

Supplementary wavelength calibration methods for SALT/RSS spectropolarimetric observations

Justin Cooper

Submitted in fulfillment of the requirements for the degree

Magister Scientiæ

in the Faculty of Natural and Agricultural Sciences

Department of Physics

University of the Free State

South Africa

Date of submission: September 2024

Supervised by: Prof. B. van Soelen, Department of Physics

Abstract

TODO:

- Done last
- Flow from use of SALT and pipeline and basics of its science implementations into why a more streamlined wavelength calibration is an improvement.
- Give summary of results.
- Aim for a paragraph (~ 600) without going too in-depth into anything specific.
- Brian's comment: Abstract should summarize paper. Include results, conclusions, etc.

Keywords: STOPS, POLSALT, IRAF, SALT, RSS, Development: Python, Pipeline, Calibration: wavelength, Polarization: optical, galaxies: AGN, Blazars, Spectropolarimetry, Astrophysics, Astronomy,

TODO:

- Add Keywords → look up the astronomy journal keywords
- Look up keywords for pipeline development and data reduction.

Acknowledgements

I hereby acknowledge and express my sincere gratitude to the following parties for their valuable contributions:

- **TODO: Add acknowledgements!**

Contents

1	Introduction	1
1.1	Outline	1
2	Spectropolarimetry and the SALT RSS	3
2.1	Spectroscopy	3
2.1.1	Telescope Optics	3
2.1.2	Slit	4
2.1.3	Collimator	4
2.1.4	Dispersion Element	4
2.1.5	Camera Optics	5
2.1.6	Detector	5
2.1.7	Dispersion of Light	5
2.1.8	Detector and Spectroscopic Calibrations	9
2.2	Polarimetry	15
2.2.1	Polarization	16
2.2.2	Polarization Measurement	19
2.2.3	Polarimetric Calibrations	22
2.3	Spectropolarimetry	23
2.3.1	Spectropolarimetric Measurement	24
2.3.2	Spectropolarimetric Calibrations	25
2.4	The Southern African Large Telescope	26
2.4.1	The Primary Mirror	26
2.4.2	Tracker and Tracking	27
2.4.3	SALT Instrumentation	28
3	Existing and Developed Software	31
3.1	POLSALT	31
3.1.1	Raw Image Reductions	32
3.1.2	Wavelength Calibrations	32
3.1.3	Spectral Extraction	33
3.1.4	Raw Stokes Calculations	33
3.1.5	Final Stokes Calculations	34
3.1.6	Visualization	34
3.1.7	Post-Processing Analysis	34
3.1.8	POLSALT Limitations and the Need for Supplementary Tools	35
3.2	IRAF	36

3.2.1	Identify	37
3.2.2	Reidentify	37
3.2.3	Fitcoords	39
3.2.4	Transform	40
3.3	STOPS	41
3.3.1	Splitting	41
3.3.2	Joining	43
3.3.3	Sky Line Checks	47
3.3.4	Cross Correlation	47
3.4	General Reduction Procedure	50
3.4.1	Initial Setup	50
3.4.2	POLSALT Pre-Reductions	51
3.4.3	Wavelength Calibration	52
3.4.4	POLSALT Reduction Completion	54
4	Testing and Application	57
4.1	Testing STOPS	57
4.1.1	Testing the split Method	57
4.1.2	Testing the join Method	59
4.2	Wavelength Solution Checks	61
4.2.1	Cross Correlation Checks	61
4.2.2	Sky Line Checks	63
4.3	Application of STOPS	63
4.3.1	Spectropolarimetry and Photometry of the Early Afterglow of the Gamma-ray Burst GRB 191221B	64
4.3.2	Development of STOPS: Application to the blazar 3C 279	64
4.3.3	SALT Spectropolarimetric Pipeline Comparisons	65
4.3.4	Modeling the Spectral Energy Distributions and Spectropolarimetry of Blazars - Application to 4C+01.02 in 2016 - 2017	66
5	Conclusions	69
5.1	Future Work	69
A	The Modified Reduction Process	71
B	STOPS Source Code	79
C	Proceedings	121
	Bibliography	137

List of Acronyms and Symbols

A-DC	Analog-to-Digital Converter
ADC	Atmospheric Dispersion Compensator
Ar	Argon
CCD	Charged-Coupled Device
CLI	Command Line Interface
CMOS	Complementary Metal-Oxide-Semiconductor
CuAr	Copper-Argon
Fermi LCR	Fermi Light Curve Repository
FITS	Flexible Image Transport System
FORS1	Focal Reducer and Low Dispersion Spectrograph
FORS2	Focal Reducer and Low Dispersion Spectrograph 2
FWHM	Full Width at Half Maximum
GUI	Graphical User Interface
HDU	Header and Data Unit
HEASA	High Energy Astrophysics South Africa
HET	Hobby-Eberly Telescope
HgAr	Mercury-Argon
HRS	High Resolution Spectrograph
IRAF	<i>Image Reduction and Analysis Facility</i>
L+45°	Linear +45° Polarized
L-45°	Linear -45° Polarized
LCP	Left Circularly Polarized
LHP	Linear Horizontally Polarized
LVP	Linear Vertically Polarized
Ne	Neon-Argon
NIR	Near Infra-Red
NIRWALS	Near Infra-Red Washburn Labs Spectrograph
POLSAIT	<i>Polarimetric reductions for SALT</i>
RCP	Right Circularly Polarized
RMS	Root Mean Square
RSS	Robert Stobie Spectrograph
S/N	Signal-to-Noise Ratio
SAAO	South African Astronomical Observatory
SAC	Spherical Aberration Corrector
SALT	Southern African Large Telescope
SALTICAM	SALT Imaging Camera

STOPS	<i>Supplementary Tools for POLSALT Spectropolarimetry</i>
ThAr	Thorium-Argon
UV	Ultraviolet
VLT	Very Large Telescope
VPH	Volume Phase Holographic
Xe	Xenon

Chapter 1

Introduction

TODO: Very short intro to Spectroscopy, Polarization, and Spectropolarimetry and their importance in astronomy

TODO: Focus on AGN implications and implementations such as the types of objects and a short history for each type of object, Blazar focus with specification on BL Lacs and FSRQs, the Unified Model, The Blazar sequence.

Brian's comment: Highlight importance of polarimetry for understanding emission and how that plays a role in AGN.

TODO: Basics of modelling (Different energy/wavelength ranges used and what the models tell us about emission processes/structure) so that Hester's results can be noted for applications of the pipeline.

1.1 Outline

TODO: Problem Statement, VERY IMPORTANT, roughly a sentence but problem thoroughly fleshed out. **TODO:** General layout of Dissertation

Chapter 2

Spectropolarimetry and the SALT RSS

This chapter gives an overview of the basics of spectropolarimetry (§ 2.3), and how it functions, following from the principles of both spectroscopy (§ 2.1) and polarimetry (§ 2.2). Further, it is discussed how these techniques are practically implemented for Southern African Large Telescope (SALT) (§ 2.4), using the Robert Stobie Spectrograph (RSS) (§ 2.4.3), and how the spectropolarimetric reduction process is completed (§ 2.4.3).

2.1 Spectroscopy

Spectroscopy originated in its most basic form with Newton's examinations of sunlight through a prism (Newton and Innys, 1730) but came to prominence as a field of scientific study with Wollaston's improvements to the optics elements (Wollaston, 1802), Fraunhofer's use of a diffraction grating instead of a prism (der Wissenschaften, 1824), and Bunsen and Kirchoff's classifications of spectral features to their respective chemical elements (Kirchhoff and Bunsen, 1861).

The simplest spectrometer schematic, as shown in Figure 2.1, consists of incident light collected from the telescope's optics, labelled A, being focused onto a slit, B, and passed through a collimator, C. The collimator collimates the light allowing a dispersion element, D, to disperse the light into its constituent wavelengths. The resultant spectrum is focused by camera optics, E, onto a focal plane, F. Viewing optics are situated at the focal plane in the case of a spectroscope and a detector is situated at the focal plane in the case of a spectrograph.

2.1.1 Telescope Optics

The telescope optics refers simply to all the components of a telescope necessary to acquire a focal point at the spectrometer entrance, labelled B. The focal point in most traditional telescope designs is fixed relative to the telescope and so the spectrometer may be mounted at that point. In cases where the telescope is designed to have a moving focal point relative to the telescope (see Buckley et al., 2006; Cohen, 2009; Ramsey et al., 1998), the spectrometer, or a signal transfer method such as a fibre feed to the spectrometer, must also move along the telescope's focal path.



Figure 2.1: Layout depicting the light path through a spectrometer. Diagram adapted from Birney et al. (2006).

2.1.2 Slit

The slit's function is to control the amount of incident light entering a spectrometer and, along with the exposure time of the detector, prevents over-exposures of bright sources on highly sensitive detectors (Tonkin, 2013). If a source is spatially resolvable, or larger than the seeing conditions, the slit additionally acts to spatially limit the source to increase the spectral resolution, resulting in sharper features in the resultant spectrum. Without the slit the spectral resolution would be determined by the projected width of the source on the detector, or the seeing if the source was a star-like point source. Increasing the spectral resolution comes with the trade-off of decreasing the light collected from the source and thus acquiring a less intense resultant spectrum. Multiple spectra may be acquired simultaneously when the slit is positioned such that collinear sources lie along the slit.

The spectrometer is usually situated at the focal point. In cases where this is not feasible due to restrictions, for example restrictions of weight or size, a fibre feed may be situated behind the slit on the telescope. This allows the signal to be routed away from the telescope to a controlled environment with only minuscule losses.

2.1.3 Collimator

The collimators function is to collimate the focused light from the telescope, ensuring that all light rays run parallel before reaching the dispersion element. The focal ratio of the collimator (f_c/D_c , where f refers to the focal length and D refers to the diameter) should ideally match the focal ratio of the telescope (f_T/D_T).

2.1.4 Dispersion Element

Including a dispersion element in the optical path is what defines a spectrometer. As the name suggests, a dispersion element disperses the light incident on it into its constituent wavelengths and produces a spectrum. There are two types of dispersion elements, namely the prism and the diffraction grating, which operate on different principles, as discussed in § 2.1.7.

2.1.5 Camera Optics

The lens functions similarly to that of the telescope's optics but in this case focuses the dispersed light onto a receiver situated at the focal plane. As mentioned previously, an eye piece is fixed to the focal point for a spectroscope while a spectrograph employs a detector.

2.1.6 Detector

The two most prevalent detector types in spectroscopy are the Charged-Coupled Device (CCD) and Complementary Metal-Oxide-Semiconductor (CMOS) detectors. In astronomical spectroscopy however, sources are fainter and exposure times are much longer and so the CCD detectors are by far the preferred detector as their output has a higher-quality and lower-noise when compared to CMOS cameras under the same conditions (Janesick et al., 2006).

The CCD is a detector composed of many thousands of pixels which can store a charge so long as a voltage is maintained across the pixels. Each pixel detects incoming photons using photo-sensitive capacitors through the photoelectric effect and converts the photons to a charge (Buil, 1991). There are also thermal agitation effects which introduce noise to the charge accumulated by a pixel, further discussed in § 2.1.8. Once the exposure is finished the accumulated charge is read column by column, row by row, through an Analog-to-Digital Converter (A-DC) which produces a two-dimensional array of ‘counts’.

2.1.7 Dispersion of Light

Light can be broken up into its constituent wavelengths through two different physical phenomena, namely dispersion and diffraction, which dispersive elements use to create spectra. Dispersive prisms and diffractive gratings each have their strengths and weaknesses and a wide spectrum of instruments exist which implement either, or both, concepts. Regardless of the specific element, dispersive elements all have a resolving power, R , and an angular dispersion. Generally, while the angular dispersion is a more involved process to determine, the resolving power of a spectrograph can be measured as:

$$R = \frac{\lambda}{FWHM}, \quad (2.1)$$

where λ is the wavelength of an incident monochromatic beam and Full Width at Half Maximum (FWHM) refers to the width of the feature on the detector at half of its maximum intensity.

Prism

The prism operates on the principle that the refractive index of light, n , varies as a function of its wavelength, λ . Prisms were the only dispersive elements available for early spectroscopic studies, but they were not without flaw. The angular dispersion of a prism is given by:

$$\frac{\partial\theta}{\partial\lambda} = \frac{B}{a} \frac{dn}{d\lambda}, \quad (2.2)$$

where θ is the angle at which the refracted light differs from the incident light, λ is the wavelength of the incident light, B is the longest distance the beam would travel through



Figure 2.2: Geometry of a prism refracting an incident monochromatic beam at a minimum deviation angle. Diagram adapted from Birney et al. (2006).

the prism. $a = L \sin(\alpha)$ is the maximal beam width that would fit onto a prism with a transmissive surface of length L for a given angle, α , at which a beam would strike the transmissive surface, as shown in Figure 2.2.

The refractive index of a material as a function of its wavelength, $n(\lambda)$, can be approximated by Cauchy's equation:

$$n(\lambda) = A_C + \frac{B_C}{\lambda^2} + \frac{C_C}{\lambda^4} + \dots, \quad (2.3)$$

where A_C, B_C, C_C are the Cauchy coefficients and have known values for certain materials. Cauchy's equation is a much simpler approximation of the refractive index that remains very accurate at visible wavelengths (Jenkins and White, 1976). Taking only the first term of the derivative of the Cauchy equation allows us to approximate the angular dispersion of a prism,

$$\frac{\partial \theta}{\partial \lambda} = -\frac{B}{a} \frac{2B_C}{\lambda^3} \propto -\lambda^{-3}, \quad (2.4)$$

which shows that the angular dispersion of a prism is wavelength dependent and furthermore that longer wavelengths are dispersed less than shorter wavelengths (Birney et al., 2006; Hecht, 2017). The dependence of the angular dispersion, $d\theta/d\lambda$, on the wavelength, λ , is crucial for the formation of a spectrum but this cubic, non-linear, relation results in a non-linear spectrum. Since prisms rely on the refractive index of the material they are made of, they have low angular dispersions.

Multiple prisms can be used to increase the angular dispersion but as the dispersion is non-linear it becomes increasingly more difficult to calibrate. The more material and material boundaries the light must pass through, the more its intensity decreases due to attenuation effects and Fresnel losses. Even so, the transmittance of modern prisms for their selected wavelength range is generally very high due to improved manufacturing methods as well as improved transmitting materials.¹

¹See manufacturers technical specifications, THORLABS, or Edmund Optics for example.



Figure 2.3: Geometry of a reflective blazed grating refracting an incident monochromatic beam. Diagram adapted from Birney et al. (2006).

Diffraction Grating

The alternative dispersing element is a diffraction grating, which operates on the principle that as light interacts with a grating where the groove size is comparable to the light's wavelength, the light is dispersed through constructive and destructive interference. This interference results in multiple diffracted beams m , called orders, either side of a central reflected, or transmitted, beam such that $m \in \mathbb{Z}$, where $m = 0$ is the non-dispersed, or reflected, beam.

An example of a reflective blazed grating is illustrated in Figure 2.3. Here a monochromatic beam is incident on the grating at an angle of α from the grating normal. Due to the interference, a diffracted beam of wavelength λ is found at an angle of β from the grating normal. The relation between the incident and diffracted beams is given by the grating equation:

$$m\lambda = \sigma(\sin(\alpha) \pm \sin(\beta)), \quad (2.5)$$

where σ is the groove spacing of the grating and m is the order of the diffracted beam being considered. The grating equation also applies to transmission gratings, though care should be taken for the signs of α and β .

Equation 2.5 also shows that different diffracted beams may share an angle of dispersion for beams not in the same order. The regions of an order that do not overlap with another order are called free spectral ranges. An order-blocking filter may be used to account for the overlaps and increase the free spectral range. A diffraction grating can also be blazed by an angle θ , as illustrated in Figure 2.3. Blazing refers to the fact that the grooves on the surface of the grating are not symmetrical. The asymmetry of the grooves diffracts the incident beam such that most of the beam's intensity is found in a reflected, zeroth order, beam. The wavelength at which a blazed spectrograph is most effective is called the blaze wavelength, λ_b , which is determined by:

$$\begin{aligned} m\lambda_b &= 2\sigma \sin(\theta) \cos(\alpha - \theta), \text{ where} \\ 2\theta &= \alpha + \beta. \end{aligned} \quad (2.6)$$

Taking the derivative of Equation 2.5 with respect to λ while keeping α constant, allows us to determine the angular dispersion of a diffraction grating,

$$\frac{\partial \beta}{\partial \lambda} = \frac{m}{\sigma \cos(\beta)}. \quad (2.7)$$



Figure 2.4: Diagram of a grism for an incident monochromatic beam of light and a diffracted beam of order $m = 1$. Diagram adapted from Birney et al. (2006).

Substituting m/σ with the grating equation results in

$$\frac{\partial \beta}{\partial \lambda} = \frac{\sin(\alpha) + \sin(\beta)}{\lambda \cos(\beta)} \propto \lambda^{-1}. \quad (2.8)$$

Similar to the dispersion of a prism, Equation 2.8 shows that the dispersion of a grating is wavelength dependent, but this dependence is only inversely proportional and thus more uniform across a wavelength range than that of a prism. Furthermore, shorter wavelengths are refracted less than longer wavelengths since there is no negative relation between the angular dispersion and the wavelength (Birney et al., 2006; Hecht, 2017).

Alternate Diffraction Elements

As mentioned before, multiple subgroups exist for both dispersive prisms and diffractive gratings. For prisms, along with the single and multiple prism setups mentioned, there also exists grisms and immersed gratings. A grism (Grating Prism), as shown in Figure 2.4, refers to a transmissive grating etched onto one of the transmissive faces of a prism and allows a single camera to capture both spectroscopic and photometric images without needing to be moved, with and without the grism in the path of the beam of light, respectively. An immersed grating refers to a grism modified such that the transmissive grating is coated with reflective material. The primary source of dispersion for both grisms and immersive gratings is the grating and any aberration effects from the prism are negligible in comparison.

Other types of gratings include the Volume Phase Holographic (VPH) grating as well as the echelle grating. The VPH grating consists of a photoresist, which is a light-sensitive material, sandwiched between two glass substrates. Diffraction is possible since the photoresist's refractive index varies near-sinusoidally perpendicularly to the gratings lines, as seen in Figure 2.5. This allows for sharper diffraction orders and low stray light scattering as compared to more traditional gratings but since blazing is not possible the

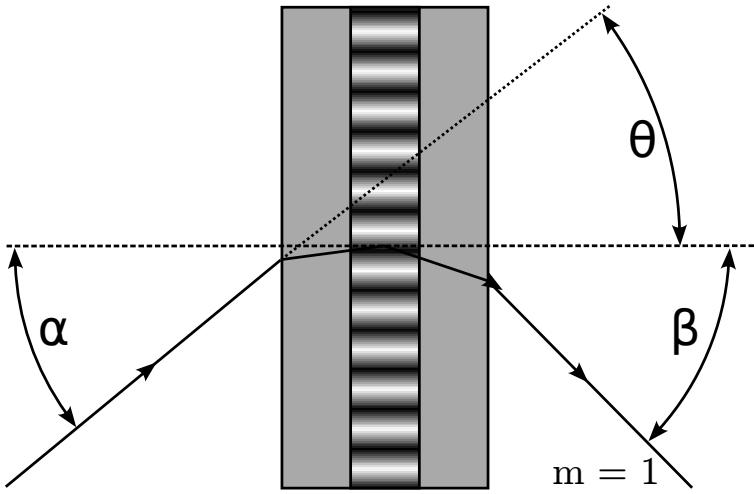


Figure 2.5: Diagram of a VPH grating for an incident monochromatic beam of light. Diagram adapted from Birney et al. (2006).

efficiency is decreased. An echelle grating refers to a diffraction grating with higher groove spacing which is optimized for use at high orders. The high order of the diffracted beam allows for greater angular dispersion which is most useful when combined with another dispersion element to cross-disperse a spectrum, resulting in a high resolution spectrum.

2.1.8 Detector and Spectroscopic Calibrations

Acquiring a spectrum from observations is more involved than simply reading out the data recorded on the CCD. A raw science image, which is the raw counts of the observed source read from the CCD with no calibrations applied, has on it a combination of useful science data as well as noise. The noise is a combination of random noise introduced through statistical processes and systematic noise introduced through the instrumentation and the observation conditions the source was observed under. This noise causes an uncertainty in the useful data and can be minimized, predominantly by calibrating for the systematic noise, but never fully removed (Howell, 2006).

The dominant source of noise in a raw image is detector noise. CCDs are manufactured to have a small base charge in each pixel, called the ‘bias’ current which allows the readout noise, a type of random noise, to better be sampled. There is also an unintentional additional charge which is linearly proportional to the exposure time and originates from thermal agitation of the CCD material, called the ‘dark’ current. The dark current can be minimized and possibly ignored if the CCD is adequately cooled. These types of noise add to the charge held by a pixel and are thus considered additive.

The CCD is not a perfect detector and the efficiency of it and the optics of the telescope also contribute noise to the image. The efficiency of a CCD is referred to as the Quantum Efficiency, and it is a measure of what percentage of light striking the detector is actually recorded and converted to a charge. The efficiency of the CCD and telescope optics is also wavelength dependent and so the noise that results from them is more complex than that of additive noise. This type of noise is referred to as multiplicative noise.

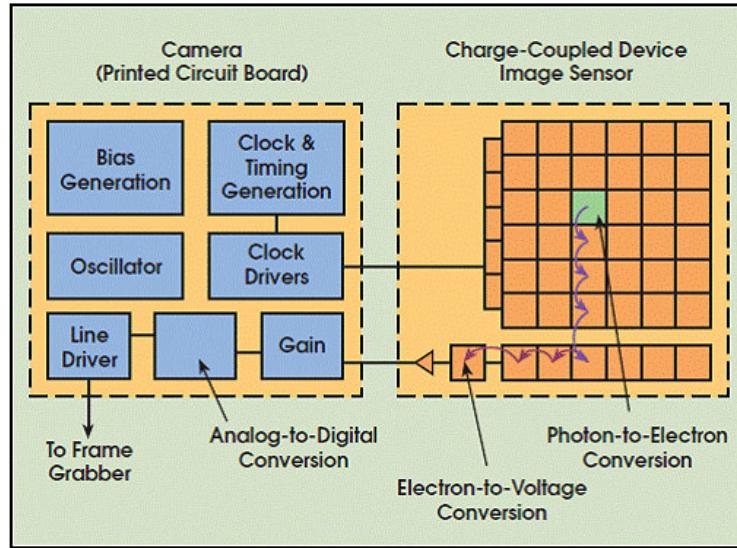


Figure 2.6: Diagram of the inner logic of a CCD. Figure adapted from Litwiller (2001).

Additive noise, such as bias and dark currents, is inherent to CCD images, and as such needs to be subtracted out first when performing calibrations. Bias currents can be found by taking a bias image or by adding an overscan region to each image. A bias image is an image where the charges on the CCD are reset and then immediately read off without exposing anything on the detector, effectively taking an image with zero exposure time. Alternatively, to save time during an observational run, overscan regions may be added to the images. An overscan region refers to adding a few cycles to the readout of each column of the CCD such that the base current is read out and appended to each image.

Dark currents can be found by taking an image with nothing exposed onto the detector for a certain exposure time. This resultant dark image can then be scaled to the science images exposure time since the dark current should be linearly proportional to exposure time. When the detector is capable of being held at precise temperatures, dark images may be taken over multiple hours during the day to produce a high quality master dark image that may then be scaled and subtracted from all subsequent images.

Next, multiplicative noise, such as a CCD’s pixel-to-pixel response, should be accounted for. This pixel-to-pixel response should be uniform across the image and to achieve this an average response may be divided out. The average response is referred to as a ‘flat’ image or flat-field and may be acquired by observing a uniformly illuminated surface to determine the pixel-to-pixel response.

Dome flats are images taken of a relatively flat surface, usually the inside of a telescopes dome, and are used in both photometry and spectroscopy. The surface is uniformly and indirectly illuminated by a projector lamp, ideal for flat-field images. Alternate flat-fielding methods, such as night sky and twilight flats, are available but are suited solely for photometry.

Night sky flats are produced from science images containing mostly sky. The science images are combined using the ‘mode’ statistic which removes any celestial objects at the cost of a low Signal-to-Noise Ratio (S/N) flat-field.

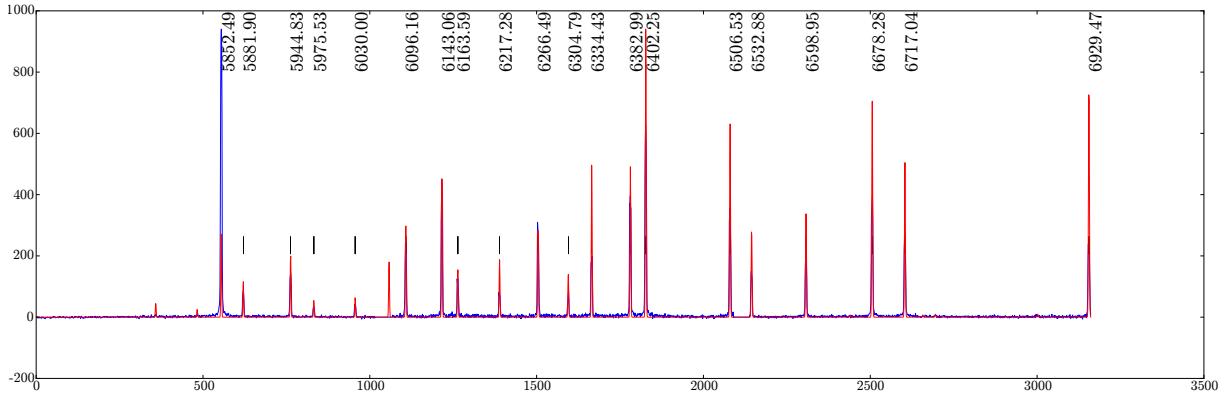


Figure 2.7: Example of an arc spectrum for NeAr taken with SALT’s RSS using the PG1800 grating at a grating angle of 34.625° , an articulation angle of 69.258° , and covering a wavelength range of $\sim 5600 - 6900 \text{ \AA}$. Plot adapted from SALT’s published Longslit Line Atlases, (2023).²

Twilight flats are produced from images of the twilight (or dawn) sky. They are taken when the Sun has just set, in the opposite direction, at $\sim 20^\circ$ from zenith and provide a better S/N at the cost of careful timing of the images.

A flat-field must be normalized before being used to correct any science images since it only acts to account for the pixel-to-pixel response and not for the additive errors. A normalized spectroscopic flat image, $F_\lambda^n(x, y)$, can be calculated as:

$$F_\lambda^n(x, y) = \frac{F_\lambda(x, y) - B(x, y) - (\frac{t_S}{t_D})D(x, y)}{\text{med}_{lp}(F_\lambda(x, y) - B(x, y) - (\frac{t_S}{t_D})D(x, y))}, \quad (2.9)$$

where $F_\lambda(x, y)$ is the non-corrected flat image, $B(x, y)$ is the bias image, $D(x, y)$ is the dark image which is scaled by the exposure time of the science image, t_S , and the dark image, t_D . med_{lp} is a low-pass median filter which smoothes out any rapid changes in the pixel-to-pixel response, removing the illumination contribution.

The calibrated science image, $S_\lambda^*(x, y)$, which accounts for the bias and dark currents as well as the flat fielding can then be calculated as:

$$S_\lambda^*(x, y) = \frac{S_\lambda(x, y) - B(x, y) - (\frac{t_S}{t_D})D(x, y)}{F_\lambda^n(x, y)}. \quad (2.10)$$

When multichannel CCDs are used, which consist of multiple CCDs or a CCD with multiple output amplifiers, additional calibrations, specifically cross-talk corrections and mosaicking, are required. Cross-talk noise refers to contamination that occurs during readout in one channel from another channel with a high signal and occurs because the signals can not be completely isolated from one another. Cross-talk corrections therefore account for this signal contamination between channels being read out at the same time (Freyhammer et al., 2001). Mosaicking is necessary for multichannel CCDs since the digitized signal read out from the detector has no reference of the physical location of the pixel it was detected at. Mosaicking, therefore, correctly orients the data acquired from a multichannel detector so that a single correctly oriented image is produced.

²NeAr plot sourced from <https://astronomers.salt.ac.za/data/salt-longslit-line-atlas/>



Figure 2.8: The first seven Chebyshev polynomials (T_0 through T_6) as defined by Equation 2.12 over the region $[-1, 1]$ for which they are orthogonal. Plot adapted from (Press et al., 2007) (2023)³

Wavelength Calibration

Finally, since the dispersion element breaks the incident light into its constituent wavelengths non-linearly (§ 2.1.7), the relation between the pixel on a detector and the wavelength of the light incident on it is unknown. Ideally, the spectrometer's optics would be modelled to produce a reliable pixel to wavelength calibration (see E.g. Liu and Hennelly, 2022), but this becomes increasingly more difficult for spectrometers with complex, non-sedentary, optical paths.

Alternatively, a source with well-defined spectral features, with said features evenly populating the wavelength region of interest, such as in Figure 2.7 may be observed. The observed frame is commonly referred to as an ‘arc’ frame, after the arc-lamps used to acquire the spectra, and should be observed alongside the science frames over the course of an observation run.

It is important that the arc frame is observed at the same observing conditions and parameters as the science frames since the optical path will vary over the course of an observing run and for different observing parameters, invalidating previously acquired arc frames. The wavelength calibrations then consist of defining a two-dimensional pixel-to-wavelength conversion function from the arc frame which may later be applied to calibrate the science frames. The two most common approximations for wavelength calibrations are the Chebyshev and Legendre polynomial approximations.

³Excellent resources on Chebyshev and Legendre polynomials are available digitally at www.numerical.recipes/book.

Chebyshev Polynomials The Chebyshev polynomials are defined explicitly as:

$$T_n(x) = \cos(n \cos^{-1}(x)), \quad (2.11)$$

or recursively as:

$$\begin{aligned} T_0(x) &= 1, \\ T_1(x) &= x, \text{ and} \\ T_{n+1}(x) &= 2xT_n(x) - T_{n-1}(x), \text{ for } n \geq 1, \end{aligned} \quad (2.12)$$

where T is a Chebyshev polynomial of order n .⁴ An important property of Chebyshev polynomials is that they are orthogonal polynomials. This means that the inner product of any two differing Chebyshev polynomials, $T_i(x)$ and $T_j(x)$, over the range $[-1, 1]$ is zero, as shown by:

$$\int_{-1}^1 T_i(x) T_j(x) \frac{1}{\sqrt{1-x^2}} dx = \begin{cases} 0, & i \neq j \\ \pi/2, & i = j \neq 0 \\ \pi, & i = j = 0 \end{cases}, \quad (2.13)$$

where $1/\sqrt{1-x^2}$ is the weighting factor for Chebyshev polynomials. This property is important because it means that the coefficients in the Chebyshev polynomial expansion are independent of one another, allowing for a unique solution when approximating an unknown function (Arfken and Weber, 1999; Press et al., 2007). A Chebyshev approximation of an unknown wavelength calibration function is given by:

$$f(x) \approx \sum_{i=0}^N c_i T_i(u), \text{ or} \quad (2.14)$$

$$F(x, y) \approx \sum_{i=0}^N \sum_{j=0}^M c_{ij} T_i(u) T_j(v), \quad (2.15)$$

for a one- or a two-dimensional wavelength surface function, respectively. Here N and M are the desired x and y orders, and c_i and c_{ij} are the Chebyshev polynomial coefficients (Florinsky and Pankratov, 2015; Leng, 1997). Since the orthogonality property of the Chebyshev polynomials only holds true over the range $[-1, 1]$, the $(x, y) \in ([0, a], [0, b])$ pixel coordinates must be remapped to $u, v \in [-1, 1]$ following the relation:

$$(u, v) = \frac{2(x, y) - a - b}{b - a}. \quad (2.16)$$

The Chebyshev polynomials are more suited for wavelength calibrations than standard polynomials since they are orthogonal and have minima and maxima located at $[-1, 1]$, as seen in Figure 2.8. This means that the Chebyshev approximation is exact when $x = x_n$, where x_n are the positions of the $n - 1$ x -intercepts of $T_N(x)$. These properties greatly minimize the error in the Chebyshev approximation, even at lower orders (Arfken and Weber, 1999).

⁴Chebyshev polynomials are denoted T as a hold-over from the alternate spelling of ‘Tchebycheff’.



Figure 2.9: The first six Legendre polynomials (P_0 through P_5) as defined by Equation 2.20 over the region $[-1, 1]$ for which they are orthogonal. Plot adapted from Geek3, CC BY-SA 3.0, via Wikimedia Commons (2023).

Legendre Polynomials Similar to the Chebyshev polynomials, the Legendre polynomials may be defined explicitly as:

$$P_n(x) = \frac{1}{2^n n!} \frac{d^n}{dx^n} (x^2 - 1)^n, \quad (2.17)$$

or recursively as:

$$\begin{aligned} P_0(x) &= 1, \\ P_1(x) &= x, \text{ and} \\ (n+1)P_{n+1}(x) &= (2n+1)xP_n(x) - nP_{n-1}(x), \text{ for } n \geq 1, \end{aligned} \quad (2.18)$$

where P is a Legendre polynomial of order n . Legendre polynomials also hold the property of orthogonality. This means that the inner product of any two differing Legendre polynomials, $P_i(x)$ and $P_j(x)$, over the range $[-1, 1]$ is zero, as shown by:

$$\int_{-1}^1 P_i(x) P_j(x) dx = \begin{cases} 0, & i \neq j \\ \frac{2}{2n+1}, & i = j \end{cases}, \quad (2.19)$$

where a weight of 1 is the weighting factor for Legendre polynomials (Dahlquist and Björck, 2003; Press et al., 2007). A Legendre approximation of an unknown wavelength calibration function is given by:

$$f(x) \approx \sum_{n=0}^N a_n P_n(u), \text{ or} \quad (2.20)$$

$$F(x, y) \approx \sum_{i=0}^N \sum_{j=0}^M a_{ij} P_i(u) P_j(v), \quad (2.21)$$

for a one-dimensional wavelength function or a two-dimensional surface function, respectively. Here N and M are the desired x and y orders, u and v are the same mapping variable as in Equation 2.16, and a_{ij} are the Legendre polynomial coefficients.

Legendre polynomials benefit from having the orthogonality condition with no weight necessary ($w = 1$) which makes their coefficients computationally easier to compute but increases the error in a Legendre approximation when compared to that of the error in a Chebyshev approximation for functions of the same order, N (Ismail, 2005).

Regardless of which method of polynomial approximation is chosen, the polynomials are fit by varying the relevant coefficients using the least squares method. The resultant minimized function may then be used to convert the science frames from an (x -pixel, y -pixel) coordinate system to a (λ , y -pixel) coordinate system.

2.2 Polarimetry

Both Huygens and Newton came to the conclusion that light demonstrates transversal properties (Huygens, 1690; Newton and Innys, 1730), which was later further investigated and coined as ‘polarization’ by Malus (Malus, 1809). Malus also investigated the polarization effects of multiple materials including some of which were birefringent, such as optical calcite, which he referred to as Iceland spar after Bartholinus’ investigations of the material (Bartholinus, 1670).

Fresnel built on Malus’ work showing that two beams of light, polarized at a right angle to one another, do not interfere, conclusively proving that light is transversal in nature, opposing the widely accepted longitudinal nature of light due to the prevalent belief in the ether. He later went on to correctly describe how polarized light is reflected and refracted at the surface of optical dielectric interfaces, without knowledge of the electromagnetic nature of light. Fresnel’s equations for the reflectance and transmittance, R and T , are defined as:

$$\begin{aligned} R_s &= \left| \frac{Z_2 \cos \theta_i - Z_1 \cos \theta_t}{Z_2 \cos \theta_i + Z_1 \cos \theta_t} \right|^2, \\ R_p &= \left| \frac{Z_2 \cos \theta_t - Z_1 \cos \theta_i}{Z_2 \cos \theta_t + Z_1 \cos \theta_i} \right|^2, \\ T_s &= 1 - R_s, \text{ and} \\ T_p &= 1 - R_p, \end{aligned} \tag{2.22}$$

where s and p are the two polarized components of light perpendicular to one another, Z_1 and Z_2 are the impedance of the two media, and θ_i , θ_t , and θ_r are the angles of incidence, transmission, and reflection, respectively (Fresnel, 1870).

Nicol was the first to create a polarizer, aptly named the Nicol prism, where the incident light is split into its two perpendicular polarization components, namely the ordinary and extraordinary beams. Faraday discovered the phenomenon where the polarization plane of light is rotated when under the influence of a magnetic field, known as the Faraday effect. Brewster calculated the angle of incidence, $\theta_B = \arctan n_2/n_1$, at which incident polarized light is perfectly transmitted through a transparent surface, with refractive indexes of n_1 and n_2 , while non-polarized incident light is perfectly polarized when reflected and partially polarized when refracted.

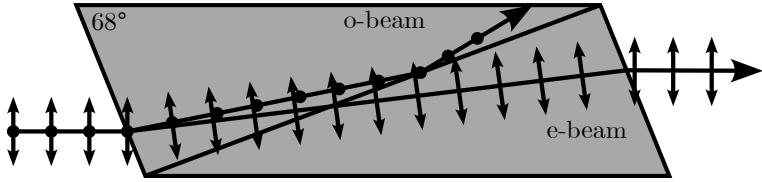


Figure 2.10: Diagram of a Nicol prism for incident non-polarized light. Diagram adapted from Fred the Oyster, CC BY-SA 4.0, via Wikimedia Commons (2023).

Stokes' work created the first consistent description of polarization and gave us the Stokes parameters which describe an operational approach to measuring polarization (discussed further in § 2.2.1) (Stokes, 1852). Hale was the first to apply polarization to astronomical observations, using a Fresnel rhomb and Nicol prism as a quarter-wave plate and polarizer, respectively (Hale, 1908, 1979). Wollaston also created a prism, similarly named the Wollaston prism, which allowed simultaneous observation of the ordinary and extraordinary beams due to the smaller deviation angle (Wollaston, 1802). Finally, Chandrasekhar's work furthered our understanding of astrophysical polarimetry by explaining the origin of polarization observed in starlight as well as mathematically modeling the polarization of rotating stars, which came to be named Chandrasekhar polarization (Chandrasekhar, 1950).

2.2.1 Polarization

Maxwell's equations for an electromagnetic field propagating through a vacuum are given as:

$$\begin{aligned} \nabla \cdot \mathbf{E} &= 0, \\ \nabla \cdot \mathbf{B} &= 0, \\ \nabla \times \mathbf{E} &= -\frac{1}{c} \frac{\partial \mathbf{B}}{\partial t}, \text{ and} \\ \nabla \times \mathbf{B} &= \frac{1}{c} \frac{\partial \mathbf{E}}{\partial t}, \end{aligned} \quad (2.23)$$

where \mathbf{E} and \mathbf{B} are the electric and magnetic field vectors, and c is the speed of light. In a right-handed (x, y, z) coordinate system, a non-trivial solution of an electromagnetic wave following Maxwell's Equations propagating along the z -axis, towards a hypothetical observer, is described by:

$$\begin{aligned} \mathbf{E} &= E_x \cos(kz - \omega t + \Phi_x) \hat{x} + E_y \cos(kz - \omega t + \Phi_y) \hat{y}, \text{ and} \\ \mathbf{B} &= \frac{1}{c} E_y \cos(kz - \omega t + \Phi_y) \hat{x} + \frac{1}{c} E_x \cos(kz - \omega t + \Phi_x) \hat{y}, \end{aligned} \quad (2.24)$$

where E_x , E_y , Φ_x , and Φ_y are all parameters describing the amplitude and phase of the electric field vector in the (x, y) plane, and with the magnetic field vector proportional and perpendicular to the electric field vector (Griffiths, 2005).

Considering only the electric field component and rewriting Equation 2.24 using complex values allows us to simplify the form of the solution to:

$$\mathbf{E} = \Re(\mathbf{E}_0 e^{-i\omega t}), \quad (2.25)$$

where we only consider the real part of the equation, and where \mathbf{E}_0 is defined as:

$$\mathbf{E}_0 = E_x e^{i\Phi_x} \hat{x} + E_y e^{i\Phi_y} \hat{y}, \quad (2.26)$$

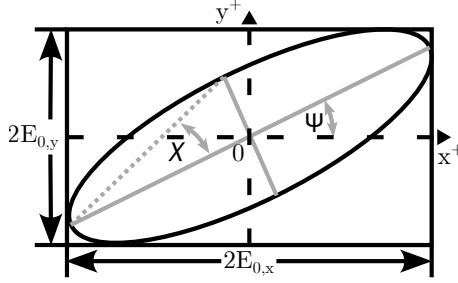


Figure 2.11: The polarization ellipse for an electric field vector propagating through free space. Diagram adapted from Inductiveload, PDM 1.0, via Wikimedia Commons (2023).

and is referred to as the polarization vector since it neatly contains the parameters responsible for the polarization properties (Degl’Innocenti, 2014).

For an electric field vector with oscillations in some combination of the x and y axes, the tip of the vector sweeps out an ellipse, as depicted in Figure 2.11. This ellipse is referred to as the polarization ellipse and has the form:

$$\left(\frac{\mathbf{E}_x}{\mathbf{E}_{0,x}}\right)^2 + \left(\frac{\mathbf{E}_y}{\mathbf{E}_{0,y}}\right)^2 - \frac{2\mathbf{E}_x\mathbf{E}_y}{\mathbf{E}_{0,x}\mathbf{E}_{0,y}} \cos \Phi = \sin^2 \Phi, \quad (2.27)$$

where $\Phi = \Phi_x - \Phi_y$ is the phase difference between the x and y phase parameters. The degree of polarization for the polarization ellipse is related to the eccentricity of the ellipse and the angle at which it is rotated relates to the polarization angle. Since $\mathbf{E}_{0,x}$, $\mathbf{E}_{0,y}$, Φ_x , and Φ_y describe the wave, the polarization ellipse that results from these parameters is fixed as the wave continues to propagate.

Since observations consist of images taken over a desired exposure time, time averaging of Equation 2.27 over the exposure time is necessary. Given the periodical nature and high frequencies of the fields, the time averaging may be found over a single oscillation using:

$$\langle \mathbf{E}_i \mathbf{E}_j \rangle = \lim_{dt \rightarrow \infty} \frac{1}{T} \int_0^T \mathbf{E}_i \mathbf{E}_j dt, \quad \text{for } i, j \in (x, y), \quad (2.28)$$

where T is the total averaging time over the electric field vectors \mathbf{E}_i and \mathbf{E}_j (Collett, 2005). Applying the time averaging to Equation 2.27 and simplifying results in:

$$(E_{0x}^2 + E_{0y}^2)^2 - (E_{0x}^2 - E_{0y}^2)^2 - (2E_x E_y \cos \Phi)^2 = (2E_x E_y \sin \Phi)^2. \quad (2.29)$$

The expressions inside the parentheses can be found through observation and may also be represented as:

$$\mathbf{S} = \begin{pmatrix} S_0 \\ S_1 \\ S_2 \\ S_3 \end{pmatrix} = \begin{pmatrix} I \\ Q \\ U \\ V \end{pmatrix} = \begin{pmatrix} E_{0x}^2 + E_{0y}^2 \\ E_{0x}^2 - E_{0y}^2 \\ 2E_{0x}E_{0y} \cos \Phi \\ 2E_{0x}E_{0y} \sin \Phi \end{pmatrix}, \quad (2.30)$$

where S_0 to S_3 are referred to as the Stokes (polarization) parameters. The parameters describe the: S_0 , total intensity (often normalized to 1); S_1 , ratio of the Linear Horizontally Polarized (LHP) to Linear Vertically Polarized (LVP) light; S_2 , ratio of the Linear $+45^\circ$ Polarized ($L+45^\circ$) to Linear -45° Polarized ($L-45^\circ$) light; and S_3 , ratio of the Right Circularly Polarized (RCP) (clockwise) to Left Circularly Polarized (LCP)



Figure 2.12: The Poincaré sphere describing the polarization properties of a wave-packet propagating through free space. Diagram adapted from Inductiveload, PDM 1.0, via Wikimedia Commons (2023).

(counter-clockwise) light. When the intensity is normalized, the Stokes parameters range from 1 to -1 , based on the dominating component of the parameter (Chandrasekhar, 1950; Stokes, 1852).

