

BoloCalc User Manual

Version 1.0

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This document is still under construction and will be updated frequently

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1 Quick Start

This section covers how to get BoloCalc up and running.

1.1 System Requirements

BoloCalc runs on a Linux and has the following dependencies:

1. Python v2.7.2
2. Numpy v1.8 or above

1.2 Code location

BoloCalc is stored on Charles Hill's GitHub account `chill90/` within the repository `chill90/BoloCalc`. To clone a local copy of the calculator, make sure that you have Git installed on your machine (I've been using Git version 2.13.5) and run

```
1 git clone https://github.com/chill90/BoloCalc.git
```

All executables are stored at in the top-most directory

```
1 ~/BoloCalc/
```

and all libraries are stored in the `src/` directory

```
1 ~/BoloCalc/src/
```

1.3 Running mappingSpeed.py

The primary BoloCalc executable is called `mappingSpeed.py`, which calculates optical power, noise-equivalent power (NEP) contributions from photon noise, bolometer thermal carrier noise, readout noise, per-detector noise-equivalent temperature (NET), array NET, mapping speed, and map depth, given some instrument configuration. `mappingSpeed.py` takes as input an experiment configuration directory, which is stored in `BoloCalc/Experiments` and has the following structure:

```
1 CHillCalc/  
2 |-- Experiements/  
3 |   |-- [Experiment1]/  
4 |   |   |-- [ExperimentDesign1]/  
5 |   |   |-- [ExperimentDesign2]/  
6 |   |   |-- ...  
7 |   |-- [Expierment2]/  
8 |   |-- ...
```

where the square brackets indicate the directories that can be named by the user.

To calculate the sensitivity an experiment, pass its `[ExperimentDesign]` directory to `mappingSpeed.py`. BoloCalc comes with an example experiment that can be used to get started:

```
1 BoloCalc$ python2.7 mappingSpeed.py Experiments/ExampleExperiment/V0/
```

By default, `mappingSpeed.py` will calculate one experiment realization, which should be shown by the status bar that appears upon execution. The simulation parameters can be tuned in the `BoloCalc/config/simulationInputs.txt`

For more information about the simulation input parameters, see Section ??.

After it has finished running, `mappingSpeed.py` generates several tables that display the calculated quantities for the simulated instrument. These tables are stored in three files named `sensitivity.txt`, with each file located at a different level of the experiment directory structure and containing different information:

```

***** Simulation Parameters *****
#
#-----
Parameter | Value | Description
#-----
Multiprocess | False | Whether or not to use multiprocessing
#-----
Cores | 1 | Number of cores to use for parallel processing. Positive integer
#-----
Verbosity | 1 | Verbosity for printing logging messages to standard output. 0 for low, 1 for moderate, 2 for high
#-----
Experiments | 1 | Number of Experiment MC realizations. Positive integer
#-----
Observations | 1 | Number of Observation MC realizations per experiment realization. Positive integer.
#-----
Detectors | 1 | Number of Detector MC realizations per observation realization. Positive Integer.
#-----
Resolution | 0.100 | Spectral resolution for integration over bands and spectra in GHz. Positive floating point value.
#-----
Foregrounds | False | Include Foregrounds? True or False
#-----
Correlations | True | Include white noise correlations? True or False
#-----
#

```

Figure 1: Example of simulation input file `simulationInputs.txt`

- `Experiments/ExampleExperiment/V0/sensitivity.txt`: contains the sensitivity combined by observation frequency of all telescopes in this experiment.
- `Experiments/ExampleExperiment/V0/Tel/sensitivity.txt`: contains the sensitivity combined by observation frequency of all cameras in this telescope.
- `Experiement/ExampleExperiment/V0/Tel/Cam/sensitivity.txt`: contains the sensitivity for each observation frequency in this camera.

For more information about the contents of the `sensitivity.txt` files, see Section ??.

1.4 Running `mappingSpeed_vary.py`

The other BoloCalc executable is called `mappingSpeed_vary.py` and essentially calculates mapping speed for a set of instrument configurations determined by a set of specified parameters varied over a linearly-spaced array. To run `mappingSpeed_vary.py`, pass it the instrument configuration at the command line

```
1 BoloCalc$ python2.7 mappingSpeed_vary.py Experiments/ExampleExp/V0/
```

The input parameter variations are stored in `BoloCalc/config/paramsToVary.txt`. An example of the `paramsToVary.txt` is shown below:

# Telescope	Camera	Channel	Optic	Parameter	Minimum	Maximum	Step Size
#							
Tel	Cam	1	Primary	Spillover	0.000	0.050	0.010
#							
Tel	Cam	1	Aperture	Temperature	1.000	5.000	1.000

Parameters are defined by row and have the following descriptors, which are organized by column:

- Column 1: Telescope in which the parameter is defined
 - Leave blank if the parameter is defined in `config/foregrounds.txt`
- Column 2: Camera in which the parameter is defined
 - Leave blank if the parameter is defined in `Tel/config/telescope.txt`
- Column 3: Channel for which the parameter is defined

- Leave blank if the parameter is defined in `Tel/config/telescope.txt` or `Cam/config/camera.txt`
- Column 4: Optic for which the parameter is defined
 - Leave blank **unless** the parameter is defined in `Cam/config/optics.txt`
- Column 5: Parameter name
- Column 6: Minimum parameter value to be calculated
- Column 7: Maximum parameter value to be calculated
- Column 8: Step size with which the parameter will be calculated over the range [Minimum, Maximum]

By default upon download, `mappingSpeed_vary.py` will calculate one experiment realization for each and every parameter combination (see Section ?? for details on simulation inputs). The status bar will show progression towards calculating all combinations of parameters defined over the ranges. `mappingSpeed_vary.py` saves three quantities¹ to files within the directory `BoloCalc/Experiments/ExampleExp/V0/paramsToVary/`

- Optical power P_{opt} : `mappingSpeedVary_Popt_[Tel]_[Cam]_[Channel]_[Param].txt`
- Photon NEP NEP_{ph} : `mappingSpeedVary_NEPph_[Tel]_[Cam]_[Channel]_[Param].txt`
- Array NET NET_{arr} : `mappingSpeedVary_NETarr_[Tel]_[Cam]_[Channel]_[Param].txt`

where `[Tel]_[Cam]_[Channel]_[Param]` is a string that contains the specifier for every parameter varied in the calculation. For example, if we looped over the following two parameters

- Telescope “Tel,” camera “Cam,” channel “1,” optic “Mirror,” Parameter “Spillover”
- Telescope “Tel,” camera “Cam,” channel “1,” Parameter “Pixel Size”

the P_{opt} output file would have the filename

```
1 mappingSpeedVary_Popt_Tel_Cam_1_Mirror_Spillover_Tel_Cam_1_PixelSize.txt
```

For information about the contents of the `mappingSpeed_vary.py` output file, see Section ??.

1.5 Contact and Citation

Please send any comments, requests, questions, or bugs to Charles Hill at chill@lbl.gov. If you use this code for published results, please cite the SPIE proceeding from 2018 “BoloCalc: a sensitivity calculator for the design of Simons Observatory,” which can be found at <https://arxiv.org/abs/1806.04316>.

2 Sensitivity Calculation

The primary function of BoloCalc is to import low-level instrument parameters and use them to estimate the instrument’s noise-equivalent CMB temperature (NET). This calculation is outlined by the following steps:

1. Collect input parameters and construct an instrument model
2. Calculate per-detector noise-equivalent power (NEP) due to contributions from
 - Photon shot and wave noise
 - Bolometer thermal carrier noise
 - Readout noise
3. Convert NEP to NET
4. Calculate the array-averaged NET, where the noise contributions from each detector are inverse-variance weighted to estimate the instantaneous sensitivity of the full camera

¹More quantities can be added to this list upon request, or as needed.

5. Calculate mapping speed (MS), a quantification of instrument noise in the power spectrum domain

We detail the assumptions and inputs for the calculation in the following subsections.

2.1 Optical power

BoloCalc assumes an array of single-moded bolometers² within an instrument that is stationary in time. The propagation of optical power from the sky to the focal plane is represented by a one-dimensional chain of blackbody absorbers/emitters in thermal equilibrium. The power deposited on the detectors is then an analytic integral over the summation of each optical element's Planck spectra modified by its frequency-dependent efficiency and emissivity. Explicitly, the optical power on a detector is given as

$$P_{\text{opt}} = \int_0^\infty \left[\sum_{i=1}^{N_{\text{elem}}} p_i(\nu) \right] B(\nu) d\nu, \quad (1)$$

where ν is frequency, $p_i(\nu)$ is the power spectral density of optical element i referred to the detector input, the summation contains all N_{elem} optical elements in the sky/telescope/camera and runs from the CMB to the focal plane, and $B(\nu)$ is the detector bandpass.

