

A MONTE CARLO SIMULATION OF THE STARS OF THE GALAXY

IMF: CHABRIER SINGLE 2005

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CHABRIER S 2005

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INTRODUCTION

Monte Carlo statistical model

**Simulated stars using IMF of
Chabrier 2005**

**Diverse sample of stars:
100, 1.000, 10.000, 100.000 and
1 million**

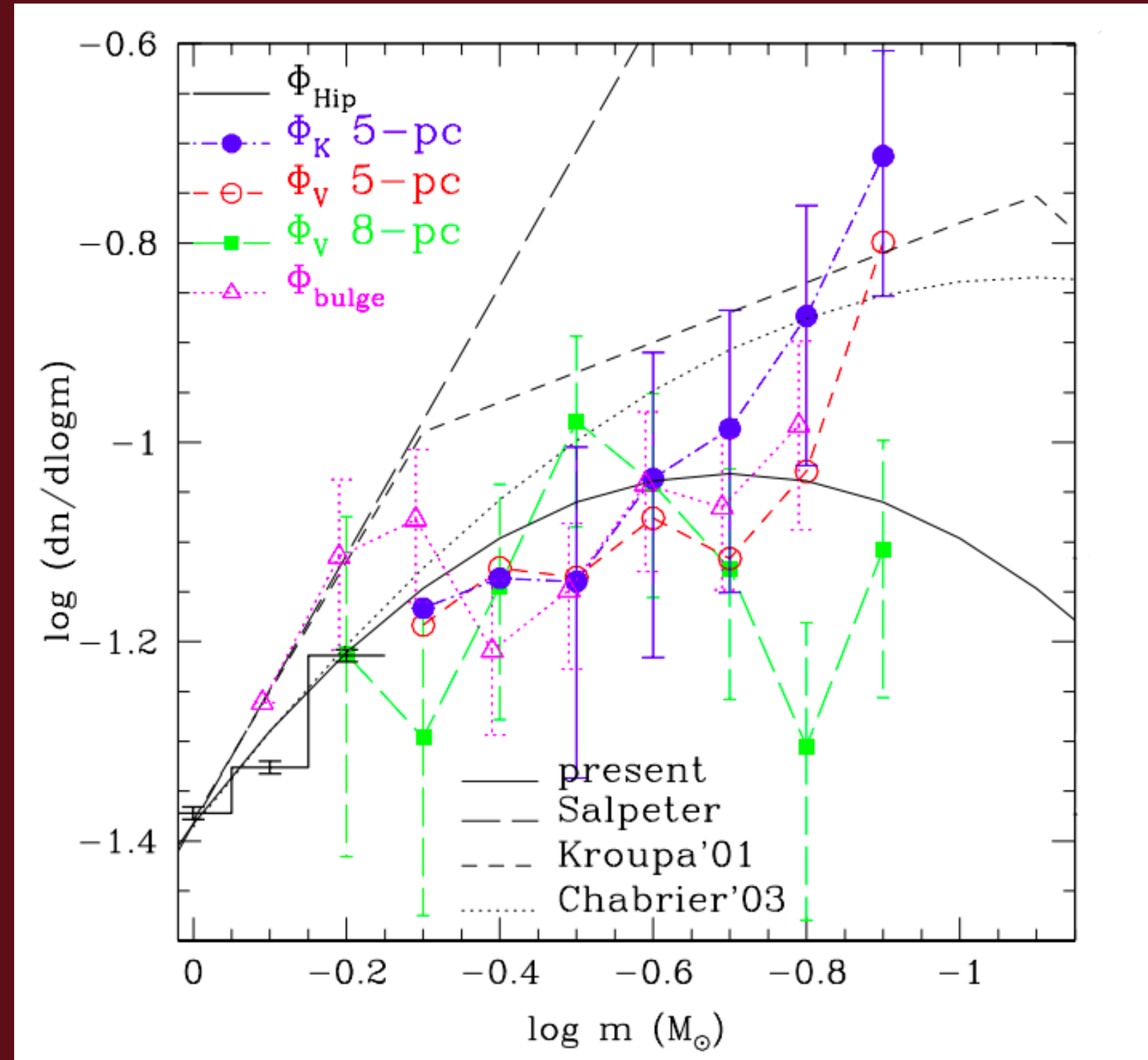
**Study of stars on the MS and
the object they will become
upon leaving it**

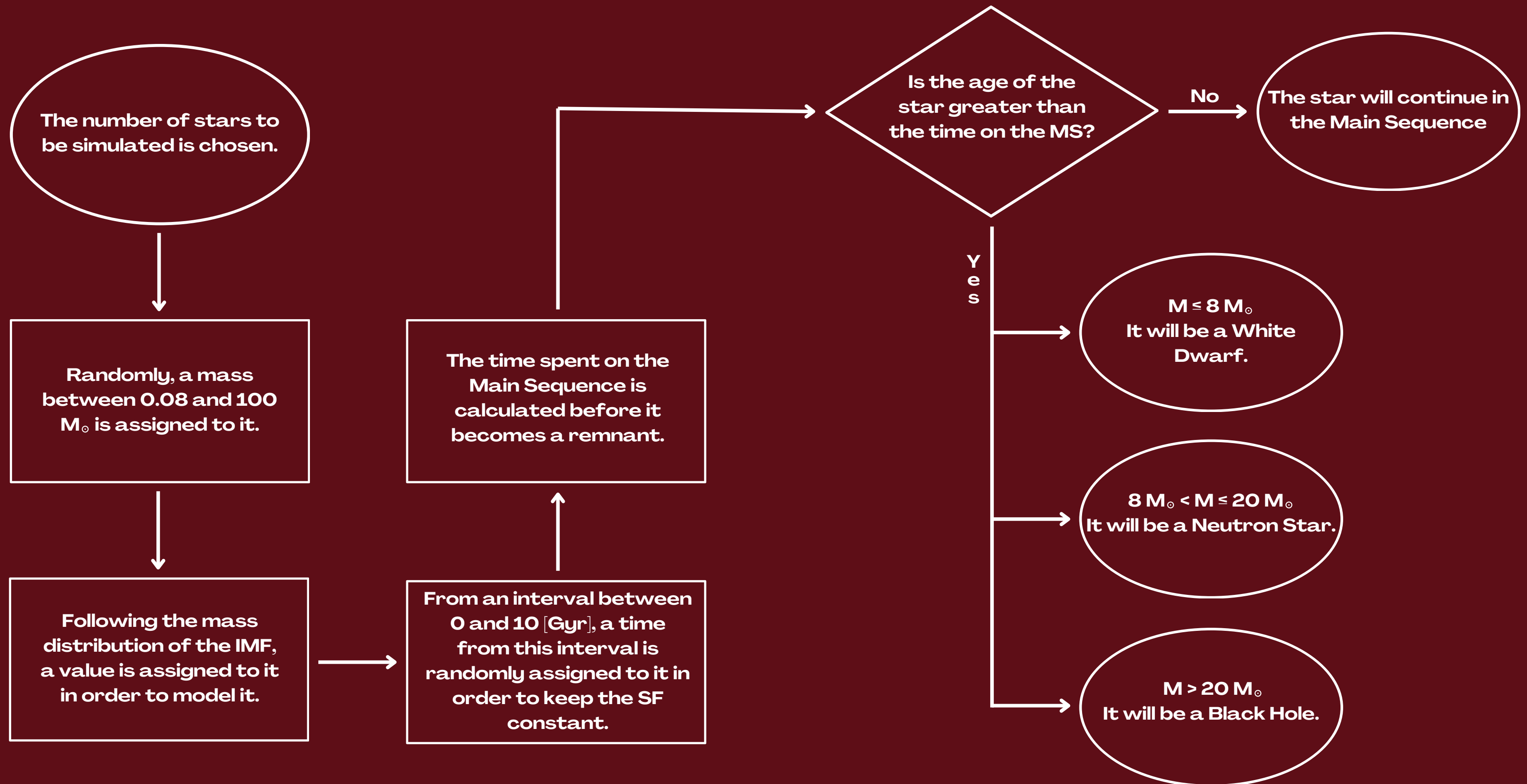
Initial Mass Function (IMF)

Chabrier S 2005

- Focused on individual stars
- It arises with the aim of properly characterizing brown dwarfs.
- Its mass distribution is more constant than that of the other IMFs.
- The log-normal form for low masses is well supported observationally

$$\xi(\log m) = 0.093 \times \exp \left[\frac{-(\log m - \log 0.2^2)}{2 \times (0.55)^2} \right], \quad m \leq 1M_{\odot}$$
$$= 0.041 m^{-1.35}, \quad m \geq 1M_{\odot}$$

