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ShOpt.jl | A Julia Library for Empirical Point Spread Function Characterization of JWST NIRCam Data

Edward Berman [1¶] and **Jacqueline McCleary** [1¶]

1 Northeastern University, USA ¶ Corresponding author

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Summary

Introduction

When astronomers capture images of the night sky, several factors – ranging from diffraction and optical aberrations to atmospheric turbulence and telescope jitter – affect the incoming light. The resulting distortions are summarized in the image's point spread function (PSF), a mathematical model that describes the response of an optical system to an idealized point of light. The PSF can obscure or even mimic the astronomical signal of interest, making its accurate characterization essential. By effectively modeling the PSF, we can predict image distortions at any location and proceed to deconvolve the PSF, ultimately reconstructing distortion-free images.

The PSF characterization methods used by astronomers fall into two main classes: forward-modeling approaches, which use physical optics propagation based on models of the optics, and empirical approaches, which use stars as fixed points to model and interpolate the PSF across the rest of the image. (Stars are essentially point sources before their light passes through the atmosphere and telescope, so the shape and size of their surface brightness profiles define the PSF at that location.) Empirical PSF characterization proceeds by first cataloging the observed stars, separating the catalog into validation and training samples, and interpolating the training stars across the field of view of the camera. After training, the PSF model can be validated by comparing the reserved stars to the PSF model's prediction.

Shear Optimization with `ShOpt.jl` introduces modern techniques, tailored to James Webb Space Telescope (JWST) imaging, for empirical PSF characterization across the field of view. ShOpt has two modes of operation: approximating stars with analytic profiles, and a more realistic pixel-level representation. Both modes take as input a catalog with image cutouts – or "vignettes" – of the stars targeted for analysis.

Statement of need

Empirical PSF characterization tools like PSFEx (Bertin, 2011) and PIFF (Jarvis et al., 2020) are widely popular in astrophysics. However, the quality of PIFF and PSFEx models tends to be quite sensitive to the parameter values used to run the software, with optimization sometimes relying on brute-force guess-and-check runs. PIFF is also notably inefficient for large, well-sampled images, taking hours in the worst cases. The James Webb Space Telescope's (JWST) Near Infrared Camera (NIRCam) offers vast scientific opportunities (e.g., (Casey, Kartaltepe, et al., 2023)); at the same time, this unprecendented data brings new challenges for PSF modeling:

(1) Analytic functions like Gaussians are incomplete descriptions of the NIRCam PSF, as evident from Figure 1. This calls for well-thought-out, non-parametric modeling and

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diagnostic tools that can capture the full dynamic range of the NIRCam PSF. ShOpt provides these models and diagnostics out of the box.

(2) The NIRCam detectors have pixel scales of 0.03 (short wavelength channel) and 0.06 (long wavelength channel) arcseconds per pixel (Beichman et al., 2012a; Rieke et al., 2003, 2005). At these pixel scales, star vignettes need to be at least 131 by 131 pixels across to fully capture the wings of the PSFs (4-5 arcseconds). These vignette sizes are 3-5 times larger than the ones used in surveys such as DES (Jarvis et al., 2020) and SuperBIT (McCleary et al., 2023) and force us to evaluate how well existing PSF fitters scale to this size. ShOpt has been designed for computational efficiency and aims to meet the requirements of detectors like NIRCam.

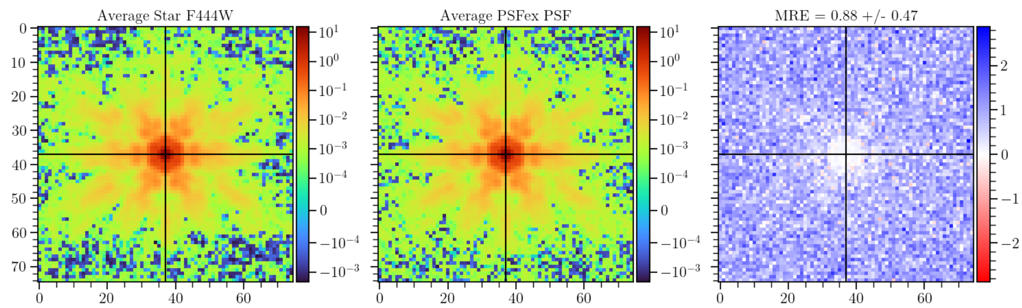

Figure 1: The plot on the left shows the average cutout of all stars in a supplied catalog. The plot in the middle shows the average point spread function model for each star. The plot on the right shows the average normalized error between the observed star cutouts and the point spread function model.

