

Scientific Research: Report Assignment

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# **Report Assignment**

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# 1.

In an non-interacting electron gas, the electrons fill up the momentum space from the lower energy states up to the Fermi energy  $E_F$ .

$$N(E_F) = N_{\text{tot}} \quad (1)$$

At the Fermi energy, the electrons have a corresponding wave number  $k_F$ :

$$E_F = \frac{\hbar^2 k_F^2}{2m} \iff k_F = \frac{\sqrt{2mE_F}}{\hbar} \quad (2)$$

and the total number of electrons in the gas is given by

$$N_{\text{tot}} = \int_{|k| < k_F} dN = \int_0^{k_F} dk \frac{V}{\pi^2} k^2 = \frac{V}{3\pi^2} k_F^3 \quad (3)$$

The energy of a relativistic particle is given by

$$E(p) = \sqrt{p^2 c^2 + m^2 c^4} \quad (4)$$

in terms of the wave number,

$$E(k) = \sqrt{(\hbar k)^2 c^2 + m^2 c^4} \quad (5)$$

So the total energy of the electron gas is given by

$$E_{\text{tot}} = \int_0^N dN E(k) \quad (6)$$

$$= \int_0^{k_F} dk \frac{V}{\pi^2} k^2 E(k) \quad (7)$$

$$\implies \varepsilon_e = \frac{1}{\pi^2} \int_0^{k_F} dk k^2 \sqrt{(\hbar k)^2 c^2 + m_e^2 c^4} \quad (8)$$

now set  $k = \frac{m_e c}{\hbar} \sinh \theta$  (i.e.  $x_F = \sinh \theta_F$ ), and so

$$dk = \frac{m_e c}{\hbar} \cosh \theta d\theta, \quad E(\theta) = m_e c^2 \cosh \theta \quad (9)$$

then the integral becomes

$$\varepsilon_e = \frac{1}{\pi^2} \int_0^{\theta_F} d\theta \left( \frac{m_e c}{\hbar} \cosh \theta \right) \left( \frac{m_e c}{\hbar} \sinh \theta \right)^2 (m_e c^2 \cosh \theta) \quad (10)$$

$$= \frac{m_e^4 c^5}{\pi^2 \hbar^2} \int_0^{\theta_F} d\theta \cosh^2 \theta \sinh^2 \theta \quad (11)$$

To calculate this integral, let us use a few formulas for hyperbolic functions:

$$\sinh x \cosh x = \frac{1}{2} \sinh(2x), \quad \sinh^2 x = \frac{1}{2} (\cosh(2x) - 1), \quad \cosh(2x) = 1 + 2 \sinh^2 x \quad (12)$$

using these, we can rewrite the integral as

$$\int_0^{\theta_F} d\theta \cosh^2 \theta \sinh^2 \theta = \frac{1}{8} \int_0^{\theta_F} d\theta \cosh(4\theta) - 1 \quad (13)$$

$$= \frac{1}{8} \left[ \frac{1}{4} \sinh(4\theta) - \theta \right]_0^{\theta_F} \quad (14)$$

$$= \frac{1}{8} \left[ \frac{1}{4} \sinh(4\theta_F) - \theta_F \right] \quad (15)$$

and the expressino inside the brackets can be simplified further:

$$\frac{1}{4} \sinh(4\theta) = \frac{1}{2} \sinh(2\theta) \cosh(2\theta) \quad (16)$$

$$= \sinh \theta \cosh \theta (1 + 2 \sinh^2 \theta) \quad (17)$$

here, since  $x_F = \sinh \theta_F$ ,

$$\cosh \theta_F = \sqrt{1 + x_F^2}, \quad \theta_F = \operatorname{arsinh} x_F = \ln(x_F + \sqrt{1 + x_F^2}) \quad (18)$$

thus,

$$\varepsilon_e = \frac{m_e^4 c^5}{8\pi^2 \hbar^2} \left( \frac{1}{4} \sinh(4\theta_F) - \theta_F \right) \quad (19)$$

$$= \frac{m_e^4 c^5}{8\pi^2 \hbar^2} \left( \sinh \theta_F \cosh \theta_F (1 + 2x_F^2) - \ln(x_F + \sqrt{1 + x_F^2}) \right) \quad (20)$$

## 2.

Table 1: Properties of 12 sequential nuclides in the outer crust of a neutron star.