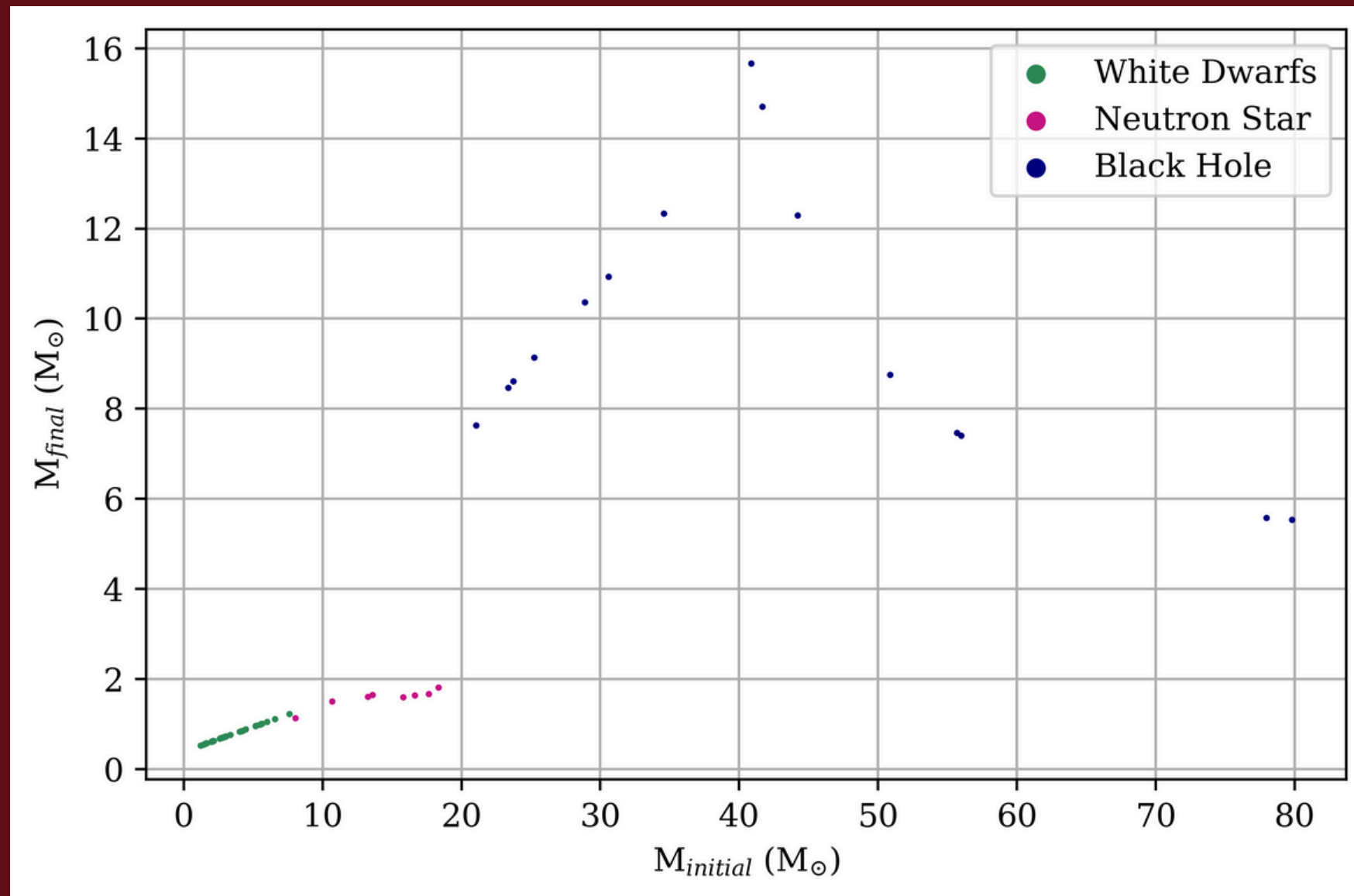




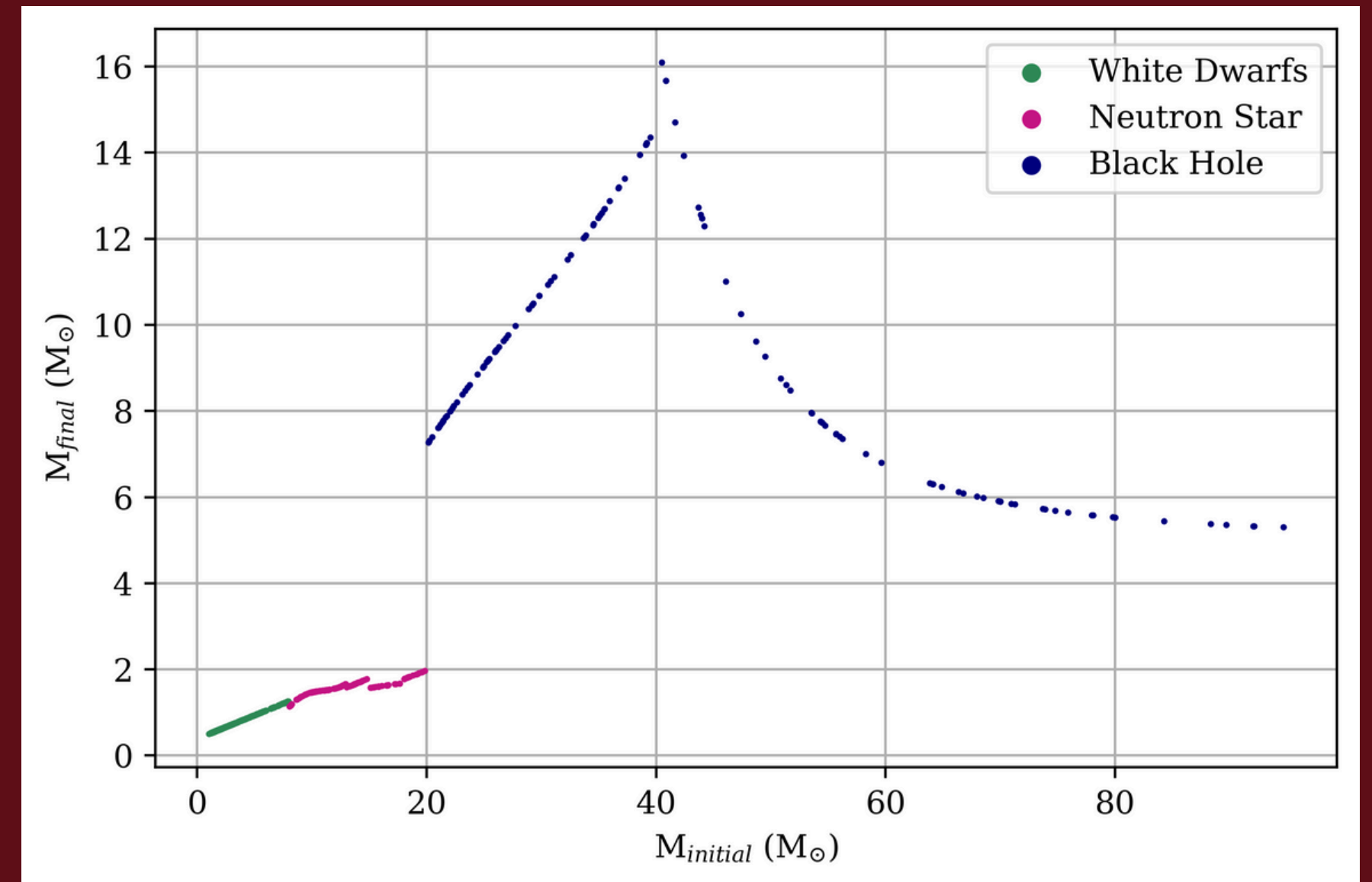
FLOWCHART

OUR RESULTS

RESIDUAL MASSES

