

Calculation of Moon Phases and 24 Solar Terms

月相和二十四節氣的計算

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This document explains the method used to compute the times of the moon phases and 24 solar terms. These times are important in the calculation of the Chinese calendar. See this page for an introduction to the 24 solar terms, and this page for an introduction to the Chinese calendar calculation.

Computation of accurate times of moon phases and 24 solar terms is complicated, but today all the necessary resources are freely available. Anyone familiar with numerical computation and computer programming can follow the procedure outlined in this document to do the computation.

Before stating the procedure, it is useful to have a basic understanding of the basic concepts behind the computation. I assume that readers are already familiar with the important astronomy concepts mentioned on this page. In Section 1, I briefly introduce the barycentric dynamical time (TDB) used in modern ephemerides, and its connection with terrestrial time (TT) and international atomic time (TAI). Readers who are not familiar with general relativity do not have to pay much attention to the formulas there. Section 2 introduces the various coordinate systems used in modern astronomy. Section 3 lists the formulas for computing the IAU 2006/2000A precession and nutation matrices.

One important component in the computation of accurate times of moon phases and 24 solar terms is an accurate ephemeris of the Sun and Moon. I use the ephemerides developed by the Jet Propulsion Laboratory (JPL) for the computation. Section 4 introduces the JPL ephemerides and describes how they can be downloaded and used to compute the positions and velocities of solar system objects. One of the main goals of the JPL ephemerides is for spacecraft navigation. Positions and velocities of the solar system objects are given in the international celestial reference system (ICRS). Hence, these data need to be transformed to the ecliptic coordinate system suitable for the computation of the moon phases and 24 solar terms. Section 5 describes the light-time correction and aberration of light. Section 6 provides a step-by-step procedure for computing the apparent geocentric longitude of the Sun and Moon from the data computed from the JPL ephemerides. Section 7 describes the numerical method for computing the TDB times of the moon phases and 24 solar terms from the apparent geocentric longitudes of the Sun and Moon. Finally, Section 8 describes the conversion of times between TDB and UTC+8 necessary for calendar calculation. For a more detailed and comprehensive introduction to the concepts mentioned in this document, I recommend the book *Explanatory Supplement to the Astronomical Almanac* by [Urban & Seidelmann 2013].

1 Ephemeris Time and Dynamical Time

From the 17th century to the late 19th century, planetary ephemerides were calculated using time scales based on Earth's rotation. It was assumed that Earth's rotation was uniform. As the precision of astronomical measurements increased, it became clear that Earth's rotation is not uniform. Ephemeris time (ET) was introduced to ensure a uniform time for ephemeris calculations. It was defined by the orbital motion of the Earth around the Sun instead of Earth's spin motion. However, a more precise definition of times is required when general relativistic effects need to be included in ephemeris calculations.

In general relativity, the passage of time measured by an observer depends on the spacetime trajectory of the observer. To calculate the motion of objects in the solar system, the most convenient time is a coordinate time, which does not depend on the motion of any object but is defined through the spacetime metric. In the *Barycentric Celestial Reference System* (BCRS), the spacetime coordinates are (t, x^i) ($i = 1, 2, 3$). Here the time coordinate t is called the *barycentric coordinate time* (TCB). The spacetime metric for the solar system can be written as [see Eq. (2.38) in [Urban & Seidelmann 2013]]

$$ds^2 = - \left(1 - \frac{2w}{c^2} + \frac{2w^2}{c^4} \right) d(ct)^2 - \frac{4w_i}{c^3} d(ct) dx^i + \delta_{ij} \left[1 + \frac{2w}{c^2} + O(c^{-4}) \right] dx^i dx^j, \quad (1)$$

where sum over repeated indices is implied. The scalar potential w reduces to the Newtonian gravitational potential $-\Phi$ in the Newtonian limit, where

$$\Phi(t, \mathbf{x}) = -G \int d^3x' \frac{\rho(t, \mathbf{x}')}{|\mathbf{x} - \mathbf{x}'|} \quad (2)$$

and ρ is the mass density. The vector potential w_i satisfies the Poisson equations with the source terms proportional to the momentum density.

TCB can be regarded as the proper time measured by an observer far away from the solar system and is stationary with respect to the solar system barycenter. The equations of motion for solar system objects can be derived in the post-Newtonian framework. The result is a system of coupled differential equations and can be integrated numerically. In this framework, TCB is the natural choice of time parameter for planetary ephemerides. However, since most measurements are carried out on Earth, it is also useful to set up a coordinate system with origin at the Earth's center of mass. In the *geocentric celestial reference system* (GCRS), the spacetime coordinates are (T, X^i) , where the time parameter T is called the *geocentric coordinate time* (TGC). GCRS is comoving with Earth's center of mass in the solar system, and its spatial coordinates X^i are chosen to be kinematically non-rotating with respect to the barycentric coordinates x^i . The coordinate time TCG is chosen so that the spacetime metric has a form similar to equation (1). Since GCRS is comoving with the Earth, it is inside the potential well of the solar system. As a result, TCG elapses *slower* than TCB because of the combined effect of gravitational time dilation and special relativistic time dilation. The relation between TCB and TCG is given by [equation (3.25) in [Urban & Seidelmann 2013]]

$$\text{TCB} - \text{TCG} = c^{-2} \left[\int_{t_0}^t \left(\frac{v_e^2}{2} - \Phi_{\text{ext}}(\mathbf{x}_e) \right) dt + \mathbf{v}_e \cdot (\mathbf{x} - \mathbf{x}_e) \right] + O(c^{-4}), \quad (3)$$

where \mathbf{x}_e and \mathbf{v}_e are the barycentric position and velocity of the Earth's center of mass, and \mathbf{x} is the barycentric position of the observer. The external potential Φ_{ext} is the Newtonian gravitational

potential of all solar system bodies apart from the Earth. The constant time t_0 is chosen so that TCB=TCG=ET at the epoch 1977 January 1, 0h TAI.

Since the definition of TCG involves only external gravity, TCG's rate is faster than TAI's because of the relativistic time dilation caused by Earth's gravity and spin. The *terrestrial time* (TT), formerly called the terrestrial dynamical time (TDT), is defined so that its rate is the same as the rate of TAI. The rate of TT is slower than TCG by $-\Phi_{\text{eff}}/c^2$ on the geoid (Earth surface at mean sea level), where $\Phi_{\text{eff}} = \Phi_E - v_{\text{rot}}^2/2$ is the sum of Earth's Newtonian gravitational potential and the centrifugal potential. Here v_{rot} is the speed of Earth's spin at the observer's location. The value of Φ_{eff} is constant on the geoid because the geoid is defined to be an equipotential surface of Φ_{eff} . Thus, $d\text{TT}/d\text{TCG} = 1 - L_G$ and L_G is determined by measurements to be $L_G = 6.969290134 \times 10^{-10}$. Therefore, TT and TCG are related by a linear relationship [equation (3.27) in [Urban & Seidelmann 2013]]:

$$\text{TT} = \text{TCG} - L_G(\text{JD}_{\text{TCG}} - 2443144.5003725) \cdot 86400 \text{ s}, \quad (4)$$

where JD_{TCG} is TCG expressed as a Julian date (JD). The constant 2443144.5003725 is chosen so that TT=TCG=ET at the epoch 1977 January 1 0h TAI (JD = 2443144.5003725). Since the rate of TT is the same as that of TAI, the two times are related by a constant offset:

$$\text{TT} = \text{TAI} + 32.184 \text{ s}. \quad (5)$$

The offset arises from the requirement that TT match ET at the chosen epoch.

TCB is a convenient time for planetary ephemerides, whereas TT can be measured directly by atomic clocks on Earth. The two times are related by equations (3) and (4) and must be computed by numerical integration together with the planetary positions. TT is therefore not convenient for planetary ephemerides. The *barycentric dynamical time* (TDB) is introduced to approximate TT. It is defined to be a linear function of TCB and is set as close to TT as possible. Since the rates of TT and TCB are different and are changing with time, TT cannot be written as a linear function of TCB. The best we can do is to set the rate of TDB the same as the rate of TT averaged over a certain time period, so that there is no long-term secular drift between TT and TDB over that time period. The resulting deviation between TDB and TT has components of periodic variation caused by the eccentricity of Earth's orbit and the gravitational fields of the Moon and planets. TDB is now defined by the IAU 2006 resolution 3 as

$$\text{TDB} = \text{TCB} - L_B(\text{JD}_{\text{TCB}} - 2443144.5003725) \cdot 86400 \text{ s} - 6.55 \times 10^{-5} \text{ s}, \quad (6)$$

where $L_B = 1.550519768 \times 10^{-8}$ and JD_{TCB} is TCB expressed as a Julian date (JD). The value of L_B can be regarded as $1 - d\text{TT}/dt$ averaged over a certain time period.

