

# DISCO Summary

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## 1 Solutions to the Disco Primer

Fluids are defined by spacetime dependent quantities: mass density  $\rho$ , velocity  $v$ , and pressure  $P$ . In case of the accretion disks of a black hole, we have 5 real variables:  $\rho$ ,  $P$ , and 3 components of  $v$ . We also have 5 equations for the 5 variables along with conservation laws:

1. Conservation of particle number
2. Conservation of momentum
3. Conservation of energy.

These conservation laws give us the equations needed to solve: the Continuity Equation, Euler Equations, and the Energy Equation.

$$\begin{aligned}\partial_t \rho + \nabla \cdot (\rho v) &= 0 \\ \partial_t (\rho v) + \nabla \cdot (\rho v v + P) &= \rho g \\ \partial_t \left( \frac{1}{2} \rho v^2 + e \right) + \nabla \cdot \left[ \left( \frac{1}{2} \rho v^2 + e + P \right) v \right] &= \rho v \cdot g - \dot{q}\end{aligned}\tag{1}$$

The goal for DISCO code is to solve these equations for a hydrodynamic system. Below I have documented the process of which I preliminarily solve these equations using analytical methods to give rise to a simple disk model.

### 1.1 Steady and axisymmetric

We assume the disk model to be long-lived and axisymmetric. That is, the disk will look the same from any angle at any moment of time. Thus, we set derivatives of time and angle  $\phi$  in cylindrical coordinates to 0. Here we will have four ordinary differential equations:

$$\begin{aligned}
\text{Continuity: } & \frac{1}{r} \partial_r (r \Sigma V_r) = 0 \\
\text{Radial Momentum: } & \frac{1}{r} \partial_r \left( r \Sigma V_r^2 + r \Pi - 2r \Sigma \nu \left( \frac{2}{3} \partial_r V_r - \frac{1}{3} \partial_\phi \Omega \right) \right) = 0 \\
\text{Angular Momentum: } & \frac{1}{r} \partial_r \left( r^3 \Sigma V_r \Omega - 2r \Sigma \nu \left( \frac{1}{2} r^2 \partial_r \Omega \right) \right) = 0 \\
\text{Energy: } & \frac{1}{r} \partial_r \left( r \left( \frac{1}{2} \Sigma v^2 + \Sigma \epsilon + \Pi \right) V_r - 2r \Sigma \nu (\sigma \cdot v) r \right) = \Sigma v \cdot g - \dot{q} \quad (2)
\end{aligned}$$

where  $\nu$  is viscosity and  $\dot{q}$  is a cooling term,  $\Omega$  is angular velocity, and

$$\begin{aligned}
v \cdot g &= -\frac{GM}{r^2} \cdot v \\
\Sigma &= \int \rho dz. \quad (3)
\end{aligned}$$

## 1.2 The razor thin limit

We are assuming that the disk is razor thin and inviscid. Here we take the approximations of  $H \simeq 0$  and  $\nu = 0$ . We also set the sounds speed  $c_s = 0$ , as well as  $\Pi$  and  $\epsilon$ . Here we find the solutions of  $V_r(r)$  and  $\Omega(r)$ .

We can do so by taking the insides of the first and third equation as constants, and solve for the other two equations using results obtained from it. Eventually, we get one of the solutions similar to Saturn's rings:  $V_r = 0$  and

$$\Omega = \sqrt{\frac{GM}{r^3}}. \quad (4)$$

As we take the radial velocity as 0, we are assuming that the gas stays in a steady orbit without moving inward or outward. This makes sense in our razor thin and axisymmetric approximation but it is unrealistic since accretion does not occur with  $V_r = 0$ .

