

N-Body Simulations with REBOUND

Lab course protocol

Group 3

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Abstract

This is optional, but never longer than half a page.

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1 Introduction

In the experiment, we will use the 80 cm telescope of the IAAT to study photometry and spectroscopy. These two topics are fundamental to observational astronomy and are used to study the properties of stars, galaxies and other bodies. In photometry, we aim to determine the brightness of several stars in the star cluster M37 in the B and V filters which will be used to determine the age and distance to the cluster. In spectroscopy, we will look at the spectrum of one (or both) of the stars of the Castor binary system and identify the $H\alpha$ spectral line (and more).

2 Theory

2.1 Telescope

Telescopes are used in observing the EM radiation emitted by distant sources. Telescopes collect light from a certain direction and are used to apply an angular magnification to the image. The two major types of telescopes are the lens telescope (the refractor) or the mirror telescope (the reflector). The amount of light that a telescope can gather depends on the ‘aperture’ of the telescope D . The ratio of this and the focal length F is the aperture ratio F/D . A telescope with a higher aperture ratio will produce a brighter image. The magnification of the telescope is determined by the ratio of the main mirror focal length F and the eye piece focal length f . i.e. $M = F/f$.

The resolving power of a telescope is its ability to distinguish between two nearby sources and is defined by the Rayleigh criterion:

$$\alpha = 1.22 \frac{\lambda}{D} \quad (1)$$

Unfortunately due to atmospheric effects, the resolving power is slightly worse and is limited to $\alpha = \lambda/D$. For a telescope with an aperture of 80 cm, the resolving power at $\lambda = 5700\text{\AA}$ is

2.1.1 The 80 cm Telescope

The 80 cm telescope is located in Sand, Tübingen and was built in 2003. It is a mirror telescope in Cassegrain configuration with Nasmyth focus. It has an aperture of 80 cm and a focal length of 6.4 m. For slewing, the telescope is mounted on an equatorial open fork mount and can move to focus on any point in the sky in its line of sight.

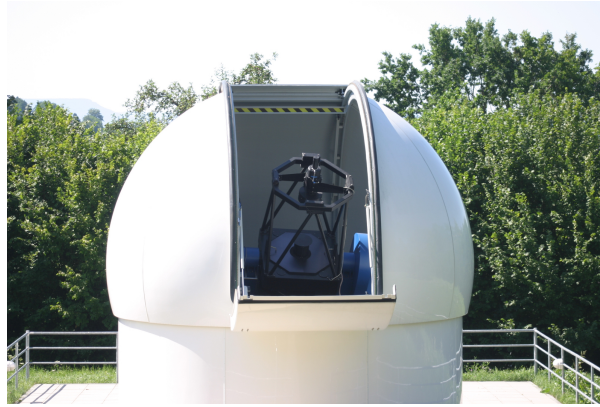


Figure 1: The 80 cm telescope at the IAAT.

In figure 2 we can see the path light has to travel to reach the focus points. The starlight enters from above and is reflected by the primary mirror towards the secondary mirror. It reflects back the light towards a plane mirror which focuses the light on either the left or the right focus. This configuration (Cassegrain-Nasmyth) allows a smaller size of the telescope (around 3 m) for a larger focal length of 6.4 m. The Nasmyth focus also is beneficial for mounting heavy instruments that don't need to be moved with the declination axis.

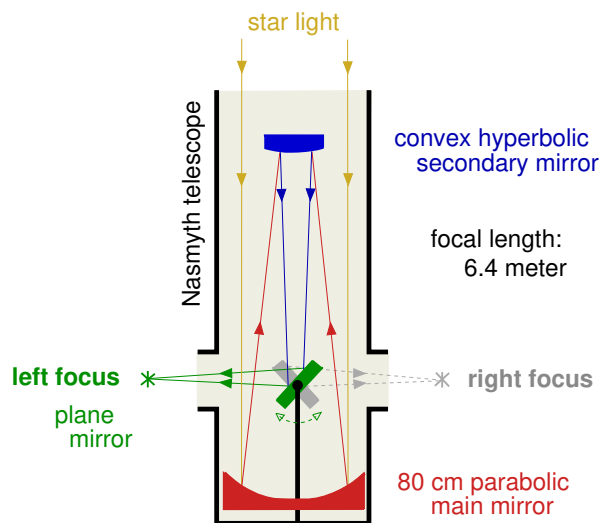


Figure 2: The light path of the 80 cm telescope.

2.2 CCD

2.3 Photometry

2.4 Spectroscopy

Studying the spectrum of stars can reveal information like their temperature, composition and their physical properties. The absorption and emission lines present in a spectra

caused by certain transitions in the elements present in the star and help in identifying its chemical composition. These lines are used to group stars into different spectral classes.