TDB is a successor of ET. It is practically equivalent to the JPL ephemeris time argument T_{eph} used by the Jet Propulsion Laboratory to calculate high-precision ephemerides of the Sun, Moon and planets. The relationship between TT and TDB can be written as (Figure 3.2 in [Urban & Seidelmann 2013])

$$\begin{aligned} \text{TDB} = & \text{TT} + 0.001658 \text{ s} \sin(g + 0.0167 \sin g) \\ & + \text{lunar and planetary terms of order } 10^{-5} \text{ s} \\ & + \text{daily terms of order } 10^{-6} \text{ s}, \end{aligned} \quad (7)$$

where g is the mean anomaly of Earth in its orbit (and hence $g + 0.0167 \sin g$ is the eccentric anomaly since 0.0167 is Earth's orbital eccentricity). A more detailed expression is given by equation (2.6) in [Kaplan 2005]. The difference between TDB and TT remains under 2 ms for several millennia around the present epoch. Thus, I treat them as the same for the calendar calculation.

2 Celestial Coordinate Systems

2.1 International Celestial Reference System (ICRS)

As mentioned in Section 1, the solar system metric can be written in the Barycentric Celestial Reference System (BCRS). The origin of the BCRS spatial coordinates is at the solar system barycenter, i.e. the center of mass of the solar system. However, BCRS is a dynamical concept. The statement of “we use BCRS” in general relativity is equivalent to the statement “we use barycentric inertial coordinates” in Newtonian mechanics. BCRS does not define the orientation of the coordinate axes.¹

The International Celestial Reference System (ICRS) is a kinematical concept. Its origin is at the solar system barycenter. The ICRS axes are intended to be fixed with respect to space. They are determined based on hundreds of extra-galactic radio sources, mostly quasars, distributed around the sky. The ICRS axes are aligned with the equatorial system based on the J2000.0 mean equator and equinox to within 17.3 milliarcseconds. The x -axis of the ICRS points in the direction of the mean equinox of J2000.0. The z -axis points very close to the mean celestial north pole of J2000.0, and the y -axis is 90° to the east of the x -axis on the ICRS equatorial plane. So the ICRS is a right-handed rectangular coordinate system. The ICRS can be transformed to the equatorial system of J2000.0 by the *frame bias matrix*, which will be discussed below.

It is assumed that the distant quasars and extragalactic radio sources do not rotate with respect to asymptotically flat reference systems like BCRS. In principle this assumption should be checked by testing if the motion of the solar system objects is compatible with the equation of motion based on BCRS, with no Coriolis and centrifugal forces. So far no deviations have been noticed.

2.2 Geocentric Celestial Reference System (GCRS)

The origin of Geocentric Celestial Reference System (GCRS) is at the center of mass of the Earth. Ignoring general relativistic correction, the GCRS spatial coordinates X^i are related to the BCRS spatial coordinates x^i by

$$X^i = x^i - x_E^i, \quad (8)$$

where x_E^i are the BCRS coordinates of Earth’s center of mass. Relativistic correction adds terms of order $(v_E/c)^2 \sim 10^{-8} = 0.002''$. The Sun moves along the ecliptic with a rate of about $0.04''$ per second. The Moon moves faster, with a rate of about $0.5''$ per second. Thus ignoring $(v_E/c)^2$ correction will lead to an error of about 0.05 seconds in the computation of the times of 24 solar terms and about 0.004 seconds in the times of moon phases. These errors are much smaller than the one-second accuracy required by the official document [GB/T 33661-2017].

Like BCRS, GCRS is a dynamical concept. In the following, I will also use GCRS to refer to the coordinate system whose origin is at the geocenter and whose axes are oriented to the same directions as those of the ICRS.

¹IAU 2006 Resolution B2 recommends that the BCRS definition is completed with the following: “For all practical applications, unless otherwise stated, the BCRS is assumed to be oriented according to the ICRS axes. The orientation of the GCRS is derived from the ICRS-oriented BCRS.”

2.3 Equatorial and Ecliptic Coordinate Systems

The orientation of the coordinate axes described above are fixed in space. This is convenient for describing the positions of stars and planets. However, observations are made on Earth and hence coordinate systems with axes defined by Earth's spin or defined by Earth's orbital plane are often used to describe positions of celestial objects. The coordinate systems whose axes are defined by Earth's spin are called the *equatorial coordinate systems*, and the coordinate system whose axes are defined by Earth's orbital plane are called the *ecliptic coordinate systems*.

2.3.1 Equator, Ecliptic and Equinoxes

The *Celestial Intermediate Pole* (CIP) is the mean rotation axis of the Earth whose motion in space contains aperiodic components as well as periodic components with periods greater than two days. The motion of CIP is described by precession and nutation (see below).

The *true equator* is defined to be the plane perpendicular to the CIP that passes through Earth's center of mass. Thus, the true equator is constantly changing as a result of precession and nutation. The *mean equator* is the moving equator whose motion is prescribed only by precession.

Ecliptic generally refers to Earth's orbital plane projected onto the celestial sphere. However, Earth's orbital plane is changing because of planetary perturbation. To reduce uncertainties in the definition of the ecliptic, the International Astronomical Union (IAU) have recommended that the ecliptic be defined as the plane perpendicular to the mean orbital angular momentum vector of the Earth-Moon barycenter passing through the Sun in the BCRS.

Equator and ecliptic intercepts at two points, called the *vernal equinox* and *autumnal equinox*. The *true equinoxes* are the two points at which the true equator and true ecliptic intercepts. The *mean equinoxes* are the two points at which the mean equator and mean ecliptic intercepts.

2.3.2 Precession, Nutation and Polar Motion

Earth's spin axis changes its orientation in space because of luni-solar and planetary torques on the oblate Earth. Earth's spin axis also moves relative to the crust. This is called the *polar motion*.

The motion of Earth's spin axis is composed of *precession* and *nutation*. Precession is the components that are aperiodic or have periods longer than 100 centuries. Nutation is the components that are of shorter periods and its magnitude is much smaller. Motion with periods shorter than two days cannot be distinguished from components of polar motion arising from the tidal deformation of the Earth. They are considered as components of polar motion. Therefore, nutation is defined as the periodic components in the motion of Earth's spin axis with periods longer than two days but shorter than about 100 centuries.

The major component of precession is the rotation of Earth's spin axis about the ecliptic pole with a period of about 26,000 years. This causes the vernal equinox to move westward by 50.3'' per year. The principal period of nutation is 18.6 years, which is caused by the Moon's orbital plane precesses around the ecliptic. The amplitude of nutation is about 9''. The amplitude of the polar motion is about 0.3''.

In addition, the orbital plane of the Earth around the Sun also moves slowly because of planetary perturbation. Hence the ecliptic moves slowly in space. This is called the *precession of the ecliptic*, to be distinguished from the *precession of the equator*².

²Precession of the equator was formerly called the luni-solar precession, and precession of the ecliptic was formerly called the planetary precession. They are renamed because the terminologies are misleading. Planetary

Polar motion is not relevant in the calculation of the Chinese calendar because Chinese calendar is based on the *geocentric* positions of the Sun and Moon, i.e. positions relative to Earth's center of mass.

2.3.3 Equatorial “Of Date” Coordinates

Equatorial coordinates are based on the equator and equinox. The x -axis points to the vernal equinox. The y -axis lies in the equatorial plane and is 90° to the east of the x -axis. The z -axis points to the celestial pole. Since equator and equinoxes are moving, an epoch must be specified (e.g. J2000.0) to the coordinate system. Equatorial coordinates based on the mean equator and equinox are called the *mean equatorial coordinates*, and those based on the true equator and equinox are called the *apparent equatorial coordinates*.

One commonly used equatorial coordinate system is based on the mean equator and equinox of J2000.0 (i.e. mean equator and equinox at TDB noon on January 1, 2000). As mentioned above, there is a small misalignment between the coordinate axes of the ICRS and the axes of the J2000.0 mean equatorial system. The coordinates in the two systems are related by the *frame bias matrix* \mathbf{B} . Let \mathbf{x}_{ICRS} denotes a column vector representing the ICRS coordinates and \mathbf{x}_{2000} denotes a column vector representing coordinates of J2000.0 mean equatorial system. Then

$$\mathbf{x}_{2000} = \mathbf{B}\mathbf{x}_{\text{ICRS}}. \quad (9)$$

The frame bias matrix is given by Equation (4.4) in [Urban & Seidelmann 2013]:

$$\mathbf{B} = \begin{pmatrix} 1 - \frac{1}{2}(d\alpha_0^2 + \xi_0^2) & d\alpha_0 & -\xi_0 \\ -d\alpha_0 - \eta_0\xi_0 & 1 - \frac{1}{2}(d\alpha_0^2 + \eta_0^2) & -\eta_0 \\ \xi_0 - \eta_0d\alpha_0 & \eta_0 + \xi_0d\alpha_0 & 1 - \frac{1}{2}(\eta_0^2 + \xi_0^2) \end{pmatrix}, \quad (10)$$

where $d\alpha_0 = -14.6$ milliarcseconds, $\xi_0 = -16.617$ milliarcseconds, and $\eta_0 = -6.8192$ milliarcseconds. All of them have to be converted to radians. Substituting the numbers to the formula gives

$$\mathbf{B} = \begin{pmatrix} 0.999999999999999425 & -7.078279744 \times 10^{-8} & 8.05614894 \times 10^{-8} \\ 7.078279478 \times 10^{-8} & 0.999999999999999695 & 3.306041454 \times 10^{-8} \\ -8.056149173 \times 10^{-8} & -3.306040884 \times 10^{-8} & 0.9999999999999996208 \end{pmatrix}. \quad (11)$$

To convert \mathbf{x}_{2000} to the coordinates with respect to the true equator and equinox of date requires the multiplication of the precession matrix $\mathbf{P}(t)$ and nutation matrix $\mathbf{N}(t)$:

$$\mathbf{x}_{\text{eq}} = \mathbf{N}(t)\mathbf{P}(t)\mathbf{x}_{2000} = \mathbf{N}(t)\mathbf{P}(t)\mathbf{B}\mathbf{x}_{\text{ICRS}}. \quad (12)$$

The formulas for $\mathbf{P}(t)$ and $\mathbf{N}(t)$ will be given in Section 3 below.

