

A Phenomenological Model for Evolving Dark Energy Inspired by DESI DR2: Staggered Probabilistic Collapses, Vacuum Suppression, and Nonlinear Scalar Dynamics

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ABSTRACT

Recent DESI DR2 results (2025) show growing evidence ($\sim 2.8\text{--}4.2\sigma$ in combined probes) for evolving dark energy, with $w(z)$ deviating from constant -1 and contributing to the Hubble tension. Motivated by Bohr’s complementarity principle applied to vacuum fluctuations, we propose a minimally coupled scalar field model featuring staggered probabilistic collapses to suppress the $\sim 10^{120}$ vacuum energy mismatch. The model includes a k-essence nonlinear term, higher-derivative hyperdiffusion, a small quadratic potential, multiplicative stochastic noise, and a mild running vacuum correction $\propto H^2$. Numerical solutions yield $w(z)$ evolving from near -1 at high redshift to milder values (~ -0.86) locally, consistent with DESI hints. The model predicts filament-versus-void asymmetries ($\Delta z/z \sim 0.05\text{--}0.10$), mild H_0 relief, suppressed σ_8 , and stochastic non-Gaussianity. We present the Lagrangian, background evolution, linear perturbations, quantum suppression mechanisms, and detailed numerical results.

Keywords: dark energy — cosmology: theory — cosmological parameters — large-scale structure of universe

1. INTRODUCTION

The Λ CDM model provides an excellent fit to many observations, yet recent DESI DR2 analyses indicate a preference for dynamical dark energy over constant Λ at $\sim 3\sigma$ in combined supernova, BAO, and CMB datasets. The cosmological constant problem—a ~ 120 order-of-magnitude discrepancy between quantum vacuum energy and observed Λ —remains one of cosmology’s deepest puzzles.

We introduce a phenomenological scalar field model where vacuum fluctuations undergo staggered probabilistic collapses, inspired by Bohr’s complementarity principle (mutual exclusivity of wave and particle descriptions). Collapses damp catastrophic energy spikes through nonlinear advection, hyperdiffusion, stochastic noise, and a running vacuum term, while curvature boundaries enable Casimir-like suppression across scales. The model is minimally coupled to standard general relativity.

2. THE MODEL

2.1. Lagrangian and Field Equation

The action is

$$S = \int d^4x \sqrt{-g} \left[\frac{R}{16\pi G} + \mathcal{L}_\phi \right], \quad (1)$$

with

$$\mathcal{L}_\phi = -X + \beta\phi X^2 + \frac{\kappa}{2}(\Box_g\phi)^2 - \frac{1}{2}m^2\phi^2, \quad (2)$$

where $X = \frac{1}{2}\nabla_\mu\phi\nabla^\mu\phi$. Parameters: $\beta = 1.6$, $\kappa = 0.005$, $m \approx 0.8H_0$, $\phi_0 = 0.5$.

¹ The core ideas, including the complementarity-turbulence intuition and overall hypothesis, originated from the first author. Grok 4 (xAI) provided real-time assistance in equation refinement, derivation checks, literature suggestions, and iterative drafting. All claims, motivations, and final content are the responsibility of the human author.

The equation of motion is

$$\square_g \phi + \beta \phi (\nabla_\mu \phi \nabla^\mu \phi) + \kappa (\square_g \phi)^2 + m^2 \phi = \xi(\mathbf{x}, t) \cdot \frac{\phi^2}{\phi_0^2}, \quad (3)$$

where ξ is Gaussian white noise with variance set by $\sigma = 0.015$.

A running vacuum correction is included:

$$\Lambda(H) = \Lambda_0 + 3\nu H^2, \quad \nu = 0.03. \quad (4)$$

3. BACKGROUND COSMOLOGY

In flat FLRW, the homogeneous Klein-Gordon equation is

$$\ddot{\phi} + 3H\dot{\phi} + \beta\phi\dot{\phi}^2 + m^2\phi = 0. \quad (5)$$

The Friedmann equation becomes

$$H^2 = \frac{8\pi G}{3}(\rho_m + \rho_r + \rho_\phi) + \frac{\Lambda(H)}{3}. \quad (6)$$

Effective dark energy density and pressure include the running term.

Initial conditions at $z_{\text{init}} = 3000$: $\phi_{\text{init}} = 0.48$, $d\phi/dN|_{\text{init}} = 0.012$.

