





**Figure 2.** Plot showing all converted OC and their respective ages

This makes OC, with a known age, ideal for the data used in this modelling process if we know the period and mass of the stars contained. The vast majority of OC data for this project was gathered from (PLANET BOOK et al. YEAR) as well as some new additions to already existing OC M37 (Chang et al. 2017) and some unpublished data generously provided by Jason Curtis. Although all these catalogs contained values for period, mass was not often provided. Instead, mass was effectively given by the photometry values of each star, with each catalog providing different bands for their stars. This made it difficult to convert all the stars to mass as different conversion are needed for different bands and different mass ranges, with some conversions not stretching as low as the lowest mass.

A conversion was possible, however, it did not use conventional functions to map photometry to mass. Instead I used the MESA Isochrone and Stellar Track (MIST) tables (MISTTHING et al. 2004), which are simulations that provide information on the properties of stars, for a range of masses, evolved through time. The converted OC can be seen in Figure 2

These tables start with discrete mass steps (e.g 0.1, 0.15...1.35, 1.40  $M_{\odot}$  etc). These masses, evolving at different rates, are very likely to have a degeneracy in their photometry, meaning they may cross each others "photometry path", and as a consequence the conversion was not as simple as choosing the two closest photometries and interpolating. Instead I had to restrict the available pool of photometries, based on the closest ages to that star, and from that pool choose the closest photometry to interpolate between. The discontinuity of the tracks, due to discrete time steps, means there will be an inherent error in choosing the pool of ages, this has not currently been addressed, however may be implemented into error

propagation in a future model if a Bayesian network approach is used.

### 2.1.1 M37 shift?

Perhaps due to metallicity

## 2.2 Unsupervised Clustering?

Initial attempt to separate the fast and slow rotators however problematic due to it not being a "two group" problem. Transitional stars need to be considered, otherwise subjecting the transition to a dirac delta.

## 2.3 Polynomial Ridge Regression

slow rotators fit using a polynomial fit of order 4. Sigmoid function overlapped and optimised to change poly term on and off at switch point. FIGURE OF CURRENT FIT

Through the coefficients, it was thought that the progression of this transition from one population to the other could be understood, however, so far this time evolution has not been assessed. though polynomial fits have been generated from the coefficients, the problem of degeneracy still presents itself with this method. A potential model to address this could have an error associated with it, where this line is the mean. This model would have confidence intervals as a function of mass and age. The closer to the "overlap area" the prediction was, the larger the uncertainty becomes, then subsequently dropping after the lone fast rotators have been reached.

This spread/overlap of the data can be explained in the next section.

## 2.4 Initial Period and other parameters

Expand of the effect initial period and perhaps metallicity. Other parameters could allow for deeper understanding of the "overlap" sections of the open clusters. FIGURE OF HOW INITIAL PERIOD CAUSES MULTIPLE LINES TO OVERLAP AND MAKE THE TRANSITION "BLURRY"

The difficulty in these predictions of period, for a given age and mass, stem from the lack of information on the initial period distribution. In a hypothetical OC, whose stars are born of a single initial rotation period, we think the distribution would look like figure XXXXX.

INSERT FIGURE HERE

However, an OC produces a range of masses and initial periods, only the former of which can be measured. Therefore the best estimates for initial period of the system can be assumed to be the youngest clusters distribution that is no longer under the influence of disc effects, such as H Persei (NGC 869).

## 3 FUTURE WORK?

Since there is overlap, a polynomial fit will give poor predictive results and without initial period is not purely deterministic to the degree we want. To remedy this a probabilistic model will be built that can be used to sample and generate a synthetic population of stars at a given age and a range of masses.

## **ACKNOWLEDGEMENTS**

The Acknowledgements section is not numbered. Here you can thank helpful colleagues, acknowledge funding agencies, telescopes and facilities used etc. Try to keep it short.

## **APPENDIX A: SOME EXTRA MATERIAL**

If you want to present additional material which would interrupt the flow of the main paper, it can be placed in an Appendix which appears after the list of references.

This paper has been typeset from a  $\text{\LaTeX}$  file prepared by the author.