2.3.4 Ecliptic “Of Date” Coordinates

The calculation of Chinese calendar requires the positions of the Sun and Moon in ecliptic coordinates of date. The ecliptic coordinate systems are based on the ecliptic and equinox. The x -axis points to the direction of the vernal equinox. The y -axis lies in the ecliptic plane and is 90° to the east of the x -axis. The z -axis points to the ecliptic pole. Therefore, the ecliptic coordinates are related to the equatorial coordinates by a rotation about the x -axis by an angle ϵ , which is

perturbation also contributes to the precession of the equator, although the magnitude is much smaller.

called the *obliquity of the ecliptic* and is the angle between the ecliptic pole and the CIP. Note that $\epsilon = \epsilon(t)$ is a function of time because of precession and nutation. The value of ϵ can be computed by equation (30) below. Let \mathbf{x}_{eq} be the column vector representing the equatorial coordinates with respect to the true equator and equinox of date, and \mathbf{x}_{ec} be the column vector representing the ecliptic coordinates with respect to the true ecliptic and equinox of date. Then

$$\mathbf{x}_{\text{ec}} = \mathbf{R}_1(\epsilon(t))\mathbf{x}_{\text{eq}} = \mathbf{R}_1(\epsilon(t))\mathbf{N}(t)\mathbf{P}(t)\mathbf{B}\mathbf{x}_{\text{ICRS}}, \quad (13)$$

where the rotation matrix is

$$\mathbf{R}_1(\epsilon(t)) = \begin{pmatrix} 1 & 0 & 0 \\ 0 & \cos \epsilon(t) & \sin \epsilon(t) \\ 0 & -\sin \epsilon(t) & \cos \epsilon(t) \end{pmatrix}. \quad (14)$$

The *ecliptic longitude* λ is defined as $\arg(x_{\text{ec}} + iy_{\text{ec}})$, where $\arg(z)$ denotes the argument of the complex number z . In other words, $\lambda = \tan^{-1}(y_{\text{ec}}/x_{\text{ec}})$ with the angle in the appropriate quadrant. In many programming languages, there is an arctangent function (e.g. `atan2` in FORTRAN, C and python) that returns the angle in the correct quadrant.

3 Precession and Nutation

3.1 Precession Matrix

Denote $\mathbf{X} = (X \ Y \ Z)^T$ the equatorial coordinates based on the mean equator and equinox at TDB time t and $\mathbf{X}_0 = (X_0 \ Y_0 \ Z_0)^T$ the equatorial coordinates based on the mean equator and equinox at J2000.0, where the superscript T denotes transpose. So \mathbf{X} and \mathbf{X}_0 are column vectors, and they are related by a 3D rotation described by the precession matrix $\mathbf{P}(t)$:

$$\mathbf{X} = \mathbf{P}(t)\mathbf{X}_0. \quad (15)$$

In August 2006, the 26th General Assembly for the International Astronomical Union passed a resolution recommending that the P03 precession theory of [Capitaine et al 2003] be used for the precession matrix. This model is referred to as the IAU 2006 precession theory in [IERS Conventions 2010]. According to this theory, the precession matrix is given by

$$\mathbf{P}(t) = \mathbf{R}_3(\chi_A)\mathbf{R}_1(-\omega_A)\mathbf{R}_3(-\psi_A)\mathbf{R}_1(\epsilon_0), \quad (16)$$

where the rotation matrix \mathbf{R}_1 is given by equation (14) above and the rotation matrix \mathbf{R}_3 is given by

$$\mathbf{R}_3(\theta) = \begin{pmatrix} \cos \theta & \sin \theta & 0 \\ -\sin \theta & \cos \theta & 0 \\ 0 & 0 & 1 \end{pmatrix}. \quad (17)$$

The angle $\epsilon_0 = 84381.406''$ is the inclination angle between the mean ecliptic and mean equator at J2000.0. The angles ψ_A , ω_A and χ_A can be found in equation (5.7) in [Kaplan 2005] or (5.39) and (5.40) in [IERS Conventions 2010]. They are given by

$$\begin{aligned} \psi_A &= 5038.481507''T - 1.0790069''T^2 - 0.00114045''T^3 + 0.000132851''T^4 - 9.51'' \times 10^{-8}T^5 \\ \omega_A &= 84381.406'' - 0.025754''T + 0.0512623''T^2 - 0.00772503''T^3 - 4.67'' \times 10^{-7}T^4 \end{aligned}$$

$$\begin{aligned}
& +3.337'' \times 10^{-7} T^5 \\
\chi_A = & 10.556403'' T - 2.3814292'' T^2 - 0.00121197'' T^3 + 0.000170663'' T^4 - 5.60'' \times 10^{-8} T^5,
\end{aligned} \tag{18}$$

where $T = (JD - 2451545)/36525$ is the Julian century from J2000.0 and JD is the TDB Julian date number. The result of the matrix multiplication in equation (16) is also written out explicitly in equation (5.10) in [Kaplan 2005]:

$$\begin{aligned}
P_{11}(t) &= C_4 C_2 - S_2 S_4 C_3 \\
P_{12}(t) &= C_4 S_2 C_1 + S_4 C_3 C_2 C_1 - S_1 S_4 S_3 \\
P_{13}(t) &= C_4 S_2 S_1 + S_4 C_3 C_2 S_1 + C_1 S_4 S_3 \\
P_{21}(t) &= -S_4 C_2 - S_2 C_4 C_3 \\
P_{22}(t) &= -S_4 S_2 C_1 + C_4 C_3 C_2 C_1 - S_1 C_4 S_3 \\
P_{23}(t) &= -S_4 S_2 S_1 + C_4 C_3 C_2 S_1 + C_1 C_4 S_3 \\
P_{31}(t) &= S_2 S_3 \\
P_{32}(t) &= -S_3 C_2 C_1 - S_1 C_3 \\
P_{33}(t) &= -S_3 C_2 S_1 + C_3 C_1
\end{aligned} \tag{19}$$

where

$$\begin{aligned}
S_1 &= \sin \epsilon_0 & C_1 &= \cos \epsilon_0 \\
S_2 &= \sin(-\psi_A) & C_2 &= \cos(-\psi_A) \\
S_3 &= \sin(-\omega_A) & C_3 &= \cos(-\omega_A) \\
S_4 &= \sin \chi_A & C_4 &= \cos \chi_A
\end{aligned} \tag{20}$$

3.2 Nutation Matrix

Nutation is computed according to the IAU 2000A Theory of Nutation with slight IAU 2006 adjustments. To construct the nutation matrix $\mathbf{N}(t)$, we need to first compute the following 14 arguments given by equations (5.43) and (5.44) in [IERS Conventions 2010].

$$\begin{aligned}
F_1 \equiv l &= \text{Mean Anomaly of the Moon} \\
&= 134.96340251^\circ + 1717915923.2178'' T + 31.8792'' T^2 \\
&\quad + 0.051635'' T^3 - 0.00024470'' T^4 \\
F_2 \equiv l' &= \text{Mean Anomaly of the Sun} \\
&= 357.52910918^\circ + 129596581.0481'' T - 0.5532'' T^2 \\
&\quad + 0.000136'' T^3 - 0.00001149'' T^4 \\
F_3 \equiv &= L - \Omega = \text{Mean Longitude of the Moon} - \Omega \\
&= 93.27209062^\circ + 1739527262.8478'' T - 12.7512'' T^2 \\
&\quad - 0.001037'' T^3 + 0.00000417'' T^4
\end{aligned} \tag{21}$$

$$\begin{aligned}
F_4 \equiv D &= \text{Mean Elongation of the Moon from the Sun} \\
&= 297.85019547^\circ + 1602961601.2090'' T - 6.3706'' T^2 \\
&\quad + 0.006593'' T^3 - 0.00003169'' T^4
\end{aligned}$$

$$F_5 \equiv \Omega = \text{Mean Longitude of the Ascending Node of the Moon}$$

$$= 125.04455501^\circ - 6962890.5431''T + 7.74722''T^2 \\ + 0.007702''T^3 - 0.00005939''T^4$$

