

Wind accretion in Supergiant X-ray binaries

II. Disc formation

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ABSTRACT

The turbulent twilight of massive stellar binaries stars determines their final fate. In Supergiant X-ray binaries, mass transfer proceeds through the accretion of a fraction of the dense and fast wind from an evolved donor star by an orbiting compact object, generally a neutron star. The accretor both perturbs the stellar wind and provides a moving X-ray source to probe its internal structure : the characteristics of the flow constrain the X-ray emission and absorption along the line-of-sight, while the X-ray ionizing feedback alters the wind acceleration. In order to consistently monitor the flow from the stellar surface down to the accretor, we designed a multi-scale model which includes recent insights on the properties of massive star winds. We performed 3D numerical simulations to evaluate the mass and angular momentum accretion rates onto the compact object, depending on the wind to orbital speed ratio and to the cooling efficiency within the shocked region. We identified conditions favorable to the formation of a disc-like structure beyond the neutron star magnetosphere, in spite of the low angular momentum carried by the wind, and analyzed its properties. These conditions, compatible with the currently known parameters of the Supergiant X-ray binary Vela X-1 and to a certain extent with Cygnus X-1, indicate the possible presence of a limited disc-structure in the former case and account for an extended one in the latter case.

Key words: accretion, accretion discs – X-rays: binaries – stars: neutron, supergiants, winds, outflows – methods: numerical

1 INTRODUCTION

Our current understanding of single-star evolution has proved consistent with the observational surveys carried on by contemporary missions. However, few stars are deprived of a gravitationally bound stellar companion, with more high mass stars showing a higher multiplicity frequency (?). Some of them undergo a phase of interaction with their companion tied enough to significantly alter their subsequent evolution. One of the stages important to provide a more complete evaluation of the impact of binarity on stellar evolution are high mass X-ray binaries (HMXB). They represent the turbulent twilight of the entangled evolution of two massive stars, one having already collapsed into a compact object. They are believed to be progenitors of compact object binaries whose final coalescence has been observed by the LIGO/VIRGO collaboration (?).

The main characteristic manifestation of binarity is through mass transfer between the two components. It can lead to chemical contamination (e.g. in the case of Barium and Carbon-enhanced metal-poor stars ??) or to the stripping of the outer envelope of an evolved star, leaving a naked Helium-rich core (e.g. for some hot subdwarf stars ??). The transfer of mass goes hand in hand with a transfer of angular momentum which can produce fast rotators like millisecond pulsars (?) or Be-stars (?).

Supergiant X-ray Binaries (sgXB) are systems where a compact object, generally a neutron star, orbits an O/B Supergiant. Mass transfer proceeds through the capture of a fraction of the dense and fast line-driven stellar wind by the accreting body. As the flow accelerates to the compact accretor, it heats up and emits copious amounts of X-rays. This highly non-conservative mechanism, also called wind accretion, has been put in opposition to the better understood Roche lobe overflow mechanism, although elements of both can co-exist in hybrid models (??). The central dis-

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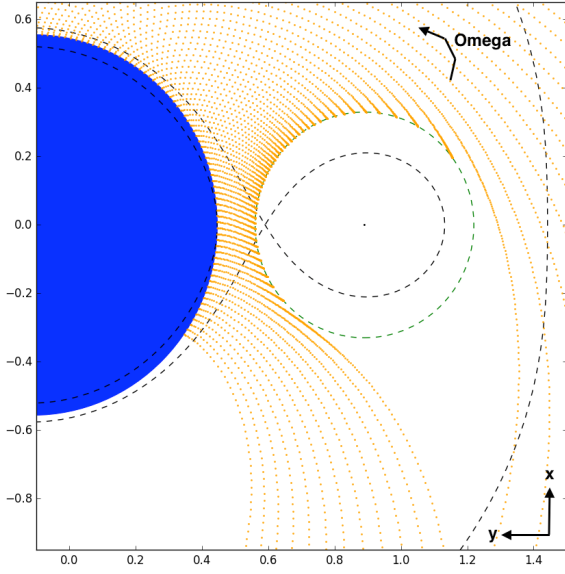


Figure 1. Computed streamlines (orange dots) from the star (dark blue) to the Roche lobe of the accretor on the right (green dashed circle), in the orbital plane. The black dashed lines represent the critical Roche surface passing by the first Lagrangian point.

tion lies in the presence of a large and permanent disc in the latter case, while the conditions for such a structure in the former case are still unclear. Our current knowledge on angular momentum wind accretion builds up on models of asymmetric Bondi-Hoyle-Lyttleton accretion (?) but we still miss a fully consistent frame to follow the flow from the stellar surface down to the X-ray emitting region, in the immediate vicinity of the accretor. Until now, the six orders-of-magnitude or so separating the two scales has precluded any bold numerical attempt to follow the flow all along its journey.

