

A Catalogue of 1.58 Million Clusters of Galaxies from the DESI Legacy Survey

Z. L. Wen and J. L. Han (2024)

Background

(In which Joe speed reviews 3 older papers)

Context

- Clusters are big, biggest virialised things going
- We need to be able to find and characterise clusters
- This is an optical approach
- Culmination of over a decade of work

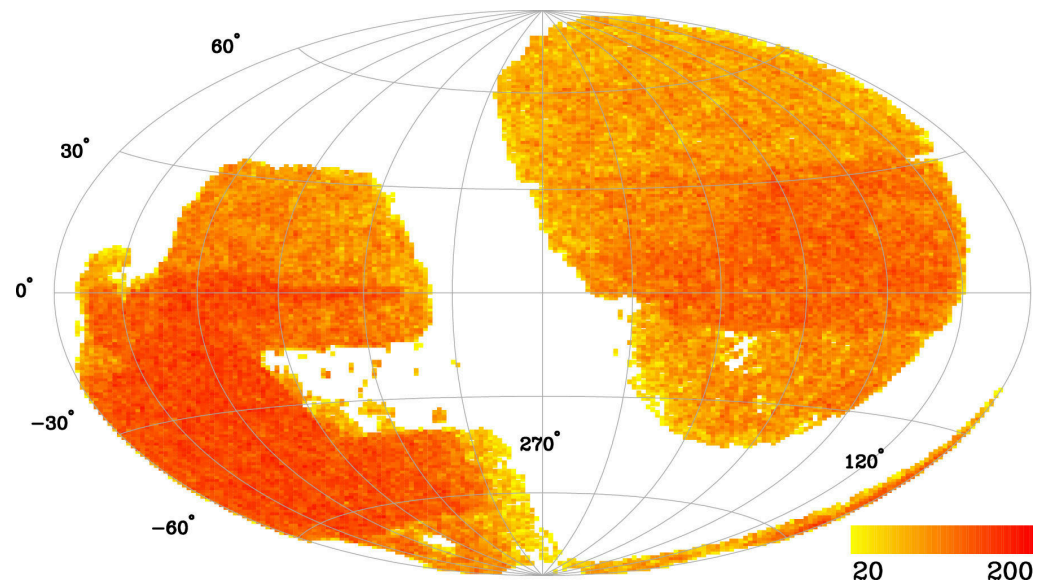


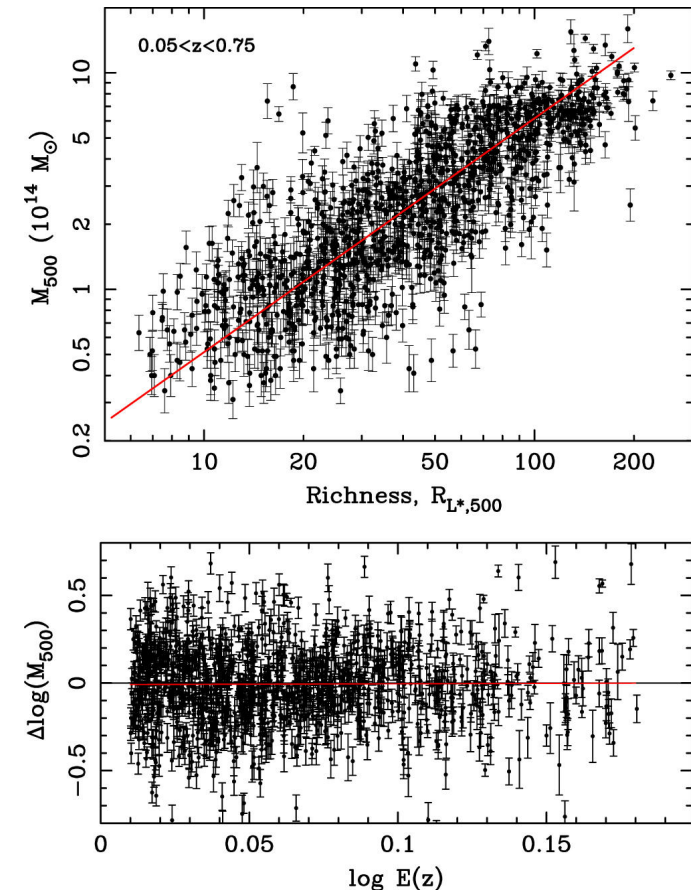
Figure 1: Density map of clusters from Wen and Han (2024, Fig. 6)