The rest of the arguments are the mean longitudes of the 8 planets and general precession. They are given in radians as

$$\begin{aligned} F_6 &\equiv L_{\text{Mercury}} = 4.402608842 + 2608.7903141574T \\ F_7 &\equiv L_{\text{Venus}} = 3.176146697 + 1021.3285546211T \\ F_8 &\equiv L_{\text{Earth}} = 1.753470314 + 628.3075849991T \\ F_9 &\equiv L_{\text{Mars}} = 6.203480913 + 334.0612426700T \\ F_{10} &\equiv L_{\text{Jupiter}} = 0.599546497 + 52.9690962641T \\ F_{11} &\equiv L_{\text{Saturn}} = 0.874016757 + 21.3299104960T \\ F_{12} &\equiv L_{\text{Uranus}} = 5.481293872 + 7.4781598567T \\ F_{13} &\equiv L_{\text{Neptune}} = 5.311886287 + 3.8133035638T \\ F_{14} &\equiv p_A = 0.02438175T + 0.0000538691T^2 \end{aligned} \tag{22}$$

Next, the nutation in longitude $\Delta\psi$ and nutation in obliquity $\Delta\epsilon$ are given by the expressions

$$\Delta\psi = \sum_{i=1}^{1320} [A_i \sin \theta_i^A + A_i'' \cos \theta_i^A] + \sum_{i=1}^{38} [A_i' \sin \theta_i^{A'} + A_i''' \cos \theta_i^{A'}]T \tag{23}$$

$$\Delta\epsilon = \sum_{i=1}^{1037} [B_i \cos \theta_i^B + B_i'' \sin \theta_i^B] + \sum_{i=1}^{19} [B_i' \cos \theta_i^{B'} + B_i''' \sin \theta_i^{B'}]T, \tag{24}$$

where the arguments of the sine and cosine functions are given by

$$\theta_i^A = \sum_{j=1}^{14} C_{ij}^A F_j, \quad \theta_i^{A'} = \sum_{j=1}^{14} C_{ij}^{A'} F_j, \quad \theta_i^B = \sum_{j=1}^{14} C_{ij}^B F_j, \quad \theta_i^{B'} = \sum_{j=1}^{14} C_{ij}^{B'} F_j. \tag{25}$$

The coefficients A_i , A_i' , A_i'' , A_i''' , C_{ij}^A and $C_{ij}^{A'}$ are listed in the table on the IERS ftp site <ftp://tai.bipm.org/iers/conv2010/chapter5/tab5.3a.txt>. Values of A_i and A_i'' are given in the second and third columns in the first 1320 rows; A_i' and A_i''' are given in the second and third columns in the last 38 rows; C_{ij}^A are given in columns 4 to 17 in the first 1320 rows; $C_{ij}^{A'}$ are given in columns 4 to 17 in the last 38 rows. Coefficients B_i , B_i' , B_i'' , B_i''' , C_{ij}^B and $C_{ij}^{B'}$ are given in the table on the IERS ftp site <ftp://tai.bipm.org/iers/conv2010/chapter5/tab5.3b.txt>. To make sure the tables are read correctly, I list the first few terms for $\Delta\psi$ and $\Delta\epsilon$:

$$\begin{aligned} \Delta\psi &= -17.20642418'' \sin \Omega + 0.003386'' \cos \Omega \\ &\quad -1.31709122'' \sin(2F - 2D + 2\Omega) - 0.0013696'' \cos(2F - 2D + 2\Omega) + \dots \\ \Delta\epsilon &= 0.0015377'' \sin \Omega + 9.2052331'' \cos \Omega \end{aligned} \tag{26}$$

$$-0.0004587'' \sin(2F - 2D + 2\Omega) + 0.5730336'' \cos(2F - 2D + 2\Omega) + \dots \quad (27)$$

The nutation matrix \mathbf{N} is given by the expression

$$\mathbf{N} = \mathbf{R}_1(-\epsilon) \mathbf{R}_3(-\Delta\psi) \mathbf{R}_1(\epsilon_A), \quad (28)$$

where ϵ_A is the obliquity of the mean ecliptic of date and ϵ is the obliquity of the true ecliptic of date. They are given by the equations

$$\begin{aligned} \epsilon_A &= 84381.406'' - 46.836769''T - 0.0001831''T^2 + 0.00200340''T^3 \\ &\quad - 0.000000576''T^4 - 0.0000000434''T^5 \end{aligned} \quad (29)$$

$$\epsilon = \epsilon_A + \Delta\epsilon \quad (30)$$

If we are only interested in computing the components of the position and velocity associated with the true ecliptic and equinox of date, the computation can be simplified. From equation (13) we see that the transformation involves $\mathbf{R}_1(\epsilon)\mathbf{N}$. It follows from (28) and $\mathbf{R}_1(\epsilon)\mathbf{R}_1(-\epsilon) = I$ (identity matrix) that

$$\mathbf{R}_1(\epsilon)\mathbf{N} = \mathbf{R}_3(-\Delta\psi)\mathbf{R}_1(\epsilon_A) = \begin{pmatrix} \cos \Delta\psi & -\sin \Delta\psi \cos \epsilon_A & -\sin \Delta\psi \sin \epsilon_A \\ \sin \Delta\psi & \cos \Delta\psi \cos \epsilon_A & \cos \Delta\psi \sin \epsilon_A \\ 0 & -\sin \epsilon_A & \cos \epsilon_A \end{pmatrix}. \quad (31)$$

This means that it is not necessary to calculate the nutation in obliquity $\Delta\epsilon$.

The nutation in longitude $\Delta\psi$ is expressed by a sum over 1000 terms. An accuracy of $0.04''$ is necessary in order to compute the times of 24 solar terms to one-second accuracy required by the official document [GB/T 33661-2017]. It is not necessary to include all 1000+ terms since many of them are very small. One possibility is to use the IAU 2000B nutation model, which is an abridged nutation model. IAU 2000B model contains fewer than 80 terms and its deviation from IAU 2000A model is less than 1 milliarcsecond during the period 1995–2050. However, I include all of the 1000+ terms in IAU 2000A model in the calculation. Even though they substantially slow down the code, it still takes only about 50 seconds to compute the times of all moon phases (new moon, first quarter, full moon and third quarter) and all 24 solar terms in the years from 1600 to 3500 on my somewhat dated computer at home. This is still tolerable considering that the computation of the TDB times only needs to be carried out once for a given ephemeris.

4 Jet Propulsion Laboratory Development Ephemeris

Jet Propulsion Laboratory (JPL) development ephemeris is often abbreviated as JPL DE(number) or just DE(number). It refers to a particular ephemeris model developed by the JPL based on numerical integration. The main purpose of the ephemerides is for spacecraft navigation and astronomy. JPL has been improving their ephemerides since 1960s. The most recent, major upgrade was DE430 and DE431 ([Folkner et al 2014]) created in 2013. DE430 has been the basis of the *Astronomical Almanac* since 2015. DE431 is currently used on JPL's HORIZONS Web-Interface to generate ephemerides for solar-system bodies. They are among the most accurate ephemerides currently available. I use DE431 in my calculation of moon phases and 24 solar terms for my Chinese calendar website.

The DE ephemerides take into account gravitational perturbations from 343 relatively large-mass asteroids. General relativistic effects are included by using dynamical equations derived from

a parametrized post-Newtonian n -body metric. Additional accelerations arising from non-spherical effects of extended bodies including the Earth, Moon and Sun are also included.

The major difference between DE430 and DE431 is that DE430 includes a damping term between the Moon’s liquid core and solid mantle that gives the best fit to observation data but that is not suitable for backward integration of more than a few centuries. DE431 is similar to DE430 but without the core/mantle damping term, so the lunar position is less accurate than in DE430 for times near the current epoch, but is more suitable for times more than a few centuries in the past. DE431 covers years from -13,200 to +17,191, whereas DE430 covers years from 1550 to 2650.

4.1 Downloading and Reading the JPL Ephemerides

JPL’s DE ephemerides are released in files containing the Chebyshev polynomials fit to the rectangular positions and velocities of the Sun, Moon and planets. The coefficients of Chebyshev polynomials have a time dependence, which is first broken into 32-day intervals, and then further into a variable number of sub-intervals, depending on the object. The time intervals and orders of the Chebyshev polynomials are selected so that the deviations between the interpolated coordinate values and values provided by the ephemerides are less than 0.5 millimeters ([Newhall 1989]). This level of interpolation precision is higher than the estimated accuracy of the ephemerides.

