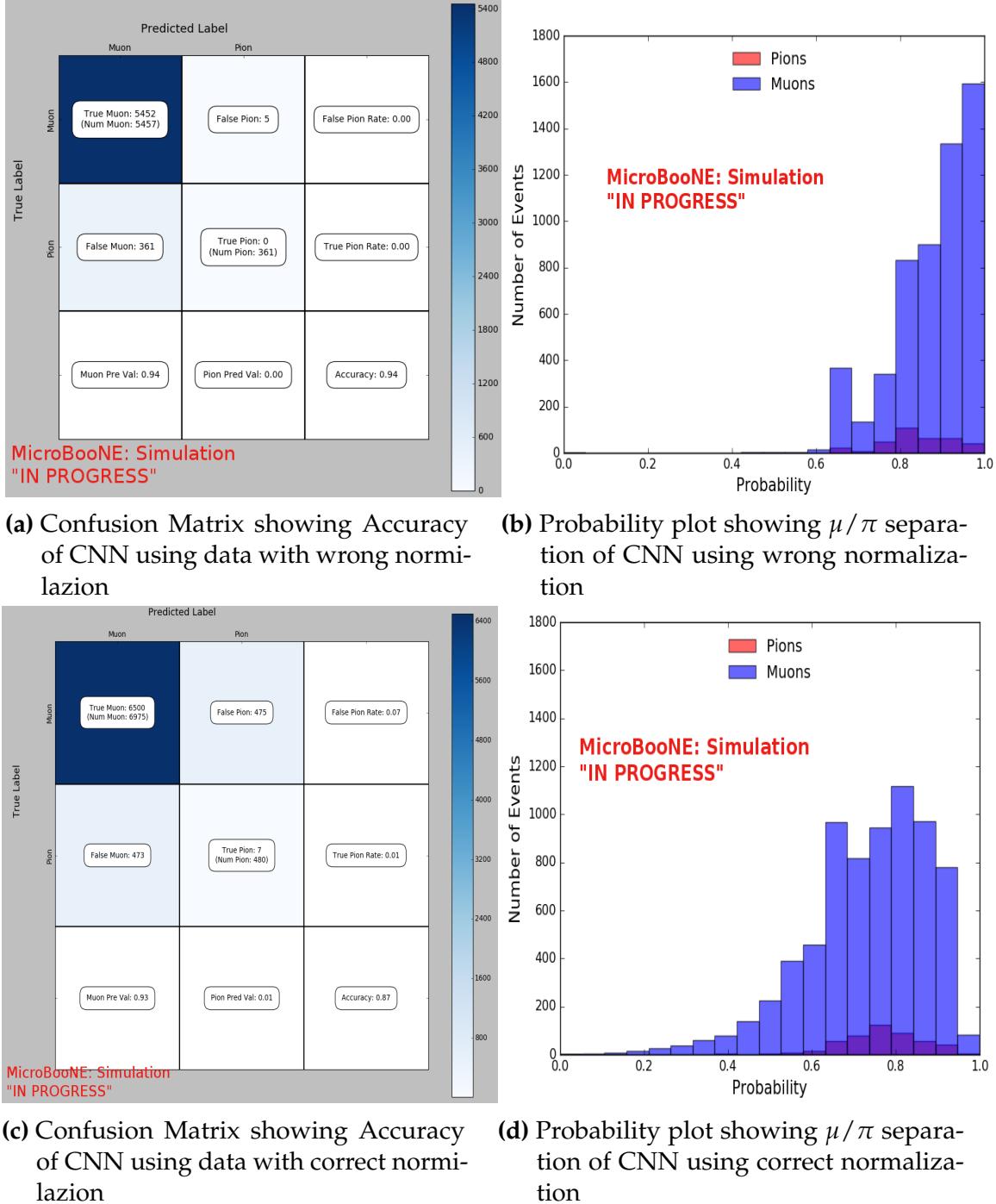
<sup>1342</sup> The next step that was taken was to use CNN10000 to classify track candidate images  
<sup>1343</sup> that were identified by the selection I original cc-inclusive filter described in [?].  
<sup>1344</sup> Passing rates for each cut in cc-inclusive filter are show in figure 8.1. For the incorrect  
<sup>1345</sup> image making normalization dataset, out of 188,880 events, 7438 passed the cut right  
<sup>1346</sup> before 75 cm track length cut which is 3.9% of total data. Discrepancies in passing rates  
<sup>1347</sup> are due to grid submission issues, however, this dataset is used to check if changes  
<sup>1348</sup> in image making normalization affects  $\mu/\pi$  separation probability due to CNN10000  
<sup>1349</sup> being trained with incorrectly image making normalized data. For the second dataset  
<sup>1350</sup> with correct image making normalization, out of 188,880 events, 9552 events passed the  
<sup>1351</sup> cut right before the 75 cm track length cut which is 5.1% passing rate and is comparable  
<sup>1352</sup> to figure 8.1. In time cosmics were also run over for efficiency and purity calculations.  
<sup>1353</sup> Out of 14395 in time cosmic events, 175 passed the cut right before the 75 cm track  
<sup>1354</sup> length cut which is a passing rate of 1.2% compared to 1.3% shown in figure 8.1.

Table 3: **Selection I: Original** The table shows passing rates for the above described event selection. Numbers are absolute event counts and Cosmic background is not scaled appropriately. The BNB+Cosmic sample contains all events, not just  $\nu_\mu$  CC inclusive. The selected events are further broken up in the following subsection. The numbers in brackets give the passing rate wrt the step before (first percentage) and wrt the generated events (second percentage). In the BNB+Cosmic MC Truth column shows how many true  $\nu_\mu$  CC inclusive events (in FV) are left in the sample. This number includes possible mis-identifications where a cosmic track is picked by the selection instead of the neutrino interaction in the same event. The cosmic only sample has low statistics, but please note that it is not used in any plots, it is just for illustrating the cut efficiency. The last column Signal:Cosmic only gives an estimate of the  $\nu_\mu$  CC events wrt the cosmic only background at each step. For this number, the cosmic background has been scaled as described in chapter 2. Note that this numbers is not a purity, since other backgrounds can't be determined at this step.

