Zade Johnston
Project Writeup

Week 1: 11 Nov - 13 Nov

Started this week learning about how to use MPI in Athena++. Followed the 3D blast wave tutorial and managed to get it working on Thunderbird with HDF5. Jono also recommended learning a text editor, so I spent an hour learning the basics of Vim. **Note:** HDF5 is better to use than .vtk with parallel computing as it allows all the processors being used to write the data in one file compared to a file for each processor that need to be joined (less hassle).

Jono then gave me an Athena++ hydrodynamic turbulence input script to play around with, and some MATLAB scripts that analyse the energy spectrum of the fluid in question. At the moment, we model the fluid in a cube. There are two different modes that we're wanting to focus on: decaying turbulence (disturb the fluid initially then leave it to its own devices) and continuously driven turbulence. Ran the Athena++ code with 3 different grid sizes for both modes; see screenshots below.

(a) Grid Size: 32   (b) Grid Size: 64   (c) Grid Size: 128

Figure 1: Face-on view of the 3D forced turbulence simulations with different grid sizes; density plotted

These were just the simulations with none of the parameters changed in the input script. Using the MATLAB scripts I also obtained the energy spectrum of the fluid; was very rough due to the low resolution. Wanted to try start a 256^3 grid size simulation before I left on Wednesday but didn't have enough time to set up; will try again later. Thursday and Friday of this week were spent at the Otago Software Carpentry Workshop.

Next Week: Play around with parameters. Want to run the larger grid size simulations to obtain a better energy spectrum that fits the $k^{-5/3}$ law, will add screenshots of spectrum then. Plot the total energy over time for both modes; should observe fluctuations in the energy for forced turbulence.

Week 2: 18 Nov - 22 Nov

This week I ran bigger simulations for both decay and forced turbulence from last week in order to be able to plot the energy spectrum and time evolution. The grid size ranged from 32 to 256, and runs over 30 seconds. All simulations left the parameters in `athinput.turb` unchanged.

Calculated the turnover time $\tau \sim L/u_l$ of eddies on the scale of the box to get an idea of the timescales involved in the energy cascade. This is important in the decaying case as all the input energy dissipates within a few turnover times, so this allowed me to get an approximate time range to average the energy spectrum over. For the decay simulations I used $L = 1$ and $u_L = \sqrt{u_x^2 + u_y^2 + u_z^2}$ taken at the start of the simulation; for the continuous simulations the average turnover time in the saturated state was used. The data was gotten from the `Turb.hst` file.

Continuous Forcing: For the continously forced case, the 256 grid size gave the best result. This is expected as it is able to simulate smaller scale eddies compared to lower resolution simulations, allowing more of the energy cascade to be observed. We see that the energy spectrum does follow the $k^{-5/3}$ law (shown in Fig. 2a) for a given range of wavevectors.

The energy evolution (Fig. 2b) shows an increase in kinetic energy until it levels out after a few turnover times. This leveling out is due to the energy dissipation rate matching the energy input rate from the forcing. There is still some variation around the mean value, which arises from fluctuations due to turbulence.

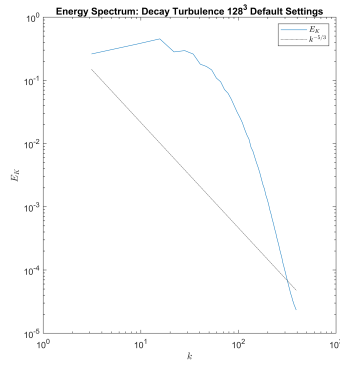(a) Averaged Spectrum
(b) Energy Evolution

Figure 2: Plots showing the energy spectrum and time evolution of the 256 grid size continuous forcing simulation

Decay: The energy spectrum (Fig 3a), averaged over the first two turnover cycles, is not

as well defined as the continuous case. I think this is because the energy depends on time instead of being approximately constant. The energy evolution (Fig. 3b) shows that the total kinetic energy decreases over time due to dissipation from viscosity.