I downloaded the DE files from JPL’s ftp site <ftp://ssd.jpl.nasa.gov/pub/eph/planets/Linux/>. The main files are in binary format and require software to read and to perform interpolation. I use Project Pluto’s C source code for reading the JPL ephemerides and performing computation. For each DE ephemeris, JPL’s ftp site provides a file for checking values of the position and velocity coordinates of different solar system objects at various different times.

I downloaded the ephemerides DE405, DE406, DE430 and DE431 from JPL’s ftp site. The sizes of the binary files are 53.3MB for DE405, 190MB for DE406, 85.5MB for DE430 and 2.6GB for DE431. For each ephemeris I downloaded, I went through the data in the check file and confirmed that Project Pluto’s C code reproduces the data in the file to machine roundoff precision.

4.2 Computing the Geometric Position and Velocity in GCRS

Project Pluto’s software contains a function that can be used to calculate the rectangular coordinates of the positions and velocities of the Sun, Moon, and planets relative to a specified object at any given TDB time using the coefficients of the Chebyshev polynomials from the appropriate time interval. The axes of the rectangular coordinate system are aligned with that of the ICRS. The raw data of the ephemerides are the rectangular coordinates of the positions and velocities of solar system objects in BCRS with axes aligned with those of the ICRS. I denote the rectangular coordinates of the position (in BCRS) of a particular object at a given TDB time t by the vector $\mathbf{x}(t)$, and the velocity by $\mathbf{v}(t)$. The position and velocity of the object relative to a target object are simply computed from $\mathbf{X}(t) = \mathbf{x}(t) - \mathbf{x}_t(t)$ and $\mathbf{V}(t) = \mathbf{v}(t) - \mathbf{v}_t(t)$, where $\mathbf{x}_t(t)$ and $\mathbf{v}_t(t)$ are respectively the position and velocity of the target object in BCRS. For the calendar calculation, the relevant target is the Earth. The resulting vectors \mathbf{X} and \mathbf{V} represent the rectangular coordinates of the geocentric position and velocity of the object. These are called the *geometry* position and velocity. They need to be converted to the *apparent* position and velocity by taking into account the combined effect of light-time correction and aberration of light.

5 Light-Time Correction and Aberration of Light

Since the speed of light is finite, the observed position of an object is the object's *retarded position*, i.e. the position at which light emitted earlier has just reached the observer. This is called the *light-time correction*. Let $\mathbf{x}(t)$ be the BCRS geometric position of the object at time t , and $\mathbf{x}_E(t)$ be Earth's BCRS geometric position at time t . Then the observed GCRS position at time t is $\mathbf{X}_{\text{obs}}(t) = \mathbf{x}(t_r) - \mathbf{x}_E(t)$, where the retarded time t_r satisfies the equation

$$t_r = t - \frac{|\mathbf{x}(t_r) - \mathbf{x}_E(t)|}{c} \quad (32)$$

and $c = 299792.458$ km/s is the speed of light. Note that the retarded time t_r appears on both sides of equation (32). This means that it has to be solved iteratively. However, for the Sun and Moon, it is sufficient to use

$$t_r \approx t - \frac{|\mathbf{x}(t) - \mathbf{x}_E(t)|}{c}. \quad (33)$$

By making this approximation, the error in $\mathbf{X}_{\text{obs}}(t)$ is of order $(v/c)^2$, where $v = |\dot{\mathbf{x}}|$ is the speed of the object in BCRS. Since the Sun contains 99.86% of total mass in the solar system, the motion of the Sun in BCRS is very small. Thus, $\mathbf{x}(t_r) \approx \mathbf{x}(t)$ for the Sun and $(v/c)^2$ is very tiny for the Sun. The Moon's orbital speed around the Earth is about 1 km/s, which is much smaller than the orbital speed of the Earth-Moon barycenter (about 30 km/s). Hence the speed of the Moon in BCRS is about 30 km/s and $(v/c)^2 \sim 10^{-8}$ for the Moon. The error associated with the Moon's position is therefore only 0.002'' and can be safely ignored.

Aberration of light also arises from the finite speed of light. Suppose an observer sees an object in the direction \mathbf{n} , the direction of the object relative to another observer moving with velocity \mathbf{v} will be in a direction \mathbf{n}' . The unit vectors \mathbf{n} and \mathbf{n}' are related by the Lorentz transformation. In our case, $\mathbf{n} = \mathbf{X}_{\text{obs}}/|\mathbf{X}_{\text{obs}}|$ and $\mathbf{v} = \dot{\mathbf{x}}_E$ is Earth's velocity in BCRS. Since Earth's orbital speed in BCRS is about 30 km/s, it is sufficient to calculate the aberration of light to first order in v/c . When expanded to order v/c , the expression in the Lorentz transformation is the same as in Newtonian kinematics:

$$\mathbf{n}' = \frac{\mathbf{n} + \boldsymbol{\beta}}{|\mathbf{n} + \boldsymbol{\beta}|}, \quad (34)$$

where $\boldsymbol{\beta} = \dot{\mathbf{x}}_E/c$. We see that the aberration of light shifts an object's position by an amount of $\sim v_E/c \approx 10^{-4} \approx 20.5''$. Failure to take into account the aberration effect will result in an error in the times of the 24 solar terms by about 8 minutes, which is about the time it takes for light to travel from the Sun to Earth. This is not a coincidence. As explained in Section 7.2.3.5 of [Urban & Seidelmann 2013], there is a simple method to calculate the combined effect of light-time correction and aberration of light, also known as the *planetary aberration*, to first order in v/c . I provide a derivation here. It follows from $\mathbf{n} = \mathbf{X}_{\text{obs}}/|\mathbf{X}_{\text{obs}}|$ and equation (34) that

$$\begin{aligned} \mathbf{n}' &\propto \frac{\mathbf{X}_{\text{obs}}}{|\mathbf{X}_{\text{obs}}|} + \frac{\dot{\mathbf{x}}_E}{c} \\ &= \frac{\mathbf{x}(t_r) - \mathbf{x}_E(t)}{|\mathbf{x}(t_r) - \mathbf{x}_E(t)|} + \frac{\dot{\mathbf{x}}_E}{c} \\ &\propto \mathbf{x}(t_r) - \left[\mathbf{x}_E(t) - \frac{|\mathbf{x}(t_r) - \mathbf{x}_E(t)|}{c} \dot{\mathbf{x}}_E \right] \end{aligned}$$

$$\begin{aligned}
&\approx \mathbf{x}(t_r) - \mathbf{x}_E \left(t - \frac{|\mathbf{x}(t_r) - \mathbf{x}_E(t)|}{c} \right) \\
&= \mathbf{x}(t_r) - \mathbf{x}_E(t_r).
\end{aligned} \tag{35}$$

Hence to first order in v/c , the combined effect of light-time correction and aberration of light results in the *apparent GCRS position* given by

$$\mathbf{X}_{\text{apparent}}(t) = \mathbf{x}(t_r) - \mathbf{x}_E(t_r). \tag{36}$$

For the Sun, t_r is the time it takes for light to travel from the Sun to Earth. This is the reason that the aberration effect causes a correction of about 8 minutes. For the Moon, $t_r = 1.3$ seconds, which is also substantial. Equation (36) is very convenient in calculating the combined effect of light-time correction and aberration of light since there is a function in Project Pluto's software that can be used to compute $\mathbf{x} - \mathbf{x}_E$ for any solar system object at any given TDB time.

In addition to the apparent GCRS position, it is also useful to compute its time derivative. Taking the time derivative of equation (36) yields

$$\dot{\mathbf{X}}_{\text{apparent}}(t) = [\dot{\mathbf{x}}(t_r) - \dot{\mathbf{x}}_E(t_r)] \frac{dt_r}{dt}. \tag{37}$$

Denote $D(t_r) = |\mathbf{x}(t_r) - \mathbf{x}|$ as the retarded distance and using (32), I obtain

$$\frac{dt_r}{dt} = 1 - \frac{\dot{D}(t_r)}{c} \frac{dt_r}{dt} \Rightarrow \frac{dt_r}{dt} = \frac{1}{1 + v_r(t_r)/c}, \tag{38}$$

where $v_r = dD/dt$ is the radial velocity of the object relative to Earth. Equation (37) can be written as

$$\dot{\mathbf{X}}_{\text{apparent}}(t) = \frac{\dot{\mathbf{x}}(t_r) - \dot{\mathbf{x}}_E(t_r)}{1 + v_r(t_r)/c} \tag{39}$$

The denominator in the above expression is the cause of the apparent superluminal motion of jets in some active galaxies: if a jet moving close to the speed of light is moving at a very small angle towards the observer, $1 + v_r/c \ll 1$ and it is possible to have the apparent speed greater than the speed of light. However, v_r/c is negligible for the Sun and Moon because both Earth's orbit around the Sun and Moon's orbit around Earth are nearly circular. The radial speed $|v_r| \sim 0.5$ km/s for the Sun and $|v_r| \sim 0.05$ km/s for the Moon. Hence the fractional error in ignoring the v_r/c term is of order 10^{-6} for the Sun and 10^{-7} for the Moon. As will be explained in Section 7, this error has almost no effect on the accuracy of the times of moon phases and 24 solar term. I therefore ignore it and set

$$\dot{\mathbf{X}}_{\text{apparent}}(t) = \dot{\mathbf{x}}(t_r) - \dot{\mathbf{x}}_E(t_r) = \mathbf{v}(t_r) - \mathbf{v}_E(t_r). \tag{40}$$

This is also a very convenient equation to compute $\dot{\mathbf{X}}_{\text{apparent}}$ since there is a function in Project Pluto's software that can be used to compute $\mathbf{v} - \mathbf{v}_E$ for any solar system object at any given TDB time.

