

Measuring precise radial velocities and cross-correlation function line-profile variations using a Skew Normal density [★]

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ABSTRACT

Context. Stellar activity is one of the primary limitations to the detection of low-mass exoplanets using the radial-velocity (RV) technique. Stellar activity can be probed by measuring time dependent variations in the shape of the cross-correlation function (CCF), often estimated using different parameters of the modeled CCF. Therefore estimating the moments of the CCF with high precision is essential to de-correlate the signal of an exoplanet from spurious RV signals originating from stellar activity.

Aims. We propose to estimate the parameters of the CCF by fitting a Skew Normal (SN) density shape which. Unlike the commonly employed Normal density, the SN density includes a skewness parameter to capture the asymmetry of the CCF induced by stellar activity and also the natural asymmetry induced by convective blueshift. Moreover, the SN distribution allows for different location parameters beyond the mean, such as the median of the CCF.

Methods. The performances of the proposed method are compared to the commonly employed Normal density, using both simulations and real observations, with different levels of activity and signal-to-noise ratio (SNR) levels.

Results. When considering real observations, the correlation between the RVs and the asymmetry of the CCF and the correlation between the RVs and the width of the CCF are stronger when using the parameters estimated with the SN rather than the ones obtained with the commonly employed Normal density. In particular the strongest correlations have been obtained when using as location parameter the mean of the SN. This suggests that the asymmetry of the CCF and the width of the CCF estimated using a SN may be more sensitive to stellar activity, which can be helpful when estimating stellar rotational periods and generally for characterizing stellar activity signals. The estimated uncertainties in the measured RVs using the proposed SN approach and as location parameter the median of the SN are on average 10% smaller than the uncertainties calculated on the mean of the Normal. The estimated uncertainties on the asymmetry parameter of the SN are on average 15% smaller than the uncertainties measured on the Bisector Inverse Slope Span (BIS SPAN), which is the commonly parameter used to evaluate the asymmetry of the CCF.

Conclusions. We strongly encourage the use of the SN distribution to retrieve the different parameters of the CCF. The correlation between the mean of the SN and the asymmetry parameter of the CCF and the correlation between the mean of the SN and the width parameter of the CCF probe better stellar activity signals than when using the parameters retrieved with the Normal distribution. Using the SN distribution leads to smaller uncertainties for both the median of the SN and the asymmetry of the CCF than using the commonly employed Normal density. Finally, all the parameters are available in one single fit to the CCF.

Key words. techniques: radial velocities – planetary systems – stars: activity – methods: data analysis

1. Introduction

When working with radial-velocities data (RVs), one of the main limitations to the detection of small-mass exoplanets is no longer the precision of the instruments used, but the different sources of variability induced by the stars (e.g. Feng et al. 2017; Dumusque et al. 2017; Rajpaul et al. 2015; Robertson et al. 2014). Stellar oscillations, granulation phenomena, and stellar activity can all induce apparent RV signals that are above the meter-per-second precision (e.g. Saar & Donahue 1997; Queloz et al. 2001; Desort et al. 2007; Dumusque et al. 2011; Dumusque 2016) reached by the best high-resolution spectrographs (HARPS, HARPS-N, Mayor et al. 2003; Cosentino et al. 2012). It is therefore mandatory to better understand stellar signals and to develop methods

to correct for them, if in the near future we want to detect or confirm an Earth-twin planet using the RV technique. This is even more true now that instruments like the Echelle Spectrograph for Rocky Exoplanet and Stable Spectroscopic Observations (ESPRESSO) (Pepe et al. 2014) and the EXtreme PREcision Spectrometer (EXPRES) (Fischer et al. 2016) should reach the stability to detect such signals. However, if solutions are not found to mitigate the impact of stellar activity, the detection or confirmation of potential Earth-twins will be extremely challenging and false detections could plague the field.

One of the most challenging stellar signal to characterize and to correct for is the signal induced by stellar activity. Stellar activity is responsible for creating magnetic regions on the surface of stars, and those regions change locally the temperature and the convection, which can induce spurious RVs variations (e.g. Meunier et al. 2010; Dumusque et al. 2014; Borgniet et al. 2015). In theory, it should be easy to differentiate between the pure Doppler-shift induced by a planet, which shifts the entire stellar spectrum, and stellar activity, which modifies the shape

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of spectral lines and by doing so create a spurious shift of the stellar spectrum (Saar & Donahue 1997; Hatzes 2002; Kurster et al. 2003; Lindegren & Dravins 2003; Desort et al. 2007; Lagrange et al. 2010; Meunier et al. 2010; Dumusque et al. 2014). However, on quiet GKM dwarfs, the main target for precise RVs measurements, stellar activity can induce signals of a few m s^{-1} . This corresponds physically to variations smaller than 1/100th of a pixel on the detector making the changing shape of the spectral lines challenging to detect. In order to measure such tiny variations, a common approach is to average the information of all the lines in the spectrum by cross correlating the stellar spectrum with a synthetic or an observed stellar template (Baranne et al. 1996; Pepe et al. 2002; Anglada-Escudé & Butler 2012). The result of this operation gives us the cross-correlation function (CCF). To measure the Doppler-shift between different spectra and therefore to retrieve the RVs of a star as a function of time, the variations of the CCF barycenter are calculated. The barycenter is generally estimated by fitting a Normal density to the CCF and retaining its mean. Variations in line shape between different spectra, which indicate the presence of signals induced by stellar activity, are measured by analyzing the different moments of the CCF. Usually, the width of the CCF is estimated using the full-width half-maximum (FWHM) of the fitted Normal density, and its asymmetry by calculating the CCF bisector and measuring the bisector inverse slope span (BIS SPAN, Queloz et al. 2001).

If an apparent RV signal is induced by activity, generally a strong correlation will be observed between the RV and chromospheric activity indicators like $\log(R'_{HK})$ or H- α (Boisse et al. 2009; Dumusque et al. 2012; Robertson et al. 2014), but also between the RV and the FWHM of the CCF or its BIS SPAN (Queloz et al. 2001; Boisse et al. 2009; Queloz et al. 2009; Dumusque 2016). It is therefore common now, that when fitting a Keplerian signal to a set of RVs to look for a planet, the model includes in addition linear dependancies with the $\log(R'_{HK})$, the FWHM and the BIS SPAN (Dumusque et al. 2017; Feng et al. 2017). It is also common to add a Gaussian process to the model to account for the correlated noise induced by stellar activity. The hyperparameters of the Gaussian process can be trained on different activity indicators (Haywood et al. 2014; Rajpaul et al. 2015) or directly on the RVs (Faria et al. 2016). It is therefore essential for mitigating stellar activity to obtain activity indicators that are the most correlated with the RVs but also for which we can obtain the best precision.

Several indicators have been developed that are more sensitive to line asymmetry than the BIS SPAN. In Boisse et al. (2011), the authors develop V_{span} , which is the difference between the RV measured respectively by fitting a Normal density to the upper and the bottom part of the CCF. This CCF asymmetry parameter is shown to be more sensitive than the BIS SPAN at low signal-to-noise ratio (SNR). Figueira et al. (2013) studied the use of new indicators, BIS-, BIS+, bi-Gauss and V_{asy} . The authors were able to show that when using bi-Gauss, the amplitude in asymmetry is 30% larger than when using BIS SPAN, therefore allowing the detection of lower levels of activity. They also demonstrated that V_{asy} seems to be a better indicator of line asymmetry at high SNR, as its correlation with RV is more significant than any of the previously proposed asymmetry indicators.

In all the methods described above, except bi-Gauss, the RV and the FWHM are derived using a Normal density fitted to the CCF, and the asymmetry is estimated using another approach. In this paper we propose to use a Skew Normal (SN) density to estimate with a single fit of the CCF, the RV, the FWHM and

the asymmetry of the CCF, as this function includes a skewness parameter (Azzalini 1985).

The paper is organized as follow. In Sec. 2 we introduce the SN density, describe its applicability for modeling the CCF, and study how the SN parameters relate to the RV, FWHM and BIS SPAN of the CCF. In Sec. 3 we propose an expanded linear model to correct for stellar activity signals in RVs, which extends the linear models previously proposed for this purpose (e.g. Dumusque et al. 2017; Feng et al. 2017). In Sec. 4 the performance of the SN fit to the CCF is investigated using simulations coming from the Spot Oscillation And Planet 2.0 code (SOAP 2.0, Dumusque et al. 2014), followed by an analysis of real observations in Sec. 5. Sec. 6 considers derived error bars for the different estimated CCF parameters, and finally a discussion of the results and conclusions are included in Secs. 7 and 8, respectively.

2. The Skew Normal distribution

The Skew Normal (SN) distribution is a class of probability distributions which includes the Normal distribution as a special case (Azzalini 1985). The SN distribution has, in addition to a location and a scale parameter analogous to the Normal distribution's mean and standard deviation, a third parameter which describes the skewness (i.e. the asymmetry) of the distribution. Considering a random variable $Y \in \mathbb{R}$ (where \mathbb{R} is the real line) which follows a SN distribution with location parameter $\xi \in \mathbb{R}$, scale parameter $\omega \in \mathbb{R}^+$ (i.e., the positive real line), and skewness parameter $\alpha \in \mathbb{R}$, its density at some value $y \in Y$ can be written as

$$SN(y; \xi, \omega, \alpha) = \frac{2}{\omega} \phi\left(\frac{y - \xi}{\omega}\right) \Phi\left(\frac{\alpha(y - \xi)}{\omega}\right), \quad (1)$$

where ϕ and Φ are respectively the density function and the distribution function of a standard Normal distribution¹. The skewness parameter α quantifies the asymmetry of the SN. Examples of SN densities under different skewness parameter values and the same location and scale parameters ($\xi = 0$ and $\omega = 1$) are displayed in Fig. 1. A usual Normal distribution is the special case of the SN distribution when the skewness parameter α is equal to zero². For reasons related to the interpretation of the parameters in Eq. 1 and computational issues with estimating α near 0, a different parametrization is used in this work, which is referred to as the *centered parametrization* (CP). This CP is much closer to the parametrization of a Normal distribution, as it uses a mean parameter μ , a variance parameter σ^2 and a skewness parameter γ . In order to define the CP, we need to express the CP parameters (μ, σ^2, γ) as a function of (ξ, ω^2, α). This can be done using the following relations:

$$\mu = \xi + \omega\beta, \quad \sigma^2 = \omega^2(1 - \beta^2), \quad \gamma = \frac{1}{2}(4 - \pi)\beta^3(1 - \beta^2)^{-3/2}, \quad (2)$$

where $\beta = \sqrt{\frac{2}{\pi}} \left(\frac{\alpha}{\sqrt{1 + \alpha^2}} \right)$ (e.g. Arellano & Azzalini 2010).

By using Eq. 2, the new set of parameters (μ, σ^2, γ) provides a clearer interpretation of the behavior of the SN distribution.

¹ A standard Normal distribution is a Normal distribution with a mean of 0 and a standard deviation of 1.