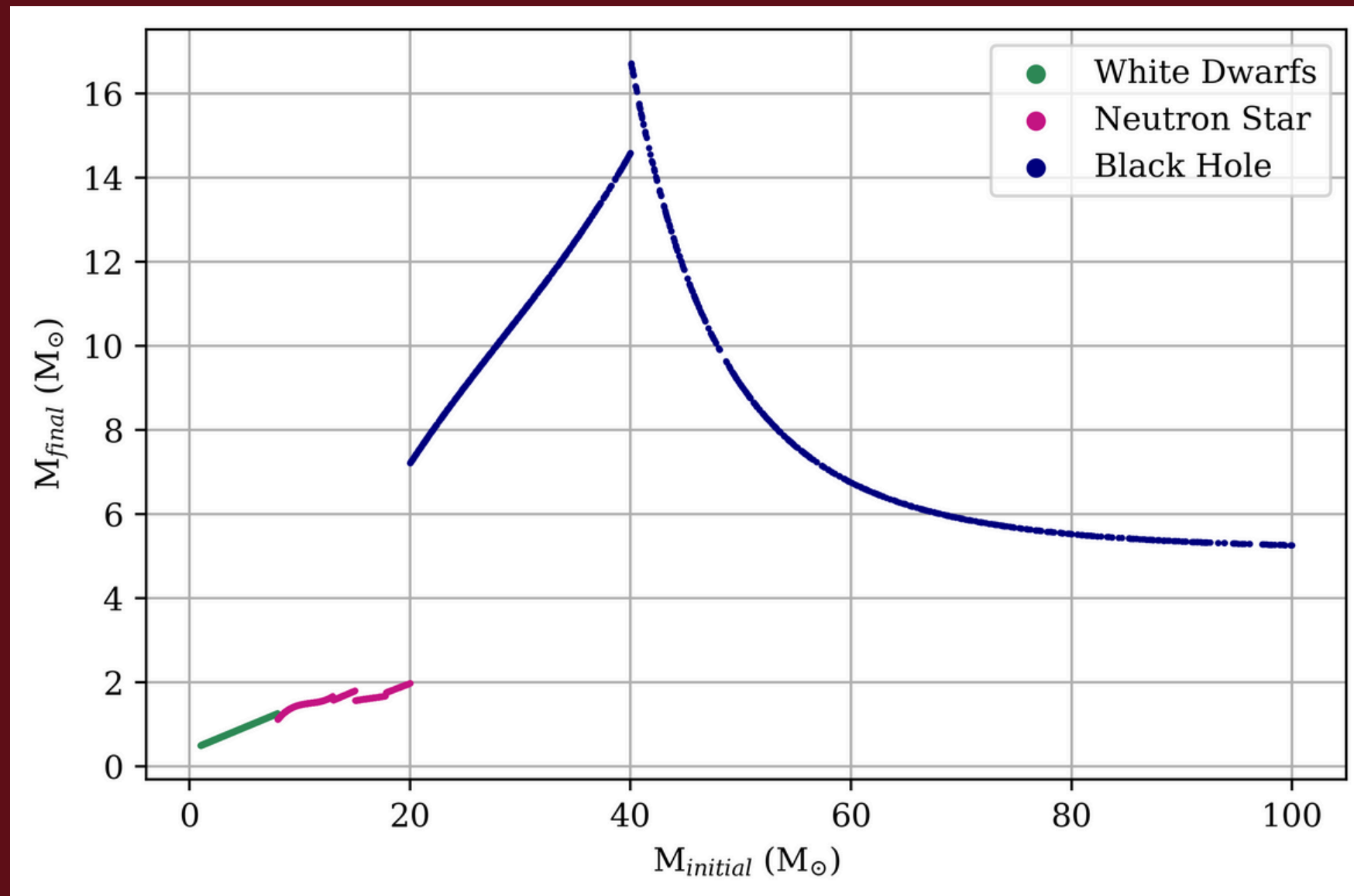


100 Stars

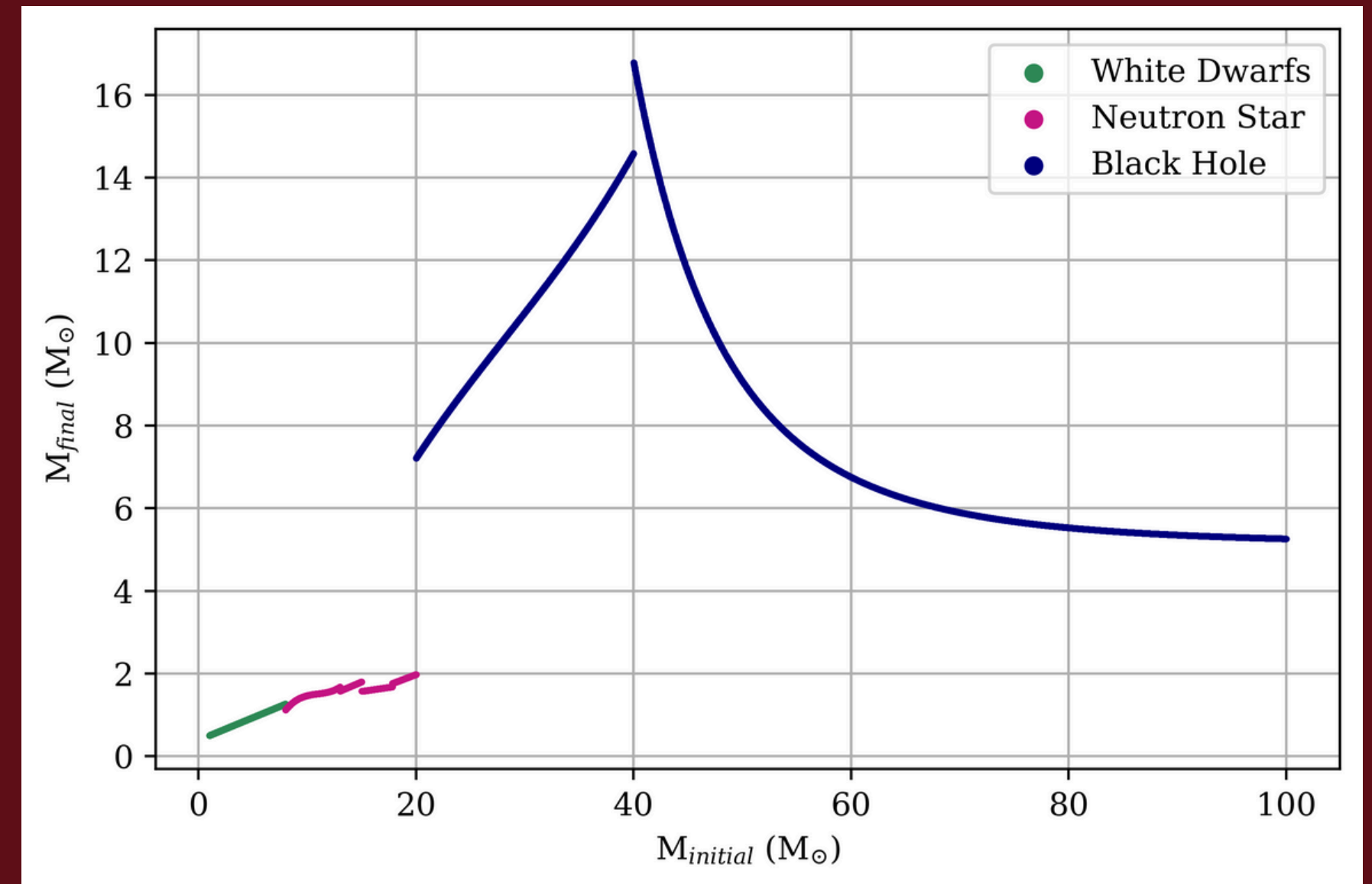


1.000 Stars

RESIDUAL MASSES