2.4.1 Star Classification

Using the Morgan and Keenan system, stars are classified depending on their spectrum and their luminosity. Both these parameters are required for classification since stars at different stages of their life can have the same spectral class, while their sizes may strongly differ. Therefore to make their classification unambiguous ‘luminosity class’ is also used. The luminosity classes are defined in the figure 3.

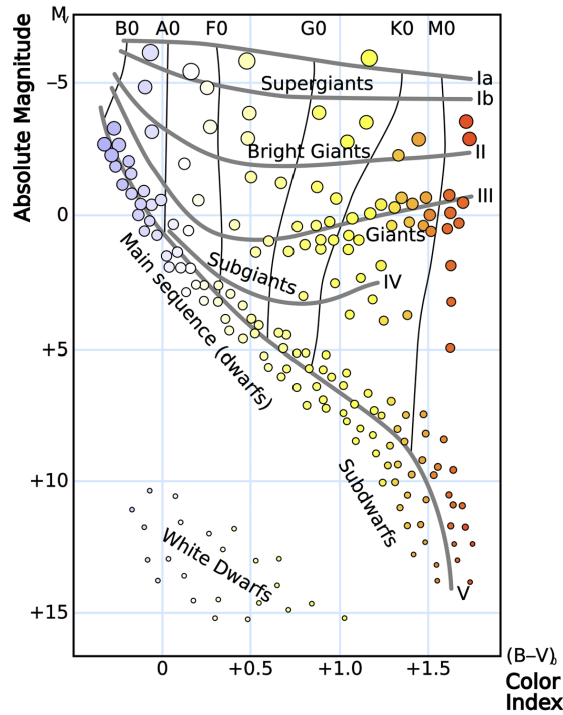


Figure 3: HRD diagram with Luminosity Classification

Stars are classified into a ‘spectral class’ based on the strength of their various spectral lines. The spectral class of a star is also a direct measure of the temperature of the star. The spectral classes are labelled ‘O’, ‘B’, ‘A’, ‘F’, ‘G’, ‘K’, ‘M’, ‘L’ and ‘T’ in decreasing order of temperature. Each is further split into 10 sub classes from 0 to 9.

Spectral Type	Temperature Range (K)	Distinguishing Features
O	$> 25,000\text{K}$	H; HeI; HeII
B	10,000-25,000K	H; HeI; HeII absent
A	7,500-10,000K	H; CaII; HeI and HeII absent
F	6,000-7,500K	H; metals (CaII, Fe, etc)
G	5,000-6,000K	H; metals; some molecular species
K	3,500-5,000K	metals; some molecular species
M	$\leq 3,500\text{K}$ metals;	molecular species (TiO!)
C	$\leq 3,500\text{K}$ metals;	molecular species (C2!)

Table 1: Temperature and Absorption lines at different spectral Classes[]

Stellar plasma below the photosphere creates a black body spectrum. The temperature at which this peaks is the effective temperature of the star. As the temperature in the outer regions of the photosphere is lower, different elements absorb parts of the light coming from deeper in the star and create the absorption lines. These lines are highly dependent on the effective temperature and the elements present in the photosphere. The most abundant absorption lines are usually the H_α , H_β , H_γ and H_δ lines of the Balmer series. These are caused by ionized hydrogen, which is the most abundant elements in stars. Hydrogen lines are most abundant in A type stars. For G type and cooler stars, Calcium lines are more dominant. Figure 4 shows this in more detail.

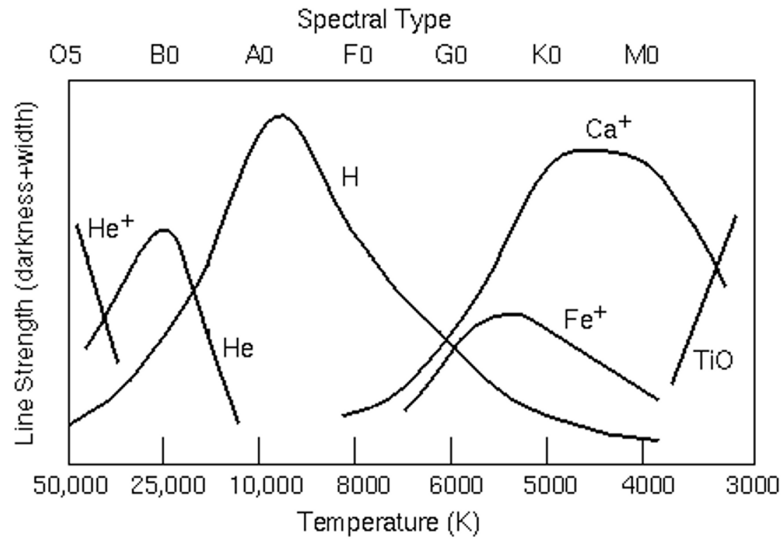


Figure 4: Strength of absorption lines at different temperatures

Figure 5 shows the spectra of stars of different sizes. Most of the spectral features are found in UV, hence spectroscopy is very popular in UV astronomy.