From Equation 2.29 and 2.30, the polarization parameters are related by:

$$I^2 = Q^2 + U^2 + V^2, \quad (2.31)$$

for entirely polarized light. Only beams of completely polarized light could be accounted for before Stokes' work on polarization. Using the Stokes parameters, we can now account for partially polarized light such that:

$$I^2 \geq Q^2 + U^2 + V^2, \quad (2.32)$$

where I, Q, U , and V are the normalized polarization parameters, often symbolized as

$$\bar{Q} = \frac{Q}{I}, \quad \bar{U} = \frac{U}{I}, \quad \text{and} \quad \bar{V} = \frac{V}{I}. \quad (2.33)$$

Similar to the polarization ellipse, the Stokes parameters may be depicted using the Poincaré sphere in spherical coordinates $(IP, 2\Psi, 2\chi)$, such that:

$$\begin{aligned} I &= S_0, \\ P &= \frac{\sqrt{S_1^2 + S_2^2 + S_3^2}}{S_0}, \text{ for } 0 \leq P \leq 1, \\ 2\Psi &= \arctan \frac{S_3}{\sqrt{S_1^2 + S_2^2}}, \text{ and} \\ 2\chi &= \arctan \frac{S_2}{S_1}, \end{aligned} \quad (2.34)$$

where I denotes the total intensity, P denotes the degree of polarization, or the ratio of polarized to non-polarized light in the wave-packet, χ denotes the polarization angle, and Ψ denotes the ellipticity angle of the polarization ellipse.



Figure 2.13: A diagram of an ideal polarimeter. Diagram adapted from Degl'Innocenti and Landolfi (2004).

2.2.2 Polarization Measurement

Except for polarimetry in the radio-wavelength regime, the polarization of a beam can not be directly measured. The polarization properties may, however, be recovered from the beam through the manipulation of the four parameters given in Equation 2.24. This so-called manipulation is achieved by passing the beam through optical elements which vary the beam for differing amplitudes and phases. These matrix operations may be represented by their corresponding Mueller matrices.

For ideal components, the resultant beam \mathbf{S}' after passing through an optical element is given by $\mathbf{S}' = \mathbf{MS}$, where \mathbf{S} is the beam incident on the optical element and \mathbf{M} represents the 4×4 Mueller matrix representing the optical element. Mueller matrices are especially useful when dealing with paths through optical elements as they observe the ‘train’ property (Priebe, 1969). This means that an incoming beam \mathbf{S} passing, in order, through elements with known Mueller matrices ($\mathbf{M}_0, \dots, \mathbf{M}_N$) results in an outgoing beam \mathbf{S}' such that:

$$\mathbf{S}' = \mathbf{M}_N \dots \mathbf{M}_0 \mathbf{S}. \quad (2.35)$$

Some Mueller Matrices are given below with angles related to those in Figure 2.13, measured counter-clockwise in a right-handed coordinate system.

General Rotation The Mueller matrix for coordinate space rotations about the origin by an angle θ ,

$$\mathbf{R}(\theta) = \begin{bmatrix} 1 & 0 & 0 & 0 \\ 0 & \cos 2\theta & \sin 2\theta & 0 \\ 0 & -\sin 2\theta & \cos 2\theta & 0 \\ 0 & 0 & 0 & 1 \end{bmatrix}. \quad (2.36)$$

General Linear Retardance The Mueller matrix for retardance where α is the angle between the incoming vector and fast axis, and δ is the retardance introduced by the retarder,

$$\mathbf{W}(\alpha, \delta) = \begin{bmatrix} 1 & 0 & 0 & 0 \\ 0 & \cos^2 2\alpha + \sin^2 2\alpha \cos \delta & \cos 2\alpha \sin 2\alpha(1 - \cos \delta) & \sin 2\alpha \sin \delta \\ 0 & \cos 2\alpha \sin 2\alpha(1 - \cos \delta) & \cos^2 2\alpha \cos \delta + \sin^2 2\alpha & -\cos 2\alpha \sin \delta \\ 0 & -\sin 2\alpha \sin \delta & \cos 2\alpha \sin \delta & \cos \delta \end{bmatrix}. \quad (2.37)$$

The retarder is often referred to by this retardance, e.g. if the retardance is $\delta = \pi$ or $\pi/2$, the retarder is referred to as a half- or quarter-wave plate, respectively.

General Linear Polarization The Mueller matrix for linear polarization where β is the angle between the incoming vector and transmission axis,

$$\mathbf{P}(\beta) = \frac{1}{2} \begin{bmatrix} 1 & \cos 2\beta & \sin 2\beta & 0 \\ \cos 2\beta & \cos^2 2\beta & \cos 2\beta \sin 2\beta & 0 \\ \sin 2\beta & \sin 2\beta \cos 2\beta & \sin^2 2\beta & 0 \\ 0 & 0 & 0 & 0 \end{bmatrix}. \quad (2.38)$$

These matrices in combination with Equation 2.35 allow us to describe how the incoming Stokes parameters would change when passing through the various optical elements. For a setup similar to Figure 2.13, the detected Stokes parameters can be described by:

$$\begin{aligned} S'(\alpha, \beta, \gamma) \propto \frac{1}{2} \{ & I + [Q \cos 2\alpha + U \sin 2\alpha] \cos(2\beta - 2\alpha) \\ & - [Q \sin 2\alpha + U \cos 2\alpha] \sin(2\beta - 2\alpha) \cos \gamma \\ & + V \sin(2\beta - 2\alpha) \sin \gamma \}, \end{aligned} \quad (2.39)$$

where the retardance angle, α , polarization angle, β , for a wave plate with a relative phase difference, γ , may be varied to acquire a system of equations that can be solved to retrieve the Stokes polarization parameters (Bagnulo et al., 2009).

Several or more frames taken under differing configurations may be used to reduce a system of equations to extract all four Stokes polarization parameters, but it is possible to extract the I , Q and U polarization parameters using only four frames, or two dual-beam frames, for well-chosen configurations and assuming ideal components. This ideal configuration varies the retarder angle such that $\Delta\alpha = \pi/8$ while keeping the polarizer stationary. More frames for additional retarder angles are advisable and often necessary, however, as they correct for any differences in sensitivity, such as may arise in a polarized flat field and which is further discussed in § 2.2.3 (Patat and Romaniello, 2006). From Equation 2.39 we see that the linear retarder element is the driving element of a polarizer as the first three Stokes parameters (S_{0-2} , or I , Q , and U) may be found by changing only the angle of retardance, α .

Wave Plates Wave plates, also commonly referred to as retarders, are generally made from optically transparent birefringent crystals. A wave plate has a fast and slow axis, which are perpendicular to one another and both perpendicular to an incident beam. Due to the birefringence of the wave plate medium, the phase velocity of the beam polarized

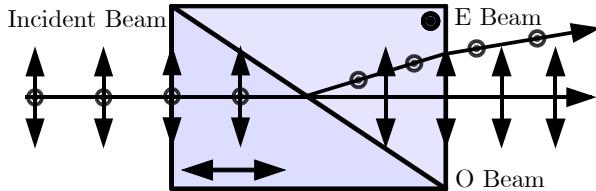


Figure 2.14: Diagram of a Rochon prism. Included in the diagram are the optical axes of the differing sections of the birefringent material as well as polarizing directions of the incident beam, denoted using the \leftrightarrow and \odot symbols, for the O - and E -beams, respectively. Figure adapted from ChrisHodgesUK, CC BY-SA 3.0, via Wikimedia Commons (2023).

parallel to the fast axis, namely the extraordinary beam, slightly increases while that of the beam polarized parallel to the slow axis, namely the ordinary beam, remains unaffected. This difference in the perpendicular component's phase velocities introduces a relative phase difference between the two beams, γ , which is given by:

$$\gamma = \frac{2\pi\Delta n L}{\lambda_0} \quad (2.40)$$

where Δn and L refer to the birefringence and thickness of the wave plate medium, respectively, and λ_0 refers to the vacuum wavelength of the beam (Hecht, 2017).

This relative phase difference determines the name of the wave plate, such that the $\gamma = m(\pi/2)$ and $\gamma = m(\pi/4)$ phase differences, for $m \in \mathbb{Z}^+$, refer to the half- and quarter-wave plates (which are the most common wave plate phases), respectively. Phase differences with an integer multiple of one another relate to the same phase difference and are referred to as multiple-order wave plates, while wave plates with a phase difference less than an integer multiple are referred to as zero-order wave plates. Several multiple-order wave plates can be combined by alternatively aligning the fast axis of one to the slow axis of another to create a compound zero-order wave plate (Hale and Day, 1988).

Polarizers Polarizers are typically made from two prisms, of a birefringent material, cemented together with an optically transparent adhesive. The actual effect of separating the perpendicular polarization components is achieved using varying effects, namely through:

- absorption of one of the polarized components, such as in Polaroid polarizing filters,
- total internal reflection of a single polarized component, such as in a Nicol prism (Figure 2.10),
- Refraction of a single polarized component, such as in a Rochon prism (Figure 2.14), or
- Refraction of both polarization components in differing directions, such as in a Wollaston prism (Figure 2.15).

Wollaston Prisms The Wollaston prism consists of two right-angle prisms consisting of a birefringent monoaxial material, cemented together with an optically transparent adhesive along their hypotenuses with their optical axes orthogonal, as seen in Figure 2.15. The Wollaston prism is a common optical polarizing element in astrophysical polarimetry which separates an incident beam into two linearly polarized O - and E -beams, orthogonal to one another, and deviated from their common axis equally. The deviation angle of the



Figure 2.15: Diagram of a Wollaston prism. Included in the diagram are the optical axes of the differing sections of the birefringent material as well as polarizing directions of the incident beam, denoted using the \leftrightarrow and $\downarrow\uparrow$ symbols, for the O - and E -beams, respectively. Diagram adapted from fgalore, CC BY-SA 3.0, via Wikimedia Commons (2023).

polarized beams is determined by the wedge angle which is defined as the angle from the common hypotenuse to that of the outer transmission face of either prism.

Wollaston prisms benefit over simpler elements (such as those listed in the polarizer paragraph) since a single frame allows for the observation of both orthogonal polarization components. This halves the observational time required to collect enough data to calculate the Stokes parameters, at the cost of an increase in calibration and reduction difficulty (Simon, 1986).

2.2.3 Polarimetric Calibrations

The raw science images acquired during polarimetric observations contain a combination of useful science data as well as noise. Corrections and calibrations related to the detector remain unchanged from those described in § 2.1.8, while those related to correcting for the optical elements relate to corrections for spurious polarization effects.

Flat Fielding

Once the CCD calibrations have been completed, the polarization intrinsic to the optical elements needs to be accounted for such that the pixel-to-pixel response is made uniform. Flat-fielding is, once again, used to correct for this. The flats taken for polarimetry, however, introduce an additional challenge as the targets for conventional flats are polarized, such as twilight and dome flats which are polarized by light scattering in the atmosphere and the reflective surface of the dome, respectively.

If no unpolarized flat images can be taken for flat field calibrations then, when possible due to the polarimeter design, the wave plate may be constantly rotated to act as a depolarizing element; this is effective so long as the wave plate rotation period is much faster than the flat's exposure time. Alternatively, polarized flats may be taken at the same set of half-wave plate angles used for science observations and averaged together to achieve a similar depolarizing effect.

Observing additional ‘redundant’ exposures for the science and flat images increases the depolarizing effect up to the maximum of 16 half-wave plate positions, where exposures with a half-wave plate angle differing by $\pi/4$ from another are considered redundant due to the O - and E -beams swapping between the related exposures.

Increasing the amount of redundant observations proportionally increases the time needed to observe all the exposures, which in turn introduces time-dependent effects such as fringing or intensity variations of the flat source. As such, a middle ground must be found for the amount of redundant frames observed. (Patat and Romanielo, 2006; Peinado et al., 2010).

Dual-Beam Extraction and Alignment

After calibrations for the CCD and light path are accounted for, the *O*- and *E*-beams can be extracted and further reduced. The extraction depends heavily on the layout of the polarimeter but often a simple cropping of the differing sections is enough to separate the two images.

After extracting the *O*- and *E*-beams for a specific half-wave plate angle, the images need to be aligned such that the sources present in them overlap. The Wollaston prism needs to be corrected for as it introduces a beam deviation which differs across both images. The aligning of the *O*- and *E*-beams is crucial as the comparison of the dual images is what allows for the calculation of the polarization properties.

Sky Subtraction

The polarization introduced by the sky introduces a difference in the intensity of the background sky and needs to be removed as it will influence the polarization results of the target source. Thankfully, the background polarization is an additive type of noise and may be subtracted out across the frames. This subtraction is done independently for both beams in a frame and for each frame since the background intensity of all observed polarimetric beams will differ based on the observational parameters.

2.3 Spectropolarimetry

As the name suggests, spectropolarimetry is the measurement of the polarization of light for a chosen spectral range and provides polarimetric results as a function of wavelength. As spectropolarimetry is so closely reliant on both spectroscopy and polarimetry, advancements in spectropolarimeters have always been gated by the advancements of spectrometers and polarimeters (as described in § 2.1 and § 2.2).

The most notable historical contributions of spectropolarimetry are those of spectropolarimetric studies instead of instrumental developments. Spectropolarimetry provides further insights into a materials physical structure, chemical composition, and magnetic field, allowing spectropolarimetry to be useful across multiple disciplines. In astronomy in particular, spectropolarimetry has been used to study the magnetic field, chemical composition, and underlying structure and emission processes of multiple types of celestial objects (see for example Antonucci and Miller, 1985; Donati et al., 1997; Wang and Wheeler, 2008).

Along with common points of consideration when developing any instrumentation for observational astronomy, such as resolution and sensitivity, spectropolarimeters need also consider the spectral response of the polarimetric components as well as the polarization



Figure 2.16: A spectropolarimetric target exposure as observed by the SALT RSS in spectropolarimetry mode.

response of the spectroscopic components as both are simultaneously in the light-path during observations and have noticeable affects on one another. Time is another constraint for spectropolarimetry as the incident light is separated both by wavelength and by polarization states. This division of the incident light results in increased exposure times for both target observations and observations necessary for calibrations.

Figure 2.16 illustrates a typical science image taken with a spectropolarimeter. The image contains the O - and E -beams which are both dispersed into their spectra. Spectropolarimetric results are acquired from measurements and calibrations of these images alongside any necessary calibration images.

2.3.1 Spectropolarimetric Measurement

The derived relations given in § 2.2.1, such as the Stokes parameters, describe polarization in general and are valid for both polarimetry and spectropolarimetry. Due to the time averaging of the observed light (Equation 2.28), any minor temporal variation, partial polarization, or monochromatic nature of the spectropolarimetric polarization parameters are accounted for.

For linear spectropolarimetry using a dual-beam polarizing element, an exposure measures the O - and E -beam wavelength dependent intensities, $f_{O,i}(\lambda)$ and $f_{E,i}(\lambda)$, for a given wave plate angle θ_i at angle i . These intensities thus relate to the wavelength dependent Stokes parameters as:

$$\begin{aligned} f_{O,i}(\lambda) &= \frac{1}{2}[I(\lambda) + Q(\lambda) \cos(4\theta_i) + U(\lambda) \sin(4\theta_i)], \text{ and} \\ f_{E,i}(\lambda) &= \frac{1}{2}[I(\lambda) - Q(\lambda) \cos(4\theta_i) - U(\lambda) \sin(4\theta_i)]. \end{aligned} \quad (2.41)$$

At least four linear equations are required to solve for three variables in a system of linear equations and thus at least two exposures must be taken to solve for the linear ($I(\lambda)$, $Q(\lambda)$, and $U(\lambda)$) polarization parameters (Degl'Innocenti et al., 2006; Keller, 2002).

The first Stokes parameter, $I(\lambda)$, may be recovered for each dual-beam exposure using

$$I_i(\lambda) = f_{O,i}(\lambda) + f_{E,i}(\lambda). \quad (2.42)$$

By calculating the $I_i(\lambda)$ Stokes parameter for each wave plate position i , the variation of the target over the course of observation may be corrected for, resulting in the $I(\lambda)$ Stokes parameter.

Next, the $Q(\lambda)$ and $U(\lambda)$ Stokes parameters are found by first defining the normalized difference in relative intensities, $F_i(\lambda)$, as:

$$F_i(\lambda) \equiv \frac{f_{O,i}(\lambda) - f_{E,i}(\lambda)}{f_{O,i}(\lambda) + f_{E,i}(\lambda)}, \quad (2.43)$$

which allows Equation 2.41 to be written, as

$$F_i(\lambda) = \bar{Q}(\lambda) \cos(4\theta_i) + \bar{U}(\lambda) \sin(4\theta_i) = P \cos(4\theta_i - 2\chi), \quad (2.44)$$

in terms of the normalized Stokes parameters, or, alternatively, the degree of polarization, P , and polarization angle, χ (as described in Equation 2.33 and 2.34).

The optimal change in wave plate angle is $\Delta\theta_i = \pi/8$ as it allows the normalized Stokes polarization parameters to be calculated as:

$$\begin{aligned} \bar{Q}(\lambda) &= \frac{2}{N} \sum_{i=0}^{N-1} F_i(\lambda) \cos\left(\frac{\pi}{2}i\right), \text{ and} \\ \bar{U}(\lambda) &= \frac{2}{N} \sum_{i=0}^{N-1} F_i(\lambda) \sin\left(\frac{\pi}{2}i\right), \end{aligned} \quad (2.45)$$

where N is the number of exposures taken, limited such that $N \in [2, 16]$ (Patat and Romaniello, 2006).

2.3.2 Spectropolarimetric Calibrations

Just as the elements of a spectropolarimeter are an amalgamation of both a spectrometer and polarimeter, it naturally follows that the calibrations necessary to reduce spectropolarimetric data are a combination of the calibrations needed for spectroscopy and polarimetry, discussed further in § 2.1.8 and § 2.2.3. Even though the spectrometer and polarimeter components both have an effect on an incident beam following the light-path through the spectropolarimeter, the calibration procedures for both methods remain mostly independent of one another and as such need not be repeated here.

Spectropolarimetric calibrations are, however, more involved when compared to the same calibrations for either spectroscopy or polarimetry. Minor deviations in the calibrations across both the spectra and the polarized beam compound, especially when dealing with the wavelength calibration, resulting in poor Signal-to-Noise Ratio (S/N)'s. Generally, more exposures over longer timespans are required to acquire enough redundancy and signal for the calculation of the Stokes parameters on top of the time necessary for calibrations to be completed. It should therefore be noted just how important the calibrations are when dealing with spectropolarimetry.



Figure 2.17: The tracker, supporting structure, and primary mirror of SALT. Figure adapted from the SALT call for proposals (2022).⁵

2.4 The Southern African Large Telescope

Southern African Large Telescope (SALT) is a 10 m class optical/near-infrared telescope situated at the South African Astronomical Observatory (SAAO) field station near Sutherland, South Africa (Burgh et al., 2003). The operational design was based on the Hobby-Eberly Telescope (HET) situated at McDonald Observatory, Texas, which limits the pointing of the telescope's primary mirror to a fixed elevation (37° from zenith in the case of SALT) while still allowing for full azimuthal rotation (Ramsey et al., 1998). Both SALT and HET utilize a spherical primary mirror which is stationary during observations and a tracker housing most of the instrumentation that tracks the primary mirrors spherically shaped focal path. Figure 2.17 depicts SALT's tracker (top left), supporting structure, and primary mirror (bottom right).

2.4.1 The Primary Mirror

The primary mirror is composed of 91 individual 1 m hexagonal mirrors which together form an 11 m segmented spherical mirror. Each mirror segment can be adjusted by actuators allowing the individual mirrors to approximate a single monolithic spherical mirror. The fixed elevation means that SALT's primary mirror has a fixed gravity vector allowing for a lighter, cost-effective supporting structure when compared to those of a more traditional altitude-azimuthal mount but with the trade-off that the control mechanism and tracking have increased complexity (Buckley et al., 2006).

⁵http://pysalt.salt.ac.za/proposal_calls/current/ProposalCall.html

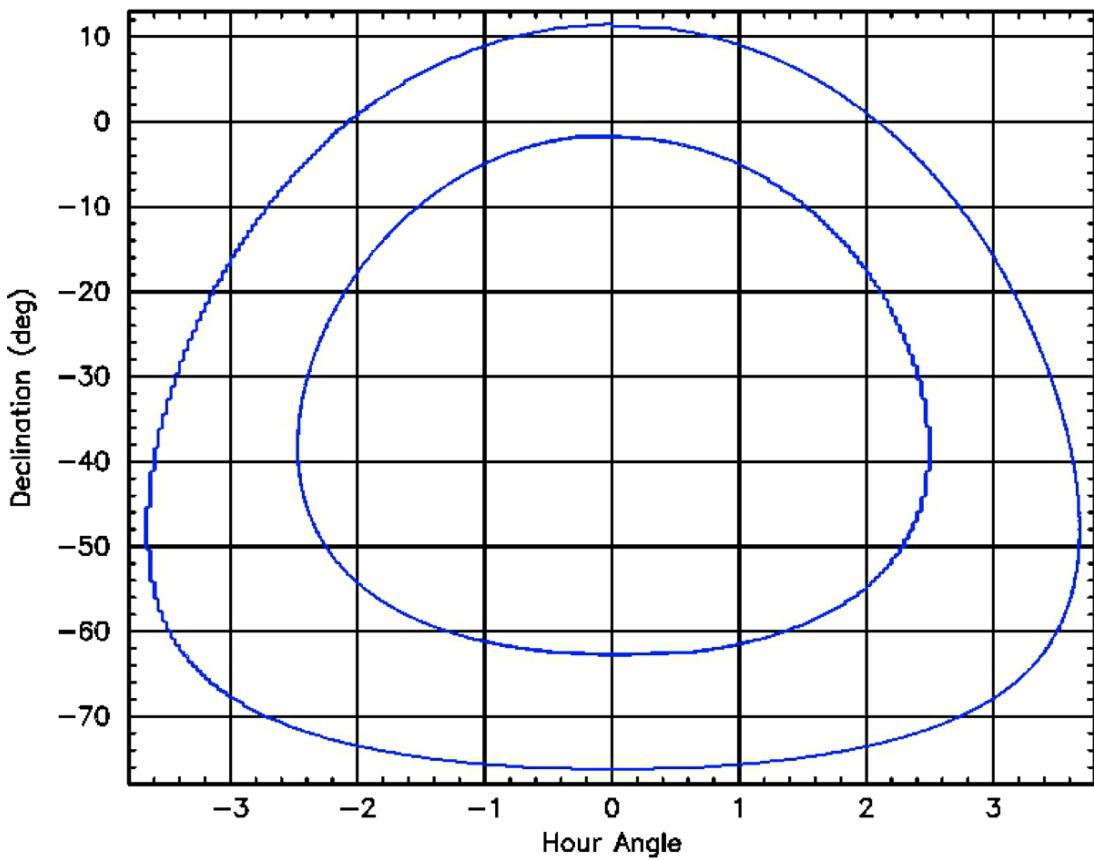


Figure 2.18: The visibility annulus of objects observable by SALT. Figure adapted from the SALT call for proposals (2013).⁶

2.4.2 Tracker and Tracking

During observations the primary mirror is stationary and the tracker tracks celestial objects across the sky by moving along the primary focus. The tracker is capable of 6 degrees of freedom with an accuracy of $\sim 5 \mu\text{m}$ and is capable of tracking $\pm 6^\circ$ from the optimal central track position. Targets at declinations from 10.5° to -75.3° , as shown in Figure 2.18 are accessible during windows of opportunity. As the tracker moves along the track the effective collecting area varies and thus SALT has a varying effective diameter of $\sim 7 \text{ m}$ to 9 m when the tracker is furthest and closest to the optimal central position, respectively.

The tracker is equipped with a Spherical Aberration Corrector (SAC) (O'Donoghue, 2000), and an Atmospheric Dispersion Compensator (ADC) (O'Donoghue, 2002), which corrects for the spherical aberration caused by the geometry of the primary mirror and allows access to wavelengths as short as 3200 \AA . These return a corrected flat focal plane with an $8'$ diameter field of view at prime focus on to the science instruments, with a $1'$ annulus around it used by the tracker in a closed-loop guidance system. The tracker also houses the calibration system which contains the Ar, CuAr, HgAr, Ne, ThAr, and Xe wavelength calibration lamps (Buckley et al., 2008).

⁶https://pysalt.salt.ac.za/proposal_calls/2013-2/



Figure 2.19: The optical path of the SALT RSS. Figure adapted from the SALT call for proposals (2023).⁷

2.4.3 SALT Instrumentation

SALT is equipped with the SALT Imaging Camera (SALTICAM) and the RSS science instruments onboard the tracker, and the High Resolution Spectrograph (HRS) and Near Infra-Red Washburn Labs Spectrograph (NIRWALS) science instruments which are fibre-fed from the tracker to their own climate controlled rooms. The RSS is currently the only instrument used for spectropolarimetry.

NIRWALS

The Near Infra-Red Washburn Labs Spectrograph (NIRWALS) is currently being commissioned and will have a wavelength coverage of 8000 to 17000 Å, providing medium resolution spectroscopy at $R = 2000$ to 5000 over Near Infra-Red (NIR) wavelengths (Brink et al., 2022; Wolf et al., 2022). NIRWALS is fibre-fed from its integral field unit, containing 212 object fibers, along with a separate sky bundle, containing 36 fibers, housed in the SALT fibre instrument feed. It is ideally suited for studies of nearby galaxies.

HRS

The High Resolution Spectrograph (HRS) echelle spectrograph was designed for high resolution spectroscopy at $R = 37000 - 67000$ covering a wavelength range of 3700 - 8900 Å and consists of a dichroic beam splitter and two VPH gratings (Nordsieck et al., 2003). This instrument is capable of stellar atmospheric and radial velocity analysis.

⁷https://pysalt.salt.ac.za/proposal_calls/current/ProposalCall.html

Grating Name	Wavelength Coverage (Å)	Usable Angles (°)	Bandpass per tilt (Å)	Resolving Power (1.25'' slit)
PG0300 ⁸	3700 – 9000		3900/4400	250 – 600
PG0700 ⁸	3200 – 9000	3.0 – 7.5	4000 – 3200	400 – 1200
PG0900	3200 – 9000	12 – 20	~ 3000	600 – 2000
PG1300	3900 – 9000	19 – 32	~ 2000	1000 – 3200
PG1800	4500 – 9000	28.5 – 50	1500 – 1000	2000 – 5500
PG2300	3800 – 7000	30.5 – 50	1000 – 800	2200 – 5500
PG3000	3200 – 5400	32 – 50	800 – 600	2200 – 5500

Table 2.1: Gratings available for use with the RSS. Table adapted from the SALT call for proposals (2023).

SALTICAM

The SALT Imaging Camera (SALTICAM) functions as the acquisition camera and simple science imager with various imaging modes, such as full-mode and slot-mode imaging, and supports low exposure times, down to 50 ms (O'Donoghue et al., 2006). This enables photometry of faint objects, especially at fast exposure times.

RSS

The Robert Stobie Spectrograph (RSS) functions as the primary spectrograph on SALT and can operate in long-slit spectroscopy and spectropolarimetry modes, a narrowband imaging mode, and multi-object and high resolution spectroscopy modes (for an in-depth discussion on operational modes see Kobulnicky et al., 2003, or the latest call for proposals).

The Detector The RSS detector consists of a mosaic of 3 CCD chips with a total pixel scale of 0.1267'' per unbinned pixel with varying readout times depending on the binning and readout mode. The mosaicking results in a characteristic double ‘gap’ in the frames and resultant spectra taken with the RSS, as seen in Figure 2.16.

The Available Gratings The RSS is equipped with a rotatable magazine of six VPH gratings, as listed in Table 2.1. Observations may be planned using simulator tools provided by SALT and are performed in the first order only. The RSS has a clear filter, as well as three Ultraviolet (UV) (with differing lower filtering ranges) and one blue order blocking filter available, used in conjunction with the various gratings to block out contamination from the second order.

RSS Spectropolarimetry Spectropolarimetry using the RSS is currently commissioned for long-slit linear spectropolarimetry, (I, Q, U), where observations are taken following the waveplate pattern lists as in Table 2.2. Circular, (I, V), and all-Stokes, (I, Q, U, V), spectropolarimetry modes are in commissioning with observations including redundant half-wave plate pairs to be commissioned thereafter.⁹

⁸The PG0300 surface relief grating has been replaced with the PG0700 VPH grating as of November 2022 but has been included here as observations using the PG0300 are used in later sections.

⁹Commission status sourced from the latest ‘Polarimetry Observers Guide’ (2024).

Linear ($^{\circ}$)		Linear-Hi ($^{\circ}$)		Circular ($^{\circ}$)		Circular-Hi ($^{\circ}$)		All Stokes ($^{\circ}$)	
$\frac{1}{2}$	$\frac{1}{4}$	$\frac{1}{2}$	$\frac{1}{4}$	$\frac{1}{2}$	$\frac{1}{4}$	$\frac{1}{2}$	$\frac{1}{4}$	$\frac{1}{2}$	$\frac{1}{4}$
0	-	0	-	0	45	0	45	0	0
45	-	45	-	0	-45	0	-45	45	0
22.5	-	22.5	-			22.5	-45	22.5	0
67.5	-	67.5	-			22.5	45	67.5	0
	-	11.25	-			45	45	0	45
	-	56.25	-			45	-45	0	-45
	-	33.75	-			67.5	-45		
		78.75	-			67.5	45		

Table 2.2: Spectropolarimetry waveplate patterns defined for the RSS. The stated angles refer to the angle of the half ($\frac{1}{2}$ -) and quarter ($\frac{1}{4}$ -) waveplate's optical axis from the perpendicular of the dispersion axis. Table adapted from the SALT call for proposals (2023).

Chapter 3

Existing and Developed Software: An Overview of POLSALT, IRAF, and STOPS

This chapter contains an overview of *Polarimetric reductions for SALT* (POLSALT) and the limitations faced during POLSALT wavelength calibrations (§ 3.1), a brief overview of the *Image Reduction and Analysis Facility* (IRAF) tasks relevant for spectropolarimetric wavelength calibrations (§ 3.2), and an overview of *Supplementary Tools for POLSALT Spectropolarimetry* (STOPS), the software developed to supplement the POLSALT reduction process (§ 3.3). Finally, a discussion of the updated reduction process, an example of which may be found in Appendix A, is included (§ 3.4).

3.1 POLSALT - *Polarimetric reductions for SALT*

The POLSALT (*Polarimetric reductions for SALT*) pipeline is the official reduction pipeline for spectropolarimetric data taken using the SALT RSS.¹ The newest version of the software, aptly named the ‘beta version’ (‘version’ 23 January 2020), was the version adapted in this study. It includes a GUI, depicted in Figure 3.1, which allows for limited interactivity during key steps in the reduction process.²

The steps that make up the POLSALT reduction pipeline include raw image reductions, wavelength calibrations, background subtraction and spectral extraction, raw Stokes calculations, final Stokes calculations, and visualization of the results. Accurate reductions at each step are crucial for accurate results and are thus briefly discussed below. Further details for the reduction process may be found at the POLSALT GitHub wiki.³

¹POLSALT is made freely available via the POLSALT GitHub repository, available at <https://github.com/saltastro/polsalt>. It is strongly advised to follow the wiki for installation instructions.

²Installation files and instructions for the ‘beta version’ utilizing the GUI are available at <http://www.sao.ac.za/~ejk/polsalt/code/> in a TAR GZIP file.

³The GitHub wiki for POLSALT is available at <https://github.com/saltastro/polsalt/wiki>.



Figure 3.1: The layout of the `POLSALT` Graphical User Interface (GUI), including the contents of the reduction steps accessible via the dropdown menu. Note that there is no trailing forward slash after the ‘Top level data directory’. Figure created from a local instance of the `POLSALT` GUI.

3.1.1 Raw Image Reductions

Raw image reductions are run via `imred.py` and apply the necessary basic reductions to the raw data before any calibrations are applied. These reductions include overscan subtractions, gain corrections, crosstalk corrections, and mosaicking as well as attaching the bad pixel maps and pixel variance information. Files with raw image reductions performed have “`mxgbp`” prepended to their names. As of February 2022, raw image reductions are automatically run for all RSS spectropolarimetric observations as part of the default SALT basic reduction pipeline that is run daily.

3.1.2 Wavelength Calibrations

Wavelength calibration and cosmic-ray rejection is performed via `specpolwavmap.py` and separately calibrates the *O*- and *E*-beams, based on the arc frames, and applies a simple cosmic-ray rejection for all science frames. This step is interactive and allows the user to individually fit wavelength calibration maps to each beam. The importance of an accurate correlation between both beams has been touched on previously (§ 2.3.2) and will be further discussed in § 3.1.8. The wavelength calibrated results are saved as an additional ‘`WAV`’ wavelength extension to each science FITS file, which are prefixed with a “`w`”, and the *O*- and *E*-beams of the extensions are split into their own sub-extensions.



Figure 3.2: The layout of the interactive POLSALT spectra extraction GUI after selecting the ‘update tilt correction and windows’ button along the bottom border of the window. Figure created from a local instance of the POLSALT GUI.

3.1.3 Spectral Extraction

Background subtraction and spectral extraction is run via `specpoleextract_dev.py` which corrects for the beam-splitter distortion and tilt, performs sky subtraction, and extracts a one dimensional wavelength dependent spectrum for each beam sub-extension. This step is interactive with Figure 3.2 showing the interactive window used for spectral extraction. The user, using the brightest trace in the image as a reference, defines regions which span the wavelength axis which define the background and trace regions for the sky subtraction and spectral extraction. Files with background and geometric corrections applied are saved with “c” prepended to their names and files which contain the extracted one dimensional spectrum have “e” further prepended to their names.

3.1.4 Raw Stokes Calculations

The raw Stokes calculations are performed via `specpolrawstokes_dev.py` and identify waveplate pairs for which the intensity, I , and a ‘raw Stokes’ signal, S , are calculated as:

$$I = \frac{1}{2}(O_1 + O_2 + E_1 + E_2), \text{ and} \quad (3.1)$$

$$S = \frac{1}{2} \left[\left(\frac{O_1 - O_2}{O_1 + O_2} \right) - \left(\frac{E_1 - E_2}{E_1 + E_2} \right) \right]. \quad (3.2)$$

The raw Stokes signal is calculated as the normalized difference of the O - and E -beams, for a waveplate pair, taken perpendicular to one another. The created files contain the raw Stokes information and use a very specific naming style; most notably the indexes of the related waveplate pairs, from Table 2.2, are included in the file names.



Figure 3.3: The layout of the interactive POLSALT visualization GUI after selecting the ‘Plot results’ button along the bottom border of the window. Figure created from a local instance of the POLSALT GUI.

3.1.5 Final Stokes Calculations

The Final Stokes calculations are performed via `specpolfinalstokes.py` and, using the waveplate pattern along with the raw Stokes signals, calibrates for the polarimetric zero-point and waveplate efficiency, and calculates the final Stokes parameters. Before the final Stokes calculations are performed, and if a sufficient number of redundant exposures were taken, the raw Stokes signals are culled to eliminate outlier signals which may arise from, for example, temporary atmospheric conditions affecting the signal. The culling is performed by comparing observation cycles against one another, comparing the deviation of the signal means which estimate the baseline systematic polarization fluctuations (due to imperfections in repeatability), and performing a χ^2 analysis to eliminate any statistical outliers.

3.1.6 Visualization

Plotting the results of the spectropolarimetric reduction process is done using `specpolview.py`, which generates a plot of the Intensity, Linear Polarization (%), and Equatorial Polarization Angle ($^\circ$) against a shared wavelength axis, as seen in Figure 3.4. This step is interactive allowing the user control over various options, such as the wavelength range, binning, etc., with the GUI shown in Figure 3.3.

3.1.7 Post-Processing Analysis

Generally, the plot of the spectropolarimetric results is the stopping point for most reduction procedures as it contains or creates the desired results. However, additional tools exist which may be used after the polarization reductions, and which are not represented in the GUI, namely, flux calibration and synthetic filtering.



Figure 3.4: A typical plot resulting from the reduction process. Figure adapted from (Cooper et al., 2022).

Flux-calibrations are performed via `specpolflux.py` and are only intended for shape corrections of the spectrum. Additionally, a flux database file must exist for the observed standard and must be included in the working science directory.

Synthetic filtering is calculated via `specpolfilter.py` and computes the synthetically filtered polarization results. Any wavelength dependent throughput filter curve may be synthesized when defined by the user, but a few pre-defined filter curves are available, namely: the SALTICAMs *U*, *B*, *V*, *R*, and *I* Johnson-Cousins filter curves.

3.1.8 POLSALT Limitations and the Need for Supplementary Tools

The creation of supplementary tools for POLSALT spectropolarimetric reductions stemmed from the limitations of the wavelength calibration process and a need to compare wavelength solutions across the perpendicular *O* and *E* polarization beams. The process of calibrating wavelength solutions using the POLSALT pipeline is time-consuming for the average user, and often results in unexpected program crashes when receiving erroneous inputs or key presses. Due to the time-consuming process of recalibrating the wavelength solutions it is not feasible to perform the wavelength calibrations time and time again for anything larger than a handful of observations. This is particularly true for observations performed with the SALT PG0300 grating as the sparse spectral features of the Ar arc lamp are not handled well by the POLSALT pipeline.

Since PG0300 provided the widest wavelength range and highest throughput, it was almost exclusively used for observations of flaring blazars, resulting in a large backlog of unanalyzed data. The only arc available for the PG0300 grating with a close enough articulation and grating angle ($\sim 10.68^\circ$ and $\sim 5.38^\circ$, respectively), was the Ar arc lamp which displays sparse spectral features with large gaps over the wavelength range at these grating and articulation angles (Figure 3.5). This often led the POLSALT pipeline to create inconsistent wavelength solutions, or to fail to create a wavelength solution altogether,



Figure 3.5: One of many Ar arc lamp spectra as provided by SALT for line identification. Plot adapted from SALT’s published Longslit Line Atlases (as of 2024), resized to fit within the document margins but otherwise unchanged.⁴

since minor deviations of identified spectral features resulted in large deviations in regions with no spectral features. To only further compound the difficulty of the wavelength calibrations, the spectrum of the Ar arc lamp contains a partial overlap of a higher order at longer wavelengths (§ 2.1.7, Equation 2.5).

The chosen solution to overcome the limitations of the wavelength calibration process was to use a well established wavelength calibration software which allowed for rapid recalibrations and provided a familiar interface. IRAF provides this familiar environment and reliability, in part thanks to its continued community development.

Unfortunately, IRAF is unable to natively parse the data structure implemented by POLSALT ‘as is’ and so the files must be restructured. This restructuring works both ways as once the IRAF reductions are complete the data structure must be restructured to match that of the POLSALT `wavelength calibration` output such that the reduction process may be completed in POLSALT.

3.2 IRAF - *Image Reduction and Analysis Facility*

The IRAF (*Image Reduction and Analysis Facility*) software is a collection of software designed specifically for the reduction and analysis of astronomical images and spectra (Tody, 1986, 1993). The software consists of many tasks which perform specific operations and which are grouped into relevant packages. Only a brief overview of the tasks will be provided here. Help documentation for any of the IRAF tasks may be found online or through the IRAF Command Line Interface (CLI) through the `?` or `:.help` ‘cursor commands’ when running interactive tasks, with more specific help documentation provided in the relevant section.⁵

Useful IRAF tasks that deserve a brief mention are: the `mkscript` task in the `system` package which allows a user to create and save a task along with the defined parameters

⁴The ‘low resolution’ Ar plot sourced from <https://astronomers.salt.ac.za/data/salt-longslit-line-atlas/>

⁵Online help for IRAF is available at <https://iraf.net/irafdocs/>

as a file which can later be called as a script,⁶ the `implot` task in the `plot` package which allows the rows or columns of an image to be interactively displayed,⁷ and the `eparam` task in the `language` package which allows the parameters of a task to be edited within the IRAF CLI.⁸

For the wavelength calibration of SALT spectropolarimetric data, the relevant tasks are the `identify` and `reidentify` tasks located in the `noao.onedspec` package, and the `fitcoords` and the (optional) `transform` tasks located under the `noao.twodspec.longslit` package. These tasks produce a two-dimensional wavelength solution which must be obtained separately for the *O*- and *E*-beam.

3.2.1 Identify

The `identify` task is used to interactively determine a one-dimensional wavelength function across a chosen row of an arc exposure by identifying features in the spectrum with known wavelengths.⁹ The task creates the first approximation of the wavelength solution (see Figure 3.6) as well as a local database in which the solution is saved (see Listing 3.1). The initial solution is built on in subsequent tasks, and it is, therefore, imperative that the initial solution is well-fit to minimize errors further along the calibration process.

The execution of `identify` consists of identifying known features spanning the entire wavelength range and then removing any features which negatively impact the wavelength solution. A balance must be found between the number of identified features, the parameters of the fit, and the deviation of the fit from the known features.

3.2.2 Reidentify

The `reidentify` task is used to run the `identify` task autonomously and repeatedly across the entirety of the arc frame at defined (row) intervals.¹¹ The task uses the one-dimensional wavelength solution stored in the database created by the initial `identify` call and refits the positions of the relevant spectral features. The task may fail based on a number of conditions, most common of which is the loss of features as the task moves further from the row at which the user manually ran `identify`.

⁶Help documentation for the `mkscript` task may be found at https://astro.uni-bonn.de/~sysstw/lfa_html/iraf/system.mkscript.html.

⁷Help documentation for the `implot` task may be found at https://astro.uni-bonn.de/~sysstw/lfa_html/iraf/plot.implot.html.

⁸Help documentation for the `eparam` task may be found at https://astro.uni-bonn.de/~sysstw/lfa_html/iraf/language.eparam.html.

⁹Help documentation for the `identify` task may be found at https://astro.uni-bonn.de/~sysstw/lfa_html/iraf/noao.onedspec.identify.html.

¹⁰See also <https://iraf.net/irafdocs/formats/identify.php> for an explanation of the database contents.

¹¹Help documentation for the `reidentify` task may be found at https://astro.uni-bonn.de/~sysstw/lfa_html/iraf/noao.onedspec.reidentify.html.

