

# **Realistic fire rendering**

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# Contents

<b>1</b>	<b>Introduction</b>	<b>3</b>
<b>2</b>	<b>Previous Work</b>	<b>6</b>
2.1	Simulation . . . . .	6
2.2	Rendering . . . . .	10
<b>3</b>	<b>Methodology</b>	<b>13</b>
3.1	Fundamental Concepts . . . . .	13
3.2	Radiative Transport Equation . . . . .	14
3.3	Scattering . . . . .	16
3.4	Black Body Radiation . . . . .	16
3.5	Absorption coefficients . . . . .	17
3.6	Refraction . . . . .	18
3.7	Visual Adaptation . . . . .	23
<b>4</b>	<b>Implementation details</b>	<b>24</b>
4.1	Application Overview . . . . .	24
4.2	Mental Ray® Rendering Approach . . . . .	25
4.3	Simplifications . . . . .	26
4.4	Shaders . . . . .	27
4.5	Miscellaneous . . . . .	31

<b>5</b>	<b>Results</b>	<b>33</b>
<b>6</b>	<b>Conclusions and Future work</b>	<b>34</b>

# Chapter 1

## Introduction

*“A case that can be made for fire being, next to the life process, the most complex of phenomena to understand.” - Hoyt Hottle*



(a) Paper fire [[FireImg1, 2015](#)].



(b) Bone fire [[FireImg2, 2015](#)].

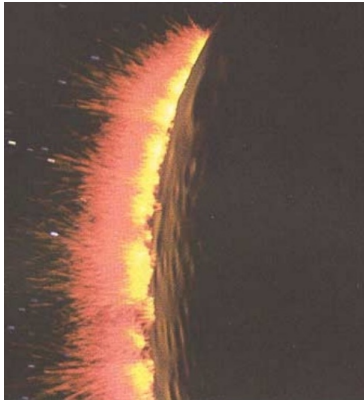
Figure 1-1: Examples of a real fire.

Computer simulations are widely used to model a diverse range of natural phenomena. Their popularity lies in their inexpensive nature, predictive capabilities given good mathematical models, parameters can be changed at will to produce new results, the equations governing the system can be modified to create scenarios that are not physically attainable, etc. Computer generated graphics are used to display the results of the simulations. Rendering the data on the screen involves transforming a scene into an image that can be displayed using some light transport and illumination model.

For centuries humans has been fascinated with fire due to its attractive presence and its dangerous nature, real fire examples are shown in Figure 1-1. Combustion phenomena are prevalent in daily life, candles, camp fires, explosions, car engines, cooking appliances, etc. Simulating and visualizing fire related processes has

many applications, for example it is widely used for visual effects in the film industry, simulated as part of the virtual environment in the computer games industry; or in the engineering community, where modelling engine combustion and fire safety evaluations are frequently demanded.

Computer generated examples in films include, a planet explosion in Star Trek II, as shown in Figure 1-2a, where a particle-based technique by [Reeves, 1983] was used; Shrek featured a dragon exhaling fire, as shown in Figure 1-2b, where parametric curves were used to drive the flames [Lamorlette and Foster, 2002]; or the more recent work of [Horvath and Geiger, 2009] based on 2D screen projections for the film Harry Potter and the Deathly Hallows, as shown in Figure 1-2c. In these and in many other applications, using real flames is an expensive and hazardous endeavour.



(a) Explosion in Star Trek II [Reeves, 1983].



(b) Dragon fire in Shrek [Lamorlette and Foster, 2002].



(c) Flames used in Harry Potter and the Deathly Hallows [Horvath and Geiger, 2009].

Figure 1-2: Examples of visual effects with fire in the film industry.

The computer graphics community has intensively researched the fluid behaviour of water and smoke. Fire can also be modelled as a fluid, however due to its multiphase flow, chemical reactions and radiative heat transport, the techniques used for water or smoke cannot be directly applied to flames. As a result of the aforementioned complexity and the interdisciplinary nature of the problem, fire simulation is still an open

problem in computer graphics.

A great deal of work done in the area has sacrificed complexity for interactiveness, therefore producing simplified models which hope to deceive the observer by exploiting the chaotic behaviour present in fire motion. Nevertheless, physically-based simulations incorporate the intrinsic processes that occur in a combustion scenario in order to be able to produce realistic results. In general the following effects are taken into consideration.

**Flame motion:** under normal gravity conditions, convection, external wind forces, front propagation and gas expansion drive the shape and position of the flame.

**Fuel erosion:** when the fuel reaches a certain temperature, it is vaporized into a gaseous state, which rises under the influence of buoyancy.

**Flame colour:** the chemical species present in the fuel and the byproducts of the combustions emit energy in various wavelengths due to black body radiation.

In this report a short review of fire simulation and rendering methods is presented in Chapter 2. In Chapter 3, a state-of-the-art physically-based model to render fire, proposed by [Pegoraro and Parker, 2006], is discussed extensively. Implementation details of a Mental Ray® shader for the the given model are outlined in Chapter 2, and results are discussed in Chapter 5. Lastly, conclusions and future research directions are considered in Chapter 6.

## Chapter 2

# Previous Work

In order to display a realistic fire scene in a computer generated world, two differentiated stages are needed. Firstly, the fire dynamics have to be collected, this can either be done through a data capture session or simulated using fluid-based solvers. Present techniques in the area are discussed below. Secondly, the previously gathered data will be visualized on the screen using some rendering method, a brief state-of-the-art review of fire rendering procedures is presented in Section 2.2. We refer the interested readers to the more detailed survey on the topic, which has been recently presented by [Huang et al., 2014].

Fire simulation has received significantly more attention from the computer graphics community than fire rendering. As the output data of the simulation methods is used by the rendering stage, an outline of the different simulation techniques is given below. Still, the main focus of this report lies in the rendering phase of the combustion simulation pipeline.

### 2.1 Simulation

The techniques presented in this section are divided adopting the same criteria that was used by [Huang et al., 2014]. Since the categories are roughly chronological, the evolution of the field is easier to follow.

**Particle-based** methods were the first approach to simulate the visual animation of fire, as shown in Figures 1-2a and 2-1. A number of particles are emitted from certain locations, each particle has a set of attributes such as shape, velocity, color or lifetime. The first model with particle systems was presented by [Reeves, 1983], the particles speed and colour were perturbed with a Gaussian noise at each time step, and the colour was subject to an additional linear perturbation on its lifetime. Two particle systems were used in a hierarchy, one would control fire spread and the other a single explosion effect. An extension was proposed by [Perry and Picard, 1994]; the authors modified the particle system such that each particle shape would be defined by a series of non-overlapping coplanar triangles. The transparency of each triangle increases towards the outer vertices, thus providing an improved visual effect.

**Noise-based** methods focus on synthesizing the high fluctuation present in fire procedurally, as shown in Figure 2-2. The objective is to approximate the turbulence present in fire with an appropriate statistical

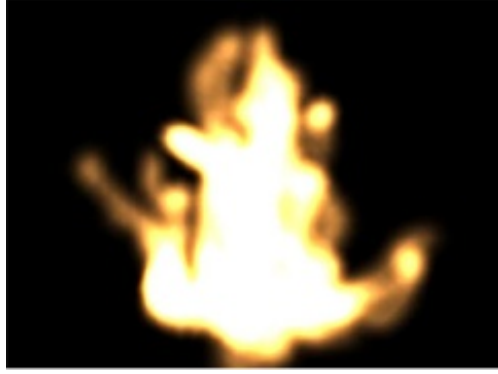
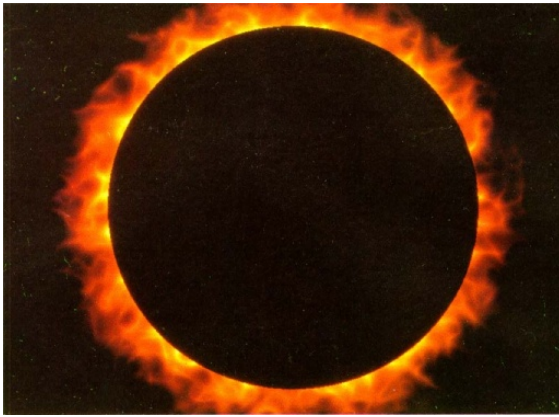
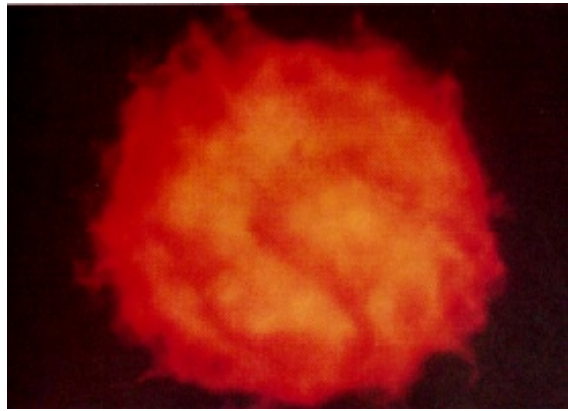


Figure 2-1: Particle-based fire [Perry and Picard, 1994].

model. Using a variation of Perlin noise, [Perlin, 1985] presented images of a corona of flames. However, the method is limited to 2D, where the color is a combination of non-linear arbitrary functions. This work was extended by [Perlin and Hoffert, 1989] to 3D, where they used volumetric rendering to achieve improved results.



(a) [Perlin, 1985].



(b) [Perlin and Hoffert, 1989].

Figure 2-2: Noise-based fire examples.