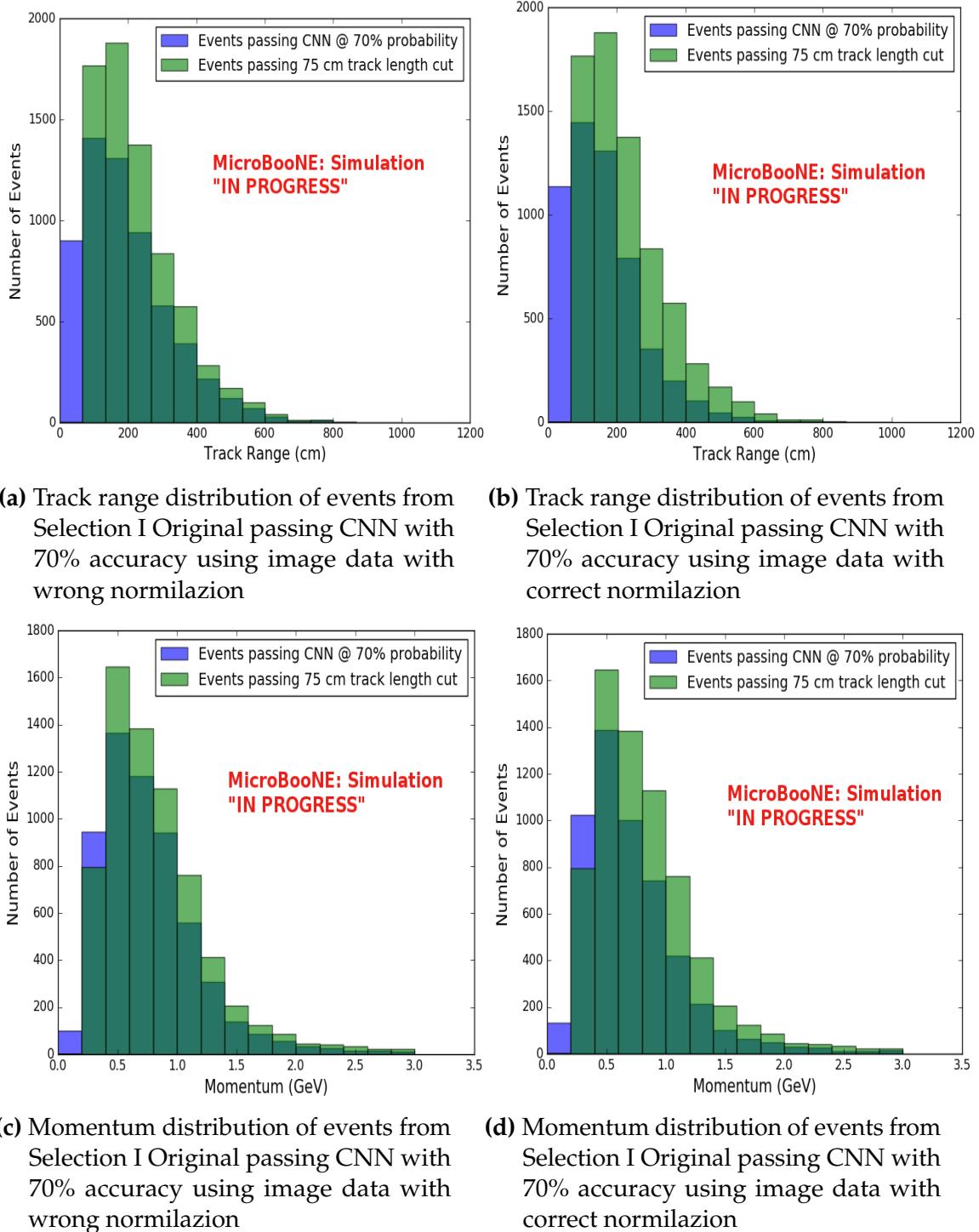
	BNB + Cosmic Selection	MC-Truth	Cosmic only	Signal: Cosmic only
Generated events	191362	45273	4804	1:22
$\geq 1$ flash with $\geq 50$ PE	136219 (71%/71%)	44002 (97%/97%)	2979 (62%/62%)	1:14
$\geq 1$ vertex in FV	131170 (96%/69%)	43794 (99%/97%)	2805 (94%/58%)	1:13
$\geq 1$ track within 5 cm of vertex	129784 (99%/68%)	43689 (99%/97%)	2756 (98%/58%)	1:13
flash matching of longest track	44775 (34%/23%)	23647 (54%/52%)	647 (23%/13%)	1:5.7
track containment	10114 (23%/5.3%)	6882 (29%/15%)	61 (9.4%/1.3%)	1:1.9
track $\geq 75$ cm	7358 (73%/3.8%)	5801 (84%/13%)	31 (51%/0.6%)	1:1.1

**Figure 8.1:** Snapshot of passing rates of Selection I from CC-Inclusive Filter

1355 Figures 8.2a, 8.2b, 8.2c and 8.2d show the accuracy and  $\mu/\pi$  separation of both the  
 1356 correct and incorrect normalized images. The confusion matrices are only composed  
 1357 of  $\mu/\pi$  data. Other particles passed the cc-inclusive filter before the 75 cm track length  
 1358 cut and were all mis-id'ed as muons. Since CNN10000 has not seen any particles  
 1359 other than muons and pions, it makes sense that those get mis-id'ed. Figures 8.2b  
 1360 and 8.2d don't have  $\mu/\pi$  separation comparable to 7.2c, but 8.2b does skew to higher  
 1361 probabilities compared to 8.2d. This is to be expected and further work on quantifying  
 1362 the performance of CNN10000 should use the incorrect image making normalization. It  
 1363 is also expected that the separation isn't as defined as the testing dataset for CNN10000.  
 1364 CNN10000 was trained and tested using single particle muons and pions and the track  
 1365 candidate dataset come from BNB+Cosmic events, not to mentions all track candidates  
 1366 have passed the cc-inclusive filter that tags "muon-like" tracks therefore the pions  
 1367 in this sample look much closer in muon topology than the network has seen. Also,  
 1368 these images were made from wire and time ticks associated to hits from the track  
 1369 candidate that passed the cc-inclusive filter. This is different from the training images  
 1370 where a bounding box was drawn over the total  $\mu$  or  $\pi$  interaction. Spurious energy  
 1371 deposition from a  $\pi - Ar$  interaction is most likely not included in the BNB+Cosmic  
 1372 images due to the tracking algorithm. To remedy this, the neural network needs to  
 1373 see more "muon-like" pions and muons and pions from a neutrino interaction passing  
 1374 the cc-inclusive filter as well as a larger particle variety including protons, photons