The power spectral density $p_i(\nu)$ for optical element i is determined by its blackbody temperature T_i , the transmission efficiency of all optics between it and the focal plane $[\eta_{i+1}(\nu), \dots, \eta_{N_{\text{elem}}}(\nu)]$, its emissivity $\epsilon_i(\nu)$, its spillover coefficient β_i , the effective temperature by which its spilled power is absorbed $T_{\beta;i}$, its scattering coefficient δ_i , and the effective temperature by which its scattered power is absorbed $T_{\delta;i}$

$$p_i(T_i, [\eta_{i+1}(\nu), \dots, \eta_{N_{\text{elem}}}(\nu)], \epsilon_i(\nu), \beta_i, T_{\beta;i}, \delta_i, T_{\delta;i}, \nu) = \prod_{j=i+1}^{N_{\text{elem}}} \eta_j(\nu) [\epsilon_i(\nu) S(T_i, \nu) + \beta_i S(T_{\beta;i}, \nu) + \delta_i S(T_{\delta;i}, \nu)] . \quad (2)$$

The power spectral density function $S(T, \nu)$ of the emitted and scattered power from each element is given by the Planck spectral density³ for a diffraction-limited, single-moded polarimeter

$$S(T, \nu) = \frac{h\nu}{\exp\left[\frac{h\nu}{k_B T}\right] - 1} . \quad (3)$$

2.2 Photon noise

Photon noise in bolometric detection is the result of fluctuations in the arrival times of photons at the absorbing element

$$\text{NEP}_{\text{ph}} = \sqrt{2 \int_0^\infty \left[h\nu \sum_{i=1}^{N_{\text{elem}}} p_i(\nu) + \left(\sum_{i=1}^{N_{\text{elem}}} p_i(\nu) \right)^2 \right] B(\nu) d\nu} . \quad (4)$$

There are two contributions to NEP_{ph} . The first term represents shot noise NEP_{shot} , which dominates when the photon occupation number $\ll 1$ (e.g. optical wavelengths) and is $\propto \sqrt{P_{\text{opt}}}$.

²If there is a demand, **BoloCalc** can be expanded to include other detector architectures, such as KIDs.

³**BoloCalc** can be expanded to incorporate non-thermal spectral densities if there is a demand for them.

The second term represents wave noise NEP_{wave} , which dominates when the photon occupation number is $\gg 1$ (e.g. radio wavelengths) and is $\propto P_{\text{opt}}$. For ground-based experiments, the photon occupation number at ~ 100 GHz is ~ 1 , and therefore a careful handling of both terms is necessary for an accurate NET estimate.

2.3 Bolometer thermal carrier noise

Thermal carrier noise in bolometers arises due to fluctuations in heat flow between the absorbing element and the bath to which it is weakly connected

$$\text{NEP}_g = \sqrt{4 k_B F_{\text{link}} T_{\text{oper}}^2 G} , \quad (5)$$

where T_{oper} is the operating temperature of the bolometer, G is the thermal conductance from the absorbing element to the bath, and F_{link} is a numerical factor that depends on the link's thermal conduction index n . F_{link} can be theoretically calculated via the equation

$$F_{\text{link}} = \frac{n+1}{2n+3} \frac{1 - (T_{\text{bath}}/T_{\text{oper}})^{2n+3}}{1 - (T_{\text{bath}}/T_{\text{oper}})^{n+1}} , \quad (6)$$

where T_{bath} is the bath temperature. However, NEP_g can vary depending on the specifics of the bolometer geometry, composition, and fabrication. For example, transition-edge sensors (TES's) have known pathological noise sources, such as flux flow noise and non-equilibrium Johnson noise, that increase the measured NEP_g beyond that of Mather's theoretical prediction. Therefore, BoloCalc provides an option for F_{link} to be set independent of T_{bath} and n , allowing NEP_g to be tuned phenomenologically.

Thermal conductance can be parameterized in terms of n , T_{bath} , and the bolometer saturation power P_{sat} —or the power conducted from the bolometer to the bath—as

$$G = P_{\text{sat}}(n+1) \frac{T_{\text{oper}}^n}{T_{\text{oper}}^{n+1} - T_{\text{bath}}^{n+1}} . \quad (7)$$

Therefore, $\text{NEP}_g \propto \sqrt{P_{\text{sat}}}$, making the tuning of saturation power important to optimizing detector sensitivity.

2.4 Readout noise

Modern CMB detectors are low-impedance, voltage-biased bolometers read out using superconducting quantum interference device (SQUID) transimpedance amplifiers. Therefore, amplifier-induced readout noise can be modeled as a noise-equivalent current (NEI), referred to the power at the detector by the inverse of the detector responsivity. For a voltage-biased bolometer operating with negative feedback and high loop gain $\mathcal{L} \gg 1$, responsivity is $\approx 1/V_{\text{bias}}$, and readout NEP is given by

$$\text{NEP}_{\text{read}} = \sqrt{P_{\text{elec}} R_{\text{bolo}}} \times \text{NEI} , \quad (8)$$

where the bias power $P_{\text{elec}} = P_{\text{sat}} - P_{\text{opt}}$ and R_{bolo} is the bolometer operating resistance.

2.5 Noise-equivalent CMB temperature

A CMB bolometer is built to measure fluctuations in incident power due to fluctuations in CMB temperature. Therefore, it is useful to convert bolometer NEP into a noise-equivalent CMB temperature (NET). The total noise in the bolometer output is the quadrature sum of photon noise, thermal carrier noise, and readout noise, and the conversion to NET is given by

$$\text{NET}_{\text{det}} = M \frac{\sqrt{\text{NEP}_{\text{ph}}^2 + \text{NEP}_{\text{g}}^2 + \text{NEP}_{\text{read}}^2}}{\sqrt{2} (dP/dT_{\text{CMB}})}, \quad (9)$$

where the M is a “margin” applied to the expected per-detector NET, and the $\sqrt{2}$ arises due to a unit conversion from output bandwidth $1/\sqrt{\text{Hz}}$ to integration time $\sqrt{\text{s}}$. The conversion factor from optical power to CMB temperature is defined as

$$dP/dT_{\text{CMB}} = \xi \int_0^\infty \left[\prod_{i=1}^{N_{\text{elem}}} \eta_i(\nu) \frac{1}{k_{\text{B}}} \left(\frac{h\nu}{T_{\text{CMB}} (\exp[h\nu/k_{\text{B}}T_{\text{CMB}}] - 1)} \right)^2 \exp[h\nu/k_{\text{B}}T_{\text{CMB}}] \right] B(\nu) d\nu, \quad (10)$$

and has units of $\text{W}/\text{K}_{\text{CMB}}$. ξ is an optical coupling factor that quantifies SNR degradation associated with image degradation at the focal plane.

When reconstructing the sky during analysis, data from each detector are co-added in the map domain to improve signal-to-noise. To quantify this SNR increase in the time domain, we define “array NET” as the inverse-variance-weighted average of the NETs of all yielded detectors within the camera

$$\text{NET}_{\text{arr}} = \frac{\text{NET}_{\text{det}}}{\sqrt{Y N_{\text{det}}}} \Gamma, \quad (11)$$

where N_{det} is the number of deployed bolometers, Y is the yield, and Γ is a factor ≥ 1 that quantifies the degree to which white noise is correlated between detector pixels on the focal plane.

2.6 Map Depth

Finally, array NET—white noise in the time domain—is converted from units of $\text{K}\sqrt{\text{s}}$ to units of K-arcmin —white noise in the map domain—using the equation

$$\sigma_S = \sqrt{\frac{4\pi f_{\text{sky}} \text{NET}_{\text{arr}}^2}{\eta_{\text{obs}} t_{\text{obs}}}} \left(\frac{10800 \text{ arcmin}}{\pi} \right), \quad (12)$$

where f_{sky} the fraction of sky observed, t_{obs} is the observation time of the experiment, and η_{obs} is the observation efficiency.

2.7 Mapping speed

While NET is useful for estimating noise in the time and map domain, mapping speed (MS) is useful for describing instrument performance in the power spectrum domain

$$\text{MS} = \frac{1}{\text{NET}_{\text{arr}}^2} \quad (13)$$

and has units of $\text{K}^{-2} \text{s}^{-1}$.

2.8 Auxiliary Calculations

There are a few calculations that are carried out “under the hood” if specific sets of parameters are defined in the input files (for more information about parameter definitions and dependencies, see . Ultimately, as shown in Equation , the calculate needs a temperature, emissivity, and efficiency in order to calculate the power from each element.

2.8.1 Synchrotron Emission

Synchrotron power in units of W/Hz is given by the following equation

$$p_{\text{synch}}(\nu) = A_s \nu^{n_s} \quad (14)$$

where A_s is the amplitude, ν is the observation frequency, and n_s is the spectral index. This quantity is integrated over the observation band when calculating the in-band optical power and signal attenuation referenced to the detector.

2.8.2 Dust Emission

Thermal dust power in units of W/Hz is given by the following equation

$$p_{\text{dust}}(\nu) = A_d \left(\frac{\nu}{\nu_d} \right)^{n_d} S(T_d, \nu) \quad (15)$$

where A_d is the amplitude, ν is the observation frequency, ν_d is the pivot frequency, n_d is the power law index, and T_d is the effective blackbody temperature. This quantity is integrated over the observation band when calculating the in-band optical power and signal attenuation referenced to the detector.

2.8.3 Atmosphere Emission

BoloCalc utilizes atmospheric simulations of the atmosphere at the Atacama and South Pole sites⁴ generated by the AM atmospheric modeling code⁵, which uses data from the MERRA-2 meteorological reanalysis⁶ as input. The output from AM produces results consistent with measured sky loading in existing Atacama experiments. The range of input elevations handled by BoloCalc is from 20–90 deg, and the range of input PWV values is from 0–8 mm.