² This can be seen from Eq. 1. If $\alpha = 0$ then $\Phi\left(\frac{\alpha(y - \xi)}{\omega}\right) = \Phi(0) = 0.5$ and therefore $SN(y; \xi, \omega, 0) = \frac{1}{\omega} \phi\left(\frac{y - \xi}{\omega}\right)$ which is the density of a Normal distribution. Note that $\Phi(0) = 0.5$ because $\Phi(0)$ is the probability that a standard Normal random variable is less than or equal than 0.



Fig. 1. Density function of a random variable Y following the SN distribution $SN(\xi, \omega^2, \alpha)$ with location parameter $\xi = 0$, scale parameter $\omega = 1$ and different values of the skewness parameter α indicated by different colors and line types. Note that the solid black line has an $\alpha = 0$, making it a Normal distribution.

For the α values used in Fig. 1, the corresponding values of (μ, σ^2, γ) are displayed in Table 1. In particular, μ and σ^2 are the actual mean and variance of the distribution, rather than simply a location and scale parameter, and γ provides an measure of the skewness of the SN. Along with the mean of the SN, we consider the median of the distribution as a measure of its barycenter. See Table 1 for the medians of the SN densities displayed in Fig. 1.

Table 1. CP values (μ, σ^2, γ) along with the median corresponding to the α values shown in Fig. 1, with location parameter $\xi = 0$ and scale parameter $\omega = 1$. Values are rounded to three decimal places.

α	μ	σ^2	γ	Median
-3	-0.757	0.427	-0.667	-0.672
0	0.000	1.000	0.000	0.000
2	0.714	0.491	0.454	0.655
6	0.787	0.381	0.891	0.674
10	0.794	0.370	0.956	0.674

Further details about the parametrization from Eq. 1, called the *Direct Parametrization* or DP, the CP, and general statistical properties of the SN are treated in rigorous mathematical and statistical viewpoints in the book by Azzalini & Capitanio (2014).

2.1. Fitting the Skew Normal density to the CCF

To fit the CCF using a SN density shape, we use a least-squares algorithm and the following model:

$$f_{CCF}(x_i) = C - A \times SN(x_i; \mu, \sigma^2, \gamma), \quad i = 1, \dots, n \quad (3)$$

where C is an unknown offset for the continuum of the CCF, A is the unknown amplitude of the CCF, some times referred to as the CCF contrast, and μ , σ^2 and γ are the mean, variance and skewness of the SN as defined above. The values x_1, \dots, x_n are the different values of the x-axis of the CCF, generally in velocity units (e.g. m s^{-1}).

When fitting a Normal density to the CCF, the estimated mean of the model is used as the estimated RV, the FWHM of the Normal density³ represents the width of the CCF. Because

the Normal density is symmetric, the skewness is always equal to 0 so a separate approach is needed to estimate the skewness of the CCF. An estimated skewness parameter is generally obtained by calculating the BIS SPAN of the CCF (see Sect. 1, and e.g. Queloz et al. 2001).

With the proposed SN approach, we propose two estimators of the RV: the mean and median of the SN model fit (referred to as SN mean RV and SN median RV, respectively), and present advantages and limitations for both of these choices in Sec. 5 and Sec. 6. The width of the SN, SN FWHM, is defined in the same way as for the Normal density⁴, and finally the skewness of the CCF is estimated by the γ parameter.

To evaluate the strength of the correlation between the estimated RVs and the different stellar activity indicators, we calculated the Pearson correlation coefficient, R , which in its general form is defined as:

$$R(x, y) = \frac{\text{cov}(x, y)}{\sigma(x)\sigma(y)}, \quad (4)$$

where x and y are two quantitative variables, $\text{cov}(x, y)$ indicates the covariance between x and y , and $\sigma(x)$ and $\sigma(y)$ represent their standard deviations. A p -value for the statistical test having null hypothesis $H_0 : R = 0$ is provided, along with a 95% confidence interval for R when needed.

3. Radial Velocity correction for stellar activity

Exoplanets only produce a pure RV signal. On the contrary, stellar activity, in particular the presence of active regions on the stellar photosphere, do not produce blueshifts or redshifts of the entire stellar spectrum but can create spurious RV signals by modifying the shape of spectral lines. To track these variations in the shape of the spectral lines, the general approach consists in using the FWHM, the BIS SPAN or other indicators such as those introduced in Boisse et al. (2011) or Figueira et al. (2013), which provide information on the width and asymmetry of the CCF. A strong correlation between the estimated RVs and one or more of these parameters provides an indication that stellar activity signals may be affecting the measurements.

When fitting for planetary signals in RV data, it is common to include linear dependencies with the BIS SPAN and the FWHM to take into account the signal induced by stellar activity (e.g. Dumusque et al. 2017; Feng et al. 2017). We propose to add additional parameters in the model to correct for stellar activity: first the amplitude parameter A of the CCF, generally referred to as the CCF contrast, and the interaction between the BIS SPAN and the FWHM (or γ and SN FWHM in the SN case). The stellar activity correction we propose can therefore be written as:

$$RV_{\text{activity}} = \beta_0 + \beta_1 A + \beta_2 \gamma + \beta_3 \text{SN FWHM} + \beta_4 (\gamma \text{SN FWHM}) + \epsilon, \quad (5)$$

where β_0 is the intercept and ϵ is the error with mean equal to 0 and covariance matrix equal to $\sigma^2 I$ (I defined as the identity matrix). The contrast parameter A accounts for the presence of a spot on the stellar surface, which produces a change in the amplitude of the CCF and not only on its asymmetry or width (see e.g. Fig. 2 in Dumusque et al. 2014). The benefits of including a variable that quantifies the interaction between γ and SN FWHM (or BIS SPAN and FWHM) will be better understood

⁴ Note that SN FWHM does not correspond to the width of the SN density at half maximum like in the Normal case.

³ FWHM = $2\sqrt{2\ln 2}\sigma$ with standard deviation σ

through the results of the examples presented in Sec.4. This interaction term can account for possible interactions between SN FWHM (or FWHM) and γ (or BIS SPAN), meaning that each variables' association with the response, RV_{activity} , depends also on the other variable.

The proposed model is analyzed using statistical tests on the parameters $\beta_0, \beta_1, \beta_2, \beta_3$ and β_4 where the null hypothesis is $H_0 : \beta_i = 0$, for $i = 0, \dots, 4$. The significance level for the tests are set at 0.05. The coefficient of determination, R^2 , is used to assess how well the proposed linear combination of variables accounts for the variability of RV_{activity} .

The proposed function defined in Eq. 5 is the result of statistical and astronomical considerations. In particular we checked that the correlations between the proposed parameters were not approaching one: if it was the case, the matrix needed to calculate the estimates would be singular, hence non invertible. This problem is known in statistics with the term multicollinearity. A detailed discussion of the topic can be found in the book by Belsley (1991). In the analysis of real data presented in this work, we never observed a correlation coefficient exceeding 0.66 between the asymmetry and width parameters and therefore, the problem of multicollinearity is avoided. Note that we investigated the statistical significant of the interaction term between A and the width, and A and the asymmetry of the CCF, however, those interaction were not relevant for accounting for stellar signal.

4. Simulation Study

In order to evaluate the performance of the proposed SN approach for modelling the CCF and the benefit of using the proposed correction for stellar activity (See Eq. 5), we begin by considering a simulation study using spectra generated from the Spot Oscillation And Planet 2.0 code (SOAP 2.0, Dumusque et al. 2014).

For a given configuration of spots and faculae on the stellar surface, SOAP 2.0 gives as output the simulated CCF as a function of rotational phase. The code also returns the RV and the FWHM by fitting a Normal density to the CCF, and the BIS SPAN by calculating the bisector of the CCF. SOAP 2.0 gives noiseless CCFs affected by stellar activity, which are used to compare the benefits of a SN density fit to the CCF compared to a Normal density fit.

For the simulations discussed below, a star similar to the Sun was modeled, with a solar disc of one solar radius seen equator-on, and with a stellar rotational period set to 25.0 days. The stellar effective temperature is set to 5778 K, and a quadratic limb-darkening relation with linear and quadratic coefficients 0.29 and 0.34 are used, respectively (Oshagh et al. 2013; Claret & Bloemen 2011). In order to make the result of the simulations more comparable to real data obtained with the HARPS spectrograph discussed in Sect. 5, the SOAP 2.0 CCFs were generated with a width of 40 km s^{-1} and considering initial spectra with a spectral resolution of $R=115'000$.

4.1. Faculae

To see the impact of a facula on the different parameters of the CCF, we simulated the effect of an equatorial faculae of size 3% relative to the visible stellar hemisphere. The faculae is face-on when the phase equals to 0. Note that a 3% faculae is relatively large for the Sun; at maximum activity, big faculae have generally a size of 1% (e.g. Borgniet et al. 2015). In Fig. 2, we compare the barycentric variation of the CCF as measured when

fitting a Normal density and using its mean (N mean RV), and when fitting a SN density and taking its mean (SN mean RV) or its median (SN median RV). We see that all the different estimates of the CCF barycenter present a signal of similar amplitude, however the signal obtained with SN mean RV is different from the two others with a maximum amplitude happening at a different phase.

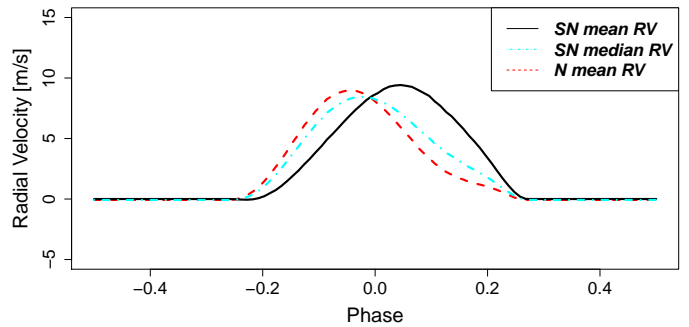


Fig. 2. RV estimates for N mean RV (red dashed line), SN mean RV (black line) or SN median RV (cyan dots-dashed line). In this case, the CCFs have been generated using SOAP 2.0, considering an equatorial 3% faculae on the simulated Sun. The star does one full rotation between phase -0.5 and 0.5, with the facula being seen face-on for phase 0. The variations observed in SN mean RV are quite different from the variations measured in SN median RV and N mean RV.

Correlations between the different RV estimates and the different CCF asymmetry or width estimates are displayed in Fig. 3. The strength of the correlation between γ and SN mean RV, and γ and SN median RV are stronger than the correlations between BIS SPAN and RV, with Pearson correlation coefficient values of $R=0.46, -0.67$ and -0.09 , respectively. For the width barycenter correlations, there is a stronger correlation between SN FWHM and SN mean RV compared to the one between FWHM and N mean RV, $R = 0.98$ and 0.84 , respectively. In this case however, the correlation between SN FWHM and SN median RV is smaller with $R = 0.50$. This first analysis shows that in the case of a facula, using some parameters from the SN can lead to stronger correlation than the usual Normal parameters and therefore, the SN parameters may better probe stellar activity. We investigate this feature further in the next sections where we consider simulated data with a single spot and a spot plus a planet, and in Sec 5 with real observations.