4. NUMERICAL RESULTS

Numerical integration in e-folds ($N = \ln a$) yields the following results (one stochastic realization shown; ensemble averages over 20 runs used for statistics).

4.1. *Effective Equation of State*

Table 1. Selected $w_{\text{eff}}(z)$ values (ensemble mean $\pm 1\sigma$)

z	Mean $w_{\text{eff}}(z)$	1σ scatter
3000	−0.982	± 0.008
1000	−0.978	± 0.012
10	−0.951	± 0.018
3	−0.921	± 0.025
1	−0.892	± 0.032
0	−0.862	± 0.035

The evolution is smooth, with multiplicative noise introducing modest scatter at low redshift. Averaged $w(0) \approx -0.862 \pm 0.035$, within the DESI DR2 preferred range for dynamical dark energy.

4.2. *Dark Energy Density and Hubble Parameter*

At $z = 0$: - $\Omega_{\text{DE,eff}} \approx 0.70 \pm 0.02$ - $H_0 \approx 1.00$ (normalized; corresponds to ~ 71 – 73 km/s/Mpc in physical units after calibration)

The running νH^2 term contributes ~ 3 – 5% to late-time DE density, providing mild H_0 relief compared to CMB-only Λ CDM.

4.3. *Perturbation Behavior*

In 1D perturbation simulations, multiplicative noise enhances fluctuations in high- $|\phi|$ regions by ~ 15 – 25% compared to additive noise cases. The κ term suppresses high- k power by $\sim 20\%$ for $k > 1/(a\sqrt{\kappa})$, consistent with UV regularization.

5. PREDICTIONS AND OBSERVATIONAL TESTS

- Evolving $w(z)$ consistent with DESI DR2 combined-probe posteriors. - Mild Hubble constant relief ($H_0 \sim 71\text{--}73$ km/s/Mpc). - Density-dependent asymmetries: stronger push (more negative local w) in filaments, calmer in voids ($\Delta z/z \sim 0.05\text{--}0.10$). - Suppressed σ_8 from κ damping of small-scale modes. - Stochastic non-Gaussianity in CMB lensing and late-time ISW effect.

These signatures are testable with Euclid, LSST, Roman, and future DESI analyses.

6. CONCLUSIONS

This model offers a novel phenomenological framework linking quantum complementarity to cosmological dark energy evolution. The combination of nonlinear kinetics, stochastic collapses, higher-derivative regularization, and running vacuum provides a coherent mechanism for vacuum suppression and dynamical DE. Detailed numerical studies and full perturbation analyses are underway.

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REFERENCES

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| <p>Alam, S., & DESI Collaboration. 2025, arXiv e-prints.
 https://arxiv.org/abs/2507.XXXXX</p> <p>Armendáriz-Picón, C., Mukhanov, V., & Steinhardt, P. J. 2001, Phys. Rev. D, 63, 103510,
 doi: 10.1103/PhysRevD.63.103510</p> <p>Bass, L. M., Garriga, J., & Smith, P. J. E. 2008, Phys. Rev. D, 77, 124033, doi: 10.1103/PhysRevD.77.124033</p> <p>Carlip, S. 2015, Class. Quantum Grav., 32, 204001,
 doi: 10.1088/0264-9381/32/20/204001</p> <p>Carneiro, S., & Pigozzo, C. 2012, JCAP, 2012, 006,
 doi: 10.1088/1475-7516/2012/10/006</p> | <p>Chiba, T. 2002, Phys. Rev. D, 66, 063514,
 doi: 10.1103/PhysRevD.66.063514</p> <p>DESI Collaboration. 2025, arXiv e-prints.
 https://arxiv.org/abs/2507.XXXXX</p> <p>Ghirardi, G. C., Rimini, A., & Weber, T. 1986, Phys. Rev. D, 34, 470, doi: 10.1103/PhysRevD.34.470</p> <p>Pearle, P. 1989, Phys. Rev. A, 39, 1463,
 doi: 10.1103/PhysRevA.39.1463</p> <p>Šola, J., Karimkhani, E., & Kazemi, A. 2015, JCAP, 2015, 029, doi: 10.1088/1475-7516/2015/12/029</p> <p>Starobinsky, A. A. 2007, JETP Lett., 86, 157,
 doi: 10.1134/S002136400715004X</p> |
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