Listing 3.1: An example of the identify database contents.¹⁰

```
# Thu 15:19:16 13-May-2021
begin    identify arc00057[*,237]
  id arc00057
  task      identify
  image     arc00057[*,237]
  units    Angstroms
  features   35
      53.61 5944.74989  5944.834 13.0 1 1 15257
      140.19 6029.9793  6029.997 13.0 1 1 4652
      185.30 6074.34644  6074.338 13.0 1 1 13396
      207.49 6096.15873  6096.161 13.0 1 1 21700
      255.23 6143.0493  6143.063 13.0 1 1 33330
      276.13 6163.56995  6163.594 13.0 1 1 11344
      330.89 6217.29293  6217.281 13.0 1 1 13705
      381.10 6266.51524  6266.495 13.0 1 1 21747
      420.21 6304.8113  6304.789 13.0 1 1 10226
      450.49 6334.45415  6334.428 13.0 1 1 36235
      500.18 6383.04826  6382.991 13.0 1 1 35824
      519.85 6402.26802  6402.248 13.0 1 1 70163
      626.70 6506.56147  6506.528 13.0 1 1 46165
      653.73 6532.91083  6532.882 13.0 1 1 21413
      721.60 6598.98642  6598.953 13.0 1 1 26396
      803.21 6678.31069  6678.277 13.0 1 1 51338
      843.15 6717.0732  6717.043 13.0 1 1 36780
      1099.95 6965.36335  6965.431 13.0 1 1 5618.4 ar
      1169.57 7032.38598  7032.413 13.0 1 1 100000
      1317.05 7173.89814  7173.938 13.0 1 1 5000 decrease
      1391.52 7245.11148  7245.167 13.0 1 1 73545
      1537.20 7383.93022  7383.981 13.0 1 1 5557.5 ar
      1595.02 7438.83545  7438.898 13.0 1 1 15000 decrease
      1663.64 7503.86263  7503.869 13.0 1 1 30000 ar; increase
      1697.46 7535.84584  7535.774 13.0 1 1 8000 increase
      1802.64 7635.07335  7635.106 13.0 1 1 20000 ar; decrease
      2209.19 8014.79559  8014.786 13.0 1 1 3000 ar; decrease
      2604.58 8377.66137  8377.607 13.0 1 1 14543
      2734.54 8495.41423  8495.359 13.0 1 1 8765
      2763.48 8521.52355  8521.442 13.0 1 1 4537.5 ar
      2840.92 8591.20799  8591.258 13.0 1 1 2000 decrease
      2889.39 8634.67334  8634.647 13.0 1 1 3059
      2911.42 8654.39264  8654.383 13.0 1 1 3000 decrease
      2926.56 8667.93501  8667.944 13.0 1 1 702.5 ar
      3135.72 8853.77575  8853.867 13.0 1 1 1820

  function legendre
  order 4
  sample *
  naverage 1
  niterate 0
  low_reject 3.
  high_reject 3.
  grow 0.
  coefficients   8
      2.
      4.
      53.60757446289061
      3135.715576171875
      7425.420339270724
      1457.513831286474
      -26.15751926622308
      -3.000903509842187
```

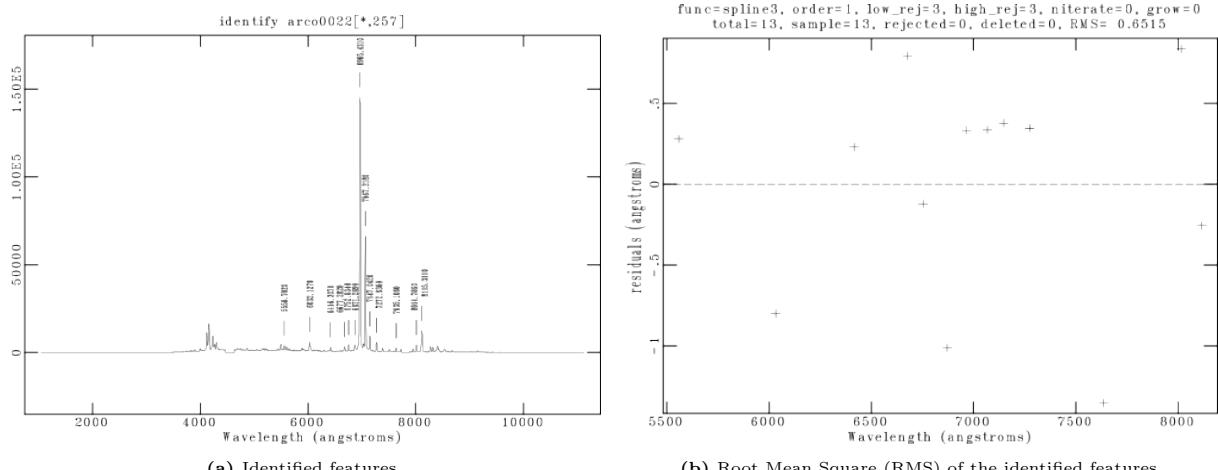


Figure 3.6: A plot and the RMS of the identified features found using the IRAF identify task. Figures created using the IRAF identify task.

Listing 3.2: An example of the `fitcoords` database contents.¹³

```
# Thu 15:26:55 13-May-2021
begin    arc00057
task      fitcoords
axis      1
units     angstroms
surface   33
      1.
      5.
      5.
      1.
      1.
      3199.
      1.
      474.
      7419.096745914063
      1510.03933621895
      -21.10886852752348
      -2.079553916887794
      0.06772631420528228
      0.7720164913117386
      -1.506773900054024
      0.1341878190232142
      -0.01659697703758917
      0.0251087019569153
      -3.318493303995171
      -0.3612632489821799
      0.003270665801371641
      -0.0157962041414068
      -0.003073690871589242
      0.007533453962924031
      0.02839687304474069
      -0.003233465769521899
      0.00174111456659807
      0.00645177595090841
      0.0105080093855621
      -0.01157827440314294
      -0.007789479002470706
      -0.006562085282926231
      -0.002321476801926803
```

When running `reidentify` non-interactively, it is recommended to set the `verbose` parameter to ‘yes’ as this will provide immediate confirmation if the task quit early. Regardless of whether the task quit successfully, the newly defined wavelength solutions are appended to the local database following the `identify` task database format, an example of which is given in Listing 3.1.

3.2.3 Fitcoords

The `fitcoords` task is used to find a two-dimensional surface function from the one-dimensional wavelength solutions found for specific rows in the previous steps.¹² The usage of `fitcoords` is similar to that of `identify` and consists of examining the distribution of identified points and eliminating any points that `reidentify` may have misidentified (see Figure 3.7).

By eliminating outliers with bad residuals and modifying the two-dimensional surface function’s type and degree, the overall error of the fit is decreased, aligning more closely to what the ‘true’ wavelength solution is. This surface function is the final two-dimensional wavelength solution for each two-dimensional spectrum. It is saved using the `fitcoords` database format, an example of which is given in Listing 3.2, as the list of parameters and function coefficients required to recreate the closest two-dimensional model. The IRAF wavelength solution is used by the `STOPS join` method to create the ‘WAV’ extension required by POLSALT, further described in § 3.3.2.

¹²Help documentation for the `fitcoords` task may be found at https://astro.uni-bonn.de/~sysstw/lfa_html/iraf/noao.twodspec.longslit.fitcoords.html.

¹³See also <https://iraf.net/irafdocs/formats/fitcoords.php> for an explanation of the database contents.



Figure 3.7: A plot and the RMS of the features identified across the exposure using the IRAF `fitcoords` task. Figures created using the IRAF `fitcoords` task.

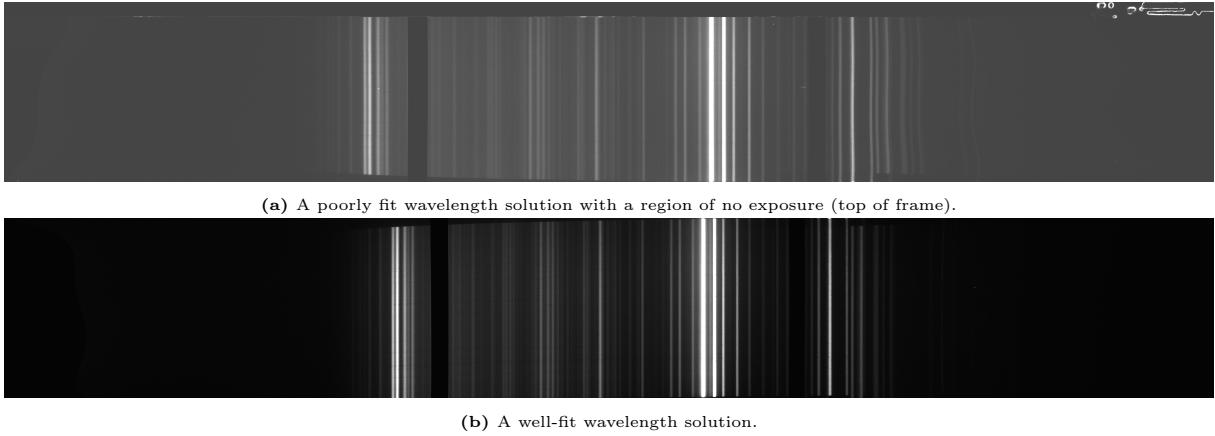


Figure 3.8: Examples of a poor fit (a) and well-fit (b) wavelength solution applied to the *O*- and *E*-beams of an arc image. The contrast of the figures were scaled to best capture any deviation of the arc lines. Figures created by the IRAF `transform` task.

3.2.4 Transform

The `transform` task is the optional final step in the IRAF wavelength calibration process.¹⁴ Simply put, `transform` converts the (x_p, y_p) pixel units of an exposure to (λ, y_p) wavelength units which allows for an immediate check of whether the wavelength solution is consistent across the frame. Any general error in the wavelength solution may be spotted in the transformed images; ranging from minor errors, such as the arc lines or sky lines not being purely vertical across the frame, to more major errors, such as an incorrect wavelength solution skewing the exposure beyond recognition.

For example, Figure 3.8a shows a good fit to the wavelength solution, as after transformation all the sky lines run exactly vertical. Figure 3.8b, on the other hand, shows a seemingly good fit, but closer inspection reveals that the sky lines (especially towards the right of the frame) deviate from the vertical, indicating a poor fit to the wavelength solution.

¹⁴Help documentation for the `transform` task may be found at https://astro.uni-bonn.de/~sysstw/lfa_html/iraf/noao.twodspec.longsit.transform.html.

3.3 STOPS - *Supplementary Tools for POLSALT Spectropolarimetry*

Supplementary Tools for POLSALT Spectropolarimetry (STOPS) provides supplementary tools which convert the POLSALT and IRAF formats back and forth, allowing IRAF to be used for wavelength calibrations of SALT spectropolarimetric data. It also provides additional tools to check the accuracy of the wavelength calibration. STOPS is written in, and requires, Python 3 (3.11+) to run, as well as **Astropy** (6.0.0+) (Astropy Collaboration et al., 2013, 2018, 2022), **ccdproc** (2.4.1+) (Craig et al., 2017), **Matplotlib** (3.5.2+) (Hunter, 2007), **NumPy** (1.26.4+) (Harris et al., 2020), and **SciPy** (1.13.0+) (Virtanen et al., 2020).

The parsing of POLSALT data into an IRAF usable format and the reformatting of the IRAF wavelength calibrated data back into a POLSALT usable format, referred to as *splitting* and *joining*, is performed by the STOPS **split** and **join** methods, respectively.

Methods to verify the validity of the wavelength calibrations were also added to STOPS. The **skyline** method checks the sky line wavelength (*x*) positions across the frame as well as the variation of the sky lines across the positional (*y*) axis of the frame. The **correlate** method checks the correlation of the *O*- and *E*-beams either within a given Flexible Image Transport System (FITS) file or across multiple files (comparing only the *O*- and *E*-beams for each). With these two additional methods, a user is able to verify that the wavelength solutions do not conflict across the *O*- and *E*-beams and that no unexpected deviations are included in the wavelength solutions.

Help on the usage of STOPS in a CLI can be viewed by running:

```
$ python ~/STOPS --help
# OR
$ python ~/STOPS [split|join|correlate|skylines] --help
```

which retrieves and prints the help documentation to the CLI from Listing B.1 (in Appendix B), such as how to enable logging or increase the verbosity, or change default parameters of the various methods. Finally, help documentation for the specific STOPS methods may be found within this section (Listing 3.3 to 3.6) or in Appendix B.

3.3.1 Splitting

As mentioned previously, the format of the FITS file created by POLSALT after basic CCD reductions and the format expected by IRAF to be used for the wavelength calibrations are incompatible. Basic POLSALT CCD reductions return FITS files which contain a primary header along with extensions for the science, variance, and ‘BPM’ images. These extensions carry the image of the trace (see Figure 3.9), the variance of the image, and a map of the pixels to be masked out, split into sub-extensions for both polarimetry beams, respectively.

While IRAF is capable of dealing with multiple traces in an extension or lists of input files, it is not as capable when dealing with multiple wavelength solutions contained in a single extension (as expected by the POLSALT **wavelength calibration**) or extensions containing sub-extensions (as expected by the POLSALT **spectral extraction**).

Listing 3.3: The ‘docstring’ for `split.py`

```

24 """
25 The `Split` class allows for the splitting of `polsalt` FITS files
26 based on the polarization beam. The FITS files must have basic
27 `polsalt` pre-reductions already applied (`mxbp...` FITS files).
28
29 Parameters
30 -----
31 data_dir : str / Path
32     The path to the data to be split
33 fits_list : list[str], optional
34     A list of pre-reduced `polsalt` FITS files to be split
35     within `data_dir`.
36     (The default is None, `Split` will search for `mxbp*.fits` files)
37 split_row : int, optional
38     The row along which to split the data of each extension in
39     the FITS file.
40     (The default is SPLIT_ROW (See Notes), the SALT RSS CCD's middle row)
41 no_arc : bool, optional
42     Decides whether the arc frames should be recombined.
43     (The default is False, `polsalt` has no use for the arc after
44     wavelength calibrations)
45 save_prefix : dict[str, list[str]], optional
46     The prefix with which to save the O & E beams.
47     Setting `save_prefix` = ``None`` does not save the split O & E beams.
48     (The default is SAVE_PREFIX (See Notes))
49
50 Attributes
51 -----
52 arc : str
53     Name of arc FITS file within `data_dir`.
54     `arc` = ```` if `no_arc` or not detected in `data_dir` .
55 o_files, e_files : list[str]
56     A list of the `O`- and `E`-beam FITS file names.
57     The first entry is the arc file if `arc` defined.
58 data_dir
59 fits_list
60 split_row
61 save_prefix
62
63 Methods
64 -----
65 split_file(file: os.PathLike)
66     -> tuple[astropy.io.fits.HDUList]
67         Handles creation and saving the separated FITS files
68 split_ext(hdulist: astropy.io.fits.HDUList, ext: str = 'SCI')
69     -> astropy.io.fits.HDUList
70         Splits the data in the `ext` extension along the `split_row`
71 crop_file(hdulist: astropy.io.fits.HDUList, crop: int = CROP_DEFAULT)
72     -> tuple[np.ndarray]
73         Crops the data along the edge of the frame, that is,
74         `O`-beam cropped as [crop:], and
75         `E`-beam cropped as [-crop].
76 update_beam_lists(o_name: str, e_name: str)
77     -> None
78         Updates `o_files` and `e_files`.
79 save_beam_lists(file_suffix: str = 'frames')
80     -> None
81         Creates (Overwrites if exists) and writes the `o_files` and `e_files`
82         to files named `o_{file_suffix}` and `e_{file_suffix}`, respectively.
83 process()
84     -> None
85         Calls `split_file` and `save_beam_lists` on each file in `fits_list`
86         for automation.
87
88 Other Parameters
89 -----
90 **kwargs : dict
91     keyword arguments. Allows for passing unpacked dictionary to
92     the class constructor.
93
94 Notes
95 -----
96 Constants Imported (See utils.Constants):
97     SAVE_PREFIX
98     CROP_DEFAULT
99     SPLIT_ROW
100 """

```

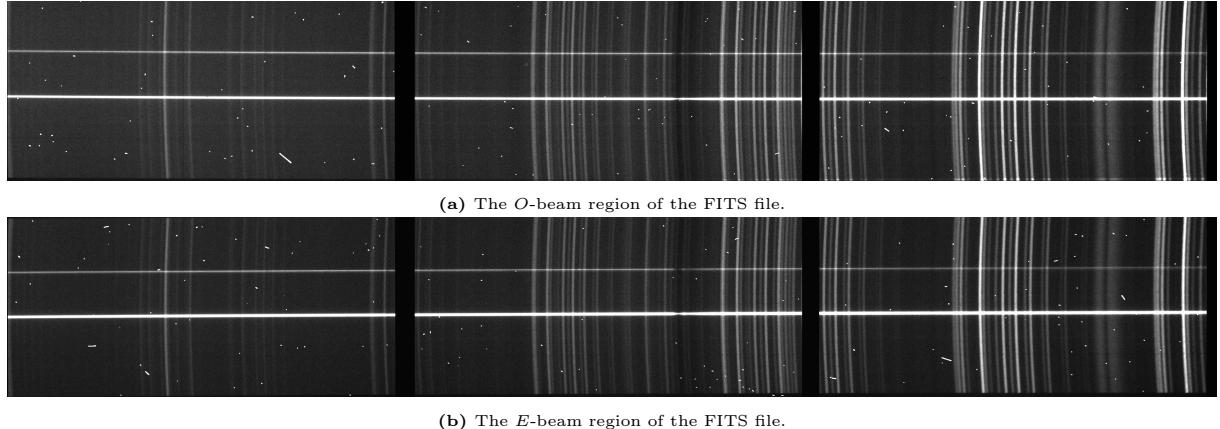


Figure 3.9: The split *O*- (a) and *E*-beams (b) as handed to **IRAF**. Figure created from the **STOPS** **split** method output.

To simplify the **IRAF** reduction procedure it was decided to separate the perpendicular polarization beams into their own files.

The files with POLSALT pre-reductions applied, namely FITS files with an ‘mxgbp’ prefix (§ 3.1), are used as the starting point for the supplementary tool’s **split** method. Running **split** finds all the FITS files for wavelength calibration within the working directory, creates two empty Header and Data Unit (HDU) structures for each FITS file (i.e. for both the *O*- and *E*-beam), and appends all header and science data necessary for wavelength calibration to the relevant HDU structure. Otherwise, defaults, such as which row to split the image along to separate the beams, were kept as close to the POLSALT pipeline as possible.

As the intent was always to parse the wavelength function back into POLSALT it was decided to keep these temporary FITS files as small as possible by only including the header and ‘SCI’ science extension. To aid the scripting of the **IRAF** wavelength calibration process, the **split** method also performs row cropping to exclude CCD regions which are not exposed to light, and creates files listing the split *O*- and *E*-beam FITS files which may be passed to the **IRAF** task inputs. Row cropping was decided on as **IRAF** does not handle rows with no exposure well, specifically when it comes to the autonomous **reidentify** task. The full **STOPS** **split** class docstring is given in Listing 3.3.

3.3.2 Joining

After the **IRAF** **fitcoords** task has been successfully run for both the *O*- and *E*-beams, the **STOPS** **join** method is used to extract and parse the wavelength solution from the **IRAF** database, and to create the wavelength calibrated FITS files required by the POLSALT pipeline. More specifically, the **join** method performs the following steps:

First, **join** parses the wavelength database file, described in § 3.2.3, for either a ‘Chebyshev’ or ‘Legendre’ solution, and calculates the wavelength at each (x_p, y_p) pixel position. This new image containing the corresponding wavelength values, seen in Figure 3.10, is appended to the wavelength calibrated FITS file as the ‘WAV’ extension.

Listing 3.4: The ‘docstring’ for join.py

```

32 """
33 The `Join` class allows for the joining of previously split files
34 and the appending of an external wavelength solution in the
35 `polsalt` FITS file format.
36
37 Parameters
38 -----
39 data_dir : str / Path
40     The path to the data to be joined
41 database : str, optional
42     The name of the `IRAF` database folder.
43     (The default is "database")
44 fits_list : list[str], optional
45     A list of pre-reduced `polsalt` FITS files to be joined
46     within `data_dir`.
47     (The default is ``None``, `Join` will search for `magbp*.fits` files)
48 solutions_list: list[str], optional
49     A list of solution filenames from which the wavelength solution
50     is created.
51     (The default is ``None``, `Join` will search for `fc*` files within
52     the `database` directory)
53 split_row : int, optional
54     The row along which the data of each extension in the FITS file
55     was split.
56     Necessary when Joining cropped files.
57     (The default is 517, the SALT RSS CCD's middle row)
58 save_prefix : dict[str, list[str]], optional
59     The prefix with which the previously split 'O'- & 'E'-beams were saved.
60     Used for detecting if cropping was applied during
61     the splitting procedure.
62     (The default is SAVE_PREFIX (See Notes))
63 verbose : int, optional
64     The level of verbosity to use for the Cosmic ray rejection
65     (The default is 30, I.E. logging.INFO)
66
67 Attributes
68 -----
69 fc_files : list[str]
70     Valid solutions found from `solutions_list`.
71 custom : bool
72     Internal flag for whether `solutions_list` uses the `IRAF`
73     or a custom format.
74     See Notes for custom solution formatting.
75     (Default (inherited from `solutions_list`) is False)
76 arc : str
77     Deprecated. Name of arc FITS file within `data_dir`.
78 data_dir
79 database
80 fits_list
81 split_row
82 save_prefix
83
84 Methods
85 -----
86 get_solutions(wavlist: list / None, prefix: str = "fc") ->
87     (fc_files, custom): tuple[list[str], bool]
88     Parse `solutions_list` and return valid solution files and if they are
89     non-`IRAF` solutions.
90 parse_solution(fc_file: str, x_shape: int, y_shape: int) ->
91     tuple[dict[str, int], np.ndarray]
92     Loads the wavelength solution file and parses keywords necessary for
93     creating the wavelength extension.
94 join_file(file: os.PathLike) -> None
95     Joins the files,
96     attaches the wavelength solutions,
97     performs cosmic ray cleaning,
98     masks the extension,
99     and checks cropping performed in `Split`.
100    Writes the FITS file in a `polsalt` valid format.
101 check_crop(hdu: pyfits.HDUList, o_file: str, e_file: str) -> int
102     Opens the split 'O'- and 'E'-beam FITS files and returns the amount of
103     cropping that was performed.
104 process() -> None
105     Calls `join_file` on each file in `fits_list` for automation.
106
107 Other Parameters
108 -----
109 no_arc : bool, optional
110     Deprecated. Decides whether the arc frames should be processed.
111     (The default is False, `polsalt` has no use for the arc after
112     wavelength calibrations)
113 **kwargs : dict
114     keyword arguments. Allows for passing unpacked dictionary to the
115     class constructor.
116
117 Notes
118 -----
119 Constants Imported (See utils.Constants):
120     DATADIR, SAVE_PREFIX, SPLIT_ROW, CR_PARAMS
121
122 Custom wavelength solutions must be formatted containing:
123     'x', 'y', *coefficients...
124 where a solution are of order ('x' by 'y') and
125 must contain x*y coefficients, all separated by newlines.
126 The name of the custom wavelength solution file must contain
127 either "cheb" or "leg" for Chebyshev or Legendre wavelength solutions,
128 respectively.
129
130 Cosmic ray rejection is performed using lacosmic [1]-
131 implemented in ccproc via astroscrappy [2]_.
132
133 References
134 -----
135 .. [1] van Dokkum 2001, PASP, 113, 789, 1420
136     (article : https://adsabs.harvard.edu/abs/2001PASP..113.1420V)
137

```

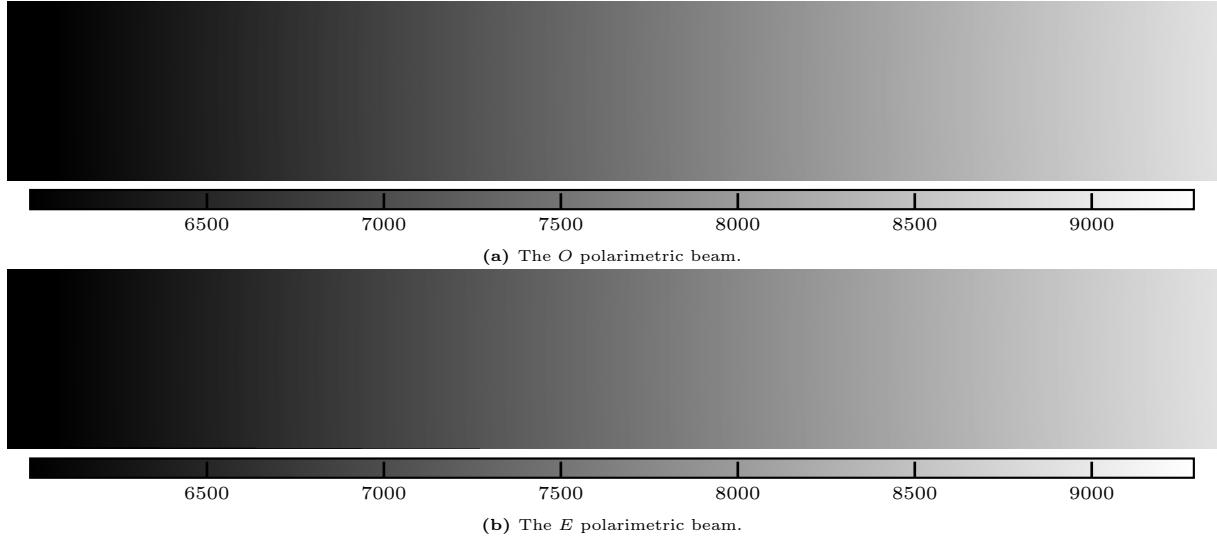


Figure 3.10: A representative ‘WAV’ extension of a FITS file, for the *O*- (a) and *E*-beam (b), ready to be processed by the POLSALT pipeline. The color bars show the wavelength in Å. Note that regions which fall outside the exposed region are masked by setting the corresponding pixel values of the wavelength and ‘BPM’ extensions to 0. Figure created from the `STOPS join` method output.

Second, cosmic-ray cleaning is performed on the ‘SCI’ extension using the `ccdproc` implementation of the `lacosmic` Python package which implements the L.A. Cosmic algorithm, based on Laplacian edge detection. The read noise and gain parameters used for cosmic ray cleaning were chosen based on the properties of the RSS, while the rest of the parameters were left as the default, following the publication and suggestions¹⁵ by the algorithm’s creator, as well as the implementation of the algorithm in the python `ccdproc` package (McCully et al., 2018; van Dokkum, 2001). The chosen parameters work well for most cosmic rays, as can be seen when comparing Figure 3.9 to Figure 3.11, but may be modified as needed.

Next, `join` updates the headers to be near-identical to those created by the POLSALT `wavelength calibration`, most notably updating the data shape, ‘CTYPE3’, and data type, ‘BITPIX’, keywords. The only difference in the header is the ‘NAXIS2’ keyword, due to the cropping performed by `split`. The cropped region could be reintroduced but would be masked out and further discarded in the following POLSALT `spectra extraction` process, making it redundant.

Finally, the ‘WAV’ extension is masked to remove any uncalibrated wavelength regions as well as masked for the skewing of the trace introduced by the wollaston element. The masking of the wollaston skewing is necessary since POLSALT introduces a wollaston correction in the `spectra extraction` process. The ‘BPM’ extension is masked to reflect the valid wavelength calibrated region, and the files are saved with the POLSALT wavelength calibrated ‘wmxgbp’ prefix. The full STOPS `join` class docstring is given in Listing 3.4.

Listing 3.5: The ‘docstring’ for `skylines.py`

```

40 """
41     """
42     Class representing the Skylines object.
43
44     Parameters
45     -----
46     data_dir : Path
47         The directory containing the data files.
48     filenames : list[str]
49         The list of filenames to be processed.
50     beam : str, optional
51         The beam mode, by default "OE".
52     plot : bool, optional
53         Flag indicating whether to plot the continuum, by default False.
54     save_prefix : Path / None, optional
55         The prefix for saving the data, by default None.
56     **kwargs
57         Additional keyword arguments.
58
59     Attributes
60     -----
61     data_dir : Path
62         The directory containing the data files.
63     fits_list : list[str]
64         The list of fits file paths.
65     beams : str
66         The beam mode.
67     can_plot : bool
68         Flag indicating whether to plot the continuum.
69     save_prefix : Path / None
70         The prefix for saving the data.
71     wav_unit : str
72         The unit of wavelength.
73
74     Methods
75     -----
76     checkLoad(self, path: str) -> np.ndarray:
77         Checks and loads the data from the given path.
78     transform(self, wav_sol: np.ndarray, spec: np.ndarray) -> np.ndarray:
79         Transforms the input wavelength and spectral data based on
80         the given wavelength solution.
81     rmvCont(self) -> np.ndarray:
82         Removes the continuum from the spectrum.
83     process(self) -> None:
84         Placeholder method for processing the data.
85
86     See Also
87     -----
88     matplotlib custom style:
89         https://matplotlib.org/stable/users/explain/customizing.html
90
91     Notes
92     -----
93     Constants imported (see utils.Constants):
94         SAVE_SKY:
95             The default save name for the skyline plot.
96         FIND_PEAK_PARAMS:
97             The default parameters for finding peaks in the spectrum.
98 """

```



Figure 3.11: A representative ‘SCI’ extension of a FITS file, for the *O*- (a) and *E*-beam (b), ready to be processed by the POLSALT pipeline. The observed intensity is displayed via the grayscale value at each pixel. Figure created from the `STOPSjoin` method output.

3.3.3 Sky Line Checks

The `skyline` method has been implemented to compare the position of the sky lines on the ‘SCI’ extension, or arc lines in the arc exposure, to the known positions of the sky lines, or arc lines, as measured by SALT, respectively.¹⁶ This provides an additional check of the accuracy of the wavelength solution across the frame. This method accepts both the IRAF `transform` FITS file or the ‘wmxgbp’ FITS files created by the `join` method as the input.

The `skyline` method loads the wavelength calibrated files, masks the traces present in the frames, transforms the frames from (x_p, y_p) pixel to $(\text{\AA}, y_p)$ wavelength units if the frame was not transformed by `transform`,¹⁷ and compares the peak wavelength position of the sky lines to the reference sky lines as measured by SALT. Finally, a figure is created containing a plot of the background spectra (offset by the respective legend entries superscript) with the known sky lines marked with a vertical line, and a plot of the closest identified peaks of said spectra.

Minor variations in the comparison of the sky lines are expected, such as those seen in Figure 3.12, but any uniform trends, such as those in Figure 3.13 (bottom left), indicate an underlying poor fit across the horizontal axis of the wavelength solution. The full STOPS `skyline` class docstring is given in Listing 3.5.

3.3.4 Cross Correlation

The `skyline` method allows for confirmation of a single wavelength solution, but has no means for comparing how the wavelength solutions of two polarization beams differ from each other. As the Stokes results, and thus final polarization results, are determined by the difference between the *O*- and *E*-beams, a direct comparison is not appropriate.

¹⁵Suggested parameters for the `lacosmic` algorithm may be found at <http://www.astro.yale.edu/dokkum/lacosmic/pars.html>.

¹⁶Both sky and arc lines are available at <https://astronomers.salt.ac.za/data/salt-longslit-line-atlas/>.

¹⁷The transformation applied by the `skyline` method uses linear interpolation and is thus less accurate at flux conservation than the transformation applied by the `transform` method.

Listing 3.6: The ‘docstring’ for cross_correlate.py

```

34
35 """
36 Cross correlate allows for comparing the extensions of multiple
37 FITS files, or comparing the $0$- and $E$-beams of a single FITS file.
38
39 Parameters
40 -----
41 data_dir : str / Path
42     The path to the data to be cross correlated
43 filenames : list[str]
44     The ecumzgbp*.fits files to be cross correlated.
45     If only one filename is defined, correlation is done against the two
46     polarization beams.
47 split_ccd : bool, optional
48     Decides whether the CCD regions should each be individually
49     cross correlated.
50     (The default is True, which splits the spectrum up into its separate
51     CCD regions)
52 cont_ord : int, optional
53     The degree of a chebyshev to fit to the continuum.
54     (The default is 11)
55 plot : bool, optional
56     Decides whether the continuum fitting should be plotted
57     (The default is False, so no continua plots are displayed)
58 save_prefix : str, optional
59     The name or directory to save the figure produced to.
60     "." saves a default name to the current working. A default name is
61     also used when save_prefix is a directory.
62     (The default is None, I.E. The figure is not saved, only displayed)
63
64 Attributes
65 -----
66 data_dir
67 fits_list
68 beams : str
69     The mode of correlation.
70     'OE' for same file, and 'O' or 'E' for different files but same
71     extension.
72 ccds : int
73     The number of CCD's in the data.
74     Used to split the CCD's if split_ccd is True.
75 cont_ord : int
76     The degree of the chebyshev to fit to the continuum.
77 can_plot : bool
78     Decides whether the continuum fitting should be plotted
79 offset : int, DEPRECATED
80     The amount the spectrum is shifted, mainly to test the effect of the
81     cross correlation.
82     (The default is 0, I.E. no offset introduced)
83 save_prefix
84 wav_unit : str
85     The units of the wavelength axis.
86     (The default is Angstroms)
87 wav_cdelet : int
88     The wavelength increment.
89     (The default is 1)
90 alt : Callable
91     An alternate method of cross correlating the data.
92     (The default is None)
93
94 Methods
95 -----
96 load_file(filename: Path) -> tuple[np.ndarray, np.ndarray, np.ndarray]
97     Loads the data from a FITS file.
98 get_bounds(bpm: np.ndarray) -> np.ndarray
99     Finds the bounds for the CCD regions.
100 remove_cont(spec: list, wav: list, bpm: list, plot_cont: bool) -> None
101     Removes the continuum from the data.
102 correlate(filename1: Path, filename2: Path / None = None) -> None
103     Cross correlates the data.
104 ftcs(filename1: Path, filename2: Path / None = None) -> None
105     Cross correlates the data using the Fourier Transform.
106 plot(spec, wav, bpm, corrdb, lagsdb) -> None
107     Plots the data.
108 process() -> None
109     Processes the data.
110
111 Other Parameters
112 -----
113 offset : int, optional
114     The amount the spectrum is shifted, mainly to test the effect of the
115     cross correlation.
116     (The default is 0, I.E. no offset introduced)
117 **kwargs : dict
118     keyword arguments.
119     Allows for passing unpacked dictionary to the class constructor.
120     ftcs : bool, optional
121     Boolean whether to use Fourier Transform for cross correlation.
122
123 See Also
124 -----
125 scipy:
126     https://docs.scipy.org/doc/scipy/reference/generated/
127         correlation:
128             scipy.signal.correlate.html
129 matplotlib custom style:
130     https://matplotlib.org/stable/users/explain/customizing.html
131
132 Notes
133 -----
134 Constants Imported (See utils.Constants):
135     SAVE_CORR:
136         The default save name for the correlation plot.
137     OFFSET:
138         The vertical offset of spectra in the output plot.
139 """

```



Figure 3.12: An example of a well calibrated wavelength solution, specifically shown for science images. Figure created via the `STOPS skyline` method.

Any observed unpolarized light, however, will reflect equally in both polarization beams and so the general trend of the two spectra may reasonably be expected to follow one another. The `correlate` method was created to allow for comparisons between the wavelength solutions of the *O*- and *E*-beams of a single exposure or the *O*- or *E*-beams of differing exposures by cross correlating the spectra.

The `correlate` method loads the `POLSALT spectra extraction` FITS files, removes the continuum and separates the CCD regions. The relevant CCD regions are cross correlated and the correlation peak is plotted and specified in the plot legend, as seen in Figure 3.14.

Sources under spectropolarimetric observation are generally expected to vary over time and, as such, the ratio of polarized to unpolarized light is also expected to vary. The accuracy of correlation may decrease as features with differences in the polarized component of the polarization beams change, and it is up to the user to determine the validity of the correlation result taking into consideration the two spectra being correlated. The differences in the features of the different spectra are often negligible when compared to the overall continuum of the spectra and are generally only reflected in a change in the intensity of said features when the continuum is removed. The full `STOPS split` class docstring is given in Listing 3.6.

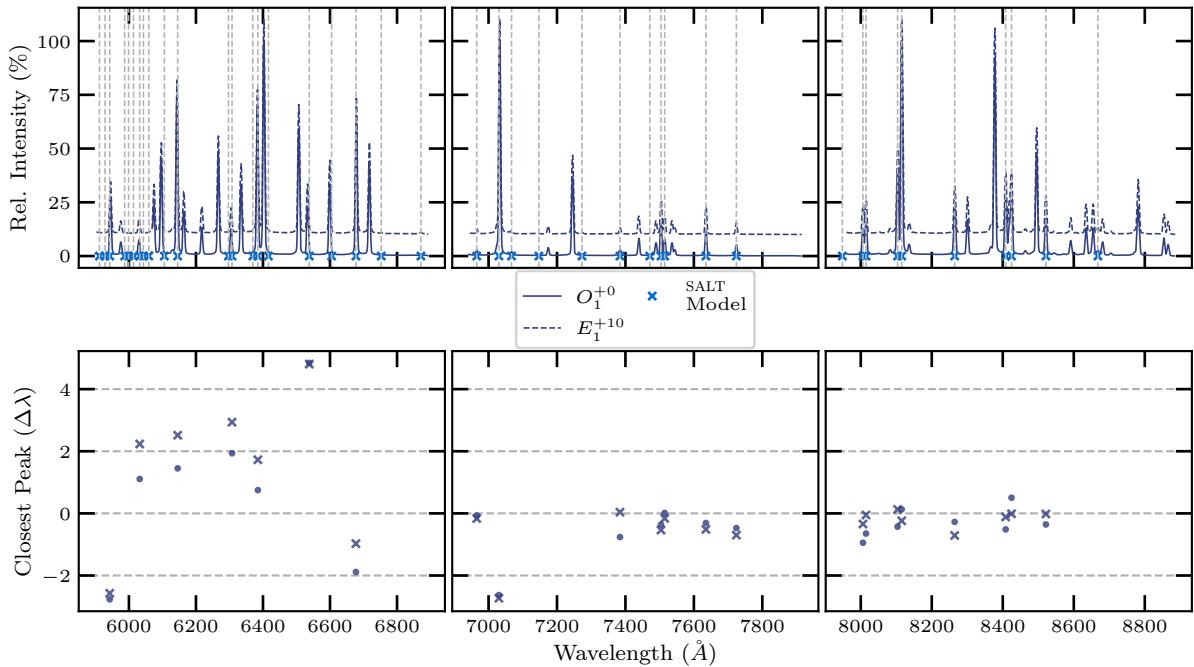


Figure 3.13: An example of a wavelength solution with a poor fit at shorter wavelengths (bottom left), specifically shown for an arc image. Arc lines are Figure created via the `STOPS` `skyline` method.

3.4 General Reduction Procedure

This section aims to provide a comprehensive discussion of the modified reduction procedure, an example of which is provided in Appendix A. As users all employ a variety of operating systems, language environments, and software setups, not much emphasis will be placed on how to get the software running or the managing of files; instead, the general order, seen in Figure 3.15, and commands necessary to complete each step of the reduction process are discussed, assuming that the software is running as intended.

3.4.1 Initial Setup

It is important to note that while POLSALT was developed in Python 2 (2.7), the STOPS supplementary tools were developed in, and require, Python 3 (3.11+), as well as the other requirements mentioned in § 3.3. While managing multiple versions of Python introduces some extra complication, it would not have been reasonable to develop STOPS in Python 2, as it has been deprecated, nor would it have been reasonable to update POLSALT to Python 3.

It is therefore recommended that the different versions of Python are managed using separate virtual environments. While the `anaconda` package manager was used in this study and is recommended, any package manager may be used. The `anaconda` environments are aliased ‘`salt`’ for Python 2.7 and ‘`stops`’ for Python 3.11. When Listings are provided (see for example Listing A.1 or the Listings in § 3.4.2 below), the `anaconda` environment is activated at the start of the Listing, otherwise it is assumed the previously specified environment is still active.



Figure 3.14: The resultant output plot of the STOPS `correlate` method. Figure created via the STOPS `correlate` method.

It is recommended to use POLSALT through the GUI as it provides a user-friendly environment while also sequentially listing each step of the reduction process in a dropdown menu, as seen in Figure 3.1. Reductions are possible, however, purely through the CLI using the POLSALT ‘beta’ scripts.

3.4.2 POLSALT Pre-Reductions

The POLSALT reduction process requires a file structure such that the raw data received from SALT is located in a folder named using the observing date with a sub-folder named `raw`, following the format `YYYYMMDD/raw/`. This directory structure allows POLSALT to create a ‘working’ directory following the format `YYYYMMDD/sci/` which contains all the files modified during the reduction process. Multiple reduction procedures using the same data may therefore be separated by simply renaming the `sci/` sub-folder.

The POLSALT GUI may be launched by opening a CLI and running the commands given in Listing A.1. Once the window, depicted in Figure 3.1, has launched, ensure that the first two paths at the top of the window point to the POLSALT and working directories, as seen in Figure 3.1. The ‘raw image reduction’ entry may then be selected from the dropdown menu and the pre-reductions run.

Alternatively, if the data already includes ‘mxgbp’ FITS files in the `YYYYMMDD/sci/` working directory, a CLI may be used to complete the initial pre-reductions using

```
$ cd <OBSDATE>/sci
$ conda activate salt
$ python ~/polsalt/scripts/reducepoldata_sc.py <OBSDATE>
```

which will start the full POLSALT reduction process. This process is quit once the POLSALT `wavelength calibration` GUI opens, and the alternate wavelength calibration procedure is followed.



Figure 3.15: A general workflow for data reductions using a combination of POLSALT, IRAF, and STOPS. Diagram adapted from Cooper et al. (2022).

3.4.3 Wavelength Calibration

The wavelength calibrations may now be completed in IRAF. This section concerns the procedure for parsing the FITS files to and from both IRAF and POLSALT, as well as the relevant task names and methods to be run to complete the calibrations. A base working case of each of the tasks and methods are presented in Listing A.2 to A.8, but it should be noted that the art of wavelength calibration consists of modifying the parameters to achieve a well-fit calibration function.

Preparing the Data for IRAF

Splitting the data is presented in Listing A.2. The STOPS `split` method may take multiple parameters, as seen in § 3.3, but default parameters should be used wherever possible. The most notable parameters are the directory, which defaults to the current working directory of the CLI, the split row, which defaults to POLSALT’s default center row, and the save prefix, which defaults to ‘`obeam`’ and ‘`ebeam`’.

IRAF Wavelength Calibrations

The IRAF wavelength calibrations are performed using the tasks described in § 3.2, namely the `identify`, `reidentify`, `fitcoords`, and optionally `transform` tasks.

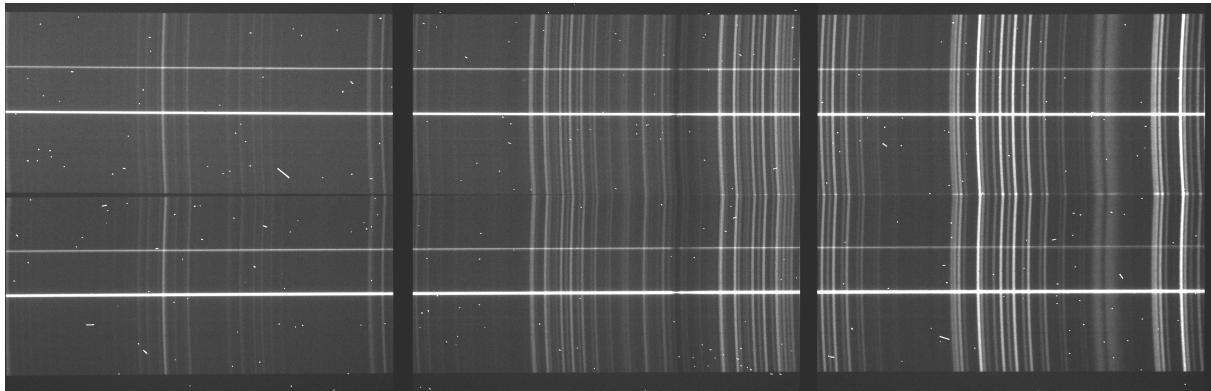


Figure 3.16: The ‘SCI’ extension of a typical spectropolarimetric FITS file taken with the SALT RSS, after basic POLSALT CCD reductions have been completed. Figure created from the `STOPS split` output.

In general, these tasks are run directly in the IRAF terminal using:¹⁸

```
cl> identify arc_files
cl> reidentify arc_ref arc_files
cl> fitcoords arc_files fit_2d
cl> transform files tr_file fit_2d
```

where ‘arc_files’ refers to a list or file containing the FITS files relevant to the task, ‘arc_ref’ refers to the FITS file previously identified, ‘fit_2d’ refers to the name to be used for the final two-dimensional wavelength solution, and ‘tr_file’ refers to the new name for the transformed input ‘files’.

The interactive tasks take up the bulk of the reduction time as this is where the fine-tuning of the reduction is done, through the use of cursor (or colon) commands, which allow modification of the parameters mid-reduction. Task parameters may, however, be edited beforehand within the IRAF terminal using the `eparam` task, and optionally saved, and quit or run using a combination of `:w`, and `:q` or `:go` cursor commands, respectively.

The reduction process in Appendix A, namely Listing A.4 to A.7, describes how the tasks may be scripted and saved for posterity. It is recommended to create an IRAF Command Language (cl) script for each task to keep track of which parameters were used and for simple recalibrations. The scripts are created using the `mkscript` task which interactively asks for a task to script and parameters to use. Multiple tasks may be appended to an IRAF script, allowing for the parameters of both beams to be tracked.

Running an IRAF script may be done by running:

```
cl> cl < script_name.cl
```

but is not suggested for interactive scripts, which run best when simply copied from the `<...>/sci/script_name.cl` file to the IRAF terminal.

¹⁸Please see the IRAF help docs, available at https://astro.uni-bonn.de/~sysstw/lfa_html/iraf/iraf.html, on the relevant tasks for a comprehensive discussion of the parameters available.

Preparing the Data for POLSALT

After the wavelength calibrations have been completed, the wavelength solution is parsed back into the format expected by POLSALT. Joining the separate beams with their respective wavelength solutions is performed in the CLI following Listing A.8.

Similar to the `split` procedure, the `join` procedure has the same defaults defined. The onus of keeping track of any previously changed default parameters falls to the user, but logging is implemented in STOPS (see the discussion on help documentation in § 3.3) which allows for later reference of any changed parameters.

Sky Line Checks

The optional IRAF `transform` task and STOPS `skylines` method are used to confirm the wavelength solution across the frame (see § 3.3.3) by transforming and comparing known and observed sky line wavelength positions, respectively.