Wen and Han (2015) – Calibration

- Calibrated a relationship between r_{500} and $L_{1 \text{ Mpc}}$
- Established **richness** as an optical mass proxy:

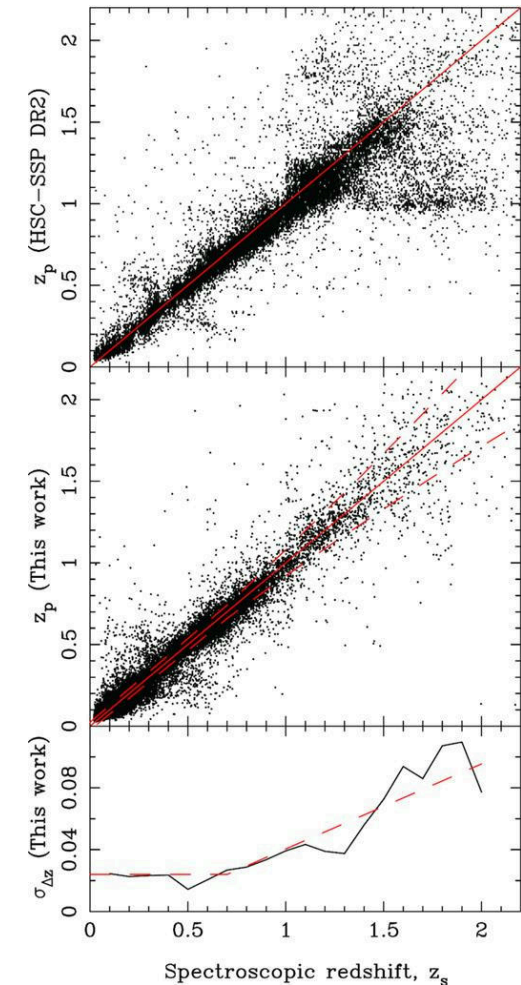
$$\lambda_{*,500} = \frac{L_{500}}{L_*} E(z)^{1.4}$$

- This is redshift independent & a good proxy



Wen and Han (2021) – Redshifts

- Combines spectroscopic and multi-band imaging surveys
- Places galaxies with spectro- z in colour space
- Uses a **nearest neighbour** algorithm to estimate the photo- z of galaxies only in imaging survey



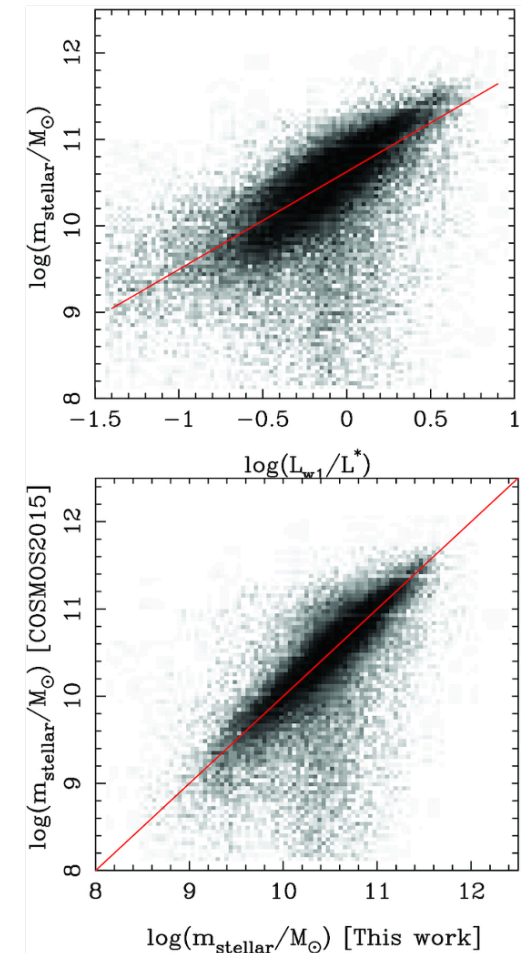
Wen and Han (2021) – Masses

- Links stellar mass and luminosity:

$$\log\left(\frac{m_{\text{stellar}}}{M_{\odot}}\right) = \gamma \log\left(\frac{L_{\text{W1}}}{L_{*}}\right) + f(z, Z)$$