**Geometry skeleton** are driven by rendering primitives that represent the main features of fire, the high frequency details are built on top of the geometry structure, as shown in Figures 1-2b and 2-3. [Lee et al., 2001] presented a technique to animate fire on meshes, the flame front propagation is based on a geodesic flow field on the surface. Merging and separating of multiple fronts is achieved, however the animators have no control over the fire evolution. [Lamorlette and Foster, 2002] proposed a B-Spline curves interpolation method, details are added from a library of images which are then cast onto the profile of the flame. The authors' technique supports a range of flame behaviour, such as spreading, flickering, and flame separation and merging. In order to achieve real time simulations, [Bridault-Louchez et al., 2006] modelled the flame shape on NURBS surfaces on which a transparent 2D texture were mapped.

**Data driven** methods are use to simulate fire with data from real flames, as shown in Figure 2-4. The quality of the animation depends directly on the data, and by avoiding the simulation stage the computation





(a) [Lee et al., 2001].

(b) [Bridault-Louchez et al., 2006].

Figure 2-3: Geometry skeleton fire examples.

overhead is reduced significantly. However, there are severe limitations when producing new animations and interactions with other objects. [Rushmeier et al., 1995] captured a heptane pool fire with water-cooled nitrogen-purged probes. Using two images from orthographic camera views, [Hasinoff and Kutulakos, 2003] reconstructed 3D volumetric fire using tomographic techniques. [Ihrke and Magnor, 2004] extended the work by [Hasinoff and Kutulakos, 2003], the authors presented a method to reconstruct fire volumes using trilinear basis functions in scenes recorded with eight cameras. [Zhang et al., 2011] presented a method to synthesize new animation sequences based on previously simulated data. A flow graph whose nodes connect resembling simulation states is built, similarity between nodes is measured based on flow pathline distances. New frames with smooth transitions are created by traversing the graph through different paths. [Okabe et al., 2015] proposed a reconstruction technique using a set of sparse images, that can also be applied to generate novel animations. The authors approach is to define an optimization problem in terms of colour histograms and steerable coefficients, which they can use to retrieve an approximation of the appearance of the volume at new viewing angles.



(a) [Hasinoff and Kutulakos, 2003].

(b) [Ihrke and Magnor, 2004].

(c) [Zhang et al., 2011].

(d) [Okabe et al., 2015].

Figure 2-4: Data driven fire examples.

**Physically based** methods simulate the fire combustion processes, including flame propagation or the chemical reactions that convert fuel into gaseous products, as shown in Figures 1-2c and 2-5. Incompressible flow equations were used by [Stam and Fiume, 1995] to drive a fire simulation. Given initial fuel conditions, the fire spread is advected on a grid using an advection-diffusion type equation. Building on a semi-Lagrangian fluid solver of [Stam, 1999], a model which includes gaseous fuel and gaseous byproducts was proposed by [Nguyen et al., 2002]. In order to include the characteristics of the noise-based methods, [Hong et al., 2007] combined the previous model with a set of third-order equations from detonation shock dynamics presented by [Yao and Stewart, 1996]. As with the noise-based methods, this addition is visually attractive, yet it is not physically based. Capitalizing on the recent advances in GPUs parallel processing power, [Horvath and Geiger, 2009] proposed a fixed camera model. Particle properties are computed on a three-dimensional coarse grid, which are then projected into several view-dependant two-dimensional slices. The authors' model is based on the assumption that fine variations, which are perpendicular to the projection plane, are not individually visible and, they do not significantly affect the overall flow.

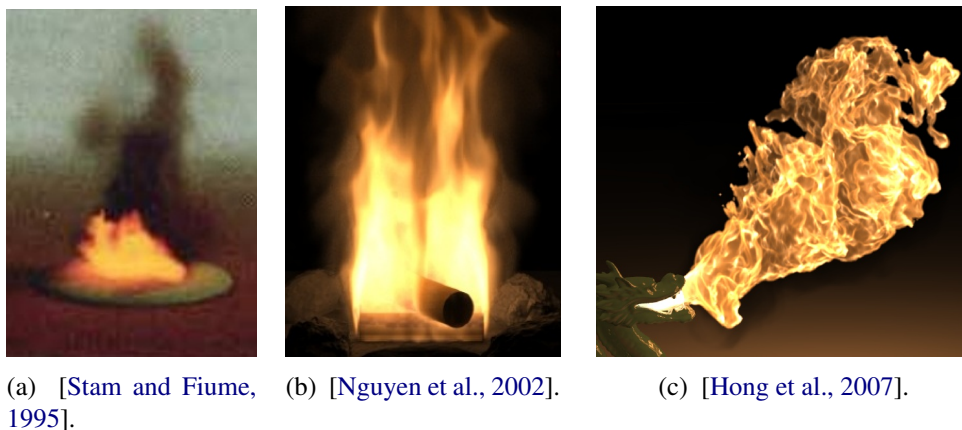


Figure 2-5: Physically-based fire examples.

**Other effects** directly related to fire have also been explored. [Feldman et al., 2003] presented a model to simulate suspended particles during explosions. An incompressible fluid model drives the motion of air and hot gases, and the suspended particles follow their movements. [Melek and Keyser, 2005] simulated the decomposition of an object as it undergoes combustion. Heat is transferred between the different parts of the object and the air; the decomposition is modelled as a moving boundary spreading in the distance field of the object, where level set methods are used to compute the moving interfaces. Credible sound is an important factor to increase the believability of a finished fire animation. [Chadwick and James, 2011] proposed a method to automatically generate plausible noise given for a given fire simulation. Low frequency sound is estimated using a physical model whose two given inputs are the flame front and heat release. A data driven sound synthesis approach, based on the work by [Wei and Levoy, 2000], is applied to generate the high frequency content.

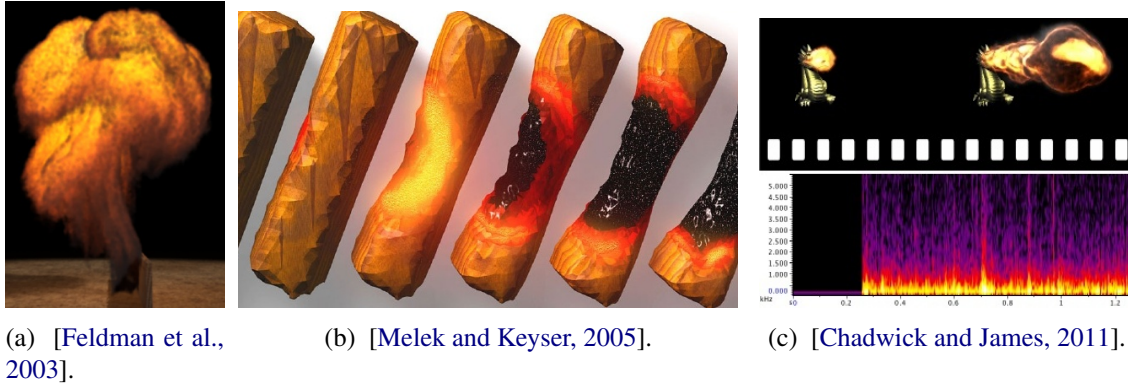


Figure 2-6: Other effects with fire.

## 2.2 Rendering

Rendering fire is more challenging than rendering other types of participating media, the main reason being that fire is a light-emitting source. The luminescence radiated by a flame is generated by black body radiation due to the high temperature of the particles during the combustion process. In order to realistically render fire, light absorption and scattering in the media, including air, has to be taken into consideration. A succinct review of the current methods used by the computer graphics community for fire rendering is presented below.

### 2.2.1 Raster-Based Methods

Raster-based techniques sacrifice quality in the interest of interactive frame rates, renderings are shown in Figure 2-7. Generally, some form of texture mapping is applied, only the surface of the flame is considered, and the illumination of the scene is approximated using parameters which can be easily computed such as the fire height.

[Reeves, 1983] applied a linear colour assignment to each particle in their simulation based on their lifetime. [Lee et al., 2001] applied the same technique to render fires on mesh surfaces, where a particle would begin with a light yellow colour, evolve to red and finish in black at the end of their lifetime. [Lamorlette and Foster, 2002] presented a technique where a flame picture is mapped onto the two-dimensional flame profile with a base colour. For each particle an intermediate emitting value is computed, and the final colour profile is super-sampled from an approximation of the cross-sectional area of the flame, as it would appear from the camera. [Westover, 1990] introduced texture splats as a new rendering primitive, [Wei et al., 2002] rendered fire using Westover's technique. Direct illumination is approximated by rendering the fire particles separately to create a light image, which is texture projected into the surrounding objects. [Zhao et al., 2003] also proposed a method with flame particle splats, which are added with splats from other objects and sorted based on their distance to the screen plane. The splats are projected in front-to-back order and alpha-blended to produce the end result. [Ihrke and Magnor, 2004] render fire realistically using a technique based on tomography reconstructions of real flames. [Bridault-Louchez et al., 2006] used a spectrophotometer to capture photometric distributions of candles. The intensities are stored on a texture and changes in illumina-

tion over time are introduced with an attenuation factor which is proportional to the size of the flame. [Zhang et al., 2011] proposed a method where fire particles and their attributes are first projected onto a set of slicing planes, which are orthogonal to the camera direction. The planes are then blended in back-to-front order, and a one-dimensional colour texture is used as a transfer function to convert flow attributes to colours and opacities. [Jamriška et al., 2015] recently proposed a texture synthesis method to generate fire animations. The method requires a hand made motion field and an alpha mask of the desired result, both are used to generate the new sequence using data from an existing video exemplar.