The `skyline` method is run in the CLI following Listing A.9. As with the rest of STOPS, default parameters describe the overplotting behavior for the *O*- and *E*-beams, the skylines provided by SALT, and the calculated variation of the wavelength axis of a frame.

Cross Correlation Checks

The `correlate` method is run in the CLI following Listing A.11. The input of the `correlate` method takes the output of the POLSALT `spectra extraction` and is thus only run thereafter, but is mentioned here as the completion of the POLSALT reductions is not discussed in much depth. If the user wishes to compare the *O*- and *E*-beams of a single file then only that image name is to be provided, otherwise it is assumed that the user wishes to compare the same polarization beam across each file provided.

Cleaning Up the IRAF and STOPS Output

Before the final POLSALT reductions, it is recommended that the user ‘clean up’ the `sci/` directory of all IRAF and STOPS files since the ‘wmxgbp’ FITS files are all that is expected by POLSALT. The POLSALT methods use wildcard file collection and as such any errant detections of files added by the user will result in unexpected crashes. It is suggested to move any additional files to a new subfolder following Listing A.10, but they may also be removed using:

```
$ rm beam*.fits arc*.fits frame* <any user created files>
```

3.4.4 POLSALT Reduction Completion

Reductions may now be completed using POLSALT. The reduction process consists of correcting for the wollaston tilt, extracting the spectra, creating the Stokes files, and displaying the results. The ‘beta’ version of POLSALT provides access to a GUI but may also be handled entirely through a CLI as scripts.

Listing 3.7: The modified `reducepoldata_sc.py` script file.

```

import os, sys, glob, shutil
poldir = '/home/justin/polsalt-beta/'                                     # Will differ according to user
reddir=poldir+'polsalt/'
scrdir=poldir+'scripts/'
sys.path.extend((reddir,scrdir,poldir))

datadir = reddir+'data/'
import numpy as np
from astropy.io import fits as pyfits
from specpolview import printstokes
from imred import imred
from specpolwavmap import specpolwavmap
from specpolextract_sc import specpolextract_sc
from specpolrawstokes import specpolrawstokes
from specpolfinalstokes import specpolfinalstokes

print sys.argv
obsdate = sys.argv[1]
print obsdate
os.chdir(obsdate)
if not os.path.isdir('sci'): os.mkdir('sci')
shutil.copy(scrdir+'script.py','sci')
os.chdir('sci')

# basic image reductions
infilelist = sorted(glob.glob('../raw/P*fits'))
imred(infilelist, '.', datadir+'bpm_rss_11.fits', crthresh=False, cleanup=True)

# basic polarimetric reductions
logfile='specpol'+obsdate+'.log'                                         # The following lines may be removed or commented out as below
# wavelength map
# infilelist = sorted(glob.glob('m*fits'))
# linelistlib=""
# specpolwavmap(infilelist, linelistlib=linelistlib, logfile=logfile)

# background subtraction and extraction
infilelist = sorted(glob.glob('wm*fits'))
extract = 10.      # star +/-5, bkg= +/- (25-35) arcsec: 2nd order is 9-20 arcsec away
locate = (-120.,120.)    # science target is brightest target in whole slit
#locate = (-20.,20.)

specpolextract_sc(infilelist,logfile=logfile,locate=locate,extract=extract)
#specpolextract_sc(infilelist,logfile=logfile,locate=locate,extract=extract, docomp=True, useoldc=True)

# raw stokes
infilelist = sorted(glob.glob('e*fits'))
specpolrawstokes(infilelist, logfile=logfile)

# final stokes
infilelist = sorted(glob.glob('*_h*.fits'))
specpolfinalstokes(infilelist, logfile=logfile)

```

POLSALT Beta in the GUI

The reduction process using the POLSALT GUI is completed by selecting and, when applicable, interactively modifying the reduction step through the interactive windows, one-by-one, from the GUI's dropdown menu, as explained in Appendix A (p. 76 onwards).¹⁹

POLSALT Beta through a CLI

Both GUI and CLI implementations of the POLSALT beta pipeline access the same script files. Although the GUI is more user-friendly, the CLI offers a more streamlined approach to the reduction process, allowing the reduction process to be automated once the IRAF wavelength solution is known and parsed into the 'wmxgbp' FITS file format. A modified version of the POLSALT beta `reducepoldata_sc.py` script (see Listing 3.7) is used to run the entire reduction process without needing to select which process to run next, using:

```
$ python reducepoldata_sc.py YYYYMMDD
```

where the only modification made to the `reducepoldata_sc.py` script file is the removal of a call to the `specpolwavmap` method.

¹⁹See the official POLSALT wiki or alternative online resources such as the SALT workshop slides.

The POLSALT beta `reducepoldata_sc.py` copies a `script.py` file into the science working directory, ‘YYYYMMDD/sci/’, which provides analysis scripts for analysis and modification of the POLSALT beta results. These tools consist of data culling for the final Stokes calculations, text and plot output, relative flux calibration corrections, and synthetic filtering of polarization results.

The POLSALT analysis scripts may be run using:

```
$ python script.py
```

followed by `specpolfinalstokes.py`, `specpolview.py`, `specpolflux.py`, or `specpolfilter.py`, for the different analysis modes, respectively.²⁰

²⁰Please see <https://github.com/saltastro/polsalt/wiki/Linear-Polarization-Reduction---Beta-version> for a comprehensive discussion of the POLSALT beta analysis scripts.

Chapter 4

Testing and Application

This chapter contains an overview of the testing performed for the development of STOPS (§ 4.1) and the checking of the replaced wavelength solutions (§ 4.2), as well as the application of STOPS on observations of both spectrophotometric standards and science targets 4.3.

4.1 Testing STOPS

The main challenge faced when developing STOPS was ensuring that the software was compatible with both the POLSALT and IRAF file structures. As development is an iterative process, STOPS was continually checked to ensure compatibility such that the varying STOPS method inputs were correctly parsed, and that their outputs were parsable by the relevant IRAF tasks or POLSALT methods.

To this end, observations which were verified to have been accurately reduced were duplicated for testing purposes, allowing for continual checks of the STOPS pipeline to be made during the development process. As the STOPS `split` and `join` methods are designed to convert between the POLSALT and IRAF file structures, greater emphasis was made to ensure that the output of both methods provided accurate and consistent results.

4.1.1 Testing the split Method

The STOPS `split` method requires any POLSALT pre-reduced ('mxgbp'- prefixed) FITS files as input and outputs IRAF compatible ('(arc|beam)(O|E)'- prefixed) FITS file structures. As no 'split' FITS files are created during pure POLSALT reductions, the STOPS `split` method was tested by comparing the pre-reduced POLSALT files to the `split` method's output files, ensuring the correct structure and data integrity of the files handed off to IRAF.

Table 4.1 shows the FITS file information for the files before and after splitting. The split FITS files contain the split 'SCI' extension data and the 'Primary' header from the pre-reduced files, with any Header or Data differences mentioned below.

Filename	No.	Name	Type	Cards	Dimensions	Format
POLSLT	0	‘Primary’	PrimaryHDU	161	()	
	1	‘SCI’	ImageHDU	19	(3199, 1028)	float32
	2	‘VAR’	ImageHDU	8	(3199, 1028)	float32
	3	‘BPM’	ImageHDU	8	(3199, 1028)	uint8
STOPS split ‘O’	0	‘Primary’	PrimaryHDU	162	(3199, 474)	float32
STOPS split ‘E’	0	‘Primary’	PrimaryHDU	162	(3199, 474)	float32

Table 4.1: A comparison of the contents of a POLSLT pre-reduced FITS file to the STOPs split *O*- and *E*-beam FITS files. Table created using the `Astropy fitsinfo` CLI tool.



(a) The difference in the ‘SCI’ extensions, for the ‘O’ polarization beam.



(b) The difference in the ‘SCI’ extensions, for the ‘E’ polarization beam.

Figure 4.1: The difference between the POLSLT pre-reduced (‘mxgbp’- prefixed) FITS files and the STOPs ‘split’ (‘arc|beam)(O|E’- prefixed) files. Figures created using both the `POLSLT Raw image reduction` and `STOPs split` method outputs.

The header is left mostly untouched, and is only updated to represent the new data type and shape: the ‘BITPIX’ value is updated, from 8 to –32, and the ‘NAXIS’ value is updated, from 0 to 2; the ‘NAXIS1’ and ‘NAXIS2’ keywords are added, and their values are set to the new split ‘SCI’ data shape; and the ‘EXTEND’ keyword is removed.¹ This accounts for the discrepancy in the ‘Cards’ between the POLSLT and STOPs file header entries in Table 4.1.

Figure 4.1 shows that the POLSLT ‘SCI’ data is unmodified when copying the data to the STOPs FITS file, but only includes half of the data, for the relevant *O*- or *E*-polarization beam, Figure 4.1a and Figure 4.1b, respectively, with a cropping which defaults to 40 pixels (see § 3.3.1), introduced to the top- and bottom-most rows of the POLSLT data. This accounts for the discrepancy in the ‘Dimensions’ between the POLSLT and STOPs files in Table 4.1.

This output file structure was chosen for IRAF compatibility, and was tested over multiple grating and articulation angles, as well as with various data sets to ensure that the `split` method was robust and reliable.

¹The ‘EXTEND’ keyword indicates that the FITS file contains multiple extensions while the ‘NAXIS1’ and ‘NAXIS2’ keywords indicate the shape and size of the data stored in the relevant extension.

Filename	No.	Name	Type	Cards	Dimensions	Format
POLSLT	0	'Primary'	PrimaryHDU	161	()	
	1	'SCI'	ImageHDU	21	(3199, 514, 2)	float32
	2	'VAR'	ImageHDU	10	(3199, 514, 2)	float32
	3	'BPM'	ImageHDU	10	(3199, 514, 2)	uint8
	4	'WAV'	ImageHDU	21	(3199, 514, 2)	float32
STOPs join	0	'Primary'	PrimaryHDU	161	()	
	1	'SCI'	ImageHDU	21	(3199, 474, 2)	float32
	2	'VAR'	ImageHDU	10	(3199, 474, 2)	float32
	3	'BPM'	ImageHDU	10	(3199, 474, 2)	uint8
	4	'WAV'	ImageHDU	21	(3199, 474, 2)	float32

Table 4.2: A comparison of the POLSLT wavelength calibrated FITS file to the (IRAF wavelength calibrated) STOPs join FITS file. Table created using the Astropy `fitsinfo` CLI tool.

4.1.2 Testing the join Method

The `join` method requires both an IRAF database with wavelength solutions (or a custom wavelength solution) for both polarimetric beams and the POLSLT pre-reduced files as input and outputs POLSLT spectra extraction compatible ('wmxgbp'- prefixed) FITS file structures. Ensuring that the output format was correct was paramount as the POLSLT spectra extraction method is unable to process the files otherwise, thus halting the reduction process. Thankfully, the `join` method output could be compared to the POLSLT wavelength calibration method output files, ensuring that any changes introduced by the STOPs pipeline were well characterized.

Table 4.2 shows the FITS file information for both the POLSLT and STOPs wavelength calibrated files. Other than the 'Dimensions' of each 'ImageHDU' extension,² the FITS files are identical in structure.

Although the 'Cards' count is the same, minor differences across the headers are present. The 'HISTORY' keyword, which contains the POLSLT 'CRCLEAN' parameters and which default to 'upper= 4.0, lower= 1.5, sigmaveto= 2.0', is left as 'None' in the STOPs file.³ Although STOPs performs cosmic ray cleaning (see § 3.3.2), the parameters are not stored in the header as POLSLT and STOPs implement different methods for cosmic ray cleaning. Other minor differences such as the date-times stored in the 'SALTLM' and 'SMOSAIC' keywords may also differ as they contain the date-times relating to the completion of the POLSLT pre-reductions. This accounts for the differences in the 'Cards' between the POLSLT and STOPs file header entries in Table 4.2.

Figure 4.2 shows the differences in the data between the POLSLT and STOPs wavelength calibrated files. It can be seen that the 'VAR' extensions (Figure 4.2b) are identical. The 'SCI' extensions (Figure 4.2a) differ only in that the cosmic ray cleaning has been applied to the STOPs data, whereas the POLSLT data applies a mask to the cosmic rays using the 'BPM' extension (Figure 4.2c). The 'BPM' and 'WAV' extensions (Figure 4.2d)

²The 'Dimensions' differ due to the before mentioned cropping of the top- and bottom-most rows of the data.

³The POLSLT pipeline performs cosmic ray cleaning using a 10σ spike to cull cosmic rays. See the POLSLT source code for more information.

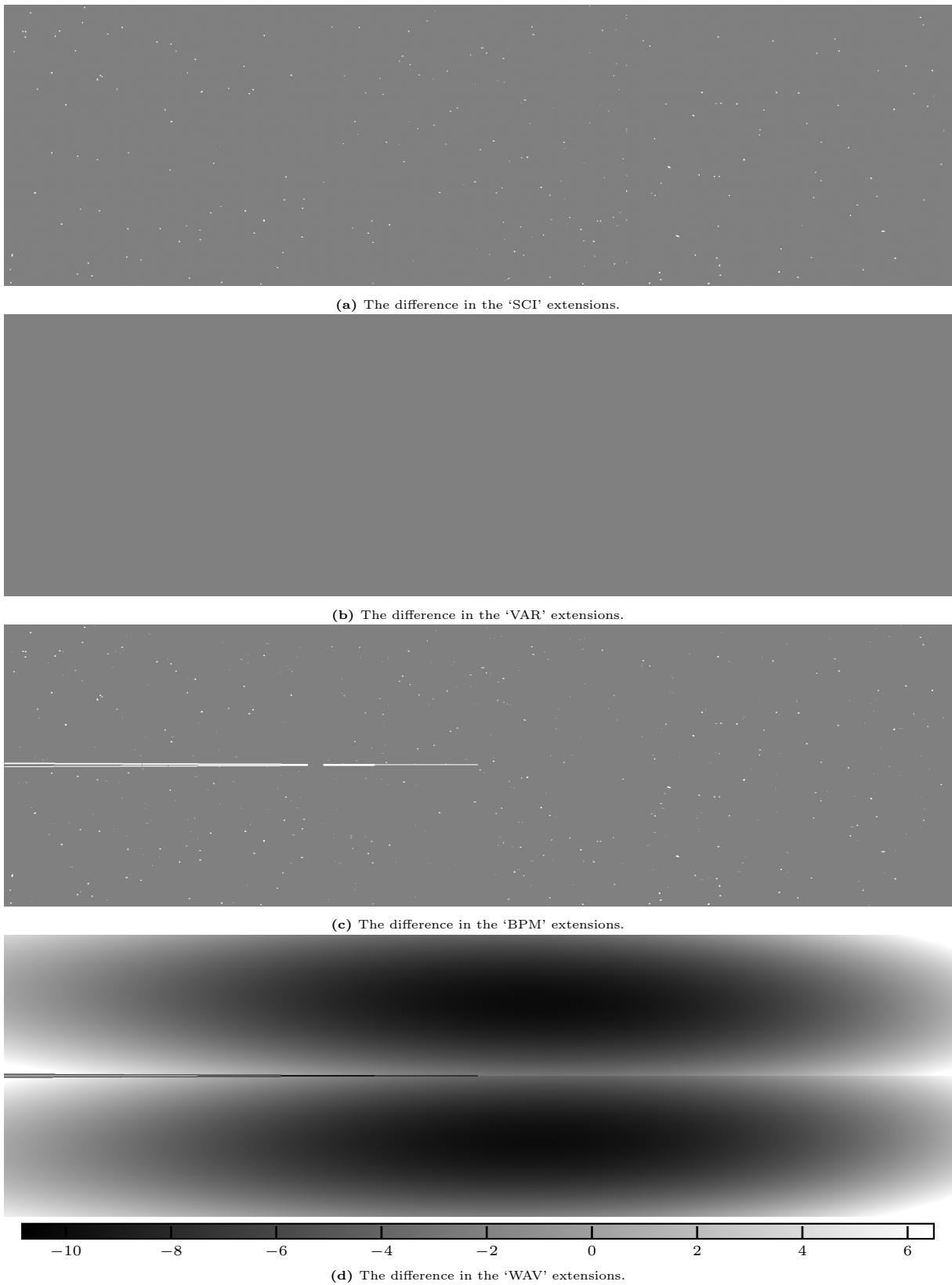


Figure 4.2: The difference of the FITS file extensions between the `POLSALT` and `STOPS` ('wmxgbp'- prefixed) wavelength calibrated files. Figures created using both the `POLSALT` and `STOPS` versions of the `POLSALT spectral extraction` input.

are also masked to account for the valid wavelength calibrated region. The ‘WAV’ extensions contain the differing wavelength solutions and as such naturally differ. This accounts for the differences in the data between the POLSALT and STOPS files.

Finally, the STOPS `join` method was tested to ensure compatibility and correctness of the output data in comparison to POLSALT. This involved testing the `join` method with various data sets to ensure that the output files were accurate and consistent.

4.2 Wavelength Solution Checks

The secondary challenge encountered when developing STOPS was ensuring that the wavelength solutions parsed by STOPS were unaffected by the pipeline and that they were similar to those created by POLSALT. This was achieved through the `correlate` and `skylines` methods, which were designed to validate the wavelength solutions produced by IRAF, but were later modified to parse both the IRAF and POLSALT wavelength solutions, allowing for further inspection of the two-dimensional wavelength solution.

Before the POLSALT wavelength calibrations were replaced with the IRAF wavelength calibrations, the accuracy of the new wavelength solutions needed to be validated. This was done both through the IRAF tasks, ensuring an accurate wavelength solution, and through the STOPS `correlate` and `skylines` methods, allowing the integration of the wavelength solutions to be validated.

TODO: Add RMS results (Table / Figures / Both?) from wavelength solution checks (IRAF RMS, POLSALT RMS) to quantify differences.

TODO: Plot IRAF vs POLSALT RMS for sources?

4.2.1 Cross Correlation Checks

The `correlate` method returns plots validating the wavelength solutions and so only has to accept the POLSALT `spectra extraction` ('ecwmxgbp'- prefixed) method output files as input.

The STOPS `correlate` method was tested by cross correlating generated *O*- and *E*-beam spectra with known offsets, with the aim of reacquiring said offsets, as shown in Figure 4.3. The spectra were generated with a feature in each CCD region, randomly offset in both the wavelength and intensity axes, Figure 4.3a.

Through cross correlation, Figure 4.3b, the introduced offsets, or ‘max lag’, were reacquired. For spectral regions with few features or features not much more significant than the continuum noise (such as the left most CCD region of Figure 4.3b), correlation may fail to determine the correct offset. It is clear that the returned ‘max lag’ is incorrect when the ‘max lag’ peak is not significantly larger than any noise of the continuum in the correlation plot.

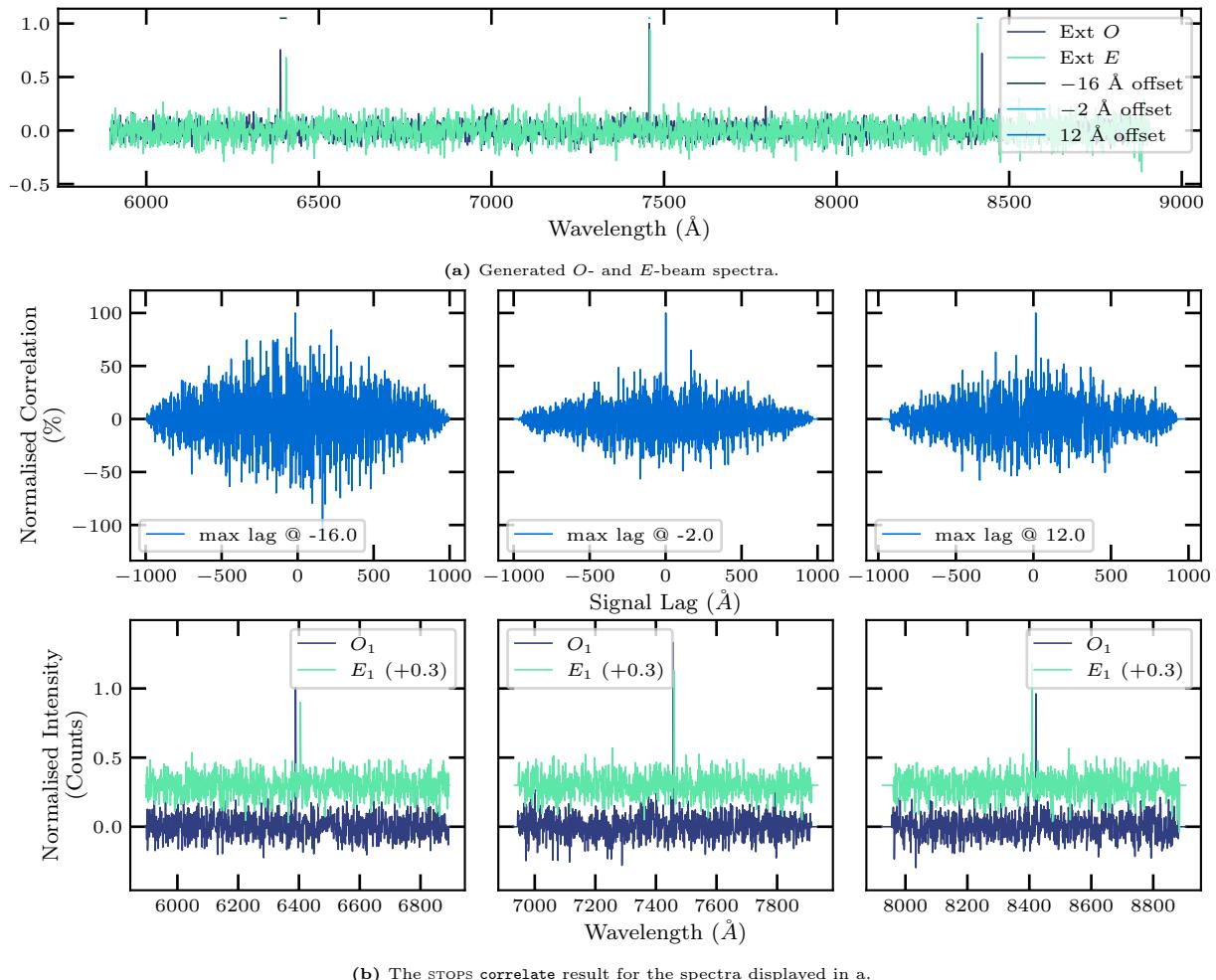


Figure 4.3: Reacquisition of synthetic offsets introduced to the O - and E -beam spectra (a) by cross-correlation (b, bottom row).

Source	Observation Date	Grating	Grating Angle (°)	Articulation Angle (°)	Exposure Time (sec)
Hiltner 600	2017-02-24	PG0900	12.500	24.91	135.60
			19.625	39.16	135.60
3C 279	2017-03-28	PG0900	12.498	24.91	480.81
			19.625	39.16	480.80
	2017-04-01	PG0900	12.500	24.92	480.78
			19.625	39.17	480.85
	2017-05-17	PG0300	5.375	10.68	720.84
			12.500	24.92	720.83
			19.625	39.17	901.05
	2017-05-21	PG0300	5.375	10.66	1200.81
	2018-06-05	PG0300	5.375	10.72	1200.81
Hiltner 652	2022-06-10	PG0900	12.878	25.71	48.50

Table 4.3: Sources discussed from proceedings and publications within this section. Table adapted from Cooper and van Soelen (2023).

4.2.2 Sky Line Checks

The `skylines` method returns plots validating the wavelength solutions and so only has to accept either the `IRAF transform` task or `STOPS join` method output files as input.

TODO: Compare skylines from STOPS to ‘poor’ POLSALT spectral extraction (I.E. spectral extraction with no trace in the ‘target’ window and the ‘background’ window on region with no skylines).

TODO: Test ‘skylines’ using known spectral sky lines / telluric lines.

TODO: Insert figure illustrating skyline identification accuracy.

4.3 Application of STOPS

The STOPS pipeline has been utilized in the reduction of a number of science targets, specifically focused on blazars. The STOPS pipeline has been utilized in the reduction of a number of sources, both for calibration tests using spectropolarimetric standards, and for science observations of transient blazars. The following sections discuss works from which observations were completed using, in part, the STOPS pipeline with `IRAF` wavelength calibrations, or from which the STOPS pipeline was developed.

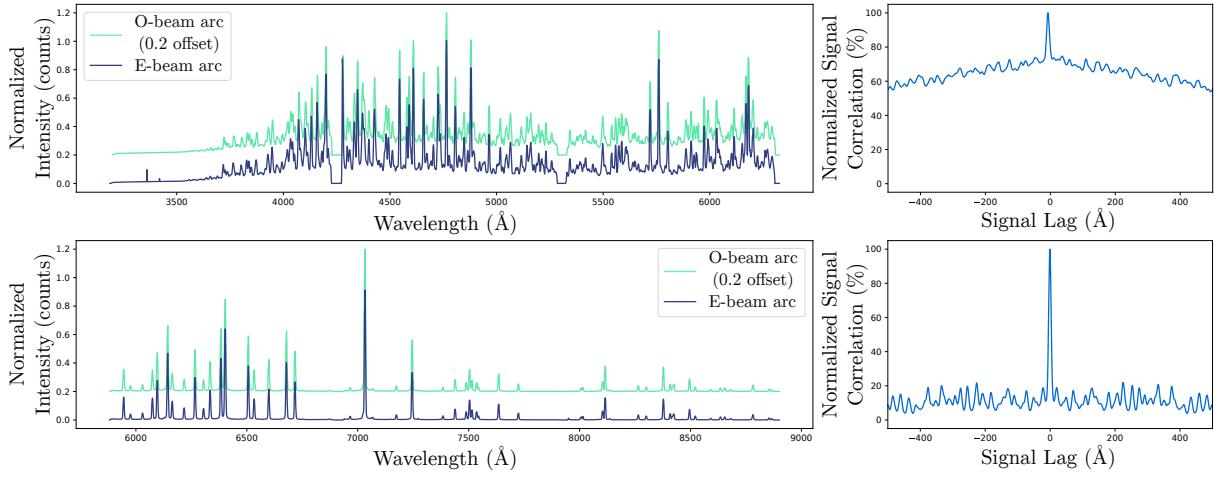


Figure 4.4: The spectra and cross correlation of the *O*- and *E*-beams of the ThAr (top) and Ne (bottom) arc lamps. Figure adapted from (Cooper et al., 2022).

4.3.1 Spectropolarimetry and Photometry of the Early Afterglow of the Gamma-ray Burst GRB 191221B (Buckley et al., 2021)

4.3.2 Development of Tools for the SALT/RSS Spectropolarimetry Reductions: Application to the Blazar 3C 279 (Proceedings of Science, HEASA 2021)

As discussed in (Cooper et al., 2022, see also Appendix C), it was shown that alternative wavelength solutions, such as those created using IRAF, are capable of being applied to SALT spectropolarimetric data. The ‘additional tools’ mentioned therein, which were the precursor to the STOPS pipeline, were used during the reduction of the blazar 3C 279.

The blazar 3C 279 was observed in linear spectropolarimetry mode, on 2017 May 17, using the PG0900 grating with two different grating angles (12.5° and 19.5°). The grating and articulation angles used for the observations, as shown in Table 4.3, matched those of observations of the spectrophotometric standard star, Hiltner 600, observed on 2017 February 24, allowing the intensity to be relatively flux calibrated.

Figure 4.4 shows the spectra and cross correlation of the *O*- and *E*-beams of the ThAr and Ne arc lamps, also with grating angles of 12.5° and 19.5° , respectively. The cross correlation of the arc lamps perpendicular polarization beams show clear peaks at 0 lag, indicating that the wavelength solutions are consistent across the two polarization beams.

Figure 4.5 shows the relative flux calibrated intensity, percentage of linear polarization, and polarization angle of the blazar 3C 279, across the visible spectrum. The spectra display a good overlap between the two grating angles, especially when considering the variable nature of blazars as well as the fact that the spectra are only relatively flux calibrated. Both the percentage of linear polarization and the polarization angle agree well across the grating angle overlap.

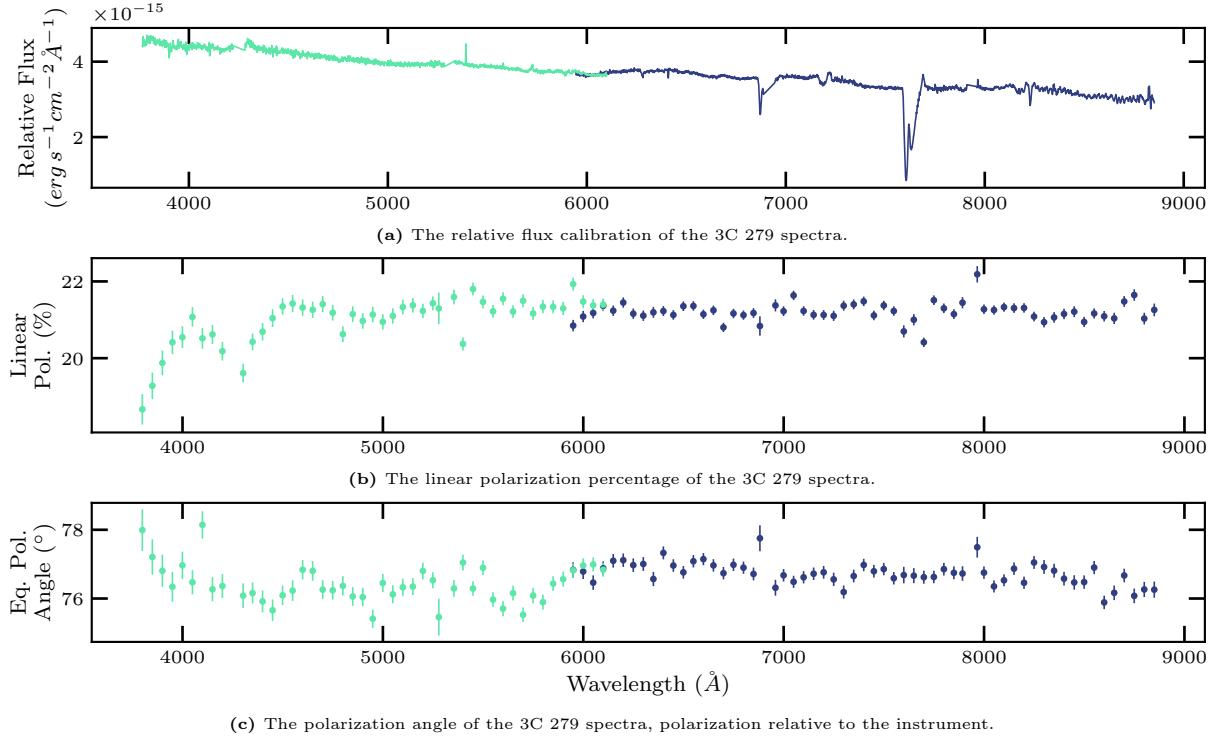


Figure 4.5: The relative flux calibrated spectra (a), linear polarization percentage (b), and polarization angle (c) of the blazar 3C 279 observed on 2017 May 17. Figure adapted from Cooper et al. (2022).

4.3.3 SALT Spectropolarimetric Pipeline Comparisons (Proceedings of Science, HEASA 2022)

As discussed in (Cooper and van Soelen, 2023, see also Appendix C), as well as expanding on the work presented in Cooper et al. (2022), it was shown that the STOPS pipeline has no noticeable effect on the polarization properties of spectropolarimetric data and that the measure of the goodness of fit for the wavelength solutions could be measured, both through correlation and measuring the position of skyline features for both the *O*- and *E*-beams. The blazar 3C 279, the spectrophotometric standard star Hiltner 600, and the spectropolarimetric standard star Hiltner 652, were observed on the dates tabulated in Table 4.3. Observations of 3C 279 were taken during periods of enhanced and flaring γ -ray activity.

Figure 4.6 shows the photon flux of 3C 279 for energies of $0.1 - 100$ GeV as provided by the Fermi Light Curve Repository (Fermi LCR) (Abdollahi et al., 2023). The figure shows that during the optical observations of 3C 279, the source was likely at γ -ray fluxes of 1.59×10^{-6} , 2.87×10^{-6} , and $2.00 \times 10^{-6} \text{ ph. cm}^{-2}\text{s}^{-1}$, as compared to a quiescent ‘median’ flux of $0.46 \times 10^{-6} \text{ ph. cm}^{-2}\text{s}^{-1}$.

Figure 4.7 shows the spectropolarimetric results for 3C 279 during the periods of enhanced and flaring γ -ray activity as mentioned above. The figure shows, from top to bottom, the normalized, relative flux calibrated, spectra for both grating angles (12.5° (observed first at each epoch) and 19.625°), the percentage of linear polarization, and the polarization angle.

⁴https://fermi.gsfc.nasa.gov/ssc/data/access/lat/LightCurveRepository/source.html?source_name=4FGL_J1256.1-0547

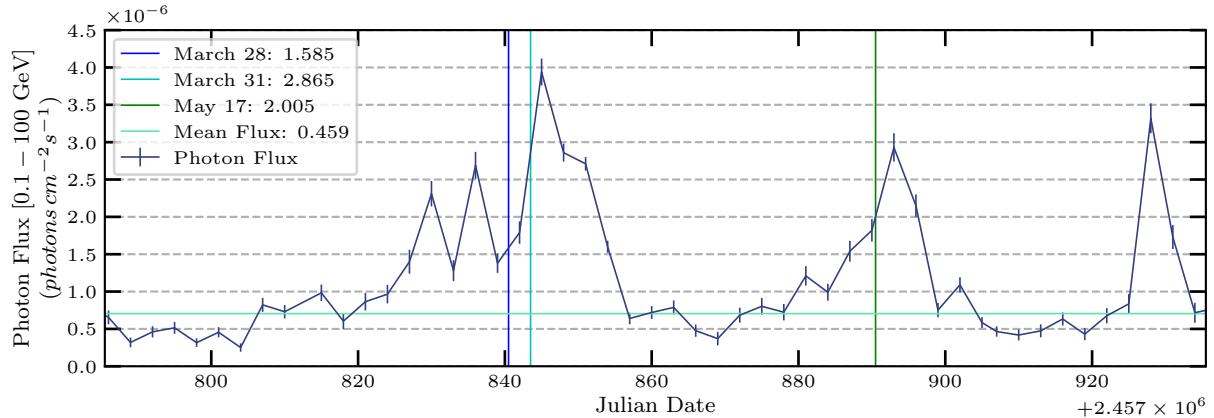


Figure 4.6: The photon flux of 3C 279 as measured during the presented observation dates. Also included are the ‘mean’ γ -ray fluxes for the observation dates, calculated using linear interpolation, as well as the non-flaring ‘mean’ γ -ray photon flux. Figure adapted from data provided via the Fermi LCR.⁴

Although the spectral shape does not vary greatly across the wavelength region, the level of γ -ray activity during observation has a clear effect on the spectra. This is most notable for the 2017 March 31 observation which displays heightened spectral intensities in the blue. The linear polarization percentages and polarization angles also show variation over the course of the observations, varying by $\sim 12.23\%$ and $\sim 24.22^\circ$, respectively. There is also good agreement between the polarization properties across the differing grating angles for the same observation dates. The difference of the polarization percentage and polarization angle between the two grating angles is $\sim 0.25\%$ and $\sim 1.25^\circ$, respectively. The most notable disagreement of the polarization properties is seen, once again, during the 2017 March 31 observation, but may be explained due to the flaring state of the blazar, the asynchronous times for the two observations, and a lower S/N near the edges of the RSS.

Figure 4.8 shows the reduced Q and U Stokes parameters for Hiltner 652 as observed with the Focal Reducer and Low Dispersion Spectrograph (FORS1) and FORS2 mounted at the Very Large Telescope (VLT) (Cikota et al., 2017; Fossati et al., 2007), as well as observed with SALT. The figure shows that the Stokes parameters of Hiltner 652 are consistent across the different instruments, mostly falling within a standard deviation of historical results. Minor discrepancies are present, but are likely due to multiple compounding effects, such as differing instrumental setups, varying atmospheric conditions, as well as Hiltner 652 being classed as a spectrophotometric, and not spectropolarimetric, standard.

4.3.4 Modeling the Spectral Energy Distributions and Spectropolarimetry of Blazars - Application to 4C+01.02 in 2016 - 2017 (Schutte et al., 2022)

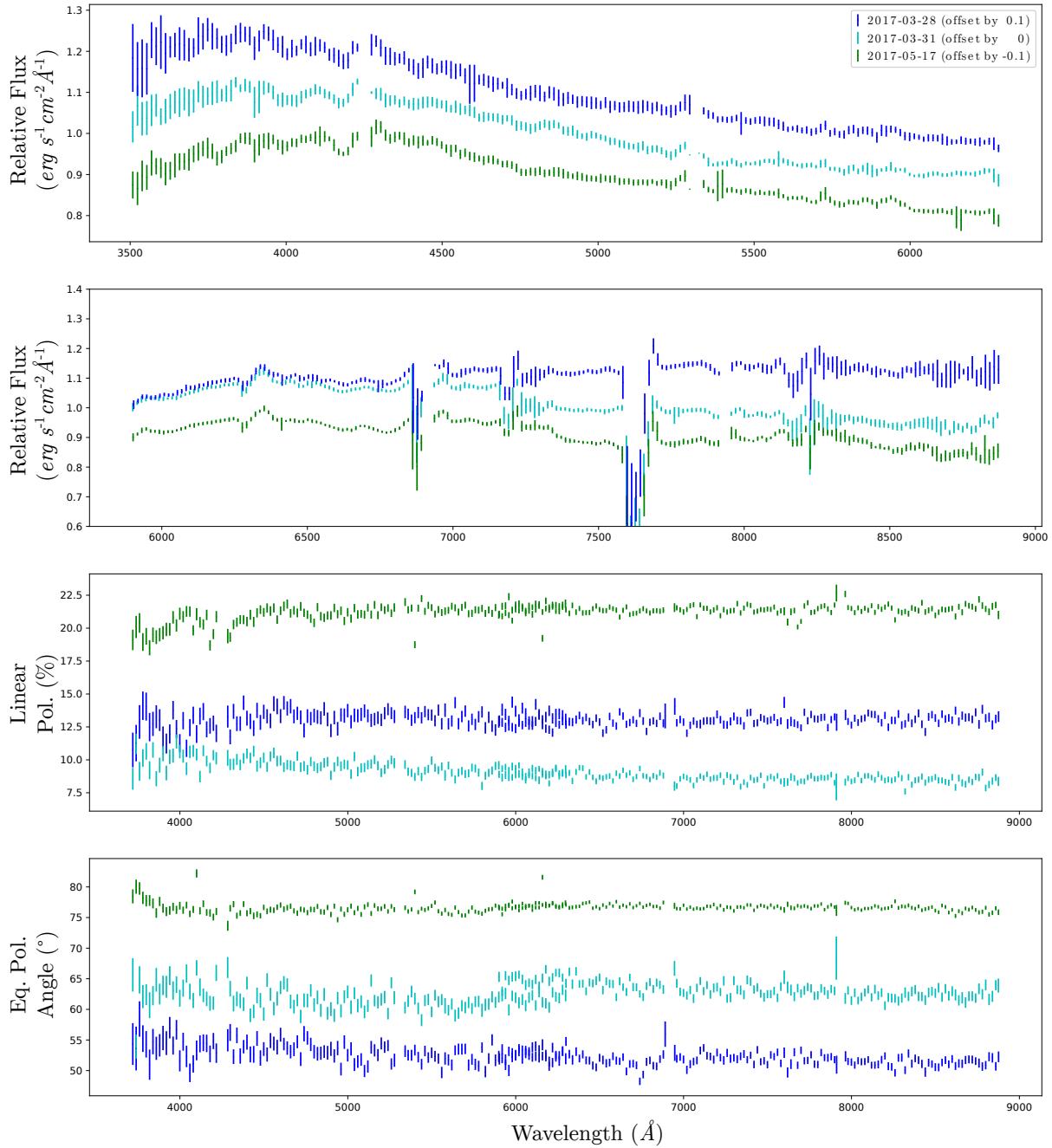


Figure 4.7: Spectropolarimetric results for 3C 279 during two distinct periods of flaring. The normalized relative flux is presented for both grating angles, 12.5° (top) and 19.625° (second from top). Figure sourced from (Cooper and van Soelen, 2023).

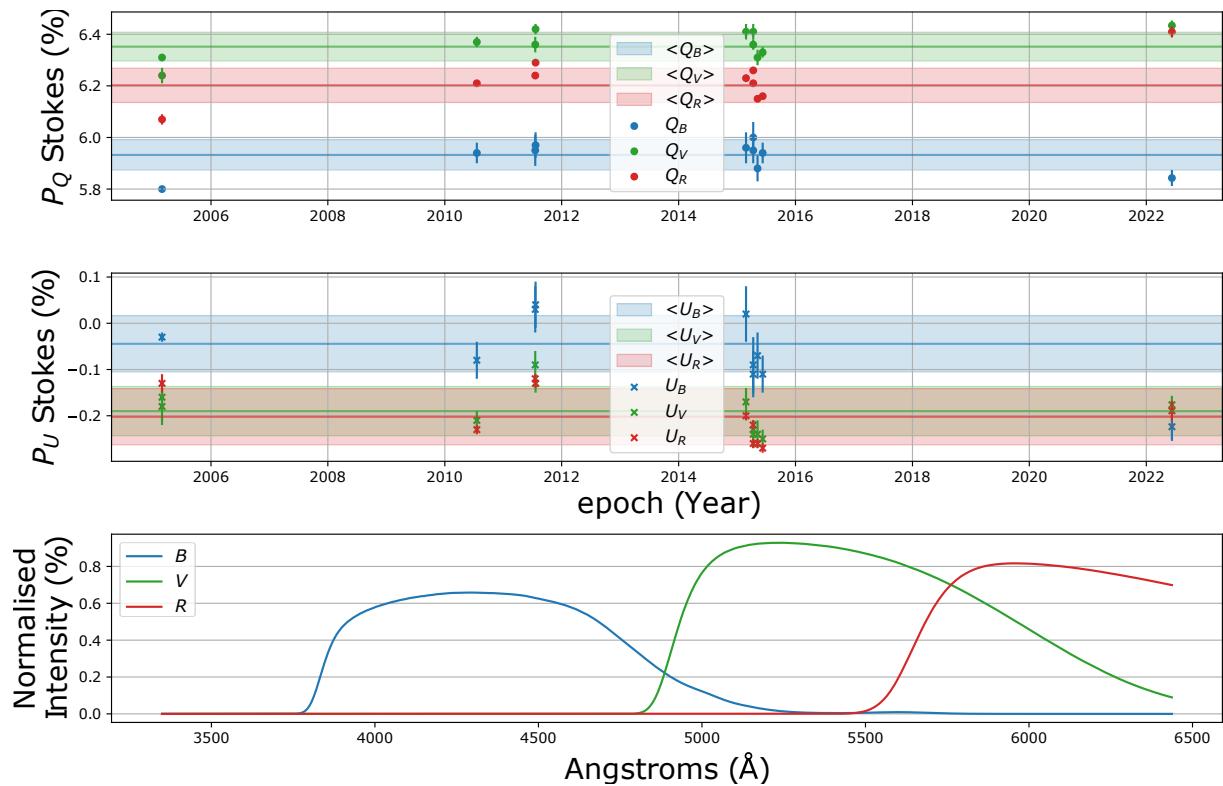


Figure 4.8: The reduced Q and U Stokes parameters for Hiltner 652, pre–2006 for FORS1, 2010 – 2016 for FORS2, and post–2022 for SALT. Figure sourced from (Cooper and van Soelen, 2023).

Chapter 5

Conclusions

TODO: A summary of the dissertation, main focus on the results and that the supplementary pipeline is a success since it allows an alternate method using IRAF to wavelength calibrate the polsalt data.

5.1 Future Work

TODO: Edit paragraph below to mention python wavelength solutions implemented to ‘future-proof’ the pipeline.

Another option to perform the wavelength calibration is Python which allows for a more modern and flexible approach, but is not discussed here. What will be discussed, however, is the structure of the wavelength solutions created through Python to be later reintroduced to the POLSALT pipeline. The solutions must be stored such that the ‘ x ’ and ‘ y ’ orders of the solution, as well as all the coefficients (C_{00} to C_{xy}) making up the solution, separated by new lines, are included. The only limitations to the names of the solution files is that they must make mention of the specific O - or E -beam as well as the wavelength solution type (e.g. ‘Chebyshev’, ‘Legendre’, etc.).

