

The Spitzer *c2d* Survey of Large, Nearby, Interstellar Clouds

VIII. Serpens Observed with MIPS

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ABSTRACT

We present maps of 1.5 square degrees of the Serpens dark cloud at 24, 70, and 160 μ m observed with the Spitzer Space Telescope MIPS Camera. We describe the observations and briefly discuss the data processing carried out by the c2d team on these data. More than 2400 compact sources have been extracted at 24 μ m, nearly 100 at 70 μ m, and 4 at 160 μ m. We estimate completeness limits for our 24 μ m survey from Monte Carlo tests with artificial sources inserted into

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the Spitzer maps. We compare source counts, colors, and magnitudes in the Serpens cloud to two reference data sets, a 0.50 deg^2 set on a low-extinction region near the dark cloud, and a 5.3 deg^2 subset of the SWIRE ELAIS N1 data that was processed through our pipeline. These results show that there is an easily identifiable population of young stellar object candidates in the Serpens Cloud that is not present in either of the reference data sets. We also show a comparison of visual extinction and cool dust emission illustrating a close correlation between the two, and find that the most embedded YSO candidates are located in the areas of highest visual extinction.

Subject headings: infrared: general — clouds: star forming regions

1. Introduction

The Spitzer Space Telescope Legacy project “From Molecular Cores to Planet-forming Disks” includes IRAC and MIPS mapping of five large star-forming clouds (Evans et al. 2003). The Serpens cloud covers more than 10 square degrees as mapped by optical extinction (Cambr  sy 1999), but for reasons of practicality the c2d project was only able to observe 1.5 deg^2 with the MIPS instrument on Spitzer (further Spitzer observations of a larger area of Serpens are planned as part of an extended survey of the Gould Belt, Allen 2007, in prep.). At an assumed distance of 260 pc (Straizys, Cernis, & Bartasiute 1996), the area mapped by c2d corresponds to $\sim 4.5 \times 7 \text{ pc}$. This paper is one in a series describing the IRAC and MIPS observations of each of the c2d clouds. Previous papers include those on IRAC observations of Serpens (Harvey et. al. 2006), Chamaeleon (Porr  s et al. 2007), and Perseus (Jorgensen et al. 2006), as well as MIPS observations of Chamaeleon (Young et al. 2005), Perseus (Rebull et al. 2007), Lupus (Chapman et al. 2007), and Ophiuchus (Padgett et al. 2007).

Our observations of Serpens cover an area that includes the well studied “core” cluster region, Cluster A, together with the newly discovered Cluster B (Harvey et. al. 2006; Djupvik et al. 2006) to the south, as well as the Herbig Ae/Be star, VV Ser. Significant portions of this cloud have been studied by previous space infrared missions, including IRAS (Zhang et al. 1988; Zhang, Laureijs, & Clark 1988) and ISO (Kaas et al. 2004; Djupvik et al. 2006). The much higher sensitivity and longer wavelength capability of the Spitzer MIPS instrument, however, allows us to detect both very low luminosity infrared-excess objects and to map very cool diffuse dust emission in the region. Our results are also complementary to the 1.1mm mapping of the same region by Enoch et al. (2007). The combined results on Serpens using both the MIPS and IRAC observations are discussed in a companion paper

where we also give detailed object lists (Harvey et al. 2007).

In §2 we describe details of the observations obtained from the MIPS instrument for Serpens and the data processing pipeline used to reduce the observations. In §3 we describe a number of results from our MIPS observations and correlations between them and the 2MASS catalog (Skrutskie et al. 2006). We show in §3.1 that there is an excellent correlation between the coolest dust that we can observe which emits at $160\mu\text{m}$ and the optical extinction in Serpens. We investigate the possibility of time variability at $24\mu\text{m}$ in our two-epoch data set in §3.2. In §3.3 we discuss our results statistically in terms of source counts and compare these to predictions of models of the Galaxy as well as to the counts in the reference fields. We present color-color and color-magnitude plots of the population of infrared sources in §3.4 and discuss the separation of likely cloud members from the extensive background population of stars and extragalactic objects. In the final part of §3 we briefly describe some details of individual sources of particular interest.

2. Observations and Data Reduction

The MIPS observations cover an area where $A_v > 6$ in the contour maps of Cambr  sy (1999). In addition, a nearby off-cloud region of 0.5 square degrees was mapped for comparison with the cloud region. A summary of the regions observed is listed in Table 1 with the AOR (Astronomical Observation Request) number to facilitate access from the Spitzer archive. The regions covered at $24\mu\text{m}$ are outlined in Figure 1 against the $25\mu\text{m}$ IRAS sky. The observing strategy and basic MIPS data analysis for the c2d star-forming clouds have been described in detail by Rebull et al. (2007), but we summarize here the most important details. Fast scan maps were obtained at two separate epochs with a spacing between adjacent scan legs of $240''$ in each epoch. The second epoch observations were offset by $125''$ from the first in the cross-scan direction to fill in the $70\mu\text{m}$ sky coverage that would otherwise have been missed due to detector problems. The second epoch scan was also offset $80''$ from the first in the scan direction to minimize missing $160\mu\text{m}$ data. For some of the c2d clouds, these offsets together with sky rotation were sufficient to give essentially complete one-epoch coverage at $160\mu\text{m}$, but for Serpens there were still small gaps between every two scan lines. Table 2 lists the sky coverage at each wavelength. The two observation epochs were separated in time by ~ 6 hours to allow identification of asteroids in the images; over this time period asteroids will typically move $0.3 - 2$ arcminutes. Because of Serpens’ relatively large ecliptic latitude, ~ 24 degrees, only a very small number of asteroids were seen, all of which were removed by requiring 2-epoch detection in our final source lists. Typical integration times are 30 seconds at $24\mu\text{m}$, 15 seconds at $70\mu\text{m}$ and 3 seconds at $160\mu\text{m}$. Additional GTO

observations east of the region of highest emission are not included in this analysis because a different observing strategy was used. Those observations could, however, be added to ours in order to construct a somewhat larger mosaic of the region.

Figure 2 shows the three individual images produced for the MIPS bands as well as a false color image of the three together. Harvey et al. (2007) show an additional image combining the $24\mu\text{m}$ data with IRAC observations as well as enlargements of the two main clusters observed. Note that unlike the IRAC instrument, the three wavelengths of MIPS all have diffraction limited spatial resolution which means the resolution varies dramatically between $24\mu\text{m}$ ($\sim 6''$) and $160\mu\text{m}$ ($\sim 40''$).

Our data reduction is described in detail by Evans et al. (2007) but we summarize the important details here. In addition, previous versions of the c2d pipeline, some of which still apply to these data, have been described in more detail by Rebull et al. (2007) and Young et al. (2005). We began our data reduction with the BCD images, processed in this case by the standard SSC S13.2 pipeline. Following this the three MIPS channels underwent slightly different processing paths in our c2d reduction. The $24\mu\text{m}$ data were mosaicked with the SSC’s Mopex software (Makovoz & Marleau 2005) after processing in the c2d pipeline to reduce artifacts, e.g. “jailbars” near bright sources. Point sources were extracted with “c2dphot” (Harvey et al. in prep.), a source extractor based on “Dophot” (Schechter, Mateo, & Saha 1993), which utilizes the mosaics for source identification but the stack of individual BCD’s for each identified object to provide the photometry and position information. We have estimated our completeness limit at $24\mu\text{m}$ in a manner similar to that described for our IRAC photometry (Harvey et. al. 2006). We inserted a number of artificial sources into the $24\mu\text{m}$ mosaic at random positions over a range of brightness covering the range $2 < [24] < 12$ mag. and then tested whether they were properly extracted. We also produced a mosaic with only artificial sources (no real ones) but a noise level comparable to that in the observed image, and tested the completeness of extraction from that artificial image to estimate the effects of confusion in this relatively high source density region. Figure 3 shows the results from these tests. Clearly at the fainter flux levels, the effects of high source density are important to the true completeness level in Serpens, e.g. $[24] > 9.5$ mag.