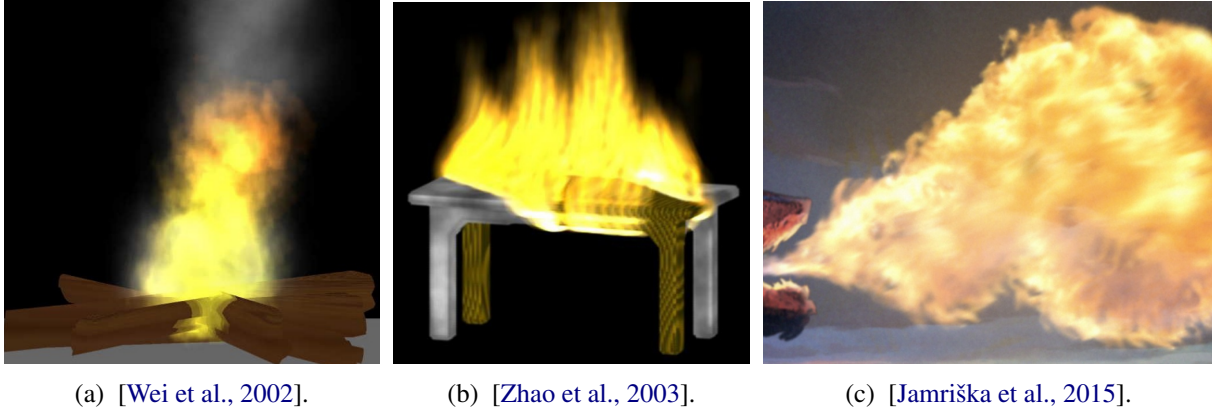


Figure 2-7: Raster-based techniques.

### 2.2.2 Ray-Tracing-Based Methods

Volume ray-tracing techniques offer astonishing results, however the associated computational costs are considerable. Rays are shot from the view plane and evaluated at small increments; the total radiance at the origin of the ray is computed by integrating the radiance at each step size. A further drawback of ray tracing techniques, in comparison to raster-based methods, is the lack of a standard ray-tracing pipeline.

[Rushmeier et al., 1995] presented a method to perform accurate ray casting on sparse measured data. The fire was modelled as a series of stacked cylindrical rings, where each ring has uniform properties. The total radiance at each point is integrated using a Monte Carlo method, summing up the measured irradiances at sample locations. [Nguyen et al., 2002] proposed a ray marching technique to solve the Radiative Transport Equation, see Chapter 3. The emitted light is computed using Planck’s formula of black body radiation, scattering in the media and visual adaptation to the fire spectra are modelled. [Feldman et al., 2003] also included black body radiation in their animation of fire with suspended particles, however the mapping to RGB was manually adjusted to match the images of real explosions. Direct illumination shadows were computed using deep shadow maps [Lokovic and Veach, 2000], while scattering and illumination by other objects in the scene used the technique proposed by [Jensen and Buhler, 2002]. An extension to [Nguyen et al., 2002] was presented by [Pegoraro and Parker, 2006], the authors’ model has physically-based absorption, emission and scattering properties. The spectroscopy characteristics of different fuels are achieved by modelling the electronic transitions between states in the molecules. Non-linear trajectories of light in the medium due to light refraction effects are included as well. In order to minimize the effects induced by the limitations of the RGB colour space, the visual adaptation process is presented as a post-processing effect. [Horvath and

[Geiger, 2009](#)] proposed a rendering method whose main objective was user-friendliness for artists. Using the fixed camera slices described in Section 2.1, the authors perform simple volume rendering to join them in a single image. Black body radiation is used for the light, the images are motion-blurred with a filter based on the velocities in the slices, and the heat distortion is added as post-processing filter defined by the user.

## Chapter 3

# Methodology

Not only explain in more detail every equation, what it means, comparison with other papers and ideally improvements or were it fails at least. Should I explain what is and how to do ray marching? What is the spectrum? How to integrate it to RGB coefficients? Add all the values for the constants here or in implementation details??

Given simulated or captured fire data, that could come from any of the methods discussed in Section 2.1, we want to generate a photo-realistic image from a camera view of the scene. From this point onwards we will assume that we have a volumetric data structure, which holds the relevant input data for the fire. An overview of the pipeline based on the work by [Pegoraro and Parker, 2006] is shown in Figure 3-1. The inputs are volumetric data of density and temperature, from those values a series of intermediate constants for each voxel for the current frame are computed, and with the previous data a ray marcher is used to generate the final image.

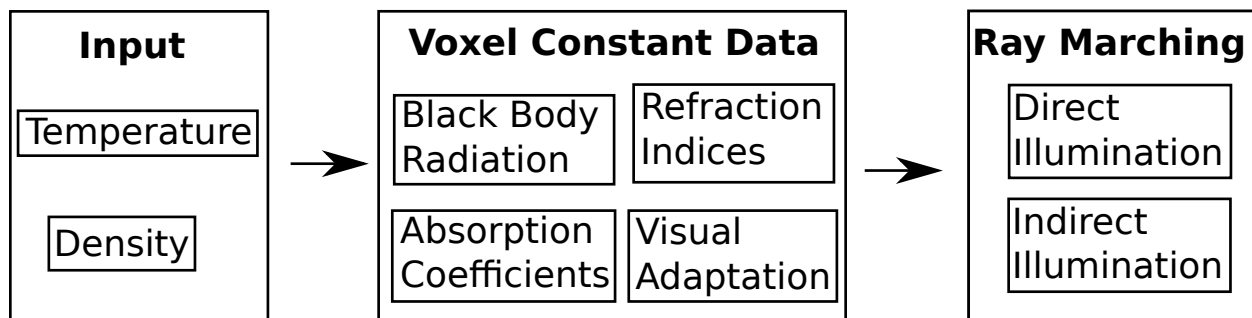


Figure 3-1: An overview of the fire rendering pipeline.

### 3.1 Fundamental Concepts

In this section an overview of the fundamentals concepts of volume rendering is presented, readers familiar with the area may proceed to Section 3.2. As the main objective is to give a intuitive introduction into the



area, the steps and concepts presented here will not be formally derived. A simple framework to generate images when light interacts with a participating media will be described.

The world is modelled as scene were there are a number of elements, the camera, fire volumes, lights and other objects. The **camera** is defined as a point and orientation in space, with a projection matrix. A **voxel** is defined a cube in space with some data associated to it. A **voxel dataset** is a collection on voxels in space, typically a cubic section of the scene is divided uniformly in voxels. There could also be other **lights** in the scene, several types are supported, point, area, directional and cone lights. **Other objects** may also be part of the scene, such as polygonal or nurbs meshes.

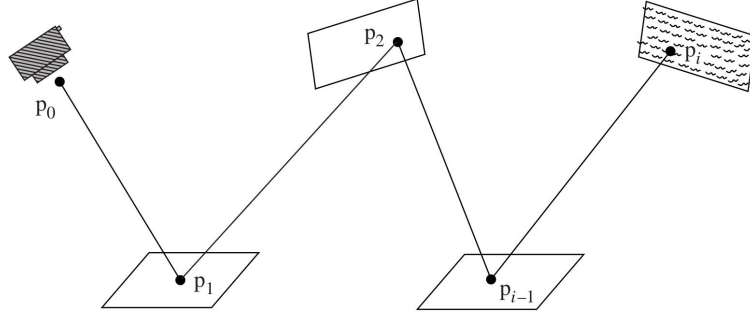


Figure 3-2: Computing the colour of  $p_1$  using path tracing [Pharr and Humphreys, 2010].

In order to generate an image from the camera view, path tracing is used to compute the colour of each pixel, as shown in Figure 3-2. For each pixel in the camera view, a **ray** is shot into the scene; whose direction is defined by the camera origin and the pixel position. When the ray hits a primitive in the scene, it will be reflected and the colour will depend on the contribution of further intersections along the path. The final point in the path will be directly connected to a light in the scene, in Figure 3-2,  $p_{i-1}$  is the final intersection and  $p_i$  is the light position.

## 3.2 Radiative Transport Equation

The Radiative Transport Equation (RTE) [Howell and Siegel, 2002] models the variation of spectral radiance  $(\omega \nabla) L(\lambda, \mathbf{x}, \omega)$  in the medium, where  $\lambda$  is a given wavelength in metres,  $\mathbf{x}$  is the point of interest in space, and  $\omega$  is a vector that points towards the viewing direction. Formally, the objective is to get a solution for the RTE for each visible point in the scene, practically, an approximate solution will be computed.

The integro-differential Radiative Transport Equation is defined as

$$\begin{aligned}
 (\omega \nabla) L(\lambda, \mathbf{x}, \omega) = & -\sigma_a(\lambda, \mathbf{x}) L(\lambda, \mathbf{x}, \omega) + \sigma_a(\lambda, \mathbf{x}) L_e(\lambda, \mathbf{x}, \omega) \\
 & -\sigma_s(\lambda, \mathbf{x}) L(\lambda, \mathbf{x}, \omega) + \sigma_s(\lambda, \mathbf{x}) L_i(\lambda, \mathbf{x}, \omega),
 \end{aligned} \tag{3.1}$$

where  $L_i$  is defined by

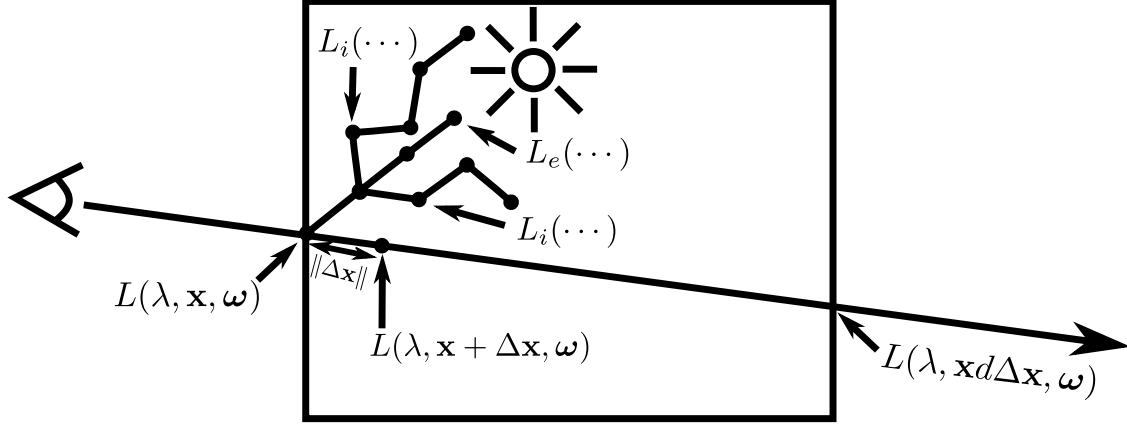


Figure 3-3: A visual representation of the RTE.