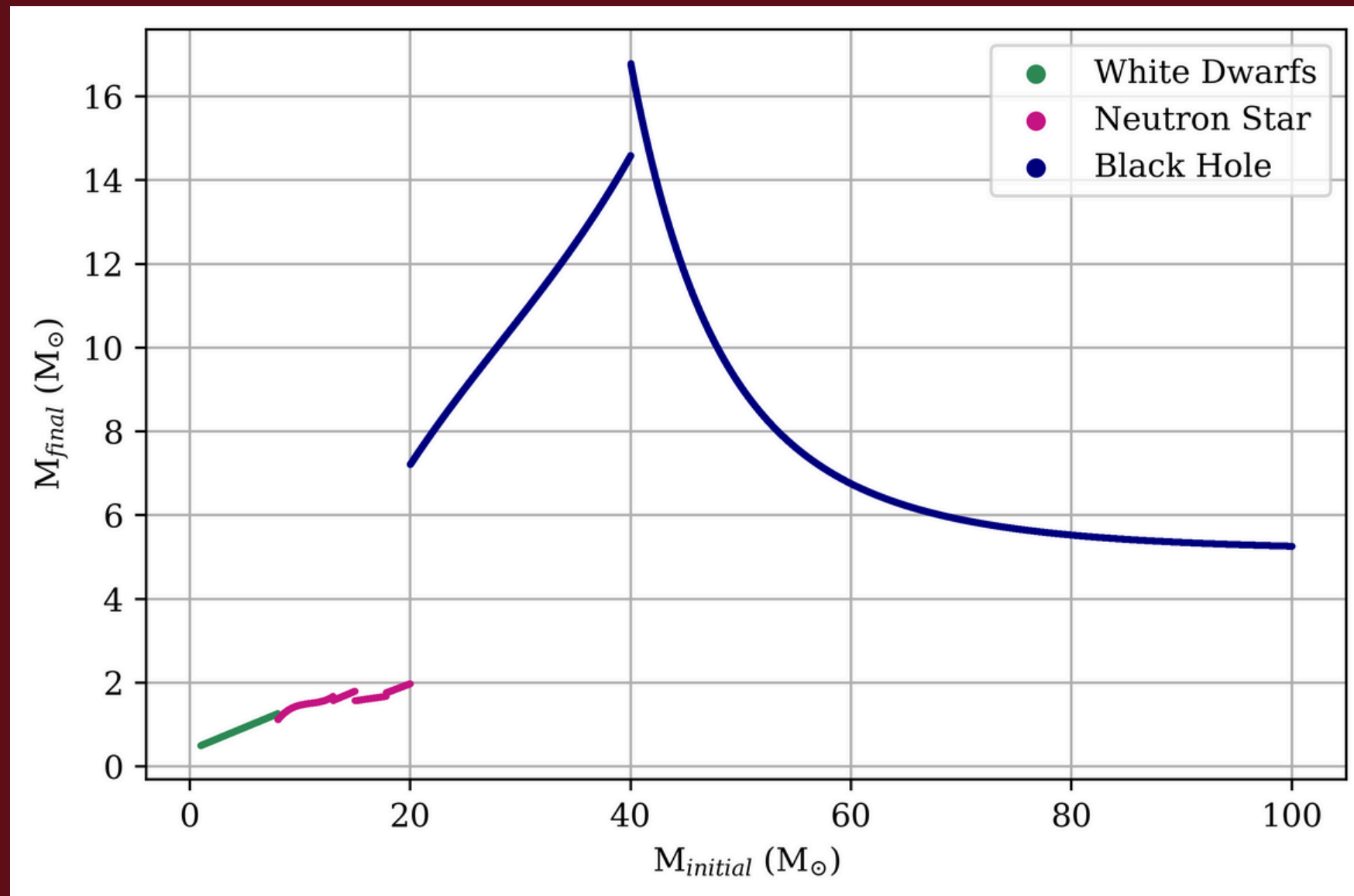


10.000 Stars



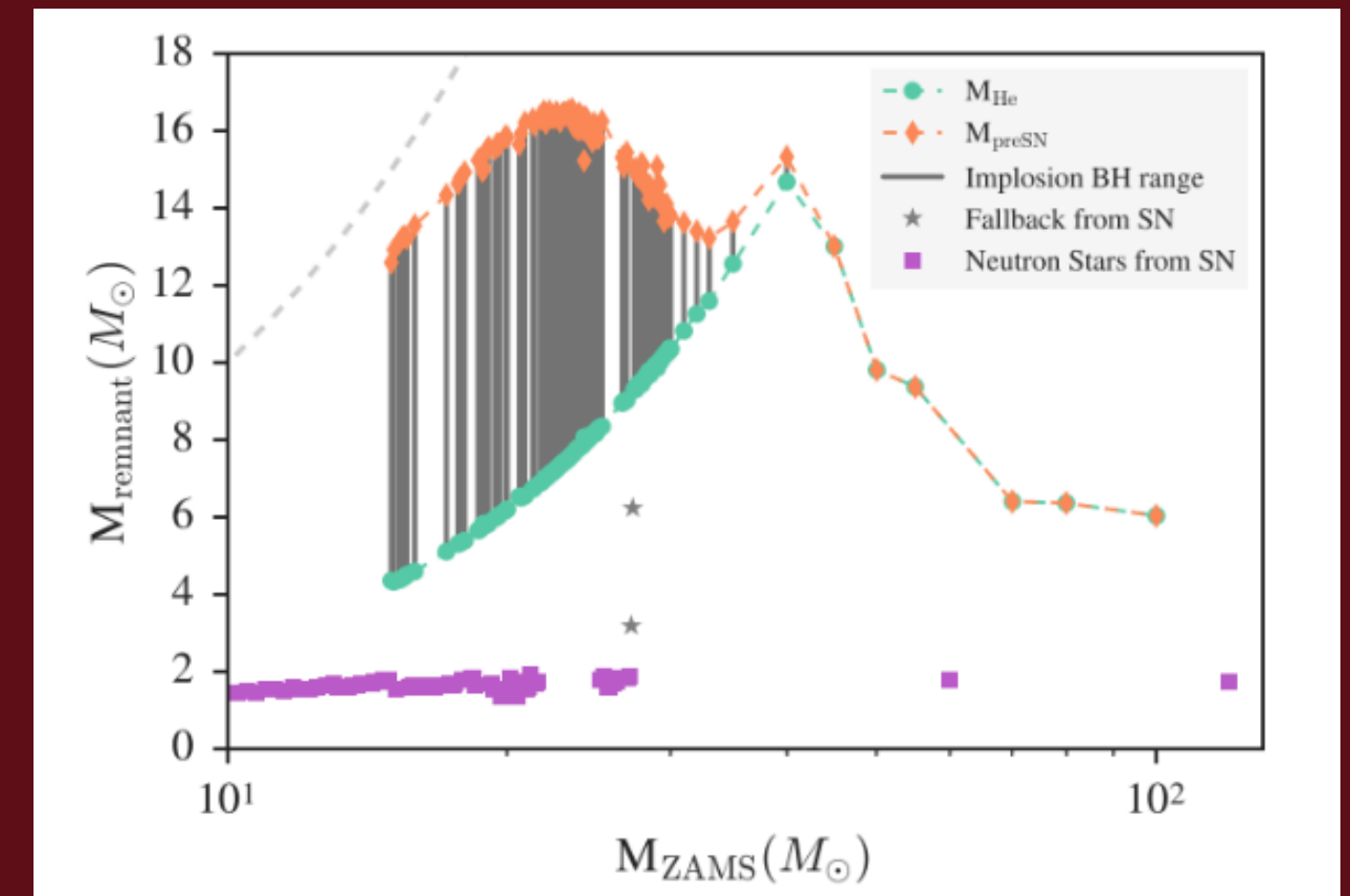
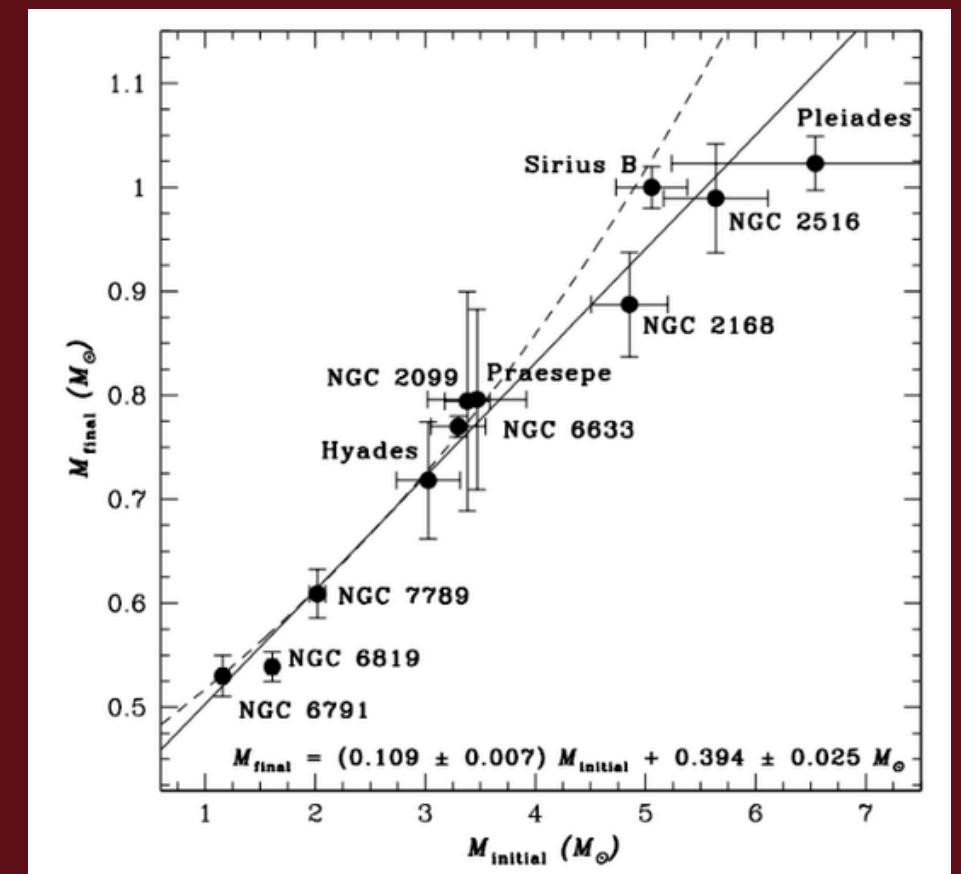
100.000 Stars