**Figure 8.2:** Results of CNN10000 classification of track candidate images output from cc-inclusive filter.



**Figure 8.3:** CNN10000 distributions of track candidate images output from Selection I Original cc-inclusive filter with different image data normalizations

and electrons. Although  $\mu/\pi$  separation is lacking, CNN10000 does an excellent job of classifying muons and using higher CNN probability can increase purity. Figures 8.3a, 8.3b, 8.3c and 8.3d show the track and momentum distributions for these two datasets. In both sets you have an increase in data in the bin below 75 cm and at bins below 0.5 GeV. These distributions were made with events classified with 70% probability of being a muon regardless of true particle type.

### 8.1.2 Classification of MC data using Selection I Modified CC-Inclusive Filter

Table 5: **Selection I: Modified** The table shows passing rates for the above described event selection. Numbers are absolute event counts and Cosmic background is not scaled appropriately. The BNB+Cosmic sample contains all events, not just  $\nu_\mu$  CC inclusive. The selected events are further broken up in the following subsection. The numbers in brackets give the passing rate wrt the step before (first percentage) and wrt the generated events (second percentage). In the BNB+Cosmic Mc Truth column shows how many true  $\nu_\mu$  CC inclusive events (in FV) are left in the sample. This number includes possible mis-identifications where a cosmic track is picked by the selection instead of the neutrino interaction in the same event. The cosmic only sample has low statistics, but please note that it is not used in any plots, it is just for illustrating the cut efficiency. The last column Signal:Cosmic only gives an estimate of the  $\nu_\mu$  CC events wrt the cosmic only background at each step. For this number, the cosmic background has been scaled as described in chapter 2. Note that this numbers is not a purity, since other backgrounds can't be determined at this step.

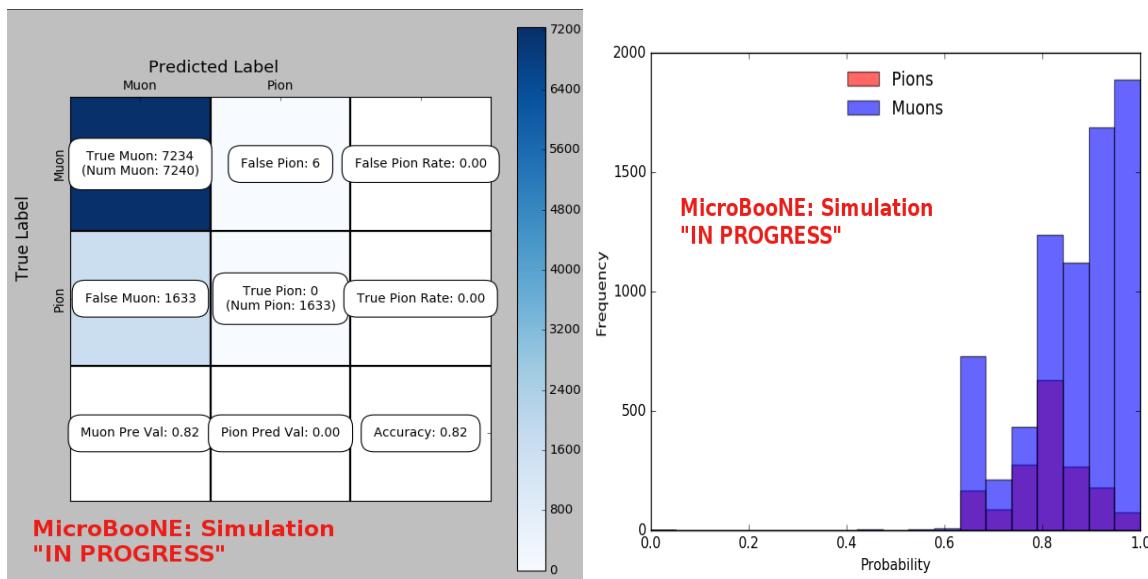
	BNB + Cosmic Selection	MC-Truth	Cosmic only	Signal: Cosmic only
Generated events	191362	45273	4804	1:22
$\geq 1$ flash with $\geq 50$ PE	136219 (71%/71%)	44002 (97%/97%)	2979 (62%/62%)	1:14
$\geq 1$ track within 5 cm of vertex	135830 (99%/71%)	43974 (99%/97%)	2975 (99%/62%)	1:14
vertex candidate in FV	79112 (58%/41%)	34891 (79%/77%)	1482 (50%/31%)	1:8.9
flash matching of longest track	40267 (51%/21%)	25891 (74%/57%)	340 (23%/7.1%)	1:2.8
track containment	19391 (48%/10%)	11693 (45%/26%)	129 (38%/2.7%)	1:2.3
track $\geq 75$ cm	6920 (36%/3.6%)	5780 (49%/13%)	17 (13%/0.4%)	1:0.6

Figure 8.4: Snapshot of passing rates of all cuts from Selection I Modified cc-inclusive filter

CNN10000 was also used to classify track candidate images that were identified by the selection I modified cc-inclusive filter described in [?]. Passing rates for each cut in this filter are shown in figure 8.4. As seen in section 8.1.1, wrong image normalization had a higher muon classification probability so all work done using selection I modified cc-inclusive filter was done using this normalization. Out of 188,880 events, 19,112 passed the cut right before the 75 cm track length cut which is a 10.1% passing rate and comparable to the 10% passing rate shown in figure 8.4. In time cosmics were also run over, out of 14,606 in time cosmics events, 302 passed the cut right before the 75 cm track length cut which is a 2.1% passing rate comparable to the 2.7% passing rate in the cc-inclusive tech-note. Figures 8.5a and 8.5b show the accuracy and  $\mu/\pi$  separation. Both plots are only composed of muons and pions and like selection I original data,