- Uses this to get a mass based **richness** similar to Wen and Han (2015):

$$\lambda_{500} = m_{500,\text{stellar}} \frac{(1+z)^{0.21}}{m_{*,\text{stellar}}}$$



Wen and Han (2022) – Extending Deeper

- Takes what they were doing before and uses **DES** to find clusters to $z = 1.5$
- ...
- Not much else different but proves validity of methods to deeper data

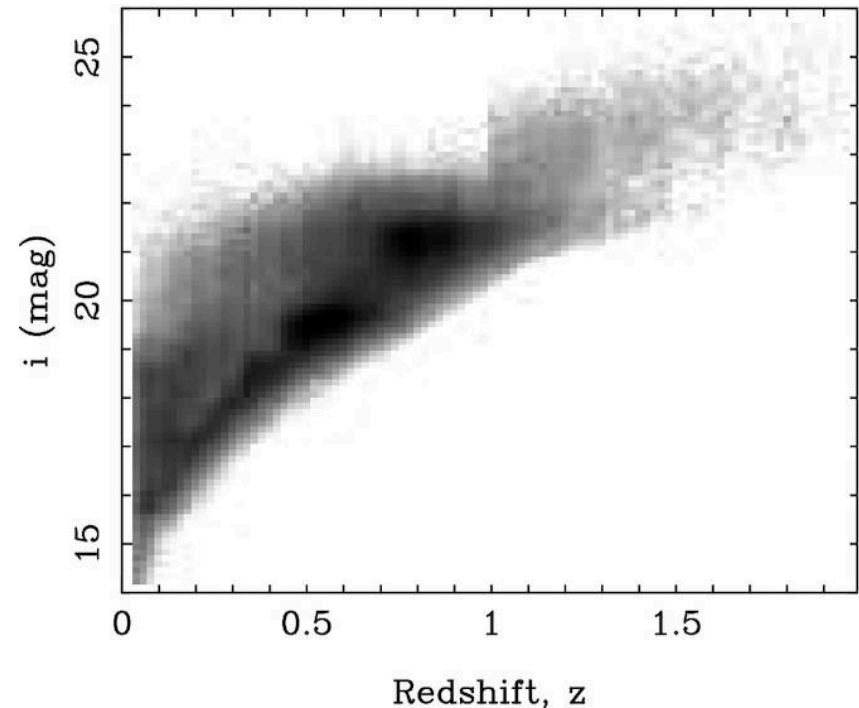


Figure 5: i -band magnitudes of the training sample as a function of redshift. Taken from Wen and Han (2022, Fig. 1)

The Actual Paper

*(Trust me, it's **definitely** a pre-print)*

The Initial Data Processing

- Using **DESI** Legacy Imaging Surveys as the photometric base
- Same processes as before for finding redshifts, with spectro- z from past work
- Slight tweak to finding m_{stellar} , using $r - z_m$ colour instead of W1 luminosity

