

# taurex-emcee: automated, parallelized atmospheric retrievals with TauREx 3.1 and the emcee sampler

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## Summary

taurex-emcee is a plugin for the TauREx 3.1 atmospheric retrieval framework ([A. F. Al-Refaie et al., 2021](#)) that extends the choice of sampling methods available to the user. The plugin provides an interface to the `emcee` sampler ([Foreman-Mackey et al., 2013](#)), a popular affine-invariant ensemble sampler widely used in the astronomy community. Running the sampler to convergence is automated through the `autoemcee` package, which also supports parallelization with MPI. Thus, the taurex-emcee plugin allows users to easily launch parallelized retrievals of atmospheric spectra with emcee. This enables reliable, efficient, and fast retrievals, especially when coupled with TauREx's GPU-accelerated forward models ([A. Al-Refaie et al., 2020](#)).

taurex-emcee is released under the BSD 3-Clause license and is available on [GitHub](#). The plugin can be installed from the source code or from [PyPI](#), so it can be installed as `pip install taurex-emcee`. The documentation is available on [readthedocs](#), including a quick-start guide, a tutorial, a description of the software functionalities, and guidelines for developers. The documentation is continuously updated and is versioned to match the software releases.

## Benchmark

We benchmarked the taurex-emcee plugin against the MultiNest sampler ([F. Feroz et al., 2009; Farhan Feroz et al., 2019](#)), natively implemented in TauREx 3, on a synthetic transmission spectrum of the hot Jupiter HD 209458b. The aim being first to assess the computational time of the two samplers on a controlled case, second to assess the consistency of the results from the retrievals. The synthetic spectrum and the retrievals were performed on a single node of the Sapienza University of Rome Melodie server, equipped with one NVIDIA A40 GPU. The high-resolution forward spectrum was generated assuming stellar and planetary parameters of HD 209458b from Edwards et al. (2019), reported in [Table 1](#).

**Table 1:** Summary of the selected target's properties.

$R_p$ [ $R_J$ ]	$M_p$ [ $M_J$ ]	$T_p$ [K]	P [d]	$R_s$ [ $R_\odot$ ]	Mag K	$T_s$ [K]	$M_s$ [ $M_\odot$ ]
1.35	0.71	1613	3.52	1.18	6.31	6,086	1.18

The stellar spectrum is simulated with the [Phoenix stellar models](#) ([Baraffe et al., 2015](#)), and the planetary atmosphere is gaseous, with hydrogen and helium at a ratio  $H_2/He = 0.172$ ,

and has an isothermal temperature profile. We consider five molecular species as atmospheric trace gases: H<sub>2</sub>O (100 ppm), CH<sub>4</sub> (10 ppm), CO (1 ppm), CO<sub>2</sub> (0.1 ppm), and NH<sub>3</sub> (0.01 ppm). Molecular abundances are assumed constant with altitude. We utilize cross sections at a resolution of 15,000 for all species, as given in [Table 2](#). Collision-induced absorption (CIA) with H<sub>2</sub>–H<sub>2</sub> and H<sub>2</sub>–He and Rayleigh scattering are included in the calculation. We also include fully-opaque gray clouds with a cloud-top pressure of 1000 Pa. Finally, 100 pressure layers are used to sample the atmosphere, from 1 Pa to 10<sup>6</sup> Pa, uniformly in log-space.

**Table 2:** Cross sections and CIA used in the simulations.

Opacity	Reference(s)
H <sub>2</sub> -H <sub>2</sub>	Abel et al. ( <a href="#">2011</a> ), Fletcher et al. ( <a href="#">2018</a> )
H <sub>2</sub> -He	Abel et al. ( <a href="#">2012</a> )
H <sub>2</sub> O	Polyansky et al. ( <a href="#">2018</a> )
NH <sub>3</sub>	Coles et al. ( <a href="#">2019</a> )
CO	Li et al. ( <a href="#">2015</a> )
CO <sub>2</sub>	Rothman et al. ( <a href="#">2010</a> )
CH <sub>4</sub>	Yurchenko et al. ( <a href="#">2017</a> )

To simulate the observation of the transmission spectrum, we assume an observation by the Ariel space mission ([Tinetti et al., 2021, 2018](#)) in Tier 2 mode ([Edwards et al., 2019](#)). We utilize radiometric estimates of the total noise on one observation of HD 209458b obtained with ArielRad ([Mugnai et al., 2020](#)), the Ariel radiometric simulator now available at the online [exoddb.space](#) website. ArielRad is based on the generic point source radiometric model ExoRad2 ([Mugnai et al., 2023](#)), adapted with the Ariel payload configuration. The software versions used are: ExoRad2 v2.1.113, ArielRad-payloads v0.0.17, and ArielRad v2.4.26.