RESIDUAL MASSES



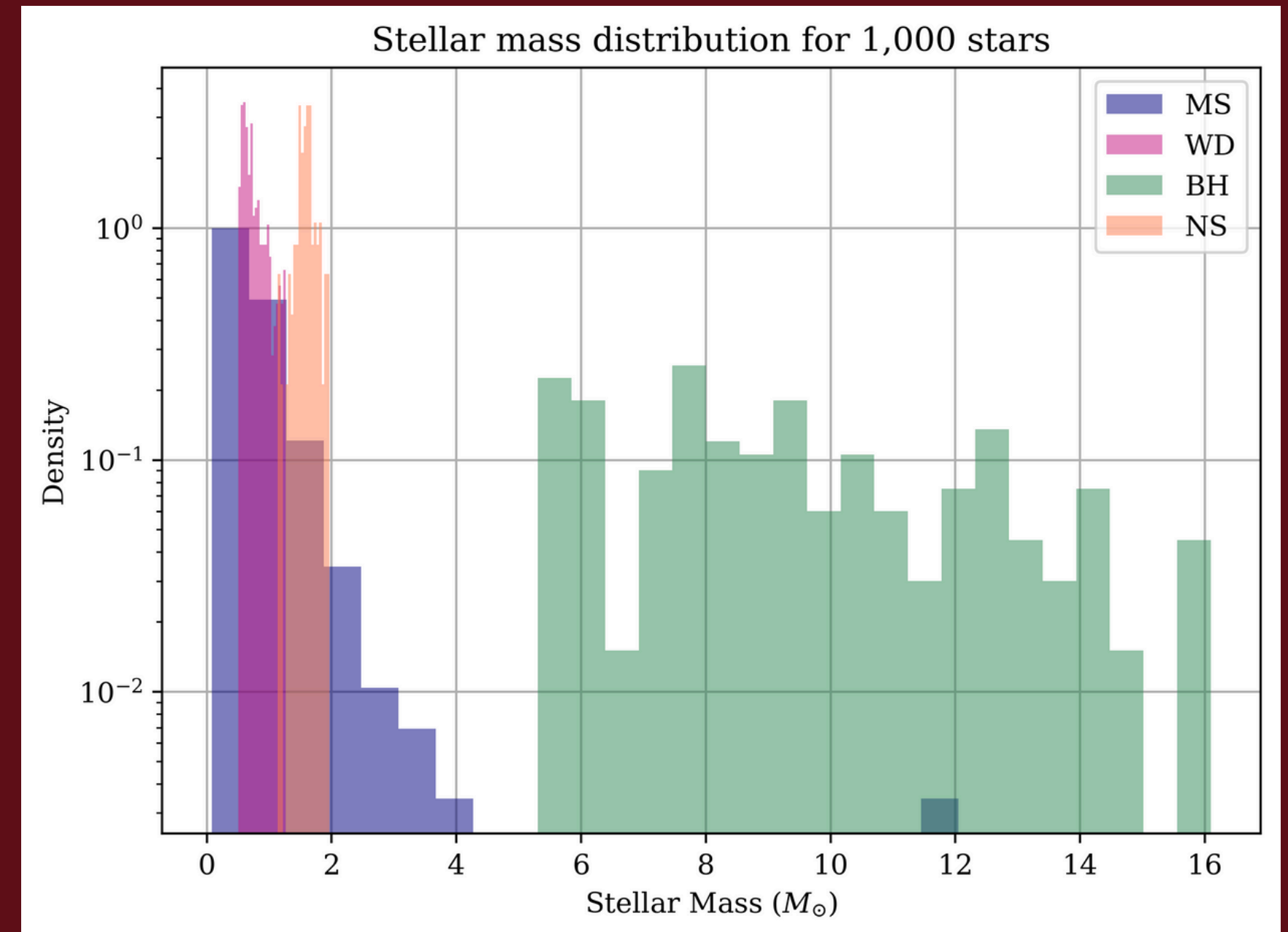
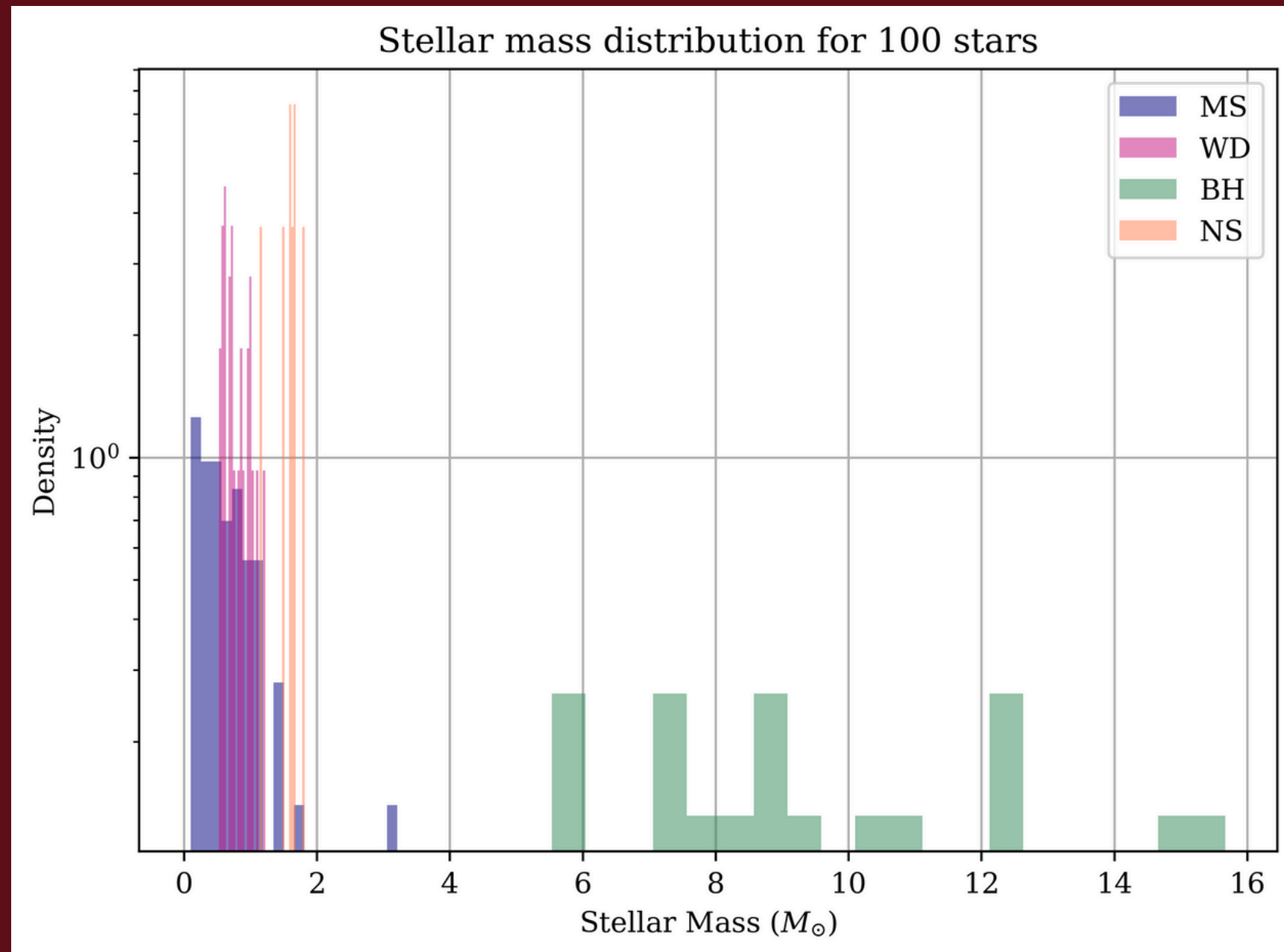
1.000.000 Stars

Kalirai et al. (2008)

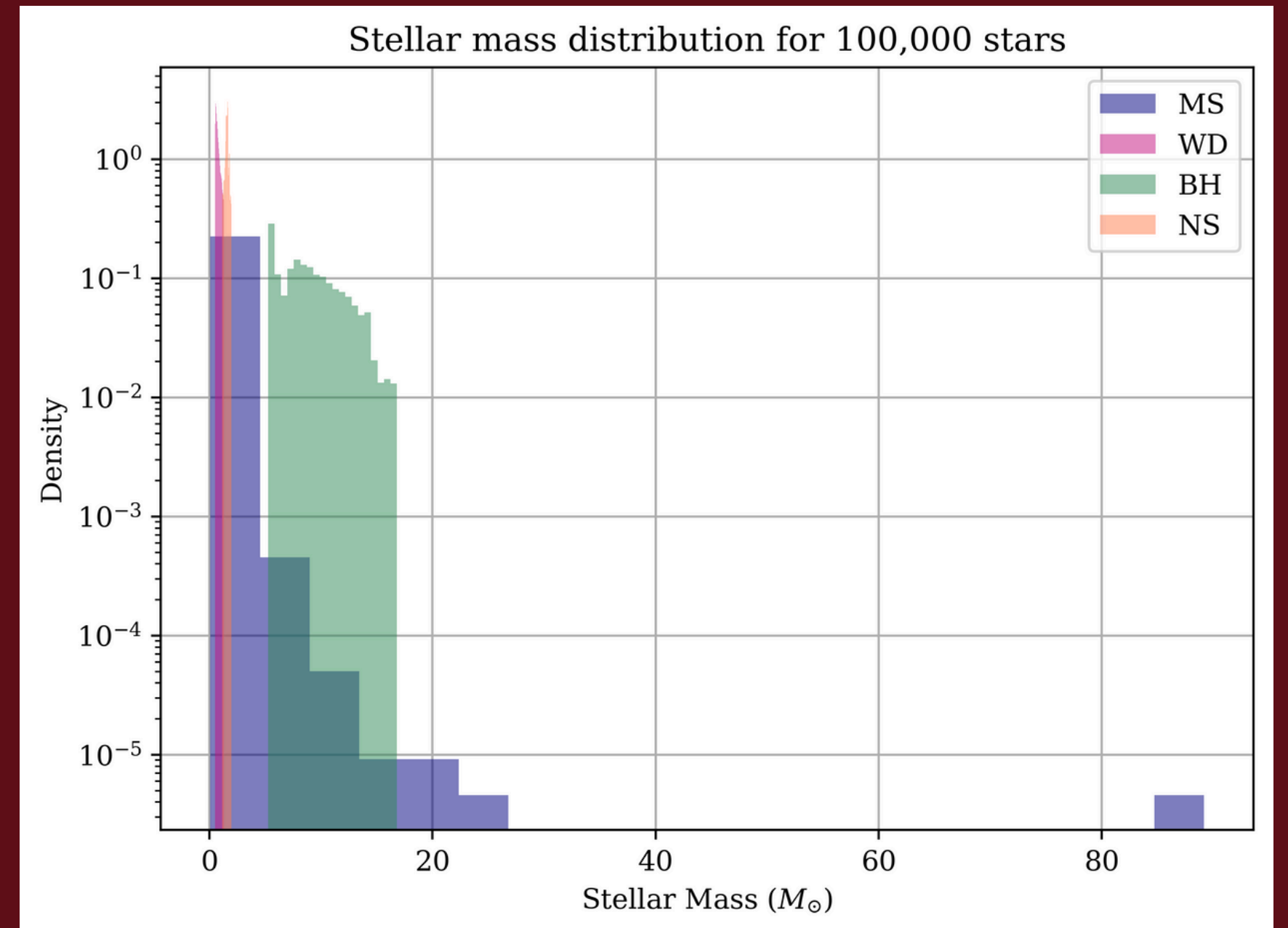
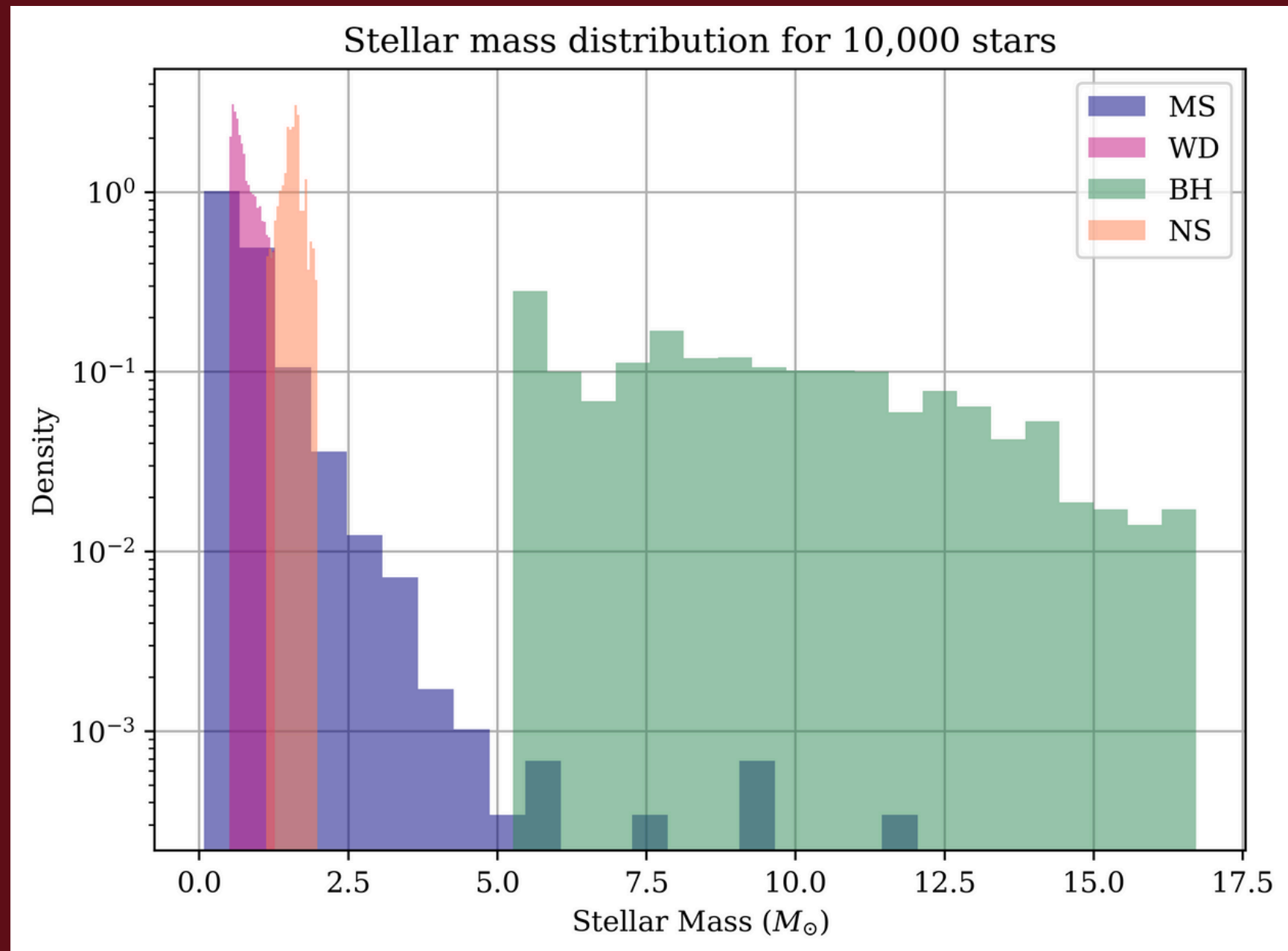