$$L_i(\lambda, \mathbf{x}, \omega) = \int_{4\pi} L(\lambda, \mathbf{x}, \omega_i) \Phi(\lambda, \omega, \omega_i) d\omega_i, \quad (3.2)$$

where  $\sigma_a$  is an absorption coefficient,  $\sigma_s$  is a scattering coefficient,  $L_e$  is the emitted spectral radiance at the point,  $L_i$  is the in-scattering radiance,  $\Phi$  is a scattering phase function and  $\omega_i$  is a scattering sampling direction. In order to get an analytical solution to the aforementioned equation, the properties of the medium are assumed to be homogeneous over a small segment  $\|\Delta\mathbf{x}\|$  in space,

$$L(\lambda, \mathbf{x} + \Delta\mathbf{x}, \omega) = e^{-\sigma_t(\lambda, \mathbf{x})\|\Delta\mathbf{x}\|} L(\lambda, \mathbf{x}, \omega) + (1 - e^{-\sigma_t(\lambda, \mathbf{x})\|\Delta\mathbf{x}\|}) \frac{\sigma_a(\lambda, \mathbf{x})L_e(\lambda, \mathbf{x}, \omega) + \sigma_s(\lambda, \mathbf{x})L_i(\lambda, \mathbf{x}, \omega)}{\sigma_t(\lambda, \mathbf{x})}, \quad (3.3)$$

where  $\sigma_t = \sigma_a + \sigma_s$  is the extinction coefficient, a visual depiction on this solution is shown in Figure 3-3. To be able to calculate the terms in Equations 3.2 and 3.3, we will explain in more detail the following factors:

- Light **scattering** in the media  $L_i(\lambda, \mathbf{x}, \omega)$  with a phase function  $\Phi(\lambda, \omega, \omega_i)$ .
- Emitted light  $L_e(\lambda, \mathbf{x}, \omega)$  due to **black body radiation**.
- Electromagnetic **absorption**,  $\sigma_a$ , of certain wavelengths.
- Due to spatially varying **refractive** indices, photons may follow non linear paths  $\Delta\mathbf{x}$ .
- **Visual adaptation** processes in the human eye affect the perceived final radiance  $L(\lambda, \mathbf{x}, \omega)$ .



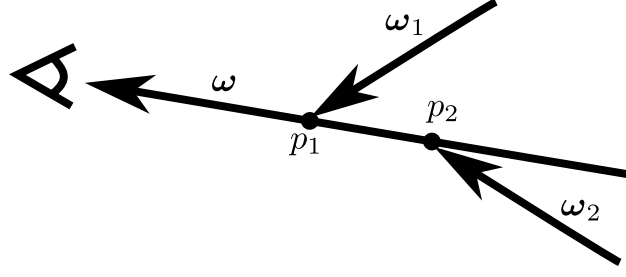


Figure 3-4: Photons scattering in the direction of interest.

### 3.3 Scattering

Some photons travelling along arbitrary directions  $\omega_i$ , might be scattered in the direction of interest  $\omega$ . A simple diagram of the process is shown in Figure 3-4, where two photons travelling in directions  $\omega_1$  and  $\omega_2$ , are scattered in the direction of interest  $\omega$ , at the points  $p_1$  and  $p_2$  respectively. The number of photons which are subjected to this process is controlled by a coefficient  $\sigma_s$ , the radiance carried by the photons is denoted by  $L_i(\lambda, \mathbf{x}, \omega)$  and follows a spherical distribution  $\Phi(\lambda, \omega, \omega_i)$ , which was first proposed by [Henye and Greenstein, 1941] to model diffuse radiation in astronomy,

$$\Phi(\lambda, \omega, \omega_i) = \frac{1 - g(\lambda)^2}{4\pi (1 + g(\lambda)^2 - 2g(\lambda)\omega\omega_i)^{\frac{3}{2}}}, \quad (3.4)$$

where  $g$  is an asymmetry coefficient which can be a function of the wavelength, although in most cases it is chosen to be constant. The value of  $g$  must be in the range  $(-1, 1)$ , where  $g < 0$  corresponds to backwards scattering,  $g = 0$  to isotropic scattering, and  $g > 0$  to forwards scattering. Monte Carlo techniques can be used to sample the light scattering directions  $\omega_i$ .

### 3.4 Black Body Radiation

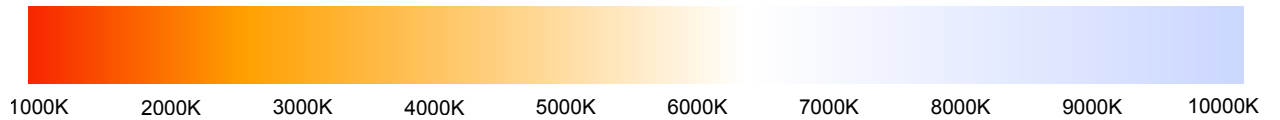


Figure 3-5: Black body emission colours normalised to ignore brightness [BlackBodyRadiation, 2015].

Emitted light from fire, which will illuminate the flame itself and other objects in the scene, is denoted as  $L_e(\lambda, \mathbf{x}, \omega)$ . Radiation from an idealized physical body that absorbs all incident electromagnetic radiation can be modelled with Planck's equation for black body radiation, as shown in Figure 3-5. This equation characterizes the electromagnetic radiation  $B(T, \lambda, \eta)$  emitted by a black body in thermal equilibrium at a given temperature  $T$ .

$$L_e(\lambda, \mathbf{x}, \omega) = B(T, \lambda, \eta) = \frac{2hc^2}{\lambda^5(e^{\frac{hc}{\lambda kT}} - 1)}, \quad (3.5)$$

where  $h = 6.62606957 \times 10^{-34} \text{ J/s}$  is Planck's constant,  $k = 1.3806488 \times 10^{-23} \text{ J/K}$  is Boltzmann constant,  $\eta$  is the refraction index of the medium and  $c = c_0/\eta$  is the speed of light in the current medium, where  $c_0 = 299792458 \text{ m/s}$  is the speed of light in a vacuum. The refraction index  $\eta$  varies across the medium, the procedure to compute it will be described in Section 3.6.

### 3.5 Absorption coefficients

As fire is not an idealized black body radiator, certain wavelengths are not present in the emitted radiance. This behaviour can be modelled using absorption coefficients for each chemical that would mask the undesired wavelengths, such that

$$\sigma_a L_e(\lambda, \mathbf{x}, \omega) = \sum_{i=0}^N \sigma_{ai} B_i(T, \lambda, \eta), \quad (3.6)$$

where  $N$  is the number of different chemical species in the mix,  $\sigma_{ai}$  and  $B_i(T, \lambda, \eta)$  are the absorption coefficient and light emission of the  $i^{th}$  chemical. In the interest of simplicity, the absorption and radiation effects for **soot** will be treated separately from those caused by **other chemical species**.

#### 3.5.1 Soot Absorption

Light generated by incandescent soot accounts for most of the total radiance emitted by a flame. The spectral absorption coefficient of soot is defined as

$$\sigma_a(\lambda, \mathbf{x}) = \frac{48N(\mathbf{x})\pi R^3 nm}{\lambda^{\alpha(\lambda)}((n^2 - m^2 + 2)^2 + 4n^2 m^2)}, \quad (3.7)$$

where  $N(\mathbf{x})$  is the number density, density per unit volume,  $R$  is the radius of a soot particle,  $n$ ,  $m$  and  $\alpha(\lambda) = 1.39$  are optical constants for different types of soot. In Table 3.1, values for the optical constants  $n$  and  $m$  are provided for several materials, the data was obtained from [Dalzell and Sarofim, 1969]. The radius of soot particles was determined in the range  $R \in \{50 \text{ \AA} \dots 800 \text{ \AA}\}$  by [Dalzell and Sarofim, 1969]. Since our data is defined with soot densities, we have chosen the radius to be the mean value  $R = 425 \times 10^{-10}$  metres.

Table 3.1: Absorption constants for propane and acetylene, wavelengths are in nanometres.

		Wavelengths			
		435.8	450	550	650
<b>Propane</b>	n	1.57	1.56	1.57	1.56
	nm	0.46	0.5	0.53	0.52
<b>Acetylene</b>	n	1.56	1.56	1.56	1.57
	nm	0.46	0.48	0.46	0.44

### 3.5.2 Absorption From Other Chemical Species

Flames of more uncommon colours are generally produced when using salts of different chemicals. The spectral line emitted by an particular atom depends of the electronic transitions between different different energy levels. The absorption coefficients associated for a given spectral frequency and a known atomic element can be computed as

$$\sigma_a(\lambda, \mathbf{x}) = \frac{\phi(\lambda)N_2A_{21}\lambda^4(e^{\frac{hc}{\lambda kT}} - 1)}{8\pi c} \quad (3.8)$$

where  $\phi(\lambda)$  is the normalized spectral line,  $N_2$  is the number density of elements,  $A_{21}$  is an Einstein coefficient measuring the transition probabilities of spontaneous emission.