The processing of the $70\mu\text{m}$ data followed a path similar to that at $24\mu\text{m}$ with two exceptions. At $70\mu\text{m}$ the SSC produces two sets of BCD’s, one of which is simply calibrated and another that is filtered spatially and temporally in a manner that makes point source identification easier but which does not conserve flux for brighter sources nor for diffuse emission. We produced mosaics of both the unfiltered and the filtered products using Mopex on the native BCD pixel scale. Point sources were extracted using APEX (Makovoz & Marleau 2005). Source reality was checked by hand inspection and comparison with the $24\mu\text{m}$ source

list. Generally the filtered mosaics were used for point source extraction, but above $F(70) \sim 2$ Jy, we used the unfiltered data. Above $F(70) \sim 23$ Jy, sources begin to be saturated. At these very high flux levels we used a procedure to fit the wings of the source profile; these data have been assigned a higher uncertainty of because of the inherent uncertainties in this procedure.

Complete tables of source positions and flux densities for likely cloud members in Serpens are given by Harvey et al. (2007) for our $3.6 - 70\mu\text{m}$ observations. At $160\mu\text{m}$ our processing was limited to producing a native pixel scale mosaic using interpolation to fill in missing pixels and point source extraction from the unfiltered mosaic. We extracted four nominal point sources in the entire mapped area. Two of these are associated with obvious multiple clumps of $24/70\mu\text{m}$ sources. The other two, SSTc2dJ1829167+0018225 (associated with IRAS 18267+0016) and SSTc2dJ18293197+0118429 (associated with source 159 of Kaas et al. (2004)) are likely powered mostly by single, shorter wavelength sources. Table 3 lists the positions and flux densities of these four nominal point sources with short comments, since their $160\mu\text{m}$ photometry is not described in any of our other publications on Serpens. None of these is in the core area of either of the main clusters. This is because large areas in those clusters are saturated, and the close spacing of many bright sources leads to the complicated, extended structure seen in Figure 2 at $160\mu\text{m}$, without obvious point-like sources.

After extraction, the source lists were bandmerged with our IRAC source lists for Serpens (Harvey et. al. 2006) and the 2MASS catalog of J, H, and K_s photometry (Skrutskie et al. 2006) as described by Evans et al. (2007). The radius for source matching with shorter wavelength detections was $4''$ at $24\mu\text{m}$ and $8''$ at $70\mu\text{m}$. Table 4 lists the number of sources extracted at 24 and $70\mu\text{m}$, and some examples of statistics of numbers identified with shorter wavelength sources. In addition to bandmerging, sources undergo a classification process based on the available photometry, 2MASS, IRAC, and MIPS. For the purposes of this paper the most important classification is that of “star” which implies a spectral energy distribution that is well-fit as a reddened stellar photosphere without requiring any excess infrared emission from possible circumstellar dust. The data reported here consist of a subset of all the sources extracted in Serpens. The entire catalog is available from the SSC website (<http://ssc.spitzer.caltech.edu/legacy/all.html>). For this paper we have limited our discussion to sources with a signal-to-noise ratio greater than 5 and to sources found in both epochs of observation to eliminate asteroids. These limits lead to a very high reliability for the objects reported here, probably greater than 98%.

In addition to our reduction of the Serpens Cloud and off-cloud data, we have also processed a 5.3 deg^2 portion of the SWIRE Spitzer Legacy data (Surace et al. 2004) from

the ELAIS N1 field through our c2d pipeline. Since this field is almost entirely populated by Galactic stars and extragalactic objects, it provides an additional control field against which to compare our Serpens Cloud population as discussed below. Note that the SWIRE observations go approximately a factor of 4 deeper than c2d due to increased integration time.

3. Results

3.1. Extended Emission

The $160\mu\text{m}$ emission traces the coolest and most extended dust seen with MIPS. Figure 4 shows an image of the $160\mu\text{m}$ emission together with contours of the optical extinction. Also shown are the locations of the two main clusters of young stellar objects in Serpens, the core Cluster A, and Cluster B (also called the G3-G6 cluster by Djupvik et al. (2006)). The optical extinction has been estimated by our fitting of the objects that were well characterized as extinguished stellar photospheres. This figure shows a very close correlation between the coolest dust and the dust that is associated with optical extinction. The figure also clearly shows that the two high-stellar-density clusters, Cluster A and B, are located in areas of maximum extinction, as we discuss further in §3.5.

3.2. $24\mu\text{m}$ Time Variability

Since many pre-main-sequence stars exhibit variable optical emission, we conducted a simple examination of the $24\mu\text{m}$ fluxes from the two observed epochs, similar to that in Perseus by Rebull et al. (2007) and for the IRAC data in Serpens (Harvey et al. 2007). As shown in Table 1, the time difference between the two epochs of observation was of order 4 hours. Figure 5 shows the ratio of the $24\mu\text{m}$ flux density between the two epochs for all the extracted sources whose signal-to-noise ratio was above 5 that were detected in both epochs of observation. Although there are a few outliers beyond the limits expected on the basis of the signal-to-noise ratios, these are all readily explained as due to poor photometry near the edges of the mosaic or problems due to source confusion or adjacency to bright sources. This is consistent with the findings of Rebull et al. (2007) for the Perseus $24\mu\text{m}$ sources and by Harvey et al. (2007) for the Serpens IRAC sources. Although there are undoubtedly some variable sources in these clouds, the observing techniques of the c2d program were not designed to enable reliable detection of modestly variable objects.

3.3. Source Counts

Because the Serpens star-forming cloud is so close to the Galactic plane, $b \sim 5$ degrees, the vast majority of the sources detected at the shorter wavelengths are background stars in the Galaxy. At fainter flux levels, background extragalactic objects constitute a significant population. In order to estimate the background Galactic star numbers we have used the Wainscoat et al. (1992) model provided by J. Carpenter (private communication). Figure 6 shows the predicted star counts from the model together with the observed counts at $24\mu\text{m}$ for both the Serpens Cloud and the off-cloud region. Also shown in the figure are the source counts from the c2d-processed SWIRE ELAIS N1 field which are largely extragalactic for fluxes below a few mJy. This figure shows that contamination by Galactic stars at the brighter fluxes and by extragalactic sources at the faint end is a significant problem for identifying Serpens Cloud members. To address this problem we discuss our use of several color and flux criteria in the following section. It is also apparent that there is an excess of bright ($F > 300$ mJy) sources relative to the expected background counts. This excess is, in fact, real and represents the bright end of the YSO candidate population discussed in the following section.