ShOpt bridges the speed of PSFex with the features of PIFF using fewer configurable hyperparameters. ShOpt employs a myriad of techniques to optimize the speed of the program. First and foremost, ShOpt is equipped with support for multithreading. Polynomial interpolation is used for handling PSF variations across the field of view. The polynomials given to each basis element of the PSF are independent of one another and therefore can be distributed to different CPU threads to be run in parralel. ShOpt also introduces new methods for fitting both analytic profiles and pixel based profiles. If an analytic profile is used to model the PSF, then there are 3 basis elements parameterizing the model. We use the Limited Memory Broyden–Fletcher–Goldfarb–Shanno algorithm (LBFGS) to find these paramters. This is faster but more memory intensive than Conjugate Gradient, the algorithm used in PIFF. Moreover, the 3 basis elements are constrained to the manifold $B_2(r) \times \mathbb{R}_+$. We constructed a function that maps any point in $\mathbb{R}^3$ into $B_2(r) \times \mathbb{R}_+$. The LBFGS algorithm uses successive iterations to converge to a solution for the 3 basis elements, and so we use this function to ensure that our update steps in LBFGS do not leave the constraint. For the pixel basis, both PIFF and PSFex approximate the PSF by minimizing the reduced χ^2 between a grid of pixels and a star vignette. PCA can quickly achieve the same purpose of approximating the input vignette without overfitting to background noise. We also provide autoencoder mode, which uses deep learning to reconstruct the image. The weights and biasses are not reset between stars, so the knowledge of how to reconstruct one star is transfered to the next. This in turn leads to less training iterations. Finally, ShOpt is written in Julia. Julia uses a just in time compiler which makes it faster than intepreted languages such as Python and has shown to be a good choice for performance critical code (Stanitzki & Strube, 2021).

State of the Field

The JWST captures images at high resolution and at wavelengths of light that have been previously unexplored (Gardner et al., 2006). With these images we are seeing farther into the early universe than we ever have before. The difficulties of producing good PSF models for the JWST are emblamatic of a larger problem: Our data sets are getting bigger and existing software was not built to scale. That is to say, the advancements in software are falling behind the advances in instrumentation. Not only does `ShOpt` produces PSF models for JWST NIRCam images, it also sets the precdent for designing software that scales.

There are several existing empirical PSF fitters, in addition to a forward model of the JWST PSFs developed by STScI (Jarvis et al., 2020 ; Bertin, 2011; Perrin et al., 2014 ; Perrin et al., 2012). We describe them here and draw attention to their strengths and weaknesses to further motivate the development of `ShOpt.jl`. As described in the statement of need, `PSFex` was one of the first precise and general purpose tools used for empirical PSF fitting. However, the Dark Energy Survey collaboration reported small but noticeable discrepancies between the sizes of `PSFex` models and the sizes of observed stars. They also reported ripple-like patterns in the spatial variation of star-PSF residuals across the field of view (Jarvis et al., 2020), which they attributed to the astrometric distortion solutions for the Dark Energy Camera.

These findings motivated the Dark Energy Survey's development of `PIFF` (Point Spread Functions in the Full Field of View). `PIFF` works in sky coordinates on the focal plane, as opposed to image pixel coordinates used in `PSFex`, which minimized the ripple patterns in the star-PSF residuals and the PSF model size bias. (Based on the DES findings, `ShOpt` also works directly in sky coordinates.) `PIFF` is written in Python, a language with a large infrastructure for astronomical data analysis, for example Astropy (Astropy Collaboration et al., 2022) and Galsim (Rowe et al., 2015). The choice of language makes `PIFF` software more accessible to programmers in the astrophysics community than `PSFex`, which was first written in `C` twenty-five years ago and much less approachable for a community of open source developers. One of the motivations of `ShOpt` was to write astrophysics specific software in `Julia`, because `Julia` provides a nice balance of readability and speed with its high-level functional paradigm and just-in-time compiler (Stanitzki & Strube, 2021). Julia ranks behind Python, IDL, Matlab, and Fortran in full-text mentions in astronomical literature (Collaboration et al., 2022). We are optimistic that `ShOpt` will demonstrate that Julia is an appealing choice for programming in astronomy despite its low adoption to date.

While WebbPSF provides highly precise forward models of the JWST PSF, these models are defined for single-epoch exposures (Perrin et al., 2014, 2012). Much of the NIRCam science is accomplished with image mosaics – essentially, the combination of single exposure detector images into a larger, deeper image. The rotation of the camera between exposures, the astrometric transformations and resampling of images before their combination into a mosaic, and the mosaic's large area all make the application of WebbPSF models to mosaics a non-trivial procedure. Additionally, some recent work being done to generate hybrid PSF models, which add an empirical correction to forward-model PSFs, for single-epoch exposures (Lin et al., 2023). At the time of writing, there is no widely available software to do this.

The COMOS-Web survey is the largest JWST extragalactic survey according to area and prime time allocation (Casey, Kartaltepe, et al., 2023), and takes up $0.54\ deg^2$ (Beichman et al., 2012b; Rieke et al., 2023). This is a large enough portion of the sky that we should prepare to see significant PSF variation across the field of view because of astrometric distortions. Thus, COSMOS-Web data provides `ShOpt` with an oppurtunity to validate PIFF's correction for handling PSF variations and test how impactful (or not impactful) PSFex's size bias is. Additionally, the COMOS-Web survey hopes to use the JWST to detect of thousands of galaxies in the Epoch of Reionization $(6 \sim z \sim 11)$ to create one of the highest resolution mass maps of the early universe ever (Casey, Kartaltepe, et al., 2023). JWST data has also been used to pick out active galactic nuclei from host galaxies (Zhuang & Shen, 2023) and indentify 15 candidate galxies whose luminosities push the limits of our ΛCDM galaxy formation models (Casey, Akins, et al., 2023). These science cases neccesitate good PSF modeling and underscore the importance of `ShOpt`.

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