2.8.4 Cold Stop Spillover Efficiency

Aperture efficiency can be specified explicitly, or, if it is left as “NA,” it can be derived from a combination of other parameters. The equation for aperture efficiency assumes a Gaussian beam which is parameterized by the ratio of the pixel diameter to the beam waist $w_f = D/w_0$

$$\eta_{\text{stop}}(\lambda) = 1 - \exp \left[-\frac{\pi^2}{2} \left(\frac{D}{F\lambda w_f} \right)^2 \right], \quad (16)$$

where D is the pixel diameter, F is the F/# at the focal plane, and λ is the observation wavelength. This quantity is integrated over the observation band when calculating the in-band optical power and signal attenuation referenced to the detector.

2.8.5 Dielectric Absorption/Emission

Absorption in an refractive optical element can be specified explicitly, or, if it is left as “NA,” it can be derived using the following equation

$$\epsilon(\lambda) = 1 - \exp[-2\pi t n \tan \delta / \lambda] \quad (17)$$

⁴More site profiles exist and can be added if requested.

⁵<https://www.cfa.harvard.edu/~spaine/am/>

⁶<https://gmao.gsfc.nasa.gov/reanalysis/MERRA-2/>

where t is the thickness of the substrate, n is the refractive index, $\tan \delta$ is the loss tangent, and λ is the observation wavelength. This quantity is integrated over the observation band when calculating the in-band optical power and signal attenuation referenced to the detector.

2.8.6 Reflector Absorption/Emission

Absorption in a reflective optical element can be specified explicitly, or, if it is left as “NA,” it can be derived using the following equation

$$\epsilon(\nu) = 4 \sqrt{\frac{\pi \nu \mu_0}{\sigma_c}} \frac{1}{Z_0} \quad (18)$$

where ν is the observation frequency, μ_0 is the permeability of free space, $Z_0 = \sqrt{\mu_0/\epsilon_0}$ is the impedance of free space, and σ_c is the reflector conductivity at the reflector’s operating temperature. This quantity is integrated over the observation band when calculating the in-band optical power and signal attenuation referenced to the detector.

2.8.7 Ruze Scattering

Scattering of a reflective or refractive optical element can be specified explicitly, or, if it is left as “NA,” it can be derived using Ruze’s equation

$$\delta(\lambda) = \exp \left[\left(\frac{4\pi\sigma_r}{\lambda} \right)^2 \right] \quad (19)$$

where σ_r is the RMS surface roughness of the optical element and λ is the observation wavelength. This quantity is integrated over the observation band when calculating the in-band optical power and signal attenuation referenced to the detector.

2.8.8 Efficiency

Efficiency can be defined explicitly for any optic or detector by passing a band file (see Section ?? and ?? for more details about band files). However, if a band file is not passed, efficiency is derived using the following equation

$$\eta(\nu) = 1 - r(\nu) - \epsilon(\nu) - \beta - \delta \quad (20)$$

where $r(\nu)$ is the reflection, $\epsilon(\nu)$ is the dielectric absorption, β is the spillover coefficient, and δ is the scattering coefficient. This quantity is integrated over the observation band when calculating the in-band optical power and signal attenuation referenced to the detector.

2.8.9 Pixel Elevation

In order to calculate the expected loading on each detector, we need to know the elevation of its projected detector pixel on the sky. This is calculated as the sum of the telescope boresight elevation, the camera boresight elevation w.r.t. telescope boresight, and the pixel elevation w.r.t. camera boresight. Therefore, the pixel elevation is given by

$$\theta_{\text{elv}} = \theta_{\text{tel}} + \theta_{\text{cam}} + \theta_{\text{pix}} \quad (21)$$

2.8.10 Saturation Power Factor

BoloCalc has an option to set saturation power P_{sat} (see Section 2.3 for more details) as a fraction of the optical power (as defined in Equation 1) if the parameter “Saturation Power” is set to “NA”

$$P_{\text{sat}} = P_{\text{opt}} \times f_{\text{psat}} \quad (22)$$

where f_{psat} is the saturation power factor.

2.8.11 Operating Temperature Fraction

BoloCalc has an option to set the operating temperature T_{oper} (see Section 2.3 for more details) as a fraction of the camera base temperature T_{base} if the parameter “Tc” is set to “NA”

$$T_{\text{oper}} = T_{\text{base}} \times f_{\text{oper}} \quad (23)$$

where f_{oper} is the operating temperature fraction.

2.8.12 Fractional Readout Noise

BoloCalc has an option to set readout noise not explicitly but instead as a fractional increase on the total NEP. If either R_{bolo} or NEI from Equation 8 are “NA,” and if the parameter “Read Noise Frac” is not “NA” in the `channels.txt` file (see Section 3.5.3 for more detail), then the following equation is used to calculate readout noise:

$$NEP_{\text{read}} = \sqrt{(1 + \Delta_{\text{read}})^2 - 1} \times \sqrt{NEP_{\text{ph}}^2 + NEP_{\text{g}}^2} \quad (24)$$

where Δ_{read} is the fractional increase in the total NEP due to readout noise.

2.8.13 Number of Detectors

The number of total detectors in a given channel N_{det} is determined by the number of detectors per wafer $N_{\text{det/waf}}$, the number of wafers per optics tube (or equivalently, per camera) $N_{\text{waf/OT}}$, and the number of optics tubes per telescope N_{OT} .

$$N_{\text{det}} = N_{\text{det/waf}} \times N_{\text{waf/OT}} \times N_{\text{OT}} \quad (25)$$

3 Defining an Experiment

BoloCalc has a modular object-oriented structure, which allows for arbitrary mixtures of sites, telescopes, cameras, optics, focal planes, and detectors. A BoloCalc project has the parent-child structure shown in Figure 2 and is built with four layers: experiments, telescopes, cameras, and channels, which are defined in Tab. 1. Each experiment can have an arbitrary set of telescopes (at different sites), each telescope an arbitrary set of cameras, and each camera an arbitrary set of channels.

Layer	Definition
Experiment	An assemblage of CMB telescopes.
Telescope	A platform that carries and points one or more cameras. It observes at a specified site with a specified observation strategy and can include warm reflectors.
Camera	A cryostat that houses cryogenic optics, filters, and detectors. Multiple cameras can be mounted on the same telescope.
Channel	A frequency band observed by some set of detectors within a camera. A multichroic camera will have multiple channels.

Table 1: Definitions of the layers used to build a BoloCalc project.

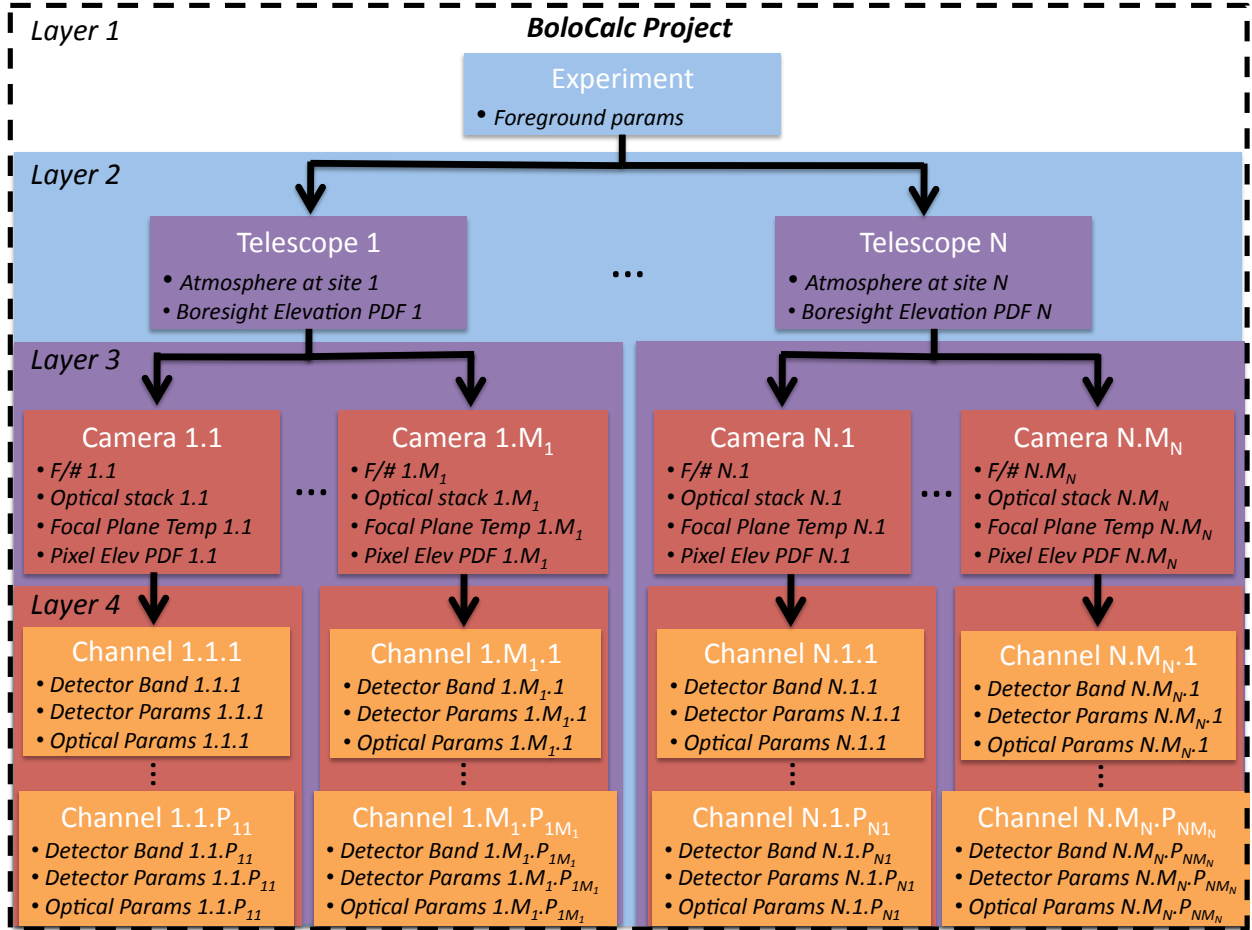


Figure 2: Layout of a BoloCalc project.