For the spectrum evolution (Figure 3c) I took snapshots of the spectrum at 3 second intervals, with a snapshot at 0.1 seconds to observe the energy input (the curve sharply peaked at $k \sim 10$). We see that the energy cascades down through the length scales directly after input (at 6 seconds), and then decreases as it is dissipated as heat at the micro scale (shown by the "sinking" of the spectrum). This is expected as we saw that the total energy in Fig. 3b is decreasing.

(a) Averaged Spectrum

(b) Energy Evolution

(c) Spectrum Evolution

Figure 3: Plots showing the energy spectrum and time evolution of the 128 grid size decay simulation

The magnetic and thermal energy had no relevance to these simulations; it's just part of the MATLAB script that I forgot to remove.

I learnt that it's good to run the simulations at different resolutions starting with the lowest as it allows you to get a rough idea of what is going to happen without the expense of computing time. It also helps as you can compare with the model to see whether the configuration used is worth investigating. The higher resolutions could differ as turbulence depends strongly on all length scales in the inertial subrange, which are included in the bigger simulations, but it still helps to see what could happen.

Week 3: 25 Nov - 29 Nov

Ran the linear wave problem in Athena to observe Alfvén waves. The pressure and density of the plasma should be constant as these waves travel, which was observed in the simulation (Fig. 4). Alfvén waves have the following dispersion relationship:

$$\omega = k_{\parallel} v_A$$

where $k_{\parallel}$ is the component of the wavevector parallel to the magnetic field and $v_A = B_0/\sqrt{4\pi\rho}$ is the Alfvén velocity at which the waves move, where B_0 and ρ are the magnetic field strength and fluid density.

The linear wave problem has the default settings (I'm not sure about units): $\rho = 1$, $\mathbf{B}_0 = (1, \sqrt{2}, 0.5) \implies B_0 = \sqrt{13}/2$. This gives a value for the Alfvén velocity of $v_A \approx 0.5086$. The next thing to do is to see whether the wave gives us this value through the dispersion relationship above.

(a) Velocity Magnitude   (b) Pressure   (c) Density

Figure 4: Plots of different properties of the fluid as a linear wave passes through

I wrote a small Python script to calculate and compare the simulation and theoretical Alfvén velocities. The angular frequency was calculated by measuring the (simulated) time it took for a peak of the wave to travel through the box, as it has periodic boundary conditions, and dividing this by 2π. The magnitude of the wavevector was gotten by dividing 2π by the wavelength. All quantities were eyeballed as only a quick check was needed. The result is seen in Fig. 5, showing the dispersion relationship holds in the simulation.

```
print("Alfven speed:", v_A) Alfven speed: 0.5085536181410275
print("Calculated speed:", v_pred) Calculated speed: 0.5045045045045046
print("Ratio:", v_pred / v_A) Ratio: 0.9920379808695018
```

Figure 5: Output of the Python script I wrote

Questions (answered):

- Is **k** automatically parallel to **B** when the `ang_3_vert` variable is set to `True`?

  No, it just rotates the axes such that **k** is along the y-axis. Can make it parallel to **B** through code.

- Why are the density and pressure constant for an Alfvén wave?

  Because the Alfvén wave is a **shear wave** i.e. no compression at all. This means $\nabla \cdot \mathbf{u} = 0 \implies D\rho/Dt = 0$ i.e. the density is constant in time and thus so is the pressure.

- Why is a rectangle used insted of a square for the region of interest?

  Just to make the numbers work out right such that the waves are continuous over the periodic boundaries. Can make it a square if needed.

The next step was simulating turbulence in the case of MHD. Essentialy the same code was used as the hydrodynamic turbulence case, with a couple of updates to allow for the initialisation of a mean magnetic field in the region of interest.

