## A Sneak Peak

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#### Abstract

This set of notes is made with reference to the Astrophysics Coursebook by The University of Edinburgh (Introductory Astrophysics, Course PHYS08050) . I have gained permission from Professor Catherine Heymans to put my modified notes in my website. It is also available in my personal website.

This is a special version of the notes (A Sneak Peak!) If you want the original copy, please email henry36c@gmail.com or ask Henry in person (Well he's quite a nice guy)

### 1 Some Black Holes

• If a small mass  $\Delta m$  starts from a distance and falls onto a large mass M, it gains kinetic energy:

$$E = \frac{GM\Delta m}{R} \tag{1}$$

- For gases however, friction and collisions randomise the energy and turn it into heat
- For a black hole:

$$R_{\text{Event Horizon}} = \frac{2GM}{c^2}$$
 (2)

• By substitution,

$$E = \frac{GM\Delta m}{\frac{2GM}{c^2}} \tag{3}$$

$$=\frac{1}{2}\Delta mc^2\tag{4}$$

- The efficiency,  $\mu$ , is 0.5, which is much higher than  $\mu_{\rm nuc} = 0.007$
- This is obviously, a fallacy

#### • Incorrect Assumptions

- Radial free fall is unlikely
- At larger distances, the accreting material will always have some angular momentum, and end up forming a rotating disc around the black hole
- Friction between neighbouring radial annuli then allows the material to slowly spiral inwards, forming a gradually heated accretion disc
- Efficiency  $\mu$  is reduced by the following 2 effects:
  - \* Effective radius is **not** the Event Horizon. For non-rotating black-holes, the ISCO is located at:

$$r_{\rm ms} = 6 \frac{GM}{c^2} \tag{5}$$

$$=3R_{\rm EH}\tag{6}$$

\* I.e. There is no stable orbit  $< 3R_{\rm EH}$ 

- \* And therefore the thermal energy gained by the gas decreases
- \* Some energy (half) is converted to rotational energy, as:

$$\Delta K = \frac{dK}{dR} \Delta R \tag{7}$$

$$= -\frac{GM\Delta m}{2R^2}\Delta R \tag{8}$$

$$\Delta U = \frac{dU}{dR} \Delta R \tag{9}$$

$$= -\frac{GM\Delta m}{R^2}\Delta R \tag{10}$$

$$\frac{\Delta K}{\Delta U} = \frac{-\frac{GM\Delta m}{2R^2} \Delta R}{-\frac{GM\Delta m}{R^2} \Delta R} \tag{11}$$

$$=\frac{1}{2}\tag{12}$$

Combining with other effects, we have  $\mu \approx 0.1$ 

## 1.1 Eddington Luminosity

- A photon has energy  $E = \frac{hc}{\lambda}$  and momentum  $p = \frac{h}{\lambda} = \frac{E}{c}$ . So you can visualize momentum flux being  $\frac{S}{c}$ , where S is the radiation flux
- Most of that flux may pass straight through, but some of it will scatter on the electrons inside atoms. That scattering produces a force on the electrons, which drag the atoms with them. The scattering process has a cross-section  $\sigma_e = 6.6510^{-29} m^2$
- The rate <sup>1</sup> of momentum transfer (or in other words, radiation force on each atom) is therefore:

$$\frac{dp}{dt} = \frac{S\sigma_e}{c} \tag{13}$$

• Combining with  $L = \frac{S}{4\pi R^2}$  we obtain:

$$F_{\text{radiation}} = \frac{L\sigma_e}{4\pi R^2 c} \tag{14}$$

• Assuming all the gas are hydrogen and each hydrogen atom composes of 1 proton, which has mass of  $m_p$ . As  $m_e \ll m_p$ , we have:

$$F_{\text{gravitation}} = \frac{GMm_p}{R^2} \tag{15}$$

• At the limiting luminosity, the outward radiation force is equal to the inwards gravitational force, *i.e.*  $F_{\text{radiation}} = F_{\text{gravitation}}$ . Combining everything gives us:

$$L_{\text{Max}} = \frac{4\pi G m_p c}{\sigma_e} \times M \tag{16}$$

## 2 Distance-redshift relation

• This section resulted in a Nobel Prize in 2011. They used distant supernova as standard candles to probe the Universe

<sup>&</sup>lt;sup>1</sup>This implies differentiation

• We have the formula  $D = \frac{cz}{H_0}$ . There is an extension that works for more distant galaxies:

$$D_L \approx \frac{c(1+z)(z - \frac{1+q}{2}z^2)}{H_0} \tag{17}$$

• Where  $q=-\left(1+rac{\dot{H}}{H^2}
ight)$  and  $\dot{H}$  is the rate of change of the Hubble parameter

# 3 Concordant cosmology

- Baryonic content  $\Omega_b \approx 0.05$
- The total dark and baryonic matter content  $\Omega_m\approx 0.3$
- Dark energy content  $\Omega_{\Lambda}\approx 0.7$
- The university is flat and has critical density:  $\Omega_b + \Omega_m + \Omega_\Lambda = 1$