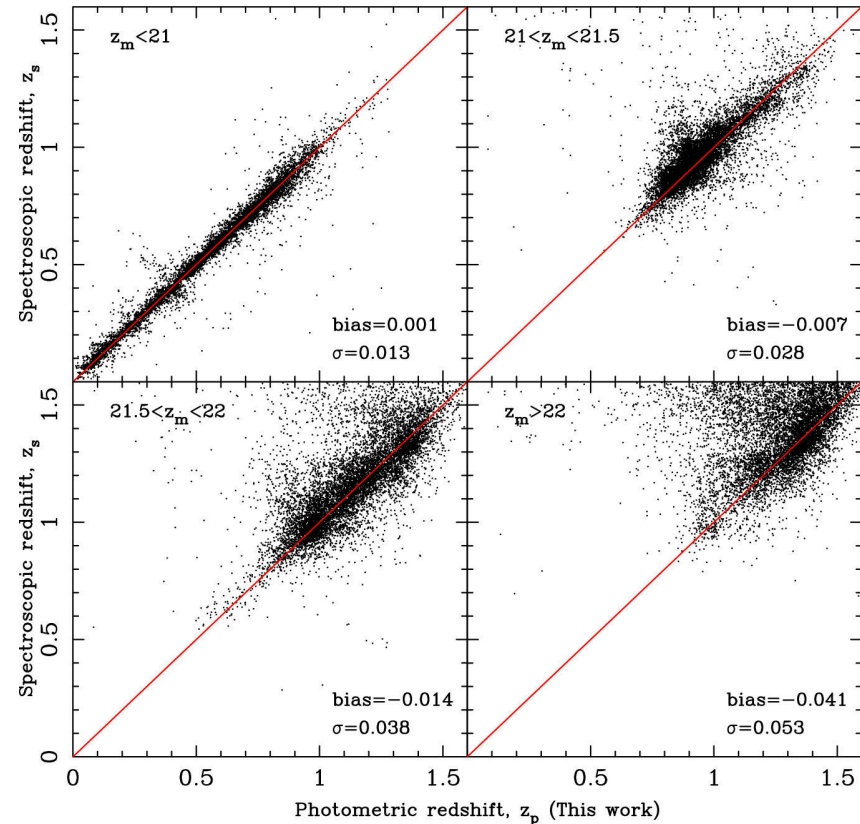


Figure 6: z_m -band magnitude binned comparisons of spectro- and photo- z s. From Wen and Han (2024, Fig. 1)