3.1 Experiment Directory Layout

The experiment directory is meant to be easily generalizable to various kinds of experiments and telescopes. For information about how experiments are stored within the `BoloCalc/Experiments/` directory, see Section 1.3. Below is an overview of the file structure within an experiment directory.

Each layer of an experiment has several parameters that define the input experiment. In this section, we go through them by input file.

```

1 BoloCalc/Experiements/[Experiment Name]/[Experiment Design]
2 |-- config
3 |   |--foregrounds.txt
4 |--Telescope1/
5 |   |-- config/
6 |   |   |-- telescope.txt
7 |   |   |-- elevation.txt
8 |   |   |-- [Custom Atmospheric Profile] (optional)
9 |   |-- Camera1/
10 |   |   |-- config/
11 |   |   |   |-- camera.txt
12 |   |   |   |-- channels.txt
13 |   |   |   |-- optics.txt
14 |   |   |   |-- elevation.txt
15 |   |   |   |-- Bands/
16 |   |   |       |-- [Optic Band File 1] (optional)
17 |   |   |       |-- [Optic Band File 2] (optional)
18 |   |   |       |-- ...
19 |   |   |       |-- [Detector Band File 1] (optional)
20 |   |   |       |-- [Detector Band File 2] (optional)
21 |   |   |       |-- ...
22 |   |-- Camera2/
23 |   |-- ...
24 |-- Telescope2/
25 |-- ...

```

As stated in Section 3, an experiment definition has four layers:

1. Experiment
 - Composed of (multiple) telescopes
 - All telescopes in a given experiment see the same galactic foregrounds
2. Telescopes
 - Composed of (multiple) cameras
 - All cameras in a given telescope see the same atmosphere
 - All cameras in a given telescope have the same scan strategy
3. Cameras
 - Composed of (multiple) frequency channels
 - All frequency channels in a given camera have the same camera parameters
 - All frequency channels in a given camera see the same optical chain
4. Channels
 - Each channel has independent detector and optical element parameters

These layers allow telescopes to be added to an experiment in an arbitrary location, each with arbitrary cameras, each with arbitrary frequency channels.

3.2 Parameter Definitions

All parameters in BoloCalc are defined with a central value and an error bar. Therefore, all parameter variation is assumed to have an underlying Gaussian distribution, and all parameters are assumed to be uncorrelated.

Parameters defined in `foregrounds.txt`, `telescope.txt`, `camera.txt`, and `channels.txt` have the following format:

```

1 Mean +/- Std Dev

```

The following formatting rules should be followed for the code to work robustly and the tables to render neatly:

1. All central values and standard deviations should have five characters. For instance, 1.000 and 10.00 are valid, while 10.000 is not.
2. All central values and standard deviations should be separated by \pm , noting the spaces on either side of the \pm sign.

Parameters defined in `optics.txt` are defined for multiple bands in multi-chroic cameras and therefore follow a different format:

```
1 [Mean for Band 1, Mean for Band 2, ...] +/- [Std Dev for Band 1, Std Dev for
   Band 2, ...]
```

In a similar way as with the single values, the following formatting rules should be followed for the code to work robustly and the tables to render neatly:

1. All central values and standard deviations should have five characters. For instance, 1.000 and 10.00 are valid, while 10.000 is not.
2. All arrays of central values and arrays of standard deviations should be separated by \pm , noting the spaces on either side of the \pm sign

3.3 Experiment Parameters (Layer 1)

The experiment is comprised of an arbitrary set of telescopes, which are assumed to see the same the same celestial loading, including power from the CMB and galactic foregrounds.

3.3.1 `foregrounds.txt`

The below list overviews the parameters defined in

`[Experiment Name]/[Experiment Version]/config/foregrounds.txt`, the file that defines the galaxy through which the telescope must observe the CMB. This galactic parameters will be the same for all telescopes.

- **Synchrotron Spec Index**
 - *Definition*: “ n_s ” in Equation 14
 - *Description*: Spectral index of the synchrotron power law spectrum
 - *Allowed Values*: Floating point value between $(-\infty, +\infty]$
 - *Default Value*: 2×10^{-3}
- **Synchrotron Amplitude**
 - *Definition*: “ A_s ” in Equation 14
 - *Description*: Amplitude of the synchrotron power law spectrum
 - *Allowed Values*: Floating point value between $(0.0, +\infty)$ W/Hz
 - *Default Value*: 6×10^3 W/Hz
- **Dust Temperature**
 - *Definition*: “ T_d ” in Equation 15
 - *Description*: Dust modified blackbody temperature
 - *Allowed Values*: Floating point value between $(0.0, +\infty)$ K
 - *Default Value*: 19.70 K
- **Dust Spec Index**
 - *Definition*: “ n_d ” in Equation 15
 - *Description*: The spectral index of the dust modified blackbody spectrum
 - *Allowed Values*: Floating point value between $(-\infty, +\infty)$

- *Default Value:* 1.50
- **Dust Amplitude**
 - *Definition:* “ A_d ” in Equation 15
 - *Description:* Amplitude of the dust modified blackbody spectrum
 - *Allowed Values:* Floating point value between $(0.0, +\infty)$
 - *Default Value:* 2×10^{-3}
- **Dust Scale Frequency**
 - *Definition:* “ ν_d ” in Equation 15
 - *Description:* Scale frequency for the dust modified blackbody spectrum
 - *Allowed Values:* Floating point value between $(0.0, +\infty)$
 - *Default Value:* 353.0 GHz

3.4 Telescope Parameters (Layer 2)

The following sections overview the input parameters for each of the telescope input files.

3.4.1 `telescope.txt`

The below list overviews the parameters defined in `telescope.txt`. These parameters are used to define the site, atmospheric profile, and patch coverage parameters for the telescope.

- **Site**
 - *Description:* Site at which the telescope observes, which defines the atmospheric conditions.
 - *Allowed Values:* Either “Atacama” or “Pole”
 - *Default Value:* “Atacama”
- **Elevation**
 - *Definition:* θ_{tel} defined in 21
 - *Description:* Elevation at which the telescope observes.
 - *Allowed Values:* $[20.0, 90.0]$ deg above the horizon
 - *Default Value:* 50.0 deg
 - If “NA,” the elevation distribution defined in `Telescope/config/elevation.txt`, will be assumed.
- **PWV**
 - *Description:* Precipitable water vapor of the atmosphere through which the telescope observes.
 - *Allowed Values:* $[0.0, 8.0]$ mm
 - *Default Value:* 1.0 mm
 - If “NA,” the PWV distribution defined in `Telescope/config/pwv.txt`, will be assumed.
- **Observation Time**
 - *Definition:* t_{obs} in Equation 12
 - *Description:* How long this telescope will observe
 - *Allowed Values:* Floating point value between $(0.0, +\infty)$ years
 - *Default Value:* 3.0
- **Sky Fraction**
 - *Definition:* f_{sky} in Equation 12
 - *Description:* What fraction of the sky this telescope will observe
 - *Allowed Values:* Floating point value between $(0.0, 1.0]$
 - *Default Value:* 0.7

- **Observation Efficiency**

- *Definition:* η_{obs} in Equation 12
- *Description:* What fraction of the observation time will the telescope be actually observing
- *Allowed Values:* Floating point value between (0.0, 1.0]
- *Default Value:* 0.8

- **NET Margin**

- *Definition:* M in Equation 9
- *Description:* Agnostic factor which multiplies the detector NETs for this telescope
- *Allowed Values:* Floating point value between (0.0, $+\infty$)
- *Default Value:* 1.0

3.4.2 elevation.txt

`elevation.txt` is an optional file that defines the elevation probability distribution function (PDF), which is determined by the scan strategy of the telescope. The first column is the elevation above the horizon in degrees, and the second column is the fraction of time spent at that elevation. This file is only used if the parameter “Elevation” in `Telescope/config/telescope.txt` is defined to be “NA” (see Section 3.4.1 for more information about `telescope.txt`). The probabilities should add up to one, and if they do not, the fraction will be determined by the provided probability divided by the sum of the provided probabilities.