## 3.6 Refraction

Heat emanating from the flame is transferred to the air which surrounds it. Since the refractive index of a medium depends on its temperature, light transiting the space near the flame will follow non linear paths due to varying refractive indices. Fresnel equations describe in detail both reflection and refraction effects, as shown in Figure 3-6, when a wave moves between media with different refractive indices. Schlick equations [Schlick, 1994] are widely used by the computer graphics community to approximate the Fresnel coefficients.

The fraction of light which is reflected is assumed to be negligible for combustion phenomena; thus, the refraction angles for the rays can be easily computed using Snell's law

$$\frac{\sin \theta_1}{\sin \theta_2} = \frac{\eta_2}{\eta_1}, \quad (3.9)$$

where  $\theta_1$  is the incident angle,  $\theta_2$  is the refracted angle,  $\eta_1$  is the index of refraction of the media which the ray is coming from and  $\eta_2$  is the index of refraction of the media which the ray is going to. The values for all the constants in this section are defined in Table 3.2.

[Ciddor, 1996] proposed a method to compute the refractive indices of air

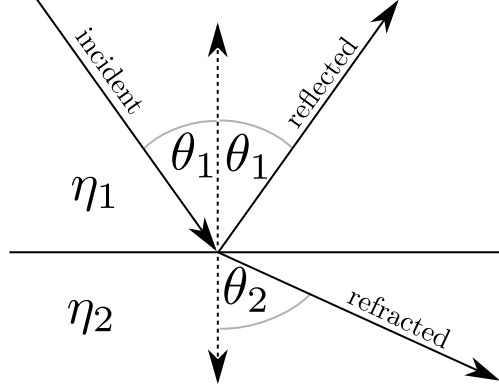


Figure 3-6: Ray being refracted and reflected in the boundary between two media.

$$\eta_{air} = 1 + \frac{\rho_a \eta_{axs}}{\rho_{axs}} + \frac{\rho_w \eta_{ws}}{\rho_{ws}}, \quad (3.10)$$

where  $\rho_{axs}$  is the density of dry air,  $\rho_{ws}$  is the density of pure water vapour,  $\rho_a$  and  $\rho_w$  are the equivalent quantities for dry air,  $\eta_{axs}$  is the refractive index for  $CO_2$  and  $\eta_{ws}$  is the refractive index for water vapour. The refractive index for  $CO_2$  is estimated as

$$\eta_{axs} = \eta_{as} \left( 1 + 0.534 \times 10^{-6} (x_c - 450) \right), \quad (3.11)$$

$$\eta_{as} = 10^{-8} \left( \frac{k_1}{k_0 - \lambda} + \frac{k_3}{k_2 - \lambda} \right), \quad (3.12)$$

where  $x_c \in \{300 \dots 450\}$  is the proportion in parts per million of  $CO_2$ ,  $k_0$ ,  $k_1$ ,  $k_2$  and  $k_3$  are constants, and  $\lambda$  is the wave number (reciprocal of the vacuum wavelength). The refractive index for water vapour is computed as

$$\eta_{ws} = 1.022 \times 10^{-8} (w_0 + w_1 \lambda + w_1 \lambda^2 + w_3 \lambda^3), \quad (3.13)$$

where  $w_0$ ,  $w_1$ ,  $w_2$  and  $w_3$  are constants.

The  $\rho$  densities are computed as follows

$$\rho = \frac{p m_a}{z r T} \left( 1 - x_w \left( 1 - \frac{m_w}{m_a} \right) \right), \quad (3.14)$$

where  $p$  is the pressure in pascals,  $r$  and  $m_w$  are constants,

$$m_a = 0.0289635 + 12.011 \times 10^{-8}(x_c - 400)), \quad (3.15)$$

$$z = 1 - \frac{P}{T} \left( a_0 + a_1 t + a_2 t^2 + (b_0 + b_1 t) x_w + (c_0 + c_1 t) x_w^2 \right) + \left( \frac{P}{T} \right)^2 (d + e x_w^2), \quad (3.16)$$

where  $a_0, a_1, a_2, b_0, b_1, c_0, c_1, d$  and  $e$  are constants.

$$T = t + 273.15, \quad (3.17)$$

$$x_w = \frac{f h v}{p}, \quad (3.18)$$

$$f = \alpha + \beta p + \gamma t^2, \quad (3.19)$$

where  $\alpha, \beta$  and  $\gamma$  are constants, and  $h \in \{0 \dots 1\}$  is the air relative fractional humidity, 0 for dry air and given by the user for moist air.

The pressure  $p$  can also be computed using the ideal gas law

$$pV = \frac{NRT}{V} = nT, \quad (3.20)$$

where  $N$  is the number of molecules,  $V$  is the total volume of the gas,  $T$  is the temperature of the gas,  $n$  is the number density,  $R = kN_a$ , where  $k$  is the Boltzmann constant and  $N_a$  is Avogadro constant.

Depending on the situation, there are two acceptable methods to compute the vapour pressure  $v$ , both are described below.

### 3.6.1 Davis' Saturation Vapour Pressure

In [Ciddor, 1996] paper, the author used the method proposed by [Davis, 1992] to compute the saturation of vapour pressure

$$v = e^{AT^2 + BT + C + D/T}, \quad (3.21)$$

where  $A, B, C$  and  $D$  are constants. This equation is relatively easy to compute, however incorrect results will be obtained for temperatures below 0 °C. As we are concerned with fire rendering, which entails high temperatures, this method was chosen for its simplicity.

### 3.6.2 Huang's Saturation Vapour Pressure

The International Association for the Properties of Water and Steam (IAPWS) adopted an alternative technique to [Davis, 1992], which was proposed by [Huang, 1998]. The more recent method addresses the

drawbacks of the previous technique, giving reasonable results for temperatures below 0 °C. Nevertheless, the generality comes with increased complexity in the equations

$$v = 10^6 \left( \frac{2C}{X} \right)^4, \quad (3.22)$$

$$X = -B + (B^2 - 4AC)^{\frac{1}{2}}, \quad (3.23)$$

$$\begin{bmatrix} A \\ B \\ C \end{bmatrix} = \begin{bmatrix} 1 & K_1 & K_2 \\ K_3 & K_4 & K_5 \\ K_6 & K_7 & K_8 \end{bmatrix} \begin{bmatrix} \Omega^2 \\ \Omega \\ 1 \end{bmatrix}, \quad (3.24)$$

$$\Omega = T + \frac{K_9}{T - K_{10}}, \quad (3.25)$$

where  $K_1, K_2, K_3, K_4, K_5, K_6, K_7, K_8, K_9$ , and  $K_{10}$  are constants.

### 3.6.3 Ciddor's Method Summary

The inputs of Ciddor's technique are a wavelength  $\lambda_0$ , a temperature  $t$ , a pressure  $p$ , a relative humidity value  $h$  and a  $CO_2$  concentration  $x_c$ . The values for all the constants needed for the method are defined in Table 3.2. A refraction index is computed following this method for each voxel in the volumetric data structure. The steps to compute the index of refraction are listed below

1. Precompute  $z_a = 0.9995922115$ , from Equation 3.16 with  $t = 20$  °C,  $p = 101.325$  Pa and  $x_w = 0$ .
2. Find  $\lambda = 1/\lambda_0^2$ , and  $T$  with Equation 3.17.
3. Find  $v$  with the method described in Section 3.6.1 or in Section 3.6.2.
4. Find  $x_w$  using Equations 3.19 and 3.18.
5. Find  $\eta_{as}$  using Equation 3.12.
6. Find  $\eta_{ws}$  using Equation 3.13.
7. Find  $m_a$  using Equation 3.15.
8. Find  $\eta_{axs}$  using Equation 3.11.
9. Find  $z_m$  using Equation 3.16.
10. Find  $\rho_{axs} = (p_0 m_a)/(z_a r T_0)$ , where  $p_0 = 101325$ , and  $T_0 = 288.15$ , using Equation 3.14.
11. Find  $\rho_w = (x_w p m_w)/(z_m r T)$  using Equation 3.14.
12. Find  $\rho_a = ((1 - x_w) p m_a)/(z_m r T)$  using Equation 3.14.
13. Find the index of refraction  $\eta_{air}$  using Equation 3.10.

Table 3.2: Values for the constants in the equations used to compute the index of refraction.