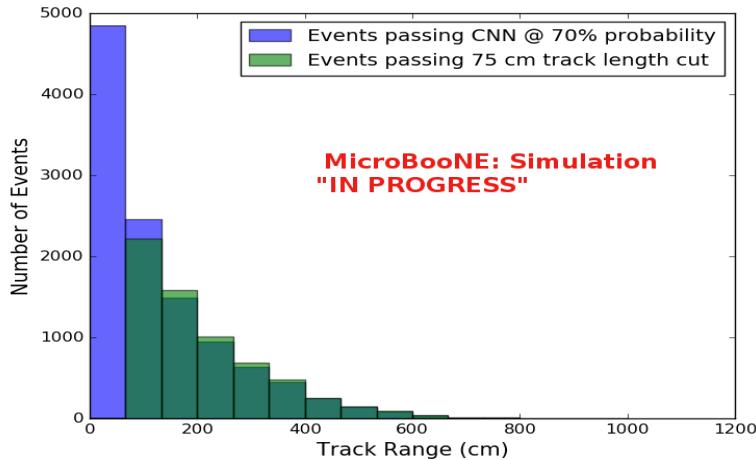
1394 all other particles were id'ed as muons. Also like selection I original data, muons are  
 1395 being identified at a very high rate. Figure 8.6a shows the track range distributions  
 1396 of all events from selection I modified being classified by the CNN as a muon with a  
 1397 probability of 70% regardless of true particle type. We get entries for the CNN curve  
 1398 in the lowest bin and none for the 75 cm curve. To see how many true CC events  
 1399 were identified by CNN10000 breaking down figure 8.6a by event type was necessary.  
 1400 Figures 8.6b and 8.6c show track range distributions separated by signal and various  
 1401 backgrounds. Particle type was not taken into consideration in these plots so true CC  
 1402 event images can be any track candidate particle passing selection I modified cut right  
 1403 before track length cut including pions and protons.

1404 To gain an even deeper understanding on how CNN10000 is performing, plotting  
 1405 these distributions with only muons and pions was done due to the fact that CNN10000  
 1406 was trained with only those particles for  $\mu/\pi$  separation. Figures 8.6d-8.7d show the  
 1407 stacked histograms of signal and background of the track range distributions with  
 1408 varying CNN probabilities starting from 70% and ending at 90% probability. With  
 1409 higher probabilities we get a purer sample in the lower bin but we end up losing  
 1410 events as well. Momentum distributions for all signal/background events are shown  
 1411 in figure 8.8.

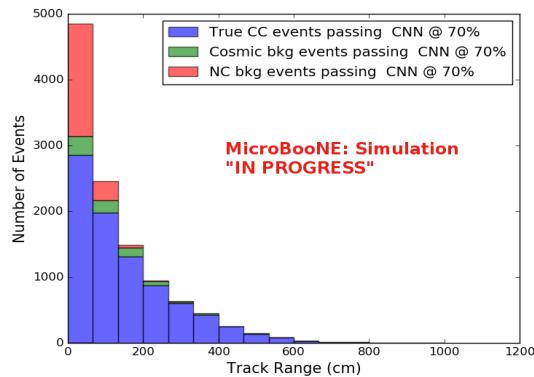


(a) Confusion Matrix for CNN10000 classified events from selection I modified      (b) Probability plot for CNN10000 classified events from selection I modified

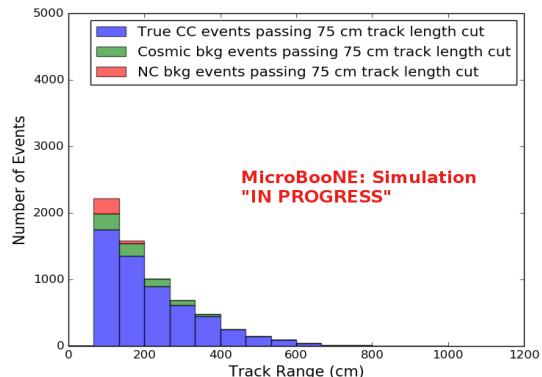
**Figure 8.5:** Confusion matrix and probability plot of events passing selection I modified cc-inclusive cuts right before 75cm track length cut



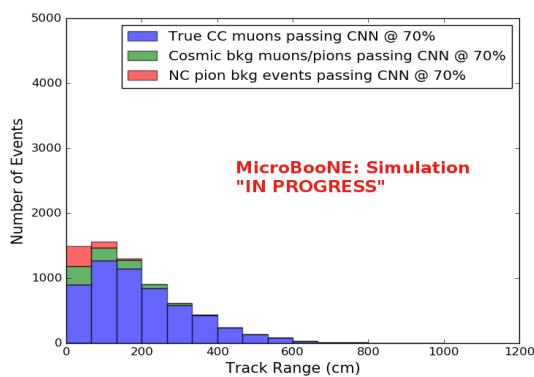
(a) Track range distribution of events from Selection I Modified passing CNN with 70% accuracy



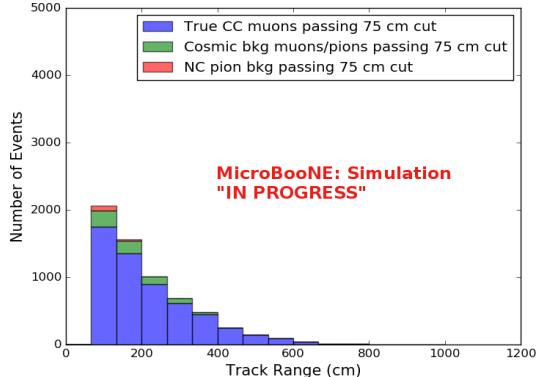
(b) Stacked signal and background track range distributions from Selection I Modified passing CNN with 70% accuracy



(c) Stacked signal and background track range distributions from Selection I Modified passing 75 cm track length cut