3.4. Color-Magnitude Diagrams

The c2d team has discussed in a number of studies how the use of color-magnitude and color-color diagrams can separate likely young cloud members with infrared excesses from reddened stars and many background extragalactic sources (Young et al. 2005; Harvey et al. 2006; Rebull et al. 2007; Harvey et al. 2007). Since nearly half of the area covered by our MIPS $24\mu\text{m}$ observations was not observed with IRAC (Harvey et al. 2006), we utilize the color and magnitude criteria developed by Young et al. (2005) and refined by Rebull et al. (2007) and Chapman et al. (2007) to isolate a candidate YSO population without requiring the existence of IRAC data. The most populated diagram is naturally the color-magnitude diagram of K_s versus $K_s - [24]$ because of the much larger number of $24\mu\text{m}$ sources than $70\mu\text{m}$ ones. Figure 7 shows the distribution of sources in this diagram for the 1453 sources with S/N above 5 at $24\mu\text{m}$ and with 2MASS K_s matches within $4''$. This distribution is very similar to that seen in other well-populated c2d clouds such as Perseus (Rebull et al. 2007). A comparison of the SWIRE results, the Serpens off-cloud results, and the Serpens Cloud data shows: 1) objects in our “star” class fall in a relatively narrow band with blue K_s -[24] colors (K_s -[24] < 1) as would be expected, and 2) the part of the diagram toward redder colors is populated by a number of sources in Serpens that are not seen in either the off-cloud region or in the SWIRE data set, except at K magnitudes fainter than $K_s \sim$

14. This allows us to assign a high probability that sources in the region $K_s < 14$ and $K_s - [24] > 2$ are Serpens Cloud YSO candidates with excess emission at $24\mu\text{m}$ probably due to circumstellar dust. Note that the off-cloud area does have a population of moderately reddened objects ($K_s - [24] < 2$), well-fit as stellar photospheres that are not seen in the SWIRE sample, simply because even the off-cloud area has more reddening than the high Galactic latitude ELAIS N1 region. In order to categorize our YSO candidates crudely in terms of evolutionary state, we have drawn lines in Figure 7 indicating where objects would fall based on the YSO source classification criteria of Greene et al. (1994) using the $K_s - [24]$ color to measure the spectral slope. Table 5 lists the number of candidates and the number in each of the four classes. Although AGB stars with substantial mass loss also exhibit mid-infrared excesses, Harvey et. al. (2006) have argued that the number expected in this area is less than or of order a half dozen (four of which have already been confirmed spectroscopically as AGB interlopers by Merin et al. (in prep.)). The positions of and photometry for the YSO candidates that are not in the area covered by IRAC are given by Harvey et al. (2007) along with those in the IRAC area.

Harvey et al. (2007) discuss the comparison between YSO's selected by the criteria used here (K_s and $24\mu\text{m}$ data only) and the more restrictive criteria possible with the combination of IRAC data. They basically find that we actually may have missed 8 or 9 YSO's in the area not covered by IRAC and included a very few, 3 or 4, that may be background extragalactic sources. But the overall conclusion is that there is a good correspondence between the YSO candidates found using only MIPS and 2MASS versus those selected with a more complete data set. It is also clear that the area mapped by both IRAC and MIPS, 0.85 deg^2 contains a much higher density of YSO's, 235 or 276 deg^{-2} than does the area only covered by MIPS/ $24\mu\text{m}$ with 51 YSO's or 54 deg^{-2} . Even if we exclude the area of the two high density clusters, the area covered by the combined IRAC/MIPS observations has a YSO density a factor of 4 higher than the area not included in the IRAC observations.

We have also plotted our photometry in two other color/magnitude spaces for comparison with other c2d clouds. Figure 8 shows the distribution of sources in K_s vs. $K_s - [70]$ space. As observed by Rebull et al. (2007) in Perseus, there are a large number of likely cloud members at much brighter K_s magnitudes than seen for SWIRE extragalactic objects. In addition, there is a small population of faint (in K_s) objects that are redder than any of the SWIRE objects in both Serpens and Perseus. The four objects redder than $K_s - [70] = 15$ are all likely to be slightly less extreme versions of the sources discussed in the next section. Two of these are located in cluster A, but tend to be around the outside of the tight cluster of very red objects. The other two are in a small grouping associated with the second of the four $160\mu\text{m}$ point sources listed in Table 3. Since all of these objects were also observed in our program with IRAC, they are also listed in the appropriate tables of Harvey et al.

(2007), and all are considered high probability YSO’s.

The final color-magnitude diagram, $[24]$ vs $[24]-[70]$ is shown in Figure 9. Again this distribution is qualitatively similar to that in Perseus, although we find many fewer sources in the area overlapping the red edge of the extragalactic distribution than did Rebull et al. (2007) for their “rest of the cloud”. The Serpens distribution is qualitatively more similar to that for the NGC1333 portion of Perseus. Since many of the sources represented in this diagram for Serpens are located in one of the two principal clusters, A and B, in Serpens, it is perhaps not surprising that they would mimic some of the properties of similar young clusters like NGC 1333.

3.5. The Most Embedded Objects

We have selected the coldest, most obscured sources from our sample by looking for objects not detected in the 2MASS survey but detected with reasonable signal-to-noise at both 24 and $70\mu\text{m}$. There are 11 such objects in our surveyed area, and these are listed in Table 6. Interestingly all 11 are located in the heart of either Cluster A or B. Additionally, as shown in Table 6 all were detected in some or all IRAC bands. Their energy distributions are all consistent with a designation of Class I even though they are not included in Figure 7 since they were not detected in the 2MASS survey. In fact, several of these objects are strongly enough peaked in the far-infrared that they have energy distributions consistent with some nominal Class 0 sources despite the fact that all were detected with IRAC. The class status of these will be discussed further using mm data by Enoch et al. (2007, in prep.). Figure 10 shows the SED’s for the two most embedded objects from Table 6. Each of these appears to be associated with an outflow in its respective cluster, and both have very similar SED’s that differ only in their absolute flux level by a factor of ~ 10 .

Table 6 shows also that the most embedded object in Cluster B (whose SED is shown in Figure 10) was not selected as a YSO by Harvey et al. (2007). The reason is that the flux at $3.6\mu\text{m}$ was too faint to meet the selection criteria of that study. The area within $15''$ of that source contains two other extracted compact sources in the c2d data set. The positions and photometry for all three are shown in Table 7 and an image of the area is shown in Figure 11. Although the source density is quite high, the $70\mu\text{m}$ contours shown in the figure are clearly centered on the northernmost source, “C”. Source “B” is a slightly extended source that may represent a separate exciting object or may just be the location of the most visible jet emission that has been discussed briefly by Harvey et al. (2007) in this region. Source “A” is a faint, but very red object about $6''$ to the west of source “C” and appears to be a point-like object in the images.

Figure 4 shows clearly that Cluster A and B are located in the highest extinction parts of the cloud. Therefore the lack of detection of the objects in Table 6 at $1 - 2.3\mu\text{m}$ may be due at least partly to the extinction of the cloud material in which they are embedded in addition to individual circumstellar material. Although the nominal extinction values in these areas range up to $A_v \sim 35 - 40$, the fact that these values result from smoothing over 90 arcseconds of the stellar distribution means that they probably underestimate the extinction in the most extreme regions. This association of the coldest objects with the highest extinction regions is similar to the correlation seen by Enoch et al. (2007) between extinction and location of dense mm cores.

4. Summary

We have described the basic observational characteristics of the c2d MIPS observations of the Serpens Cloud. In a 1.5 deg^2 area we have found 250 YSO candidates on the basis of the K_s -[24] color. An additional 11 objects can be identified on the basis of their 24 and $70\mu\text{m}$ fluxes and lack of detection by 2MASS. All of these YSO candidates will be discussed in more detail in a companion paper (Harvey et al. 2007). All the most embedded objects are found in the central area of the two main clusters of YSO's previously identified in Serpens. The images and source catalogs derived from these data are all available on the SSC website, <http://ssc.spitzer.caltech.edu/legacy/all.html>.

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Table 1. Summary of Observations

Region	AOR	Time-Date (UT)	l ^a (deg)	b ^a (deg)
Serpens	5713408	2004-04-05 23:40	31.5	5.4
	5713920	2004-04-06 04:05	31.5	5.3
	5713664	2004-04-06 00:22	31.6	5.2
	5714176	2004-04-06 04:48	30.6	5.1
Off Cloud	5716736	2004-04-06 01:26	35.2	4.4
	5716992	2004-04-06 05:52	35.2	4.3

^a l and b are listed for the center of the 24 μm AOR.

