Massive stars -> Early-type stars

Host star metallicities -> host-star metallicities or host star metallicities のどちらかに統一

Figure 4の後半. The vertical short dashed line and gray region in each mass regime the mean metallicity and its standard error over 1000 iterations, respectively. The distribution of the samples in terms of host-star metallicity and lower limit of companion mass shows an example among the 1000 calculations.

Figure 5. Distributions of host-star metallicities and masses for the common-biased samples orbiting G-type stars with masses ranging from 0.8 to 1.3 Solar mass(left) and early-type stars with masses more massive than 1.3 Solar mass (right). The symbols are same as those in Figure 4.

Figure 6. Distributions of eccentricities and masses for the common-biased samples orbiting G-type stars (left) and early-type stars (right). The symbols are same as those in Figure 4.

Figure 7の後半. Histograms of mean masses for the common-biased samples with masses less than 20 MJ orbiting G-type stars with masses ranging from 0.8 to 1.3 Msun (cyan bar) and early-type stars with masses more massive than 1.3 Msun (yellow bar) and the simulation samples with masses ranging from 0.1 to 20 MJ.

Figure 7のSemi-major axis and massの分布は必要か？

Figure 8. Distributions of host-star metallicities and companion masses for the common biases samples orbiting G-type stars (left) and early-type stars (right). The common-biased samples (black dots) were compared with expectations from core accretion and disk instability theories in terms of host-star metallicities and planetary mass distributions. The red, green, and blue regions indicate where the objects can be formed by core accretion, disk instability and binary star formation, respectively. The error bars represent the 1-sigma measurement errors.