Appendix A

The Modified Reduction Process

This section of the Appendix aims to provide a minimum working example of the commands necessary to reduce POLSALT data using STOPS and IRAF. It contains the commands necessary to activate all software and run through the reduction process but makes no attempt at discussion.

Both POLSALT and IRAF are launched from the default CLI but use independent interfaces during the reduction process. To distinguish which window is in focus, the ‘\$’ token is used for default CLI commands while the ‘c1>’ and ‘>>>’ tokens are used for IRAF’s ‘xgterm’ single- and multi-line commands, respectively.

General instructions for the reduction process which might not necessarily be line-fed commands passed to a CLI may either be discussed outside a ‘Listing’ environment or included as part of the ‘Listing’ environment with a preceding ‘#’ token. Finally, POLSALT implements a GUI and thus takes no line-fed commands. As such, the instructions when using the POLSALT GUI follow those of the general instructions with the added exception that they relate to the GUI.

As a final note, some parameters are distinguished using an ‘<angle brackets>’ notation. They signify necessary parameters that may vary from reduction to reduction. Notable uses of this notation include the date of observation, $\langle OBSDATE \rangle$ (formatted ‘YYYYMMDD’), the split science FITS files, $\langle O\text{-beam FILES} \rangle$ or $\langle E\text{-beam FILES} \rangle$, the split arc FITS files, $\langle O\text{-beam ARC} \rangle$ or $\langle E\text{-beam ARC} \rangle$, and a wildcard symbol, $\langle * \rangle$.

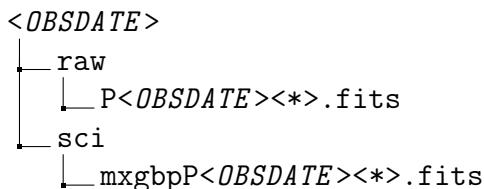


Figure A.1: The typical minimal file structure of data provided by SALT.

Ensure the data is formatted in a file structure similar to that in Figure A.1. Data located in the ‘sci’ folder is often provided by SALT but is not strictly necessary to begin the reduction process. If ‘mxgbp’ prefixed data is available, the reductions may be begun starting at Listing A.2. The POLSALT GUI is launched from the default CLI running the commands in Listing A.1.

Listing A.1: Launching the POLSALT GUI

```
$ cd ~/polsalt
$ conda activate salt
$ python -W ignore reducepoldataGUI.py &
```

Refer to Figure 3.1 for a depiction of the POLSALT GUI. To complete the POLSALT pre-calibrations, and with the GUI in focus:

- Ensure that the ‘POLSALT code directory’ is correct.
- Set the ‘Top level data directory’ to $\langle OBSDATE \rangle$.
- Ensure ‘Raw data directory’ is correct.
- Ensure ‘Science data directory’ is correct.
- Select ‘Raw image reduction’ from the ‘Data reduction step’ drop down menu.
- Check the tick boxes of all raw images to be processed (include the arc) in the display box covering the lower half of the GUI.
- Proceed with the reductions by clicking the ‘OK’ button.

The pre-calibrated ‘mxgbp’ FITS files are now available in the ‘sci’ folder. The files may be split using STOPS by running the commands in Listing A.2.

Listing A.2: Splitting data using STOPS

```
$ cd <OBSDATE>/sci
$ conda activate stops
$ python ~/STOPS . split
```

The split *O*- and *E*-beam FITS files are now available. The IRAF wavelength calibrations are now run. The IRAF xgterm CLI is launched using Listing A.3.

Listing A.3: Launching IRAF in xgterm

```
$ cd ~/iraf
$ xgterm -sb &
cl> conda activate salt
cl> cl
cl> noao
cl> twodspec
cl> longslit
cl> unlearn longslit
cl> longslit.dispaxis=1
```

The IRAF `identify` task requires an average feature width, ‘fwidth’, as a parameter. The width of a feature may be found in IRAF using the `implot` task¹ along with cursor commands, but may also be found using FITS viewing software, such as `ds9`.² The `identify` task may be run using the commands in Listing A.4.

Listing A.4: Running the IRAF `identify` task

```
cl> mkscript 01_identify.cl
cl> # Add identify to 01_identify.cl twice, for both beams
cl> # Edit the parameters of 01_identify.cl in a text editor
cl> # Paste an identify script into the CLI, resulting in:
cl>
cl> identify ("<O-beam ARC>",
>>> "", "", section="middle line", database="database",
>>> coordlist="linelists$idheneare.dat", units="", nsum="10", match=-3.,
>>> maxfeatures=50, zwidth=100., ftype="emission", fwidth=8.,
>>> cradius=5., threshold=0., minsep=2., function="spline3", order=2,
>>> sample="*", niterate=0, low_reject=3., high_reject=3., grow=0.,
>>> autowrite=no, graphics="stdgraph", cursor="", aidpars="")
```

The `identify` task will launch an interactive window. Cursor commands refer to keys that provide unique functionality to the interactive IRAF tasks. The cursor commands for `identify` allow the arc lines to be identified using ‘m’ (and typing the relevant wavelength), while ‘d’ and ‘i’ will delete a single and all identified arc lines, respectively. The ‘f’ cursor command will perform a preliminary fit which can be quit using the ‘q’ cursor command. The ‘l’ cursor command will attempt to identify any unidentified arc lines. Once complete, a figure of the identified lines may be saved using ‘:labels coord’ and ‘:.snap eps’, and the task safely quit with the ‘q’ cursor command.³ The `identify` procedure is repeated, replacing $<O\text{-beam } ARC>$ with $<E\text{-beam } ARC>$.

The `reidentify` task may be run using the commands in Listing A.5.

Listing A.5: Running the IRAF `reidentify` task

```
cl> mkscript 02_reidentify.cl
cl> # Add reidentify to 02_reidentify.cl twice, for both beams
cl> # Edit the parameters of 02_reidentify.cl in a text editor
cl> # Paste a reidentify script into the CLI, resulting in:
cl>
cl> reidentify ("<O-beam ARC>",
>>> "<O-beam ARC>", "yes", "", "", interactive="no", section="middle
>>> line", newaps=yes, override=no, refit=yes, trace=yes, step="10",
>>> nsum="10", shift="0.", search=0., nlost=0, cradius=5.,
>>> threshold=0., addfeatures=no, coordlist="linelists$idheneare.dat",
>>> match=-3., maxfeatures=50, minsep=2., database="database",
>>> logfiles="logfile", plotfile="", verbose=yes, graphics="stdgraph",
>>> cursor="", aidpars="")
```

¹See https://astro.uni-bonn.de/~sysstw/lfa_html/iraf/plot.implot.html for documentation on the `implot` task.

²See <https://sites.google.com/cfa.harvard.edu/saoimageds9> for documentation on the `ds9` software.

³See https://astro.uni-bonn.de/~sysstw/lfa_html/iraf/noao.onedspec.identify.html for documentation on the `identify` task.

The `reidentify` task will run autonomously so long as the `interactive` parameter is set to “no”.⁴ Repeat the `reidentify` procedure, replacing $\langle O\text{-beam } ARC \rangle$ with $\langle E\text{-beam } ARC \rangle$ at both the ‘reference’ and ‘image’ parameter locations.

The `fitcoords` task may be run using the commands in Listing A.6.

Listing A.6: Running the `IRAF fitcoords` task

```
cl> mkscript 03_fitcoords.cl
cl> # Add fitcoords to 03_fitcoords.cl twice, for both beams
cl> # Edit the parameters of 03_fitcoords.cl in a text editor
cl> # Paste a fitcoords script into the CLI, resulting in:
cl>
cl> fitcoords ("<O-beam ARC> (exclude the file extension)" ,
>>> fitname="", interactive=yes, combine=no, database="database",
>>> deletions="deletions.db", function="chebyshev", xorder=6, yorder=6,
>>> logfiles="STDOUT,logfile", plotfile="plotfile",
>>> graphics="stdgraph", cursor="")
```

The `fitcoords` task will launch an interactive window. The x- and y-axis being plotted may be changed using the ‘x’ or ‘y’ cursor commands followed by the desired data axis (‘x’ for the x-axis, ‘y’ for the y-axis, or ‘r’ for the residuals).⁵ Repeat the `fitcoords` procedure, replacing $\langle O\text{-beam } ARC \rangle$ with $\langle E\text{-beam } ARC \rangle$.

The `transform` task may be run using the commands in Listing A.7.

Listing A.7: Running the `IRAF transform` task

```
cl> mkscript 04_transform.cl
cl> # Add transform to 04_transform.cl twice, for both beams
cl> # Edit the parameters of 04_transform.cl in a text editor
cl> # Paste a transform script into the CLI, resulting in:
cl>
cl> transform ("@<O-beam FILES>" ,
>>> "t/@<O-beam FILES>", "<O-beam ARC> (exclude the file extension)" ,
>>> minput="", moutput="", database="database", interptype="linear",
>>> x1="INDEF", x2="INDEF", dx="INDEF", nx="INDEF", xlog="no",
>>> y1="INDEF", y2="INDEF", dy="INDEF", ny="INDEF", ylog="no",
>>> flux="yes", blank="INDEF", logfiles="STDOUT,logfile")
```

The `transform` task will run autonomously.⁶ Repeat the `transform` procedure, replacing the $\langle O\text{-beam } FILES \rangle$ and $\langle O\text{-beam } ARC \rangle$ with $\langle E\text{-beam } FILES \rangle$ and $\langle E\text{-beam } ARC \rangle$ at both parameter locations. Inspect the transformed images, most notably the arc images, using any FITS viewer as a cursory check that the wavelength calibrations were completed without error.

The ‘gain’ and ‘read noise’ are now needed as the cosmic-ray rejection of the STOPS `join` method accepts them as parameters. These parameters may be found using the ‘`GAINSET`’ and ‘`ROSPEED`’ keywords in the FITS headers. The cosmic ray rejection

⁴See https://astro.uni-bonn.de/~sysstw/lfa_html/iraf/noao.onedspec.reidentify.html for documentation on the `reidentify` task.

⁵See https://astro.uni-bonn.de/~sysstw/lfa_html/iraf/noao.twodspec.longslit.fitcoords.html for documentation on the `fitcoords` task.

⁶See https://astro.uni-bonn.de/~sysstw/lfa_html/iraf/noao.twodspec.longslit.transform.html for documentation on the `transform` task.

defaults to *GAINSET*=‘FAINT’, and *ROSPEED*=‘SLOW’. If the gain and read noise values differ from the defaults, the parameters should be updated when running `join`.⁷

The STOPS `join` method may be run using the commands in Listing A.8.

Listing A.8: Joining the data using STOPS

```
$ cd <OBSDATE>/sci
$ conda activate stops
$ python ~/STOPS . join
```

The STOPS `skylines` method may be run on any ‘joined’ or transformed FITS files, *<FILE(S)>*, using the commands in Listing A.9.

Listing A.9: Running the STOPS `skylines` method

```
$ cd <OBSDATE>/sci
$ conda activate stops
$ python ~/STOPS . skylines <FILE(S)>
```

The ‘sci/’ directory may now be slightly organized by running the commands in Listing A.10, moving all the files relevant to the wavelength calibrations into either the ‘database’ or ‘split_files’ directories.

Listing A.10: Directory cleanup for POLSALT

```
$ cd <OBSDATE>/sci
$ mkdir split_files
$ mv *beam0* *beamE* *arc0* *arcE* split_files/
$ mv *.eps *.cl *.db database/
```

The POLSALT `spectra extraction` is now run. If the POLSALT GUI was closed it should now be reopened using Listing A.1. With the GUI in focus:

- Ensure all directories are still correct.
- Select ‘Spectra extraction’ from the ‘Data reduction step’ drop down menu.
- Check the tick boxes of all wavelength calibrated images to be processed (exclude the arc) in the display box covering the lower half of the GUI.
- Proceed with the reductions by clicking ‘OK’.

The POLSALT `spectra extraction` is interactive and will launch a separate GUI for the background subtraction and spectral extraction (see Figure 3.2). The background and spectral regions to be extracted may be adjusted, noting that adjustments affect both *O*- and *E*-beams. Once both background regions contain no trace and the spectral region fully contains only the science trace, the reduction may be completed by clicking ‘OK’.

⁷The read noise and gain may be determined from http://pysalt.salt.ac.za/proposal_calls/current/ProposalCall.html, specifically Table 6.1 and Table 6.2.

The STOPS `correlate` method may now be run on any ‘joined’ FITS files by running the commands in Listing A.11.

Listing A.11: Running the STOPS `correlate` method

```
$ cd <OBSDATE>/sci
$ conda activate stops
$ python ~/STOPS . correlate <FILE(S)>
```

The POLSALT `raw Stokes calculation`, `final Stokes calculation`, and `results visualisation` may now be completed. For the last time, if the POLSALT GUI was closed it should now be reopened using Listing A.1. With the GUI in focus:

- Ensure all directories are still correct.
- Select ‘Raw Stokes calculation’ from the ‘Data reduction step’ drop down menu.
- Check the tick boxes of all the extracted spectra images to be processed in the display box covering the lower half of the GUI.
- Proceed with the `raw Stokes calculation` by clicking ‘OK’.
- Select ‘Final Stokes calculation’ from the ‘Data reduction step’ drop down menu.
- Check the tick boxes of all the “raw Stokes” images to be processed in the display box covering the lower half of the GUI.
- Proceed with the `Final Stokes calculation` by clicking ‘OK’.
- Select ‘Results visualisation - interactive’ from the ‘Data reduction step’ drop down menu.
- Check the tick boxes of the “final Stokes” image to be visualized in the display box covering the lower half of the GUI.
- Proceed with the `visualisation` by clicking ‘OK’.

The POLSALT `visualisation` is interactive and will launch a separate GUI (See Figure 3.3). The GUI may be used to change the binning and parameters of the plot before saving the plot to a PDF file.

This concludes the minimum working example of the POLSALT reduction process when substituting the POLSALT `wavelength calibrations` with those done in IRAF. Aside from the final results, the file structure after reductions should resemble something akin to that provided in Figure A.2.

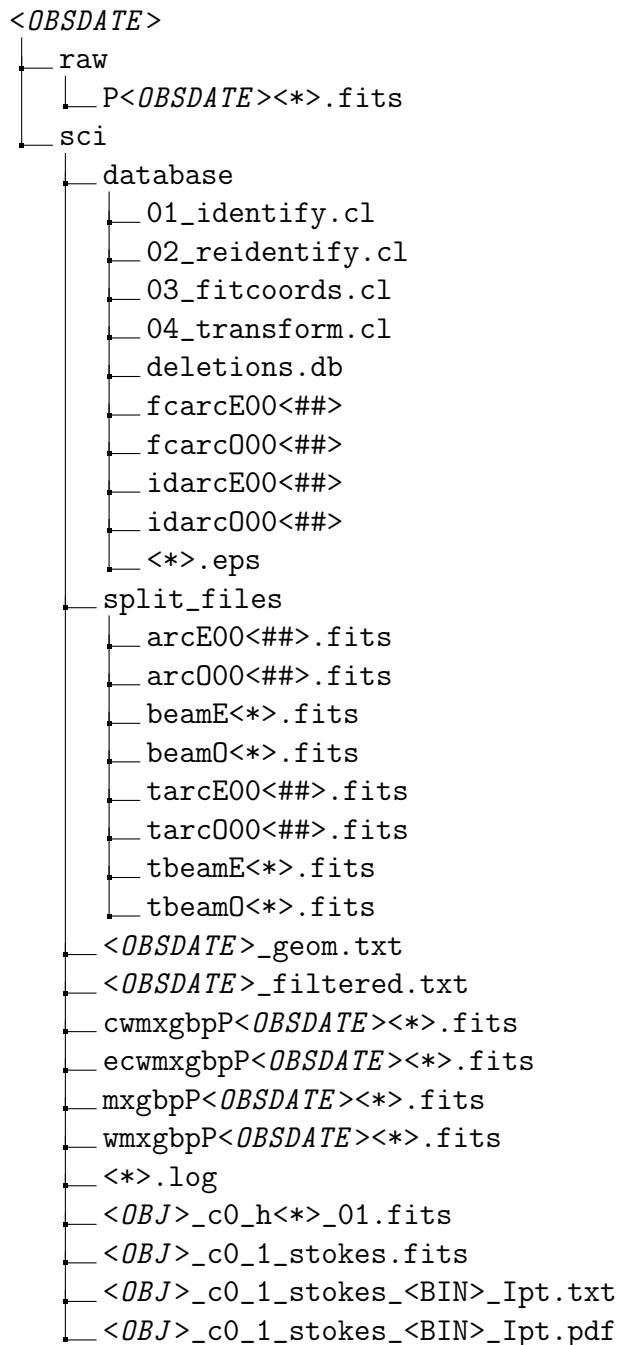


Figure A.2: The typical file structure after completing the reduction process.

Appendix B

STOPS Source Code

This section of Appendix includes all the major STOPS source code files related to the reduction process. Files such as those related to python initialization, testing directories, and other non-essential modules have been excluded for brevity and clarity.

Listing B.1: The source code for `__main__.py`

```
1 #!/usr/bin/env python3
2 # -*- coding: utf-8 -*-
3
4 """
5 The Command Line Interface (CLI) for STOPS, handling argument parsing,
6 class instantiation, processing, and logging.
7
8 Usage
9 -----
10 `python -m STOPS`
11
12 Suffix the above command with `-h` for help on usage.
13
14 Description
15 -----
16 This module is the entry point for the STOPS package. It provides
17 the CLI argument parser for the supplementary tools provided by STOPS.
18 The parser is built using the `argparse` module and is designed to be
19 user-friendly.
20
21 See Also
22 -----
23 IRAF:
24     For more information on the IRAF package, see the IRAF website:
25     https://iraf-community.github.io/
26
27 POLSALT:
28     For more information on the POLSALT pipeline, see the POLSALT website:
29     https://github.com/saltastro/polsalt
30
31 PYSALT:
32     For more information on the PYSALT package, see the PYSALT website:
33     https://astronomers.salt.ac.za/software/pysalt-documentation/
34
35 PYRAF:
36     For more information on the PYRAF package, see the PYRAF website:
37     https://pyraf.readthedocs.io/en/latest/
38
39 """
40
41 # MARK: Imports
```

```

42 import sys
43 import argparse
44 import logging
45 from pathlib import Path
46
47 from STOPS import __version__
48 from STOPS import Split, Join, CrossCorrelate, Skylines
49 from STOPS.utils import ParserUtils as Parser
50 from STOPS.utils.Constants import SPLIT_ROW, PREFIX, PARSE, SAVE_CORR, SAVE_SKY
51
52
53 # MARK: Constants
54 PROG = "STOPS"
55 DESCRIPTION = """
56 Supplementary TOols for Polsalt Spectropolarimetry (STOPS) is a
57 collection of supplementary tools created for SALT's POLSALT pipeline,
58 allowing for wavelength calibrations with IRAF. The tools provide
59 support for splitting and joining polsalt formatted data as well as
60 cross correlating complementary polarimetric beams.
61
62 Pre-release Reference:
63 DOI: 10.22323/1.401.0056
64 """
65
66
67 # MARK: Universal Parser
68 parser = argparse.ArgumentParser(
69     prog=PROG,
70     description=DESCRIPTION,
71     formatter_class=argparse.RawDescriptionHelpFormatter,
72 )
73 parser.add_argument(
74     "-V",
75     "--version",
76     action="version",
77     version=f"%(prog)s as of {__version__}",
78 )
79 parser.add_argument(
80     "-v",
81     "--verbose",
82     action="count",
83     default=PARSE['VERBOSE'],
84     help=(
85         "Counter flag which enables and increases verbosity. "
86         "Use -v or -vv for greater verbosity levels."
87     ),
88 )
89 parser.add_argument(
90     "-l",
91     "--log",
92     action="store",
93     type=Parser.parse_logfile,
94     help=(
95         "Filename of the logging file. "
96         "File is created if it does not exist. Defaults to None."
97     ),
98 )
99 parser.add_argument(
100    "data_dir",
101    action="store",
102    nargs="?",
103    default=PARSE['DATA_DIR'],
104    type=Parser.parse_path,
105    help=(
106        "Path of the directory which contains the working data. "
107        f"Defaults to the cwd -> '{PARSE['DATA_DIR']}' (I.E. '.')."
108    ),
109 )
110
111
112 # MARK: Split\Join Parent
113 split_join_args = argparse.ArgumentParser(add_help=False)
114 split_join_args.add_argument(

```

```

115     "-n",
116     "--no_arc",
117     action="store_true",
118     help="Flag to exclude arc files from processing.",
119 )
120 split_join_args.add_argument(
121     "-s",
122     "--split_row",
123     default=SPLIT_ROW,
124     type=int,
125     help=(
126         "Row along which the O and E beams are split. "
127         f"Defaults to polsalt's default -> {SPLIT_ROW}."
128     ),
129 )
130 split_join_args.add_argument(
131     "-p",
132     "--save_prefix",
133     nargs=2,
134     default=PREFIX,
135     help=(
136         "Prefix appended to the filenames, "
137         "with which the O and E beams are saved. "
138         f"Defaults to {PREFIX}."
139     ),
140 )
141
142
143 # MARK: Corr.\Sky. Parent
144 corr_sky_args = argparse.ArgumentParser(add_help=False)
145 corr_sky_args.add_argument(
146     "filenames",
147     action="store",
148     nargs="+",
149     type=Parser.parse_file,
150     help=(
151         "File name(s) of FITS file(s) to be processed."
152         "A minimum of one filename is required."
153     ),
154 )
155 corr_sky_args.add_argument(
156     "-b",
157     "--beams",
158     choices=[ "O", "E", "OE" ],
159     type=str.upper,
160     default=PARSE['BEAMS'],
161     help=(
162         "Beams to process. "
163         f"Defaults to {PARSE['BEAMS']}, but "
164         "may be given 'O', 'E', or 'OE' to "
165         "determine which beams are processed."
166     ),
167 )
168 corr_sky_args.add_argument(
169     "-ccd",
170     "--split_ccd",
171     action="store_false",
172     help=(
173         "Flag to NOT split CCD's. "
174         "Recommended to leave off unless the chip gaps "
175         "have been removed from the data."
176     ),
177 )
178 corr_sky_args.add_argument(
179     "-c",
180     "--continuum_order",
181     type=int,
182     default=PARSE['CONT_ORD'],
183     dest="cont_ord",
184     help=(
185         "Order of continuum to remove from spectra. "
186         "Higher orders recommended to remove most variation, "
187         "leaving only significant features."

```

```

188     ),
189 )
190 corr_sky_args.add_argument(
191     "-p",
192     "--plot",
193     action="store_true",
194     help="Flag for additional plot outputs.",
195 )
196
197
198 # MARK: Create subparser modes
199 subparsers = parser.add_subparsers(
200     dest="mode",
201     help="Operational mode of supplementary tools",
202 )
203
204
205 # MARK: Split Subparser
206 split_parser = subparsers.add_parser(
207     "split",
208     aliases=["s"],
209     help="Split mode",
210     parents=[split_join_args],
211 )
212 # 'children' split args here
213 # Change `split` defaults here
214 split_parser.set_defaults(
215     mode="split",
216     func=Split,
217 )
218
219
220 # MARK: Join Subparser
221 join_parser = subparsers.add_parser(
222     "join",
223     aliases=["j"],
224     help="Join mode",
225     parents=[split_join_args],
226 )
227 # 'children' join args here
228 join_parser.add_argument(
229     "-c",
230     "--coefficients",
231     dest="solutions_list",
232     nargs='*',
233     type=Parser.parse_file,
234     help=(
235         "Custom coefficients to use instead of the `IRAF` fitcoords "
236         "database. Use as either '-c <o_solution> <e_solution>' or "
237         "a regex descriptor '-c <*solution*extention>'."
238     ),
239 )
240 # Change `join` defaults here
241 join_parser.set_defaults(
242     mode="join",
243     func=Join,
244 )
245
246
247 # MARK: Correlate Subparser
248 corr_parser = subparsers.add_parser(
249     "correlate",
250     aliases=["x"],
251     help="Cross correlation mode",
252     parents=[corr_sky_args],
253 )
254 # 'children' correlate args here
255 corr_parser.add_argument(
256     "-o",
257     "--offset",
258     type=int,
259     default=PARSE['OFFSET'],
260     help=(

```

```

261     "Introduces an offset when correcting for "
262     "known offset in spectra or for testing purposes. "
263     f"Defaults to {PARSE['OFFSET']}. "
264     "(For testing, not used during regular operation.)"
265 ),
266 )
267 corr_parser.add_argument(
268     "-s",
269     "--save_prefix",
270     action="store",
271     nargs="?",
272     type=lambda path: Path(path).expanduser().resolve(),
273     const=SAVE_CORR,
274     help=(
275         "Prefix used when saving plot. "
276         "Excluding flag does not save output plot, "
277         f"flag usage of option uses default prefix {SAVE_CORR}, "
278         "and a provided prefix overwrites default prefix."
279     ),
280 )
281 # Change `correlate` defaults here
282 corr_parser.set_defaults(
283     mode="correlate",
284     func=CrossCorrelate,
285 )
286
287
288 # MARK: Skyline Subparser
289 sky_parser = subparsers.add_parser(
290     "skylines",
291     aliases=["sky"],
292     help="Sky line check mode",
293     parents=[corr_sky_args],
294 )
295 # 'children' skyline args here
296 sky_parser.add_argument(
297     "-t",
298     "--transform",
299     action="store_false",
300     help=(
301         "Flag to force transform images. "
302         "Recommended to use only when input image(s) "
303         "are prefixed 't' but are not yet transformed."
304     ),
305 )
306 sky_parser.add_argument(
307     "-s",
308     "--save_prefix",
309     action="store",
310     nargs="?",
311     type=lambda path: Path(path).expanduser().resolve(),
312     const=SAVE_SKY,
313     help=(
314         "Prefix used when saving plot. "
315         "Excluding flag does not save output plot, "
316         f"flag usage of option uses default prefix {SAVE_SKY}, "
317         "and a provided prefix overwrites default prefix."
318     ),
319 )
320 # Change `skylines` defaults here
321 sky_parser.set_defaults(
322     mode="skyline",
323     func=Skylines,
324 )
325
326
327 # MARK: Keyword Clean Up
328 args = parser.parse_args()
329
330 if len(sys.argv) == 1:
331     parser.print_help(sys.stderr)
332     sys.exit(2)
333

```

```
334 args.verbose = Parser.parse_loglevel(args.verbose)
335
336 if 'log' in args and args.log not in ['', None]:
337     args.log = args.data_dir / args.log
338
339 if "filenames" in args:
340     args.filenames = Parser.flatten(args.filenames)
341
342 if "solutions_list" in args and isinstance(args.solutions_list, list):
343     args.solutions_list = Parser.flatten(args.solutions_list)
344
345 # MARK: Begin logging
346 logging.basicConfig(
347     filename=args.log,
348     format="%(asctime)s - %(module)s - %(levelname)s - %(message)s",
349     datefmt="%Y-%m-%d %H:%M:%S",
350     level=args.verbose,
351 )
352
353 # MARK: Call Relevant Class(Args)
354 logging.debug(f"Argparse namespace: {args}")
355 logging.info(f"Mode:{args.mode}")
356 args.func(**vars(args)).process()
357
358
359 # Confirm all processes completed and exit without error
360 logging.info("All done! Come again!\n")
```

Listing B.2: The source code for `split.py`

```

1 #!/usr/bin/env python3
2 # -*- coding: utf-8 -*-
3
4 """Module for splitting ``polsalt`` FITS files."""
5
6 # MARK: Imports
7 import os
8 import sys
9 import logging
10 from copy import deepcopy
11 from pathlib import Path
12
13 import numpy as np
14 from astropy.io import fits as pyfits
15
16 from STOPS.utils.SharedUtils import find_files, find_arc
17 from STOPS.utils.Constants import SAVE_PREFIX, CROP_DEFAULT, SPLIT_ROW
18
19
20 # MARK: Split Class
21 class Split:
22
23     #-----split0-----
24
25     """
26     The `Split` class allows for the splitting of `polsalt` FITS files
27     based on the polarization beam. The FITS files must have basic
28     `polsalt` pre-reductions already applied (`mzgbp...` FITS files).
29
30     Parameters
31     -----
32     data_dir : str / Path
33         The path to the data to be split
34     fits_list : list[str], optional
35         A list of pre-reduced `polsalt` FITS files to be split
36         within `data_dir`.
37         (The default is None, `Split` will search for `mzgbp*.fits` files)
38     split_row : int, optional
39         The row along which to split the data of each extension in
40         the FITS file.
41         (The default is SPLIT_ROW (See Notes), the SALT RSS CCD's middle row)
42     no_arc : bool, optional
43         Decides whether the arc frames should be recombined.
44         (The default is False, `polsalt` has no use for the arc after
45         wavelength calibrations)
46     save_prefix : dict[str, list[str]], optional
47         The prefix with which to save the O & E beams.
48         Setting `save_prefix` = ``None`` does not save the split O & E beams.
49         (The default is SAVE_PREFIX (See Notes))
50
51     Attributes
52     -----
53     arc : str
54         Name of arc FITS file within `data_dir`.
55         `arc` = `""` if `no_arc` or not detected in `data_dir`.
56     o_files, e_files : list[str]
57         A list of the `O`- and `E`-beam FITS file names.
58         The first entry is the arc file if `arc` defined.
59     data_dir
60     fits_list
61     split_row
62     save_prefix
63
64     Methods
65     -----
66     split_file(file: os.PathLike)
67         -> tuple[astropy.io.fits.HDUList]
68         Handles creation and saving the separated FITS files
69     split_ext(hdulist: astropy.io.fits.HDUList, ext: str = 'SCI')
70         -> astropy.io.fits.HDUList
71         Splits the data in the `ext` extension along the `split_row`
```

```

72     crop_file(hdulist: astropy.io.fits.HDUList, crop: int = CROP_DEFAULT)
73         -> tuple[numpy.ndarray]
74         Crops the data along the edge of the frame, that is,
75         `O`-beam cropped as [crop:], and
76         `E`-beam cropped as [: - crop].
77     update_beam_lists(o_name: str, e_name: str)
78         -> None
79         Updates `o_files` and `e_files`.
80     save_beam_lists(file_suffix: str = 'frames')
81         -> None
82         Creates (Overwrites if exists) and writes the `o_files` and `e_files`
83         to files named `o_{file_suffix}` and `e_{file_suffix}`, respectively.
84     process()
85         -> None
86         Calls `split_file` and `save_beam_lists` on each file in `fits_list`
87         for automation.
88
89     Other Parameters
90     -----
91     **kwargs : dict
92         keyword arguments. Allows for passing unpacked dictionary to
93         the class constructor.
94
95     Notes
96     -----
97     Constants Imported (See utils.Constants):
98         SAVE_PREFIX
99         CROP_DEFAULT
100        SPLIT_ROW
101
102    #-----split1-----
103
104    # MARK: Split init
105    def __init__(
106        self,
107        data_dir: Path,
108        fits_list: list[str] = None,
109        split_row: int = SPLIT_ROW,
110        no_arc: bool = False,
111        save_prefix: Path | None = None,
112        **kwargs,
113    ) -> None:
114        self.data_dir = data_dir
115        self.fits_list = find_files(
116            data_dir=data_dir,
117            filenames=fits_list,
118            prefix="mxgbp",
119            ext="fits"
120        )
121        self.split_row = split_row
122        self.save_prefix = SAVE_PREFIX
123        if isinstance(save_prefix, dict):
124            self.save_prefix = save_prefix
125
126        self.arc = "" if no_arc else find_arc(self.fits_list)
127        self.o_files = []
128        self.e_files = []
129
130        logging.debug(f"__init__\n{repr(self)}")
131
132    return
133
134
135    def __repr__(self) -> str:
136        template = (
137            "Split(\n"
138            f"\tdata_dir={self.data_dir},\n"
139            f"\tfits_list=[\n\t\t{''}\n\t\t'.join("
140            map(str, self.fits_list)
141            )}\n\t],\n"
142            f"\tsplit_row={self.split_row},\n"
143            f"\tno_arc={self.arc == ''},\n"
144            f"\tsave_prefix={self.save_prefix}\n"

```

```

145         ") \n"
146     )
147
148     return template
149
150 # MARK: Split Files
151 def split_file(
152     self,
153     file: os.PathLike
154 ) -> tuple[pyfits.HDUList, pyfits.HDUList]:
155     """
156     Split the single FITS file into separated `O`- and `E`- FITS files.
157
158     Parameters
159     -----
160     file : os.PathLike
161         The name of the FITS file to be split.
162
163     Returns
164     -----
165     tuple[astropy.io.fits.HDUList, astropy.io.fits.HDUList]
166         Tuple containing the split O and E beam HDULists.
167     """
168
169     # Create empty HDUList
170     o_beam = pyfits.HDUList()
171     e_beam = pyfits.HDUList()
172
173     # Open file and split O & E beams
174     with pyfits.open(file) as hdul:
175         o_beam.append(hdul["PRIMARY"].copy())
176         e_beam.append(hdul["PRIMARY"].copy())
177
178     # Split specific extention
179     raw_split = self.split_ext(hdul, "SCI")
180
181     # o_beam[0].data = raw_split['SCI'].data[1]
182     # e_beam[0].data = raw_split['SCI'].data[0]
183     o_beam[0].data, e_beam[0].data = self.crop_file(raw_split)
184
185     # Handle prefix and names
186     pref = "arc" if file == self.arc else "beam"
187     o_name = self.save_prefix[pref][0] + file.name[-9:]
188     e_name = self.save_prefix[pref][1] + file.name[-9:]
189
190     # Add split data to O & E beam lists
191     self.update_beam_lists(o_name, e_name, pref == "arc")
192
193     # Handle don't save case
194     if self.save_prefix is None:
195         return o_beam, e_beam
196
197     # Handle save case
198     o_beam.writeto(o_name, overwrite=True)
199     e_beam.writeto(e_name, overwrite=True)
200
201     return o_beam, e_beam
202
203 # MARK: Split extensions
204 def split_ext(
205     self,
206     hdulist: pyfits.HDUList,
207     ext: str = "SCI"
208 ) -> pyfits.HDUList:
209     """
210     Split the data of the specified extension of `hdulist` into
211     its `O`- and `E`- beams.
212
213     Parameters
214     -----
215     hdulist : astropy.io.fits.HDUList
216         The FITS HDUList to be split.
217     ext : str, optional
218         The name of the extension to be split.

```

```

218     (Defaults to 'SCI')
219
220     Returns
221     -----
222     astropy.io.fits.HDUList
223         The HDUList with the split applied.
224     """
225     hdu = deepcopy(hdulist)
226     rows, cols = hdu[ext].data.shape
227
228     # if odd number of rows, strip off the last one
229     rows = int(rows / 2) * 2
230
231     # how far split is from center of detector
232     offset = int(self.split_row - rows / 2)
233
234     # split arc into o/e images
235     ind_rc = np.indices((rows, cols))[0]
236     padbins = (ind_rc < offset) | (ind_rc > rows + offset)
237
238     # Roll split_row to be centre row
239     image_rc = np.roll(hdu[ext].data[:rows, :], -offset, axis=0)
240     image_rc[padbins] = 0.0
241
242     # Split columns equally
243     hdu[ext].data = image_rc.reshape((2, int(rows / 2), cols))
244
245     return hdu
246
247 # MARK: Crop files
248 @staticmethod
249 def crop_file(
250     hdulist: pyfits.HDUList,
251     crop: int = CROP_DEFAULT
252 ) -> tuple[np.ndarray, np.ndarray]:
253     """
254     Crop the data with respect to the `O`/`E` beam.
255
256     Parameters
257     -----
258     hdulist : astropy.io.fits.HDUList
259         The HDUList containing the data to be cropped.
260     crop : int, optional
261         The number of rows to be cropped from the bottom and top
262         of the `O` and `E` beam, respectively.
263         (Defaults to 40)
264
265     Returns
266     -----
267     tuple[numumpy.ndarray, np.ndarray]
268         Tuple containing the cropped O and E beam data arrays.
269     """
270     o_data = hdulist["SCI"].data[1, 0:-crop]
271     e_data = hdulist["SCI"].data[0, crop:]
272
273     return o_data, e_data
274
275 # MARK: Update beam lists
276 def update_beam_lists(
277     self,
278     o_name,
279     e_name,
280     arc: bool = True
281 ) -> None:
282     """
283     Update the `o_files` and `e_files` attributes.
284
285     Parameters
286     -----
287     o_name : str
288         The filename of the O beam.
289     e_name : str
290         The filename of the E beam.

```

```

291     arc : bool, optional
292         Indicates whether the first entry should be the arc frame.
293         (Defaults to True)
294
295     Returns
296     -----
297     None
298
299     """
300     if arc:
301         self.o_files.insert(0, o_name)
302         self.e_files.insert(0, e_name)
303     else:
304         self.o_files.append(o_name)
305         self.e_files.append(e_name)
306
307     return
308
309 # MARK: Save beam lists
310 def save_beam_lists(self, file_suffix: str = 'frames') -> None:
311     with open(f"o_{file_suffix}", "w+") as f_o, \
312          open(f"e_{file_suffix}", "w+") as f_e:
313         for i, j in zip(self.o_files, self.e_files):
314             f_o.write(i + "\n")
315             f_e.write(j + "\n")
316
317     return
318
319 # MARK: Process all Listed Images
320 def process(self) -> None:
321     """
322     Process all FITS images stored in the `fits_list` attribute
323
324     Returns
325     -----
326     None
327     """
328     for target in self.fits_list:
329         logging.debug(f"Processing {target}")
330         self.split_file(target)
331
332     self.save_beam_lists()
333
334     return
335
336
337 # MARK: Main function
338 def main(argv) -> None:
339     """Main function."""
340
341     return
342
343
344 if __name__ == "__main__":
345     main(sys.argv[1:])

```

Listing B.3: The source code for `join.py`

```

1 #!/usr/bin/env python3
2 # -*- coding: utf-8 -*-
3
4 """
5 Module for joining the split FITS files with an external wavelength solution.
6 """
7
8 # MARK: Imports
9 import os
10 import sys
11 import logging
12 import re
13 from pathlib import Path
14
15 import numpy as np
16 from numpy.polynomial.chebyshev import chebgrid2d as chebgrid2d
17 from numpy.polynomial.legendre import leggrid2d as leggrid2d
18 from astropy.io import fits as pyfits
19
20 # from lacosmic import lacosmic # Deprecated: ccdproc is ~6x faster
21 from ccdproc import cosmicray_lacosmic as lacosmic
22
23 from STOPS.utils.specpolpy3 import read_wollaston, split_sci
24 from STOPS.utils.SharedUtils import find_files, find_arc
25 from STOPS.utils.Constants import DATADIR, SAVE_PREFIX, SPLIT_ROW, CR_PARAMS
26
27
28 # MARK: Join Class
29 class Join:
30
31     #-----join0-----
32
33     """
34     The `Join` class allows for the joining of previously split files
35     and the appending of an external wavelength solution in the
36     `polsalt` FITS file format.
37
38     Parameters
39     -----
40     data_dir : str / Path
41         The path to the data to be joined
42     database : str, optional
43         The name of the `IRAF` database folder.
44         (The default is "database")
45     fits_list : list[str], optional
46         A list of pre-reduced `polsalt` FITS files to be joined
47         within `data_dir`.
48         (The default is ``None``, `Join` will search for `mxbp*.fits` files)
49     solutions_list: list[str], optional
50         A list of solution filenames from which the wavelength solution
51         is created.
52         (The default is ``None``, `Join` will search for `fc*` files within
53         the `database` directory)
54     split_row : int, optional
55         The row along which the data of each extension in the FITS file
56         was split.
57         Necessary when Joining cropped files.
58         (The default is 517, the SALT RSS CCD's middle row)
59     save_prefix : dict[str, list[str]], optional
60         The prefix with which the previously split `O`- & `E`-beams were saved.
61         Used for detecting if cropping was applied during
62         the splitting procedure.
63         (The default is SAVE_PREFIX (See Notes))
64     verbose : int, optional
65         The level of verbosity to use for the Cosmic ray rejection
66         (The default is 30, I.E. logging.INFO)
67
68     Attributes
69     -----
70     fc_files : list[str]
71         Valid solutions found from `solutions_list`.
```

```

72     custom : bool
73         Internal flag for whether `solutions_list` uses the `IRAF`
74         or a custom format.
75         See Notes for custom solution formatting.
76         (Default (inherited from `solutions_list`) is False)
77     arc : str
78         Deprecated. Name of arc FITS file within `data_dir`.
79     data_dir
80     database
81     fits_list
82     split_row
83     save_prefix
84
85     Methods
86     -----
87     get_solutions(wavlist: list / None, prefix: str = "fc") ->
88         (fc_files, custom): tuple[list[str], bool]
89         Parse `solutions_list` and return valid solution files and if they are
90         non-`IRAF` solutions.
91     parse_solution(fc_file: str, x_shape: int, y_shape: int) ->
92         tuple[dict[str, int], np.ndarray]
93         Loads the wavelength solution file and parses keywords necessary for
94         creating the wavelength extension.
95     join_file(file: os.PathLike) -> None
96         Joins the files,
97         attaches the wavelength solutions,
98         performs cosmic ray cleaning,
99         masks the extension,
100        and checks cropping performed in `Split`.
101        Writes the FITS file in a `polsalt` valid format.
102    check_crop(hdu: pyfits.HDUList, o_file: str, e_file: str) -> int
103        Opens the split `O`- and `E`-beam FITS files and returns the amount of
104        cropping that was performed.
105    process() -> None
106        Calls `join_file` on each file in `fits_list` for automation.
107
108     Other Parameters
109     -----
110     no_arc : bool, optional
111         Deprecated. Decides whether the arc frames should be processed.
112         (The default is False, `polsalt` has no use for the arc after
113         wavelength calibrations)
114     **kwargs : dict
115         keyword arguments. Allows for passing unpacked dictionary to the
116         class constructor.
117
118     Notes
119     -----
120     Constants Imported (See utils.Constants):
121         DATADIR, SAVE_PREFIX, SPLIT_ROW, CR_PARAMS
122
123     Custom wavelength solutions must be formatted containing:
124         `x`, `y`, *coefficients...
125     where a solution are of order (`x` by `y`) and
126     must contain  $x \times y$  coefficients, all separated by newlines.
127     The name of the custom wavelength solution file must contain
128     either "cheb" or "leg" for Chebyshev or Legendre wavelength solutions,
129     respectively.
130
131     Cosmic ray rejection is performed using lacosmic [1]-
132     implemented in ccdproc via astroscrappy [2]|.
133
134     References
135     -----
136     .. [1] van Dokkum 2001, PASP, 113, 789, 1420
137         (article : https://adsabs.harvard.edu/abs/2001PASP..113.1420V)
138     .. [2] https://zenodo.org/records/1482019
139
140     """
141
142     #-----join1-----
143
144     # MARK: Join init

```

```

145     def __init__(  
146         self,  
147         data_dir: Path,  
148         database: str = "database",  
149         fits_list: list[str] = None,  
150         solutions_list: list[Path] = None,  
151         split_row: int = SPLIT_ROW,  
152         no_arc: bool = True,  
153         save_prefix: Path | None = None,  
154         verbose: int = 30,  
155         **kwargs,  
156     ) -> None:  
157         self.data_dir = data_dir  
158         self.database = Path(data_dir) / database  
159         self.fits_list = find_files(  
160             data_dir=self.data_dir,  
161             filenames=fits_list,  
162             prefix="mxgbp",  
163             ext="fits",  
164         )  
165         self.fc_files, self.custom = self.get_solutions(solutions_list)  
166         self.split_row = split_row  
167         self.save_prefix = SAVE_PREFIX  
168         if isinstance(save_prefix, dict):  
169             self.save_prefix = save_prefix  
170  
171         self.no_arc = no_arc  
172         self.arc = find_arc(self.fits_list)  
173  
174         self.verbose = verbose < 30  
175  
176         logging.debug(f"__init__\n{repr(self)}")  
177  
178     return  
179  
180 # MARK: Join repr  
181 def __repr__(self) -> str:  
182     template = (  
183         "Join(\n"  
184         f"\tdata_dir={self.data_dir},\n"  
185         f"\tdatabase={self.database},\n"  
186         f"\tfits_list=[\n\t\t{"'\n\t\t'.join(  
187             map(str, self.fits_list))}\n\t],\n"  
188         f"\tsolutions_list={self.fc_files},\n"  
189         f"\tsplit_row={self.split_row},\n"  
190         f"\tno_arc={self.no_arc},\n"  
191         f"\tsave_prefix={self.save_prefix},\n"  
192         f"\tverbose={self.verbose}\n"  
193         ")\n"  
194     )  
195  
196     return template  
197  
198 # MARK: Find 2D WAV Functions  
199 def get_solutions(  
200     self,  
201     wavlist: list[str] | None,  
202     prefix: str = "fc",  
203     reverse: bool = False,  
204 ) -> tuple[list[str], bool]:  
205     """  
206         Get the list of wavelength solution files.  
207  
208     Parameters  
209     -----  
210     wavlist : list[str] / None  
211         A list of custom wavelength solutions files.  
212         If ``None``, `Join` will search for wavelength solutions  
213         in the `database` directory.  
214     prefix : str, optional  
215         The prefix of the wavelength solution files.  
216         (Defaults to "fc")  
217