In order to get critical balance, we want $u_\perp \sim v_{A0} = B_0/\sqrt{4\pi\rho}$. Athena sets $4\pi \to 1$, so we get $v_{A0} = B_0/\sqrt{\rho}$. We can set $B_0 = 1$ and $\rho = 1$ such that $u_\perp \sim v_{A0} = 1$. To do this in the simulation I've set $dE/dt = 1$ as this means $u_\perp^3/L \sim 1$ (the box size has $L = 1$; letting everything work out to 1 makes it simpler).

Forgot to add the magnetic field flag so have to redo larger simulations. **A good check is to look at the first few snapshots in VisIt to see if the magnetic fields (or other parameters) are set up the way I wanted.**

Trying out Athena's Python codes for plotting and reading data files to help write the `spectrum.m` code in Python.

Theory

Kolmogorov energy cascade law: Energy that is put into a turbulent fluid at the largest relevant scale doesn't immediately dissipate as heat. Instead it is cascaded down the length scales of the fluid (from macro to micro) in the formation of increasingly smaller eddies, and is dissipated as heat at the microscopic scale due to viscous effects. This produces a kind of order out of the (seemingly) chaos of turbulence.

For a given eddy length scale l, we can find a relation between the wavenumber ($k = 2\pi/l$) and the energy contained in the fluid $E(k)$. This theory was developed by Kolmogorov and says $E(k) \sim k^{-5/3}$ for a given range of wavenumbers, called the **inertial subrange**. This law fails at small scales (large k) for the reason explained above, and at large scales (small k) as this is where the energy is being injected into the system.

Derivation using dimensional analysis: The Navier-Stokes equation is

$$\frac{\partial \mathbf{u}}{\partial t} + (\mathbf{u} \cdot \nabla)\mathbf{u} = -\nabla p + \nu \nabla^2 \mathbf{u}$$

Turbulence arises from the nonlinear term $(\mathbf{u} \cdot \nabla)\mathbf{u}$ (represents the energy cascade), and energy is dissipated via the viscosity term $\nu \nabla^2 \mathbf{u}$. Using dimensional analysis, at a given scale l with bulk velocity u_l we have $|(\mathbf{u} \cdot \nabla)\mathbf{u}| \sim u_l^2/l$ and $|\nu \nabla^2 \mathbf{u}| \sim \nu u_l/l^2$ (using $\nabla \sim 1/l$).

Reynolds number Re is the ratio of the turbulence term to the viscosity term:

$$\text{Re} \sim \frac{|(\mathbf{u} \cdot \nabla)\mathbf{u}|}{|\nu \nabla^2 \mathbf{u}|} = \frac{u_l l}{\nu}$$

The viscous term is much smaller than the nonlinear term for large l (due to the factor of $1/l^2$). This means that at large scales energy can't dissipate as heat, so the nonlinear term effectively cascades energy through eddies as it has nowhere else to go. Only when the two terms are comparable (Re ~ 1) can the energy be dissipated as heat. This also explains the self-similarity in the energy spectrum across the inertial subrange.

For a given length scale l and velocity u_l (this is the **structure function**), the energy is $E \sim u_l^2$. The turn-over time for an eddy is $\tau \sim l/u_l$. The rate of energy dissipation is then $E/\tau \sim u_l^3/l = \text{const.}$ across the inertial subrange due to self-similarity.

We then have $u_l \sim l^{1/3} \sim k^{-1/3} \equiv u_k$. The energy in all eddies with wavenumber $\geq k$ is given by

$$E = u_k^2 \sim k^{-2/3} \sim \int_k^\infty E(\kappa) \, d\kappa \implies E(k) \sim k^{-5/3}$$

giving us Kolmogorov's result.

The structure function $\langle [u_l]^2 \rangle = \langle [u(x + l) - u(x)]^2 \rangle$ measures how smooth the velocity distribution is for a given length scale. It is small at large l as velocity is continous at that scale, and large at small l as velocity can change appreciatively over that scale.