Table 2. Serpens Cloud Sky Coverage

Region	24 μm (deg ²)	70 μm (deg ²)	160 μm (deg ²)
Serpens	1.81	1.57	1.49
Off-Cloud	0.47	0.36	0.41

Table 3. 160 μ m Point Sources

RA (J200)	Dec (J200)	Flux (mJy)	Comment	YSO# ^a
18 29 32.3	+01 18 56	24000	Single 24/70 μ m Source	104
18 29 52.9	+00 36 09	18200	Cluster of four 24 μ m Sources	
18 29 16.7	+00 18 20	10000	Single 24/70 μ m Source	88
18 28 15.7	−00 03 11	6070	Cluster of four 24 μ m Sources	

^aYSO number from Harvey et al. (2007).

Table 4. Serpens Cloud Detection Statistics

Wavelength(s)	Source Number
$24\mu\text{m} > 3\sigma$	2635
$24\mu\text{m} > 5\sigma$	1494
$70\mu\text{m} > 3\sigma$	97
$70\mu\text{m} > 5\sigma$	88
$24 \text{ \& } 70\mu\text{m} > 5\sigma$	75
$24\mu\text{m} \text{ \& } 2\text{MASS } K_s > 5\sigma$	1085
$24\mu\text{m} \text{ \& any IRAC}$	1040 ^a
$70\mu\text{m} \text{ \& any IRAC}$	77

^aThe greater number of matches between $24\mu\text{m}$ and K_s versus IRAC is due to the smaller area coverage of the IRAC data.

Table 5. Classification based on $K_s-[24]$

Classification	Serpens Source Count ^a
number with $K_s-[24]>2$, $K_s<14$	250
number with $K_s-[24]>2$, $K_s<14$, and Class I $K_s-[24]$ color	15 (6%)
number with $K_s-[24]>2$, $K_s<14$, and “flat” $K_s-[24]$ color	21 (8%)
number with $K_s-[24]>2$, $K_s<14$, and Class II $K_s-[24]$ color	158 (63%)
number with $K_s-[24]>2$, $K_s<14$, and Class III $K_s-[24]$ color	56 (22%)

^aSince a 2MASS detection is required to be included in these statistics, very cold or deeply embedded sources are not present in these counts, e.g. those sources in Table 6.

Table 6. The Most Embedded Objects

Name/Position SSTc2dJ...	YSO # ^a	3.6 μ m (mJy)	4.5 μ m (mJy)	5.8 μ m (mJy)	8.0 μ m (mJy)	24.0 μ m (mJy)	70.0 μ m (mJy)	Associated Source ^b
18285404+0029299	40	5.81 \pm 0.50	27.6 \pm 2.3	44.8 \pm 2.6	56.4 \pm 3.2	918 \pm 85	11100 \pm 1040	D62/66
18285486+0029525	42	1.94 \pm 0.12	10.6 \pm 0.6	20.4 \pm 1.1	30.2 \pm 1.6	765 \pm 70	7250 \pm 675	D65
18290619+0030432	67	8.05 \pm 0.41	45.0 \pm 2.8	93.9 \pm 4.8	129 \pm 7	1320 \pm 139	7240 \pm 713	D90
18290675+0030343	68	3.27 \pm 0.21	11.7 \pm 0.7	14.9 \pm 0.8	20.7 \pm 1.2	1000 \pm 105	11400 \pm 1180	D94
18290906+0031323		< 0.12	0.29 \pm 0.03	0.40 \pm 0.09	0.31 \pm 0.08	64.6 \pm 6.0	6380 \pm 611	D101
18294810+0116449	135	1.96 \pm 0.10	6.98 \pm 0.42	12.1 \pm 0.6	16.7 \pm 0.8	219 \pm 21	14900 \pm 1420	K241, SMM9
18294963+0115219	141	0.85 \pm 0.08	2.64 \pm 0.27	2.32 \pm 0.28	3.54 \pm 0.31	1180 \pm 117	82800 \pm 7810	K258a, SMM1
18295219+0115478	150	7.38 \pm 0.41	33.0 \pm 2.1	41.3 \pm 2.2	40.0 \pm 2.6	1640 \pm 154	15200 \pm 1420	K270, SMM10
18295285+0114560	155	8.65 \pm 0.44	34.6 \pm 1.8	72.0 \pm 3.4	110 \pm 5	1040 \pm 96	5570 \pm 523	K276
18295927+0114016	195	2.72 \pm 0.28	5.76 \pm 0.44	7.78 \pm 1.16	36.0 \pm 5.4	109 \pm 19	12200 \pm 1160	SMM3
18295992+0113116	198	2.77 \pm 0.16	29.5 \pm 1.5	103 \pm 4	199 \pm 10	2620 \pm 249	6830 \pm 675	K331

^aIdentifying number from YSO table in Harvey et al. (2007).

^bReferences are: D: (Djupvik et al. 2006), K: (Kaas et al. 2004), SMM: Davis et al. (1999).

Table 7. Sources Marked In Figure 11

Marker	Name/Position	3.6 μm	4.5 μm	5.8 μm	8.0 μm	24.0 μm	70.0 μm
SSTc2dJ...	(mJy)	(mJy)	(mJy)	(mJy)	(mJy)	(mJy)	
A ^a	18290904+0031280	0.95 \pm 0.11	2.78 \pm 0.23	2.92 \pm 0.24	5.03 \pm 0.40	14.0 \pm 1.9	...
B	18290864+0031305	0.06 \pm 0.03	0.32 \pm 0.02	0.47 \pm 0.05	0.62 \pm 0.07	36.2 \pm 3.4	...
C	18290906+0031323	< 0.12	0.29 \pm 0.03	0.40 \pm 0.09	0.31 \pm 0.08	64.6 \pm 6.0	6380 \pm 611

^aThis is YSO # 75 in Harvey et al. (2007).