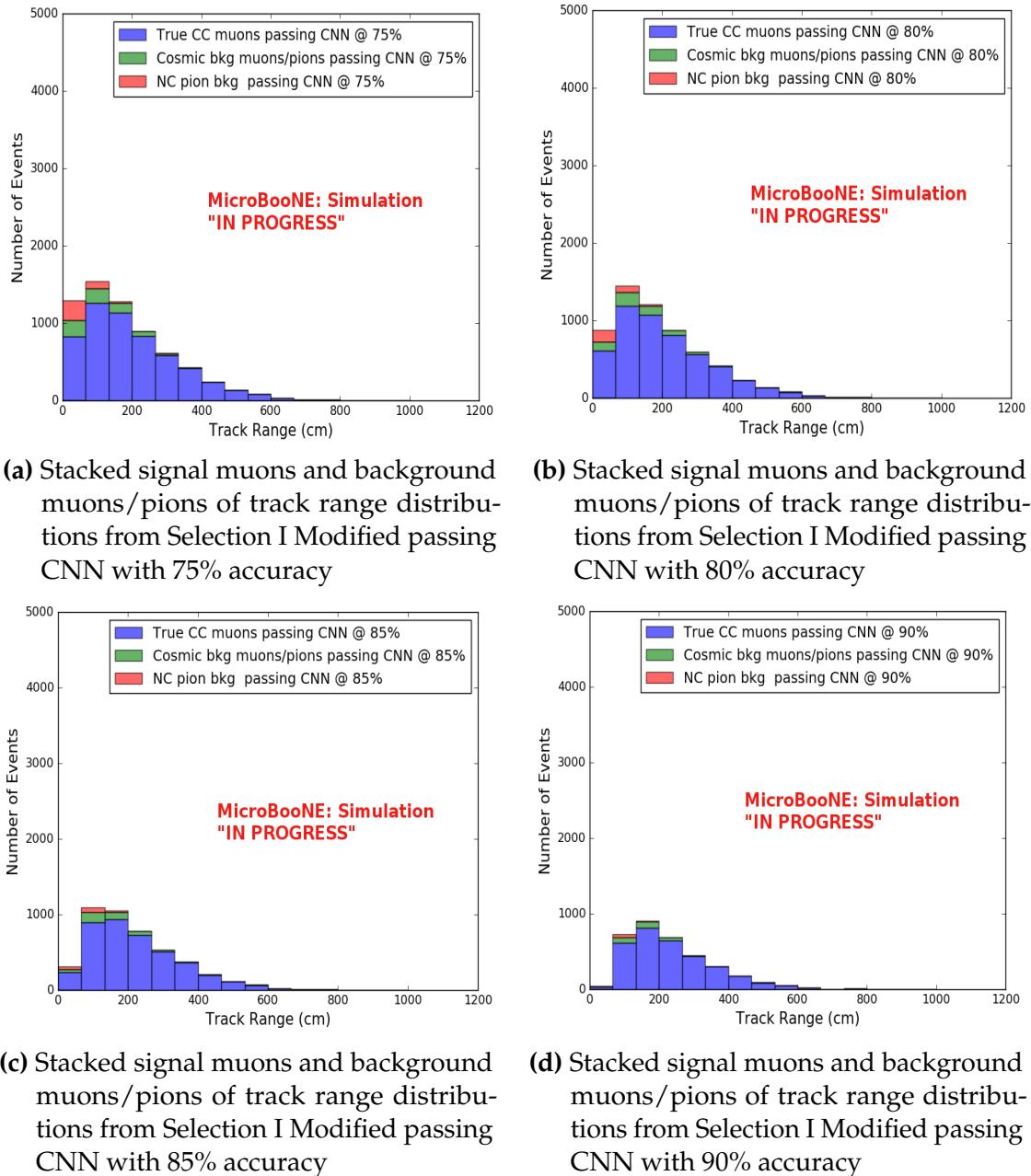


(d) Stacked signal muons and background muons/pions of track range distributions from Selection I Modified passing CNN with 70% accuracy

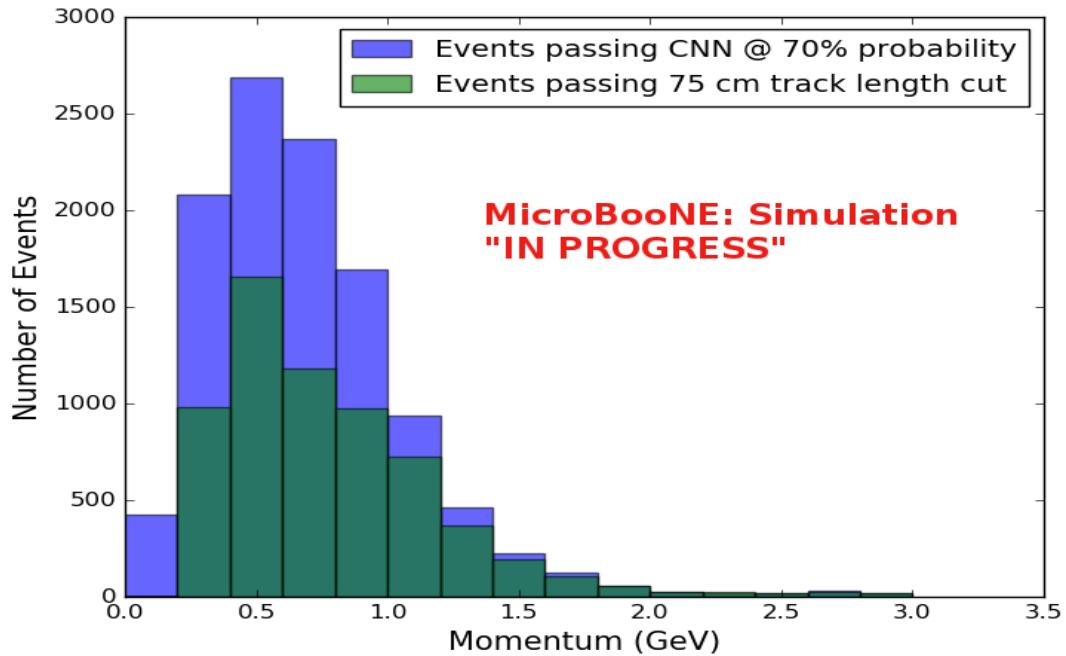


(e) Stacked signal muons and background muons/pions of track range distributions from Selection I Modified passing 75 cm track length cut