Nuclide	Z	N	A	$n_B$ [fm <sup>-3</sup> ]	$M_{\text{cal}}$ [MeV]	$n_N M_{\text{nuc}} c^2$	$\epsilon'_e$ [MeV fm <sup>-3</sup> ]	$\epsilon_L$	Total $\epsilon(n_B)$
<sup>56</sup> Fe	26	30	56	$4.75 \times 10^{-8}$	-60.61	0.043916	$1.83 \times 10^{-5}$	$-2.71 \times 10^{-7}$	0.043940
<sup>62</sup> Ni	28	34	62	$1.61 \times 10^{-6}$	-66.75	0.047870	$2.65 \times 10^{-4}$	$-3.53 \times 10^{-6}$	0.048184
<sup>64</sup> Ni	28	36	64	$7.87 \times 10^{-6}$	-67.09	0.050479	$8.84 \times 10^{-4}$	$-1.04 \times 10^{-5}$	0.051448
<sup>86</sup> Kr	36	50	86	$1.83 \times 10^{-5}$	-82.90	0.053181	$1.83 \times 10^{-3}$	$-2.04 \times 10^{-5}$	0.055171
<sup>84</sup> Se	34	50	84	$6.29 \times 10^{-5}$	-82.35	0.060155	$5.25 \times 10^{-3}$	$-5.19 \times 10^{-5}$	0.065747
<sup>82</sup> Ge	32	50	82	$1.66 \times 10^{-4}$	-79.62	0.068225	$1.18 \times 10^{-2}$	$-1.10 \times 10^{-4}$	0.080517
<sup>80</sup> Zn	30	50	80	$3.60 \times 10^{-4}$	-74.96	0.077259	$2.27 \times 10^{-2}$	$-2.02 \times 10^{-4}$	0.100913
<sup>78</sup> Ni	28	50	78	$4.80 \times 10^{-4}$	-66.86	0.081182	$2.80 \times 10^{-2}$	$-2.43 \times 10^{-4}$	0.109893
<sup>124</sup> Mo	42	82	124	$7.30 \times 10^{-4}$	-85.73	0.084478	$4.50 \times 10^{-2}$	$-3.42 \times 10^{-4}$	0.130638
<sup>122</sup> Zr	40	82	122	$1.25 \times 10^{-3}$	-84.38	0.091391	$7.04 \times 10^{-2}$	$-5.16 \times 10^{-4}$	0.163351
<sup>120</sup> Sr	38	82	120	$2.10 \times 10^{-3}$	-81.98	0.100230	$1.08 \times 10^{-1}$	$-7.53 \times 10^{-4}$	0.210080
<sup>118</sup> Kr	36	82	118	$2.53 \times 10^{-3}$	-77.41	0.104646	$1.25 \times 10^{-1}$	$-8.54 \times 10^{-4}$	0.231142

The structure of a neutron star's outer crust is determined by several important physical processes. As the density increases, nuclei become progressively more neutron-rich through neutron capture processes. As a result that can be seen in Table 1, the neutron/proton increases from 1.15 in <sup>56</sup>Fe to 2.16 in <sup>118</sup>Kr.

Regarding energy contributions, the nuclear mass energy is the largest component of the total energy density. However, the kinetic energy of the electrons grows much more rapidly with density. A much smaller component, the lattice energy, is a negative correction that accounts for the electrostatic arrangement of nuclei into a crystal lattice.

This sequence of increasingly dense and neutron-rich nuclei ends at the neutron drip point, where the neutron is no longer energetically bound to the nucleus.