6 Apparent Geocentric Longitude

The apparent geocentric longitude of the Sun and Moon are crucial in the computation of the times of moon phases and 24 solar terms. The longitude can be computed by combining the equations in Sections 2, 3 and 5. Here are the steps of the calculation for a given TDB time t :

1. Compute the geometric position of the object (Sun or Moon) in GCRS by $\mathbf{X}_{\text{geometric}}(t) = \mathbf{x}(t) - \mathbf{x}_E(t)$, where \mathbf{x} is the BCRS position of the object and \mathbf{x}_E is the BCRS position of the Earth.
2. Compute the retarded time approximately using $t_r \approx t - |\mathbf{X}_{\text{geometric}}(t)|/c$.
3. Compute the apparent geocentric position using the equation $\mathbf{X}_{\text{apparent}}(t) \approx \mathbf{x}(t_r) - \mathbf{x}_E(t_r)$. The resulting vector gives the components in the GCRS coordinate system. Recall that the origin of the GCRS is at the geocenter and its axes are aligned with those of the ICRS.
4. Calculate the precession matrix $\mathbf{P}(t)$ using equations (18), (19) and (20). Calculate the matrix $\mathbf{R}_1(\epsilon(t))\mathbf{N}(t)$ using equations (21)–(23), (25), (29) and (31).
5. Convert the GCRS apparent geocentric position $\mathbf{X}_{\text{apparent}}(t)$ to the apparent geocentric position in coordinates associated with the true ecliptic and equinox of date using the equation

$$\mathbf{X}_{\text{ec}}(t) = \mathbf{R}_1(\epsilon(t))\mathbf{N}(t)\mathbf{P}(t)\mathbf{B}\mathbf{X}_{\text{apparent}}(t), \quad (41)$$

where the frame bias matrix \mathbf{B} is given by equation (11).

6. Calculate the apparent geocentric longitude using $\lambda = \arg(X_{\text{ec}} + iY_{\text{ec}})$. That is, $\lambda = \tan^{-1}(Y_{\text{ec}}/X_{\text{ec}})$ in the appropriate quadrant.

It is also useful to calculate the time derivative of λ . Differentiating $\lambda(t) = \tan^{-1}[Y_{\text{ec}}(t)/X_{\text{ec}}(t)]$ yields

$$\dot{\lambda}(t) = \frac{X_{\text{ec}}\dot{Y}_{\text{ec}} - Y_{\text{ec}}\dot{X}_{\text{ec}}}{X_{\text{ec}}^2 + Y_{\text{ec}}^2}. \quad (42)$$

The time derivatives \dot{X}_{ec} and \dot{Y}_{ec} are obtained by differentiating equation (41) with respect to t :

$$\dot{\mathbf{X}}_{\text{ec}}(t) = \mathbf{R}_1(\epsilon(t))\mathbf{N}(t)\mathbf{P}(t)\mathbf{B}\dot{\mathbf{X}}_{\text{apparent}}(t) + \frac{d}{dt}[\mathbf{R}_1(\epsilon(t))\mathbf{N}(t)\mathbf{P}(t)]\mathbf{B}\mathbf{X}_{\text{apparent}}(t). \quad (43)$$

The first term arises from the fact that the object is moving relative to Earth; the second term arises from the fact that the coordinate axes associated with the true ecliptic and equinox of date are moving because of precession and nutation. Obviously the first term is the dominant term and the second term can be ignored. Hence,

$$\dot{\mathbf{X}}_{\text{ec}}(t) \approx \mathbf{R}_1(\epsilon(t))\mathbf{N}(t)\mathbf{P}(t)\mathbf{B}\dot{\mathbf{X}}_{\text{apparent}}(t) = \mathbf{R}_1(\epsilon(t))\mathbf{N}(t)\mathbf{P}(t)\mathbf{B}[\mathbf{v}(t_r) - \mathbf{v}_E(t_r)]. \quad (44)$$

Let's perform an order of magnitude estimate of the error in dropping the second term. The sidereal period of the Earth around the Sun is 365.2564 days. This means that the Sun completes a full circle as seen on Earth in the GCRS coordinates in 365.2564 days. So in one day, the Sun moves, on average, about $360^\circ/365.2564 \approx 1^\circ$. Hence $|\dot{\mathbf{X}}_{\text{apparent}}|/|\mathbf{X}_{\text{apparent}}| \approx 1^\circ/\text{day}$ for the Sun. The Moon's sidereal period around Earth is 27.3217 days. Similar calculation shows $|\dot{\mathbf{X}}_{\text{apparent}}|/|\mathbf{X}_{\text{apparent}}| \approx 13^\circ/\text{day}$ for the Moon. The major component of precession is the westward motion of the vernal equinox at a rate of $50.3''/\text{year}$, which is about $0.14''/\text{day}$. Hence, $|\dot{\mathbf{P}}| \sim 0.14''/\text{day}$. Here $|\dot{\mathbf{P}}|$ means the (maximum) magnitude of the components in the $\dot{\mathbf{P}}$ matrix. The major component of nutation is the 18.6 year cycle caused by Moon's orbital plane precessing around the ecliptic. This is represented by the first term in equation (26). So $|d(\mathbf{R}_1(\epsilon)\mathbf{N})/dt| \sim 17.2''|\dot{\Omega}|$ and from equation (21) we have $|\dot{\Omega}| \approx 7000000''/\text{century} \approx 10^{-3} \text{ rad/day}$. Hence

$|d(\mathbf{R}_1(\boldsymbol{\epsilon})\mathbf{N})/dt| \sim 0.02''/\text{day}$. Thus, not surprisingly precession is the dominant component in the second term, but $|\dot{\mathbf{P}}|/|\mathbf{X}_{\text{apparent}}| \sim 4 \times 10^{-5}$ for the Sun and $\sim 3 \times 10^{-6}$ for the Moon. Therefore, dropping the second term introduces a fractional error of order 4×10^{-5} for the Sun and 3×10^{-6} for the Moon. As will be explained in the next section, this error has almost no effect on the accuracy in the times of the moon phases and 24 solar terms.

7 Computation of the TDB Times of Moon Phases and 24 Solar Terms

The 24 solar terms are defined when the apparent geocentric longitude of the Sun reaches integer multiples of 15° or $\pi/12$ radians. Hereafter, all angles are assumed to be in radians. New moon (lunar conjunction) is defined when the apparent geocentric longitude of the Moon λ_M is equal to the apparent geocentric longitude of the Sun λ_S . First quarter is defined when $P(\lambda_M - \lambda_S) = \pi/2$; full moon is defined when $P(\lambda_M - \lambda_S) = -\pi$; third quarter is defined when $P(\lambda_M - \lambda_S) = -\pi/2$. Here the function P is defined as

$$P(x) \equiv x - 2\pi \left[\frac{x + \pi}{2\pi} \right] \quad (45)$$

and $[x]$ means the largest integer less than or equal to x . Thus, the operator $[]$ is the same as the `floor` function in C and python. The function P simply adds an integer multiples of 2π to bring the argument into the interval $[-\pi, \pi)$. Computation of the times of moon phases and 24 solar terms therefore boils down to finding the roots of $f(t) = 0$. The function f is listed in the tables below for computing the moon phases and 24 solar terms. Note that both λ_S and λ_M are functions of t .