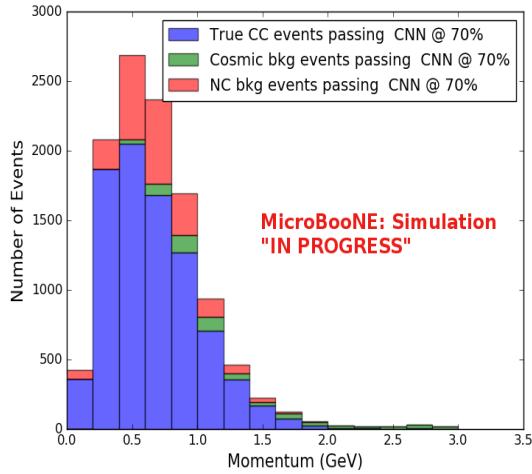
**Figure 8.6:** CNN10000 distributions of track candidate images output from Selection I Modified cc-inclusive filter



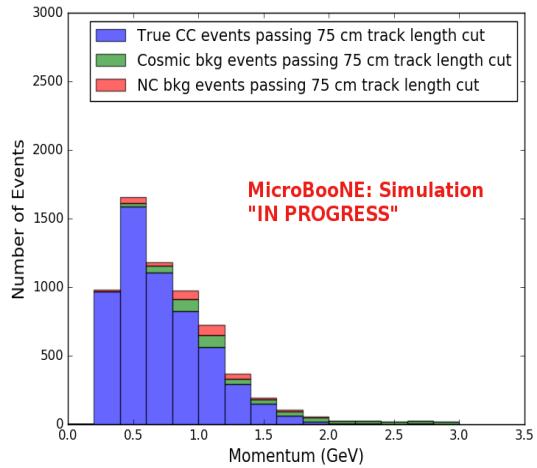
**Figure 8.7:** CNN10000 stacked signal/background track range distributions of track candidate images output from Selection I Modified cc-inclusive filter



(a) Momentum distribution of events from Selection I Modified passing CNN with 70% accuracy



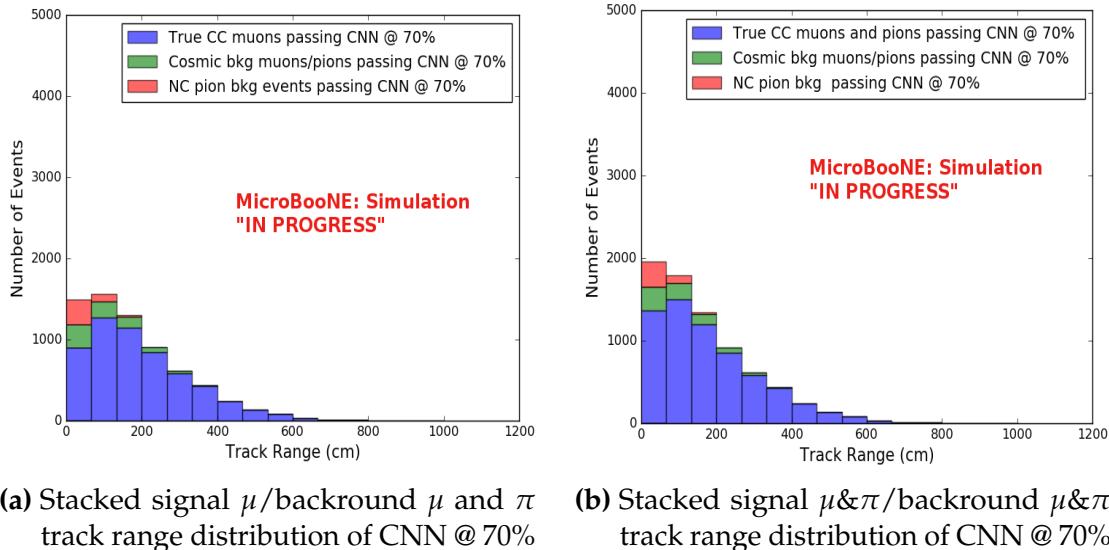
(b) Stacked signal and background momentum distributions from Selection I Modified passing CNN with 70% accuracy



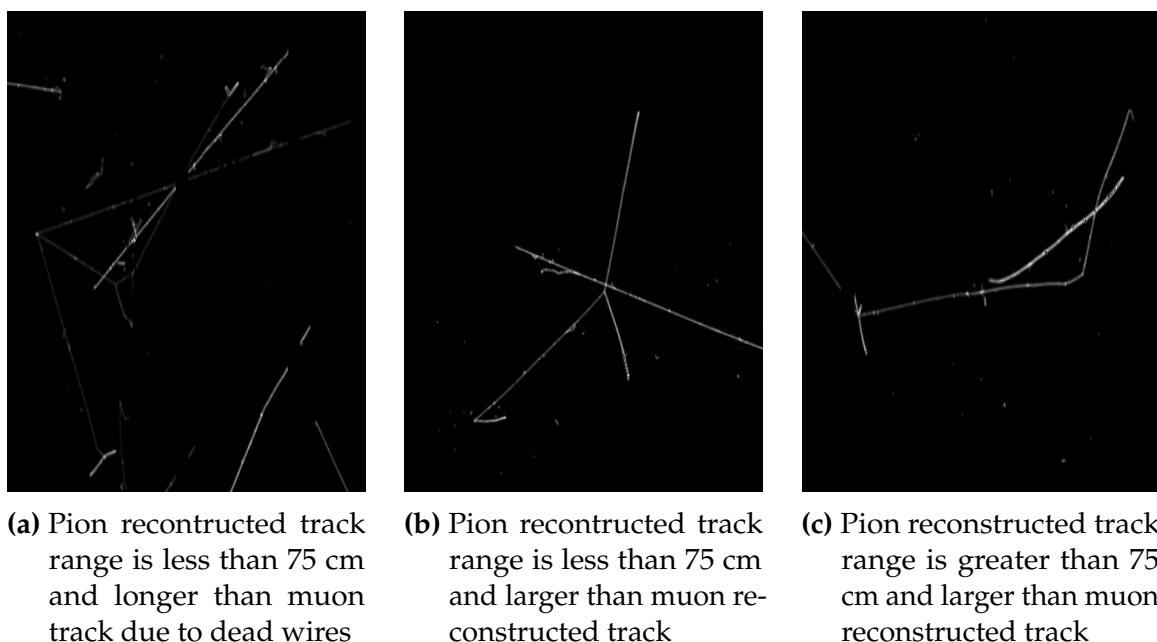
(c) Stacked signal and background momentum distributions from Selection I Modified passing 75 cm track length cut