Since the RV variation displayed in Fig. 2 is caused by only stellar activity, in this case a facula, we applied the activity correction proposed in Eq. 5 to check its performance in this setting. The results of this correction are displayed in Fig. 4 and the statistical tests on the coefficients involved in Eq. 5 are summarized in Table 2. The proposed correction for stellar activity is able to account for the majority of the activity signal created by a facula, with a R^2 of our model larger than 0.95. In addition, the rms of the different estimates of the RV reduces from about 3 m s^{-1} before correction to values below 0.15 m s^{-1} after correction. We see a slightly smaller rms after correction when using the SN parameters compared to the Normal parameters, however the difference is probably not significant. When comparing the correction proposed in Eq. 5 with what is generally used (i.e. a linear combination of only the asymmetry and width parameter), we see that the proposed correction is able to reduce the rms of the RV residuals by a factor of 2. Looking at the signifi-



Fig. 3. *Left:* Correlations between the different asymmetry parameters and their corresponding RVs estimates in the case of an equatorial 3% faculae on the simulated Sun. *Right:* Correlations between the different width parameters and their corresponding RVs estimates for the same facula. In the presence of a facula, both the shape and the width of the CCF change as the star rotates, producing statistically significant correlations.

cance of the coefficients in table 2, we observe that all the Sn or Normal parameters are relevant for the correction.

4.2. Spot

In this section, we consider the effects on the CCF parameters of an equatorial spot of size 1% relative to the visible stellar hemisphere. The spot is face on when the phase equals to 0. Note that this is a large spot for the Sun, as in general large spots are more in the regime of 0.1% (e.g. Borgniet et al. 2015). In Fig. 5, we shows the barycentric variation of the CCF induced by this simulated spot. In contrast to the case of the facula, all the different estimates of the CCF barycenter for the spot have the same shape in variation. The amplitude for SN mean RV is however slightly smaller.

Fig. 6 shows the correlations between the asymmetry parameters and the different estimates for the CCF barycenter (i.e. SN mean RV, SN median RV and N mean RV). The correlation be-

Table 2. P-values for the different coefficients used in Eq. 5 for the correction from stellar activity induced by an equatorial 3% facula on the simulated Sun. All the parameters corresponding to the Normal or SN variables are statistically significant to explain the spurious RV variations caused by a facula. The estimated R^2 show that the proposed correction for stellar activity explains the vast majority of the spurious variability seen in the different RV estimates.

Parameter	N mean RV	SN mean RV	SN median RV
β_0	0.033	0.00020	0.61
β_1	$2.22e - 16$	$2.22e - 16$	$2.22e - 16$
β_2	0.0034	$2.22e - 16$	$2.22e - 16$
β_3	0.00016	$1.091e - 6$	$9.75e - 7$
β_4	$2.22e - 16$	$2.22e - 16$	$2.22e - 16$
R^2	0.9978	0.9985	0.9981

tween γ and SN median RV is the strongest with a $R = 0.94$,

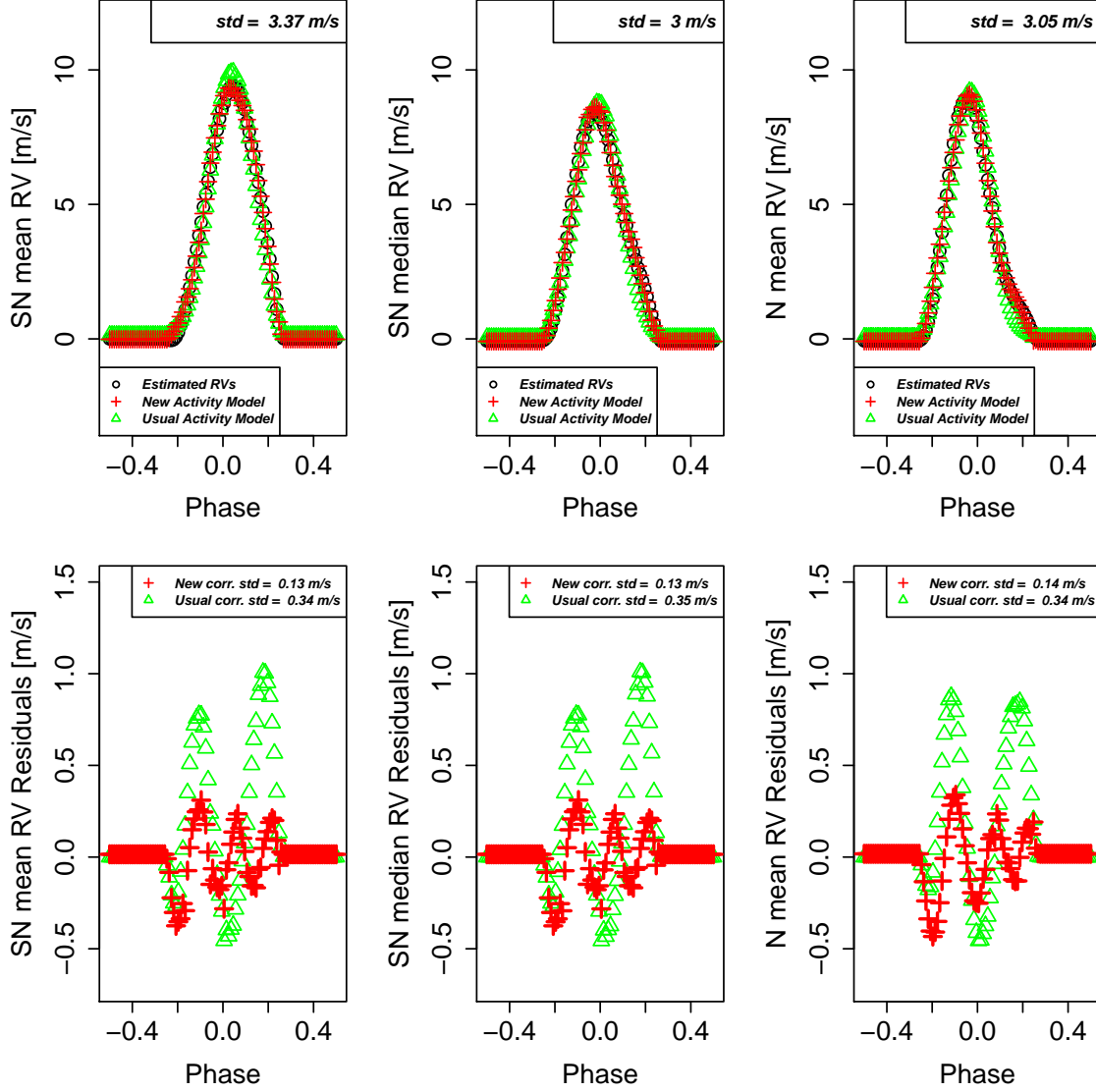


Fig. 4. (top) The spurious RVs (black dots) caused by a facula in the simulated data using a Normal and a SN fit, the estimated RVs using Eq. 5 (red crosses) and the estimated RVs using the usual correction for stellar activity (green triangles), based on $RV_{\text{activity}} = \beta_0 + \beta_1 \gamma + \beta_2 \text{SN FWHM}$ for the SN fit and on $RV_{\text{activity}} = \beta_0 + \beta_1 \text{BIS SPAN} + \beta_2 \text{FWHM}$ for the normal fit. (bottom) The residuals from the model fit using Eq. 5 (New corr. std–red crosses) and the residuals from the usual correction (Usual corr. std–green triangles). The residuals have a smaller systematic component when using the proposed model of Eq. 5 (red crosses) compared to the usual model (green triangles). The tests of statistical significance on the parameters are presented in Table 2.

followed by the correlation BIS SPAN - N mean RV and γ - SN mean RV, with $R = 0.86$. Regarding the correlation between the width and the CCF barycenter, we note that the variation is seen as a circle in this parameter space and therefore no correlation is observed. Once again, like in the case of the facula, we see that some parameters of the SN gives stronger correlations than the Normal parameters.

As before, the originally RVs were corrected by using Eq. 5. The results of the correction are displayed in Fig. 7 and in Table 3. Like for the faculae, the proposed correction is able, in the case of the spot, to almost model entirely the stellar activity signal when considering the SN or Normal parameters, with R^2 values for the linear combination larger than 0.99. Looking at Fig. 7, we see that the proposed activity correction is able to reduce the signal of a spot from a raw RV rms larger than 4.80 m s^{-1} down to a rms of 0.38 m s^{-1} , for any of the different

RV estimates. When comparing the activity correction proposed in this paper with what is commonly used, i.e only a linear dependence with the width and asymmetry of the CCF, we see that our solution is capable of reducing the RV residual rms by a factor of 3.5, which is even more than the factor 2 found in the case of the faculae.

Looking at Table 3, we see that the Normal or SN parameters appearing in Eq. 5 are all statistically significant to explain the activity signal, except the width of the CCF when the RVs are derived with the mean or median of the SN density. This is not surprising when looking at the circle shape drawn when plotting the width as a function of the RV in Fig. 6.

(top) The spurious RVs (black dots) caused by a spot in the simulated data using a Normal and a SN fit, the estimated RVs using Eq. 5 (red crosses) and the estimated RVs using the usual correction for stellar activity (green triangles), based on



Fig. 5. RV estimates for N mean RV (red dashed line), SN mean RV (black line) or SN median RV (cyan dots-dashed line). In this case, the CCFs have been generated using SOAP 2.0, considering an equatorial 1% spot on the simulated Sun. The star do one full rotation between phase -0.5 and 0.5, with the spot being seen face-on for phase 0. SN mean RV seems to have the smallest spurious variations caused by the spot.

Table 3. P-values for the different coefficients used in Eq. 5 for the correction of stellar activity induced by an equatorial 1% spot on the simulated Sun. All the parameters corresponding to the Normal or SN parameters are statistically significant to explain the spurious RV variations caused by this spot, except for the width of the CCF when using SN mean RV or SN median RV as RV estimates. The estimated R^2 show that the proposed correction for stellar activity explains the vast majority of the spurious variability seen in the different RV estimates.

Parameter	N mean RV	SN mean RV	SN median RV
β_0	0.4975	0.21	0.21
β_1	$2e-16$	$2e-16$	$2e-16$
β_2	$2e-16$	$2e-16$	$2e-16$
β_3	0.017	0.13	0.11
β_4	$2e-16$	$2e-16$	$2e-16$
R^2	0.9959	0.9936	0.9952

$RV_{\text{activity}} = \beta_0 + \beta_1\gamma + \beta_2\text{SN FWHM}$. (bottom) The residuals from the model fit using Eq. 5 (New corr. std) and the residuals from the usual correction (Usual corr. std), obtained fitting the function $RV_{\text{activity}} = \beta_0 + \beta_1\gamma + \beta_2\text{SN FWHM}$. The residuals have a smaller systematic component when using the proposed model of Eq. 5 (red crosses) compared to the usual model (green triangles). The tests of statistical significance on the parameters are presented in Table 3.