Raithel et al. (2018)

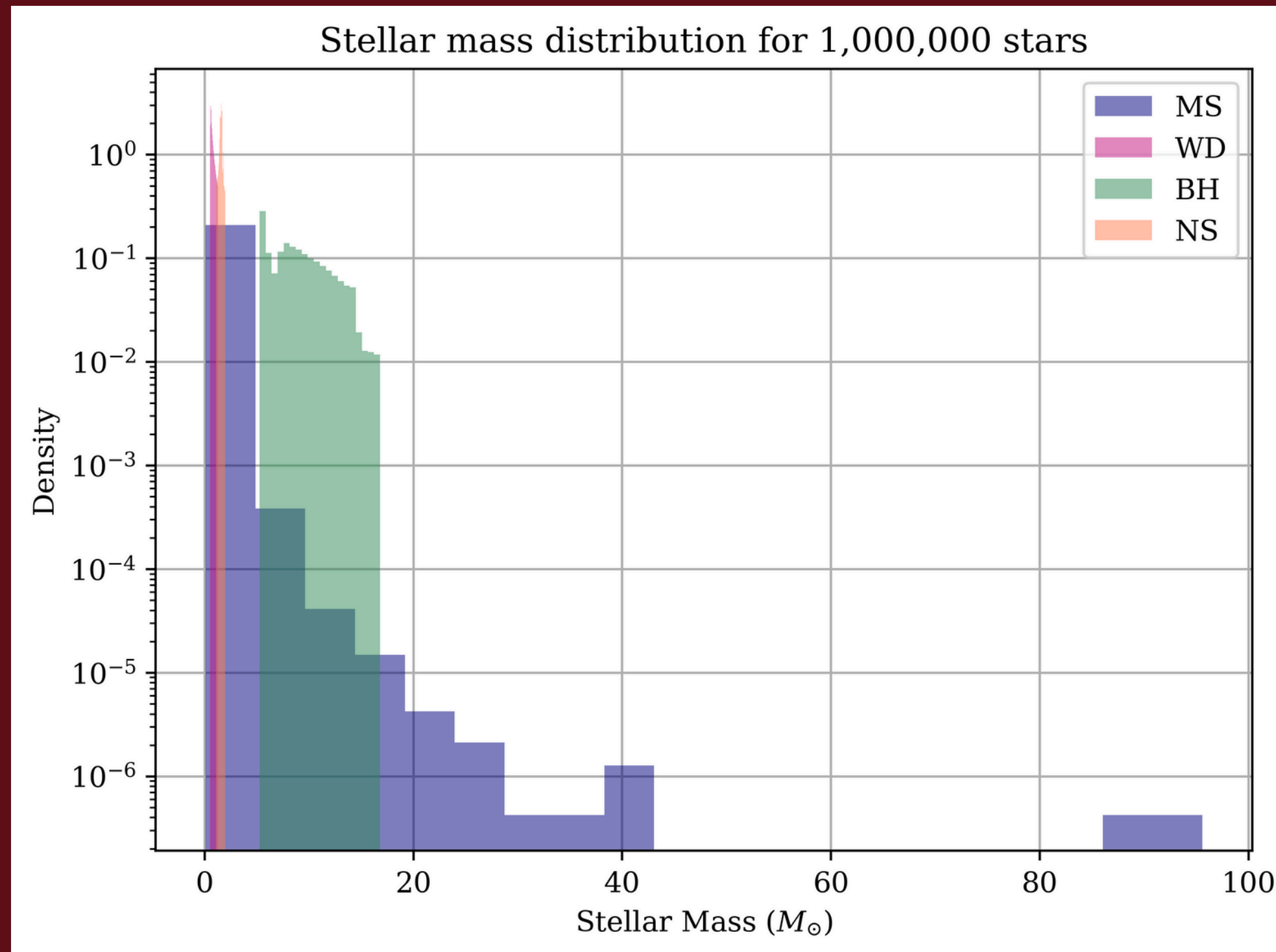
FINAL MASSES DISTRIBUTION



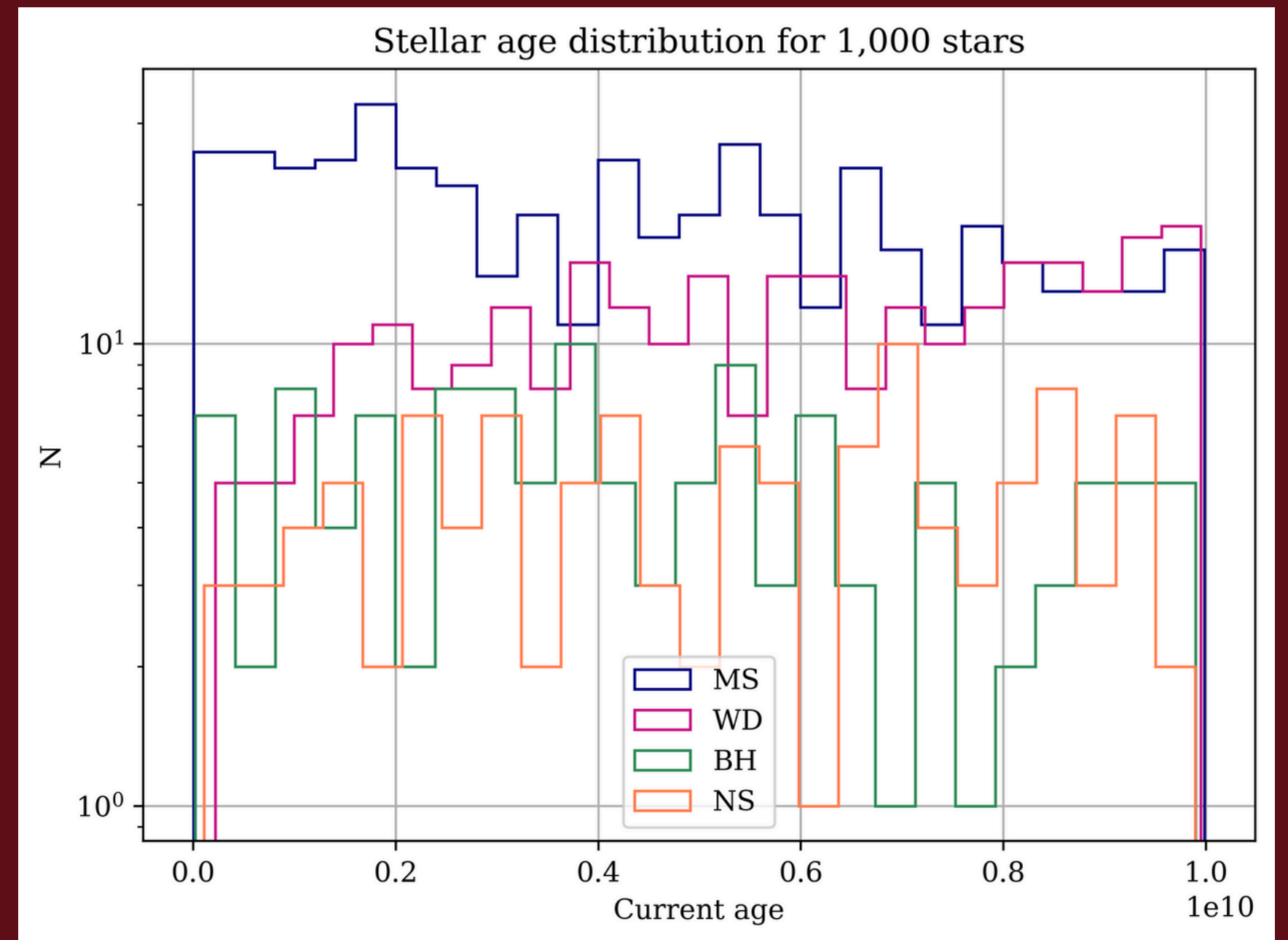
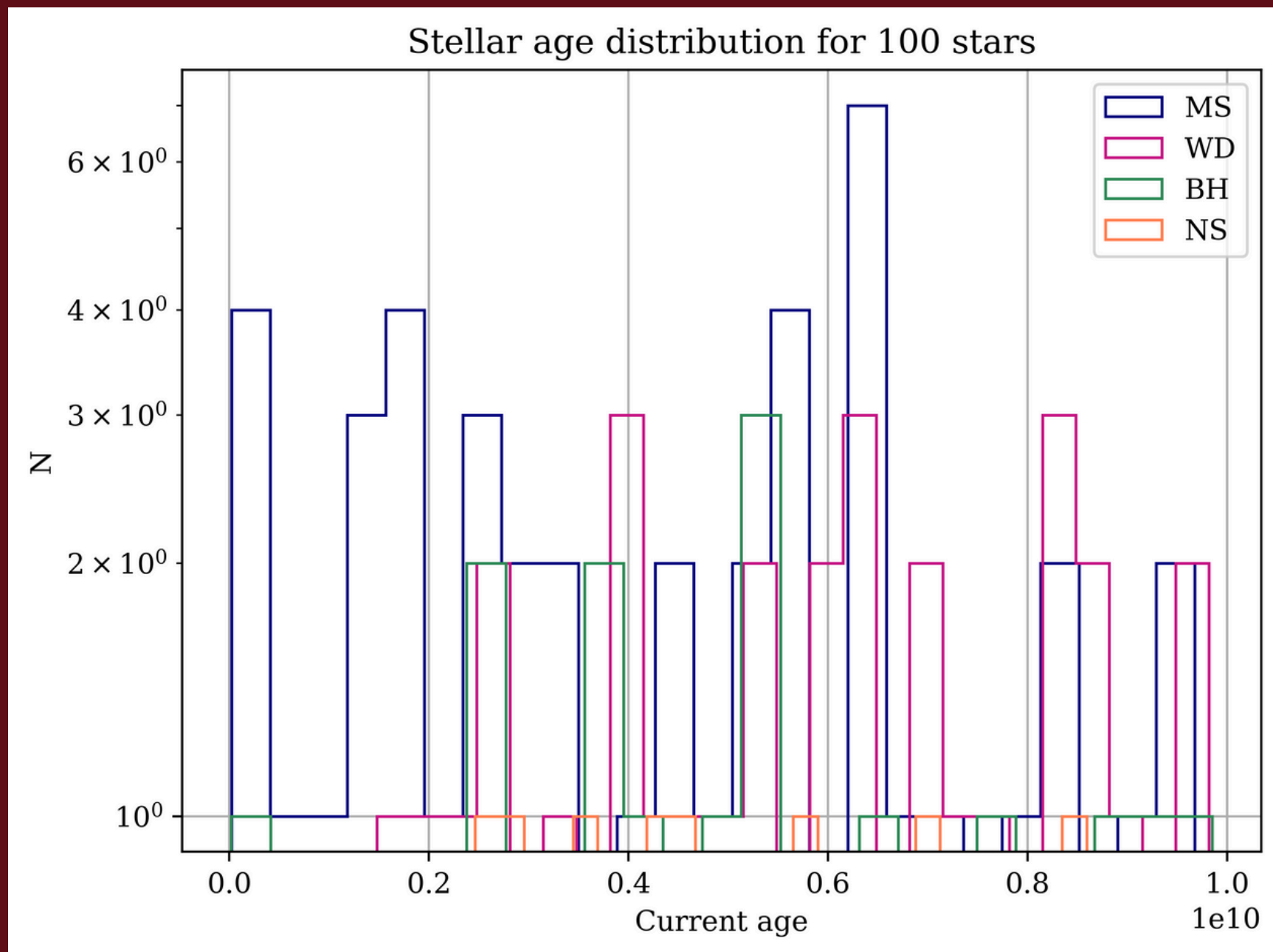