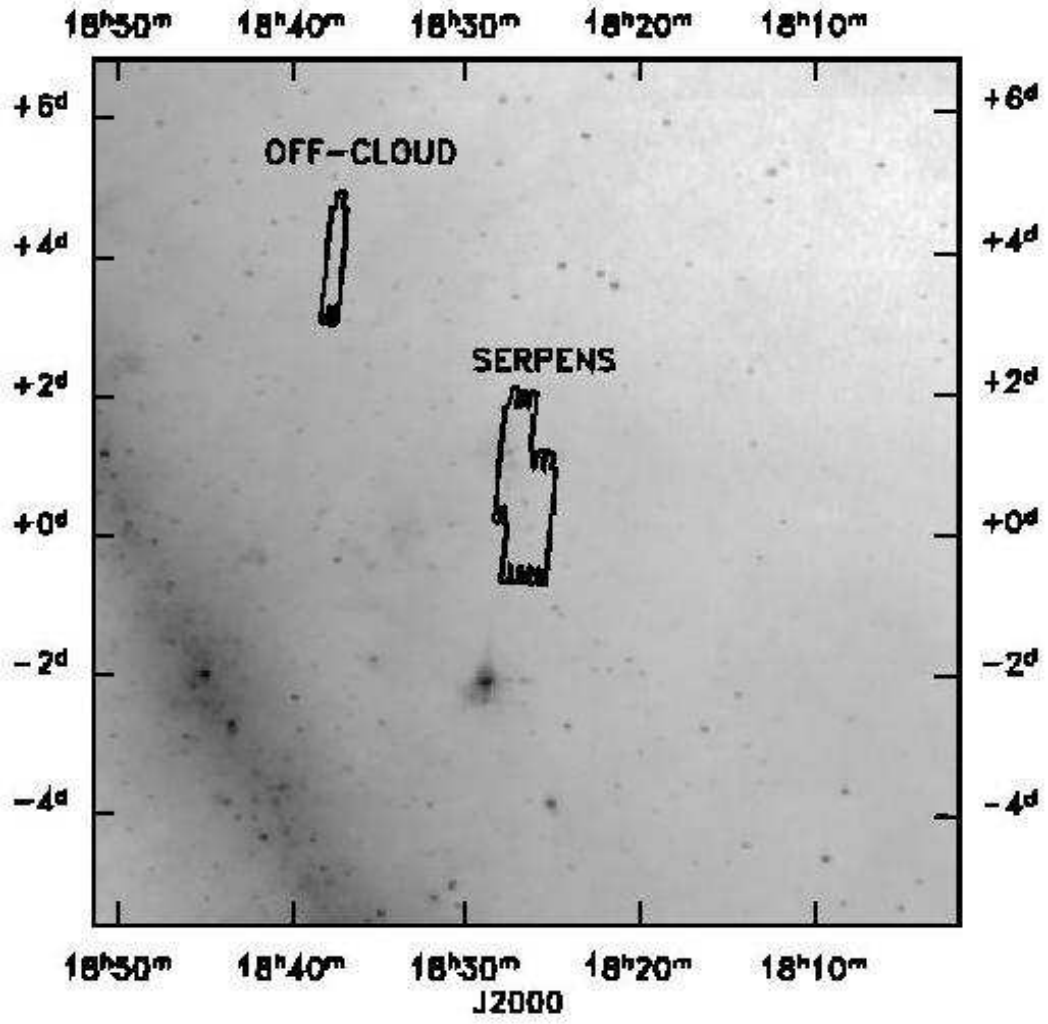


Fig. 1.— IRAS 25 μm map showing the observed $c2d$ regions in the Serpens cloud, both the star-forming region marked “SERPENS” and the low-extinction “OFF-CLOUD” area.

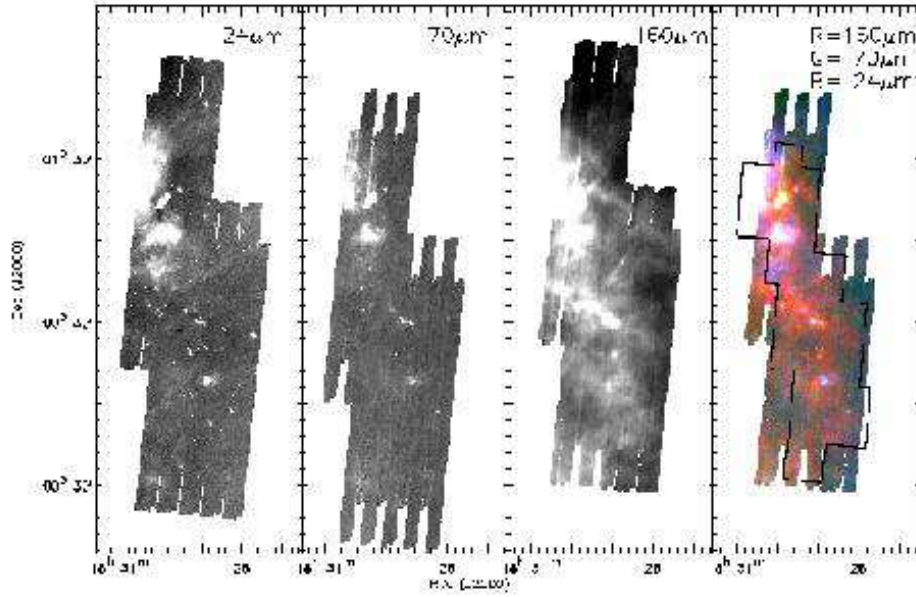


Fig. 2.— Registered Serpens 24μm, 70μm and 160μm images of the *c2d* MIPS region. The color image is a composite of all three bands, and includes only the 1.27 square degree area where data are available for each of the three bands. Colors represent red:160μm green:70μm and blue:24μm. The black outline shows the region where 4 bands of IRAC data were

Fig. 3.— Completeness test at $24\mu\text{m}$. The upper solid line shows the measured completeness fraction for artificial sources inserted into the observed $24\mu\text{m}$ mosaic image of Serpens as a function of magnitude. The slightly higher dash-dot line shows the completeness fraction for sources inserted into an artificial image with no real sources but with a noise level equal to that in the observed data. The lower solid line (mostly equal to zero) shows the fraction of “unreliable” sources, i.e. sources extracted which were not real.

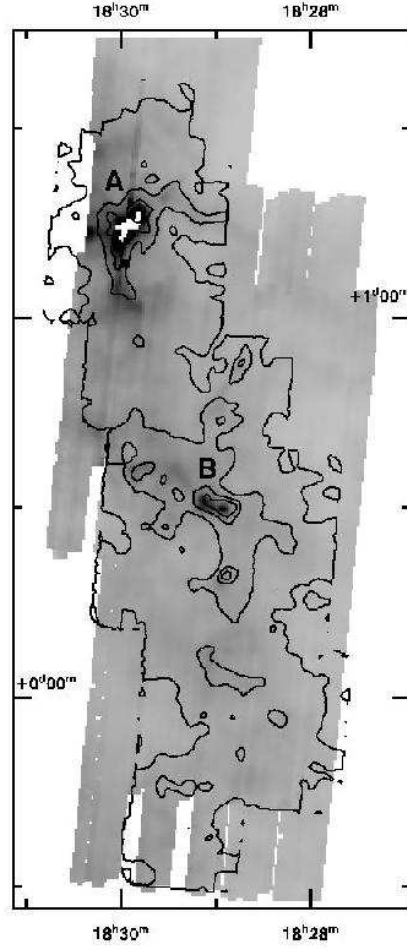


Fig. 4.— Contours of A_v at levels of 5,10,20,30 mag determined from 2MASS and Spitzer c2d IRAC data are overlaid on the Serpens 160 μ m image. The visual extinction and 160 μ m emission are quite well correlated. The locations of Cluster A and B are indicated.