The Initial Data Processing

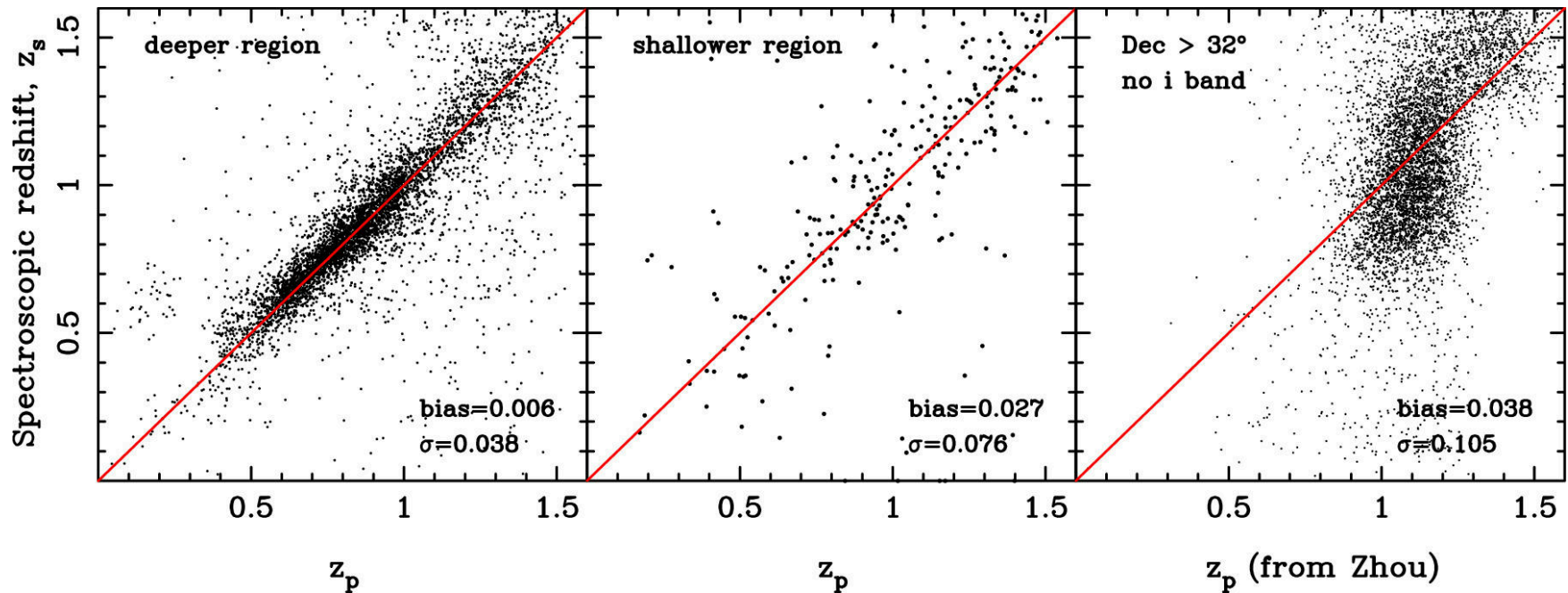


Figure 7: Comparisons of Wen and Han (2024) photo- z and those published by DESI in Zhou et al. (2021) without i -band mags

Finding Clusters

- Looking for overdensity in redshifts
- Take slices on **candidate “BCGs”** defined with half slice thickness:

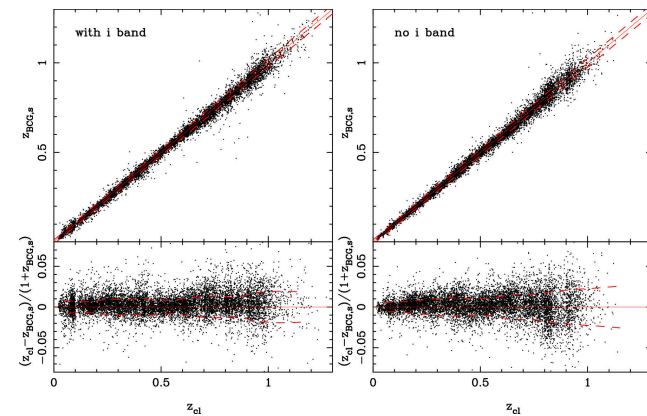
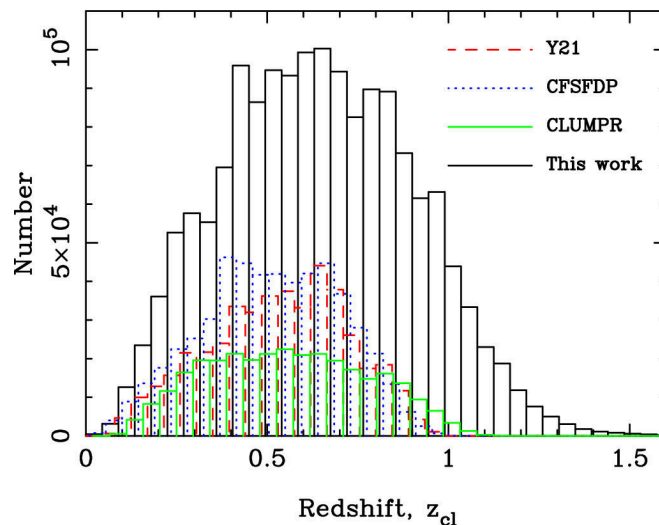
$$\Delta z = \begin{cases} 0.04(1 + z) & \text{for } z \leq 0.7 \\ 0.15z - 0.037 & \text{for } z > 0.7 \end{cases}$$

- Only using massive clusters ($M_* \geq 10^{10} M_\odot$)
- Use the equations calibrated before to find cluster radii and richness
- Define a cluster when $\lambda_{500} \geq 10$ **and** $N_{\text{gal}} \geq 6$

Cluster Redshifts

Defined in one of the following ways:

1. The **spectroscopic** redshift of the BCG, if available
2. Available spectroscopic redshifts of other galaxies, if within $0.025(1 + z)$ of cluster photo- z
3. Unclear, but I think using the average photo- z of members as in Wen and Han (2022)



Why do I care?

(Yeah, why do you? Aren't you an X-ray astronomer?)

Bibliography

Wen, Z. L., Han, J. L., 2024. A Catalog of 1.58 Million Clusters of Galaxies Identified from the DESI Legacy Imaging Surveys. The Astrophysical Journal Supplement Series 272, 39.. <https://doi.org/10.3847/1538-4365/ad409d>

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