## 1.3 The Accretion Rate and Torque

We define the accretion rate  $\dot{M}$  as the rate at which matter passes a particular radius moving toward the central object:  $\dot{M} = -2\pi r \Sigma v_r$ . From the continuity equation from step 1, we can derive

$$\partial_r (r \Sigma V_r) = 0 \quad (5)$$

and  $\dot{M} = -2\pi r \Sigma V_r = \text{constant}$ . This means that for a disk that is steady and axisymmetric, the accretion rate is constant across all radii. Later, we will use the expression of  $V_r$  to solve for other quantities.

$$V_r = \frac{\dot{M}}{-2\pi r \Sigma} \quad (6)$$

Similarly, we define the torque through the disk  $\dot{J}$  as the rate at which the central object accretes angular momentum, which is equal to  $2\pi$  times the angular momentum flux. From the angular momentum equation from step 1, we have

$$\frac{1}{r}\partial_r(r^3\Sigma V_r\Omega - r^3\Sigma\nu\partial_r\Omega) = \partial_r\dot{J} = 0. \quad (7)$$

Using  $V_r$  from above, we get

$$\dot{J} = \dot{M}r^2\Omega + 2\pi r^3\Sigma\nu\partial_r\Omega. \quad (8)$$

Rearrange, we have an expression for  $\Sigma$

$$\Sigma = \frac{\dot{J} - \dot{M}r^2\Omega}{2\pi r^3\nu\partial_r\Omega}. \quad (9)$$

#### 1.4 An accretion disk

So far, we have two algebraic equations involving  $\dot{M}$  and  $\dot{J}$  and the radial momentum and energy equations. It is time to solve the final radial momentum equation assuming the disk is thin  $H \ll r$  and the radial velocity subsonic  $|v_r| < c_s$ . The centripetal acceleration term should balance the term with the largest other term, in this case,  $r\Sigma\frac{GM}{r^2}$ . This confirms the expression for angular velocity  $\Omega = \sqrt{\frac{GM}{r^3}}$ .

After simplifying, we obtain the final solution for  $\Sigma$

$$\Sigma = \frac{-1}{3\pi\nu} \left( \frac{\dot{J}}{\sqrt{GM}r} - \dot{M} \right), \quad (10)$$

and for  $V_r$

$$V_r = -\frac{3\nu}{2r}. \quad (11)$$

From here, we have obtained a series of constants and expressions that will help the DISCO code solve the hydrodynamic equations.

We define function  $f$ :

$$f = 1 - \frac{\dot{J}}{\dot{M}\sqrt{GM}r}. \quad (12)$$

Finally, some expressions in terms of  $f$ :

$$\begin{aligned} \Sigma &= \frac{\dot{M}f}{3\pi\nu} \\ V_r &= \frac{3\nu}{-2rf} \\ \Omega &= \sqrt{\frac{GM}{r^3}}. \end{aligned} \quad (13)$$

## 2 Instructions on using the DISCO code

Now that we have analytically solved solutions of a simple disk model, it is time that we put these variables into the disco code. There are many templates for different simulations performed in disco, but in our case we only focus on thin accretion disk for singular or binary black holes.

In order to run the code, we need to provide initial conditions and the instructions for the computer to run the simulations. I have listed below some of the useful tips for this project running on disco.

The *Makefile\_optin* file contains the initial condition file we will use. We can specify the file name after

`INITIAL = {FILENAME}.`

The file we are putting into the initial section of *Makefile\_optin* is stored in the *initial* directory. In this *c* file, we can specify the disk model we are using. We can load quantities such as viscosity and adiabatic index into the file using *setICparams()* function. This is the place we can put the simple disk model we derived above into the simulation. Specifically, these formulas are in this file:

$$\begin{aligned} f &= 1 - \frac{\dot{J}}{\dot{M}\sqrt{GM}r} \\ \Sigma &= \frac{\dot{M}f}{3\pi\nu} \\ V_r &= \frac{3\nu}{-2rf} \\ \Omega &= \sqrt{\frac{GM}{r^3}}. \end{aligned} \tag{14}$$

If we want to have a circumbinary black hole simulation, we can add mass profiles in the *planet* directory.

In the file *inpar*, we can specify parameters such as the periods of rotation, time step, grid dimensions, hydro parameters etc. In our practice so far, we only run a small number of checkpoints and observe the simulation through *vdisco*, a visual platform of which the checkpoints of the simulation can be visualized.

In *vdisco*, the numbers on your keyboard switches among the variables:

- 1 Surface density ( $\rho$  or  $\Sigma$  in derivations above)
- 2 Pressure  $P$
- 3 Radial velocity  $V_r$
- 4 Angular velocity  $\Omega$
- 5 Vertical velocity  $V_z$  (0 most of the times)
- 6 Passive scalar for visualization

The checkpoints can be scrolled through pressing [ and ] keys.