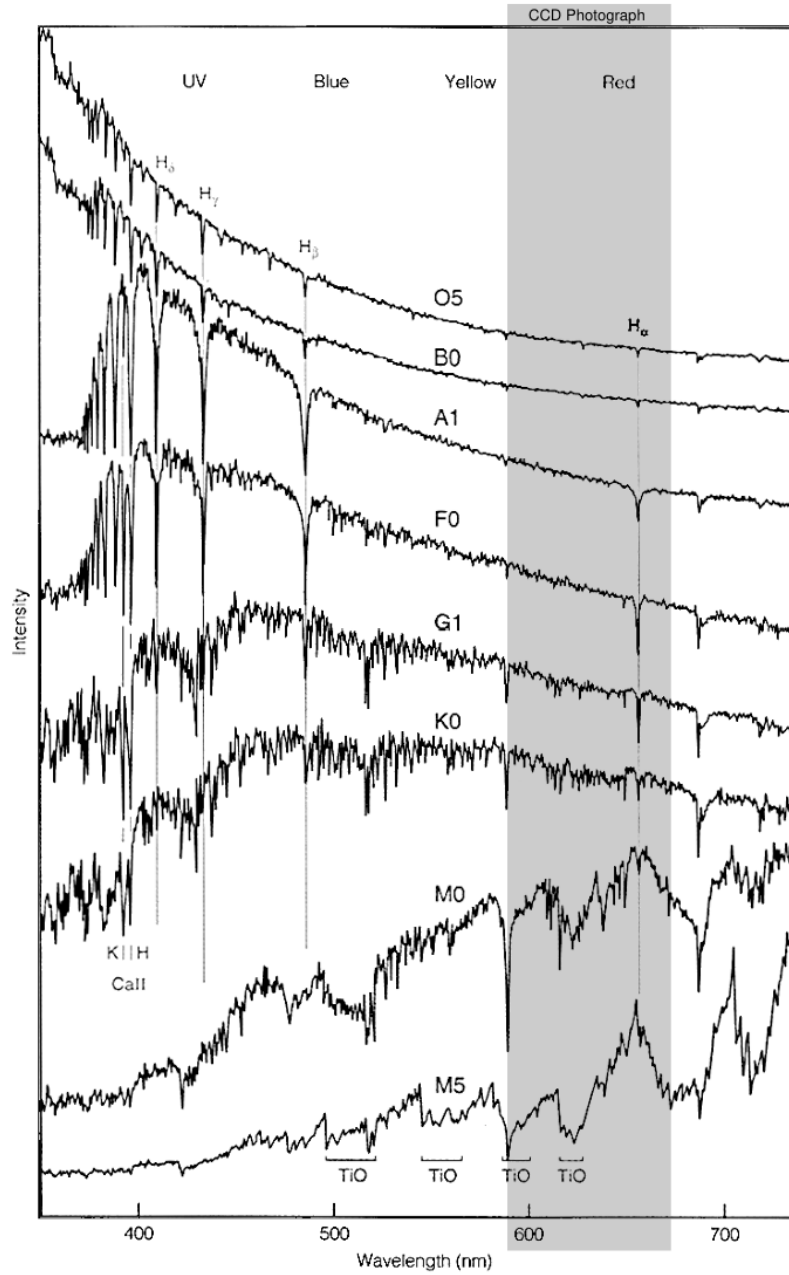


Figure 5: Spectra of stars of different spectral classes.

3 Experiment

3.1 Spectroscopy

We observed the spectrum of one (or both) stars in the Castor binary system. To observe the spectra, the 10C spectrograph was mounted on the left focus of the 80 cm telescope. The spectrograph uses a CCD camera to obtain the spectral image of the star. To control the spectrograph, ‘Maxim DL’ software is used.