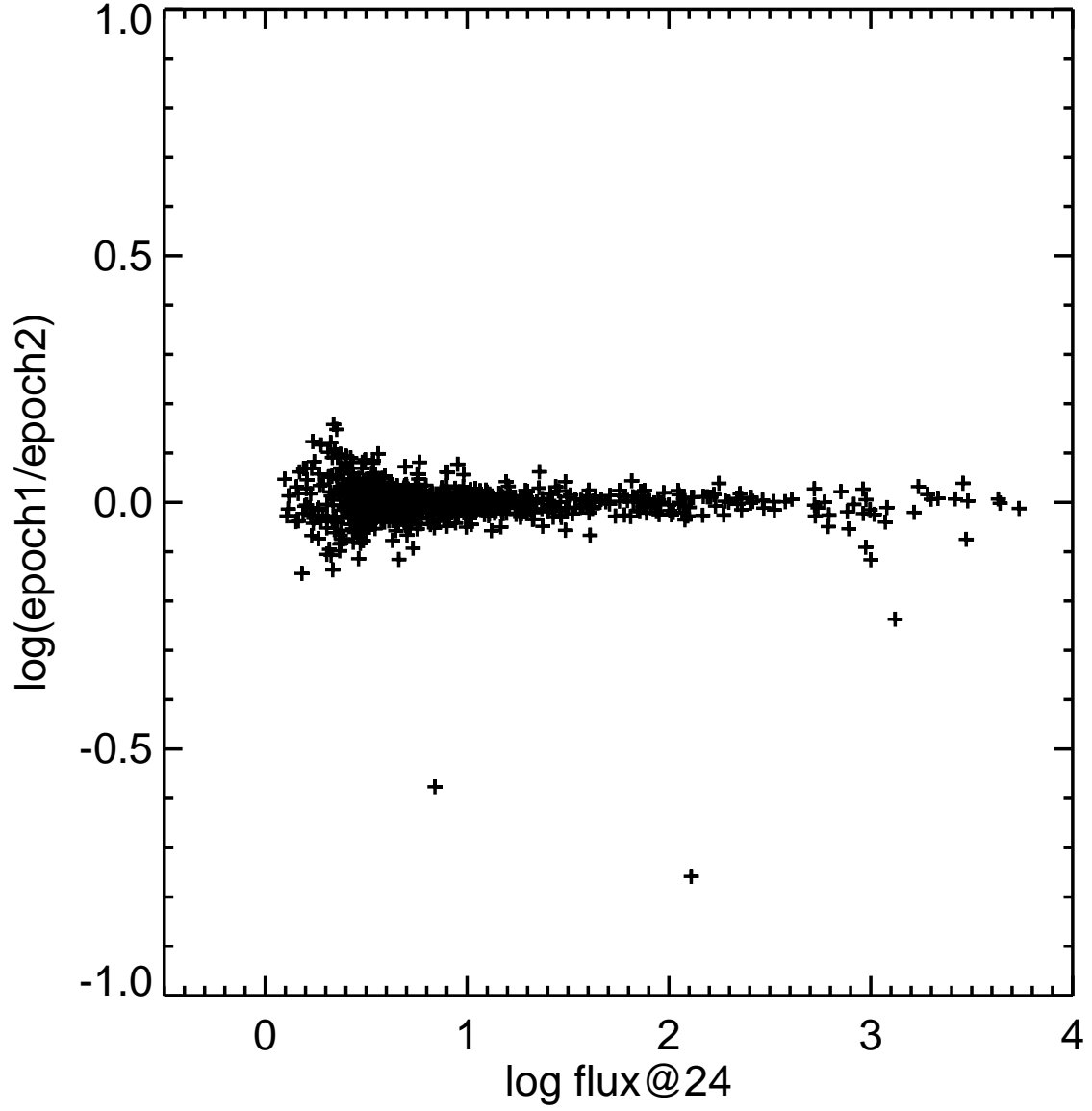


Fig. 5.— A search for time variability in the Serpens $24\mu\text{m}$ data; plot of log flux ratio of epoch1 to epoch2 versus log flux density (mJy) for the combined epoch data. There is no verifiable time variable source in the cloud based on these data.

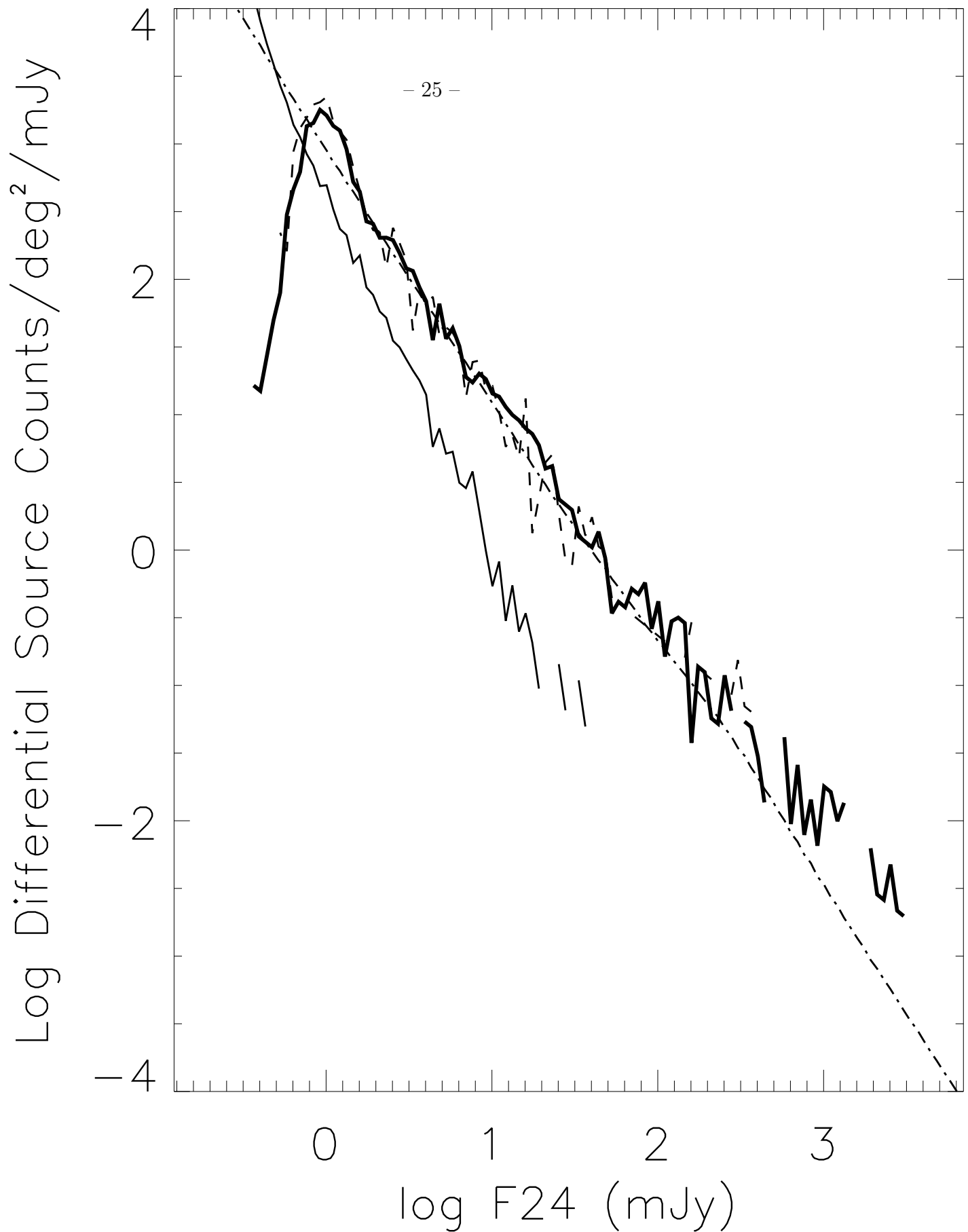


Fig. 6.— $24\mu\text{m}$ source counts in the Serpens MIPS field (dark line), and off-cloud region (dashed line). SWIRE galaxy counts (thin line) fall below the Serpens data at our flux limit of 1 mJy. The predicted source counts from the Wainscoat model at $25\mu\text{m}$ (Wainscoat et al. 1992) are shown by the dot-dash line.