```

```

218     reverse : bool, optional
219         Whether to reverse the wavelength solution files.
220         Necessary when `wavlist` ordered ['O', 'E']
221         (Defaults to False)
222
223     Returns
224     -----
225     tuple[list[str], bool]
226         A tuple containing the list of wavelength solutions files and
227         a boolean indicating whether custom solutions were provided.
228
229     """
230     # No custom solutions
231     if not wavlist:
232         # Handle finding solutions
233         ws = []
234         for fl in os.listdir(self.database):
235             if os.path.isfile(
236                 self.database / fl
237             ) and (prefix == fl[0:len(prefix)]):
238                 ws.append(fl)
239
240         if len(ws) != 2:
241             # Handle incorrect number of solutions found
242             msg = (
243                 f"Incorrect amount of wavelength solutions "
244                 f"({len(ws)} files) found in the solution "
245                 f"dir.: {self.database}"
246             )
247             logging.error(msg)
248             raise FileNotFoundError(msg)
249
250         sols = {
251             i: j for i, j in zip(['O', 'E'], sorted(ws, reverse=reverse))
252         }
253         logging.debug(f"get_solutions - Found {sols} in {self.database}")
254
255         return sorted(ws, reverse=reverse), False
256
257     # Custom solution
258     if len(wavlist) >= 2:
259         if len(wavlist) > 2:
260             logging.warning(
261                 f" Too many solutions, only {wavlist[:2]} are considered"
262             )
263             wavlist = wavlist[:2]
264
265         for fl in wavlist:
266             if not os.path.isfile(os.path.join(self.data_dir, fl)):
267                 msg = (
268                     f"{fl} not found in the "
269                     f"data directory {self.data_dir}"
270                 )
271                 logging.error(msg)
272                 raise FileNotFoundError(msg)
273
274         sols = {
275             i: j for i, j in zip(
276                 ['O', 'E'],
277                 sorted(wavlist, reverse=reverse)
278             )
279         }
280         logging.debug(f"get_solutions - Found {sols} in {self.database}")
281
282         return sorted(wavlist, reverse=reverse), True
283
284     # MARK: Parse 2D WAV Function
285     def parse_solution(
286         self,
287         fc_file: str,
288         x_shape: int,
289         y_shape: int
290     ) -> tuple[dict[str, int], np.ndarray]:

```

```

291     """
292     Parse the 2D wavelength solution function from `fc_file`.
293
294     Parameters
295     -----
296     fc_file : str
297         The filename of the wavelength solutions file.
298     x_shape : int
299         The x-order of the 2D solution.
300     y_shape : int
301         The y-order of the 2D solution.
302
303     Returns
304     -----
305     tuple[dict[str, int], np.ndarray]
306         A tuple containing a dictionary of
307         the parameters of the solution function
308         and the function coefficients.
309
310     """
311     fit_params = {}
312
313     if self.custom:
314         # Load coefficients
315         coeff = np.loadtxt(fc_file)
316
317         fit_params["xorder"] = int(coeff[0].astype(int))
318         fit_params["yorder"] = int(coeff[1].astype(int))
319         coeff = coeff[2:]
320
321         f_type = 3
322         if "cheb" in str(fc_file):
323             f_type = 1
324         elif "leg" in str(fc_file):
325             f_type = 2
326         fit_params["function"] = f_type
327
328         fit_params["xmin"], fit_params["xmax"] = 1, x_shape
329         fit_params["ymin"], fit_params["ymax"] = 1, y_shape
330
331     else:
332         # Parse IRAF fc database files
333         file_contents = []
334         with open(self.database / fc_file) as fcfile:
335             for i in fcfile:
336                 file_contents.append(re.sub(r"\n\t\s*", "", i))
337
338         if file_contents[9] != "1.":  # xterms - Cross-term type
339             msg = (
340                 "Cross-term not recognised (always 1 for "
341                 "'FITCOORDS'), redo FITCOORDS or change manually."
342             )
343             raise Exception(msg)
344
345         fit_params["function"] = int(file_contents[6][-1])
346
347         fit_params["xorder"] = round(float(file_contents[7][-1]), None)
348         fit_params["yorder"] = round(float(file_contents[8][-1]), None)
349
350         fit_params["xmin"] = int(file_contents[10][-1])
351         fit_params["xmax"] = x_shape
352         # int(file_contents[11][-1])# stretch fit over x
353         fit_params["ymin"] = int(file_contents[12][-1])
354         fit_params["ymax"] = y_shape
355         # int(file_contents[13][-1])# stretch fit over y
356
357         coeff = np.array(file_contents[14:], dtype=float)
358
359         coeff = np.reshape(
360             coeff,
361             (fit_params["xorder"], fit_params["yorder"])
362         )
363

```

```

364     return fit_params, coeff
365
366     # MARK: Join Files
367     def join_file(self, file: os.PathLike) -> None:
368         """
369             Join the `O`- and `E`-beams, attach the wavelength solutions,
370             perform cosmic ray cleaning, mask the extensions,
371             and checks cropping performed by `Split`.
372             Write the FITS file in a `polsalt` valid format.
373
374         Parameters
375         -----
376         file : os.PathLike
377             The path of the FITS file to be joined.
378
379         See Also
380         -----
381         IRAF - `fitcoords` task
382             https://iraf.net/irafdocs/formats/fitcoords.php,
383         numpy 2D grid functions
384             https://numpy.org/doc/stable/reference/generated/numpy
385             Chebyshev: + '.polynomial.chebyshev.chebgrid2d.html'">'Chebyshev: + '.polynomial.chebyshev.chebgrid2d.html'
386             Legendre: + '.polynomial.legendre.leggrid2d.html'">'Legendre: + '.polynomial.legendre.leggrid2d.html'
387
388         """
389         # Create empty wavelength appended hdu list
390         whdu = pyfits.HDUList()
391
392         # Handle prefix and names
393         pref = "arc" if file == self.arc else "beam"
394         o_file = self.save_prefix[pref][0] + file.name[-9:]
395         e_file = self.save_prefix[pref][1] + file.name[-9:]
396
397         # Open file
398         with pyfits.open(file) as hdu:
399             # Check if file has been cropped
400             cropsize = self.check_crop(hdu, o_file, e_file)
401
402             y_shape = int(hdu["SCI"].data.shape[0] / 2) - cropsize
403             x_shape = hdu["SCI"].data.shape[1]
404
405             # No differences in "PRIMARY" extention header
406             primary_ext = hdu["PRIMARY"]
407             primary_ext.header["HISTORY"] = f"CRCLEAN: {CR_PARAMS}"
408             whdu.append(primary_ext)
409
410             for ext in ["SCI", "VAR", "BPM"]:
411                 whdu.append(pyfits.ImageHDU(name=ext))
412                 whdu[ext].header = hdu[ext].header.copy()
413                 whdu[ext].header["CTYPE3"] = "O,E"
414
415                 # Create empty extention with correct order and format
416                 if ext == "BPM":
417                     whdu[ext].data = np.zeros(
418                         (2, y_shape, x_shape),
419                         dtype="uint8"
420                     )
421                     whdu[ext].header["BITPIX"] = "-uint8"
422                 else:
423                     whdu[ext].data = np.zeros(
424                         (2, y_shape, x_shape),
425                         dtype=">f4"
426                     )
427                     whdu[ext].header["BITPIX"] = "-32"
428
429                 # Fill in empty extentions
430                 if cropsize:
431                     temp_split = split_sci(
432                         hdu,
433                         self.split_row,
434                         ext=ext
435                     )[ext].data
436                     whdu[ext].data[0] = temp_split[0, cropsize:]

```

```

437         whdu[ext].data[1] = temp_split[1, 0:-cropsize]
438
439     else:
440         whdu[ext].data = split_sci(
441             hdu,
442             self.split_row,
443             ext=ext
444         )[ext].data
445     # End of hdu calls, close hdu
446
447     # MARK: Join (Wav. Ext.)
448     whdu.append(pyfits.ImageHDU(name="WAV"))
449     wav_header = whdu["SCI"].header.copy()
450     wav_header["EXTNAME"] = "WAV"
451     wav_header["CTYPE3"] = "O,E"
452     whdu["WAV"].header = wav_header
453
454     whdu["WAV"].data = np.zeros(
455         whdu["SCI"].data.shape,
456         dtype=">f4"
457     )
458
459     for num, fname in enumerate(self.fc_files):
460         params, coeffs = self.parse_solution(
461             fname,
462             x_shape,
463             y_shape
464         )
465
466         if params["function"] == 1: # Function type (1 = chebyshev)
467             # Set wavelength extention values to function
468             whdu["WAV"].data[num] = chebgrid2d(
469                 x=np.linspace(-1, 1, params["ymax"]),
470                 y=np.linspace(-1, 1, params["xmax"]),
471                 c=coeffs,
472             )
473
474         elif params["function"] == 2: # Function type (2 = legendre)
475             # Set wavelength extention values to function
476             whdu["WAV"].data[num] = leggrid2d(
477                 x=np.linspace(-1, 1, params["ymax"]),
478                 y=np.linspace(-1, 1, params["xmax"]),
479                 c=coeffs,
480             )
481
482     else:
483         msg = (
484             "Function type not recognised, please wavelength "
485             "calibrate using either Chebyshev or Legendre."
486         )
487         raise Exception(msg)
488
489     # MARK: Cosmic Ray Cleaning
490     # See utils.Constants for `CR_PARAMS` discussion
491     whdu["SCI"].data[num] = lacosmic(
492         whdu["SCI"].data[num],
493         # contrast=CR_PARAMS['CR_CONTRAST'],
494         # threshold=CR_PARAMS['CR_THRESHOLD'],
495         # neighbor_threshold=CR_PARAMS['CR_NEIGH_THRESH'],
496         # effective_gain=CR_PARAMS['GAIN'],
497         # background=CR_PARAMS['BACKGROUND'],
498         readnoise=CR_PARAMS['READ_NOISE'],
499         gain=CR_PARAMS['GAIN'],
500         gain_apply=CR_PARAMS['GAIN_APPLY'],
501         objlim=CR_PARAMS['OBJ_LIM'],
502         verbose=self.verbose,
503     )[0]
504
505     # MARK: WAV masking
506     # Left & Right Crop
507     whdu["WAV"].data[whdu["WAV"].data[:] < 3_000] = 0.0
508     whdu["WAV"].data[whdu["WAV"].data[:] >= 10_000] = 0.0
509

```

```

510     # Top & Bottom Crop (shift\tilt)
511     rpix_oc, cols, rbin, lam_c = read_wollaston(
512         whdu,
513         DATADIR + "wollaston.txt"
514     )
515
516     drow_oc = (rpix_oc - rpix_oc[:, int(cols / 2)][:, None]) / rbin
517
518     # Cropping as suggested
519     for c, col in enumerate(drow_oc[0]):
520         if np.isnan(col):
521             continue
522
523         if int(col) < 0:
524             whdu["WAV"].data[0, int(col):, c] = 0.0
525         elif int(col) > cropsize:
526             whdu["WAV"].data[0, 0: int(col) - cropsize, c] = 0.0
527
528     for c, col in enumerate(drow_oc[1]):
529         if np.isnan(col):
530             continue
531
532         if int(col) > 0:
533             whdu["WAV"].data[1, 0: int(col), c] = 0.0
534         elif (int(col) < 0) & (abs(int(col)) > cropsize):
535             whdu["WAV"].data[1, int(col) + cropsize:, c] = 0.0
536
537     # MARK: BPM masking
538     whdu["BPM"].data[0] = np.where(
539         whdu["WAV"].data[0] == 0,
540         1,
541         whdu["BPM"].data[0]
542     )
543     whdu["BPM"].data[1] = np.where(
544         whdu["WAV"].data[1] == 0,
545         1,
546         whdu["BPM"].data[1]
547     )
548
549     whdu.writeto(f"w{os.path.basename(file)}", overwrite=True)
550
551     return
552
553     # MARK: Check Crop
554     @staticmethod
555     def check_crop(
556         hdu: pyfits.HDUList,
557         o_file: str,
558         e_file: str
559     ) -> int:
560         """
561             Check if cropping is necessary when joining `O`- and `E`-beams.
562
563             Parameters
564             -----
565             hdu : astropy.io.fits.HDUList
566                 The HDUList to check for cropping.
567             o_file : str
568                 The name of the previously split `O`-beam FITS file.
569             e_file : str
570                 The name of the previously split `E`-beam FITS file.
571
572             Returns
573             -----
574             int
575                 The number of rows which were cropped by `Split`.
576
577             """
578             cropsize = 0
579
580             with pyfits.open(o_file) as o,
581                 pyfits.open(e_file) as e:
582                 o_y = o[0].data.shape[0]

```

```
583         e_y = e[0].data.shape[0]
584
585     if hdu["SCI"].data.shape[0] != (o_y + e_y):
586         # Get crop size, assuming crop same on both sides
587         cropsize = int((hdu["SCI"].data.shape[0] - o_y - e_y) / 2)
588
589     return cropsize
590
591 # MARK: Process all Listed Images
592 def process(self) -> None:
593     """Process all FITS images stored in the `fits_list` attribute"""
594     for target in self.fits_list:
595         logging.debug(f"Processing {target}")
596         self.join_file(target)
597
598     return
599
600
601 def main(argv) -> None:
602     """Main function."""
603
604     return
605
606
607 if __name__ == "__main__":
608     main(sys.argv[1:])
```

Listing B.4: The source code for `cross_correlate.py`

```

1 #!/usr/bin/env python3
2 # -*- coding: utf-8 -*-
3
4 """Module for cross correlating polarization beams."""
5
6 # MARK: Imports
7 import sys
8 import logging
9 import itertools as iters
10 from pathlib import Path
11 from importlib.resources import files
12 from typing import Callable
13
14 import numpy as np
15 from numpy.polynomial import chebyshev
16 import matplotlib.pyplot as plt
17 import matplotlib.axes
18 from astropy.io import fits as pyfits
19 from scipy import signal
20
21 from STOPS.utils.SharedUtils import find_files, continuum
22 from STOPS.utils.Constants import SAVE_CORR, OFFSET
23 import STOPS.utils
24
25 # MARK: Logging init.
26 mpl_logger = logging.getLogger('matplotlib')
27 mpl_logger.setLevel(logging.INFO)
28
29
30 # MARK: Correlate class
31 class CrossCorrelate:
32
33     #-----corr0-----
34
35     """
36     Cross correlate allows for comparing the extensions of multiple
37     FITS files, or comparing the $0$- and $E$-beams of a single FITS file.
38
39     Parameters
40     -----
41     data_dir : str / Path
42         The path to the data to be cross correlated
43     filenames : list[str]
44         The ecwmxgbp*.fits files to be cross correlated.
45         If only one filename is defined, correlation is done against the two
46         polarization beams.
47     split_ccd : bool, optional
48         Decides whether the CCD regions should each be individually
49         cross correlated.
50         (The default is True, which splits the spectrum up into its separate
51         CCD regions)
52     cont_ord : int, optional
53         The degree of a chebyshev to fit to the continuum.
54         (The default is 11)
55     plot : bool, optional
56         Decides whether the continuum fitting should be plotted
57         (The default is False, so no continua plots are displayed)
58     save_prefix : str, optional
59         The name or directory to save the figure produced to.
60         "." saves a default name to the current working. A default name is
61         also used when save_prefix is a directory.
62         (The default is None, I.E. The figure is not saved, only displayed)
63
64     Attributes
65     -----
66     data_dir
67     fits_list
68     beams : str
69         The mode of correlation.
70         'OE' for same file, and 'O' or 'E' for different files but same
71         extension.

```

```

72     ccds : int
73         The number of CCD's in the data.
74         Used to split the CCD's if split_ccd is True.
75     cont_ord : int
76         The degree of the chebyshev to fit to the continuum.
77     can_plot : bool
78         Decides whether the continuum fitting should be plotted
79     offset : int, DEPRECATED
80         The amount the spectrum is shifted, mainly to test the effect of the
81         cross correlation.
82         (The default is 0, I.E. no offset introduced)
83     save_prefix
84     wav_unit : str
85         The units of the wavelength axis.
86         (The default is Angstroms)
87     wav_cdel : int
88         The wavelength increment.
89         (The default is 1)
90     alt : Callable
91         An alternate method of cross correlating the data.
92         (The default is None)
93
94     Methods
95     -----
96     load_file(filename: Path) -> tuple[np.ndarray, np.ndarray, np.ndarray]
97         Loads the data from a FITS file.
98     get_bounds(bpm: np.ndarray) -> np.ndarray
99         Finds the bounds for the CCD regions.
100    remove_cont(spec: list, wav: list, bpm: list, plot_cont: bool) -> None
101        Removes the continuum from the data.
102    correlate(filename1: Path, filename2: Path / None = None) -> None
103        Cross correlates the data.
104    ftcs(filename1: Path, filename2: Path / None = None) -> None
105        Cross correlates the data using the Fourier Transform.
106    plot(spec, wav, bpm, corrdb, lagsdb) -> None
107        Plots the data.
108    process() -> None
109        Processes the data.
110
111    Other Parameters
112    -----
113    offset : int, optional
114        The amount the spectrum is shifted, mainly to test the effect of the
115        cross correlation.
116        (The default is 0, I.E. no offset introduced)
117    **kwargs : dict
118        keyword arguments.
119        Allows for passing unpacked dictionary to the class constructor.
120    ftcs : bool, optional
121        Boolean whether to use Fourier Transform for cross correlation.
122
123    See Also
124    -----
125    scipy:
126        https://docs.scipy.org/doc/scipy/reference/generated/scipy.signal.correlate.html
127        correlation:
128            https://matplotlib.org/stable/users/explain/customizing.html
129        matplotlib custom style:
130            https://matplotlib.org/stable/users/explain/customizing.html
131
132    Notes
133    -----
134    Constants Imported (See utils.Constants):
135        SAVE_CORR:
136            The default save name for the correlation plot.
137        OFFSET:
138            The vertical offset of spectra in the output plot.
139        """
140
141        #-----corr1-----
142
143        # MARK: Correlate init
144        def __init__(


```

```

145     self,
146     data_dir: Path,
147     filenames: list[str],
148     beams: str = "OE",
149     split_ccd: bool = True,
150     cont_ord: int = 11,
151     plot: bool = False,
152     offset: int = 0,
153     save_prefix: Path | None = None,
154     **kwargs,
155 ) -> None:
156     self.data_dir = data_dir
157     self.fits_list = find_files(
158         data_dir=self.data_dir,
159         filenames=filenames,
160         prefix="ecwmxgbp",
161         ext="fits",
162     )
163     self._beams = None
164     self.beams = beams
165
166     self.ccds = 1
167     if split_ccd:
168         # with pyfits.open(self.fits_list[0]) as hdu:
169         #     BPM == 2 near center of CCD (extract version != *_sc)
170         #     self.ccds = sum(hdu["BPM"].data.sum(axis=1)[0] == 2)
171         self.ccds = 3
172
173     self.cont_ord = cont_ord
174     self.can_plot = plot
175
176     self.offset = offset
177     if offset != 0:
178         # logging.warning("'offset' is only for testing.")
179         # # Add an offset to the spectra to test cross correlation
180         # self.spec1 = np.insert(
181         #     self.spec1, [0] * offset, self.spec1[:, :offset], axis=-1
182         # )[ :, : self.spec1.shape[-1]]
183
184     err_msg = "Offset deprecated after testing finalized."
185     logging.error(err_msg)
186     raise DeprecationWarning(err_msg)
187
188     self.save_prefix = save_prefix
189     # Handle directory save name
190     if self.save_prefix and self.save_prefix.is_dir():
191         self.save_prefix /= SAVE_CORR
192     logging.warning((
193         f"Correlation save name resolves to a directory. "
194         f"Saving under {self.save_prefix}"
195     ))
196
197     self.wav_unit = "$\\AA$"
198     self.wav_cdelt = 1
199
200     self.alt = self.ftcs if kwargs.get("ftcs") else None
201
202     logging.debug(f"__init__ - \n{repr(self)}")
203
204     return
205
206 # MARK: Correlate repr
207 def __repr__(self) -> str:
208     template = (
209         "CrossCorrelate(\n"
210         f"\tdata_dir={self.data_dir},\n"
211         f"\tfits_list=[\n\t\t{'\n\t\t'.join(
212             map(str, self.fits_list)
213         )}\n\t],\n"
214         f"\tbeams={self._beams},\n"
215         f"\tsplit_ccd={self.ccds},\n"
216         f"\tcont_ord={self.cont_ord},\n"
217         f"\tplot={self.can_plot},\n"

```

```

218         f"\t\t\t\toffset={self.offset},\n"
219         f"\t\t\t\ttsave_prefix={self.save_prefix},\n"
220         f"\t\t\t\ttwav_unit={self.wav_unit},\n"
221         f"\t\t\t\ttwav_cdelt={self.wav_cdelt},\n"
222         f"\t\t\t\ttalt={self.alt},\n"
223         ")\n"
224     )
225
226     return template
227
228 # MARK: Beams property
229 @property
230 def beams(self) -> str:
231     return self._beams
232
233 @beams.setter
234 def beams(self, mode: str) -> None:
235     if mode not in ['O', 'E', 'OE']:
236         err_msg = f"Correlation mode '{mode}' not recognized."
237         logging.error(err_msg)
238         raise ValueError(err_msg)
239
240     self._beams = mode
241
242     return
243
244 # MARK: Load file
245 def load_file(
246     self,
247     filename: Path
248 ) -> tuple[np.ndarray, np.ndarray, np.ndarray]:
249     """
250         Load the data from a FITS file.
251
252     Parameters
253     -----
254     filename : Path
255         The name of the FITS file to load.
256
257     Returns
258     -----
259     tuple[np.ndarray, np.ndarray, np.ndarray]
260         The spectrum, wavelength, and bad pixel mask.
261
262     """
263
264     # Open HDU
265     with pyfits.open(filename) as hdul:
266         spec = hdul["SCI"].data.sum(axis=1)
267         wav = (
268             np.arange(spec.shape[-1])
269             * hdul["SCI"].header["CDELT1"]
270             + hdul["SCI"].header["CRVAL1"]
271         )
272         wav = np.array((wav, wav))
273         bpm = hdul["BPM"].data.sum(axis=1)
274
275         self.wav_cdelt = float(hdul["SCI"].header["CDELT1"])
276
277         if hdul["SCI"].header["CTYPE1"] != 'Angstroms':
278             self.wav_unit = hdul["SCI"].header["CTYPE1"]
279
280     return spec, wav, bpm
281
282 # MARK: Get bounds
283 def get_bounds(self, bpm: np.ndarray, gap_length: int = 30) -> np.ndarray:
284     """
285         Find the bounds for a file based on the CCD count.
286
287     Parameters
288     -----
289     bpm : np.ndarray
290         The bad pixel mask.

```

```

291     gap_length : int, optional
292         The minimum length of a gap to be considered a CCD region.
293         (Defaults to 30)
294
295     Returns
296     -----
297     np.ndarray
298         The bounds for the CCD regions.
299
300     """
301     if not 0 <= gap_length <= 60:
302         msg = f"An invalid gap length of {gap_length} was encountered."
303         logging.error(msg)
304         raise ValueError(msg)
305
306     # Check if get_bounds is needed
307     if self.ccds == 1:
308         return np.array(
309             [(0, bpm[0].shape[-1]), (0, bpm[1].shape[-1])],
310         ).astype(int)
311
312     # bounds.shape -> (0/E, CCD's, low./up. bound)
313     bounds = np.zeros((2, self.ccds, 2))
314
315     # Check if `BPM` contains any `2`'s
316     if np.any(bpm == 2):
317         for ext, ccd in iter.product(range(2), range(self.ccds)):
318             mid = np.where(bpm[ext] == 2)[0][ccd]
319             ccds = self.ccds * 2
320             bounds[ext, ccd] = (
321                 max(mid - bpm.shape[-1] // ccds, 0),
322                 min(mid + bpm.shape[-1] // ccds, bpm.shape[-1])
323             )
324
325     return bounds.astype(int)
326
327     for ext in range(len(self._beams)):
328
329         # Find min and max vals
330         min_val, max_val = 0, bpm[ext].shape[-1]
331
332         while True:
333             if bpm[ext][min_val] == 0:
334                 break
335             min_val += 1
336
337         while True:
338             if bpm[ext][max_val - 1] == 0:
339                 break
340             max_val -= 1
341
342         # Find ranges of non zero values
343         regions: list[np.ndarray] = np.split(
344             np.where(bpm[ext], min_val: max_val == 1)[0],
345             np.where(
346                 np.diff(np.where(bpm[ext], min_val: max_val == 1)[0]) != 1
347             )[0] + 1
348         )
349
350         # If less than 2 regions, raise error
351         if len(regions) < 2:
352             msg = "Less than 2 regions found in BPM. Returning bounds."
353             logging.error(msg)
354             raise ValueError(msg)
355
356         # Find `regions` longer than `gap_length`
357         regions = [
358             region for region in regions
359             if len(region) >= gap_length
360         ]
361
362         # Ensure 2 regions are found
363         if len(regions) < 2:

```

```

364         logging.debug(
365             "get_bounds - Less than 2 regions found in BPM." +
366             f"Calling get_bounds with gap_length = {gap_length - 10}"
367         )
368         return self.get_bounds(bpm, gap_length - 10)
369
370     elif len(regions) > 2:
371         logging.debug(
372             "get_bounds - More than 2 regions found in BPM. " +
373             f"Calling get_bounds with gap_length = {gap_length + 5}"
374         )
375         return self.get_bounds(bpm, gap_length - 10)
376
377     # Ensure region order correct
378     if regions[0][0] > regions[1][0]:
379         regions = regions[::-1]
380
381     # Assign bounds from regions
382     bounds[ext] = np.array([
383         (min_val, regions[0][0]),
384         (regions[0][-1], regions[1][0]),
385         (regions[1][-1], max_val),
386     ])
387
388     # Get lower and upper bound for each ccd, save to bounds
389     # Lower -> min is zero, Upper -> max is bpm length
390
391     # for ext in range(2):
392     #     cedge = ccdcenter(bpm[ext])
393     #     bounds[ext] = np.array(cedge)
394
395     logging.debug(f"get_bounds - found bounds at \n{bounds.astype(int)}")
396
397     return bounds.astype(int)
398
399     # MARK: Remove Continua
400     def remove_cont(
401         self,
402         spec: np.ndarray,
403         wav: np.ndarray,
404         bpm: np.ndarray,
405         plot_cont: bool
406     ) -> np.ndarray:
407         """
408             Remove the continuum from the data.
409
410         Parameters
411         -----
412         spec : np.ndarray
413             The spectrum to remove the continuum from.
414         wav : np.ndarray
415             The wavelength of the spectrum.
416         bpm : np.ndarray
417             The bad pixel mask.
418         plot_cont : bool
419             Decides whether the continuum fitting should be plotted
420
421         Returns
422         -----
423         spec : np.ndarray
424
425         """
426         # Mask out the bad pixels for fitting continua
427         okwav = np.where(bpm != 1)
428
429         # Define continua
430         ctm = continuum(
431             wav[okwav],
432             spec[okwav],
433             deg=self.cont_ord,
434             plot=plot_cont,
435         )
436

```

```

437     # Normalise spectra
438     spec /= chebyshev.chebval(wav, ctm)
439     spec -= 1
440
441     return spec
442
443     # MARK: Correlate
444     def correlate(
445         self,
446         filename1: Path,
447         filename2: Path | None = None,
448         alt: Callable = None
449     ) -> tuple[np.ndarray, np.ndarray, np.ndarray, list[list], list[list]]:
450         """
451             Cross correlates the data.
452
453             Parameters
454             -----
455             filename1 : Path
456                 The name of the first FITS file to cross correlate.
457             filename2 : Path, optional
458                 The name of the second FITS file to cross correlate.
459                 (Defaults to None)
460             alt : Callable, optional
461                 An alternate method of cross correlating the data.
462                 (Defaults to None)
463
464             Returns
465             -----
466             spec, wav, bpm, lagsdb, corrdb:
467                 tuple[np.ndarray, np.ndarray, np.ndarray, list[list], list[list]]
468
469         """
470         # mode: 0E -> '01' & 'E1', 0 -> '01' & '02', E -> 'E1' & 'E2'
471         # Load data
472         spec, wav, bpm = self.load_file(filename1)
473         if filename2 and self._beams != '0E':
474             def unpack(exts, *args):
475                 return [arr[exts] for arr in args]
476
477             if self._beams == '0':
478                 spec[-1], wav[-1], bpm[-1] = unpack(
479                     0, *self.load_file(filename2)
480                 )
481
482             else:
483                 spec[0], wav[0], bpm[0] = spec[-1], wav[-1], bpm[-1]
484                 spec[-1], wav[-1], bpm[-1] = unpack(
485                     -1, *self.load_file(filename2)
486                 )
487
488         bounds = self.get_bounds(bpm)
489
490         logging.debug(
491             f"correlate - data shape:\n\tspec/wav/bpm: {spec.shape}"
492         )
493
494         corrdb = [[] for _ in range(self.ccdfs)]
495         lagsdb = [[] for _ in range(self.ccdfs)]
496         for ccd in range(self.ccdfs):
497             sig = []
498             for ext in range(2):
499                 lb, ub = bounds[ext, ccd]
500
501                 if self.cont_ord > 0:
502                     spec[ext, lb:ub] = self.remove_cont(
503                         spec[ext, lb:ub],
504                         wav[ext, lb:ub],
505                         bpm[ext, lb:ub],
506                         self.can_plot
507                     )
508
509             # Invert BPM (and account for 2); zero bad pixels

```

```

510         sig.append((
511             spec[ext, lb:ub]
512             * abs(bpm[ext, lb:ub].astype(np.int8) * -1 + 1)
513         ))
514
515     # Finally(!!!) cross correlate signals and scale max -> 1
516     corrdb[ccd] = signal.correlate(*sig) if not alt else alt(*sig)
517     corrdb[ccd] /= np.max(corrdb[ccd])
518     # noinspection PyTypeChecker
519     lagsdb[ccd] = signal.correlation_lags(
520         sig[0].shape[-1],
521         sig[1].shape[-1]
522     ) * self.wav_cdel
523
524     return spec, wav, bpm, corrdb, lagsdb
525
526     # MARK: ftcs alternate
527     def ftcs(
528         self,
529         signal1: np.ndarray,
530         signal2: np.ndarray
531     ) -> np.ndarray:
532         """
533             Cross correlates the data using the Fourier Transform.
534
535             Parameters
536             -----
537             signal1 : np.ndarray
538                 The first signal to cross correlate.
539             signal2 : np.ndarray
540                 The second signal to cross correlate.
541
542             Returns
543             -----
544             np.ndarray
545                 The correlation data using the Fourier Transform.
546
547             """
548             logging.debug(
549                 f"ftcs - data shape:\n{tspec.wav/bpm: {signal1.shape}}"
550             )
551
552             # Invert BPM (and account for 2); zero bad pixels
553             ft_spec1 = np.fft.fft(signal1)
554             ft_spec2 = np.fft.fft(signal2)
555
556             if self.can_plot:
557                 plt.plot(ft_spec1)
558                 plt.plot(ft_spec2)
559                 plt.show()
560
561             # Cross correlate signals
562             # ft_spectrum1 * np.conj(ft_spectrum2)
563             corr_entry = signal.correlate(ft_spec1, ft_spec2)
564
565             return np.fft.ifft(corr_entry)
566
567     # MARK: Plot
568     def plot(self, spec, wav, bpm, corrdb, lagsdb) -> None:
569         """
570             Plot the data.
571
572             Parameters
573             -----
574             spec : np.ndarray
575                 The spectrum.
576             wav : np.ndarray
577                 The wavelength.
578             bpm : np.ndarray
579                 The bad pixel mask.
580             corrdb : np.ndarray
581                 The cross correlation data.
582             lagsdb : np.ndarray

```

```

583     The `lags` data.
584
585     Returns
586     -----
587     None
588
589     """
590     plt.style.use([
591         files(STOPS.utils).joinpath('STOPS.mplstyle'),
592         files(STOPS.utils).joinpath('STOPS_correlate.mplstyle'),
593     ])
594     bounds = self.get_bounds(bpm)
595
596     fig, axs = plt.subplots(2, self.ccds, sharey="row")
597
598     if self.ccds == 1:
599         # Convert axs to a 2D array
600         # noinspection PyTypeChecker
601         axs: np.ndarray[matplotlib.axes.Axes] = np.swapaxes(
602             np.atleast_2d(axs), 0, 1
603         )
604
605     # for ext, ccd in iters.product(range(2), range(self.ccds)):
606
607     for ccd in range(self.ccds):
608         axs[0, ccd].plot(
609             lagsdb[ccd],
610             corldb[ccd] * 100,
611             color='C4',
612             label=f"max lag @ {lagsdb[ccd][corldb[ccd].argmax()]} - (bounds[1, ccd,
613             ↪ 0] - bounds[0, ccd, 0])",
614         )
615
616         for ext in range(2):
617             lb, ub = bounds[ext, ccd]
618             logging.debug(f"fl-{ext}: {wav[ext, lb]}:{wav[ext, ub - 1]}")
619
620             axs[1, ccd].plot(
621                 wav[ext, lb:ub],
622                 spec[ext, lb:ub]
623                 * abs(bpm[ext, lb:ub].astype(np.int8) * -1 + 1)
624                 + OFFSET * ext,
625                 label=(
626                     f"${self._beams if self._beams != 'OE' else self._beams[ext]}"
627                     f"_{ext + 1 if self._beams != 'OE' else 1}$$"
628                     f"{'(' + str(OFFSET * ext) + ')') if ext > 0 else '}'"
629                 ),
630             )
631
632         axs[0, 0].set_ylabel("Normalised Correlation\n($\\%$)")
633         for ax in axs[1:, 0]:
634             ax.set_ylabel("Normalised Intensity\n(Counts)")
635             xcol = int(self.ccds != 1)
636             axs[0, xcol].set_xlabel(f"Signal Lag ({self.wav_unit})")
637             axs[-1, xcol].set_xlabel(f"Wavelength ({self.wav_unit})")
638             for ax in axs.flatten():
639                 leg = ax.legend()
640                 leg.set_draggable(True)
641
642     plt.show()
643
644     # Handle do not save
645     if not self.save_prefix:
646         return
647
648     # Handle save
649     fig.savefig(fname=self.save_prefix)
650
651     return
652
653     # MARK: Process all listed images
654     def process(self) -> None:
655         """

```

```

655     Process the data.
656
657     Returns
658     -----
659     None
660
661     """
662     if self._beams != 'OE' and len(self.fits_list) == 1:
663         # change mode to OE with warning
664         logging.warning((
665             f"``{self._beams}`` correlation not possible for "
666             "a single file. correlation `mode` changed to 'OE'."
667         ))
668         self._beams = 'OE'
669
670     # OE `mode` (same file, diff. ext.)
671     if self._beams == 'OE':
672         for fl in self.fits_list:
673             logging.info(f"'OE' correlation of {fl}.")
674             spec, wav, bpm, corr, lags = self.correlate(fl, alt=self.alt)
675             self.plot(spec, wav, bpm, corr, lags)
676
677     return
678
679     # O/E `mode` (diff. files, same ext.)
680     for fl1, fl2 in iter.combinations(self.fits_list, 2):
681         logging.info(f"{self._beams} correlation of {fl1} vs {fl2}.")
682         spec, wav, bpm, corr, lags = self.correlate(fl1, fl2, alt=self.alt)
683         self.plot(spec, wav, bpm, corr, lags)
684
685     return
686
687
688 # MARK: Main function
689 def main(argv) -> None:
690     return
691
692
693 if __name__ == "__main__":
694     main(sys.argv[1:])

```

Listing B.5: The source code for `skylines.py`

```

1 #!/usr/bin/env python3
2 # -*- coding: utf-8 -*-
3
4 """
5 Module for analyzing the sky lines of a wavelength calibrated image.
6 """
7
8 # MARK: Imports
9 import sys
10 import logging
11 from pathlib import Path
12 from importlib.resources import files
13
14 import numpy as np
15 import matplotlib.pyplot as plt
16 import matplotlib.axes
17 from astropy.io import fits as pyfits
18 from scipy import signal
19
20 from STOPS.utils.SharedUtils import find_files, get_arc_lamp
21 from STOPS.utils.Constants import SAVE_SKY, FIND_PEAK_PARAMS
22 import STOPS.data
23 import STOPS.utils
24 import STOPS.data.RSS_arc
25
26
27 # print(
28 #     [logging.getLogger(name) for name in logging.root.manager.loggerDict]
29 # )
30 mpl_logger = logging.getLogger('matplotlib')
31 mpl_logger.setLevel(logging.INFO)
32 pil_logger = logging.getLogger('PIL')
33 pil_logger.setLevel(logging.INFO)
34
35
36 # MARK: Skylines Class
37 class Skylines:
38
39     #-----sky0-----
40
41     """
42     Class representing the Skylines object.
43
44     Parameters
45     -----
46     data_dir : Path
47         The directory containing the data files.
48     filenames : list[str]
49         The list of filenames to be processed.
50     beam : str, optional
51         The beam mode, by default "OE".
52     plot : bool, optional
53         Flag indicating whether to plot the continuum, by default False.
54     save_prefix : Path / None, optional
55         The prefix for saving the data, by default None.
56     **kwargs
57         Additional keyword arguments.
58
59     Attributes
60     -----
61     data_dir : Path
62         The directory containing the data files.
63     fits_list : list[str]
64         The list of fits file paths.
65     beams : str
66         The beam mode.
67     can_plot : bool
68         Flag indicating whether to plot the continuum.
69     save_prefix : Path / None
70         The prefix for saving the data.
71     wav_unit : str

```

```

72     The unit of wavelength.
73
74     Methods
75     -----
76     checkLoad(self, path1: str) -> np.ndarray:
77         Checks and loads the data from the given path.
78     transform(self, wav_sol: np.ndarray, spec: np.ndarray) -> np.ndarray:
79         Transforms the input wavelength and spectral data based on
80         the given wavelength solution.
81     rmvCont(self) -> np.ndarray:
82         Removes the continuum from the spectrum.
83     process(self) -> None:
84         Placeholder method for processing the data.
85
86     See Also
87     -----
88     matplotlib custom style:
89         https://matplotlib.org/stable/users/explain/customizing.html
90
91     Notes
92     -----
93     Constants imported (see utils.Constants):
94         SAVE_SKY:
95             The default save name for the skylines plot.
96         FIND_PEAK_PARAMS:
97             The default parameters for finding peaks in the spectrum.
98
99
100    #-----sky1-----
101
102    # MARK: Skylines init
103    def __init__(
104        self,
105        data_dir: Path,
106        filenames: list[str],
107        beams: str = "OE",
108        split_ccd: bool = False,
109        cont_ord: int = 11,
110        plot: bool = False,
111        transform: bool = True,
112        save_prefix: Path | None = None,
113        **kwargs,
114    ) -> None:
115        self.data_dir = data_dir
116        self.fits_list, self.arc_list = find_files(
117            data_dir=self.data_dir,
118            filenames=filenames,
119            prefix="wmxgbp", # t[o/e]beam
120            ext="fits",
121            sep_arc=True,
122        )
123        if self.arc_list:
124            self.arc_lamp = get_arc_lamp(self.arc_list[0])
125
126        self._beams = None
127        self.beams = beams
128        self.ccds = 1
129        if split_ccd:
130            # See cross_correlate for initial implementation
131            self.ccds = 3
132
133        self.cont_ord = cont_ord
134        self.can_plot = plot
135        self.must_transform = transform
136
137        self.save_prefix = save_prefix
138        # Handle directory save name
139        if self.save_prefix and self.save_prefix.is_dir():
140            self.save_prefix /= SAVE_SKY
141            logging.warning((
142                f"Skylines save name resolves to a directory. "
143                f"Saving under {self.save_prefix}"
144            ))

```

```

145     self.max_difference = 5
146
147     self.wav_unit = "$\\AA$"
148
149     logging.debug(f"__init__\n{repr(self)}")
150
151     return
152
153
154     # MARK: Skyline repr
155     def __repr__(self) -> str:
156         template = (
157             "Skyline(\n"
158             f"\tdata_dir={self.data_dir},\n"
159             f"\tfits_list=[\n\t\tt{{\"\\n\\t\\t\"}}.join(
160                 map(str, self.fits_list)
161             )}\n\t],\n"
162             f"\tbeams={self.beams},\n"
163             f"\tplot={self.can_plot},\n"
164             f"\tsave_prefix={self.save_prefix},\n"
165             f"\twav_unit={self.wav_unit}\n"
166         ")\\n"
167     )
168
169     return template
170
171     # MARK: Beams property
172     @property
173     def beams(self) -> str:
174         return self._beams
175
176     @beams.setter
177     def beams(self, mode: str) -> None:
178         if mode not in ['O', 'E', 'OE']:
179             err_msg = f"Correlation mode '{mode}' not recognized."
180             logging.error(err_msg)
181             raise ValueError(err_msg)
182
183         self._beams = mode
184
185     return
186
187     # MARK: Find Peaks
188     def find_peaks(
189         self,
190         spec: np.ndarray,
191         axis: int | None = None,
192         min_height: float = 0.5,
193         **kwargs,
194     ) -> tuple[list[np.ndarray], list[dict]]:
195         """
196             Finds the peaks in the given spectral data.
197
198             Parameters
199             -----
200             spec : np.ndarray
201                 The spectral data.
202             axis: int / None, optional
203                 The axis along which the peaks are found.
204             min_height : float, optional
205                 The minimum height of the peaks, by default 0.5.
206
207             Returns
208             -----
209             peaks, properties : tuple[list[np.ndarray], list[dict]]
210                 The peaks and their properties.
211             """
212             peaks = []
213             props = []
214             row_means = []
215
216             for ext in range(len(self.beams)):
217                 row_mean = spec[ext] if axis is None else np.mean(

```

```

218         spec[ext], axis=axis
219     )
220
221     peak, prop = signal.find_peaks(
222         row_mean,
223         prominence=min_height * np.max(row_mean),
224         width=0,
225         **kwargs,
226     )
227     peaks.append(peak)
228     props.append(prop)
229     row_means.append(row_mean)
230
231     if self.can_plot:
232         fig, axes = plt.subplots(2, 1)
233         for ext in range(len(self.beams)):
234             axes[ext].plot(
235                 row_means[ext],
236                 label=f"'E' if ext else 'O'")
237             )
238             axes[ext].plot(
239                 peaks[ext],
240                 row_means[ext][peaks[ext]],
241                 "x",
242                 label=f"'E' if ext else 'O'} peaks"
243             )
244             axes[ext].legend()
245         plt.show()
246
247     logging.debug(f"find_peaks - peaks: {[len(i) for i in peaks]}")
248     # logging.debug(
249     #     f"find_peaks - props: {[key for key in props[0].keys()]}"
250     # )
251
252     return peaks, props
253
254     # MARK: Min. of Diff. Matrix
255     @staticmethod
256     def min_diff_matrix(
257         a: np.ndarray,
258         b: np.ndarray,
259         max_diff: int = 100
260     ) -> tuple[np.ndarray, np.ndarray, np.ndarray]:
261         """
262             Find the minimum difference between the elements of two arrays.
263
264             Parameters
265             -----
266             a : np.ndarray
267                 The first 1d array.
268             b : np.ndarray
269                 The second 1d array.
270             max_diff : int, optional
271                 The maximum difference allowed, by default 100.
272
273             Returns
274             -----
275             a : np.ndarray (len(a))
276                 The elements of the first array.
277             min_vals : np.ndarray (len(a))
278                 The minimum difference between the elements of the two arrays.
279             min_idxs : np.ndarray (len(a))
280                 The indices of the minimum difference between
281                 the elements of the two arrays.
282             """
283
284         # Compute the difference matrix using transpose
285         diff = a - b[:, np.newaxis]
286
287         # Find the minimum value in each row (A) of `diff`
288         min_idxs = np.abs(diff).argmin(axis=0)
289
290         min_vals = np.array([diff[j, i] for i, j in enumerate(min_idxs)])
291         # TODO: Recalculate min_val after selecting best min_val and