Moon Phase	Function f
new moon	$P(\lambda_M - \lambda_S)$
first quarter	$P(\lambda_M - \lambda_S - \pi/2)$
full moon	$P(\lambda_M - \lambda_S - \pi)$
third quarter	$P(\lambda_M - \lambda_S + \pi/2)$

Solar Term	Function f	Solar Term	Function f
J1 (立春)	$P(\lambda_S + \pi/4)$	Z1 (雨水)	$P(\lambda_S + \pi/6)$
J2 (驚蟄)	$P(\lambda_S + \pi/12)$	Z2 (春分)	$P(\lambda_S)$
J3 (清明)	$P(\lambda_S - \pi/12)$	Z3 (穀雨)	$P(\lambda_S - \pi/6)$
J4 (立夏)	$P(\lambda_S - \pi/4)$	Z4 (小滿)	$P(\lambda_S - \pi/3)$
J5 (芒種)	$P(\lambda_S - 5\pi/12)$	Z5 (夏至)	$P(\lambda_S - \pi/2)$
J6 (小暑)	$P(\lambda_S - 7\pi/12)$	Z6 (大暑)	$P(\lambda_S - 2\pi/3)$
J7 (立秋)	$P(\lambda_S - 3\pi/4)$	Z7 (處暑)	$P(\lambda_S - 5\pi/6)$
J8 (白露)	$P(\lambda_S - 11\pi/12)$	Z8 (秋分)	$P(\lambda_S - \pi)$
J9 (寒露)	$P(\lambda_S + 11\pi/12)$	Z9 (霜降)	$P(\lambda_S + 5\pi/6)$
J10 (立冬)	$P(\lambda_S + 3\pi/4)$	Z10 (小雪)	$P(\lambda_S + 2\pi/3)$
J11 (大雪)	$P(\lambda_S + 7\pi/12)$	Z11 (冬至)	$P(\lambda_S + \pi/2)$
J12 (小寒)	$P(\lambda_S + 5\pi/12)$	Z12 (大寒)	$P(\lambda_S + \pi/3)$

7.1 Newton-Raphson Method

In our case, the best method for finding the roots of $f(t) = 0$ is the Newton-Raphson method. This is an iterative scheme. Suppose t_n is the approximate value of a root in the n th iteration. The improved approximation in the $(n + 1)$ th iteration is given by the equation

$$t_{n+1} = t_n - \frac{f(t_n)}{\dot{f}(t_n)} \quad (46)$$

This scheme requires the time derivative of f , which is $\dot{\lambda}_S$ for 24 solar terms and $\dot{\lambda}_M - \dot{\lambda}_S$ for moon phases. They can be computed using the method described in the previous section. An initial guess is required to start the iteration. This turns out to be very easy. We know that the motion of the Sun and Moon are fairly uniform because of the low eccentricity of the Earth's orbit around the Sun and the low eccentricity of the Moon's orbit around the Earth. As a result, both $\dot{\lambda}_S$ and $\dot{\lambda}_M$ do not change significantly and can be used to calculate the approximate time when λ_S or $\lambda_M - \lambda_S$ reaches a certain value. Suppose we want to find a root of $f(t) = 0$ close to a given TDB time t_0 . We can simply set $t_1 = t_0 - f(t_0)/\dot{f}(t_0)$. However, it is an approximation to a root since we know $f(t) = 0$ have multiple roots in our cases. Sometimes we may want to find a specific root, e.g. the first new moon before a particular winter solstice or the first full moon after a particular new moon. In cases like these, we want to find the first root before or after a given TDB time t_0 . We can modify the initial guess t_1 according to the following equation:

$$t_1 = t_0 - \frac{f(t_0)}{\dot{f}(t_0)} + \frac{2k\pi}{\dot{f}(t_0)}, \quad (47)$$

where k is an integer that takes the value of 0, 1 or -1 . As mentioned above, \dot{f} is close to being a constant. In the case of 24 solar terms, $2\pi/\dot{f}$ is close to a tropical year (365.2422 days). In the case of moon phases, $2\pi/\dot{f}$ is close to a synodic month (29.5306 days). Hence, the addition of $2k\pi/\dot{f}(t_0)$ is just to add/subtract a period to the original guess $t_0 - f(t_0)/\dot{f}(t_0)$. It works as follows.

Case 1: Want to find the first root *before* t_0 . In this case, we want $t_1 < t_0$. Hence if $f(t_0) > 0$ set $k = 0$; if $f(t_0) < 0$ set $k = -1$. (Note that $f(t_0) \in [-\pi, \pi)$ and $\dot{f}(t_0) > 0$ for both moon phases and 24 solar terms.)

Case 2: Want to find the first root *after* t_0 . In this case, we want $t_1 > t_0$. Hence if $f(t_0) > 0$ set $k = 1$; if $f(t_0) < 0$ set $k = 0$.

Once t_1 is determined, the iteration procedure can be started. The sequence $\{t_n\}$ converges very rapidly. I stop the process at iteration m when $|t_m - t_{m-1}| < \varepsilon$ and I set $\varepsilon = 10^{-8}$ days = 0.864 milliseconds. This convergence criterion is equivalent to setting $|f(t_{m-1})/\dot{f}(t_{m-1})| < \varepsilon$.

Newton-Raphson method turns out to be very effective in the computation of the times of moon phases and 24 solar terms. After setting t_1 , the method converges to the required accuracy within 3 or 4 iterations. As a result, it only takes about 50 seconds to compute all the times of the 4 moon phases and all 24 solar terms from years 1600 to 3500 using my somewhat old computer at home. In the early development of my code, I did not simplify the calculation of the matrix $\mathbf{R}_1(\epsilon)\mathbf{N}$ as described in Section 3 and computed the nutation in obliquity $\Delta\epsilon$ as well. It took about 90 seconds to do the same calculation. The series involving $\Delta\epsilon$ contains 1056 terms involving trigonometric functions. The series involving $\Delta\psi$ contains 1358 terms of trigonometric functions.

When I only kept a few terms in $\Delta\psi$ and $\Delta\epsilon$, the computation finished in a few seconds. From this information I conclude that most of the computation time is spent in computing $\Delta\psi$. This is probably not true if I had used a semi-analytic ephemeris such as VSOP87 or ELP/MPP02, which involves Poisson series containing thousands of terms. This is the advantage of using the JPL ephemerides since the positions and velocities are simply computed by interpolation of the tabulated values using the Chebyshev polynomials. The computation is much faster and the result is more accurate than other semi-analytic theories. One disadvantage of the JPL ephemerides is that they require large data files in order for the interpolation to produce highly accurate data. However, even the largest data file (for DE431) is about 2.6GB, which is not much in today's computation world. Another disadvantage is that it is not possible to extrapolate to times beyond the time span covered by the ephemerides. For DE431, the time span is from years -13,200 to 17,191, which should be adequate for many applications.

As mentioned above, my criterion for convergence is equivalent to setting $|f(t)/\dot{f}(t)| < \varepsilon$. Hence the accuracy is mainly determined by how close $f(t)$ is towards 0. The role of \dot{f} is to provide a conversion factor between the deviation of $|f(t)|$ from 0 to the error in the computed time. I mention in the previous sections that there is some error in the computation of \dot{f} because I drop some terms. In particular, dropping the factor dt_r/dt in equation (37) introduces a fractional error of $\sim 10^{-6}$ for the Sun and $\sim 10^{-7}$ for the Moon. Dropping the second term in equation (43) introduces a fractional error of $\sim 4 \times 10^{-5}$ for the Sun and $\sim 3 \times 10^{-6}$ for the Moon. Since the accuracy is set by $|f(t)/\dot{f}(t)| < \varepsilon$, error in \dot{f} is equivalent to changing ε . Thus, a fractional error of 10^{-4} is the same as changing ε to 1.0001ε in the worse cases. For $\varepsilon = 10^{-8}$ days = 0.864 ms, that is the difference between 0.864 ms and 0.8640864 ms. Therefore, the accuracy of the computed times is almost independent of the small error in \dot{f} . Even a 10% error in \dot{f} is harmless as long as ε is set to a much smaller value than the target accuracy. Since I know \dot{f} is nearly constant, I tried setting it to its average value just out of curiosity. So I tried $\dot{f} = 2\pi/29.5306$ days for the moon phase calculation and $\dot{f} = 2\pi/365.2422$ days for the solar term calculation. When compared to the calculation using accurate \dot{f} , I find a maximum deviation of 0.16 ms for the 24 solar terms and 0.24 ms for moon phases from years 1600 to 3500. These are all smaller than the prescribed ε . However, this does not mean that accurate values of \dot{f} are useless because even though the accuracy does not suffer, the number of iterations required to reach convergence increase. Using these inaccurate values of \dot{f} takes up to 7 iterations for the solar term calculation to converge to the required accuracy, and up to 12 iterations for the moon phase calculation to converge, causing a significant slowdown in the calculation. It is well-known that Newton-Raphson method is a second-order scheme, which converges much faster than first-order root-searching schemes. However, it is only second order if the derivative is accurate. Computing \dot{f} is very straightforward since velocities are output quantities in the JPL ephemerides.

7.2 Comparison with Different Ephemerides and with a Different Precession Model

I have computed the TDB times of the moon phases and 24 solar terms using JPL's DE431, DE430 and DE406. As mentioned above, the only difference between DE430 and DE431 is that DE430 includes a damping term between the Moon's liquid core and solid mantle that gives the best fit to observation data but not suitable for backward integration of more than a few centuries. JPL recommends using DE430 for times within few hundred years from J2000.0 and DE431 for earlier and later periods. Comparing the computed times of moon phases and 24 solar terms between

DE430 and DE431, I find a maximum difference of 0.000419 s ($< \varepsilon$) in 24 solar terms and 0.85 s in moon phases between years 1600 and 2500. The maximum differences between years 1900 and 2200 are 0.00015 s ($< \varepsilon$) in 24 solar terms and 0.2 s in moon phases. Note that the maximum deviations in the times of 24 solar terms are less than the prescribed convergence criterion $\varepsilon = 0.000864$ s and so there is essentially no difference in the times of 24 solar terms.