In our retrieval benchmarks, we attempt to constrain the abundances of the trace gases, alongside the temperature profile, the planetary radius, and the cloud-top pressure. [Table 3](#) lists each parameter, its units, and the priors (with corresponding scale) used in the retrievals.

**Table 3:** Parameters and priors used in the retrievals.

Parameter	Unit	Prior	Scale
R <sub>P</sub>	R <sub>J</sub>	±10%	linear
T	K	100; 4000	linear
P <sub>clouds</sub>	Pa	1; 10 <sup>6</sup>	log
H <sub>2</sub> O	VMR	10 <sup>-12</sup> ; 10 <sup>-1</sup>	log
CH <sub>4</sub>	VMR	10 <sup>-12</sup> ; 10 <sup>-1</sup>	log
CO	VMR	10 <sup>-12</sup> ; 10 <sup>-1</sup>	log
CO <sub>2</sub>	VMR	10 <sup>-12</sup> ; 10 <sup>-1</sup>	log
NH <sub>3</sub>	VMR	10 <sup>-12</sup> ; 10 <sup>-1</sup>	log

We perform five retrievals each with MultiNest and emcee, with increasing number of molecules included in the free parameters (the other molecules being fixed to their true values), from only H<sub>2</sub>O to all five trace gases. With MultiNest, we set the evidence tolerance to 0.5 and sample the parameter space through 1500 live points. With emcee, we utilize 100 walkers and run two independent chains, each to convergence, where we adopt the default convergence criteria of the [autoemcee](#) package, i.e. the Geweke convergence diagnostic ([Geweke, 1991](#)) is z-score < 2.0 and the Gelman-Rubin rank diagnostic ([Gelman & Rubin, 1992](#)) is  $\hat{r} < 1.01$ . When iterating the chains, we increase the number of steps by multiplying the number of steps of the previous iteration by a growth factor of 2, a parameter that we have enabled as the

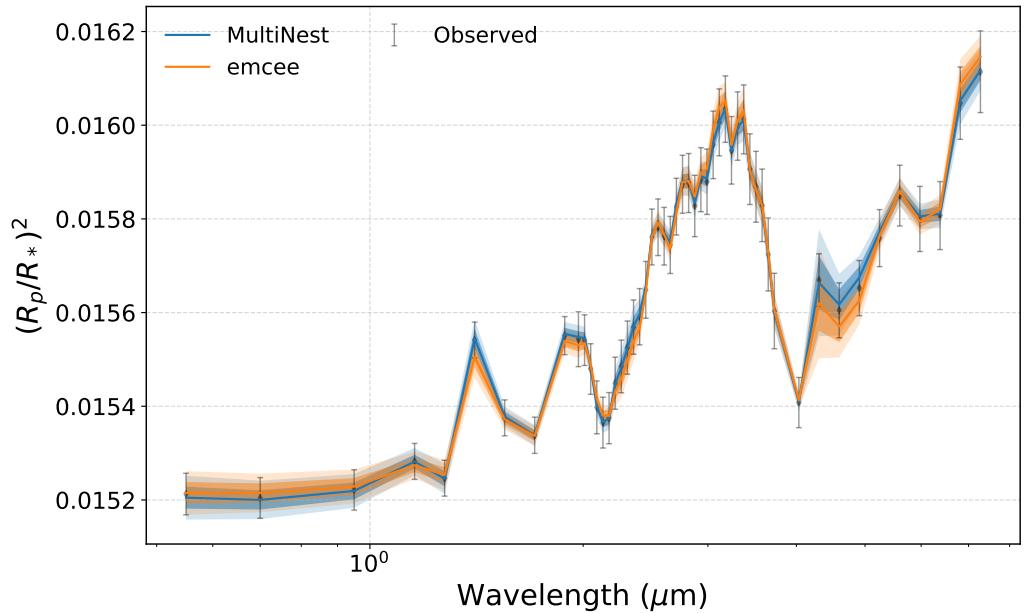
default value of 10 was deemed too conservative for our tests. As a first comparison, we report in [Table 4](#) the computational time of the retrievals with the two samplers.