```

```

291     # TODO: removing the corresponding row/column
292
293     logging.debug(
294         f"min_diff_matrix - min_vals=\n{np.round(min_vals, 2).astype(int)}"
295     )
296     logging.debug(f"min_diff_matrix - min_idxs=\n{min_idxs}")
297
298     max_mask = (min_vals <= max_diff) & (min_vals >= -1 * max_diff)
299
300     return a[max_mask], min_vals[max_mask], min_idxs[max_mask]
301
302 # MARK: Load File Data
303 def load_file_data(
304     self,
305     filename: Path
306 ) -> tuple[np.ndarray, np.ndarray, np.ndarray]:
307     """
308         Loads the data from the given file.
309
310     Parameters
311     -----
312     filename : Path
313         The path to the file to be loaded.
314
315     Returns
316     -----
317     spec, wav, bpm : tuple[np.ndarray, np.ndarray, np.ndarray]
318         The wavelength, spectral, and bad pixel mask data.
319     """
320
321     # Load data from self.beams extension
322     with pyfits.open(filename) as hdul:
323         exts = [0, 1]
324         if len(self.beams) != 2:
325             exts = 0 + int(self.beams == 'E')
326
327         spec2d = np.atleast_3d(hdul["SCI"].data[exts])
328         wav2d = np.atleast_3d(hdul["WAV"].data[exts])
329         bpm2d = np.atleast_3d(hdul["BPM"].data[exts].astype(bool))
330
331         logging.info(
332             f"load_file_data - {filename.name} - shape: {spec2d.shape}"
333         )
334
335     return spec2d, wav2d, bpm2d
336
337 # MARK: Load Sky or Arc Lines
338 @staticmethod
339 def load_lines(
340     filename: str | Path | None = None,
341     dtype: list[tuple] = [(('wav', float), ('flux', float))],
342     skip_header: int = 3,
343     skip_footer: int = 1
344 ) -> np.ndarray:
345     """
346         Loads the sky or arc lines from the given file.
347
348     Parameters
349     -----
350     filename : str / Path / None, optional
351         The path to the file to be loaded.
352         Defaults to loading the skylines from `data/sky.salt`.
353     dtype : list[tuple], optional
354         The data type of the sky lines.
355         (Default is [(('wav', float), ('flux', float))])
356     skip_header : int, optional
357         The amount of lines to skip of the `filenames` header.
358         (Default is 3)
359     skip_footer : int, optional
360         The amount of lines to skip of the `filename`'s footer.
361         (Default is 1)
362
363     Returns
364     -----

```

```

364     sky_lines : np.ndarray['wav', 'flux']
365         The sky lines from the file.
366     """
367     lines: np.ndarray = None
368
369     usecols = None
370     if filename:
371         filename = files(STOPS.data.RSS_arc).joinpath(filename)
372         usecols = (0, 1)
373     else:
374         filename = files(STOPS.data).joinpath('sky.salt')
375
376     lines = np.genfromtxt(
377         filename,
378         dtype=dtype,
379         skip_header=skip_header,
380         skip_footer=skip_footer,
381         usecols=usecols
382     )
383
384     logging.debug(
385         f"load_lines - {filename.name} - shape: {lines.shape}"
386     )
387
388     return lines
389
390 # MARK: Mask Traces
391 def mask_traces(
392     self,
393     spec: np.ndarray,
394     bpm: np.ndarray,
395     max_traces: int = 1,
396     tr_pad: int = 5,
397     bg_margin: int = 10,
398     lr_margins: list[int] = [10, 10],
399     h_min: float = 0.5,
400     h_rel: float = 1 - 0.05,
401 ) -> np.ndarray:
402     """
403         Masks the traces in the bad pixel mask.
404
405     Parameters
406     -----
407     spec : np.ndarray
408         The spectral data.
409     bpm : np.ndarray
410         The bad pixel mask.
411     max_traces : int, optional
412         The maximum number of traces to be masked.
413         (Default is 1)
414     tr_pad : int, optional
415         The amount to pad traces by.
416         (Default is 5)
417     bg_margin : int, optional
418         The margin size for the background.
419         (Default is 10)
420     lr_margins : list[int, int], optional
421         The left and right background margins at the spectrum edge.
422         (Default is [10, 10])
423     h_min: float, optional
424         The minimum height of a detected peak.
425         (Default is 0.5)
426     h_rel: float, optional
427         The relative height for the properties of a detected peak.
428         (Default is 1)
429
430     Returns
431     -----
432     bpm : np.ndarray
433         The updated bad pixel mask.
434     """
435     # Base mask
436     bpm[:, :bg_margin] = True

```

```

437     bpm[:, -bg_margin:] = True
438     bpm[:, :, :lr_margins[0]] = True
439     bpm[:, :, -lr_margins[1]:] = True
440
441     # Get the traces
442     traces, tr_props = self.find_peaks(
443         spec,
444         axis=1,
445         min_height=h_min,
446         rel_height=h_rel
447     )
448
449     for ext in range(len(self.beams)):
450         # Mask the traces
451         for i in range(len(traces[ext][:max_traces])):
452             lb = max(
453                 0,
454                 int(tr_props[ext]['left_ips'][i]) - tr_pad
455             )
456             ub = min(
457                 spec.shape[-1],
458                 int(tr_props[ext]['right_ips'][i]) + tr_pad
459             )
460             bpm[ext, lb: ub] = True
461             # TODO: Relocate targets after initial masking
462
463             logging.info(
464                 f"mask_traces - {min(max_traces, len(traces))} " +
465                 f"of {len(traces)} traces masked."
466             )
467
468     return bpm
469
470     # MARK: Transform Spectra
471     def transform(
472         self,
473         spec: np.ndarray,
474         wav_sol: np.ndarray,
475         row_max: int | None = None,
476         res_plot: bool = False,
477     ) -> tuple[np.ndarray, np.ndarray]:
478         """
479             Transforms the input wavelength and spectral data
480             based on the given wavelength solution.
481
482             Parameters
483             -----
484             spec : np.ndarray
485                 The spectral data.
486             wav_sol : np.ndarray
487                 The wavelength solution.
488             row_max : int, optional
489                 The row along which the spectral data is to be transformed.
490                 (Default is None)
491             res_plot : bool, optional
492                 Flag indicating whether to plot the results.
493                 (Default is False)
494
495             Returns
496             -----
497             spec, wav : np.ndarray
498                 The transformed wavelength and spectral data.
499
500             """
501             # Create arrays to return
502             cs = np.zeros_like(spec)
503             cw = np.zeros_like(wav_sol.mean(axis=1))
504
505             for ext in range(len(self.beams)):
506
507                 if row_max:
508                     avg_max = row_max
509                 else:

```

```

510     # Get middle row (to interpolate the rest of the rows to)
511     avg_max = np.mean(spec[ext], axis=1).argmax()
512
513     # Correct extensions based on wavelength
514     # Get wavelength values at row with most trace
515     cw[ext] = wav_sol[ext, avg_max]
516
517     # Spec ext
518     for row in range(cs.shape[1]):
519         cs[ext, row] = np.interp(
520             cw[ext],
521             wav_sol[ext, row],
522             spec[ext, row]
523         )
524         # f_2d = interpolate.interp2d(
525         #     wav_sol[ext, row],
526         #     np.arange(rows),
527         #     spec[ext],
528         # )
529         # cs[ext] = f_2d(cw[ext], np.arange(rows))
530
531     # Plot results
532     if res_plot:
533         fig, axs = plt.subplots(2, 1, figsize=[20, 4])
534         for ext in range(len(self.beams)):
535             axs[ext].imshow(
536                 cs[ext],
537                 vmax=cs[ext].mean() + 2 * cs[ext].std(),
538                 vmin=cs[ext].mean() - 2 * cs[ext].std()
539             )
540
541             logging.debug(
542                 f"'E' if ext else 'O' Average continuum = " +
543                 f"{np.median(np.median(cs[ext], axis=0)):4.3f}"
544             )
545
546             axx = axs[ext].twinx()
547             axx.hlines(
548                 np.median(np.median(cs[ext], axis=0)),
549                 0,
550                 cs[ext].shape[-1],
551                 colors='black',
552             )
553             axx.plot(
554                 cs[ext].mean(axis=0),
555                 "k",
556                 label=f"mean {'E' if ext else 'O'}"
557             )
558             axx.plot(
559                 np.median(cs[ext], axis=0),
560                 "r",
561                 label=f"median {'E' if ext else 'O'}"
562             )
563             axx.legend()
564             plt.show()
565
566             logging.info(f"transform - {cs.shape} transformed.")
567
568     return cs, cw
569
570     # MARK: Plot
571     def plot(
572         self,
573         spectra,
574         wavelengths,
575         peaks,
576         properties,
577         arc: bool = False,
578     ) -> None:
579         plt.style.use([
580             files(STOPS.utils).joinpath('STOPS.mplstyle'),
581             files(STOPS.utils).joinpath('STOPS_skylines.mplstyle'),
582         ])

```

```

583     plt.rcParams['figure.subplot.hspace'] *= len(self.beams)
584
585     def norm(vals):
586         return (vals - np.min(vals)) / (np.max(vals) - np.min(vals))
587
588     # Load known lines
589     if arc:
590         lines = self.load_lines(filename=self.arc_lamp)
591     else:
592         lines = self.load_lines()
593
594     # noinspection PyTypeChecker
595     lines = lines[
596         (lines['wav'] > wavelengths[1][0][0].min()) &
597         (lines['wav'] < wavelengths[1][0][0].max())
598     ]
599
600     # Create plot for results
601     fig, axs = plt.subplots(2, self.ccds, sharex='col', sharey='row')
602
603     # Convert axs to a 2D array if ccd count is 1
604     if self.ccds == 1:
605         # noinspection PyTypeChecker
606         axs: np.ndarray[matplotlib.axes.Axes] = np.swapaxes(
607             np.atleast_2d(axs),
608             0,
609             1
610         )
611
612     for fl in range(len(self.arc_list if arc else self.fits_list)):
613
614         # set color cycle
615         # private deprecated
616         # color = next(axs[0, 0]._get_lines.prop_cycler)['color']
617         color = axs[0, 0]._get_lines.get_next_color()
618
619         for ext in range(len(self.beams)):
620
621             ccdrange = spectra[1][fl][ext].shape[-1] // self.ccds
622             for ccd in range(self.ccds):
623                 # MARK: plot spectrum
624                 # (transformed)
625                 axs[0, ccd].plot(
626                     wavelengths[1][fl][ext][
627                         ccdrange * ccd:ccdrange * (ccd + 1)
628                     ],
629                     norm(spectra[1][fl][ext][
630                         ccdrange * ccd:ccdrange * (ccd + 1)
631                     ]) * 100 + 10 * ext + 30 * fl,
632                     color=color,
633                     linestyle='dashed' if ext else 'solid',
634                     label=f"${{{self.beams[ext]}}}_{{{fl + 1}}}^{+{10 * ext + 30 * fl}}$"
635                     if ccd == 0 else None,
636                 )
637
638                 # MARK: plot dev
639                 # noinspection PyTypeChecker
640                 sky_wavs, dev, peak_idx = self.min_diff_matrix(
641                     lines['wav'],
642                     wavelengths[1][fl][ext][peaks[1][fl][ext]],
643                     max_diff=self.max_difference,
644                 )
645
646                 # # MARK: width/width_init
647                 # width = properties[1][fl][ext]['widths'][peak_idx]
648                 # width_i = np.zeros_like(width)
649
650                 # sky_i, i_dev, i_idx = self.min_diff_matrix(
651                 #     lines['wav'],
652                 #     wavelengths[0][fl][ext][peaks[0][fl][ext]],
653                 #     max_diff=self.max_difference,
654                 # )

```

```

655     # width_i = np.array([
656     #     properties[0][fl][ext]['widths'][
657     #         np.where(wav == sky_i)[0][0]
658     #     ]
659     #     if wav in sky_i else 1000
660     #     for wav in sky_wavs
661     # ])
662     # width_ratio = (width / width_i) - 1
663     # width_ratio[width_ratio < 0] = 0
664
665     # ylolims = width_ratio > self.max_difference
666     # width_ratio[
667     #     width_ratio > self.max_difference
668     # ] = self.max_difference // 2
669
670     ok = np.where(
671         (sky_wavs >= wavelengths[1][fl][ext].data[
672             ccdrange * ccd
673         ]) &
674         (sky_wavs <= wavelengths[1][fl][ext].data[
675             ccdrange * (ccd + 1)
676         ])
677     )
678     axs[1, ccd].plot(
679         sky_wavs[ok],
680         dev[ok],
681         # yerr=(width_ratio[ok] * 0, width_ratio[ok]),
682         # lolims=ylolims[ok],
683         "."
684         # if ext else "x",
685         # fmt="."
686         # if ext else "x",
687         alpha=0.8,
688         color=color,
689         # markeredgecolor='white',
690         # markeredgewidth=0.5,
691         # label=f"${self.beams[ext]}_{{{fl + 1}}}$",
692     )
693
694     logging.debug(
695         f"plot - RMS: {np.sqrt(np.mean(dev[ok] ** 2)):2.3f}"
696     )
697
698     for ccd in range(self.ccds):
699         ccdrange = spectra[1][0][0].shape[-1] // self.ccds
700
701         # spectrum
702         # noinspection PyTypeChecker
703         ok = np.where(
704             (lines['wav'] >= wavelengths[1][0][0].data[ccdrange * ccd]) &
705             (lines['wav'] <= wavelengths[1][0][0].data[ccdrange * (ccd+1)])
706         )
707         # noinspection PyTypeChecker
708         axs[0, ccd].plot(
709             lines['wav'][ok],
710             lines['flux'][ok] * 0,
711             'x',
712             color='C4',
713             label="\text{sc}\{salt\}\nModel" if ccd == 0 else None,
714         )
715
716         # noinspection PyTypeChecker
717         for x in lines['wav'][ok]:
718             axs[0, ccd].axvline(x, ls='dashed', c='0.7')
719
720         axs[0, 0].set_ylabel("Rel. Intensity ($\\%$)")
721         axs[1, 0].set_ylabel(
722             "Closest Peak ($\\Delta\\lambda$)")
723
724         # for ax in axs[:, 0]:
725         #     ax.legend(loc='upper left', ncols=(fl + 1) * (ext + 1) + 1)
726         leg = fig.legend(
727             loc='center',
728             ncol=min(8, len(spectra[0])) + 1,
729             columnspacing=0.5,
730             bbox_to_anchor=(

```

```

728         np.mean((
729             plt.rcParams['figure.subplot.left'],
730             plt.rcParams['figure.subplot.right']
731         )),
732         np.mean((
733             plt.rcParams['figure.subplot.bottom'],
734             plt.rcParams['figure.subplot.top']
735         )))
736     ),
737 )
738 leg.set_draggable(True)
739 for ax in axs[1, :]:
740     ax.grid(axis='y')
741
742     axs[-1, 0 if self.ccds == 1 else 1].set_xlabel(
743         f"Wavelength ({self.wav_unit})"
744     )
745
746 # plt.tight_layout()
747
748 plt.show()
749
750 # Save results
751 if self.save_prefix:
752     fig.savefig(fname=self.save_prefix)
753
754 return
755
756 # MARK: Process all listed images
757
758 def process(self, arc: bool = False) -> None:
759     files = self.fits_list
760     if arc:
761         files = self.arc_list
762
763     logging.info(f"Processing '{self.beams}' lines.")
764
765     spectra = [[], []]
766     wavs = [[], []]
767     peaks = [[], []]
768     peak_props = [[], []]
769
770     for fl in files:
771         # Load data
772         spec2d, wav2d, bpm2d = self.load_file_data(fl)
773
774         # Mask traces in BPM
775         bpm2d = self.mask_traces(
776             spec2d,
777             bpm2d,
778             max_traces=0,
779             bg_margin=15,
780             h_min=0.05
781         )
782         m_spec2d = np.ma.masked_array(spec2d, mask=bpm2d) # spec2d
783         m_wav2d = np.ma.masked_array(wav2d, mask=bpm2d) # wav2d
784
785         # Initial spectra
786         spec_i = np.mean(m_spec2d, axis=-2)
787         wav_i = np.mean(m_wav2d, axis=-2)
788
789         # Transform data
790         t_spec2d, t_wav = self.transform(
791             m_spec2d,
792             m_wav2d,
793             res_plot=self.can_plot
794         )
795
796         # Final spectra
797         spec_f = np.mean(t_spec2d, axis=-2)
798         wav_f = t_wav
799
800         # Find peaks

```

```
801     peaks_i, props_i = self.find_peaks(
802         spec_i,
803         **FIND_PEAK_PARAMS
804     )
805     peaks_f, props_f = self.find_peaks(
806         spec_f,
807         **FIND_PEAK_PARAMS
808     )
809
810     spectra[0].append([*spec_i])
811     spectra[1].append([*spec_f])
812     wavs[0].append([*wav_i])
813     wavs[1].append([*wav_f])
814     peaks[0].append([*peaks_i])
815     peaks[1].append([*peaks_f])
816     peak_props[0].append([*props_i])
817     peak_props[1].append([*props_f])
818
819     # Plot results
820     self.plot(spectra, wavs, peaks, peak_props, arc=arc)
821
822     if arc:
823         return
824     elif self.arc_list:
825         self.process(arc=True)
826
827     return
828
829
830 # MARK: Main function
831 def main(argv) -> None:
832     """main function."""
833
834     return
835
836
837 if __name__ == "__main__":
838     main(sys.argv[1:])
```

Appendix C

Proceedings

Development of STOPS and it's application to spectropolarimetric observations was presented at conferences, with the following proceedings included here.

- HEASA 2021 - Cooper, van Soelen, and Britto (2022)
- HEASA 2022 - Cooper and van Soelen (2023)

Development of tools for the SALT/RSS spectropolarimetry reductions: application to the blazar 3C 279

Justin Cooper,^{a,*} Brian van Soelen^a and Richard J. Britto^a

^a*University of the Free State,
205 Nelson Mandela Drive, Bloemfontein, South Africa*

E-mail: CooperJ@ufs.ac.za, VanSoelenB@ufs.ac.za

Blazars represent a subset of AGN with relativistic jets, where the direction of the jet lies very close to our line of sight. The highly Doppler boosted emission from the blazar's jet results in high apparent luminosities, and blazars display variability on periods from less than one day up to years. At optical wavelengths, the observed emission of the blazar is a superposition of the polarised non-thermal synchrotron emission, arising from the jet, and the unpolarised thermal emission, arising from the accretion disc, broad line region, torus and host galaxy. Polarimetry observations can serve as an important tool for diagnosing the emission from blazars. The RSS spectrograph, on SALT, can operate in spectropolarimetry mode and is currently being used to undertake spectropolarimetric observations of transient blazar sources. We present additional tools developed to work in conjunction with the current SALT spectropolarimetry reduction pipeline, POLSALT, that aims to streamline the reduction of the SALT polarisation data, including the testing of the wavelength calibration of the individual O and E beams. This was applied to observations of 3C 279 during 2017.

*** *High Energy Astrophysics in Southern Africa (HEASA2021)* ***

*** *13 - 17 September, 2021* ***

*** *Online* ***

*Speaker

1. Introduction

Active Galactic Nuclei (AGN) are the centres of galaxies powered by accretion onto a supermassive black hole. The accretion can power relativistic jets which produce non-thermal emission over multiple wavelengths [1]. Blazars are a subclass of AGN where the direction of the jet lies very close to our line of sight. The highly Doppler boosted emission from the jet results in high apparent luminosities.

At optical wavelengths the observed emission is a superposition of polarised non-thermal synchrotron emission, arising from the jet, and unpolarised thermal emission, arising from the disc, broad line region, torus, and host galaxy [2].

The Robert Stobie Spectrograph (RSS) on the Southern African Large Telescope (SALT) [3] [4] is capable of long-slit spectropolarimetry and has been used to observe blazars in flaring and quiescent states. Data reduction of these observations is processed using the SALT spectropolarimetry reduction package `POLSALT`.¹ `POLSALT` allows for the full reduction, from pre-reduction to the extraction of the target spectra, and the measuring of the polarisation. However, it does not allow for much flexibility with the wavelength calibrations, nor does it provide tools to confirm that the wavelength calibrations lead to similar wavelength calibrated spectra for the O and E beams.

In order to streamline the data reduction we are developing additional tools to work in conjunction with `POLSALT`. We present the overview of this pipeline and demonstrate its application to an observation of 3C 279 taken in 2017.

2. Pipeline Overview

The tools being developed are designed to work in conjunction with `POLSALT` and provide a method to perform the wavelength calibrations using conventional `IRAF` methods [5]. They also allow easy comparisons to be made between the wavelength solutions of the O and E beam, to determine their quality and reliability.

The workflow for reductions is depicted in figure 1. It summarises the steps followed to reduce spectropolarimetric data. Steps that are performed using `POLSALT` are indicated in blue; steps performed by the new tools are shown in orange, while the `IRAF` calibration is shown in green.

If the wavelength solutions are found with `IRAF`, the tools split the multi-extension FITS files into the correct format for `IRAF` and the `IRAF` wavelength calibration is performed. The files are then written back into multi-extension FITS files in the correct format for `POLSALT` to complete the extraction and measurement of polarisation.

The cross-correlation between the O and E beam spectra, named correlation in figure 1, and relative flux calibrations are optional, though correlation is suggested. Relative flux calibrations may be performed assuming a standard and published standard spectra can be acquired. We performed the relative flux calibrations using the `ASTROPY` [6, 7] and `SCIPY`² packages.

¹<https://github.com/saltastro/polsalt>

²<http://www.scipy.org/>

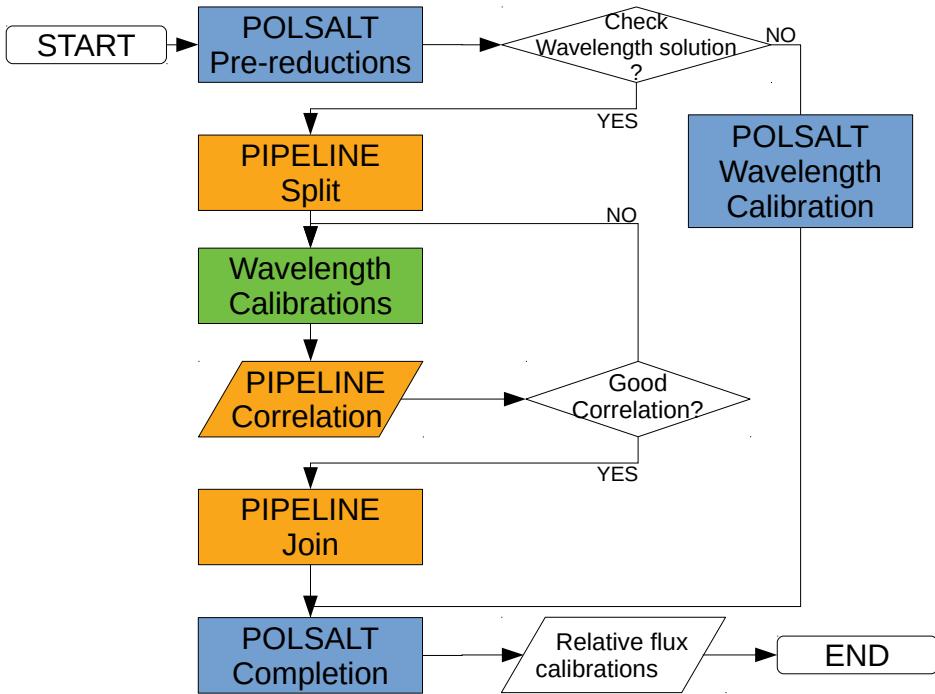


Figure 1: The flow diagram of the SALT RSS spectropolarimetry reductions steps.

POS (HEAS2021) 056

2.1 Pre-reduction

The first step in the pipeline uses **POLSLT** and this performs the pre-reductions, the over-scan subtraction, gain and cross talk corrections, and mosaicing. After the pre-reduction is performed, the next step is to either undertake the wavelength calibration with **POLSLT**, or to split the files into the correct format for **IRAF**.

2.2 FITS split

Once all pre-reductions are complete, the tool **SPLIT** converts the multi-extension FITS files into multiple, single extension O or E beam FITS files. This is done as some **IRAF** tasks do not handle multi-extension FITS files well. The O and E beam frames are cropped to remove any non-exposed rows to accommodate the standard **IRAF** tasks and the FITS files saved with updated header information.

2.3 Wavelength calibration

If the **POLSLT** pipeline is used, the wavelength calibration is performed using the wavelength calibration tools which form part of **PYSALT** [8].

If the files have been formatted for **IRAF**, standard wavelength calibration methods are followed, using the **NOAO** package and the **identify**, **re-identify**, and **fitcoords** tasks. This allows greater control of the resultant wavelength solution.

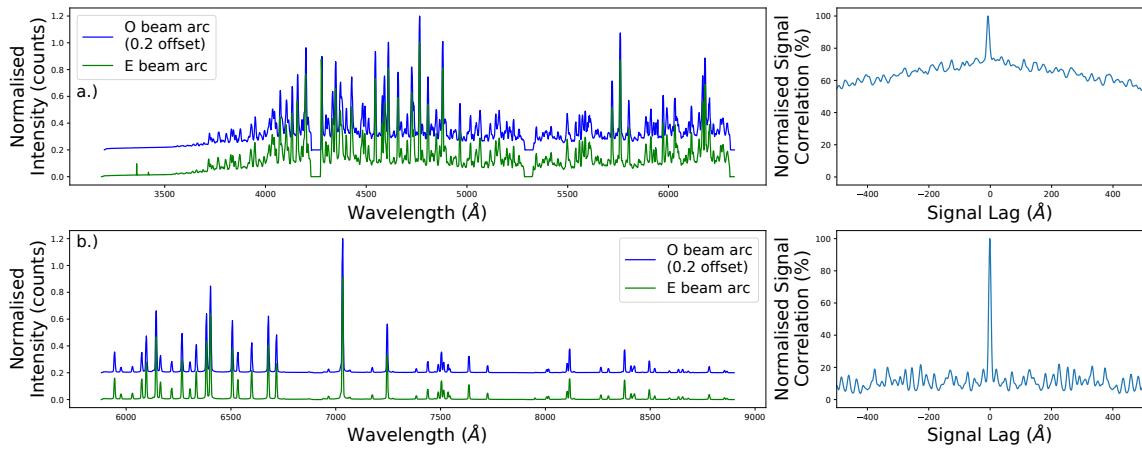


Figure 2: Comparison of the O and E beams for a.) ThAr and b.) Ne arc lamp spectra as well as the cross-correlation of their O and E beams.

2.4 Correlation

Our tool CORRELATE allows a comparison to be made between the O and E beam solutions, by performing a cross-correlation between the two arc or target observations, after the wavelength solution as been applied by `IRAF`. This is important as the polarisation is calculated from the difference between the O and E beam and, therefore, accurate wavelength solutions are required. The comparison between the O and E beams for a ThAr and Ne arc spectra and their cross-correlations, with a near zero lag, are shown in figure 2.a) and figure 2.b), respectively. Figure 2's Signal Lag plots show a clear peak at zero lag but also show an important distinction between noisy and quiet spectra in the non-peaking region.

2.5 FITS join

Once the `IRAF` wavelength calibrations are complete, the tool JOIN performs cosmic-ray cleaning using the LACOSMIC package [9], applies a wollaston correction which accounts for the physical shift of the spectra on the CCD, and then formats the single extension O and E beam FITS files into the multi-extension format required by `POLSALT`. This includes saving the wavelength solution found in `IRAF` as an additional extension. Currently only wavelength solutions fit with a Chebychev function are recognised but further development is planned to recognise Legendre functions.

2.6 Completion - extraction and polarisation measurement

After the wavelength solution is found, the final step is performed using `POLSALT`, which performs the curvature corrections, background subtraction, spectral extraction, and the raw and final stokes calculations. The binning and plotting also form part of the `POLSALT` completion.

3. Application to 3C 279

The blazar 3C 279 was observed in linear spectropolarimetry mode, on 2017 May 17, using the the PG0900 grating with two different grating angles (12.5° and 19.5°). The spectrophotometric

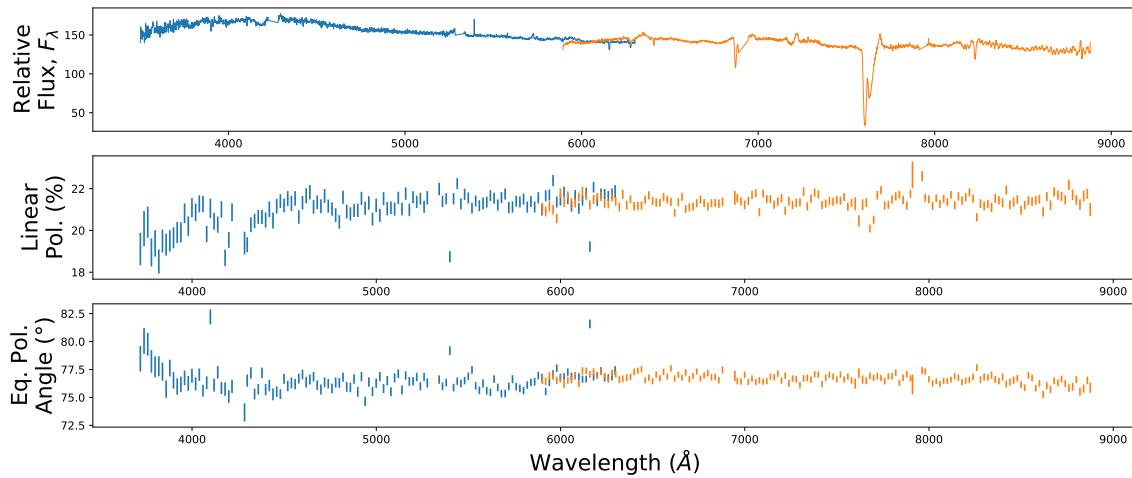


Figure 3: Reduced observations of 3C 279, taken with grating angle 12.5° (blue line) and 19.5° (orange line). Top: relative flux calibrated spectra. Middle: degree of linear polarisation. Bottom: Polarisation angle (relative to instrument).

standard star Hiltner 600 was observed on 2017 February 24, using the same configuration and was used for relative flux calibrations.

The target and standard were reduced following the steps discussed about. The flux correction was performed using `ASTROPY` and `SCIPY` python packages to correct the flux shape.

The results are shown in figure 3, indicating a good overlap between the observations taken at two different grating angles, with the polarisation measurements generally agree well where the observations overlap.

Figure 2 shows the comparison of the O and E beams for this observations. The cross-correlation is very close to zero lag, indicating the wavelength solution is similar for both the O and E beams.

4. Discussion and Conclusions

Creating an independent 2D wavelength solution for the O and E beams using `IRAF` allows for more flexibility in obtaining the wavelength solutions and provides a tool to directly compare the O and E beam solutions. This is useful for difficult cases, such as arc line identification for the PG0300 spectral grating, because there are fewer arc lines in some parts of the spectra and it is more complicated to create a consistent solution across the full wavelength range covered.

The cross-correlation plot cannot only be considered when deciding on whether to proceed onward with the reductions or to return to wavelength calibrations as depicted in figure 1. The comparison of the spectra is as important since the cross-correlation may show zero lag but lines in the spectra could still be mis-identified in both beams. The cross-correlation serves as a method to check the similarity of the wavelength solutions but not the accuracy of said solutions.

The tools developed create an alternative to the `POLSALT` wavelength calibration and add a method to check the consistency of the wavelength solutions, as can be seen with the application to 3C 279. These assist with the wavelength calibration, in particular for observations taken with

the PG0300 and PG0900 RSS gratings, that have been used for numerous observations of blazar sources.

Acknowledgments

The observational data presented was obtained with the Southern African Large Telescope (SALT) under the proposal 2016-2-LSP-001 (PI: D.A.H. Buckley). This work is based on the research supported in part by the National Research Foundation of South Africa (Grant Numbers: 116300).

References

- [1] C.M. Urry and P. Padovani, *Unified Schemes for Radio-Loud Active Galactic Nuclei*, *PASP* **107** (1995) 803 [[astro-ph/9506063](#)].
- [2] M. Böttcher, B. Van Soelen, R.J. Britto, D.A.H. Buckley, J.P. Marais and H. Schutte, *Salt spectropolarimetry and self-consistent sed and polarization modeling of blazars*, *Galaxies* **5** (2017) .
- [3] E.B. Burgh, K.H. Nordsieck, H.A. Kobulnicky, T.B. Williams, D. O'Donoghue, M.P. Smith et al., *Prime focus imaging spectrograph for the southern african large telescope: optical design*, in *Instrument Design and Performance for Optical/Infrared Ground-based Telescopes*, vol. 4841, pp. 1463–1471, International Society for Optics and Photonics, 2003.
- [4] H.A. Kobulnicky, K.H. Nordsieck, E.B. Burgh, M.P. Smith, J.W. Percival, T.B. Williams et al., *Prime Focus Imaging Spectrograph for the Southern African Large Telescope: operational modes*, in *Instrument Design and Performance for Optical/Infrared Ground-based Telescopes*, M. Iye and A.F.M. Moorwood, eds., vol. 4841, pp. 1634 – 1644, International Society for Optics and Photonics, SPIE, 2003, [DOI](#).
- [5] D. Tody, *The IRAF Data Reduction and Analysis System*, in *Instrumentation in astronomy VI*, D.L. Crawford, ed., vol. 627 of *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, p. 733, Jan., 1986, [DOI](#).
- [6] Astropy Collaboration, T.P. Robitaille, E.J. Tollerud, P. Greenfield, M. Droettboom, E. Bray et al., *Astropy: A community Python package for astronomy*, *A&A* **558** (2013) A33 [[1307.6212](#)].
- [7] Astropy Collaboration, A.M. Price-Whelan, B.M. Sipőcz, H.M. Günther, P.L. Lim, S.M. Crawford et al., *The Astropy Project: Building an Open-science Project and Status of the v2.0 Core Package*, *AJ* **156** (2018) 123 [[1801.02634](#)].
- [8] S.M. Crawford, M. Still, P. Schellart, L. Balona, D.A.H. Buckley, G. Dugmore et al., *PySALT: the SALT science pipeline*, in *Observatory Operations: Strategies, Processes, and Systems III*, D.R. Silva, A.B. Peck and B.T. Soifer, eds., vol. 7737 of *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, p. 773725, July, 2010, [DOI](#).

- [9] P.G. Van Dokkum, *Cosmic-ray rejection by laplacian edge detection*, *Publications of the Astronomical Society of the Pacific* **113** (2001) 1420.

POS (HEASA2021) 056

SALT Spectropolarimetric Pipeline Comparisons

J. Cooper^{a,*} and B. van Soelen^a

^a*University of the Free State, Physics Dept.,*

PO Box 339, Bloemfontein, 9300

E-mail: Justin.jb78@gmail.com, vanSoelenB@ufs.ac.za

Blazars are active galactic nuclei (AGN) with jets aligned very closely to our line of sight. The optical emission of blazars is often dominated by the polarised, non-thermal emission arising in the jets, with an underlying unpolarised, thermal emission component arising from the host galaxy, dusty torus, and accretion disk components.

Coupled with multi-wavelength observations, optical spectropolarimetry of blazars during both flaring and quiescent states can be used to disentangle the polarised and unpolarised components in their spectral energy distributions, providing better constraints for the non-thermal particle distribution. To this end, spectropolarimetry of blazars during different states of activity was taken with the Southern African Large Telescope (SALT) using the Robert Stobie Spectrograph (RSS). For RSS spectropolarimetry, reductions are performed using the POLSALT pipeline. In order to streamline the spectropolarimetric reductions, we have implemented supplementary interactive tools which provides additional wavelength calibration to improve the accuracy of the wavelength calibration for the O & E beam.

Here we present a brief overview of the tools and the results for Hiltner 652, a spectropolarimetric standard, as well as results for the blazars 3C 279. The reduced, P_Q and P_U , Stokes parameters of Hiltner 652 show no major deviation from previously published results which reassures us that there is no interference introduced into the Stokes parameter calculations when wavelength calibrations are handled by our supplementary tools. The $\sim 6000 - 9000 \text{ \AA}$ range of 3C 279 shows a notable dip in the normalized spectrum during a period of flaring when compared to epochs of enhanced activity. The degree of polarisation for 3C 279 of 13.2%, 9.5%, and 21.2% for the epochs 2017 March 28, 2017 March 31, and 2017 May 17, respectively, remains fairly constant across the observed wavelength ranges while still varying with the blazars' state of quiescence or flaring.

High Energy Astrophysics in Southern Africa 2022 - HEASA2022
28 September - 1 October 2022
Brandfort, South Africa

*Speaker

1. Introduction

Blazars are a radio-loud sub-set of Active Galactic Nuclei (AGN), where the jet is aligned very closely to our line of sight and whose observed emission varies over time scales from hours to decades [1]. The optical emission observed from blazars is often dominated by the polarised, non-thermal emission arising in the jets, but there is also underlying non-polarised, thermal emission arising from the host galaxy, dusty torus, and accretion disk components [2]. When a blazar varies between flaring and quiescent states, the relative strength of the thermal component to the non-thermal component changes.

The Spectral Energy Distribution (SED) of blazars shows two clear components: a lower energy component produced through leptonic synchrotron processes (covering the radio to the UV/soft X-ray regimes) and a higher energy component (covering the X-ray to the γ -ray regimes) which can be produced through either leptonic or hadronic processes [3]. Optical spectropolarimetry observations, coupled with multi-wavelength observations during both flaring and quiescent states, can be used to disentangle the polarised and non-polarised components in the blazar's SED, providing better constraints for the non-thermal particle distribution [4, 5].

To this end, spectropolarimetric observations of blazars during different states of activity were taken using SALT [6] and the RSS [7]. RSS spectropolarimetry data is reduced with the dedicated POLSALT pipeline [8]. In order to facilitate the blazar observations we have developed supplementary tools to provide a more interactive approach to the wavelength calibration, allowing for improved accuracy for the O & E beam wavelength solutions [9]. Furthermore, spectropolarimetric standards, which were also observed with SALT and the RSS, were reduced alongside the blazar observations to test that the developed tools correctly performed the wavelength calibration. Here we present an overview of the pipeline as well as the preliminary results of a spectropolarimetric standard, namely Hiltner 652, and for the blazar 3C 279.

2. Pipeline Overview

RSS spectropolarimetry data is currently reduced using the POLSALT pipeline [8]. This python package allows for a full reduction from pre-reduction to extracted spectrum, and measured polarisation. However, the wavelength calibration component of the package does not allow for flexibility, nor provide a tool to confirm that the wavelength calibrations of the O & E beams are in agreement. This flexibility is most notably missed when performing calibrations of observations taken using the PG0300 grating¹ due to the very sparse spectral features of the Argon arc lamp used by SALT as well as second order contamination which differs between the O & E beams.

The developed tools work in conjunction with the existing pipeline and provide a method to perform the wavelength calibrations using conventional IRAF [10] methods. Figure 1 shows the typical workflow for data reductions of spectropolarimetric SALT data. Items in blue refer to processing steps completed entirely using POLSALT while items in orange refer to processing steps completed using the developed tools. The single green item refers to the wavelength calibration

¹This grating has been decommissioned and replaced by the PG0700, and is not available for general use as of November 2022.

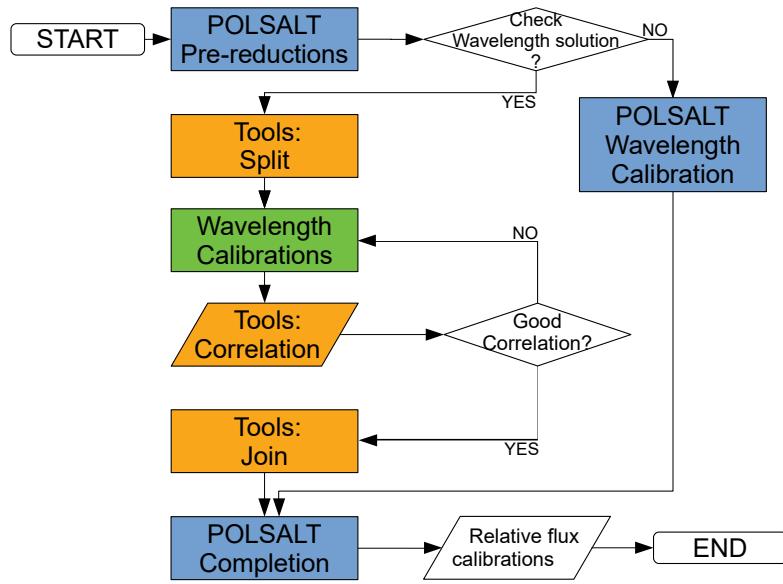


Figure 1: A general workflow for data reductions using a combination of **POLSALT** and our developed tools.

Source	Observation date	Grating angle (°)	Slit width (")	Exposure time (sec)	Wavelength range (Å)
Hiltner 652	2022 June 10	12.878	2	48.5	3350 - 6440
3C 279	2017 March 28	12.5	1.5	480.8	3300 - 6300
3C 279	2017 March 28	19.625	1.5	480.8	5800 - 8800
3C 279	2017 March 31	12.5	1.5	480.8	3300 - 6300
3C 279	2017 March 31	19.625	1.5	480.8	5800 - 8800
3C 279	2017 May 17	12.5	1.5	720.8	3300 - 6300
3C 279	2017 May 17	19.625	1.5	720.8	5800 - 8800

Table 1: All spectropolarimetric observations taken with SALT and the RSS and used as part of development and testing of the supplementary tools.

done in **IRAF**. Additional tools to cross-check that the wavelength calibration is consistent across both beams have also been implemented.²

3. Observations

The spectropolarimetric observations were taken using the RSS mounted on SALT, in **LINEAR** mode. The spectropolarimetric standard, Hiltner 652, was observed in 2022 and the blazar 3C 279, was observed in 2017, as summarized in Table 1. All observations presented here were observed using the PG0900 and were reduced following the reduction workflow as described in Figure 1.

²A further, in-depth, discussion of the developed tools can be found at [9].

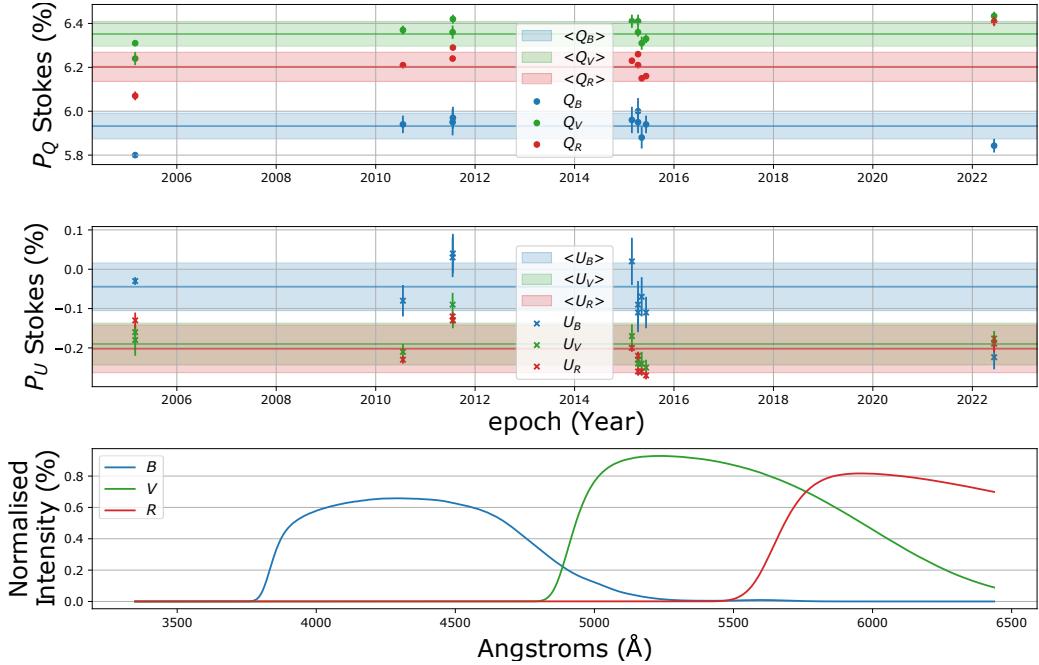


Figure 2: The reduced Q and U Stokes parameters for Hiltner 652 obtained with SALT in 2022 (post 2020) calculated by convolving the spectropolarimetry with the Johnson-Cousins filter pass-bands. This is compared to the published FORS1 [11] and FORS2 [12] observation results (top and middle panels). The average of the FORS reduced Stokes parameters as well as their standard deviation are plotted as the horizontal lines and shaded regions, respectively. Hiltner 652 was observed by SALT in 2022, by FORS1 from 1999 to 2006 (shown as average value at March 2005), and by FORS2 from 2010 to 2016. Finally, the Johnson-Cousins filter pass-band's used by POLSALT to calculate the stokes parameters in the different filters are shown in the bottom panel.