Below is a made-up example for `elevation.txt`. Notice that each row is separated by a commented line of hyphens, while each column is separated by a vertical bar.

```

1  #####Elevation Distribution#####
2  #
3  Elevation      | Probability
4  #-----
5  [deg]          | NA
6  #-----
7  30.0           | 0.100
8  #-----
9  40.0           | 0.100
10 #-----
11 50.0           | 0.300
12 #-----
13 60.0           | 0.400
14 #-----
15 70.0           | 0.100
16 #-----

```

3.4.3 pwv.txt

`pwv.txt` is an optional file that defines the precipitable water vapor (PWV) probability distribution function (PDF), which is determined by the weather conditions at the observation site. The first column is the PWV in mm, and the second column is the fraction of time spent at that PWV value. This file is only used if the parameter “PWV” in `Telescope/config/telescope.txt` is defined to be “NA” (see Section 3.4.1 for more information about `telescope.txt`). The probabilities should add up to one, and if they do not, the fraction will be determined by the provided probability divided by the sum of the provided probabilities.

Below is a made-up example for `pwv.txt`. Notice that each row is separated by a commented line of hyphens, while each column is separated by a vertical bar.


```

1  #####Site PWV Distribution#####
2  #
3  PWV          | Probability
4  #-----
5  [mm]         | NA
6  #-----
7  0.200        | 0.100
8  #-----
9  0.400        | 0.100
10 #-----
11 0.600        | 0.200
12 #-----
13 0.800        | 0.200
14 #-----
15 1.000        | 0.100
16 #-----
17 1.200        | 0.100
18 #-----
19 1.400        | 0.100
20 #-----
21 1.600        | 0.100
22 #-----

```

3.4.4 atm.txt

There is an option to override the code’s handling of the atmosphere by providing a custom atmospheric profile. The atmosphere file must be in [Telescope Name]/config/, and the filename must have the form atm.[Elevation]deg_[PWV]um.txt. For example, a valid filename is atm.60deg_1000um.txt.

The atmosphere file should have four columns, delimited by white space:

1. Column 1: Frequency [GHz]
2. Column 2: Optical Depth
3. Column 3: Planck Temperature [K]
4. Column 4: Transmission

```

1 | Frequency [GHz] | Optical Depth | Planck Temp [K] | Transmission |

```

Note that the temperature is in Planck units, not Rayleigh-Jeans. The passed frequency range in the atmosphere file needs to be large enough to cover the frequency bands observed by the telescope—which are defined in the “channels.txt” and “Bands/” files at the camera directory level (see Section 3.5.3 and Section for more details)—with at least a 15% buffer on the bandwidths of the highest and lowest bands. For instance, if your detector bands range from 240 to 300 GHz, your atmosphere file should have values that go from at least 205 to 345 GHz.

Below is an example of an atmospheric profile generated using the AM simulation code and a configuration file designed using the 10-year MERRA dataset for the Chajnantor Observatory.

```

1  1.0000000e+01  3.664259e-03  3.589557e+00  9.963424e-01
2  1.0020000e+01  3.667538e-03  3.590339e+00  9.963392e-01
3  1.0040000e+01  3.671011e-03  3.591163e+00  9.963357e-01
4  1.0060000e+01  3.674741e-03  3.592043e+00  9.963320e-01
5  1.0080000e+01  3.678826e-03  3.593001e+00  9.963279e-01
6  1.0100000e+01  3.683414e-03  3.594069e+00  9.963234e-01
7  1.0120000e+01  3.688744e-03  3.595301e+00  9.963181e-01
8  1.0140000e+01  3.695220e-03  3.596788e+00  9.963116e-01
9  1.0160000e+01  3.703583e-03  3.598698e+00  9.963033e-01

```

```

10 1.0180000e+01 3.715317e-03 3.601374e+00 9.962916e-01
11 1.0200000e+01 3.733976e-03 3.605660e+00 9.962730e-01
12 ...

```

3.5 Camera Parameters (Layers 3 and 4)

In the following sections, we will overview the input parameters for each of the camera input files.

3.5.1 camera.txt

`camera.txt` defines parameters that are the same for all frequency bands observed with this camera. The file must be located within `[Telescope Name]/[Camera Name]/config/`.

- **Boresight Elevation**

- *Definition:* θ_{cam} in Equation 21.
- *Description:* The elevation of the camera boresight with respect to the telescope boresight.
- *Allowed Values:* Floating point value between $(-40.0, 40.0]$ deg
- *Default Value:* 0.0 deg

- **Optical Coupling**

- *Definition:* ξ in Equation 10.
- *Description:* A parameter that quantifies the SNR degradation associated with how well the far-field image couples to the detector pixels on the focal plane.
- *Allowed Values:* Floating point value between $(0.0, 1.0]$
- *Default Value:* 1.0

- **F-number**

- *Definition:* F in Equation 16.
- *Description:* The focal ratio at the focal plane
- *Allowed Values:* Floating point value between $(0.0, +\infty)$
- *Default Value:* 2.0

- **Bath Temp**

- *Definition:* T_{bath} in Equation 6.
- *Description:* The temperature of the focal plane
- *Allowed Values:* Floating point value between $(0.0, +\infty)$ K
- *Default Value:* 0.100 K

3.5.2 elevation.txt

`elevation.txt` is an optional file that defines the elevation distribution with respect to the camera boresight of pixels on the focal plane. This file sets the value for θ_{mahrmelv} defined in equation 21.

This pixel elevation distribution file has the exact same structure as the file `Telescope/config/elevation.txt`, which instead defines the scan strategy and is described in section [?]. The first column is the elevation above the horizon in degrees, and the second column is the fraction of time spent at that elevation.

This file is only used if the parameter “Elevation” in `Telescope/config/elevation.txt` is defined to be “NA” (see Section 3.4.1 for more information about `telescope.txt`).

The probabilities should add up to one, and if they don not, the fraction will be determined by the provided probability divided by the sum of the provided probabilities.

Below is a made-up example of correct formatting for `elevation.txt`. Notice that each row is separated by a commented line of hyphens, while each column is separated by a vertical bar.

```

1  #####Elevation Distribution#####
2  #
3  Elevation      | Probability
4  #-----
5  [deg]          | NA
6  #-----
7  30.0           | 0.100
8  #-----
9  40.0           | 0.100
10 #-----
11 50.0           | 0.300
12 #-----
13 60.0           | 0.400
14 #-----
15 70.0           | 0.100
16 #-----

```

3.5.3 channels.txt

`channels.txt` defines the parameters of the frequency bands and detectors that are observing in this camera. The file must be located within `[Telescope Name]/[Camera Name]/config/`.

- **Band ID**
 - *Description:* The identification number for this band.
 - *Allowed Values:* An integer value between $[1, +\infty)$
 - *Default Value:* 1
 - Cannot have two bands with same Band ID on the same pixel
- **Pixel ID**
 - *Description:* The identification number for the pixel this band observes from
 - *Allowed Values:* An integer value between $[1, +\infty)$
 - *Default Value:* 1
 - Cannot have two bands with same Band ID on the same pixel
- **Band Center**
 - *Definition:* ν_c in Equation 26.
 - *Description:* The central frequency for this band.
 - *Allowed Values:* A floating point value between $(0.0, +\infty)$ GHz
 - *Default Value:* $[90, 150]$ GHz
- **Fractional BW**
 - *Definition:* f_{BW} in Equation 27.
 - *Description:* Fractional arithmetic bandwidth for this band
 - *Allowed Values:* A floating point value between $(0.0, 2.0]$
 - *Default Value:* $[0.350, 0.275]$
- **Pixel Size**
 - *Definition:* D in Equation 16.
 - *Description:* Size of the pixel via which this band is observing
 - *Allowed Values:* A positive floating point value between $(0.0, +\infty)$ mm
 - *Default Value:* 6.8 mm
- **Num Det per Wafer**
 - *Definition:* $N_{\text{det/waf}}$ in Equation 25.
 - *Description:* Number of detectors per wafer within this band.
 - *Allowed Values:* An integer value between $[1, +\infty)$

- *Default Value:* 542
- **Num Waf per OT**
 - *Definition:* $N_{\text{waf/OT}}$ in Equation 25.
 - *Description:* Number of wafers with this band per camera.
 - *Allowed Values:* An integer value between $[1, +\infty)$
 - *Default Value:* 7
- **Num OT**
 - *Definition:* N_{OT} in Equation 25.
 - *Description:* Number of cameras in this telescope observing at this frequency.
 - *Allowed Values:* An integer value between $[1, +\infty)$
 - *Default Value:* 1
- **Waist Factor**
 - *Definition:* w_f in Equation 16.
 - *Description:* The ratio of the Gaussian beam waist coming from the pixel aperture to the pixel aperture diameter for this band.
 - *Allowed Values:* A floating point value between $[2.0, +\infty)$
 - *Default Value:* 3.0
- **Det Eff**
 - *Definition:* η_{det} in Equation 26.
 - *Description:* Efficiency of the detector from the entry point of the pixel aperture to the bolometer.
 - *Allowed Values:* A floating point value between $(0.0, 1.0]$
 - *Default Value:* 0.700
- **Psat**
 - *Definition:* P_{sat} in Equation 7.
 - *Description:* Saturation power of the bolometer.
 - *Allowed Values:* A floating point value between $(0.0, +\infty]$ pW
 - *Default Value:* “NA”
 - If “NA,” use Equation 22 to calculate P_{sat}
- **Psat Factor**
 - *Definition:* f_{psat} in Equation 22.
 - *Description:* Safety factor used to calculate saturation power from optical power
 - *Allowed Values:* A floating point value between $(0.0, +\infty)$
 - *Default Value:* 3.0
 - If “Psat” is not “NA,” this parameter is ignored.
- **Carrier Index**
 - *Definition:* n in Equations 7 and 6.
 - *Description:* Thermal carrier index for bolometer leg conductivity
 - *Allowed Values:* A floating point value between $(0.0, +\infty)$
 - *Default Value:* 2.7
- **Tc**
 - *Definition:* T_{oper} in Equations 5, 6, and 7.
 - *Description:* Bolometer operating temperature
 - *Allowed Values:* A floating point value between $(0.0, +\infty)$ K
 - *Default Value:* 0.165 K
 - If “NA,” use Equation 23 to calculate operating temperature.
- **Tc Fraction**
 - *Definition:* f_{oper} in Equation 23.