3.1.1 Recording the Data

In order to obtain the spectra, the steps are listed below:

1. The telescope is powered on and the operation software is started on the ‘Telescope Computer’, ‘Navigation and Observation Computer’ and the ‘Spectral Analysis Computer’.
2. The calibration lamps and the CCD of the spectrograph are powered on.
3. Adjust the focus of the CCD. This must be set up manually directly at the spectrograph by following the markings on the camera. When mounting the spectrograph, be careful not To touch the focus dial.
4. The telescope is slewed to the target star. Out of focus, the star will appear as a donut shaped object. The focus is manually adjusted until we observe a star like object(s).
5. When the calibration lamps are on, the slit is visible between the two red discs that appear in the eyepiece. The telescope is slewed slightly until one of the stars is directly on the slit.
6. The spectral image can now be taken by the spectrograph.

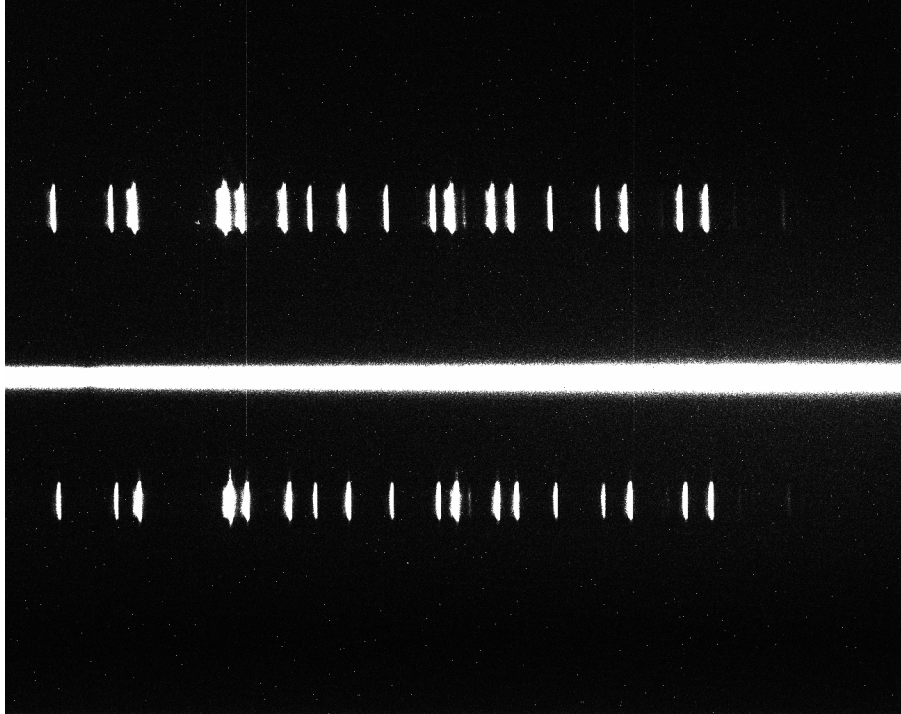


Figure 6: Spectral image of the Castor binary system obtained using the SX-HX9 CCD camera of the 10C spectrograph and the 80 cm telescope. We observe the top and bottom calibration lines and the star spectrum in the middle. Not much is directly visible in the image except possibly the H_{α} line

3.1.2 Data Analysis

The obtained spectral image is analysed using the ‘Spec-Analyzer’ software. The software interactively guides the user through the steps to analyze the image. The software is written In IDL/GDL and involves command line inputs and mouse pointer interactions. Care needs to be taken when clicking mouse buttons as it can send incorrect inputs to the software. If a mistake is made, the process needs to be started again.

In figure 6 The top and bottom spectrums are created by Neon lamps and are used for calibration of the star spectrun. Currently, the star spectrum appears as a continuous bright band in the middle and only a small decrease in intensity near what is possibly the H_α line is visible.

Using spec-analyzer, the final spectra is obtained with the following steps:

1. The software scans for three traces from the image, the top and bottom calibration traces(`cal-trace1` and `cal-trace2`) and the middle trace for the star (`spec-trace`). A ‘trace’ is a 1-D array that contains all the pixel values along the x direction for a fixed y. For ‘spec-trace’, we take average of several traces along the middle region which covers the star spectrum.
2. We now obtain the raw uncalibrated spectrum.