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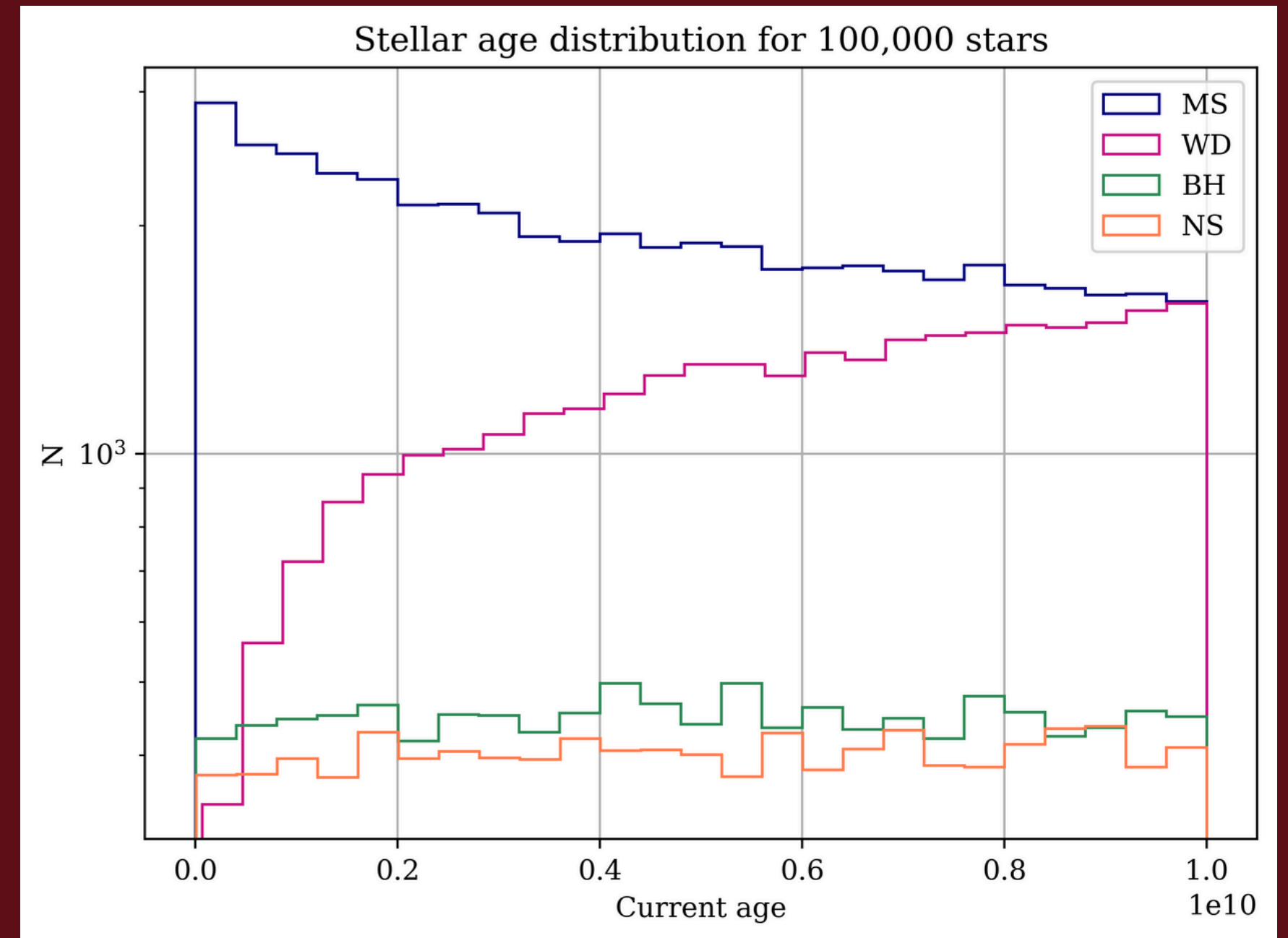
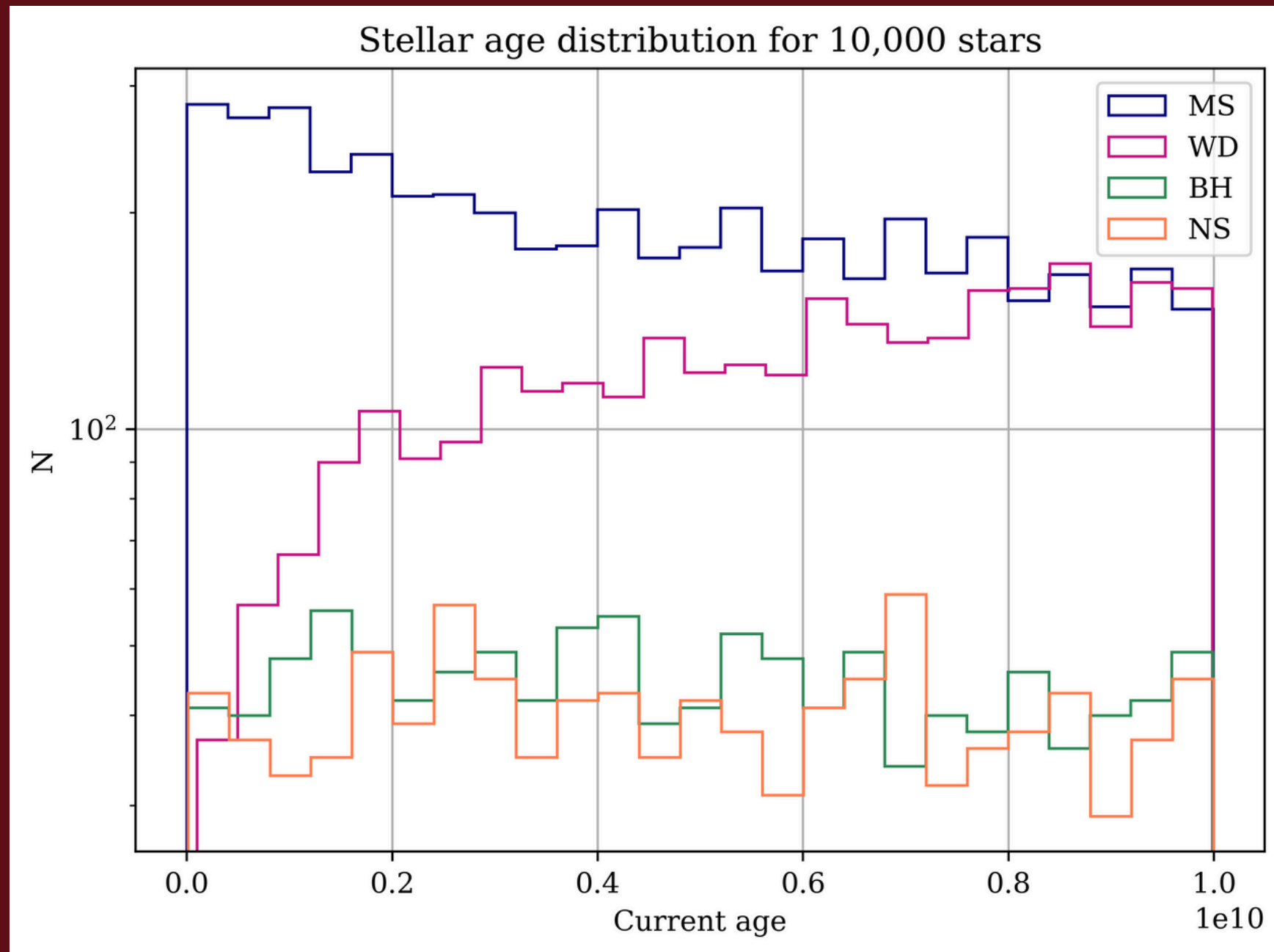
FINAL MASSES DISTRIBUTION