4. Results

The linear Stokes parameters, Q & U , normalized to the total intensity, referred to henceforth as the reduced Stokes parameters, P_Q & P_U , are relative quantities which are used to calculate the polarisation properties, and from which the fraction of linear polarisation may be easily calculated as $P_L = \sqrt{P_Q^2 + P_U^2}$ [13]. The reduced Stokes parameters of Hiltner 652, convolved against the Johnson-Cousins filter pass-bands, as performed by POLSALT, are shown in Figure 2 and are compared against published filtered reduced Stokes parameters obtained with FORS1 [11] and FORS2 [12]. The bottom panel of Figure 2 also shows the filter band-passes for the wavelength range covered by the SALT spectropolarimetry; note that the full R band is not covered by the observation setup.

The reduced Stokes parameters of Hiltner 652 fall within or near the 1σ level of the previously published reduced Stokes parameters over the course of multiple years. This shows that Hiltner 652 is not only a good standard for spectropolarimetric observations but also that using POLSALT with the supplementary wavelength calibrations correctly calculated the Stokes parameters.

The blazar 3C 279 was observed during periods of enhanced and flaring γ -ray activity for energies of 0.1 – 100 GeV and were sourced from the Fermi Light Curve Repository (Fermi LCR)

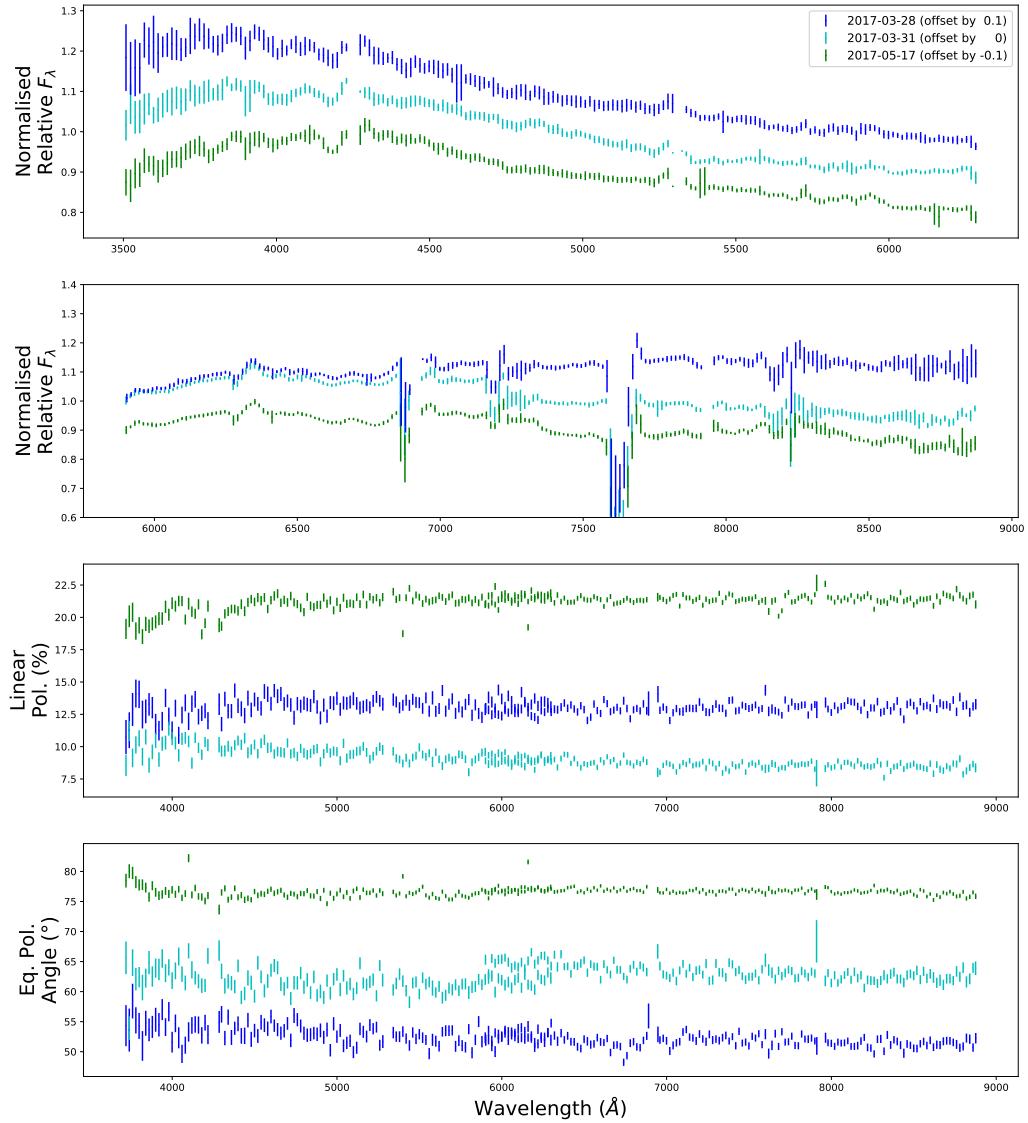


Figure 3: Optical spectropolarimetric observations of 3C 279 during two different periods of flaring. The relative flux calibrated spectra for the grating angles 12.5° (top), observed first at each epoch, and 19.625° (second from top), the degree of polarisation (third from top), and the polarisation angle (bottom).

[14]. 3C 279 was observed at fluxes of 1.58×10^{-6} , 2.87×10^{-6} , and 2.00×10^{-6} ph cm $^{-2}$ s $^{-1}$ as compared to a ‘quiescent’ median flux of 4.59×10^{-7} ph cm $^{-2}$ s $^{-1}$.

Figure 3 shows the relative flux calibrated spectra normalized to the spectrum’s mean (and offset as indicated), degree of polarisation, and polarisation angle for 3C 279. The relative flux was normalized to better compare how the shape of the spectrum changes during the differing γ -ray states. 3C 279 showed variable flux and polarisation during the three observations, with the polarisation changing by $\sim 12.23\%$ and the angle changing by $\sim 24.22^\circ$. The relative spectral shape remains similar over all three observations, however the second observation (2017-03-31) appears to show a slightly brighter spectrum in the blue. There is also a relatively good agreement

in the polarisation at the wavelength overlap for the different grating angles. The difference of the polarisation and degree of polarisation is on the order of 0.25 % and 1.25°, respectively.

5. Discussion and Conclusions

In this paper we have considered the observations of a spectropolarimetric standard, Hiltner 652, and a flaring blazar, 3C 279, to test our supplementary tools used for the reduction of spectropolarimetric data obtained with the RSS. The results serve to confirm that the supplementary tools' wavelength reductions, in combination with POLSALT, produces the correct polarisation behavior. Our results for Hiltner 652 agree well with the historical results obtained using the FOcal Reducer and low dispersion Spectrograph (FORS1 and FORS2) mounted at the Very Large Telescope (VLT).

For 3C 279, there is similarly good agreement between the polarisation for the observations taken at two grating angles where the wavelength ranges overlap. Some disagreement is found which may be due to the inherent nature of flaring blazars which cause the polarisation angle to drastically vary and the two observations at different grating angles are not taken simultaneously. Additionally, the disagreement may be due to a lower Signal-to-Noise Ratio (SNR) at the edges of the RSS, which would introduce a larger systematic error. The lower SNR could originate from either blaze effects, or from the extraction of the trace, which is currently limited to a rectangular aperture in the current POLSALT pipeline.

The results show that our supplementary wavelength calibrations for POLSALT increase the efficiency of the data reduction and produce acceptable results. This reduction procedure has been used for the reduction of a number of blazar observations undertaken as part of a long term monitoring campaign [15].

Acknowledgments

All observations taken in this study were obtained using the RSS mounted on SALT under the programme ID 2016-2-LSP-001 (PI DAH Buckley).

References

- [1] C.M. Urry and P. Padovani, *Unified Schemes for Radio-Loud Active Galactic Nuclei*, *Publications of the Astronomical Society of the Pacific* **107** (1995) 803 [[astro-ph/9506063](#)].
- [2] G. Ghisellini, F. Tavecchio, L. Foschini, G. Ghirlanda, L. Maraschi and A. Celotti, *General physical properties of bright Fermi blazars*, *Monthly Notices of the Royal Astronomical Society* **402** (2010) 497 [[0909.0932](#)].
- [3] M. Böttcher, A. Reimer, K. Sweeney and A. Prakash, *Leptonic and hadronic modeling of fermi-detected blazars*, *The Astrophysical Journal* **768** (2013) 54.

- [4] H.M. Schutte, M. Böttcher, J. Barnard, A. Dmytriiev, B. van Soelen, A. Falcone et al., *Modeling the Multi-wavelength Polarization and Spectral Energy Distributions of Blazars*, in *44th COSPAR Scientific Assembly. Held 16-24 July*, vol. 44, p. 1869, July, 2022.
- [5] H.M. Schutte, R.J. Britto, M. Böttcher, B. van Soelen, J.P. Marais, A. Kaur et al., *Modeling the spectral energy distributions and spectropolarimetry of blazars—application to 4C+01.02 in 2016-2017**, *The Astrophysical Journal* **925** (2022) 139.
- [6] E.B. Burgh, K.H. Nordsieck, H.A. Kobulnicky, T.B. Williams, D. O'Donoghue, M.P. Smith et al., *Prime Focus Imaging Spectrograph for the Southern African Large Telescope: optical design*, in *Instrument Design and Performance for Optical/Infrared Ground-based Telescopes*, M. Iye and A.F.M. Moorwood, eds., vol. 4841 of *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, pp. 1463–1471, Mar., 2003, DOI.
- [7] H.A. Kobulnicky, K.H. Nordsieck, E.B. Burgh, M.P. Smith, J.W. Percival, T.B. Williams et al., *Prime focus imaging spectrograph for the Southern African large telescope: operational modes*, in *Instrument Design and Performance for Optical/Infrared Ground-based Telescopes*, M. Iye and A.F.M. Moorwood, eds., vol. 4841 of *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, pp. 1634–1644, Mar., 2003, DOI.
- [8] S. Crawford and K. Nordsieck, “Polsalt.” <https://github.com/saltastro/polsalt>, 2020.
- [9] J. Cooper, B. van Soelen and R. Britto, *Development of tools for SALT/RSS spectropolarimetry reductions: application to the blazar 3C279*, in *High Energy Astrophysics in Southern Africa 2021*, p. 56, May, 2022, DOI.
- [10] D. Tody, *The IRAF Data Reduction and Analysis System*, in *Instrumentation in astronomy VI*, D.L. Crawford, ed., vol. 627 of *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, p. 733, Jan., 1986, DOI.
- [11] L. Fossati, S. Bagnulo, E. Mason and E. Landi Degl'Innocenti, *Standard Stars for Linear Polarization Observed with FORS1*, in *The Future of Photometric, Spectrophotometric and Polarimetric Standardization*, C. Sterken, ed., vol. 364 of *Astronomical Society of the Pacific Conference Series*, p. 503, Apr., 2007.
- [12] A. Cikota, F. Patat, S. Cikota and T. Faran, *Linear spectropolarimetry of polarimetric standard stars with VLT/FORS2*, *Monthly Notices of the Royal Astronomical Society* **464** (2017) 4146 [1610.00722].
- [13] E. Landi Degl'Innocenti, S. Bagnulo and L. Fossati, *Polarimetric Standardization*, in *The Future of Photometric, Spectrophotometric and Polarimetric Standardization*, C. Sterken, ed., vol. 364 of *Astronomical Society of the Pacific Conference Series*, p. 495, Apr., 2007, DOI [astro-ph/0610262].

- [14] S. Abdollahi, M. Ajello, L. Baldini, J. Ballet, D. Bastieri, J.B. Gonzalez et al., *The fermi-lat light curve repository*, 2023. 10.48550/ARXIV.2301.01607.
- [15] J. Barnard, B. van Soelen, J. Cooper, R. Britto, J.P. Marais, I. van der Westhuizen et al., *Optical Spectropolarimetry Observations of the BL Lac-type object PKS 0537-441 after a period of quiescence*, in *High Energy Astrophysics in Southern Africa 2021*, p. 9, May, 2022, DOI.

Bibliography

S. Abdollahi, M. Ajello, L. Baldini, J. Ballet, D. Bastieri, J. Becerra Gonzalez, R. Bellazzini, A. Berretta, E. Bissaldi, R. Bonino, A. Brill, P. Bruel, E. Burns, S. Buson, R. A. Cameron, R. Caputo, P. A. Caraveo, N. Cibrario, S. Ciprini, P. Cristarella Orestano, M. Crnogorcevic, S. Cutini, F. D'Ammando, S. De Gaetano, S. W. Digel, N. Di Lalla, L. Di Venere, A. Domínguez, V. Fallah Ramazani, S. J. Fegan, E. C. Ferrara, A. Fiori, H. Fleischhack, A. Franckowiak, Y. Fukazawa, P. Fusco, V. Gammaldi, F. Gargano, S. Garrappa, C. Gasbarra, D. Gasparrini, N. Giglietto, F. Giordano, M. Giroletti, D. Green, I. A. Grenier, S. Guiriec, M. Gustafsson, E. Hays, D. Horan, X. Hou, G. Jóhannesson, M. Kerr, D. Kocevski, M. Kuss, L. Latronico, J. Li, I. Liodakis, F. Longo, F. Loparco, L. Lorusso, B. Lott, M. N. Lovellette, P. Lubrano, S. Maldera, A. Manfreda, G. Martí-Devesa, M. N. Mazziotta, I. Mereu, M. Meyer, P. F. Michelson, T. Mizuno, M. E. Monzani, A. Morselli, I. V. Moskalenko, M. Negro, N. Omodei, E. Orlando, J. F. Ormes, D. Panque, G. Panzarini, J. S. Perkins, M. Persic, M. Pesce-Rollins, R. Pillera, T. A. Porter, G. Principe, J. L. Racusin, S. Rainò, R. Rando, B. Rani, M. Razzano, S. Razzaque, A. Reimer, O. Reimer, M. Sánchez-Conde, P. M. Saz Parkinson, Jeff Scargle, L. Scotton, D. Serini, C. Sgrò, E. J. Siskind, G. Spandre, P. Spinelli, D. J. Suson, H. Tajima, D. J. Thompson, D. F. Torres, J. Valverde, T. Ventters, Z. Wadiasingh, S. Wagner, and K. Wood. The fermi-lat lightcurve repository*. *The Astrophysical Journal Supplement Series*, 265(2):31, March 2023. ISSN 1538-4365. doi: 10.3847/1538-4365/acbb6a. URL <http://dx.doi.org/10.3847/1538-4365/acbb6a>.

R. R. J. Antonucci and J. S. Miller. Spectropolarimetry and the nature of NGC 1068. *ApJ*, 297:621–632, October 1985. doi: 10.1086/163559.

George B. Arfken and Hans J. Weber. Mathematical methods for physicists, 1999.

Astropy Collaboration, T. P. Robitaille, E. J. Tollerud, P. Greenfield, M. Droettboom, E. Bray, T. Aldcroft, M. Davis, A. Ginsburg, A. M. Price-Whelan, W. E. Kerzendorf, A. Conley, N. Crighton, K. Barbary, D. Muna, H. Ferguson, F. Grollier, M. M. Parikh, P. H. Nair, H. M. Unther, C. Deil, J. Woillez, S. Conseil, R. Kramer, J. E. H. Turner, L. Singer, R. Fox, B. A. Weaver, V. Zabalza, Z. I. Edwards, K. Azalee Bostroem, D. J. Burke, A. R. Casey, S. M. Crawford, N. Dencheva, J. Ely, T. Jenness, K. Labrie, P. L. Lim, F. Pierfederici, A. Pontzen, A. Ptak, B. Refsdal, M. Servillat, and O. Streicher. Astropy: A community Python package for astronomy. *A&A*, 558:A33, October 2013. doi: 10.1051/0004-6361/201322068.

Astropy Collaboration, A. M. Price-Whelan, B. M. Sipőcz, H. M. Günther, P. L. Lim,

S. M. Crawford, S. Conseil, D. L. Shupe, M. W. Craig, N. Dencheva, A. Ginsburg, J. T. VanderPlas, L. D. Bradley, D. Pérez-Suárez, M. de Val-Borro, T. L. Aldcroft, K. L. Cruz, T. P. Robitaille, E. J. Tollerud, C. Ardelean, T. Babej, Y. P. Bach, M. Bachetti, A. V. Bakanov, S. P. Bamford, G. Barentsen, P. Barmby, A. Baumbach, K. L. Berry, F. Biscani, M. Boquien, K. A. Bostroem, L. G. Bouma, G. B. Brammer, E. M. Bray, H. Breytenbach, H. Buddelmeijer, D. J. Burke, G. Calderone, J. L. Cano Rodríguez, M. Cara, J. V. M. Cardoso, S. Cheedella, Y. Copin, L. Corrales, D. Crichton, D. D'Avella, C. Deil, É. Depagne, J. P. Dietrich, A. Donath, M. Droettboom, N. Earl, T. Erben, S. Fabbro, L. A. Ferreira, T. Finethy, R. T. Fox, L. H. Garrison, S. L. J. Gibbons, D. A. Goldstein, R. Gommers, J. P. Greco, P. Greenfield, A. M. Groener, F. Grollier, A. Hagen, P. Hirst, D. Homeier, A. J. Horton, G. Hosseinzadeh, L. Hu, J. S. Hunkeler, Ž. Ivezić, A. Jain, T. Jenness, G. Kanarek, S. Kendrew, N. S. Kern, W. E. Kerzendorf, A. Khvalko, J. King, D. Kirkby, A. M. Kulkarni, A. Kumar, A. Lee, D. Lenz, S. P. Littlefair, Z. Ma, D. M. Macleod, M. Mastropietro, C. McCully, S. Montagnac, B. M. Morris, M. Mueller, S. J. Mumford, D. Muna, N. A. Murphy, S. Nelson, G. H. Nguyen, J. P. Ninan, M. Nöthe, S. Ogaz, S. Oh, J. K. Parejko, N. Parley, S. Pasqual, R. Patil, A. A. Patil, A. L. Plunkett, J. X. Prochaska, T. Rastogi, V. Reddy Janga, J. Sabater, P. Sakurikar, M. Seifert, L. E. Sherbert, H. Sherwood-Taylor, A. Y. Shih, J. Sick, M. T. Silbiger, S. Singanamalla, L. P. Singer, P. H. Sladen, K. A. Sooley, S. Sornarajah, O. Streicher, P. Teuben, S. W. Thomas, G. R. Tremblay, J. E. H. Turner, V. Terrón, M. H. van Kerkwijk, A. de la Vega, L. L. Watkins, B. A. Weaver, J. B. Whitmore, J. Woillez, V. Zabalza, and Astropy Contributors. The Astropy Project: Building an Open-science Project and Status of the v2.0 Core Package. *AJ*, 156(3):123, September 2018. doi: 10.3847/1538-3881/aabc4f.

Astropy Collaboration, Adrian M. Price-Whelan, Pey Lian Lim, Nicholas Earl, Nathaniel Starkman, Larry Bradley, David L. Shupe, Aarya A. Patil, Lia Corrales, C. E. Brasseur, Maximilian N"othe, Axel Donath, Erik Tollerud, Brett M. Morris, Adam Ginsburg, Eero Vaher, Benjamin A. Weaver, James Tocknell, William Jamieson, Marten H. van Kerkwijk, Thomas P. Robitaille, Bruce Merry, Matteo Bachetti, H. Moritz G"unther, Thomas L. Aldcroft, Jaime A. Alvarado-Montes, Anne M. Archibald, Attila B'odi, Shreyas Bapat, Geert Barentsen, Juanjo Baz'an, Manish Biswas, M'ed'eric Boquien, D. J. Burke, Daria Cara, Mihai Cara, Kyle E. Conroy, Simon Conseil, Matthew W. Craig, Robert M. Cross, Kelle L. Cruz, Francesco D'Eugenio, Nadia Dencheva, Hadrien A. R. Devillepoix, J"org P. Dietrich, Arthur Davis Eigenbrot, Thomas Erben, Leonardo Ferreira, Daniel Foreman-Mackey, Ryan Fox, Nabil Freij, Suyog Garg, Robel Geda, Lauren Glattly, Yash Gondhalekar, Karl D. Gordon, David Grant, Perry Greenfield, Austen M. Groener, Steve Guest, Sebastian Gurovich, Rasmus Handberg, Akeem Hart, Zac Hatfield-Dodds, Derek Homeier, Griffin Hosseinzadeh, Tim Jenness, Craig K. Jones, Prajwel Joseph, J. Bryce Kalmbach, Emir Karamehmetoglu, Mikolaj Kaluszy'nski, Michael S. P. Kelley, Nicholas Kern, Wolfgang E. Kerzendorf, Eric W. Koch, Shankar Kulumani, Antony Lee, Chun Ly, Zhiyuan Ma, Conor MacBride, Jakob M. Maljaars, Demitri Muna, N. A. Murphy, Henrik Norman, Richard O'Steen, Kyle A. Oman, Camilla Pacifici, Sergio Pascual, J. Pascual-Granado, Rohit R. Patil, Gabriel I. Perren, Timothy E. Pickering, Tanuj Rastogi, Benjamin R. Roulston, Daniel F. Ryan, Eli S. Rykoff, Jose Sabater, Parikshit Sakurikar, Jes'us Salgado, Aniket Sanghi, Nicholas Saunders, Volodymyr Savchenko, Ludwig Schwardt, Michael Seifert-Eckert, Albert Y. Shih, Anany Shrey Jain, Gyanendra Shukla, Jonathan Sick, Chris Simpson, Sudheesh

- Singanamalla, Leo P. Singer, Jaladh Singhal, Manodeep Sinha, Brigitta M. SipHocz, Lee R. Spitler, David Stansby, Ole Streicher, Jani Sumak, John D. Swinbank, Dan S. Taranu, Nikita Tewary, Grant R. Tremblay, Miguel de Val-Borro, Samuel J. Van Kooten, Zlatan Vasović, Shresth Verma, Jos'e Vin'icius de Miranda Cardoso, Peter K. G. Williams, Tom J. Wilson, Benjamin Winkel, W. M. Wood-Vasey, Rui Xue, Peter Yoachim, Chen Zhang, Andrea Zonca, and Astropy Project Contributors. The Astropy Project: Sustaining and Growing a Community-oriented Open-source Project and the Latest Major Release (v5.0) of the Core Package. *ApJ*, 935(2):167, August 2022. doi: 10.3847/1538-4357/ac7c74.
- S. Bagnulo, M. Landolfi, J. D. Landstreet, E. Landi Degl'Innocenti, L. Fossati, and M. Sterzik. Stellar spectropolarimetry with retarder waveplate and beam splitter devices. *Publications of the Astronomical Society of the Pacific*, 121(883):993, aug 2009. doi: 10.1086/605654. URL <https://dx.doi.org/10.1086/605654>.
- Erasmus Bartholinus. Experimenta crystalli islandici dis-diaclastici, quibus mira et insolita refractio detegitur (copenhagen, 1670). *Edinburgh Philosophical Journal*, 1:271, 1670.
- D. Scott Birney, Guillermo Gonzalez, and David Oesper. *Observational Astronomy - 2nd Edition*. Cambridge University Press, 2006. doi: 10.2277/0521853702.
- Janus D. Brink, Moses K. Mogotsi, Melanie Saayman, Nicolaas M. Van der Merwe, Jonathan Love, and Alrin Christians. Preparing the SALT for near-infrared observations. In Heather K. Marshall, Jason Spyromilio, and Tomonori Usuda, editors, *Ground-based and Airborne Telescopes IX*, volume 12182 of *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, page 121822E, August 2022. doi: 10.1117/12.2627328.
- D. A. H. Buckley, J. Brink, N. S. Loaring, A. Swat, and H. L. Worters. The Southern African Large Telescope (SALT) calibration system. In Ian S. McLean and Mark M. Casali, editors, *Ground-based and Airborne Instrumentation for Astronomy II*, volume 7014, page 70146H. International Society for Optics and Photonics, SPIE, 2008. doi: 10.1117/12.790385. URL <https://doi.org/10.1117/12.790385>.
- D. A. H. Buckley, S. Bagnulo, R. J. Britto, J. Mao, D. A. Kann, J. Cooper, V. Lipunov, D. M. Hewitt, S. Razzaque, N. P. M. Kuin, I. M. Monageng, S. Covino, P. Jakobsson, A. J. van der Horst, K. Wiersema, M. Böttcher, S. Campana, V. D'Elia, E. S. Gorbovskoy, I. Gorbunov, D. N. Groenewald, D. H. Hartmann, V. G. Kornilov, C. G. Mundell, R. Podesta, J. K. Thomas, N. Tyurina, D. Vlasenko, B. van Soelen, and D. Xu. Spectropolarimetry and photometry of the early afterglow of the gamma-ray burst GRB 191221B. *MNRAS*, 506(3):4621–4631, September 2021. doi: 10.1093/mnras/stab1791.
- David A. H. Buckley, Gerhard P. Swart, and Jacobus G. Meiring. Completion and commissioning of the Southern African Large Telescope. In Larry M. Stepp, editor, *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, volume 6267 of

- Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series, page 62670Z, June 2006. doi: 10.1117/12.673750.
- Christian Buil. CCD astronomy : construction and use of an astronomical CCD camera / Christian Buil ; translated and adapted from the French by Emmanuel and Barbara Davoust. Willmann-Bell, Richmond, Va, 1st english ed. edition, 1991. ISBN 0943396298.
- Eric B. Burgh, Kenneth H. Nordsieck, Henry A. Kobulnicky, Ted B. Williams, Daragh O'Donoghue, Michael P. Smith, and Jeffrey W. Percival. Prime Focus Imaging Spectrograph for the Southern African Large Telescope: optical design. In Masanori Iye and Alan F. M. Moorwood, editors, Instrument Design and Performance for Optical/Infrared Ground-based Telescopes, volume 4841 of Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series, pages 1463–1471, March 2003. doi: 10.1117/12.460312.
- Subrahmanyan Chandrasekhar. Radiative transfer, 1950.
- Aleksandar Cikota, Ferdinando Patat, Stefan Cikota, and Tamar Faran. Linear spectropolarimetry of polarimetric standard stars with VLT/FORS2. MNRAS, 464(4):4146–4159, February 2017. doi: 10.1093/mnras/stw2545.
- Marshall H. Cohen. Genesis of the 1000-foot Arecibo dish. Journal of Astronomical History and Heritage, 12(2):141–152, July 2009.
- E. Collett. Field Guide to Polarization. Field Guides. SPIE Press, 2005. ISBN 9780819458681. URL <https://books.google.co.za/books?id=51JwcCsLbLsC>.
- J. Cooper and B. van Soelen. SALT Spectropolarimetric Pipeline Comparisons. In High Energy Astrophysics in Southern Africa 2022, page 56, December 2023.
- J. Cooper, B. van Soelen, and R. Britto. Development of tools for SALT/RSS spectropolarimetry reductions: application to the blazar 3C279. In High Energy Astrophysics in Southern Africa 2021, page 56, May 2022. doi: 10.22323/1.401.0056.
- M. Craig, S. Crawford, M. Seifert, T. Robitaille, B. Sipőcz, J. Walawender, Z. Vinícius, J. P. Ninan, M. Droettboom, J. Youn, E. Tollerud, E. Bray, N. Walker, Janga V. R., C. Stotts, H. M. Günther, E. Rol, Yoonsoo P. Bach, L. Bradley, C. Deil, A. Price-Whelan, K. Barbary, A. Horton, W. Schoenell, N. Heidt, F. Gasdia, S. Nelson, and O. Streicher. astropy/ccdproc: v1.3.0.post1, December 2017. URL <https://doi.org/10.5281/zenodo.1069648>.
- G. Dahlquist and Å. Björck. Numerical Methods. Dover Books on Mathematics. Dover Publications, 2003. ISBN 9780486428079. URL <https://books.google.co.ls/books?id=armfeHpJIwAC>.
- E. Landi Degl'Innocenti, S. Bagnulo, and L. Fossati. Polarimetric standardization, 2006.

Egidio Landi Degl'Innocenti. The physics of polarization. *Proceedings of the International Astronomical Union*, 10(S305):1–1, 2014.

Egidio Landi Degl'Innocenti and M. Landolfi. *Polarization in Spectral Lines*, volume 307. Springer Dordrecht, 2004. doi: 10.1007/978-1-4020-2415-3.

Königlich Bayerische Akademie der Wissenschaften. *Denkschriften der Königlichen Akademie der Wissenschaften zu München für das Jahre 1820 und 1821*, volume 8. Die Akademie, 1824. URL <https://books.google.co.za/books?id=k-EAAAAYAAJ>.

J. F. Donati, M. Semel, B. D. Carter, D. E. Rees, and A. Collier Cameron. Spectropolarimetric observations of active stars. *MNRAS*, 291(4):658–682, November 1997. doi: 10.1093/mnras/291.4.658.

I. V. Florinsky and A. N. Pankratov. Digital terrain modeling with the chebyshev polynomials. *Machine Learning and Data Analysis*, 1(12):1647 – 1659, 2015. doi: 10.48550/ARXIV.1507.03960. URL <https://arxiv.org/abs/1507.03960>.

L. Fossati, S. Bagnulo, E. Mason, and E. Landi Degl'Innocenti. Standard Stars for Linear Polarization Observed with FORS1. In C. Sterken, editor, *The Future of Photometric, Spectrophotometric and Polarimetric Standardization*, volume 364 of *Astronomical Society of the Pacific Conference Series*, page 503, April 2007.

Augustin Fresnel. *Oeuvres completes d'Augustin Fresnel: 3*. Imprimerie impériale, 1870.

L. M. Freyhammer, M. I. Andersen, T. Arentoft, C. Sterken, and P. Nørregaard. On Cross-talk Correction of Images from Multiple-port CCDs. *Experimental Astronomy*, 12(3):147–162, January 2001. doi: 10.1023/A:1021820418263.

David J Griffiths. *Introduction to electrodynamics*, 2005.

George E. Hale. The Zeeman Effect in the Sun. *PASP*, 20(123):287, December 1908. doi: 10.1086/121847.

George E. Hale. *16. On the Probable Existence of a Magnetic Field in Sun-Spots*, pages 96–105. Harvard University Press, Cambridge, MA and London, England, 1979. ISBN 9780674366688. doi: doi:10.4159/harvard.9780674366688.c19. URL <https://doi.org/10.4159/harvard.9780674366688.c19>.

P. D. Hale and G. W. Day. Stability of birefringent linear retarders(waveplates). *Appl. Opt.*, 27(24):5146–5153, Dec 1988. doi: 10.1364/AO.27.005146. URL <https://opg.optica.org/ao/abstract.cfm?URI=ao-27-24-5146>.

C. R. Harris, K. J. Millman, S. J. van der Walt, R. Gommers, P. Virtanen, D. Cournapeau, E. Wieser, J. Taylor, S. Berg, N. J. Smith, R. Kern, M. Picus, S. Hoyer, M. H. van Kerkwijk, M. Brett, A. Haldane, J. F. del Río, M. Wiebe, P. Peterson, P. Gérard-Marchant, K. Sheppard, T. Reddy, W. Weckesser, H. Abbasi, C. Gohlke, and T. E. Oliphant.

- Array programming with NumPy. *Nature*, 585(7825):357–362, September 2020. doi: 10.1038/s41586-020-2649-2. URL <https://doi.org/10.1038/s41586-020-2649-2>.
- E. Hecht. *Optics*. Pearson Education, Incorporated, 2017. ISBN 9780133977226. URL <https://books.google.co.za/books?id=ZarLoQEACAAJ>.
- Steve B. Howell. *Handbook of CCD Astronomy*, volume 5. Cambridge University Press, 2006.
- J. D. Hunter. Matplotlib: A 2d graphics environment. *Computing in Science & Engineering*, 9(3):90–95, 2007. doi: 10.1109/MCSE.2007.55.
- Christian Huygens. Treatise on light, 1690. translated by Thompson, s. p., 1690. URL <https://www.gutenberg.org/files/14725/14725-h/14725-h.htm>.
- Mourad E. H. Ismail. *Classical and Quantum Orthogonal Polynomials in One Variable*. Encyclopedia of Mathematics and its Applications. Cambridge University Press, 2005. doi: 10.1017/CBO9781107325982.
- James Janesick, James T. Andrews, and Tom Elliott. Fundamental performance differences between CMOS and CCD imagers: Part 1. In David A. Dorn and Andrew D. Holland, editors, *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, volume 6276 of *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, page 62760M, June 2006. doi: 10.1117/12.678867.
- F.A. Jenkins and H.E. White. *Fundamentals of Optics*. International student edition. McGraw-Hill, 1976. ISBN 9780070323308. URL <https://books.google.co.za/books?id=dCdRAAAAMAAJ>.
- Christoph U. Keller. Instrumentation for astrophysical spectropolarimetry. *Astrophysical Spectropolarimetry*, 1:303–354, 2002.
- G. Kirchhoff and R. Bunsen. Chemische Analyse durch Spectralbeobachtungen. *Annalen der Physik*, 189(7):337–381, January 1861. doi: 10.1002/andp.18611890702.
- Henry A. Kobulnicky, Kenneth H. Nordsieck, Eric B. Burgh, Michael P. Smith, Jeffrey W. Percival, Ted B. Williams, and Darragh O’Donoghue. Prime focus imaging spectrograph for the Southern African large telescope: operational modes. In Masanori Iye and Alan F. M. Moorwood, editors, *Instrument Design and Performance for Optical/Infrared Ground-based Telescopes*, volume 4841 of *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, pages 1634–1644, March 2003. doi: 10.1117/12.460315.
- Gerard Leng. Compression of aircraft aerodynamic database using multivariable chebyshev polynomials. *Advances in Engineering Software*, 28(2):133–141, 1997. ISSN 0965-9978. doi: [https://doi.org/10.1016/S0965-9978\(96\)00043-9](https://doi.org/10.1016/S0965-9978(96)00043-9). URL <https://www.sciencedirect.com/science/article/pii/S0965997896000439>.

- Dave Litwiller. Ccd vs. cmos. *Photonics spectra*, 35(1):154–158, 2001.
- Dongyue Liu and Bryan M. Hennelly. Improved wavelength calibration by modeling the spectrometer. *Applied Spectroscopy*, 76(11):1283–1299, 2022. doi: 10.1177/00037028221111796. URL <https://doi.org/10.1177/00037028221111796>. PMID: 35726593.
- Etienne L. Malus. Sur une propriété de la lumière réfléchie. *Mém. Phys. Chim. Soc. d'Arcueil*, 2:143–158, 1809.
- Curtis McCully, Steve Crawford, Gabor Kovacs, Erik Tollerud, Edward Betts, Larry Bradley, Matt Craig, James Turner, Ole Streicher, Brigitta Sipocz, Thomas Robitaille, and Christoph Deil. astropy/astroscrappy: v1.0.5 zenodo release, November 2018. URL <https://doi.org/10.5281/zenodo.1482019>.
- I. Newton and W. Innys. *Opticks:: Or, A Treatise of the Reflections, Refractions, Inflections and Colours of Light*. Opticks:: Or, A Treatise of the Reflections, Refractions, Inflections and Colours of Light. William Innys at the West-End of St. Paul's., 1730. URL <https://books.google.co.za/books?id=GnAFAAAQAAJ>.
- Kenneth H. Nordsieck, Kurt P. Jaehnig, Eric B. Burgh, Henry A. Kobulnicky, Jeffrey W. Percival, and Michael P. Smith. Instrumentation for high-resolution spectropolarimetry in the visible and far-ultraviolet. In Silvano Fineschi, editor, *Polarimetry in Astronomy*, volume 4843 of *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, pages 170–179, February 2003. doi: 10.1117/12.459288.
- D. O'Donoghue, D. A. H. Buckley, L. A. Balona, D. Bester, L. Botha, J. Brink, D. B. Carter, P. A. Charles, A. Christians, F. Ebrahim, R. Emmerich, W. Esterhuyse, G. P. Evans, C. Fourie, P. Fourie, H. Gajjar, M. Gordon, C. Gumede, M. de Kock, A. Koeslag, W. P. Koorts, H. Kriel, F. Marang, J. G. Meiring, J. W. Menzies, P. Menzies, D. Metcalfe, B. Meyer, L. Nel, J. O'Connor, F. Osman, C. Du Plessis, H. Rall, A. Riddick, E. Romero-Colmenero, S. B. Potter, C. Sass, H. Schalekamp, N. Sessions, S. Siyengo, V. Sopela, H. Steyn, J. Stoffels, J. Scholtz, G. Swart, A. Swat, J. Swiegers, T. Tiheli, P. Vaisanen, W. Whittaker, and F. van Wyk. First science with the Southern African Large Telescope: peering at the accreting polar caps of the eclipsing polar SDSS J015543.40+002807.2. *MNRAS*, 372(1):151–162, October 2006. doi: 10.1111/j.1365-2966.2006.10834.x.
- Darragh O'Donoghue. Correction of spherical aberration in the Southern African Large Telescope (SALT). In Philippe Dierickx, editor, *Optical Design, Materials, Fabrication, and Maintenance*, volume 4003 of *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, pages 363–372, July 2000. doi: 10.1117/12.391526.
- Darragh O'Donoghue. Atmospheric dispersion corrector for the Southern African Large Telescope (SALT). In Richard G. Bingham and David D. Walker, editors, *Large Lenses and Prisms*, volume 4411 of *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, pages 79–84, February 2002. doi: 10.1117/12.454874.

Ferdinando Patat and Martino Romaniello. Error Analysis for Dual-Beam Optical Linear Polarimetry. *PASP*, 118(839):146–161, January 2006. doi: 10.1086/497581.

Alba Peinado, Angel Lizana, Josep Vidal, Claudio Iemmi, and Juan Campos. Optimization and performance criteria of a stokes polarimeter based on two variable retarders. *Opt. Express*, 18(10):9815–9830, May 2010. doi: 10.1364/OE.18.009815. URL <https://opg.optica.org/oe/abstract.cfm?URI=oe-18-10-9815>.

W. H. Press, S. A. Teukolsky, W. T. Vetterling, and B. P. Flannery. *Numerical Recipes* 3rd Edition: The Art of Scientific Computing. Cambridge University Press, 2007. ISBN 9780521880688. URL <https://books.google.co.za/books?id=1aA0dzK3FegC>.

J. R. Priebe. Operational form of the mueller matrices. *J. Opt. Soc. Am.*, 59(2):176–180, Feb 1969. doi: 10.1364/JOSA.59.000176. URL <https://opg.optica.org/abstract.cfm?URI=josa-59-2-176>.

Lawrence W. Ramsey, M. T. Adams, Thomas G. Barnes, John A. Booth, Mark E. Cornell, James R. Fowler, Niall I. Gaffney, John W. Glaspey, John M. Good, Gary J. Hill, Philip W. Kelton, Victor L. Krabbendam, L. Long, Phillip J. MacQueen, Frank B. Ray, Randall L. Ricklefs, J. Sage, Thomas A. Sebring, W. J. Spiesman, and M. Steiner. Early performance and present status of the Hobby-Eberly Telescope. In Larry M. Stepp, editor, *Advanced Technology Optical/IR Telescopes VI*, volume 3352 of *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, pages 34–42, August 1998. doi: 10.1117/12.319287.

Hester M. Schutte, Richard J. Britto, Markus Böttcher, Brian van Soelen, Johannes P. Marais, Amanpreet Kaur, Abraham D. Falcone, David A. H. Buckley, Andry F. Rajole-limanana, and Justin Cooper. Modeling the Spectral Energy Distributions and Spectropolarimetry of Blazars - Application to 4C+01.02 in 2016-2017. *ApJ*, 925(2):139, February 2022. doi: 10.3847/1538-4357/ac3cb5.

Maria C. Simon. Wollaston prism with large split angle. *Appl. Opt.*, 25(3):369–376, Feb 1986. doi: 10.1364/AO.25.000369. URL <https://opg.optica.org/ao/abstract.cfm?URI=ao-25-3-369>.

G. G. Stokes. On the Composition and Resolution of Streams of Polarized Light from different Sources. *Transactions of the Cambridge Philosophical Society*, 9:399, January 1852.

Doug Tody. The IRAF Data Reduction and Analysis System. In David L. Crawford, editor, *Instrumentation in astronomy VI*, volume 627 of *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, page 733, January 1986. doi: 10.1117/12.968154.

Doug Tody. IRAF in the Nineties. In R. J. Hanisch, R. J. V. Brissenden, and J. Barnes, editors, *Astronomical Data Analysis Software and Systems II*, volume 52 of *Astronomical Society of the Pacific Conference Series*, page 173, January 1993.

Stephen F. Tonkin. Practical Amateur Spectroscopy. The Patrick Moore Practical Astronomy Series. Springer London, 2013. ISBN 9781447101277. URL <https://books.google.fr/books?id=b2fgBwAAQBAJ>.

Pieter G. van Dokkum. Cosmic-Ray Rejection by Laplacian Edge Detection. PASP, 113(789):1420–1427, November 2001. doi: 10.1086/323894.

Pauli Virtanen, Ralf Gommers, Travis E. Oliphant, Matt Haberland, Tyler Reddy, David Cournapeau, Evgeni Burovski, Pearu Peterson, Warren Weckesser, Jonathan Bright, Stéfan J. van der Walt, Matthew Brett, Joshua Wilson, K. Jarrod Millman, Nikolay Mayorov, Andrew R. J. Nelson, Eric Jones, Robert Kern, Eric Larson, C J Carey, İlhan Polat, Yu Feng, Eric W. Moore, Jake VanderPlas, Denis Laxalde, Josef Perktold, Robert Cimrman, Ian Henriksen, E. A. Quintero, Charles R. Harris, Anne M. Archibald, Antônio H. Ribeiro, Fabian Pedregosa, Paul van Mulbregt, and SciPy 1.0 Contributors. SciPy 1.0: Fundamental Algorithms for Scientific Computing in Python. Nature Methods, 17:261–272, 2020. doi: 10.1038/s41592-019-0686-2.

L. Wang and J. C. Wheeler. Spectropolarimetry of supernovae. ARA&A, 46:433–474, September 2008. doi: 10.1146/annurev.astro.46.060407.145139.

Marsha J. Wolf, Matthew A. Bershadsky, Michael P. Smith, Kurt P. Jaehnig, Jeffrey W. Percival, Joshua E. Oppor, Mark P. Mulligan, and Ron J. Koch. Laboratory performance and commissioning status of the SALT NIR integral field spectrograph. In Christopher J. Evans, Julia J. Bryant, and Kentaro Motohara, editors, Ground-based and Airborne Instrumentation for Astronomy IX, volume 12184 of Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series, page 1218407, August 2022. doi: 10.1117/12.2630242.

William H. Wollaston. XII. A Method of Examining Refractive and Dispersive Powers, by Prismatic Reflection. Philosophical Transactions of the Royal Society of London Series I, 92:365–380, January 1802. doi: 10.1098/rstl.1802.0013.

Glossary

FITS extensions Extensions used in FITS files to store different types of data.

‘**BPM**’ The Bad Pixel Map extension in a FITS file.

‘**Primary**’ The Primary extension of a FITS file which contains the file Header.

‘**SCI**’ The Science data extension in a FITS file.

‘**VAR**’ The Variance data extension in a FITS file.

‘**WAV**’ The Wavelength extension in a FITS file.

Johnson-Cousins photometric system A widely-used system of broad-band photometry that uses a set of standard filters to measure the magnitudes of stars and other astronomical objects.

U The ultraviolet filter, typically centered around 3640 Å.

B The blue filter, typically centered around 4420 Å.

V The visual filter, designed to approximate human visual sensitivity and typically centered around 5400 Å.

R The red filter, typically centered around 6580 Å.

I The infrared filter, typically centered around 8060 Å.