Wind accretion turns out to be extremely sensitive to the properties of the stellar wind. O/B supergiants display dense and fast outflows, with mass loss rates up to several solar masses per million year. The underlying mechanism, unveiled by ? and ?, is the resonant line absorption of UV photons by partly ionized metal ions in the outer layers of the star. As the flow accelerates, it keeps tapping previously untouched Doppler-shifted photons .

key-ingredient

2 MODEL AND NUMERICAL METHOD

2.1 At the orbital scale

2.1.1 General principle

2.1.2 Wind acceleration

XXX ANDREAS : would you be able to summarize the principle of your calculation and the way we account for it here?XXX

2.1.3 The accretion radius

2.2 Within the Roche lobe of the accretor

2.2.1 Equations

1 refers to donor star 2 refers to accretor

normalization units

angular momentum preserving way

Using the finite volume code **MPI-AMRVAC** (?), we solve the equations of hydrodynamics under their conservative form :

$$\partial_t \rho + \nabla \cdot (\rho \mathbf{v}) = 0 \quad (1)$$

$$\partial_t (\rho \mathbf{v}) + \nabla \cdot (\rho \mathbf{v} \mathbf{v} + P \mathbb{I}) = \rho \mathbf{f} - 2\mathbf{\Omega} \wedge \mathbf{v} \quad (2)$$

$$\partial_t e + \nabla \cdot [(e + P) \mathbf{v}] = -\rho \mathbf{v} \cdot \mathbf{f} \quad (3)$$

where ρ , \mathbf{v} , P and e are the mass density, velocity, pressure and total energy density respectively. $\mathbf{\Omega}$ is the orbital angular speed vector. \mathbf{f} is the modified Roche force per mass unit given by :

$$\mathbf{f} = \alpha(r_1) \frac{q}{r_1^3} \mathbf{r}_1 - \frac{1}{r_2^3} \mathbf{r}_2 + \frac{1+q}{a^3} \mathbf{r}_\perp \quad (4)$$

with the mass ratio $q = M_1/M_2$ \mathbf{r}_1 \mathbf{r}_2 \mathbf{r}_\perp α : see section 2.1.2. Encapsulates wind acceleration process and stellar gravity. -1 with only gravity, leading to the usual Roche force per unit mass.

The energy equation 3 is adiabatic. See section 2.2.2 for way to account for cooling. EOS ideal gas monoatomic with an adiabatic index $\gamma = 5/3$ with a mean molecular weight set to 1.

2.2.2 Cooling

Cooling time scale (?) Polytropic index Γ (?). Physical meaning of Γ . Bypass the energy equation 3 where the cooling time scale is XXX times smaller than the dynamical time scale.

Provided the density is high enough, condensation of the hot shocked flow into a disc (?, and references therein) : underlying principle of two component accretion flows models.

2.2.3 Numerical setup

2.3 Physical parameters

3 RESULTS

3.1 Inhomogeneity and asymmetry of the inflow

3.2 Flow morphology

3.2.1 Without cooling

3.2.2 With cooling

FOR BLACK HOLES : Presence of a disc-like structure which does not extend as far as in RLOF-fed systems (LMXB) => no hysteresis in hardness-intensity diagram