Name	Value	Name	Value
$\rho_{ws}$	0.00985938	$a_0$	$1.58123 \times 10^{-6} KPa^{-1}$
$x_c$	450 ppm	$a_1$	$-2.9331 \times 10^{-8} Pa^{-1}$
$k_0$	$238.0185 \mu m^{-2}$	$a_2$	$1.1043 \times 10^{-10} Pa^{-1}$
$k_1$	$5792105 \mu m^{-2}$	$b_0$	$5.707 \times 10^{-6} KPa^{-1}$
$k_2$	$57.362 \mu m^{-2}$	$b_1$	$-2.051 \times 10^{-8} Pa^{-1}$
$k_3$	$167917 \mu m^{-2}$	$c_0$	$1.9898 \times 10^{-4} KPa^{-1}$
$w_0$	$295.235 \mu m^{-2}$	$c_1$	$-2.376 \times 10^{-6} KPa^{-1}$
$w_1$	$2.6422 \mu m^{-2}$	$d$	$1.83 \times 10^{-11} K^2 Pa^{-2}$
$w_2$	$-0.03238 \mu m^{-2}$	$e$	$-0.765 \times 10^{-8} K^2 Pa^{-2}$
$w_3$	$0.004028 \mu m^{-2}$	$\alpha$	1.00062
$p$	1 atm = 101325 Pa	$\beta$	$3.14 \times 10^{-8} Pa^{-1}$
$r$	$8.31451 Jmol^{-1} K^{-1}$	$\gamma$	$5.6 \times 10^{-7} ^\circ C^{-2}$
$m_w$	0.018015 kg/mol	$k$	$1.3806488 \times 10^{-23} J/K$
$N_a$	$6.02214129 \times 10^{23} mol^{-1}$	$K_4$	$1.20208247025 \times 10^4$
$A$	$1.2378847 \times 10^{-5} K^{-2}$	$K_5$	$-3.23255503223 \times 10^6$
$B$	$-1.9121316 \times 10^{-2} K^{-1}$	$K_6$	$1.49151086135 \times 10$
$C$	33.93711047	$K_7$	$-4.82326573616 \times 10^3$
$D$	$-6.3431645 \times 10^3 K$	$K_8$	$4.05113405421 \times 10^5$
$K_1$	$1.16705214528 \times 10^3$	$K_9$	$-2.38555575678 \times 10^{-1}$
$K_2$	$-7.24213167032 \times 10^5$	$K_{10}$	$6.50175348448 \times 10^2$
$K_3$	$-1.70738469401 \times 10$		

### 3.7 Visual Adaptation

The human eye presents a non-linear response to incident radiance  $L$ , as shown in Figure 3-7. If a flame is the dominant light source in a scene, it will appear with yellow-white colours, while the same flame in an environment with comparable or larger lights might appear orange-yellow.



Figure 3-7: Fire variations in colour due to eye adaptation mechanisms [Pegoraro and Parker, 2006].

A simple equation which can model this phenomena was proposed by [Naka and Rushton, 1966]

$$R(L, \tau) = \frac{L}{L + \tau}, \quad (3.26)$$

where  $\tau$  is a non-linear adaptation state. This state is determined by the visual system to maximize the perception of features for a given scene. Several approaches have been proposed to determine which features are used and how to compute the optimal state. We refer the readers to [Fairchild, 2005] for an in depth review of visual adaptation methods. The Von Kries model [Fairchild, 2005], the new values for the LMS cone responses for a given stimulus are simply multiplied by the inverse of the maximum response in the scene. The Retinex model [Fairchild, 2005], extends Von Kries' work by normalizing the the ratio of the stimulus with the average response in a retinex space. The retinex space is defined as a three dimensional space with the spectral responses of the cones. [Irawan et al., 2005] presented neural network-based model which includes time varying adaptations, and it also considers variations due to the observer characteristics, such as their age or visual impairment.



## Chapter 4

# Implementation details

In this chapter we will explain how to use the shaders which have been implemented and the minor differences between the theoretical pipeline explained in Chapter 3 and the actual software implementation. An overview of the pipeline is shown in Figure 4-1, note that modules in grey have not been implemented.

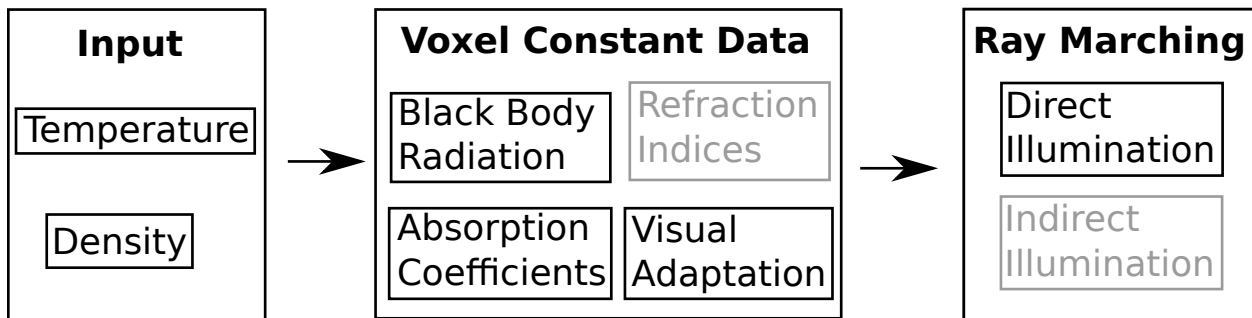


Figure 4-1: An overview of the implemented fire rendering pipeline.

### 4.1 Application Overview

The model chosen for our data is that of a three-dimensional voxel dataset, a cube in space is uniformly divided and a data value is stored for each voxel. This data is either a soot density estimate for the *fire volume shader* or a temperature estimate for the *fire light shader*. Three file formats are supported, dense and sparse float voxel data is ASCII, and sparse binary RGBA values, as shown in Figure 4-2. In the ASCII dense format, the file first line are three integers, separated by spaces, declaring the width, height and depth of the data, followed by  $w \cdot h \cdot d$  lines with a single floating point values. The ASCII sparse format, has two extra parameters in the first line,  $c$  is the number of data points in the file and  $b$  is the default value for any voxel that is not specified in the file. Each data line has the  $x, y, z$  coordinates and the data values are separated by spaces. The binary sparse format starts with a 4 byte integer declaring how many data points are in the file, each voxel has its coordinates  $x, y, z$  as 4 bytes integers and a rgb alpha colour, where

each channel is an 8 bytes floating point double value. Values not included in the file are considered to be initialized to zero.

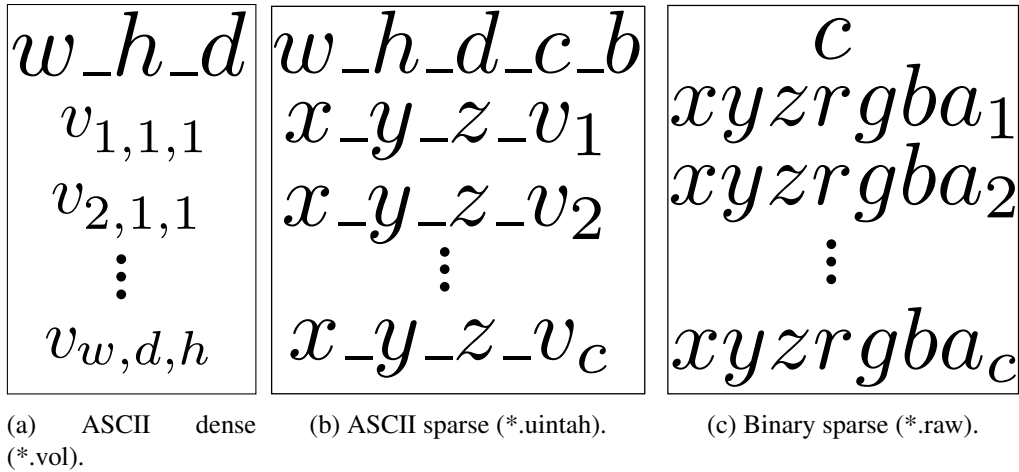


Figure 4-2: Supported file formats, \_ represent the character space.

## 4.2 Mental Ray® Rendering Approach

How mental ray shoots rays around and what is its paradigm.

Mental Ray® approach to solving the rendering equation is based on path tracing, as shown in Figure 4-3, for each pixel in the camera view, an eye ray will be shot in the scene. On an intersection with an object in the scene, its material shader will be called, this shader will shoot a light ray for each light in the scene, which in effects calls the light shader of the given light. In order to compute the irradiance at the intersection point, the light shader will probably trace a shadow ray from the light to the intersection point. Eventually, the material shader will compute the final colour with the information received from the light shader.

Under this assumptions, the steps required to solve the RTE, Equation 3.3, are:

1. Shoot a ray from the eye into the scene.
2. If the ray intersects the fire volume, perform ray marching in the volume.
  - (a) Compute direct illumination at the current point.
    - i. Get initial radiance using black body radiation.
    - ii. Attenuate with absorption coefficient.
  - (b) Compute indirect illumination at the current point.
    - i. Use phase function to get the scattered light distribution.
  - (c) Advance to next point in ray marching.
    - i. Change ray direction using refraction index.

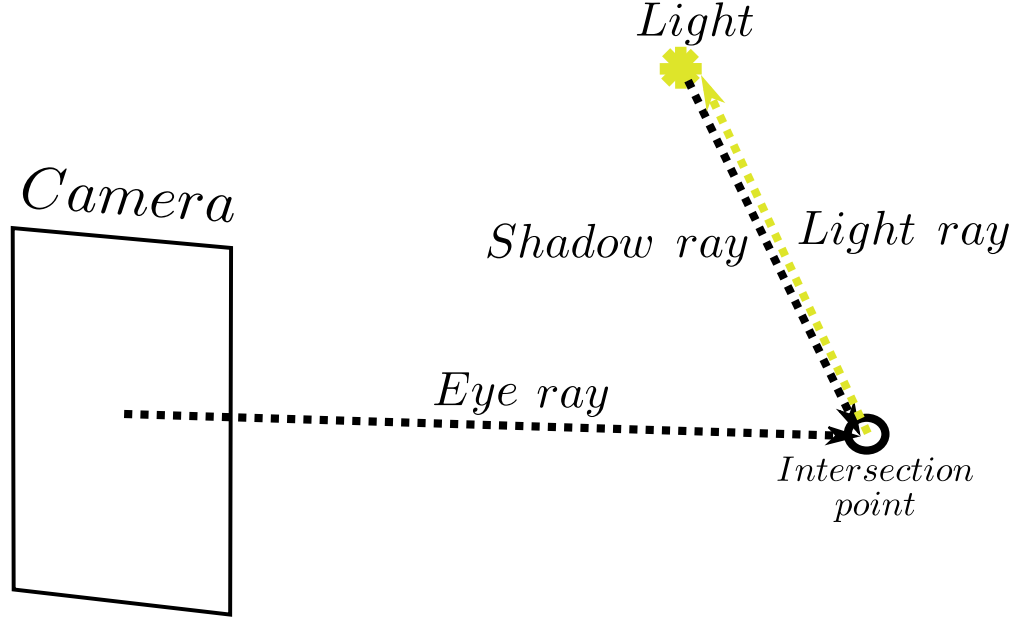


Figure 4-3: Mental Ray<sup>®</sup> simple ray casting example.