MHD Equations: Using $D/Dt = \partial/\partial t + \mathbf{u} \cdot \nabla$ as the Lagrangian derivative.

$$\frac{D\rho}{Dt} = -\rho \nabla \cdot \mathbf{u} \tag{1}$$

$$\rho \frac{D\mathbf{u}}{Dt} = -\nabla \left(c_s^2 \rho + \frac{B^2}{8\pi} \right) + \frac{(\mathbf{B} \cdot \nabla)\mathbf{B}}{4\pi} + \nu \nabla^2 \mathbf{u} \tag{2}$$

$$\frac{D\mathbf{B}}{Dt} = (\mathbf{B} \cdot \nabla)\mathbf{u} - (\nabla \cdot \mathbf{u})\mathbf{B} + \eta \nabla^2 \mathbf{B} \tag{3}$$

Equation 1 is the mass continuity equation (what comes in must go out).

Equation 2 gives the time evolution of the fluid due to forces, and is obtained from the Navier-Stokes equation:

- $-\nabla \left(c_s^2 \rho + \frac{B^2}{8\pi} \right)$: Says that the higher the mass and magnetic field density, the greater the force is on the fluid in a direction away from that density.

- $\frac{\mathbf{B} \cdot \nabla \mathbf{B}}{4\pi}$: acts as a "tension" on the magnetic field lines; i.e. if the magnetic field lines are perturbed it acts to restore to its inital condition.

- $\nu \nabla^2 \mathbf{u}$: usual viscosity term in the fluid.

Equation 3 determines the evolution of the magnetic field. Looks similar to the mass continuity equation, and is obtained from manipulating the Lorentz force equation and Maxwell's equations. Often called the "frozen in" condition as field lines move with the fluid:

- $(\mathbf{B} \cdot \nabla)\mathbf{u}$ and $(\nabla \cdot \mathbf{u})\mathbf{B}$: acts to stretch the field lines with the fluid. The first is stretching due to shear flows while the second is stretching due to compressible flows (not as important as we usually have $\nabla \cdot \mathbf{u} = 0$).

- $\eta \nabla^2 \mathbf{B}$: The "viscosity" term of the magnetic field; η is the resistance of the charged fluid. Gives a similar spectrum for the energy in the magnetic field as the one for the kinetic energy.

Critical Balance: For turbulence in MHD, critical balance is the assumption that

$$k_\perp u_{\perp k} \sim k_\parallel v_A$$

across all scales in the inertial range, where $k_\perp$ and $k_\parallel$ are the components of $\mathbf{k}$ perpendicular and parallel to the magnetic field. Unlike in hydrodynamic turbulence, eddies in MHD spread out along the magnetic field lines as energy is cascaded down.

Alfvén waves are shear waves (i.e. doesn't compress fluid $\iff \nabla \cdot \mathbf{u} = 0$) that travel through a charged fluid along magnetic field lines. Alfvén waves have the following dispersion relationship:

$$\omega = k_\parallel v_A$$

where $k_\parallel$ is the component of the wavevector parallel to the magnetic field and $v_A = B_0/\sqrt{4\pi\rho}$ is the Alfvén velocity at which the waves move, where B_0 and ρ are the magnetic field strength and fluid density. Because they travel along the magnetic field lines, viscosity doesn't affect them as much as it does sonic waves because the particles can't "talk" to each other. This means that smaller scale structures that would otherwise have been lost due to viscosity survive; this is what allowed finer details in the solar wind to be observed near Earth.

When simulating turbulence in MHD we choose a box whose size is small compared to the length scales of the forcing. We can consider the "forcing" on this box to be the energy coming from the energy cascade at larger scales. We align the box such that the longest edge is parallel to the magnetic field lines. Call this length L_x, and L_y the length of the side perpendicular to the magnetic fields. Then:

$$k_\parallel = \frac{2\pi}{L_x}, \ k_\perp = \frac{2\pi}{L_y}$$

In order to respect the assumption of critical balance, these lengths must satisfy $L_y/L_x \sim u_{\perp k}/v_A$.