**Table 4:** Runtime of the retrievals with emcee and MultiNest samplers.

Fitted molecules	Emcee [s]	MultiNest [s]
$\text{H}_2\text{O}$	16,465	5,423
$\text{H}_2\text{O}, \text{CH}_4$	6,550	7,505
$\text{H}_2\text{O}, \text{CH}_4, \text{CO}$	15,247	10,093
$\text{H}_2\text{O}, \text{CH}_4, \text{CO}, \text{CO}_2$	14,930	12,847
$\text{H}_2\text{O}, \text{CH}_4, \text{CO}, \text{CO}_2, \text{NH}_3$	30,151	17,097

MultiNest is faster than emcee for all the retrievals bar the one with  $\text{H}_2\text{O}$  and  $\text{CH}_4$ , where the computational time is around 15 minutes longer. Additionally, runtimes with MultiNest scale slower with the increase in dimensionality of the parameter space. Finally, the first retrieval with emcee is slower than the second as the chains may require more iterations to converge based on the realization of the random walkers.

Next, we compare the results of the retrievals. For brevity, we only discuss the retrieval with the full stack of trace gases included in the free parameters, as the results of the other retrievals are consistent. [Figure 1](#) shows the retrieved spectrum yielded by the two samplers, alongside the observed spectrum. The retrieved spectrum is shown as the best-fit model and the  $1\sigma$  and  $2\sigma$  confidence intervals. The retrieved spectra are consistent with each other and with the observed spectrum within the experimental uncertainties.



**Figure 1:** Best-fit spectra for the HD 209458b retrievals with MultiNest (blue) and emcee (orange). The synthetic observed spectrum is shown as the black error bars. The  $1\sigma$  and  $2\sigma$  confidence intervals are shown as the shaded regions. The error bars of the observed spectrum are generated with ArielRad ([Mugnai et al., 2020](#)) and the spectral grid corresponds to the Ariel Tier 2 mode ([Edwards et al., 2019](#)).

[Figure 2](#) shows the posteriors of the retrieved parameters, and [Table 5](#) reports the median and 16 and 84 percentiles of the posteriors.