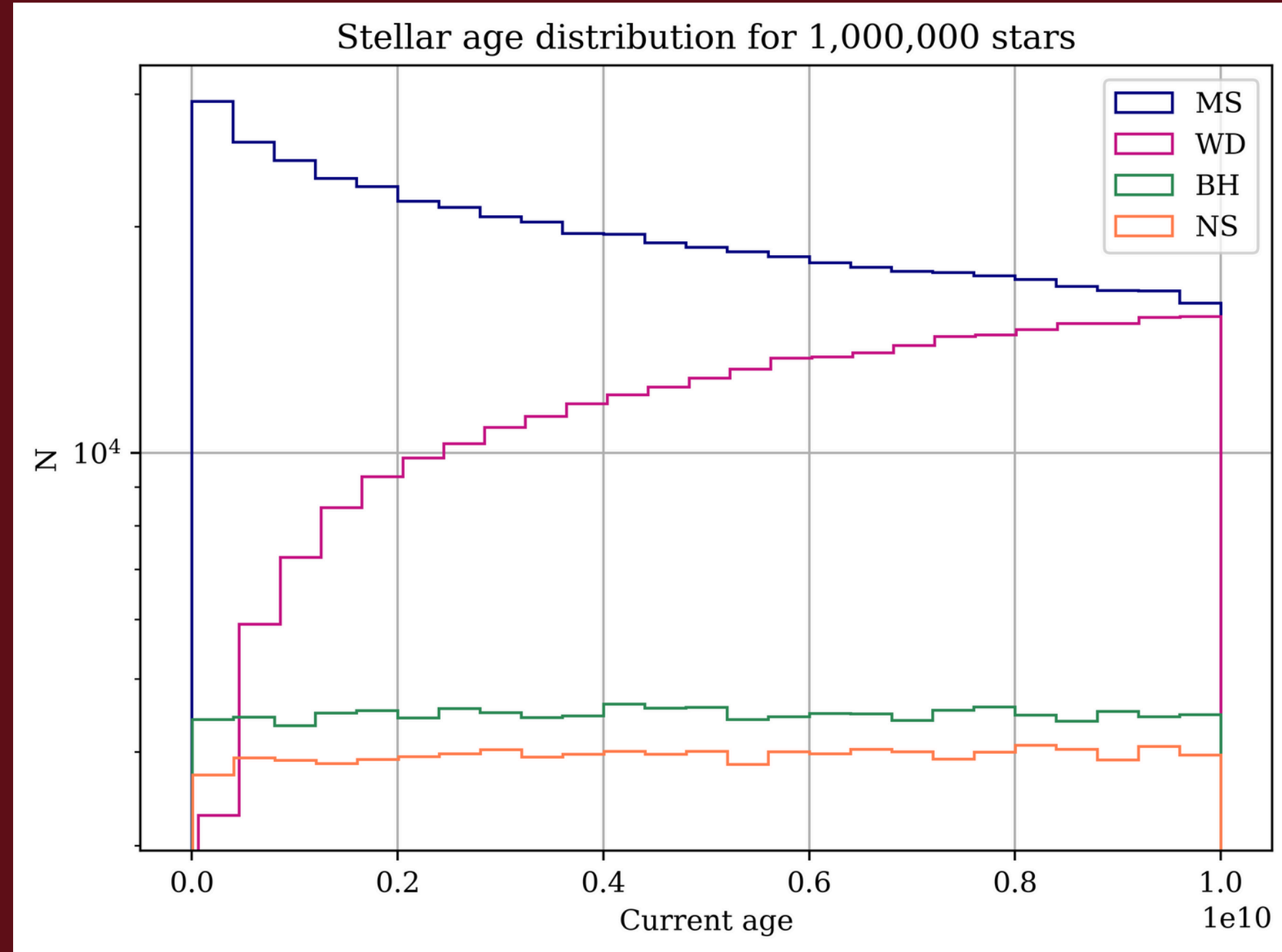
AGE DISTRIBUTION



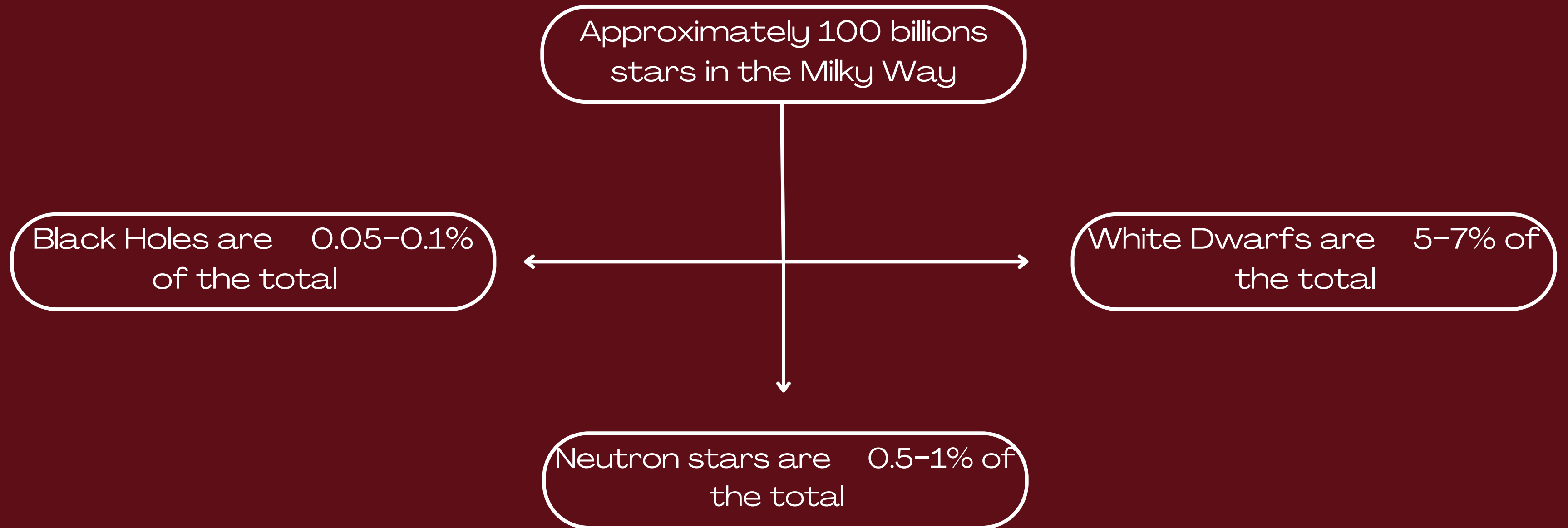
AGE DISTRIBUTION



AGE DISTRIBUTION



COMPARISON WITH THE MW



COMPARISON WITH THE MW

Number of stars	BH fraction	NS fraction	WD fraction	MS fraction
100	15.00 %	8.00%	31.00%	46.00%
1.000	12.30%	11.40%	28.10%	48.20%
10.000	11.16%	10.09	29.99%	48.76%
100.000	11.24%	10.06%	29.68%	49.02%
1.000.000	11.22 %	9.90%	29.54%	49.33%

CONCLUSIONS

Using the Monte Carlo statistical model, it was possible to generate various analyses for the stellar distribution based on the initial mass function of Chabrier (2005).

It was possible to observe the impact this IMF has on the stellar mass distribution, which is the objective for which it was created.

The results obtained indicate that it is indeed a good regime for low-mass stars, such as brown dwarfs, but not for more massive stars, ultimately predicting a much higher value than expected for BHs in the Milky Way. This value is superior to 10% in all cases.

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SCAN ME



Project 4

THANK YOU!

Questions?