4.3. Spot and planet

The final simulation presented here includes a planetary signal influencing the CCF along with the 1% spot modeled previously (see Sec. 4.2). The purpose of this example is to check if we are able to disentangle these two different sources of variations when using the parameters derived using a Normal versus a SN fit to the CCF. In this scenario the planet is injected with a semi-amplitude of 10 m s^{-1} with no eccentricity and with a period corresponding to one third of the stellar rotational period [[Jessi: (i.e. one-third of 25 days)]] [[Jessi: The previous i.e. I added is highlighted just so you can verify that indeed the rotation period is 25 days as noted at the beginning of this section]].

Fig. 8 shows the variation observed in the CCF barycenter parameters. As in the case of the spot, all RV estimates show

Table 4. P-values for the different coefficients used in Eq. 5 for the correction of stellar activity induced by an equatorial 1% spot on the simulated Sun and a planet with period one third of the rotational period and semi-amplitude 10 m s^{-1} . All the parameters corresponding to the Normal or SN variables are statistically significant to explain the spurious RV variations caused by this spot plus planet, except for the width of the CCF. Note that since nonlinear least squares was required, the residual standard error rather than the R^2 is displayed as a reference.

Parameter	N mean RV	SN mean RV	SN median RV
β_0	0.00063	$2e-16$	$1.42e-09$
β_1	$2e-16$	$2e-16$	$2e-16$
β_2	$2e-16$	$2e-16$	$2e-16$
β_3	0.067	0.40	0.38
β_4	$2e-16$	$2e-16$	$2e-16$
K	$2e-16$	$2e-16$	$2e-16$
P	$2e-16$	$2e-16$	$2e-16$
t_0	$2e-16$	$2e-16$	$2e-16$
Residuals	0.71 m s^{-1}	0.66 m s^{-1}	0.70 m s^{-1}

similar variations, with SN mean RV showing a slightly smaller amplitude.

The correlation between the different CCF parameters are displayed in Fig. 9. The correlations are weaker than in the case of the spot due to the planet inducing changes in RV without affecting the width or asymmetry of the CCF. However, the strength of the correlations between the CCF asymmetry and RV are in the same order as with the spot-only model: γ -SN median RV has the highest correlation followed by BIS SPAN-N mean RV and then γ -SN mean RV, with R values of -0.84 , -0.78 and -0.76 , respectively. The patterns seen in the width-RV phase space in Fig. 9 follow a circle similar to the spot-only model, and no correlation is observed between those two parameters.

In order to correct the estimated RVs from the spurious variation caused by the spot, the proposed model for correcting the activity is added to a signal that takes into account the RV variation caused by a injected planet. The observed RV can therefore be modeled as a combination of the activity and planetary signals:

$$RV = RV_{\text{activity}} + RV_{\text{planet}}, \quad (6)$$

where RV_{activity} can be found in Eq. 5, and RV_{planet} , in the case with no eccentricity, can be modeled by the following sinusoidal function:

$$RV_{\text{exoplanet}} = K \sin\left(\frac{2\pi}{P}(t - t_0)\right), \quad (7)$$

with amplitude K , orbital period P , and an epoch at the periaapsis t_0 . The previous three unknown parameters define the planetary orbit.

The proposed model from Eqn (6) was fit and the results of the estimated model are summarized in Table 4. Except for the width parameters with coefficient β_3 , all the other Normal or SN parameters are significantly useful to explain the RV variation induce by a spot plus a planet. It is also apparent that the residuals, once corrected for stellar activity and the presence of the planet, lead to similar results with SN mean RV giving the smallest value. [[Jessi: I'm not sure what the point of the previous sentence is - the SN mean RV residuals is similar to what?]]



Fig. 6. *Left:* Correlations between the different asymmetry parameters and their corresponding RV estimates in the case of an equatorial 1% spot on the simulated Sun. *Right:* Correlations between the different width parameters and their corresponding RV estimates for the same spot. In the presence of a spot, both the shape and the width of the CCF change as the star rotates. However, only the asymmetry produces a statistically significant correlation with the different RV estimates. In the case of the correlation between the width parameters and their corresponding RV estimates the data produces a circle, meaning that those two quantities are not varying in phase.

5. Real data application

In this Section we present the analysis conducted on Alpha Centauri B, comparing the result of fitting a CCF using the SN density defined in Sec. 2.1 with the usual approach based on fitting a Normal density to retrieve the RV and width of the CCF and calculating the bisector to derive the asymmetry parameter BIS SPAN. Four other stars have been analyzed with the proposed method and details can be found in the Appendix A. For all the stars that we considered in the present work, we selected only the CCFs that were derived from spectra that had at least a SNR at 550 nm of 10.

5.1. Comparison between the different CCF parameters derived with the Normal and the Skew Normal density

We analyzed 1808 CCFs derived from spectra of Alpha Centauri B taken in 2010 by the HARPS spectrograph. Note that more observations have been done during this year, however we selected here only the data that are not significantly affected by contamination from Alpha Cen A (see Dumusque et al. 2012). We chosen those observations as they represent probably the best sampled and most precise RV data set showing a strong solar-like activity signal (Thompson et al. 2017; Dumusque et al. 2012).

We begin the analysis by evaluating the correlation between γ and BIS SPAN. In the left panel of Fig. 10, we see that the relationship between γ and the BIS SPAN is linear, with a slope equal to 720 and a strong Pearson correlation coefficient of $R = 0.95$. This strong correlation shows that γ and BIS SPAN are measuring in a very similar way the CCF asymmetry. This

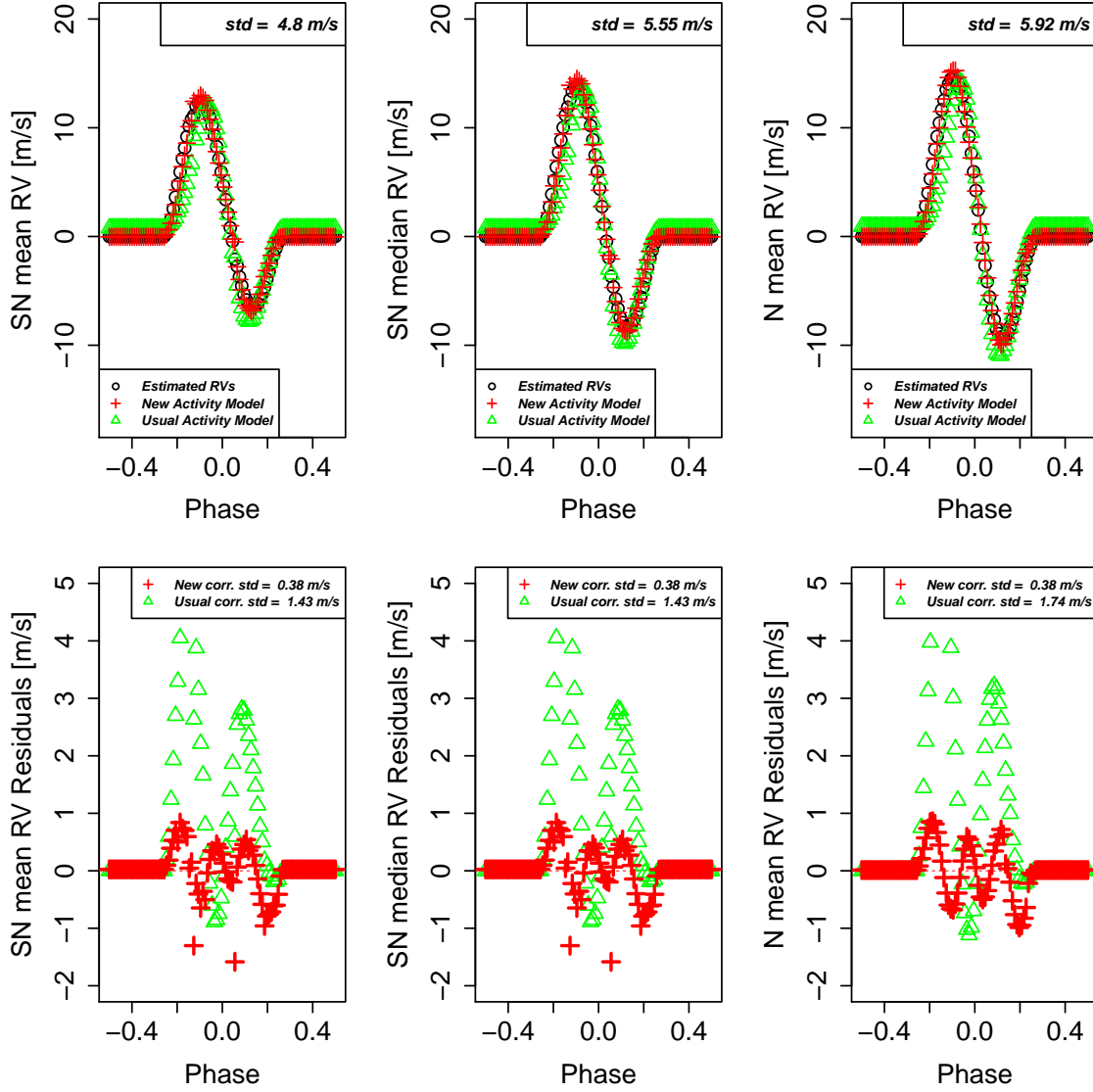


Fig. 7. (top) The spurious RVs (black dots) caused by a spot in the simulated data using a Normal and a SN fit, the estimated RVs using Eq. 5 (red crosses) and the estimated RVs using the usual correction for stellar activity (green triangles), based on $RV_{\text{activity}} = \beta_0 + \beta_1\gamma + \beta_2\text{SN FWHM}$ for the SN fit and on $RV_{\text{activity}} = \beta_0 + \beta_1\text{BIS SPAN} + \beta_2\text{FWHM}$ for the normal fit. (bottom) The residuals from the model fit using Eq. 5 (New corr. std–red crosses) and the residuals from the usual correction (Usual corr. std–green triangles). The residuals have a smaller systematic component when using the proposed model of Eq. 5 (red crosses) compared to the usual model (green triangles). The tests of statistical significance on the parameters are presented in Table 3. [\[\[Jessi: The SN Mean and Median bottom plots still show the two outliers - in the updated plots emailed out by Umberto it had looked like the outliers were at least somewhat resolved\]\]](#)

strong correlation allows as well to convert the dimensionless γ parameters into m s^{-1} using the slope of the correlation, in this case 720 m s^{-1} .

In the right of Fig. 10, we show the comparison between the RVs retrieved using the SN density and the ones obtained with the Normal density. We clearly see that the amplitude of the activity signal is stronger when using SN mean RV, and the signal measured using N mean RV or SN median RV are very similar. We observe the same behavior in the case of a faculae (see Sec. 4.1) and therefore, it seems that the activity signal seen in those data is most likely induced by faculae. This is an observation that was already made in Dumusque (2014).