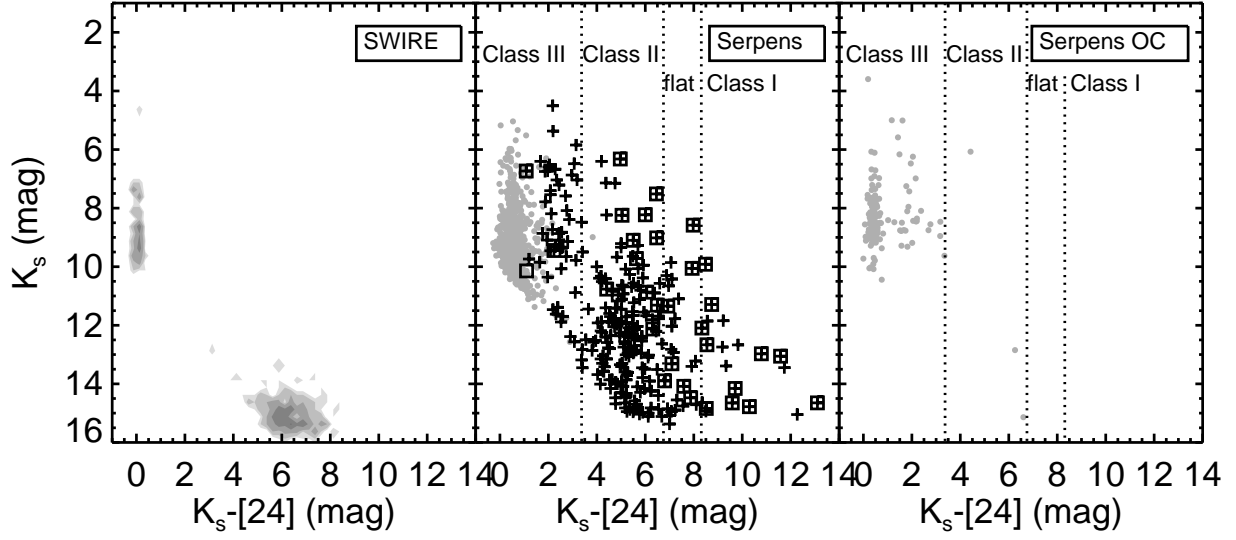


Fig. 7.— Color-magnitude diagram for K_s vs. $K_s - [24]$ for objects in SWIRE (left) and Serpens (center) and off-cloud region (right). The SWIRE counts are shown as a surface density with darker implying higher density. Objects in SWIRE are expected to be mostly galaxies (objects with $K_s \gtrsim 14$) or stellar photospheres (objects with $K_s - [24] \lesssim 1$). For the Serpens and off-cloud plots, filled gray circles are objects with SEDs resembling photospheres, and plus signs are the remaining objects. An additional box around a point denotes that it was also detected at $70\mu\text{m}$. Objects that are candidate young objects have colors unlike those objects found in SWIRE, e.g., $K_s \lesssim 14$ and $K_s - [24] \gtrsim 1$. Dashed lines denote the divisions between Class I, flat, Class II, and Class III objects; to omit foreground and background stars, we have further imposed a $K_s - [24] > 2$ requirement on our Class III objects (see text).

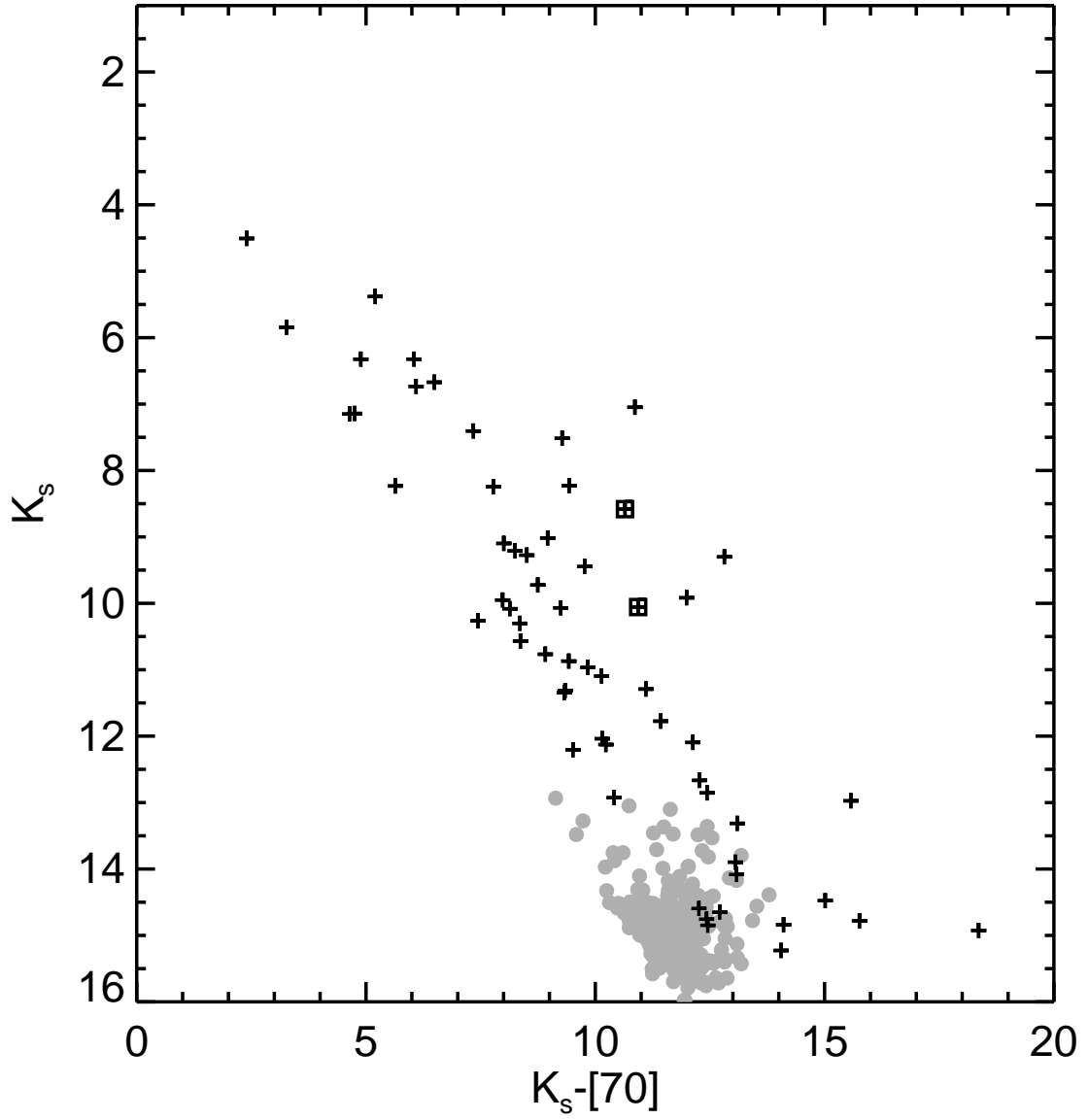


Fig. 8.— Color-magnitude diagram of K_s vs. $K_s - [70]$ for Serpens (crosses) with data from the full SWIRE survey (grey dots) included for comparison.

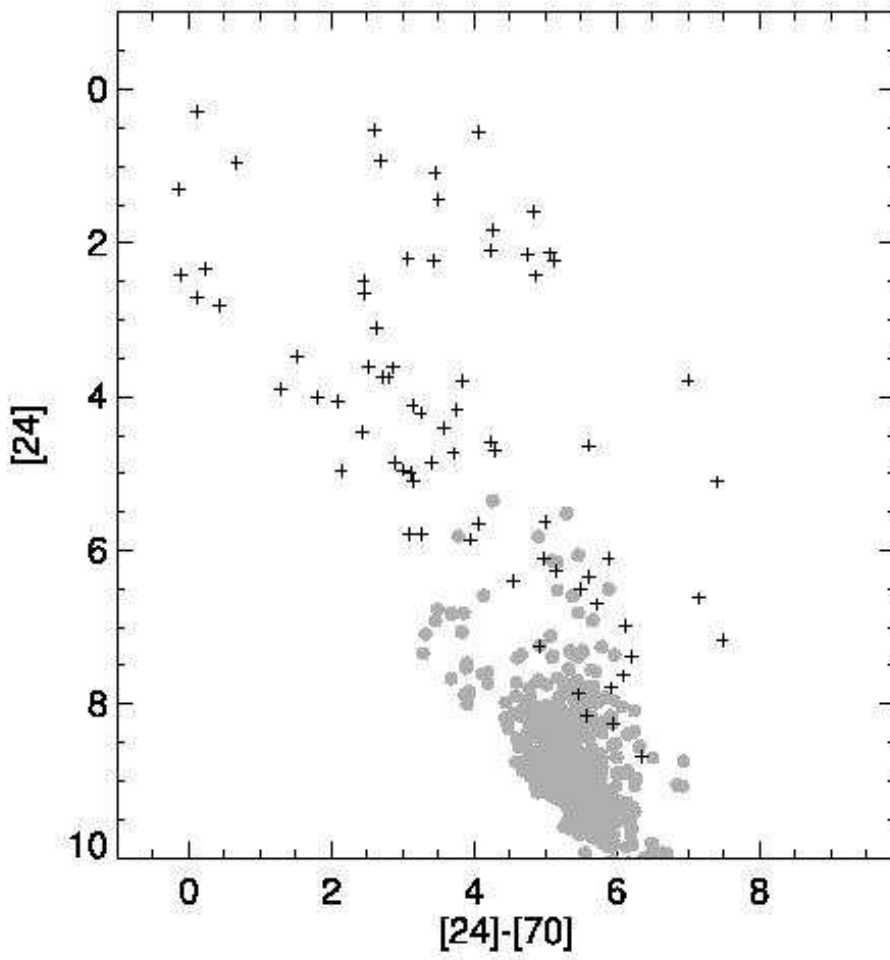


Fig. 9.— Color-magnitude diagram of $[24]$ vs. $[24] - [70]$ for Serpens (crosses) with data from the full SWIRE survey (grey dots) included for comparison.