DE405 was released in 1998 and was the basis for the Astronomical Almanac from 2003–2014. DE406 was released at the same time as DE405. It was a condensed version of DE405. The integration method used in DE406 was the same as that of DE405, but the accuracy of the interpolating polynomials has been lessened to reduce file size for the longer time span covered by the file. Comparing the computed times of moon phases and 24 solar terms between DE406 and DE431, I find maximum deviations of 0.14 s in 24 solar terms and 0.18 s in moon phases for years between 1600 and 2500.

As pointed out in [Vondrák et al 2011], the IAU 2006 precession model is only accurate within about 1000 years from J2000.0. In the paper the authors developed a new precession model that can be used within 200 millennia from J2000.0. This precession model is adopted by the popular open-source planetarium software Stellarium. I also use it for my star chart website. Since I compute times up to the year 3500, I want to investigate the difference between the new precession model and IAU 2006. I compute the times of the moon phases and 24 solar terms using DE431 and the new precession model for years from 1600 to 3500, and then compare them to the times calculated using DE431 and IAU 2006 precession model. I find maximum deviations of 0.19 s in 24 solar terms and 0.00038 s ($< \varepsilon$) in moon phases for years between 1600 and 2500, but the maximum deviations go up to 3 s in 24 solar terms and 0.0035 s in moon phases when comparing the times for years from 1600 to 3500.

Based on these comparisons, I estimate that the TDB times of the computed moon phases and 24 solar terms using DE431 and IAU 2006 precession model are probably accurate to better than 0.2 s for years between 1901 and 2200, which is the time span covered in my Chinese calendar website.

7.3 Text File for the TDB Times

Even though the computation of the TDB times of moon phases and 24 solar terms is complicated, they only need to be done once for a given ephemeris and precession-nutation model. I created a text file in ASCII format storing these times computed from DE431 with IAU 2006/2000A precession-nutation model for years between 1600 and 3500. The file is `TDBtimes.txt`, available on my GitHub repository. It is designed for use in the calculation of Chinese calendar.

The file contains 1901 rows and 87 columns. The first column is the Gregorian year. The second column is the Julian date number on January -1 of the year at 16:00 TDB, or January 0 of the year at 0:00 (TDB+8). For example, the `jd0` in the year 2000 is 2451543.166666667. This was the Julian date number on January -1, 2000 at 16:00 TDB (December 30, 1999 at 16:00 TDB). The value of `jd0` is always an integer plus 1/6 because the TDB time is always 16:00. `jd0` is the time origin from which the times in the rest of the columns are measured. That is, the Julian date numbers of the times in the rest of the columns are `jd0 + values listed in those columns`.

The third column, labelled `Z11a`, is the time of the first winter solstice *before* `jd0`, i.e. the time of the December solstice in the previous year. For example, the third column in year 2000 is `-8.343841734507215`. This means that the December solstice in 1999 occurred at the TDB Julian day `2451543.166666667 - 8.343841734507215`, which was December 22, 1999 at 7:44:52 TDB. Columns 4–27 are the 24 solar terms in the year, starting from J12 (小寒) and ending with

Z11 (labelled Z11b), the December solstice in the year. Therefore, the time listed in column Z11b is the same as the time listed in Z11a in the following year. Their values are different simply because they are measured from different time origins.

Column 28, labelled Q0_01, is the time of the first new moon occurring *before* the winter solstice listed in the Z11a column. The remaining columns (Columns 29–87) are the times of the first quarters (labelled Q1_xx), full moons (labelled Q2_xx), third quarters (labelled Q3_xx) and other new moons (labelled Q0_xx) in chronological order. Here xx ranges from 01 to 15 representing the lunation number counting from the new moon in column 28. Each year lists the four moon phases covering 15 lunations.

The data structure of the file is designed for convenience in the computation of the Chinese calendar in the *suì* from the winter solstice of the previous Gregorian year to the winter solstice of the current Gregorian year. The new moon listed in the Q0_01 column is usually but not always associated with month 11 in the Chinese calendar. Even though by definition the new moon in the Q0_02 column must occur after the winter solstice, it could happen that it is on the same day as the winter solstice if it occurs only a few hours later than the winter solstice. If that is the case, it is the new moon in the Q0_02 column that is associated with month 11. To determine if this is the case requires the knowledge of the beginning and ending of a day. The Chinese calendar uses UTC+8 to define a day, so we will need to know the conversion between TDB and UTC to determine accurately if that is the case. A *suì* can contain up to 13 Chinese months, which requires 14 new moons to determine the number of days in each of the 13 months. If month 11 starts on the date of the second new moon, it must be on the same day as the winter solstice. It is easy to show that the *suì* can only have 12 months in that case. So in all cases 14 new moons are sufficient to determine all months in a *suì*. Listing moon phases covering 15 lunations is more than enough for the calendar calculation, and there are some overlaps between the moon phases listed in a year and those listed in the previous and the following year.

8 Conversion Between TDB and UTC

The choice of separating the computation of times in TDB and conversion between TDB and UTC is deliberate. The computation of the TDB times is complicated but only need to be done once for any given ephemeris and precession-nutation model. The conversion between TDB and UTC is constantly modified as new data become available. By separating the two procedures, each time the conversion between TDB and UTC is modified, we just need to change the conversion and do not need to recompute the times of moon phases and 24 solar terms. If the UTC+8 times were given in the text file `TDBtimes.txt`, it would become outdated when a modified conversion emerges.

UTC was invented in 1960. There were several changes in UTC until it was finalized in 1972. To prevent confusion, I use UT1+8 for times before 1972. Conversion between TDB and TT is ignored because the deviations between the two times are less than 0.002 s over several millennia. For years before 1972, values of $\Delta T = TT - UT1$ are calculated using the analytic fitting formula by [Espanak and Meeus]. So UT1 is obtained by simply subtracting ΔT from TDB. For years from 1972 to the present day, I use a table of published leap seconds (e.g. https://en.wikipedia.org/wiki/Leap_second) to calculate $TT - UTC$. Specifically,

$$TT - UTC = 42.184 \text{ s} + \text{total number of leap seconds added to UTC since 1972.} \quad (48)$$

For the future years, I use the extrapolation formulas for ΔT by [Espanak and Meeus] to extrap-

olate the value of $TT - UTC$. By construction, $|UTC - UT1| < 0.9$ s, so the error in setting $TT - UTC$ to ΔT is probably less than the error in the extrapolation formulas. In the following I give two examples to demonstrate how the conversion works.

Example 1: From the text file `TDBtimes.txt`, the first new moon in the year 2018 is the new moon in the `Q0_02` column (column 32). The TDB time of the new moon is 17.42943724648089 days from January 0, 2018 at 00:00 (TDB+8), which is January 17, 2018 at 10:18:23.378 (TDB+8). According to the leap second table, the total number of leap seconds added to UTC in 2018 is 27 s. Hence $TT - UTC = 69.184$ s. Subtracting 69.184 s from the TDB+8 time above, I find the new moon occurred on January 17, 2018 at 10:17:14 (UTC+8). This is consistent with the time listed in *Astronomical Phenomena for the year 2018* (January 17, 2018 at 02:17 UTC) published jointly by The Nautical Almanac Office and Her Majesty’s Nautical Almanac Office. This is not surprising since the astronomical data of the book are based on DE430 and the time differences of the moon phases calculated by DE431 and DE430 are less than 0.2 s in the period 1901–2200 as mentioned in the previous section.

Example 2: From the text file `TDBtimes.txt`, the new moon listed in the column `Q0_11` is 272.0013125274384 days from January 0, 2057 at 00:00 (TDB+8), which is September 29, 2057 at 00:01:53.402 (TDB+8). Using the extrapolation formula by [Espanak and Meeus] I find $\Delta T = 108.875$ s in September 2057. Setting $TT - UTC \approx 108.875$ s and subtracting it from the TDB+8 time above gives the new moon occurring on September 29, 2057 at 00:00:4.5 (UTC+8). This is just 4.5 s past midnight! If the extrapolation is off by more than 5 s, this new moon may actually occur on September 28, 2057. This new moon is associated with month 9 in the Chinese calendar. Therefore, it is currently not possible to determine the exact start day of month 9 in 2057. We just have to wait for the time to come to get a better handle on the situation.

As we see, the conversion between TDB and UTC depends on the number of leap seconds added to UTC. It is very difficult, if not impossible, to predict accurately how many leap seconds will be added to UTC decades from now because of the irregularity of Earth’s rotation. The situation may also be complicated by the uncertainty of the leap-second policy in the future: there have been discussions on the abolishment of leap seconds since 2005.

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