3. Compute eye visual adaptation to fire colour.

### 4.3 Simplifications

Equation 3.3 provides a radiance value,  $L(\lambda, \mathbf{x} + \Delta\mathbf{x}, \omega)$  for the next march increment, however it is more intuitive to think about the final radiance  $L(\lambda, \mathbf{x}, \omega)$ , for the first intersection point. The radiance at the interest point is given by

$$L(\lambda, \mathbf{x}, \omega) = e^{-\sigma_t(\lambda, \mathbf{x})\|\Delta\mathbf{x}\|} L(\lambda, \mathbf{x} + \Delta\mathbf{x}, \omega) + (1 - e^{-\sigma_t(\lambda, \mathbf{x})\|\Delta\mathbf{x}\|}) \frac{\sigma_a(\lambda, \mathbf{x})L_e(\lambda, \mathbf{x}, \omega) + \sigma_s(\lambda, \mathbf{x})L_i(\lambda, \mathbf{x}, \omega)}{\sigma_t(\lambda, \mathbf{x})}. \quad (4.1)$$

In fire phenomena the effect of scattered light in the final image is practically imperceptible [Pegoraro and Parker, 2006]. Intuitively, it means that as the medium is highly emissive and thin, most of the emitted photons have their initial paths unaltered. The scattering directions  $\omega_i$  are in essence sampling directions over a sphere of scattered light  $L_i$ . In practice, the simplification will significantly decrease rendering time, as  $\omega_i$  would have to be sampled by means of a computationally expensive Monte Carlo approximation. Exploiting the previous prior knowledge, the scattering contributions can be safely ignored by setting  $\sigma_s = 0$ , which simplifies the previous equation to

$$L(\lambda, \mathbf{x}, \omega) = e^{-\sigma_a(\lambda, \mathbf{x})\|\Delta\mathbf{x}\|} L(\lambda, \mathbf{x} + \Delta\mathbf{x}, \omega) + (1 - e^{-\sigma_a(\lambda, \mathbf{x})\|\Delta\mathbf{x}\|}) L_e(\lambda, \mathbf{x}, \omega). \quad (4.2)$$

A visual representation of the implications of the aforementioned simplification is shown in Figure 4-4. Note that a single point light is used in the diagram to avoid clutter, yet in our implementation there is a point light located at the centre of each voxel whose temperature is high enough to be emissive. The non-linear trajectories of the photons, of the rays in practice, due to varying refraction indices in the media were also outside of the scope of this project. Deformations produced by the refraction phenomena can be noticeable, however there are important simplifications in the implementation if we choose to ignore them. Namely, we can reduce the recursiveness significantly, instead of starting in the first intersection and recursively solve the equation, we can start from the end point and accumulate the result backwards with a loop, as computing the end point in this situation is trivial.

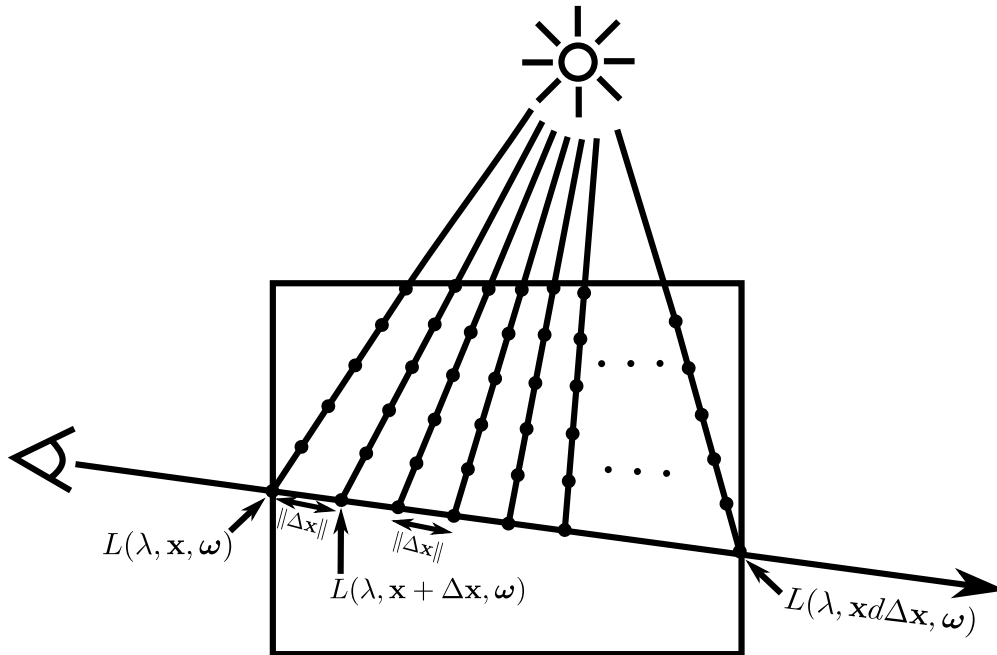


Figure 4-4: Calculating the illumination values for samples along a ray.

## 4.4 Shaders

Talk about things like which parts are written in parallel code, instance support, shader internal memory, sparse data, how the software escalates, memory consumption, Maya integration, spectrum to RGB integration, maybe more details about units for black body radiation and everything that was not explained before.

In order to be able to implement the different parts of our architecture, input data, voxel constant data, etc, without a monolithic design, lead to shader organization depicted in Figure 4-5. The temperature and density shaders will read the raw input values, the emission and absorption shaders will compute respectively, black body radiation and absorption coefficients. The fire light shader will sample and compute emitted light  $L_e$ , and the fire volume shader is in charge of adding absorption from light to interest point and to perform the main ray marching computations.

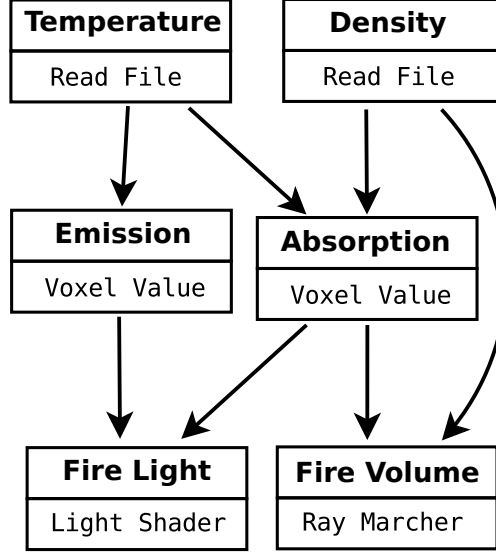
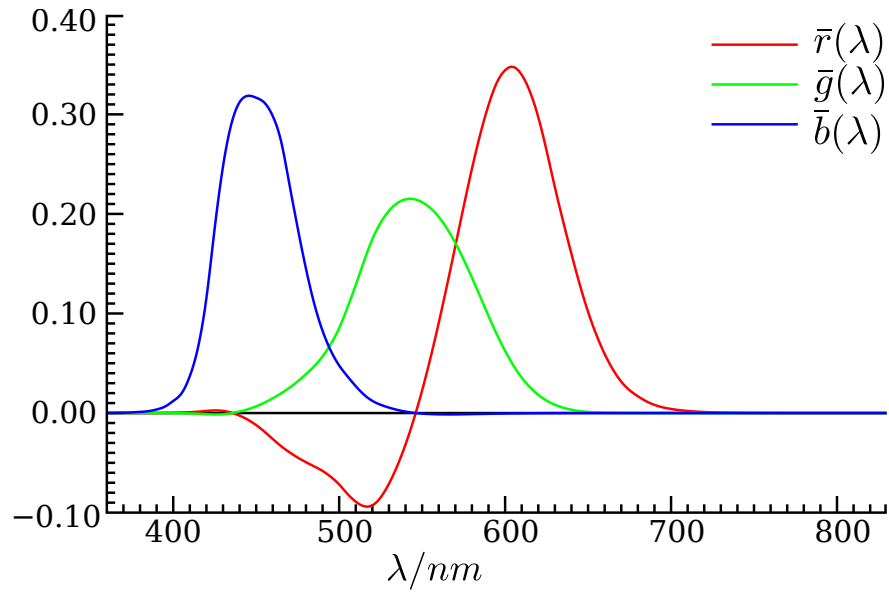


Figure 4-5: Diagram depicting the shaders and their interconnections.