- *Description*: Factor used to calculate transition temperature from the bath temperature
- *Allowed Values*: A floating point value between (1.0, $+\infty$)
- *Default Value*: “NA”
- If “Tc” is not “NA,” this parameter is ignored.
- **Yield**
 - *Definition*: Y in Equation 11.
 - *Description*: Fraction of deployed detectors that are observing. This is also commonly referred to as “end-to-end” yield.
 - *Allowed Values*: A floating point value between (0.0, 1.0]
 - *Default Value*: 0.7
- **SQUID NEI**
 - *Definition*: NEI in Equation 8.
 - *Description*: SQUID noise equivalent current.
 - *Allowed Values*: A floating point value between (0.0, $+\infty$) pA/ $\sqrt{\text{Hz}}$
 - *Default Value*: “NA”
 - If “NA,” use Equation 24 to calculate readout noise NEP_{read}
- **Bolo Resistance**
 - *Definition*: R_{bolo} in Equation 8.
 - *Description*: Bolometer operating resistance.
 - *Allowed Values*: A floating point value between (0.0, $+\infty$) Ω
 - *Default Value*: “NA”
 - If “NA,” use Equation 24 to calculate readout noise NEP_{read}
- **Read Noise Frac**
 - *Definition*: Δ_{read} in Equation 24.
 - *Description*: Fraction of the total NEP that is due to readout noise.
 - *Allowed Values*: A floating point value between [0.0, $+\infty$)
 - *Default Value*: 0.1
 - If “SQUID NEI” and “Bolo Resistance” are not “NA,” this parameter is ignored.

3.5.4 optics.txt

optics.txt defines the parameters the frequency bands and detectors that are observing in this camera. The file must be located within [Telescope Name]/[Camera Name]/config/.

- **Element**
 - *Description*: Name of optical element
 - *Allowed Values*: Any string
 - *Default Value*: “Element”
 - **RESERVED NAMES**: “Primary,” “Mirror,” “Aperture,” and “Stop” trigger additional calculations and therefore must be used purposefully
- **Temperature**
 - *Definition*: T_i in Equation 2.8.
 - *Description*: Temperature of the optical element
 - *Allowed Values*: A floating point value between (0.0, $+\infty$) K
 - *Default Value*: 4 K
- **Absorption**
 - *Definition*: ϵ_i in Equation 1.
 - *Description*: Fractional power attenuation/emission due to absorption within the optical element

- *Allowed Values:* A floating point value between [0.0, 1.0)
 - *Default Value:* 0.000
 - If “NA” ...
 - * If “Mirror” or “Primary” is in “Element,” Equation 18 is used to calculate ohmic losses in the reflective optical element.
 - * If “Mirror” or “Primary” is not in “Element,” Equation 17 is used to calculate dielectric loss in the refractive optical element.
- **Reflection**
 - *Definition:* r in Equation 20.
 - *Description:* Fractional power lost due to reflection at the optical element.
 - *Allowed Values:* A floating point value between [0.0, 1.0)
 - *Default Value:* 0.000
 - If “NA,” Equation 20 is used to calculate reflection loss
- **Thickness**
 - *Definition:* t in Equation 17
 - *Description:* Thickness of the optical element
 - *Allowed Values:* A floating point value between (0.0, $+\infty$) mm
 - *Default Value:* “NA”
 - If “Mirror” or “Primary” is in “Element,” this parameter is ignored
 - If “Absorption” is not “NA,” this parameter is ignored
- **Index**
 - *Definition:* n in Equation 17.
 - *Description:* Index of refraction of the optical element.
 - *Allowed Values:* A floating point value between [1.0, $+\infty$)
 - *Default Value:* “NA”
 - If “Mirror” or “Primary” is in “Element,” this parameter is ignored
 - If “Absorption” is not “NA,” this parameter is ignored
- **Loss Tangent**
 - *Definition:* $\tan \delta$ in Equation 17.
 - *Description:* Loss tangent of the optical element.
 - *Allowed Values:* A floating point value between [0.0, $+\infty$)
 - *Default Value:* “NA”
 - If “Mirror” or “Primary” is in “Element,” this parameter is ignored
 - If “Absorption” is not “NA,” this parameter is ignored
- **Conductivity**
 - *Definition:* σ_c in Equation 18.
 - *Description:* Electrical conductivity of the optical element.
 - *Allowed Values:* A floating point value between (0.0, $+\infty$)
 - *Default Value:* “NA”
 - If “Mirror” or “Primary” is not in “Element,” this parameter is ignored
 - If “Absorption” is not “NA,” this parameter is ignored
- **Scatter Frac**
 - *Definition:* δ_i in Equation 2.8.
 - *Description:* Fractional power lost due to scattering.
 - *Allowed Values:* A floating point value between [0.0, 1.0]
 - *Default Value:* “NA”
 - If “NA,” Equation 19 is use to calculate scattering loss.
- **Surface Rough**

- *Definition:* σ_r in Equation 19.
- *Description:* Surface roughness of the optical element.
- *Allowed Values:* A floating point value between $[0.0, +\infty)$
- *Default Value:* “NA”
- If “Scatter Frac” is not “NA,” this parameter is ignored.
- If “NA” and “Scatter Frac” is “NA,” scattering is assumed to be zero.
- **Scatter Temp**
 - *Definition:* $T_{\delta,i}$ in Equation 2.8.
 - *Description:* The effective temperature that the scattered power lands on.
 - *Allowed Values:* A floating point value between $[0.0, +\infty)$
 - *Default Value:* “NA”
 - If “NA,” scattered power is assumed to land at “Temperature,” the temperature of the optical element itself
- **Spillover**
 - *Definition:* β_i in Equation 2.8.
 - *Description:* Fractional power that spills over the optical element.
 - *Allowed Values:* A floating point value between $[0.0, 1.0)$
 - *Default Value:* “NA”
 - If “NA,” spillover is assumed be zero
- **Spillover Temp**
 - *Definition:* $T_{\beta,i}$ in Equation 2.8.
 - *Description:* The effective temperature that the spilled power lands on.
 - *Allowed Values:* A floating point value between $[0.0, +\infty)$
 - *Default Value:* “NA”
 - If “NA,” spillover is assumed to land at “Temperature,” the temperature of the element itself

3.6 Custom Bands

By default, BoloCalc assumes top hat bands for all detectors and optical elements, whose height is determined solely by the mean and standard deviation of the end-to-end optical efficiency. However, custom bands can be input for any optical element.

The general file format for all band files is the same:

```
1 | Frequency [GHz] | Mean Efficiency | Standard Deviation (optional) |
```

Each column is separated by a vertical bar, and the final column, which contains the error bars, can be omitted. If only two columns are present, all stand deviations are assumed to be zero.

An example of a made-up detector bandpass text file, which is space-delimited, is shown below:

```
1 70.      0.000  0.000
2 75.      0.500  0.100
3 80.      0.700  0.100
4 85.      0.650  0.100
5 90.      0.700  0.100
6 100.     0.800  0.100
7 105.     0.600  0.100
8 110.     0.700  0.100
9 115.     0.500  0.100
10 120.     0.000  0.000
```

3.6.1 Detectors

Detector band files are stored in

```
1 [Telescope Name]/[Camera Name]/config/Bands/
```

The band is identified using its filename, which must have the format `[Camera Name][Band ID].txt` for text files, or `[Camera Name][Band ID].csv` for comma-separated value files; both file formats work equally well. For example, to load a text file for a camera named “MF” (in the directory `[Telescope Name]/MF/` and a Band ID “1,” the file name should be `MF1.txt`. Note that for a multi-chroic camera, there should be multiple band files, one for each frequency channel.

3.6.2 Optics

Optics band files are stored in

```
1 [Telescope Name]/[Camera Name]/config/Bands/
```

The band is identified using its filename, which must have the format `[Optical Element Name].txt` for text files, or `[Optical Element Name].csv` for comma-separated value files; both file formats work equally well. For example, to load a text file for an optical element named “Lens1,” the file name should be `Lens1.txt`. Note that there is only one band file for optical elements, even in multi-chroic cameras, as all detectors are all frequencies see the same bandpass. Also note that you can apply the same band to multiple optical elements by duplicating matching names in `optics.txt`.