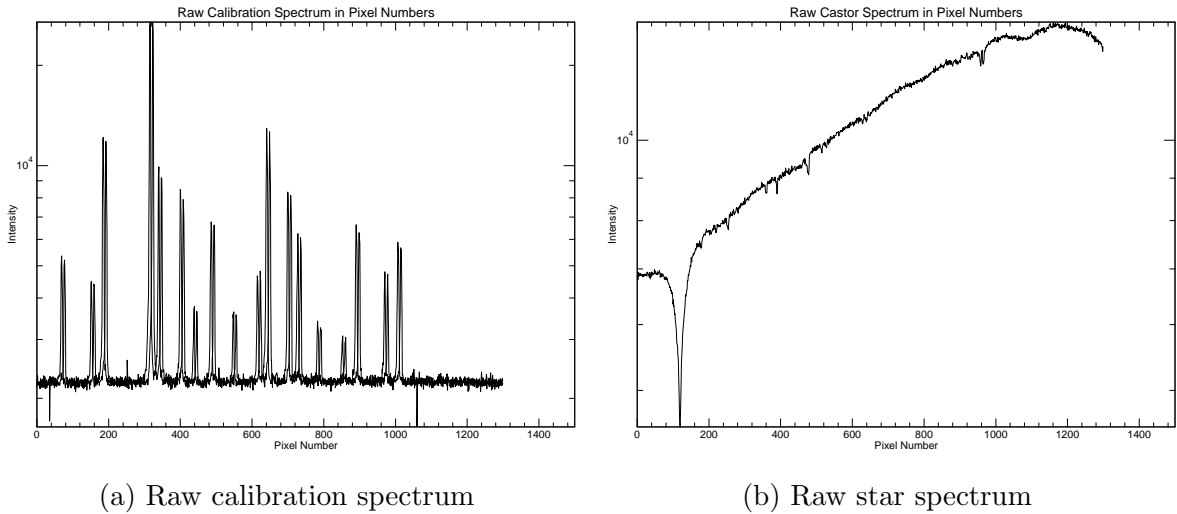


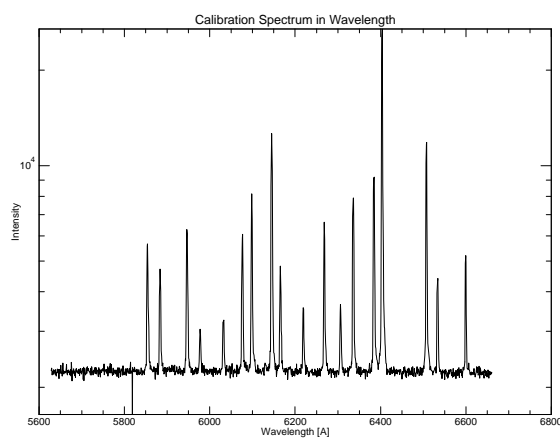
Figure 7: The raw spectrum with the pixel number on the x axis and the intensity on the y axis.

3. Due to the CCD camera being removable, there will be a skew between the image plane and the x axis of the spectrograph which needs to be corrected for. This can be done by correlating both the calibration spectrums and the maximum correlation can be observed manually and is sent as an input to the software. Here, it was found to be 0.8.
4. Now, we need to calibrate the observed calibration spectra to a reference spectrum with the same Neon lines. This is difficult to do hence we will compare the observed spectrum image with a calibrated spectrum image. This calibrated image is shown in figure 8.

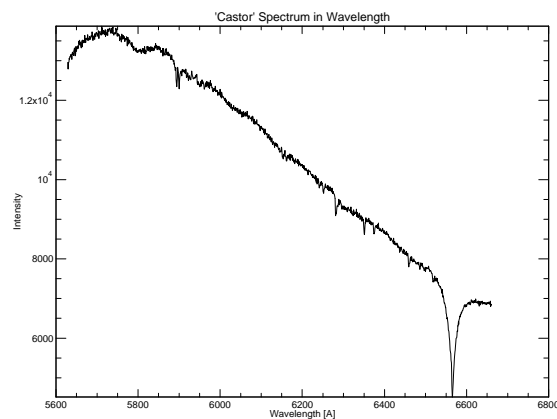


Figure 8: Calibrated spectrum image. This was recorded using earlier using the 10C spectrograph calibration lambs for lines in optical spectrum($4000\text{\AA} - 8000\text{\AA}$) These lines are compared with the observed calibration spectrum in figure 6

5. We manually compare both these images to get the wavelengths of any 5 matching lines in the observed spectrum. This will finally convert the pixel numbers to wavelength numbers. If chosen correctly, the graph of wavelength vs pixel numbers would be a straight line.
6. We now have the calibrated spectrum



(a) Calibrated spectrum for observed calibration lines



(b) Calibrated star spectrum

Figure 9: The calibrated spectrum with the wavelength on the x axis and the intensity on the y axis

4 Conclusions

An important section in which you should critically review the experiment and its results. Mention also parts that did not work out as expected, but keep a neutral to positive view. This can span from a few sentences to half a page.

References

- URL: <https://lweb.cfa.harvard.edu/~pberlind/atlas/htmls/note.html>.

Appendix

Please attach here your original handwritten notes and other documents created during the experiment.