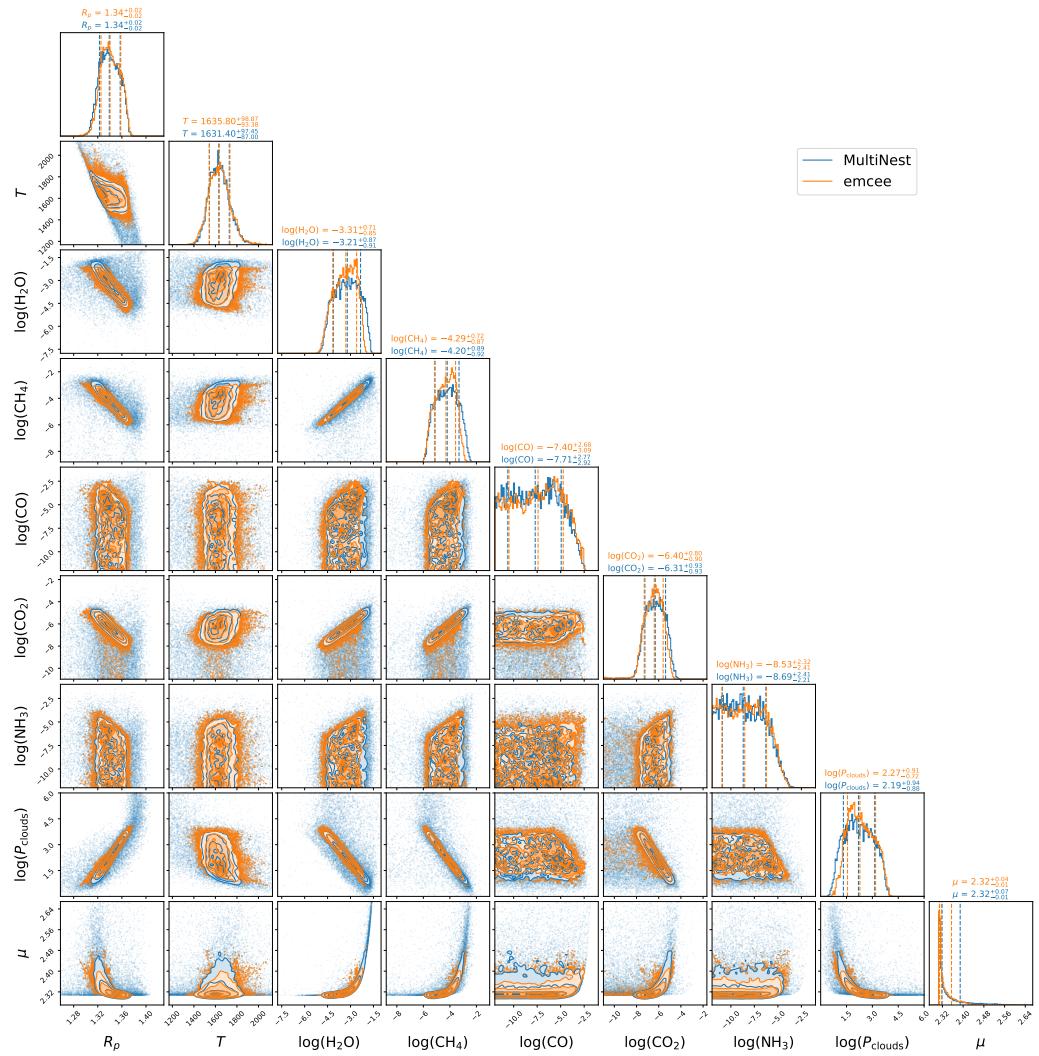


Figure 2: Caption for posteriors figure.

Table 5: Retrieved parameters and their 16 and 84 percentiles.

Parameters	Emcee	MultiNest
$R_P$	$\text{val}_{-d}^{+u}$	$\text{val}_{-d}^{+u}$
$T$	$\text{val}_{-d}^{+u}$	$\text{val}_{-d}^{+u}$
$\log(P_{\text{clouds}})$	$\text{val}_{-d}^{+u}$	$\text{val}_{-d}^{+u}$
$\log(H_2O)$	$\text{val}_{-d}^{+u}$	$\text{val}_{-d}^{+u}$
$\log(CH_4)$	$\text{val}_{-d}^{+u}$	$\text{val}_{-d}^{+u}$
$\log(CO)$	$\text{val}_{-d}^{+u}$	$\text{val}_{-d}^{+u}$
$\log(CO_2)$	$\text{val}_{-d}^{+u}$	$\text{val}_{-d}^{+u}$
$\log(NH_3)$	$\text{val}_{-d}^{+u}$	$\text{val}_{-d}^{+u}$
$\mu$ (derived)	$\text{val}_{-d}^{+u}$	$\text{val}_{-d}^{+u}$

## Statement of need

Optimized sampling methods are a key component of any retrieval code. Nested samplers (F. Feroz et al., 2009; Farhan Feroz et al., 2019) are a powerful and robust sampling method,

successfully applied to the retrieval of exoplanet atmospheric spectra (Barstow et al., 2020; Bocchieri et al., 2023; Changeat et al., 2020). TauREx 3.1 natively implements a suite of nested samplers, including the `MultiNest` sampler, or makes them available as plugins, such as the `UltraNest` sampler. The primary target of nested samplers is the efficient calculation of the Bayesian evidence, whilst the inference of the posterior is a by-product. This is regarded as a key advantage of nested samplers, as the evidence can be readily used for model selection. However, the evidence is not always required, and the interpretation of the posterior from nested samplers necessitates some care. Additionally, algorithmic assumptions of nested samplers may require to tailor the priors to explore the parameter space thoroughly.

Where the inference of the Bayesian posterior is the primary target, a well-established alternative to nested samplers are a family of Markov chain Monte Carlo methods known as affine-invariant ensemble samplers (Goodman & Weare, 2010). The implementation in `emcee` (Foreman-Mackey et al., 2013) is a popular choice in the astronomy community, as it takes care of the heavy lifting of the sampling process, is well documented, and is straightforward to utilize. To date, the `emcee` sampler is not natively implemented in the TauREx 3.1 retrieval framework, nor elsewhere in other retrieval codes, to the knowledge of the authors. To fill this gap, we developed the `taurex-emcee` plugin, which interfaces the `emcee` sampler to TauREx. Key advantages of `taurex-emcee` are that it affords a more straightforward interpretation of the posterior and is more robust to the choice of priors. However, the current implementation is not intended to explore multimodal posteriors, for which nested samplers will inevitably be more efficient. Moreover, we caveat that it may require substantial computational time to sample high-dimensional parameter spaces with `emcee`, which can be mitigated by coupling `taurex-emcee` with TauREx's GPU-accelerated forward models (A. Al-Refaie et al., 2020).

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