3.7 Default Bands

If no custom bands are defined, `BoloCalc` assumes a top-hat band function, which is determined by the detector bandpass.

3.7.1 Detectors

If no custom detector bandpass is provided, then a top-hat band is assumed with an overall efficiency factor for the detector η_{det}

$$B(\nu) = \begin{cases} \eta_{\text{det}} & \text{if } \nu_{\text{lo}} \leq \nu \leq \nu_{\text{hi}} \\ 0 & \text{otherwise} \end{cases}, \quad (26)$$

where the high and low band edges are defined as

$$\nu_{\text{lo}} = \nu_c \left(1 + \frac{f_{\text{BW}}}{2}\right) ; \quad \nu_{\text{hi}} = \nu_c \left(1 - \frac{f_{\text{BW}}}{2}\right) \quad (27)$$

where ν_c is the band central frequency and f_{BW} is the fractional bandwidth.

3.7.2 Optics

If no custom bandpass is provided for an optical element i , then the bandpass is assumed to be a top-hat with an overall efficiency factor defined in Equation 20

$$B(\nu) = \begin{cases} \eta_i & \text{if } \nu_{\text{lo}} \leq \nu \leq \nu_{\text{hi}} \\ 0 & \text{otherwise} \end{cases} \quad (28)$$

If a custom detector bandpass is defined, then the frequency range covered ν_{lo} to ν_{hi} are the highest and lowest frequencies defined in the custom band. Otherwise, ν_{lo} and ν_{hi} are set by Equation 27.

4 Monte Carlo Simulation

The Monte Carlo (MC) simulation iterates using a nest of the following structure:

- N_{exp} Experiment realizations
 - N_{obs} Observations per experiment realization
 - * N_{det} Detector realizations per observation

4.1 Simulation Inputs

Simulation inputs are stored in the file `BoloCalc/config/simulationInputs.txt`. An example of the file is shown in Figure 1. Below, we define the parameters for the simulation explicitly.

- **Multiprocess**
 - *Description*: Whether or not to use multiprocessing to calculate statistics in parallel
 - *Allowed Values*: “True” or “False”
 - *Default Value*: “False”
- **Cores**
 - *Description*: Number of cores to pool for parallel processing
 - *Allowed Values*: Positive integer between $[1, +\infty)$
 - *Default Value*: 1
- **Verbosity**
 - *Description*: Logging verbosity. 2 to print all logging output, 1 to print some output, 0 to print little output
 - *Allowed Values*: Integer between $[0, 2]$
 - *Default Value*: 1
- **Experiments**
 - *Description*: Number of experiment realizations to Monte Carlo; each time an experiment is MC-ed, its foreground, telescope, camera, and optical parameters are re-sampled.
 - *Allowed Values*: Positive integer $[1, +\infty)$
 - *Default Value*: 1
 - If 1, assume central value (i.e. ignore the spread) for all parameters set by the experiment realization.
- **Observations**
 - *Description*: Number of independently observations, which draw the PWV and telescope boresight elevation from the distributions defined in `Telescope/config/pwv.txt` and `Telescope/config/elevation.txt`.
 - *Allowed Values*: Positive integer $[1, +\infty)$
 - *Default Value*: 1
- **Detectors**
 - *Description*: Number of independently-sampled detector parameters defined in `channels.txt` per observation.
 - *Allowed Values*: Positive integer $[1, +\infty)$
 - *Default Value*: 1
 - If 1, assume central value for all detector parameters

- **Resolution**
 - *Description*: Frequency resolution assumed in the simulation.
 - *Allowed Values*: A positive floating-point value between (0.0, 20.0] GHz
 - *Default Value*: 0.1
 - Reducing the resolution can have a dramatic impact on the speed of the MC simulations.
- **Foregrounds**
 - *Description*: Whether or not to include foregrounds in the estimate of the optical loading.
 - *Allowed Values*: “True” or “False”
 - *Default Value*: “False”
 - Celestial loading doesn’t have a noticeable impact on ground-based telescopes but can be important for satellite experiments.
- **Correlations**
 - *Description*: Whether to impose white noise correlations when calculating array NET
 - *Allowed Values*: “True” or “False”
 - *Default Value*: “True”
 - This switch is useful for computing the impact of white-noise correlations on array NET.

5 Output Files

Running `mappingSpeed.py` generates several output files, which quantify the performance of the simulated experiment. Note that all examples in this document assume the Simons Observatory V3 configuration with no error bars for all parameters, with all observations at 60 deg elevation and 1 mm PWV.

5.1 Sensitivity Tables

BoloCalc produces tables of outputs related to the white noise performance of the instrument. All sensitivity tables are in files named `sensitivity.txt`, and there are multiple tables generated at multiple directory tree levels, describing the performance of the camera, telescope, and entire experiment.

5.1.1 Experiment/`sensitivity.txt`

This file lists the following parameters:

- **Chan**
 - Frequency channel name [`Camera Name`] [`Band ID`]
- **Frequency**
 - Central frequency of the frequency channel
- **Frac Bandwidth**
 - Fractional bandwidth of frequency channel
- **Num Det**
 - Total number of detectors deployed in this frequency channel, combined for all telescopes within the experiment
- **Array NET**
 - Array-averaged noise equivalent temperature of this frequency channel, combined for all telescopes within the experiment
- **Mapping Speed**

- Mapping speed of this frequency channel, combined for all telescopes within the experiment
- Map Depth
 - Map depth achieved by this frequency channel, combined for all telescopes within the experiment

Note that if identical bands are shared between telescopes, the detectors are combined in this table, with the combined Array NET of that band is found by taking the inverse-variance average of the duplicated bands.

Below is an example of `sensitivity.txt` at the Experiment directory level:

Chan	Frequency	Frac Bandwidth	Num Det	Array NET	Mapping Speed	Map Depth
	[GHz]			[uK-rtSec]	[(uK ² s) ⁻¹]	[uK-arcmin]
LF1	27.0 +/- 0.0	0.222 +/- 0.000	444	24.37 +/- 0.00	0.0017 +/- 0.0000	40.4 +/- 0.0
LF2	39.0 +/- 0.0	0.462 +/- 0.000	444	14.08 +/- 0.00	0.0050 +/- 0.0000	23.3 +/- 0.0
MF1	93.0 +/- 0.0	0.376 +/- 0.000	14092	2.42 +/- 0.00	0.1706 +/- 0.0000	4.0 +/- 0.0
MF2	145.0 +/- 0.0	0.276 +/- 0.000	14092	3.04 +/- 0.00	0.1079 +/- 0.0000	5.0 +/- 0.0
UHF1	225.0 +/- 0.0	0.267 +/- 0.000	16836	5.54 +/- 0.00	0.0326 +/- 0.0000	9.2 +/- 0.0
UHF2	278.0 +/- 0.0	0.162 +/- 0.000	16836	13.28 +/- 0.00	0.0057 +/- 0.0000	22.0 +/- 0.0
Total			62744	1.76 +/- 0.00	0.3235 +/- 0.0000	2.9 +/- 0.0

Figure 3: Simons Observatory V3 experiment sensitivity, which combines the NETs from the Large Aperture Telescope and the Small Aperture Telescope. Both telescopes share wafers, and therefore have duplicates of each frequency band. The NETs are inverse-variance averaged in this sensitivity table for the entire experiment.

5.1.2 Experiment/Telescope/sensitivity.txt

This file lists

- Chan
 - Frequency channel name [Camera Name] [Band ID]
- Frequency
 - Central frequency of the frequency channel
- Frac Bandwidth
 - Fractional bandwidth of frequency channel
- Num Det
 - Total number of detectors deployed in this frequency channel, combined for all cameras within the telescope
- Array NET
 - Aarray-averaged noise equivalent temperature of this frequency channel, combined for all cameras within the telescope
- Mapping Speed
 - Mapping speed of this frequency channel, combined for all cameras within the telescope
- Map Depth
 - Map depth achieved by this frequency channel, combined for all cameras within the telescope

Note that if identical bands are shared between cameras, the detectors are combined in this table, with the combined Array NET of that band is found by taking the inverse-variance average of the duplicated bands.

Below is an example of `sensitivity.txt` at the Telescope directory level:

Chan	Frequency	Frac Bandwidth	Num Det	Array NET	Mapping Speed	Map Depth
	[GHz]			[uK-rtSec]	[$(\text{uK}^2 \text{ s})^{-1}$]	[uK-arcmin]
LF1	27.0 +/- 0.0	0.222 +/- 0.000	222	43.21 +/- 0.00	0.0005 +/- 0.0000	71.6 +/- 0.0
LF2	39.0 +/- 0.0	0.462 +/- 0.000	222	22.14 +/- 0.00	0.0020 +/- 0.0000	36.7 +/- 0.0
MF1	93.0 +/- 0.0	0.376 +/- 0.000	6504	4.60 +/- 0.00	0.0473 +/- 0.0000	7.6 +/- 0.0
MF2	145.0 +/- 0.0	0.276 +/- 0.000	6504	5.71 +/- 0.00	0.0306 +/- 0.0000	9.5 +/- 0.0
UHF1	225.0 +/- 0.0	0.267 +/- 0.000	6084	12.63 +/- 0.00	0.0063 +/- 0.0000	20.9 +/- 0.0
UHF2	278.0 +/- 0.0	0.162 +/- 0.000	6084	30.30 +/- 0.00	0.0011 +/- 0.0000	50.2 +/- 0.0
Total			25620	3.37 +/- 0.00	0.0879 +/- 0.0000	5.6 +/- 0.0

Figure 4: Simons Observatory V3 Large Aperture Telescope sensitivity, which combines the NETs from the LF, MF, and UHF cameras.