**Figure 8.8:** CNN10000 momentum distributions of track candidate images output from Selection I Modified cc-inclusive filter

Another check was to see if any true CC pions were passing through the cut right before the 75 cm track length cut. Figure 8.9 shows the comparison of the stacked track range distribution with only true CC muon signal versus the stacked distribution with true CC muons and pions signal. As you can see, we gain more events when plotting CC events with a particle type of either muons or pions due to the CNN classifying all pions in this dataset as muons. This is an interesting scenario and a sample of topologies of these images are represented in figure 8.10, at least 3 tracks are coming out of the vertex for these types of events. With the 75 cm track length cut, the selection is cutting event topologies like this where the pion is the tagged track candidate. Figure 8.10a has a defined longer muon track, but because of dead wires through the track, the reconstructed range is 1. less than 75 cm and 2. shorter than the reconstructed pion whose length is also less than 75 cm. This is a very interesting event, but because of issues with the tracking algorithm, the 75 cm cut would get rid of this event. The CNN was able to recover this event only because it has classified all pions as muons. Figure 8.10b shows the second case to think about, the pion, while still less than 75 cm has a reconstructed track length longer than the muon. Again, the CNN recovered this event due to pions being classified as muons. Lastly, figure 8.10c shows a pion with a reconstructed track length greater than 75 cm and the muon. These three cases show that a broader question must be asked when training the network other than is it a muon or pion. There are different routes to recover interesting events like these. One route is to ask the network “Is it a CC event or is it an NC event?” and obtain an image dataset consisting of whole CC/NC events that will train the network to answer this question. The other route is to ask the network “Is this a  $\mu/\pi/p/$  from a CC event or NC event and obtain an image dataset consisting of primary particles from a CC/NC event. Both these paths will be explored in future work.



**Figure 8.9:** Track distribution comparisons of true CC muons plotted vs true CC muons and pions plotted



**Figure 8.10:** Images of true CC events where the pion was the tagged track candidate

		BNB + Cosmics		Cosmic Only	Signal: Cosmic Only
		Selection	MC Truth		
75 cm Cut passing rates	Generated Events	191362	45723	4804	1:22
	Track Containment	19391 (48%/10%)	11693 (45%/26%)	129 (38%/2.7%)	1:2.3
	track $\geq$ 75 cm	6920 (36%/3.6%)	5780 (49%/13%)	17 (13%/0.4%)	1:0.6
CNN passing rates	Generated Events	188880	44689	14606	1:21
	Track Containment	19112 ( /10%)	11554 ( /26%)	302 ( /2.1%)	1:1.73
	CNN cut @ 70% Probability	16502 (86%/8.7%)	10605 (92%/23%)	205 (68%/14%)	1:1.28
	CNN cut @ 83% Probability	7511 (46%/4.0%)	6142 (58%/14%)	32 (16%/0.2%)	1:0.4

**Table 8.1:** Comparing passing rates of CNN at different probabilities versus 75 cm track length cut: Numbers are absolute event counts and Cosmic background is not scaled appropriately. The BNB+Cosmic sample contains all events. The numbers in brackets give the passing rate wrt the step before (first percentage) and wrt the generated events (second percentage). In the BNB+Cosmic MC Truth column shows how many true  $\nu_\mu$  CC-inclusive events (in FV) are left in the sample. This number includes possible mis-identifications where a cosmic track is picked by the selection instead of the neutrino interaction in the same event. The CNN MC True generated events were scaled wrt the MC True generated events for the 75 cm cut passing rates due to only running over 188,880 generated events versus the 191362 generated events. The last column Signal:Cosmic only gives an estimate of the  $\nu_\mu$  CC events wrt the cosmic only background at each step. For this number, the cosmic background has been scaled as described in [?]. Note that these numbers are not a purity, since other backgrounds can't be determined at this step.

Signal	$\nu_\mu$ CC events with true vertex in FV	#Events(Fraction)	#Events(Fraction)
		passing CNN @ 70% Probability	passing CNN @ 83% Probability
		10605(35%)	6142(61%)
Backgrounds	Cosmics Only Events	13573(45%)	2582(26%)
	Cosmics in BNB Events	2249(7.4%)	492(4.9%)
	NC Events	3412(11%)	778(7.7%)
	$\nu_e$ and $\bar{\nu}_e$ Events	139(0.5%)	32(0.3%)
	$\bar{\nu}_\mu$ Events	97(0.3%)	67(0.7%)

**Table 8.2:** Signal and background event numbers at modified selection level with CNN cut estimated from a BNB+Cosmic sample and Cosmic only sample normalized to  $5 * 10^{19}$  PoT. The last column gives the fraction of this signal or background type to the total selected events per CNN probability.

Table 8.1 shows the passing rates for the 75 cm track length cut and the CNN cut at 70% and 83%. The passing rates at the track containment level for the 75 cm track length cut compared to the CNN are comparable with only a 0.6% difference in the in time cosmic bin which may be due in part to the larger in time cosmic statistics used for the CNN dataset. These passing rates need to be comparable to then be able to compare the passing rates after the CNN cut to the 75 cm cut. Again, the same BNB+Cosmic sample was used for both selection I modified with 75 cm cut and selection I modified with CNN cut. As it stands, a CNN cut at 83% probability has

<sup>1445</sup> a MC true CC event passing rate of 14% compared to the 13% passing rate of the 75  
<sup>1446</sup> cm track length cut. The Signal:Cosmic Only background is also reduced from 1:0.6  
<sup>1447</sup> to 1:0.4 The total passing rate is also higher than the 75 cm cut, 3.6% vs 4.0%. Table  
<sup>1448</sup> 8.2 shows the breakdown of signal and backgrounds for the CNN at the different  
<sup>1449</sup> probabilities. We have a 61% signal passing rate with the CNN cut @ 83% versus the  
<sup>1450</sup> 53.8% signal passing rate of the 75 cm cut.