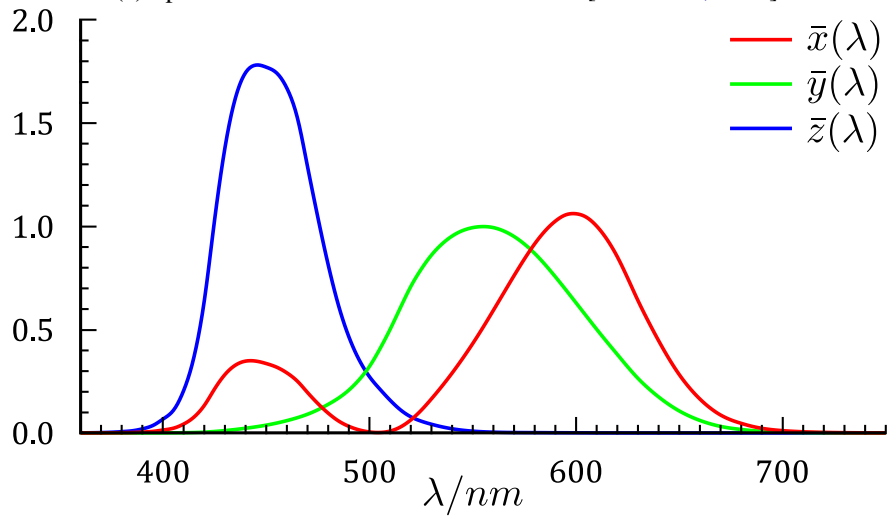
Volumetric data tends to be quite sparse, and usually we are working with volumes of  $256 \times 256 \times 256$ , which assuming a standard 32 bits floating point value per voxel, accumulates to 64 megabytes of data. With the shader architecture shown in Figure 4-5, there will be at least four copies of those structures per fire in the scene. A total of 256 megabytes per voxel dataset is unacceptable, we dealt with the problem using an open source sparse voxel dataset library [OpenVDB, 2015].

To speed up the render time, all spectrum related computations are integrated at source to RGB coefficients. Our spectrum class has a number of samples that can be set at compile time, 30 samples were used in all of our renderings. To compute the absorption coefficients that were introduced in Section 3.5, data from Table 3.1 is used. As there is not enough data to compute the 30 spectrum coefficients, a linear interpolation is used to fill the missing values. RGB coefficients are the standard colour representation in computer screens, however not every colour visible by the human eye can be reproduced in RGB space. There are further disadvantages inherent to this representation, for example certain colours visible by the human eye would require negative values for the  $\bar{r}(\lambda)$  coefficient, as shown in Figure 4-6a. The XYZ colour space is built using three sets of imaginary primaries, which have a series of favourable properties. Any colour which can be seen by the human eye can be represented with some  $x(\lambda)$ ,  $y(\lambda)$ ,  $z(\lambda)$  coefficients, the Y channel is equal to the photopic luminous efficiency function  $V(\lambda)$ , which models the variation of perceived brightness, and as shown in Figure 4-6b, the coefficients are always positive.

Given a Spectral Power Distribution (SPD)  $S(\lambda)$ , the  $x$ ,  $y$ ,  $z$  coefficients for  $S(\lambda)$  in the XYZ space are defined with respect to spectral matching curves  $X(\lambda)$ ,  $Y(\lambda)$  and  $Z(\lambda)$ , such that



(a) Spectral tristimulus values for CIE RGB [CIE.RGB, 2015].



(b) Spectral tristimulus values for CIE XYZ [CIE.XYZ, 2015].

Figure 4-6: Colour matching curves for RGB and XYZ spaces.

$$\begin{aligned}
x &= \frac{1}{\int Y(\lambda)d\lambda} \int_{\lambda} S(\lambda)X(\lambda)d\lambda, \\
y &= \frac{1}{\int Y(\lambda)d\lambda} \int_{\lambda} S(\lambda)Y(\lambda)d\lambda, \\
z &= \frac{1}{\int Y(\lambda)d\lambda} \int_{\lambda} S(\lambda)Z(\lambda)d\lambda.
\end{aligned} \tag{4.3}$$

The term  $1/\int Y(\lambda)d\lambda$  is added in each equation to act as a normalization factor for the colour brightness. As we are concerned with sampled values on a discrete domain, the integration for XYZ coefficients is approximated by a Riemann sum

$$\begin{aligned}
x &\approx \frac{1}{\int Y(\lambda)d\lambda} \sum_i X_i c_i, \\
y &\approx \frac{1}{\int Y(\lambda)d\lambda} \sum_i Y_i c_i, \\
z &\approx \frac{1}{\int Y(\lambda)d\lambda} \sum_i Z_i c_i,
\end{aligned} \tag{4.4}$$

where  $c_i$  is the  $i^{th}$  coefficient in the SPD,  $X_i$  is area for the the corresponding sample range the  $X(\lambda)$  curve and equivalently for  $Y_i$  and  $Z_i$ . Note that  $1/(\int Y(\lambda)d\lambda)$ ,  $X_i$ ,  $Y_i$  and  $Z_i$  are constants and they can be precomputed for efficiency. For the Riemann sum, the area under the samples is approximated using piecewise linear interpolation. The conversion from XYZ to RGB space is performed using

$$\begin{bmatrix} r \\ g \\ b \end{bmatrix} = \begin{bmatrix} \int R(\lambda)X(\lambda)d\lambda & \int R(\lambda)Y(\lambda)d\lambda & \int R(\lambda)Z(\lambda)d\lambda \\ \int G(\lambda)X(\lambda)d\lambda & \int G(\lambda)Y(\lambda)d\lambda & \int G(\lambda)Z(\lambda)d\lambda \\ \int B(\lambda)X(\lambda)d\lambda & \int B(\lambda)Y(\lambda)d\lambda & \int B(\lambda)Z(\lambda)d\lambda \end{bmatrix} \begin{bmatrix} x \\ y \\ z \end{bmatrix}, \tag{4.5}$$

where  $R(\lambda)$ ,  $G(\lambda)$ ,  $B(\lambda)$  are the spectral curves for the red, green and blue colours respectively. Note that all the factors in the transformation matrix are constants and they can be precomputed for efficiency.

#### 4.4.1 Fire Volume Shader

In order to achieve smoother rendering results, a trilinear interpolation is computed in each step in the ray-marching algorithm. Computing the interpolation smoothed the shape of the flames and reduced the effects of outliers in the input data. Tricubic interpolation was also considered, however it was discarded due to the significant increase in computational overhead.

The implementation of the eye visual adaptation to the colours in fire, described in Equation 3.26 is performed as described in [Nguyen et al., 2002], which uses the Von Kries model [Fairchild, 2005]

$$\begin{aligned}
\mathbf{x}_a &= \mathbf{M}^{-1} \mathbf{T}_w \mathbf{M} \mathbf{x}_i, \\
\mathbf{l}_{max} &= \mathbf{M} \mathbf{x}_{max}, \\
\mathbf{T}_w &= \begin{bmatrix} 1/l_{max} & 0 & 0 \\ 0 & 1/m_{max} & 0 \\ 0 & 0 & 1/s_{max} \end{bmatrix}, \\
\mathbf{M} &= \begin{bmatrix} 0.4002 & 0.7076 & -0.0808 \\ -0.2263 & 1.1653 & 0.0457 \\ 0 & 0 & 0.9182 \end{bmatrix},
\end{aligned} \tag{4.6}$$

where  $\mathbf{x}_a = [x_a, y_a, z_a]^T$  is a column vector with the XYZ adapted coefficients,  $\mathbf{T}_w$  is the Von Kries transformation,  $x_{max}$  are the coefficients for the maximum temperature present in the fire,  $\mathbf{l}_{max} = [l_{max}, m_{max}, s_{max}]^T$  are the LMS coefficients of  $x_{max}$ ,  $\mathbf{M}$  is a XYZ to LMS transformation matrix and,  $\mathbf{x}_i = [x_i, y_i, z_i]^T$  is a column vector with the XYZ input coefficients. The XYZ to LMS  $\mathbf{M}$  matrix is a Hunt-Pointer-Estevéz transformation, proposed by [Hunt and Pointer, 1985], which has been normalized to the CIE Standard Illuminant D65. The coefficients for  $M^{-1}$  are not provided in the CIE specification. Given that  $M$  is a small  $3 \times 3$  matrix, the coefficients of  $M^{-1}$  can be evaluated with any of the standard matrix inversion methods, in our case the default algorithm in Matlab<sup>®</sup> *inv()* function was used, and the output was hard-coded in the shader. The aforementioned equation is solved in the XYZ space, in order to delay clamping coefficients to RGB as much as possible.

#### 4.4.2 Fire Light Shader

#### 4.4.3 Voxel Dataset Shader

### 4.5 Miscellaneous

#### 4.5.1 Utilities

Several auxiliary scripts have been developed in order to aid and provide some automation while using the software. A bash render script “render.sh” is provided, this script batch renders a given Maya<sup>®</sup> scene and creates a video with the output images.

New fuel types can be added to extend the default selection using spectral lines from the NIST database [NIST, 2015]. A Matlab<sup>®</sup> script “save\_nist\_data.m” can be used to download and save new “.specline” files for new atoms. Once a new file has been added to the data folder, the files “fire\_shader.mi” and “FuelTypes.h” have to be updated to include the name of the new data.

Untah [Untah, 2015] is the combustion simulator used by [Pegoraro and Parker, 2006] as input for their rendering method. Detailed instructions for its installation on Ubuntu, how to run examples, as well as scripts to convert from their output data format to a format compatible with the shaders is provided with the source code.



## 4.5.2 Known bugs and workarounds

Batch rendering through the Maya® GUI does not automatically update the input files for the shaders. A workaround is to use the command line render utility as shown below

Listing 4.1: Batch render command

```
$ Render -r mr -perframe -s <start_frame_number> -e <end_frame_number> <path_to_scene_file>
```

Saving images, either from batch render or single frame gui render, will produce unexpected results, normally a completely white output, if the image format natively supports transparency, such as tif, tga or exr. In order to get consistent results, the scene should be surrounded by solid objects.

Any Maya® scene which uses the shaders should be saved in Maya® ASCII format, “<scene>.ma”. The saved files are not platform independent, as the input file names for the shaders use absolute paths. If the scene is to be opened in a new environment, the file paths have to be modified to match the current locations.

## **Chapter 5**

# **Results**

Images everywhere and proper analysis of what the user is seeing, what is failing and why. Include default Maya fire effect, pictures of fire from the papers mentioned in the previous work section and a historical overview on the improvements in the shaders.

## **Chapter 6**

# **Conclusions and Future work**

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