5.1.3 Experiment/Telescope/Camera/sensitivity.txt

This file lists

- Chan
 - Frequency channel name [Camera Name] [Band ID]
- Frequency
 - Central frequency of the frequency channel
- Frac Bandwidth
 - Fractional bandwidth of frequency channel
- Num Det
 - Total number of detectors deployed in this frequency channel, combined for all cameras within the telescope
- Chan
 - Frequency channel name [Camera Name] [Band ID]
- Lyot Efficiency
 - Lyot stop spillover efficiency of the frequency channel
- Optical Power
 - Total optical power on detectors within this frequency channel
- Photon NEP
 - Noise-equivalent power due to photon noise for detectors within this frequency channel
- Bolometer NEP
 - Noise-equivalent power due to bolometer thermal carrier noise for detectors within this frequency channel
- Detector NEP
 - Total noise-equivalent power due to the combinatin of photon, thermal carrier, and readout noise for detectors within this frequency channel
- Detector NET
 - Per-detector noise-equivalent temperature for detectors within this channel

- Array NET
 - Array-averaged noise equivalent temperature of this frequency channel
- Mapping Speed
 - Mapping speed of this frequency channel
- Map Depth
 - Map depth achieved by this frequency channel

for every frequency channel in the provided camera.

Below is an example of `sensitivity.txt` at the Camera directory level:

Chan	Frequency	Prac Bandwidth	Num Det	Lyot Efficiency	Optical Power	Photon NEP	Bolometer NEP	Readout NEP	Detector NEP	Detector NET	Array NET	Mapping Speed	Map Depth
	[GHz]			[%]	[pW]	[aW/rtHz]	[aW/rtHz]	[aW/rtHz]	[aW/rtHz]	[uK-rtSec]	[uK-rtSec]	[((uK ² s) ⁻¹]	[uK-arcmin]
MP1	93.0 +/- 0.0	0.376 +/- 0.000	6504	35.30 +/- 0.00	1.89 +/- 0.00	20.93 +/- 0.00	10.88 +/- 0.00	0.00 +/- 0.00	23.59 +/- 0.00	283.1 +/- 0.0	4.60 +/- 0.00	0.0473 +/- 0.0000	7.6 +/- 0.0
MP2	145.0 +/- 0.0	0.276 +/- 0.000	6504	52.80 +/- 0.00	3.49 +/- 0.00	35.79 +/- 0.00	14.77 +/- 0.00	0.00 +/- 0.00	38.72 +/- 0.00	375.2 +/- 0.0	5.71 +/- 0.00	0.0306 +/- 0.0000	9.5 +/- 0.0
Total			13008								3.98 +/- 0.00	0.0780 +/- 0.0000	5.9 +/- 0.0

Figure 5: Simons Observatory V3 Large Aperture Telescope MF Camera sensitivity.

5.2 Optical Power Table

BoloCalc also outputs tables of optical powers, including the the following information in columns:

- Element
 - Name of optical element, as defined within `optics.txt` for this camera
- Power from Sky
 - Power incident on this optical element from the sky side
- Power to Detect
 - Power emitted from this optical element which is seen by the detector
- Cumulative Eff
 - Cumulative Efficiency between this optical emenet and the detector

5.2.1 Experiment/Telescope/Camera/opticalPower.txt

Optical powers are listed for each channel in the camera, formatted as one table per frequency. Below is an example for the V3 Large Aperture Telescope Mid-Frequency camera.

***** MF 1 *****				***** MF 2 *****			
Element	Power from Sky	Power to Detect	Cumulative Eff	Element	Power from Sky	Power to Detect	Cumulative Eff
	[pW]	[pW]			[pW]	[pW]	
CMB	0.00 +/- 0.00	0.08 +/- 0.00	0.152 +/- 0.000	CMB	0.00 +/- 0.00	0.07 +/- 0.00	0.221 +/- 0.000
ATM	0.52 +/- 0.00	0.68 +/- 0.00	0.157 +/- 0.000	ATM	0.33 +/- 0.00	1.06 +/- 0.00	0.229 +/- 0.000
Primary	4.81 +/- 0.00	0.79 +/- 0.00	0.163 +/- 0.000	Primary	4.93 +/- 0.00	1.42 +/- 0.00	0.238 +/- 0.000
Mirror	9.48 +/- 0.00	0.04 +/- 0.00	0.164 +/- 0.000	Mirror	10.69 +/- 0.00	0.18 +/- 0.00	0.239 +/- 0.000
Mirror	9.72 +/- 0.00	0.04 +/- 0.00	0.164 +/- 0.000	Mirror	11.38 +/- 0.00	0.18 +/- 0.00	0.240 +/- 0.000
Window	9.96 +/- 0.00	0.11 +/- 0.00	0.166 +/- 0.000	Window	12.07 +/- 0.00	0.35 +/- 0.00	0.245 +/- 0.000
IRShader1	10.45 +/- 0.00	0.02 +/- 0.00	0.167 +/- 0.000	IRShader1	13.27 +/- 0.00	0.04 +/- 0.00	0.246 +/- 0.000
IRShader2	10.58 +/- 0.00	0.02 +/- 0.00	0.167 +/- 0.000	IRShader2	13.42 +/- 0.00	0.04 +/- 0.00	0.246 +/- 0.000
IRShader1	10.70 +/- 0.00	0.02 +/- 0.00	0.167 +/- 0.000	IRShader1	13.55 +/- 0.00	0.03 +/- 0.00	0.246 +/- 0.000
IRShader2	10.80 +/- 0.00	0.02 +/- 0.00	0.167 +/- 0.000	IRShader2	13.67 +/- 0.00	0.03 +/- 0.00	0.246 +/- 0.000
AluminaF	10.89 +/- 0.00	0.00 +/- 0.00	0.171 +/- 0.000	AluminaF	13.76 +/- 0.00	0.01 +/- 0.00	0.252 +/- 0.000
IRShader1	10.69 +/- 0.00	0.01 +/- 0.00	0.171 +/- 0.000	IRShader1	13.52 +/- 0.00	0.01 +/- 0.00	0.252 +/- 0.000
IRShader2	10.71 +/- 0.00	0.00 +/- 0.00	0.171 +/- 0.000	IRShader2	13.54 +/- 0.00	0.01 +/- 0.00	0.252 +/- 0.000
AluminaF	10.72 +/- 0.00	0.00 +/- 0.00	0.175 +/- 0.000	AluminaF	13.56 +/- 0.00	0.00 +/- 0.00	0.258 +/- 0.000
LowPass1	10.51 +/- 0.00	0.03 +/- 0.00	0.186 +/- 0.000	LowPass1	13.29 +/- 0.00	0.05 +/- 0.00	0.274 +/- 0.000
Lens1	10.06 +/- 0.00	0.00 +/- 0.00	0.188 +/- 0.000	Lens1	12.69 +/- 0.00	0.00 +/- 0.00	0.279 +/- 0.000
LowPass1	9.94 +/- 0.00	0.00 +/- 0.00	0.200 +/- 0.000	LowPass1	12.49 +/- 0.00	0.00 +/- 0.00	0.296 +/- 0.000
Lens2	9.35 +/- 0.00	0.00 +/- 0.00	0.203 +/- 0.000	Lens2	11.75 +/- 0.00	0.00 +/- 0.00	0.302 +/- 0.000
Aperture	9.23 +/- 0.00	0.02 +/- 0.00	0.574 +/- 0.000	Aperture	11.55 +/- 0.00	0.00 +/- 0.00	0.571 +/- 0.000
LowPass1	3.30 +/- 0.00	0.00 +/- 0.00	0.611 +/- 0.000	LowPass1	6.10 +/- 0.00	0.00 +/- 0.00	0.608 +/- 0.000
LowPass2	3.10 +/- 0.00	0.00 +/- 0.00	0.650 +/- 0.000	LowPass2	5.74 +/- 0.00	0.00 +/- 0.00	0.646 +/- 0.000
Lens3	2.91 +/- 0.00	0.00 +/- 0.00	0.659 +/- 0.000	Lens3	5.39 +/- 0.00	0.00 +/- 0.00	0.658 +/- 0.000
LowPass1	2.88 +/- 0.00	0.00 +/- 0.00	0.701 +/- 0.000	LowPass1	5.30 +/- 0.00	0.00 +/- 0.00	0.700 +/- 0.000
Detector	2.70 +/- 0.00	0.00 +/- 0.00	1.001 +/- 0.000	Detector	4.98 +/- 0.00	0.00 +/- 0.00	1.000 +/- 0.000

Figure 6: Simons Observatory V3 Large Aperture Telescope Mid-Frequency camera optical power tables. Each table is labeled by “[camera name] + [band name].”