<sup>1451</sup> Based on these numbers, the following performance values of the modified selec-  
<sup>1452</sup> tion with 75 cm cut versus modified selection with CNN @ 83% probability cut were  
<sup>1453</sup> calculated:

- <sup>1454</sup> • Efficiency: Number of selected true  $\nu_\mu$  CC events divided by the number of  
<sup>1455</sup> expected true  $\nu_\mu$  CC events with interaction in the FV.
  - <sup>1456</sup> – Selection I modified: 13%
  - <sup>1457</sup> – Selection I modified with CNN cut @ 83% probability: 14%
- <sup>1458</sup> • Purity: Number of selected true  $\nu_\mu$  CC events divided by sum of itself and the  
<sup>1459</sup> number of all backgrounds.
  - <sup>1460</sup> – Selection I modified: 53.8%
  - <sup>1461</sup> – Selection I modified with CNN cut @ 83% probability: 61%

<sup>1462</sup> Lastly, figure 8.12 shows a more representative performance of the CNN. Due to  
<sup>1463</sup> the fact that the CNN was trained on muons and pions, showing the performance  
<sup>1464</sup> of CC muon events versus NC pion events with respect to CNN probability gives a  
<sup>1465</sup> better picture of how the network is performing. Figure 8.12 shows that at 83% we  
<sup>1466</sup> are below the 75 cm cut NC pion threshold and still above the CC muon threshold.  
<sup>1467</sup> Using 83% probability not only reduced the NC pion background, it also dramatically  
<sup>1468</sup> reduced the in time cosmics and cosmics in the BNB.

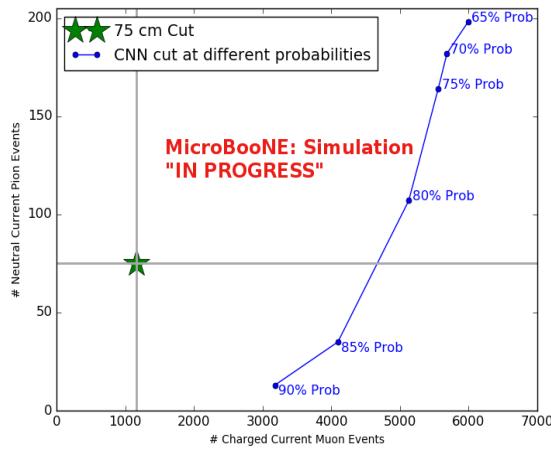
### <sup>1469</sup> 8.1.3 Conclusions and Future Work

<sup>1470</sup> It was shown that even though CNN10000 was trained with single particle generated  
<sup>1471</sup> muons and pions, it performs fairly well at classifying track candidate images from  
<sup>1472</sup> BNB+Cosmic events. Events have been regained below the 75 cm track length cut and  
<sup>1473</sup> the momentum and track range distributions have similar shapes to the distributions of  
<sup>1474</sup> Selection I original and modified. Efficiencies and purities were calculated for selection

Table 8: **Selection I: Modified** Signal and background event numbers at modified selection level estimated from a BNB+Cosmic sample and Cosmic only sample normalized to  $5 \times 10^{19}$  PoT. The last column gives the fraction of this signal or background type to the total 2189 selected events.

Signal	#Events	
$\nu_\mu$ CC events with true vertex in FV	1168	53.8%
Backgrounds		
Cosmics only events	725	33.4%
Cosmics in BNB events	144	6.6%
NC events	75	3.5%
$\nu_e$ and $\bar{\nu}_e$ events	4	0.2%
$\bar{\nu}_\mu$ events	15	0.7%
$\nu_\mu$ CC events with true vertex outside FV	40	1.8%

**Figure 8.11:** Snapshot of signal and background event numbers of Selection I modified from cc-inclusive note [?]



**Figure 8.12:** CNN performance of classified muons and pions compared to the already implemented 75 cm track length cut

<sup>1475</sup> I modified events before 75 cm track length cut with the CNN at 83% probability and  
<sup>1476</sup> are 14% and 62% respectively. Although the CNN doesn't have separation between  
<sup>1477</sup> muons and pions and although all particles passing CNN are classified as muon,  
<sup>1478</sup> increasing CNN probability allows us to increase the purity as well as maintain an  
<sup>1479</sup> efficiency comparable to the 75 cm track length cut all while recovering events below  
<sup>1480</sup> that 75 cm cut. Out of the 6142 events that passed the CNN @ 83% 1470 events were  
<sup>1481</sup> below the 75 cm cut, a recovery of 3.3% of data with a purity of 15%. Although  
<sup>1482</sup> these numbers are low, it is an improvement from the selection I modified in both total  
<sup>1483</sup> efficiency and purity and an increase in phase space by recovering these events.

## <sup>1484</sup> **8.2 Classification using CNN100000**

<sup>1485</sup> All future classifications will be done using Selection I Modified CC-Inclusive Filter  
<sup>1486</sup> because it has a higher efficiency and purity than Selection I Original CC-Inclusive  
<sup>1487</sup> Filter. To reiterate, CNN100000 was trained using 20,000 images of each  $\mu/\pi/p/\gamma/e$ .  
<sup>1488</sup> The results of using CNN100000 to classify BNB+Cosmics will be outlined below.

### <sup>1489</sup> **8.2.1 Classification of MC data using Selection I Modified 1490 CC-Inclusive Filter**

### <sup>1491</sup> **8.2.2 Classification of MicroBooNE data using Selection I Modified 1492 CC-Inclusive Filter**

### <sup>1493</sup> **8.2.3 Comparing two CC-Inclusive Cross Section Selection Filters**



<sup>1494</sup> **Chapter 9**

<sup>1495</sup> **Conclusion**

<sup>1496</sup> Your Conclusions here.

<sup>1497</sup>



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