Similar to what is done in Sec. 4, we compare the correlation between the asymmetry or the width parameters of the CCF and the RV in Fig. 11. In addition to what is done in Sec. 4, we

add the asymmetry parameters derived in Boisse et al. (2011), V_{span} and in Figueira et al. (2013), BIS-, BIS+, Bi Gauss and V_{asy} , as these authors found those asymmetry parameters more correlated to the RVs than BIS SPAN. It is clear in the case of Alpha Cen B, but also in the four other stars presented in Appendix A, that the correlation found between γ and SN mean RV is always the strongest. On alpha Cen B, the Pearson correlation coefficient reaches a value of $R = 0.74$, while it reaches at best $R = -0.42$ for all the other asymmetry–RV correlations not derived using the SN density fit. When looking at the width–RV correlation, once again, the correlation between SN FWHM and SN mean RV is the strongest for all the stars, except for the quietest star of all five, HD10700, for which the Pearson correlation coefficient is 0.53 for the Normal parameters and 0.42 for the SN parameters. This analysis shows that the parameters



Fig. 8. RV estimates for N mean RV (red dashed line), SN mean RV (black line) or SN median RV (cyan dots-dashed line). In this case, the CCFs have been generated using SOAP 2.0, considering an equatorial 1% spot on the simulated Sun in addition to a planet with a period of one third of the rotational period of the star and with an amplitude of 10 m s^{-1} . The star do one full rotation between phase -0.5 and 0.5, with the spot being seen face-on for phase 0.

derived when using a SN density are more sensitive the stellar activity, and therefore using those parameters can lead to the detection of stellar activity, when the Normal parameters do not detect anything. This is the case for the asymmetry-RV correlation of HD10700, HD215152, Corot-7, and the width-RV correlation for HD215152 (see Appendix A).

In Fig 12, we study the efficiency of the stellar activity correction proposed in Sec. 3. The first observation we can make is that the RV measured with SN mean RV presents a rms 35% larger than the RV measured when fitting a Normal density. In the case of SN median RV, we see as well a larger RV rms, however only 9% larger. Even though we see those differences in RV, once we correct for the stellar activity using Eq. 5, we find the same RV residuals rms. Therefore, although it seems that the parameters derived using the SN density are more sensitive to stellar activity, we are not able to correct better for stellar activity signal using a linear combination of the different CCF parameters.

When looking at the significance of the different parameters in the activity correction in Table 5, we see that the intercept is not significant for any of the correction, and BIS SPAN, with coefficient β_2 , is not useful to explain stellar activity when using the parameters derived from the Normal density fit. However all the other parameters in the Normal and SN cases are useful to model stellar activity. By analyzing the values of R^2 , we see that the correction is more efficient at explaining stellar activity when considering SN mean RV as the RV estimate. This is expected from the results above since for the three different RV estimates, we arrive to the same RV residual rms after correction for activity, but before correction, SN mean RV shows the largest RV rms.

5.2. Detection limits when using the RV derived with the Normal or the Skew Normal density

In the preceding section, we saw that the RV measured when considering a SN or a Normal density present different amplitudes, mainly in the case of SN mean RV. However, when correcting for stellar activity using the linear combination presented in Eq. 5, it seems that we are able to reach the same RV precision. In this section, we want to investigate if for some of the three different RV definitions presented in this paper, it is eas-

Table 5. P-values for the different coefficients used in Eq. 5 for the correction from stellar activity in Alpha Cen B data. All the variables corresponding to the Normal or SN parameters are statistically significant to explain the spurious RV variations caused by stellar activity, except for the asymmetry of the CCF when using the Normal parameters (BIS SPAN). The evaluation of the R^2 shows that the proposed linear combination better explains variations in RVs due to the stellar activity when the RVs are derived using SN mean RV.

Parameter	RV	SN mean RV	SN median RV
β_0	0.49	0.90	0.027
β_1	$2.22e - 16$	$2.22e - 16$	$2.22e - 16$
β_2	0.33	$2.22e - 16$	$1.23e - 11$
β_3	$2.22e - 16$	$2.22e - 16$	$2.22e - 16$
β_4	$2.22e - 16$	$2.22e - 16$	$2.22e - 16$
R^2	0.57	0.78	0.66

ier or more difficult to detect planetary signals when considering stellar activity. To do so, we estimated what would be the minimum amplitude of a planetray signal we could detect at different orbital periods when considering data affected by stellar activity.

To get CCFs affected by realistic stellar activity signals, we took the ones from Alpha Cen B used in the preceding section. To simulate a planetary signal, we then blue- or red-shifted those CCFs with the wanted amplitude, period and phase. We first simulated several RV data sets, with always the same stellar signal, but injecting planets with parameters corresponding to the following grid:

- period of 3, 5, 7, 9, 11, 15, 20, 25 and 30 days,
- amplitude from 0.5 to 3 m s^{-1} by steps of 0.05 m s^{-1} ,
- 10 different phases, evenly sampled between 0 and 2π .

For each simulation, 4500 in total, we applied the activity correction presented in Eq. 5 and look for signals in the residuals using a Generalized Lomb-Scargle periodogram (Zechmeister & Kürster 2009; Scargle 1982; Lomb 1976). If a signal with a p-value⁵ smaller than 1% had a period compatible with the injected planetary period considering an error of 20%, we were estimating this signal as significant, and therefore the corresponding planet as detected. Finally, for each period, we searched for the minimum amplitude for which at least 50% of the planets with different phases were detected. The results are shown in Fig. 13. As we can see, all the different RV estimates gives extremely similar detection limits. Therefore, we can use any of these estimates when searching for planetary signal in RV data contaminated by stellar activity.

Although the detection limits stay constant for all the periods below 15 days, we see that for 20 and 30 days, we have a huge increase in detection limits for all the RV estimates. This is because the periods of the simulated planets, 20 days and 30 days, are close to the first harmonic or to the rotational period of the star (36.7 days, Dumusque et al. 2012) and therefore close to the semi-periodicity of the stellar activity signal. In such a case, the activity correction absorbs part of the planetary signal. We thus need to inject a much stronger planetary signal at these periods to detect them.

6. Estimation of standard errors for the CCF parameters

In this section, we investigate how the noise of the CCF influences the CCF parameters derived either by a Normal density

⁵ estimated by bootstrap



Fig. 9. Evaluation of the correlation between the RVs and the asymmetry parameters when a spot is present on the photosphere of the star and a planet is injected. The shape of the CCF changes as the spot moves, producing statistically significant correlations only between the RVs and the asymmetry parameter. The correlations between the RVs and the width parameter of the CCF is weaker than the previous case that considers only the presence of a spot on the photosphere of the star.

or a SN density fit. Because a CCF is obtained from a cross-correlation, each point in a CCF is correlated with each other. Therefore, we cannot simply vary each point in the CCF with their respective error bar and then recalculate the best SN or Normal density fit, to see how the CCF noise influences N mean RV, SN mean RV, SN median RV, FWHM, SN FWHM, BIS SPAN and γ . We therefore started from the spectrum, for which we know that each point is totally independent of the others. The standard error on each point of the spectrum is given by the photon noise, which follows a Poisson law and is therefore obtained by taking the square root of the measured flux.

We applied the following method to estimate the error bars on the different parameters derived from the CCF. We first modify the values of all the points in the spectrum given their respective error bars. To do so we added to each point the value randomly drawn from a Gaussian distribution centered on the value of the point and with standard deviation the square root of

the flux. We then calculated the CCF from this spectrum using the method presented in Pepe et al. (2002), fitted either a Normal or SN density to this CCF and recorded the different parameters. We redo this process a hundred times, which gives us at the end, a distribution for each CCF parameter. The standard deviations of the obtained distributions is then associated to the standard errors of each CCF parameters.

We measured the standard errors of each CCF parameters for all the CCFs of HD215152, HD192310 and Corot-7. This gives us an information of how the noise on each CCF parameter varies as a function of SNR. In this case, the SNR, measured at 550 nm on the original spectra, varies between 10 and 500. The results can be seen in Fig. 14. In the top plots of this figure, we show the standard errors for the different RV, width and asymmetry estimates. Note that because BIS SPAN and γ does not have the same units, we used the slope of the correlation between those two parameters to transform γ in m s^{-1} (see Fig. 10 and Table A.1

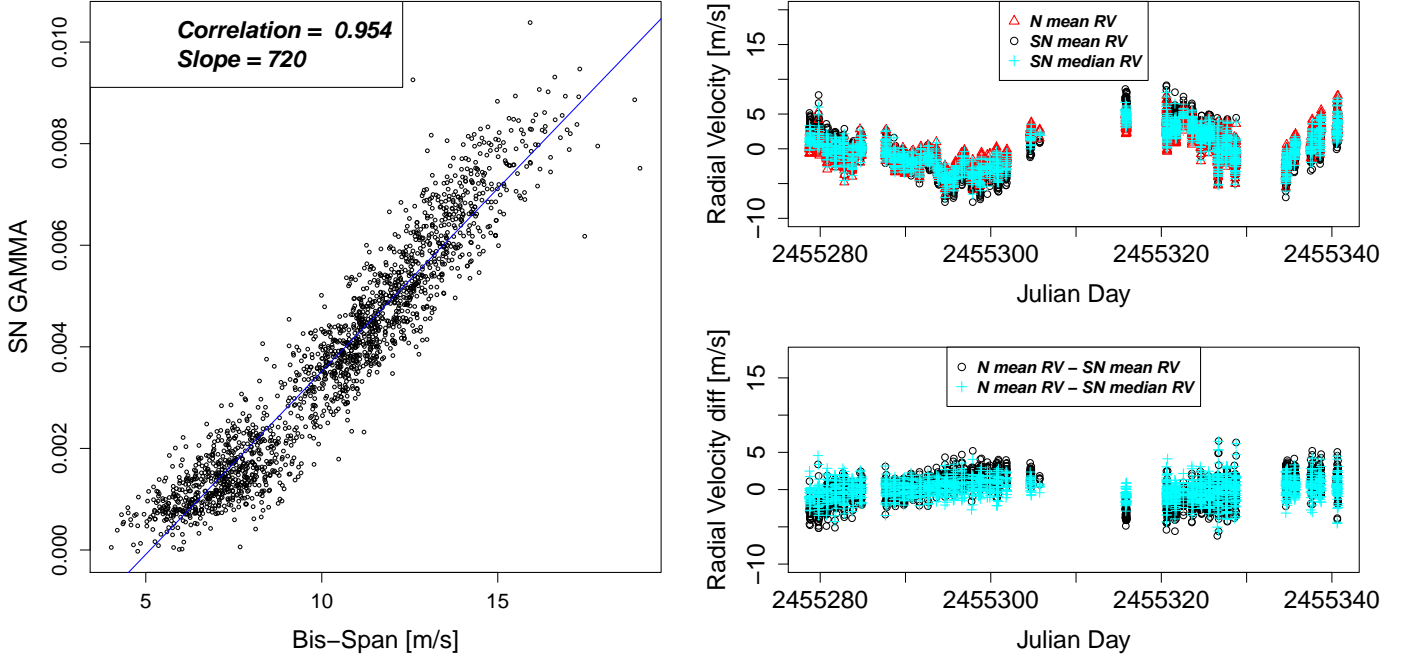


Fig. 10. *Left:* Correlation between γ and the BIS SPAN for Alpha Centauri B. As we can see with the strong correlation, those two parameters measure the CCF asymmetry in the same way. *Top right:* RVs as function of Julian Day for Alpha Centauri B in 2010. The RVs are retrieved using the mean of a Normal fitted to the CCF (red triangles), or the mean (black circles) or median (cyan crosses) of a SN density fitted to the CCF. *Bottom right:* RV differences between the RVs derived when using the SN density and the RV derived when using the Normal density.

for the value of the slope for each star). In the bottom plots, we show the ratio between the standard errors measured when using the SN parameters and the Normal parameters. Values smaller or larger than one will imply that standard errors from the SN parameters are more or less precise, respectively.