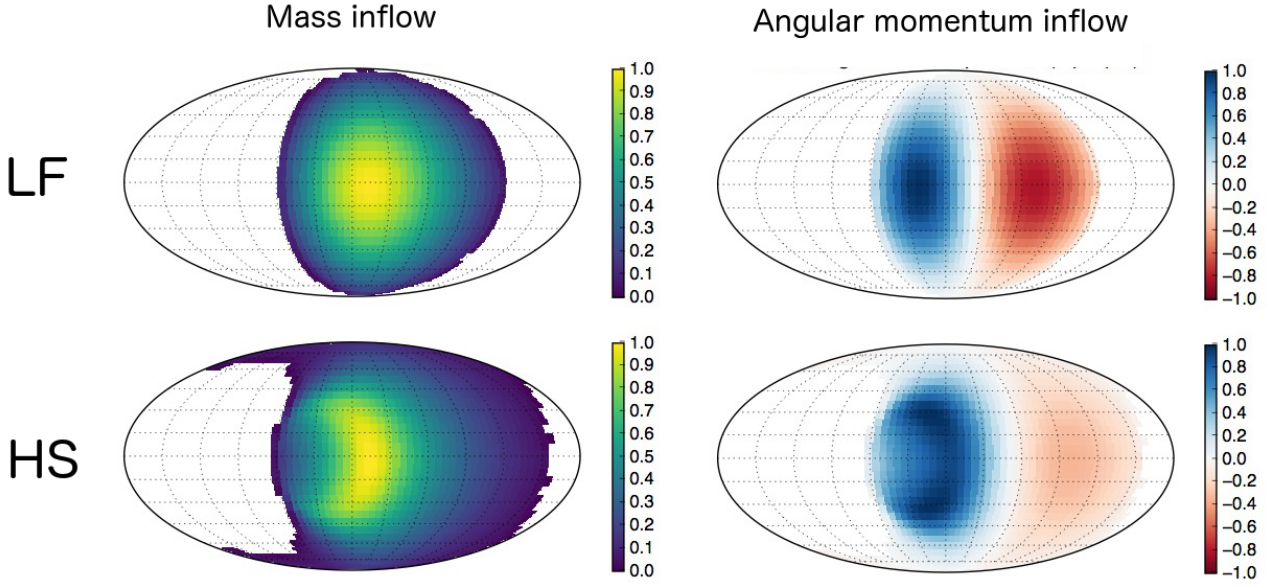


Figure 2. Mollweide projections of local mass and angular momentum inflows within the simulation space centered on the accretor (dashed green sphere on Figure 1). The upper row corresponds to the light fast (LF) case while the bottom row is for the heavy slow (HS) case. Each map is scaled to its maximum (absolute) value and centered on the axis from the accretor to the donor star. Positive (resp. negative) values of angular momentum stands for locally prograde (resp. retrograde) flow with respect to the orbital motion.

Table 1. Parameters and integrated quantities at the outer edge of the simulation space for the 2 models considered.

	LF	HS
M_1	$20.2M_\odot$	
R_1	$28.4R_\odot$	
P	8.964357 days	
\dot{M}_1	$6.3 \cdot 10^{-7} M_\odot \text{ yr}^{-1}$	
M_2	$1.5M_\odot$	$2.5M_\odot$
Boosted	Yes	No
$\dot{M}_{\text{out}}/\dot{M}_1$	XXX	XXX
$\dot{j}_{\text{out}}/\dot{j}_{\text{SL}}$	XXX	XXX
$R_{\text{circ}} / R_{\text{mag}}$	XXX	XXX

(for Cyg X-1, LMC X-1 or LMC X-3, the 3 wind-fed BH-HMXB). Indeed, the soft state might originate from : "A drop in the accretion rate affecting both flows would propagate through the halo immediately but might take up to several weeks to propagate through the disk. While the inner halo is thus temporarily depleted compared to the disk, a temporary soft state is expected." but if the disc has a much smaller outer radius (due to a much smaller angular momentum of the inflow), the viscous delay is expected to be so small that the dimming of the disc will be almost as fast as the one of the disc. (?). Explains also why no large outburst (ie a low contrast between the brightest and dimmest X-ray emission) in Cygnus X-1 (x3) (?) : without an outer cool disc, the thermal instability resulting from the ionization of Hydrogen cannot occur (REF?). No quiescent state in Cyg X-1 : it is an argument in favor of the existence of a cool disc in the hard state.

3.3 Mass and angular momentum accretion rates

If halted mass accretion, might be due to radiative heating from the X-ray source (?) : if circularization radius > 0.04

times the Bondi radius (for us, the accretion radius?), much lower accretion rate than Bondi (for us, BHL?). But their work depends also on isotropic or anisotropic X-ray source, alpha viscosity parameter, etc...

3.4 Disc mass and morphology

4 CONCLUSION

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