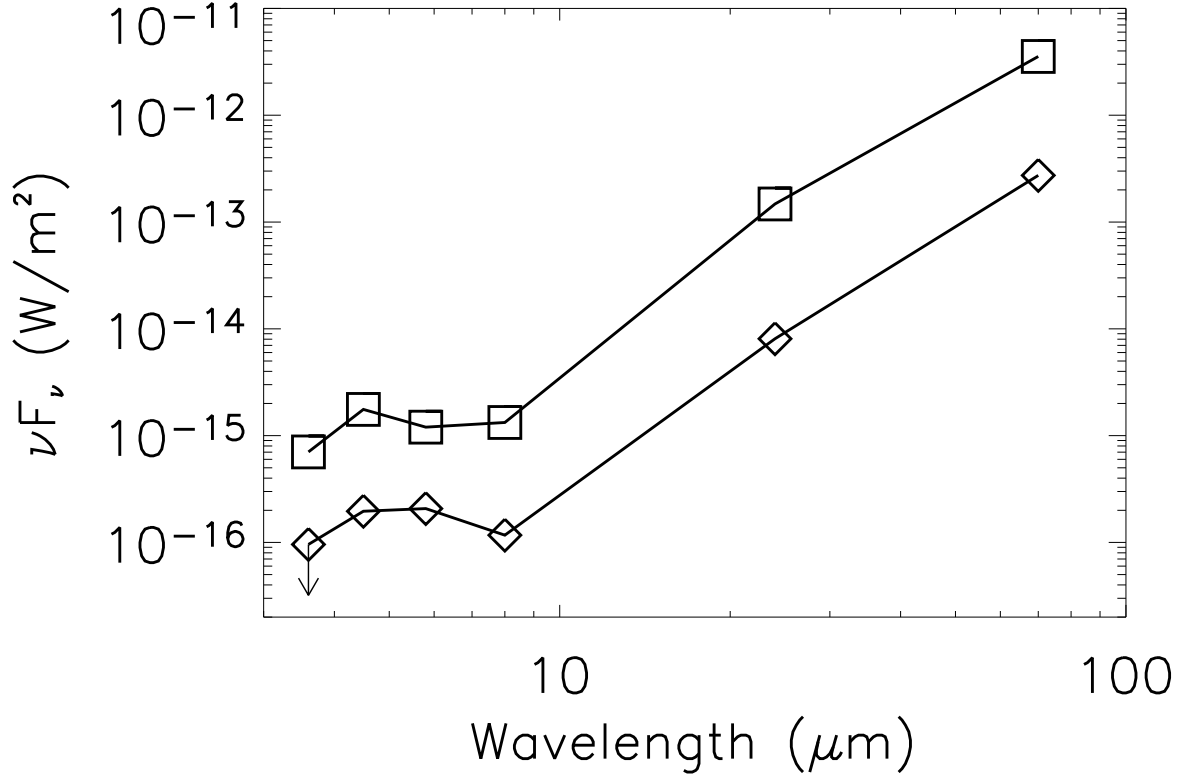


Fig. 10.— Spectral energy distribution for the two most embedded sources in Table 6, one in Cluster A (open squares, SSTc2dJ1829463+0115219) and one in Cluster B (open diamonds, SSTc2dJ18290906+0031323, source “C” in Table 7), both of which appear to be associated with outflows.

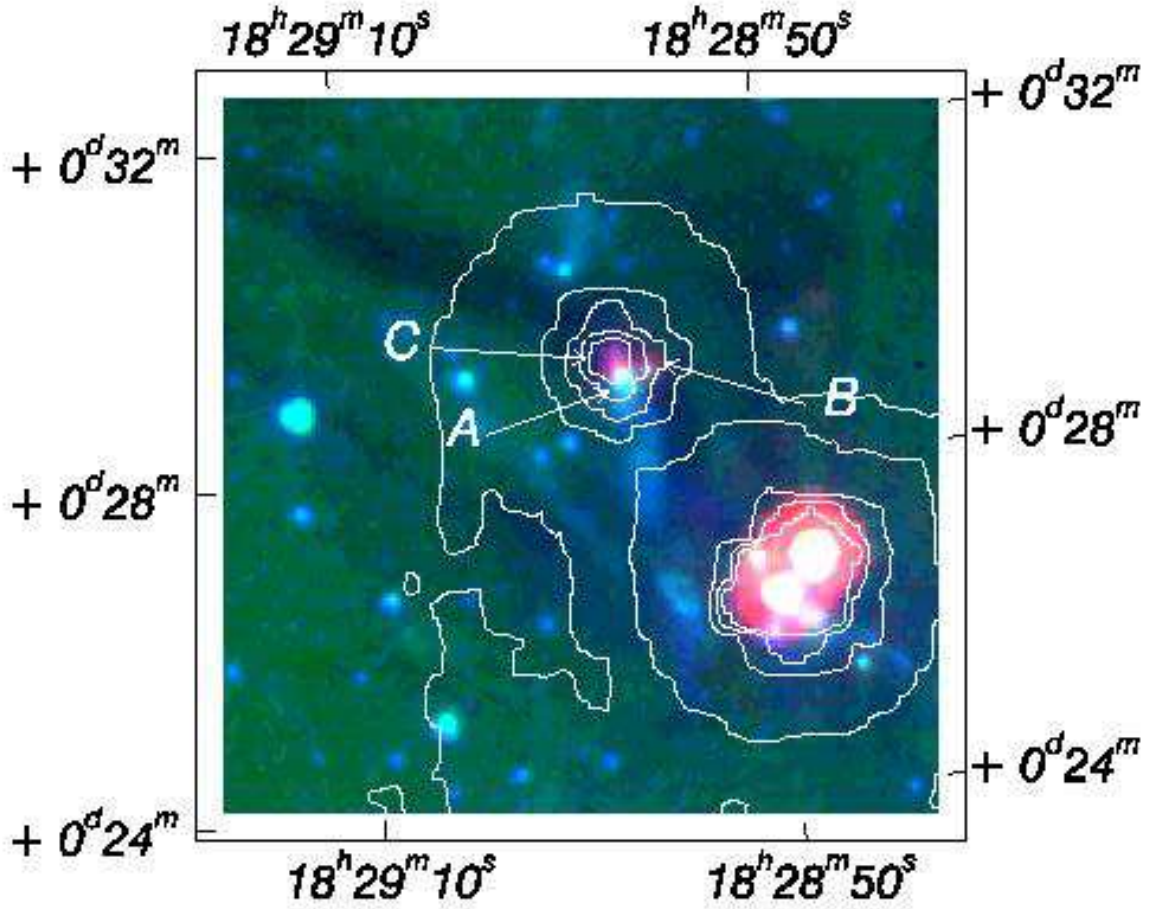


Fig. 11.— Three color image of the eastern end of Cluster B where the most embedded source, C, is located. This is the likely exciting source for an HH-like outflow visible in the IRAC data. The color scheme is: blue/4.5μm, green/8.0μm, and red/24μm. The contours of 70μm emission are also superimposed with levels at 40, 80, 160, 240, and 320 MJy/sr. Also shown are the positions of two other compact sources extracted from the images in this region. The letters correspond to positions/fluxes in Table 7.