Looking at the estimates for the RV, we see that they all follow a similar exponential decay. Although we plot the data for three different stars, we do not see any offsets in this decay, which implies that the SNR at 550 nm is the main contributor to the precision measured in RV. This is not surprising as the three stars studied here are all main sequence K-dwarfs.

When comparing the three different estimates for the RV, we see that SN mean RV presents standard errors that are 60% larger than what N mean RV gives. On the opposite, SN median RV gives errors 10% more precise than N mean RV. Regarding the parameters describing the width of the CCF, FWHM and SN FWHM have the same standard errors. Finally, for the asymmetry parameters, we see that γ , derived from the SN, is 15% more precise than BIS SPAN. In conclusion, when fitting a SN density to the CCF and using as RV estimates SN mean RV, we are able to improve by 10% the precision on the RV, 15% the precision on the asymmetry, and this without degrading the precision on the width estimate. SN mean RV should not be used to derive precise RVs, except perhaps in specific conditions described below, as the precision on this parameter is 60% worse than the precision on the RVs derived when using the mean of a Normal density fit to the CCF.

7. Discussion

When fitting a SN density to the CCF, it is possible at once to retrieve parameters that describe the RV, i.e. the CCF barycenter, the amplitude, sometimes called the CCF contrast, the width,

and the asymmetry of the CCF. To estimate the RV, we investigated the use of the mean and the median of the SN density. Then the width is derived using the variance of the SN density ($\text{SN FWHM} = 2\sqrt{2\ln(2)\sigma^2}$) and the asymmetry by using γ , the skewness parameter of the SN density.

When we tested the SN density fit to the CCFs provided by the SOAP 2.0 code, that simulate the activity signal induced by a spot or a faculae on a solar-like star, we arrived to the conclusion that the parameters derived from the SN density fit are more sensitive to activity than the parameters obtained by the usual method. By more sensitive, we mean here that the correlation between the skewness and either the mean or median of the SN density is stronger than between BIS SPAN and the mean of a Normal density. Moreover, the correlation between the FWHM and the CCF barycenter is also stronger when using the parameters from the SN compared to the parameters from the Normal. We also confirmed the behavior that the SN parameters are more correlated than the Normal parameters in the case where a planetary signal is added to the stellar activity modeled by SOAP 2.0.

We also present in this paper a new function to correct for stellar activity signals. Generally, when fitting for planetary signals, people use a model composed of one or several Keplerians to account for the planets, in addition to a linear combination with the FWHM and BIS SPAN of the CCF to account for stellar activity signals. The new function proposes to add a linear dependency with the amplitude of the CCF retrieved either by the SN or Normal density fit, and an interaction term between the asymmetry and the width parameters. On the simulated data from SOAP 2.0, this new model allows to mitigate stellar activity signals by a factor of 2 and 3.5 better compared to the model generally used, for the faculae and spot, respectively.

When moving to real data, we arrived to the same conclusions than the results of the SOAP 2.0 simulations regarding the strength of the correlation between the CCF asymmetry and the RV, as well as between the CCF width and the RV. We note however that although in the case of the simulated data, the stronger asymmetry-RV correlation was found between γ and the median of the SN, in the case of real data, it was found between γ and the mean of the SN. This can be due to the fact that in SOAP 2.0, the faculae is not modelled properly, as it uses as emerging spectrum the one from a spot. We know that the temperature between a spot and a faculae is significantly different, therefore their spectra should be different. In addition, there is always several active regions on a star, at different locations in longitude and latitude, which is not the case in our SOAP 2.0 simulations were only one active region was modelled at a time. This could also be responsible for the fact that we do not observe the same behavior for the mean and the median of the SN when we compare the SOAP

When comparing the different RV estimates obtained from the mean of the Normal density, the mean or the median of the SN density, we see that in the case of real data, the amplitude of stellar activity is larger for the mean of the SN, followed by the median of the SN and the mean of the Normal. Although there is a huge gap between the first and the second estimate, the difference between the second and the third is much smaller. The mean of the SN density is therefore more sensitive to variation in the CCF shape than the median of the SN or the mean of the Normal, which is not surprising as the SN density is able to



Fig. 12. (top) The RVs (black dots) for Alpha Centauri B estimated using a SN and a Normal fit. (bottom) The residuals from the model fit using Eq. 5 (New corr. std–black dots) and the residuals from the usual correction (Usual corr. std–blue triangles), based on $RV_{\text{activity}} = \beta_0 + \beta_1\gamma + \beta_2\text{SN FWHM}$ for the SN fit and on $RV_{\text{activity}} = \beta_0 + \beta_1\text{BIS SPAN} + \beta_2\text{FWHM}$ for the normal fit. The residuals have a smaller systematic component when using the proposed model of Eq. 5 (black dots) compared to the usual model (blue triangles).

model the asymmetry of the CCF, what the Normal density does not.

Having an estimate of the RV that is more sensitive to stellar activity, like the mean of the SN, can help us to probe better stellar rotational periods, or understand better the covariance of stellar signals when fitting a Gaussian Process to the RVs (e.g. Faria et al. 2016; Haywood et al. 2014). We saw in the preceding section that the RV estimated derived by the mean of the SN is 60% noisier than the RV derived by the mean of the Normal. This is not necessarily a negative aspect. Indeed, in the case where photon-noise is not dominant, stellar activity can be better characterized if its effect can be amplified by measuring SN mean RV, even if the white noise in the data is increased by 60%.

Even if the different RV estimates have different amplitudes, once we apply our new correction for stellar activity, we arrive exactly to the same RV rms in the residuals, meaning that for an activity correction based on a linear combination of CCF pa-

rameters, the SN density fit does not help. We note that when comparing the activity correction presented in this paper with the usual correction that only uses a linear combination of the CCF asymmetry and width, we only see a slight improvement brought by our proposed correction.

We checked if some RV estimates were better at finding planets in RV data affected by observed stellar signals. When using the new correction proposed in this paper to mitigate the effect from stellar activity, and when assessing planetary detection based on aperiodogram analysis, we find very similar detection limits for the three RV estimates presented in this paper. Therefore it seems that any of the RV estimate can be used to search for planetary signals.

Finally, we investigated the precision of each of the SN and Normal parameters including the BIS SPAN. It turns out that the mean of the SN should not be used to get precise RVs as the standard errors on this parameter is 60% greater than when



Fig. 13. Detection limits of planetary signals once we have removed from the raw RVs the linear combination of CCF parameters proposed in Eq. 5 to correct for stellar activity. As we can see, when considering the different parametrization for the RV: N mean RV, SN mean RV and SN median RV, we arrive to very similar detection limits.

taking the mean of the Normal distribution. However, the median of the SN should be used, as it is 10% more precise than the mean of the Normal. Regarding the asymmetry estimates, we observe that γ has a precision 15% better than the BIS SPAN. It is therefore clear in this case that the SN parameters are more precise to measure the RV and the asymmetry of the CCF.

8. Conclusion

When searching for small-mass exoplanets using the RV technique, it is crucial to get the best precision possible on the determination of the RV, but also to measure precisely the shape variations of the CCF, as such variations are induced by stellar activity, and not by planets. The correlations between the width of the CCF and the RV, and the asymmetry of the CCF and the RV are used as a way to assess if the measurements are affected by stellar activity signals, therefore, the stronger those correlations are, the better we can probe low level of stellar activity.

In this paper, we introduced a novel approach based on the Skew Normal (SN) density to retrieve the RV and shape variation of the CCF. This new method is compared to the usual one, where a Normal density is fitted to the CCF to retrieve the information on the RV and the width of the CCF, and a bisector analysis allows to retrieve the asymmetry of the CCF using the BIS SPAN. In addition, we also compare this novel approach to other parameterisations of the CCF asymmetry that have been shown to be more sensitive to activity signals (Boisse et al. 2011; Figueira et al. 2013).

In all the different tests we could conduct in this paper, the SN parameters were doing as well as the parameters derived when fitting a Normal distribution and calculating the BIS SPAN. Moreover, in some cases, we were getting better results when using the SN parameters. This is the case when looking at the correlations between the CCF asymmetry and the RV or the CCF width and the RV. The SN parameters γ , SN FWHM and the mean of the SN density always show stronger correlation than any of the parameters derived by the Normal density fit and the BIS SPAN, or the different asymmetry parametrization presented in Boisse et al. (2011) and Figueira et al. (2013).

Those SN parameters can therefore probe better stellar activity signals. In addition, the precision measured on the median of the SN density and the asymmetry of the SN density are 10 and 15% more precise than the mean of the Normal and BIS SPAN, respectively.

Because of these advantages of using the SN parameters compared to the Normal parameters, BIS SPAN, and all the asymmetry parameters described in Boisse et al. (2011) and Figueira et al. (2013), we strongly encourage the community of using this new parametrization of the CCF. Correlations between the asymmetry and the mean of the SN, and between the width and the mean of the SN should be used to probe better stellar activity signals in RV data, and the median of the SN should be used as the new definition for RV.

9. Acknowledgements

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Appendix A: Appendix

In this Appendix we present the same analysis as presented in Sec. 5 for those four main-sequence stars: HD192310 (K2V), HD10700 (G8V), HD215152 (K3V) and finally Corot-7 (K0V).

Table A.1 summarizes the results obtained by the SN fit and some of the results based on the Normal fit. The results are all consistent with the conclusions derived by the analyses on Alpha Centauri B. The correlation between γ and SN mean RV is stronger than the correlation between the BIS SPAN and RV or between the asymmetry parameters derived in Boisse et al. (2009) and Figueira et al. (2013) and RV for all the considered stars. The correlation between SN FWHM and SN mean RV is stronger than the correlation between FWHM and RV for three of the four stars. Also for all these stars we corrected the originally estimated RVs from spurious variations in RVs caused by stellar activity, using Eq. 5. Fig. A.2–A.6 show the resulting corrected RVs. While the Normal and SN residuals, once corrected for stellar activity, are comparable for the stars HD192310 and HD10700, the results of the analyses on HD215152 and Corot-7, whose measurements have lower SNR with respect to the previous two analyzed stars, suggest that the residuals for the Normal are 0.05 and 0.34 m s^{-1} higher than the residuals retrieved with the SN analysis, respectively.

References

- Anglada-Escudé, G. & Butler, R. P. 2012, *The Astrophysical Journal Supplement Series*, 200, 15
- Arellano, R. B. & Azzalini, A. 2010
- Azzalini, A. 1985, *Scandinavian journal of statistics*, 171
- Azzalini, A. & Capitanio, A. 2014, *The skew-normal and related families*. Institute of Mathematical Statistics Monographs
- Baranne, A., Queloz, D., Mayor, M., et al. 1996, *Astronomy and Astrophysics Supplement Series*, 119, 373
- Belsley, D. A. 1991, *Conditioning diagnostics: Collinearity and weak data in regression* No. 519.536 B452 (Wiley New York)
- Boisse, I., Bouchy, F., Hébrard, G., et al. 2011, *Astronomy & Astrophysics*, 528, A4
- Boisse, I., Moutou, C., Vidal-Madjar, A., et al. 2009, *Astronomy & Astrophysics*, 495, 959
- Borgniet, S., Meunier, N., & Lagrange, A.-M. 2015, *A&A*, 581, A133

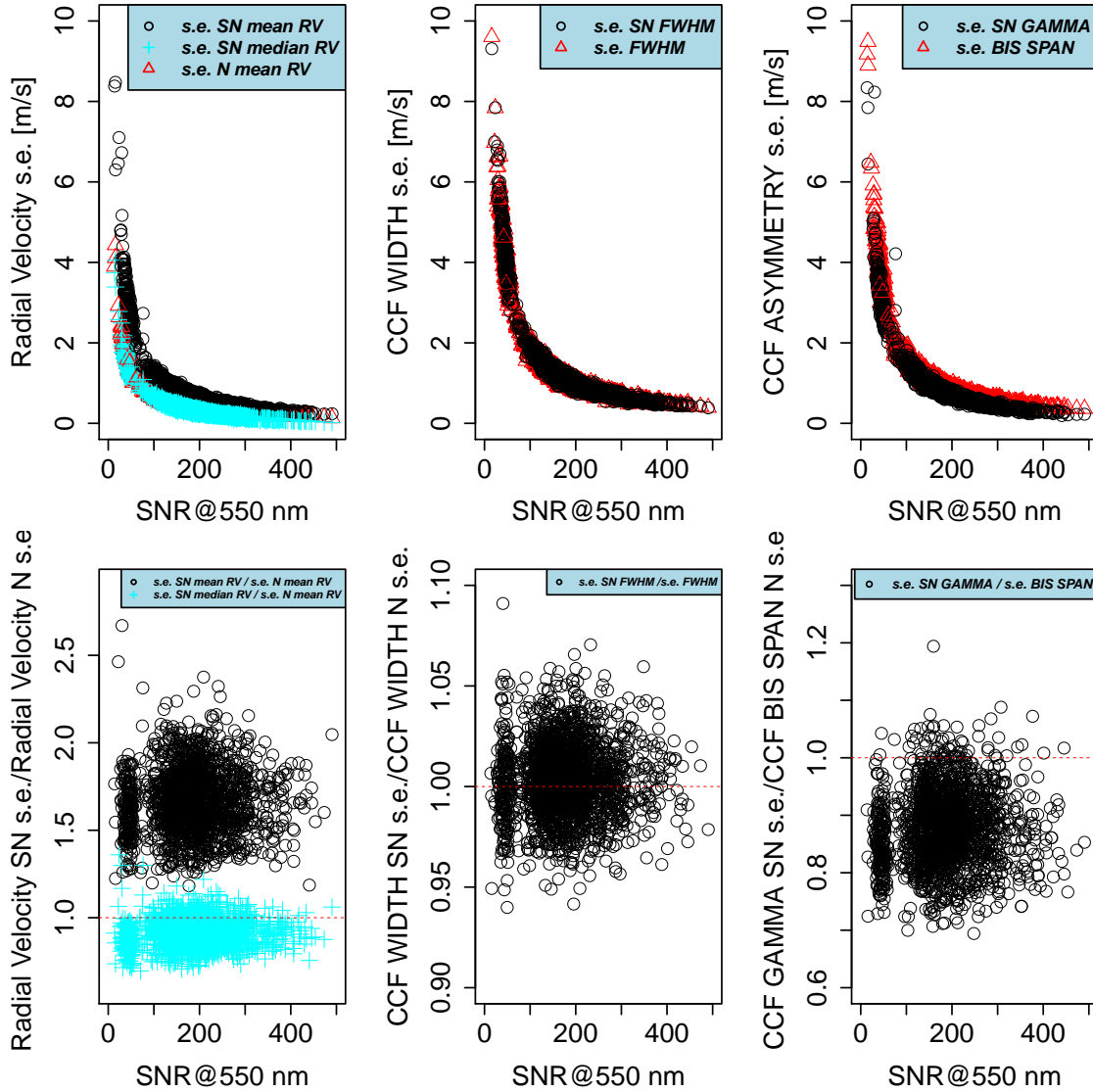


Fig. 14. Comparison between the standard errors using the bootstrap analysis for the RVs, the FWHM and the asymmetry parameters. When using SN mean RV (black circles), the standard errors are in average 60% larger than the standard errors retrieved when fitting a Normal density (red triangles). However, if using SN median RV (cyan crosses), the standard errors are in average 10% smaller than the standard errors coming from the Normal fit. The use of the asymmetry SN parameter γ leads to standard errors in average 15% smaller than the standard errors related to the BIS SPAN. Note that for the asymmetry, the error in BIS SPAN is in m s^{-1} . To be able to compare the errors in γ and BIS SPAN, we multiplied the error in γ by the slope of the correlation between γ and BIS SPAN.

Claret, A. & Bloemen, S. 2011, *A&A*, 529, A75
 Cosentino, R., Lovis, C., Pepe, F., et al. 2012, in *Proc. SPIE*, Vol. 8446, 84461V
 Desort, M., Lagrange, A.-M., Galland, F., Udry, S., & Mayor, M. 2007, *Astronomy & Astrophysics*, 473, 983
 Dumusque, X. 2014, *ApJ*, 796, 133
 Dumusque, X. 2016, *Astronomy & Astrophysics*, 593, A5
 Dumusque, X., Boisse, I., & Santos, N. 2014, *The Astrophysical Journal*, 796, 132
 Dumusque, X., Borsa, F., Damasso, M., et al. 2017, *Astronomy & Astrophysics*, 598, A133
 Dumusque, X., Pepe, F., Lovis, C., et al. 2012, *Nature*, 491, 207
 Dumusque, X., Udry, S., Lovis, C., Santos, N. C., & Monteiro, M. 2011, *Astronomy & Astrophysics*, 525, A140
 Faria, J., Haywood, R., Brewer, B., et al. 2016, *Astronomy & Astrophysics*, 588, A31
 Feng, F., Tuomi, M., & Jones, H. R. 2017, *Astronomy & Astrophysics*, 605, A103
 Figueira, P., Santos, N., Pepe, F., Lovis, C., & Nardetto, N. 2013, *Astronomy & Astrophysics*, 557, A93

Fischer, D. A., Anglada-Escude, G., Arriagada, P., et al. 2016, *Publications of the Astronomical Society of the Pacific*, 128, 066001
 Hatzes, A. P. 2002, *Astronomische Nachrichten*, 323, 392
 Haywood, R., Collier Cameron, A., Queloz, D., et al. 2014, *Monthly notices of the royal astronomical society*, 443, 2517
 Kurster, M., Endl, M., Rouesnel, F., et al. 2003, *ASTRONOMY AND ASTROPHYSICS-BERLIN*, 403, 1077
 Lagrange, A.-M., Desort, M., & Meunier, N. 2010, *Astronomy & Astrophysics*, 512, A38
 Lindegren, L. & Dravins, D. 2003, *Astronomy & Astrophysics*, 401, 1185
 Lomb, N. R. 1976, *Astrophysics and Space Science*, 39, 447
 Mayor, M., Pepe, F., Queloz, D., et al. 2003, *The Messenger*, 114, 20
 Meunier, N., Desort, M., & Lagrange, A.-M. 2010, *Astronomy & Astrophysics*, 512, A39
 Oshagh, M., Boisse, I., Boué, G., et al. 2013, *A&A*, 549, A35
 Pepe, F., Mayor, M., Galland, F., et al. 2002, *Astronomy & Astrophysics*, 388, 632
 Pepe, F., Molaro, P., Cristiani, S., et al. 2014, *Astronomische Nachrichten*, 335, 8

Table A.1. Subset of notable correlations between the asymmetry or the FWHM parameters and the RVs for four stars: HD192310, HD10700, HD215152 and Corot 7. The complete results of the analyses of the correlations for the four stars are presented in Fig. A.1–A.7.

Star	# CCFs	R(SN γ , Bis-Span)	slope(SN γ , Bis-Span)	R(SN γ , SN mean RV)
HD192310	1577	0.888	786	0.669(0.64; 0.695)
HD10700	7928	0.78	604	0.322(0.302; 0.342)
HD215152	273	0.763	794	0.571(0.485; 0.646)
Corot 7	173	0.814	607	0.561(0.450; 0.656)

Star	R(Bis-Span, N mean RV)	R(FIG BiGaussian, N mean RV)	R(SN FWHM, SN mean RV)	R(FWHM, N mean RV)
HD192310	0.329(0.285; 0.373)	−0.333(−0.376; −0.289)	0.666(0.637; 0.692)	0.476(0.4367; 0.514)
HD10700	−0.073(−0.095; −0.0051)	0.127(0.105; 0.148)	0.421(0.403; 0.439)	0.529(0.513; 0.545)
HD215152	−0.067(−0.184; 0.052)	0.269(0.155; 0.376)	0.210(0.094; 0.321)	−0.138(−0.253; −0.020)
Corot 7	0.092(−0.058; 0.238)	−0.335(−0.228; −0.082)	−0.709(0.626; 0.776)	0.595(0.489; 0.683)

Queloz, D., Bouchy, F., Moutou, C., et al. 2009, *Astronomy & Astrophysics*, 506, 303

Queloz, D., Henry, G., Sivan, J., et al. 2001, *Astronomy & Astrophysics*, 379, 279

Rajpaul, V., Aigrain, S., Osborne, M. A., Reece, S., & Roberts, S. 2015, *Monthly Notices of the Royal Astronomical Society*, 452, 2269

Robertson, P., Mahadevan, S., Endl, M., & Roy, A. 2014, *Science*, 1253253

Saar, S. H. & Donahue, R. A. 1997, *The Astrophysical Journal*, 485, 319

Scargle, J. D. 1982, *ApJ*, 263, 835

Thompson, A., Watson, C., de Mooij, E., & Jess, D. 2017, *Monthly Notices of the Royal Astronomical Society: Letters*, 468, L16

Zechmeister, M. & Kürster, M. 2009, *A&A*, 496, 577

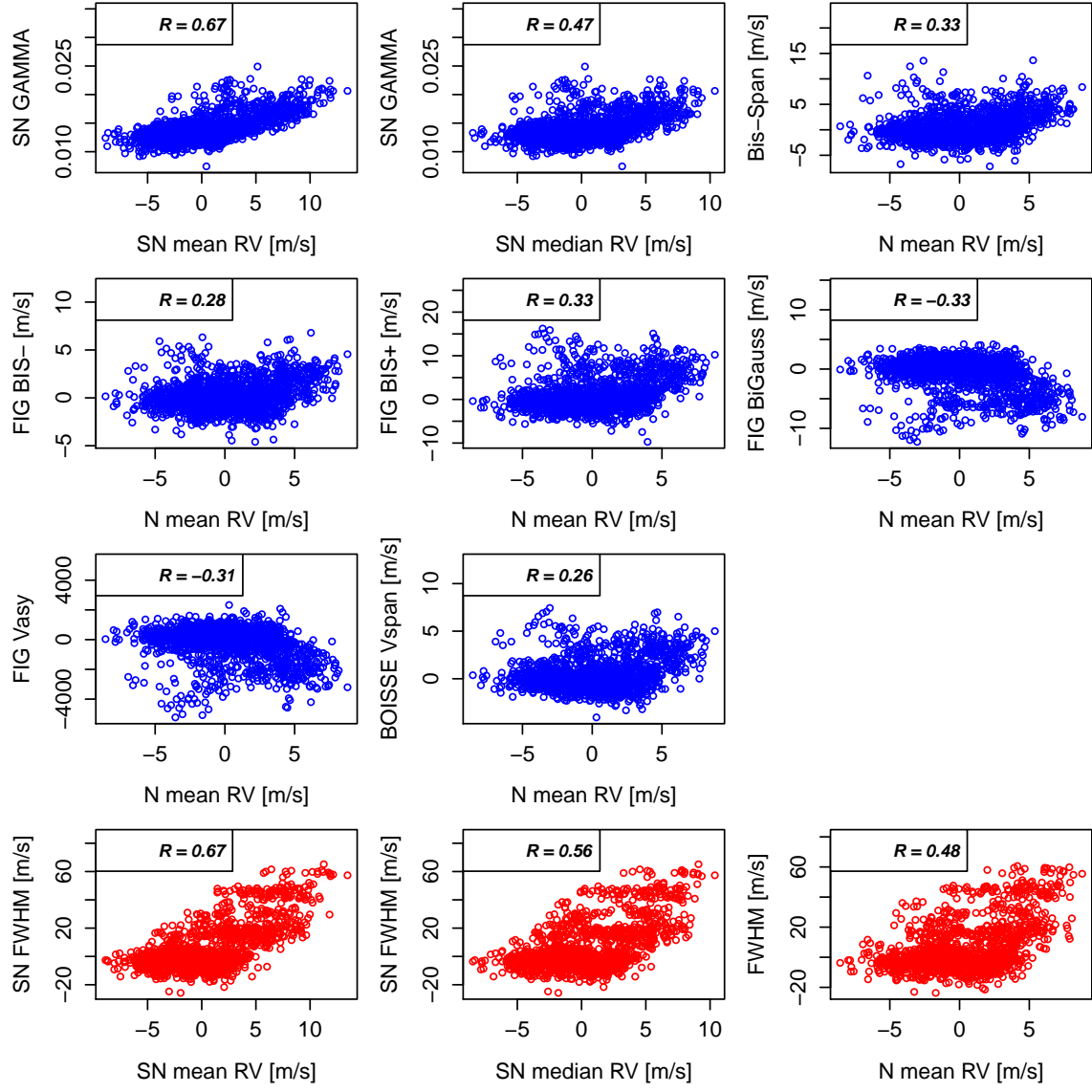


Fig. A.1. Correlations between the asymmetry parameters and their corresponding RVs for HD192310. The last three plots show the correlation between the FWHM and the RVs for HD192310 using respectively the SN (SN mean RV and SN median RV) and the Normal fits. The correlations are always stronger when using parameters derived from the SN fit than the Normal one. The p-values associated with each R are statistically different from 0.



Fig. A.2. (top) The RVs (black dots) for HD192310 estimated using a SN and a Normal fit. (bottom) The residuals from the model fit using Eq. 5 (New corr. std—black dots) and the residuals from the usual correction (Usual corr. std—blue triangles), based on $RV_{\text{activity}} = \beta_0 + \beta_1 \gamma + \beta_2 \text{SN FWHM}$ for the SN fit and on $RV_{\text{activity}} = \beta_0 + \beta_1 \text{BIS SPAN} + \beta_2 \text{FWHM}$ for the normal fit. The residuals for both the proposed correction from stellar activity are comparable.



Fig. A.3. Correlations between the asymmetry parameters and their corresponding RVs for HD10700. The last three plots show the correlation between the FWHM and the RVs for HD10700 using respectively the SN (SN mean RV and SN median RV) and the Normal fits. The correlations between the asymmetry parameters and their corresponding RVs are stronger when using the parameter SN mean RV obtained from the SN fit than the Normal one. The correlation between the FWHM and the RVs is stronger when using the parameter retrieved with the Normal distribution. The p-values associated with each R are statistically different from 0, except for the correlation between FIG BIS- and RV (p-values=0.36).



Fig. A.4. (top) The RVs (black dots) for HD10700 estimated using a SN and a Normal fit. (bottom) The residuals from the model fit using Eq. 5 (New corr. std–black dots) and the residuals from the usual correction (Usual corr. std–blue triangles), based on $RV_{\text{activity}} = \beta_0 + \beta_1 \gamma + \beta_2 \text{SN FWHM}$ for the SN fit and on $RV_{\text{activity}} = \beta_0 + \beta_1 \text{BIS SPAN} + \beta_2 \text{FWHM}$ for the normal fit. The residuals for both the proposed correction from stellar activity are comparable.

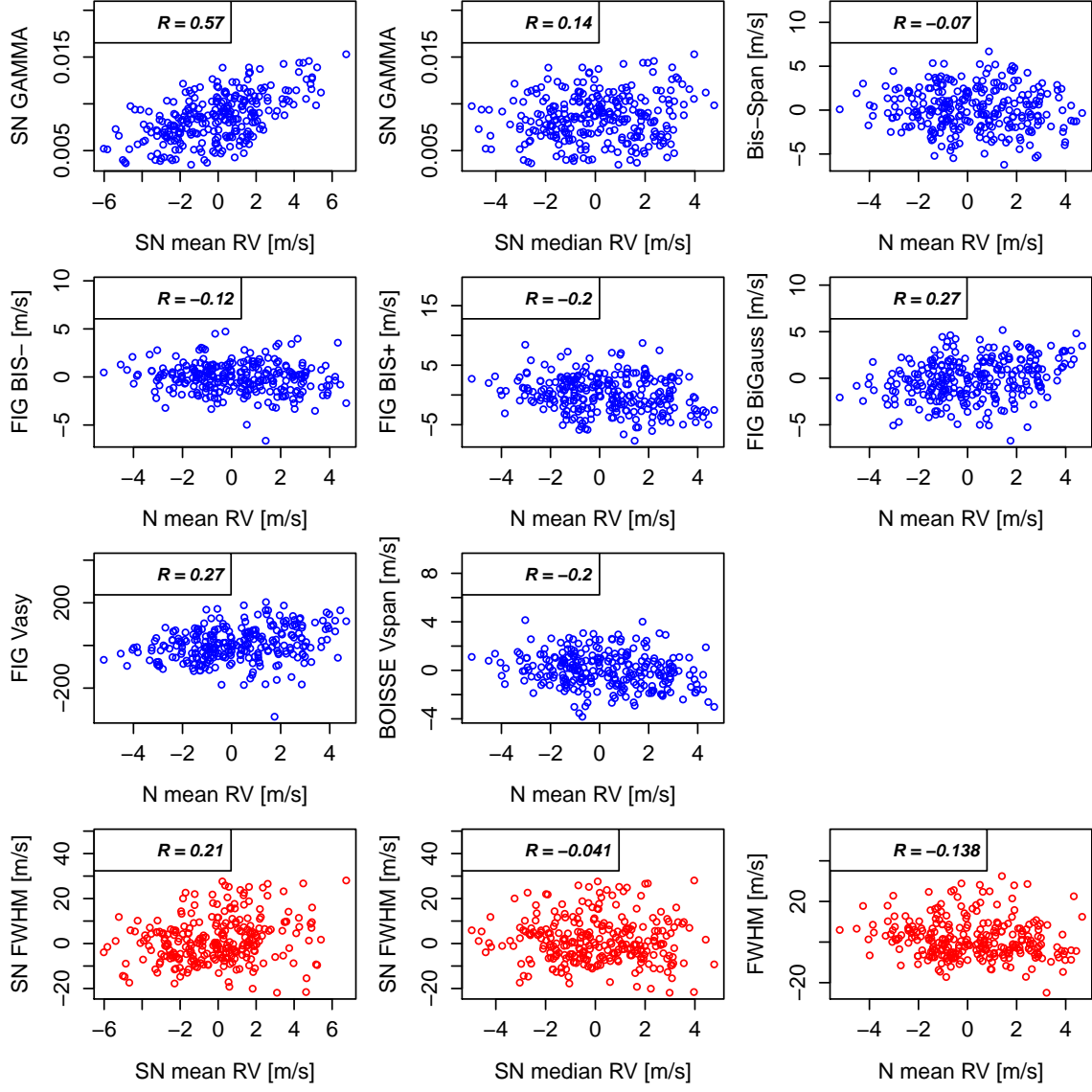


Fig. A.5. Correlations between the asymmetry parameters and their corresponding RVs for HD215152. The last three plots show the correlation between the FWHM and the RVs for HD215152 using respectively the SN (SN mean RV and SN median RV) and the Normal fits. The correlations between the asymmetry parameters and their corresponding RVs are stronger when using the parameter SN mean RV obtained from the SN fit than the Normal one. The p-values associated with each R are, from the top to the bottom, from the left to the right: $2.2e-16$, 0.02, 0.27, 0.05, $7e-4$, $1e-5$, $1e-5$, 0.0007, 0.0004, 0.5, 0.2.



Fig. A.6. (top) The RVs (black dots) for HD215152 estimated using a SN and a Normal fit. (bottom) The residuals from the model fit using Eq. 5 (New corr. std–black dots) and the residuals from the usual correction (Usual corr. std–blue triangles), based on $RV_{\text{activity}} = \beta_0 + \beta_1\gamma + \beta_2\text{SN FWHM}$ for the SN fit and on $RV_{\text{activity}} = \beta_0 + \beta_1\text{BIS SPAN} + \beta_2\text{FWHM}$ for the normal fit. The residuals for both the proposed correction from stellar activity are comparable.



Fig. A.8. (top) The RVs (black dots) for Corot 7 estimated using a SN and a Normal fit. (bottom) The residuals from the model fit using Eq. 5 (New corr. std–black dots) and the residuals from the usual correction (Usual corr. std–blue triangles), based on $RV_{\text{activity}} = \beta_0 + \beta_1 \gamma + \beta_2 \text{SN FWHM}$ for the SN fit and on $RV_{\text{activity}} = \beta_0 + \beta_1 \text{BIS SPAN} + \beta_2 \text{FWHM}$ for the normal fit. The residuals have a smaller systematic component when using the proposed model of Eq. 5 (black dots) compared to the usual model (blue triangles). Moreover, once corrected from stellar activity using Eq. 5, the systematic component using the SN is 0.334 m s^{-1} smaller than the one retrieved with the Normal analysis.