

Notes

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Contents

1 Self-gravitating thin disk with bulb

1

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Consider a self gravitating disk with surface density Σ and gravitational potential Φ . Let \mathbf{v} be its velocity field and P its pression field. Then it is described by the system

$$\frac{\partial \Sigma}{\partial t} + \nabla \cdot (\Sigma \mathbf{v}) = 0, \quad (1)$$

$$\frac{\partial \mathbf{v}}{\partial t} + (\mathbf{v} \cdot \nabla)(\mathbf{v}) = -\frac{1}{\Sigma} \nabla P - \nabla \Phi, \quad (2)$$

$$\Delta \Phi = 4\pi G \Sigma \delta(z). \quad (3)$$

Here, $\Phi = \Phi_{\text{bulb}} + \Phi_{\text{disk}} + \Phi_{\text{DH}} = \Phi_{\text{bulb}} + \Phi_{\text{sg}}$. Furthermore, Φ_{bulb} is such that it doesn't yield any contribution to the surface density, i.e. $\Delta \Phi_{\text{bulb}} = 0$, nor the pressure field, hence $\Sigma = \Sigma_{\text{disk}} + \Sigma_{\text{DH}}$ and $P = P_{\text{disk}} + P_{\text{DH}}$. Letting $q = \Sigma_{\text{disk}}/\Sigma$, we have

$$\Sigma = \Sigma_{\text{disk}} + \Sigma_{\text{DH}} = \Sigma_{\text{disk}} \left(1 + \frac{\Sigma_{\text{DH}}}{\Sigma_{\text{disk}}} \right) = \frac{1}{q} \Sigma_{\text{disk}}. \quad (4)$$

Suppose that we have a polytrope gas such that $P = \kappa \Sigma^\Gamma$ and let

$$\psi = \int \frac{dP(\Sigma)}{\Sigma} \Leftrightarrow \nabla \psi = \frac{\nabla P}{\Sigma}. \quad (5)$$

Then

$$\psi = \frac{\kappa \Gamma}{\Gamma - 1} \Sigma^{\Gamma-1}. \quad (6)$$

Letting $\Psi = \Phi + \psi$, we obtain

$$\frac{\partial \mathbf{v}}{\partial t} + (\mathbf{v} \cdot \nabla)(\mathbf{v}) = -\nabla \Psi. \quad (7)$$

If $\Sigma_{\text{disk}} \propto \Sigma_{\text{DH}}$ then $q \in [0, 1]$ is constant and this system of equation can be rewritten as

$$\frac{\partial \Sigma_{\text{disk}}}{\partial t} + \frac{1}{r} \frac{\partial (r \Sigma_{\text{disk}} v_r)}{\partial r} + \frac{1}{r} \frac{\partial (\Sigma_{\text{disk}} v_\theta)}{\partial \theta} = 0, \quad (8)$$

$$\frac{\partial v_r}{\partial t} + v_r \frac{\partial v_r}{\partial r} + \frac{v_\theta}{r} \frac{\partial v_r}{\partial \theta} - \frac{(v_\theta)^2}{r} = -\frac{\partial \Psi}{\partial r}, \quad (9)$$

$$\frac{\partial v_\theta}{\partial t} + v_r \frac{\partial v_\theta}{\partial r} + \frac{v_\theta}{r} \frac{\partial v_\theta}{\partial \theta} + \frac{v_r v_\theta}{r} = -\frac{1}{r} \frac{\partial \Psi}{\partial \theta}, \quad (10)$$

$$\Delta \Phi_{\text{sg}} = 4\pi G \frac{1}{q} \Sigma_{\text{disk}} \delta(z). \quad (11)$$

Let M be the total mass of the galactic disk+bulb. Let $\Phi_{\text{bulb}} = -GM(1-x)/\sqrt{c^2+r^2}$ with $x \in [0, 1]$. An equilibrium state is given by the Plummer equilibrium, such that, letting $\xi = (r^2 - a^2)/(r^2 + a^2)$, that is, $r/a = \sqrt{(1+\xi)/(1-\xi)}$,

$$\begin{aligned} v_r^0 &= 0, & (v_\theta^0)^2 &= r \frac{\partial \Psi^0}{\partial r}, \\ \psi^0 &= \frac{\kappa \Gamma}{\Gamma - 1} (\Sigma^0)^{\Gamma-1} = \frac{\kappa \Gamma}{(\Gamma - 1)q^{\Gamma-1}} (\Sigma_{\text{disk}}^0)^{\Gamma-1}, \\ \Sigma_{\text{disk}}^0 &= \frac{xM}{2\pi a^2} \frac{1}{(1 + (r/a)^2)^{3/2}} = \frac{xM}{2\pi a^2} \left(\frac{1-\xi}{2} \right)^{3/2}, \\ \Phi_{\text{sg}}^0 &= -\frac{GMx}{aq} \frac{1}{\sqrt{1 + (r/a)^2}} = -\frac{GMx}{aq} \left(\frac{1-\xi}{2} \right)^{1/2}. \end{aligned}$$

Therefore, letting

$$\varepsilon = \frac{\text{sg internal energy}}{|\text{total sg energy}|} = \frac{3a\kappa\Gamma}{GM} \left(\frac{M}{2\pi a^2} \right)^{\Gamma-1},$$

we obtain

$$\begin{aligned} \Psi^0 &= -\frac{GM(1-x)}{\sqrt{c^2+r^2}} - \frac{GMx}{aq} \left(\frac{1-\xi}{2} \right)^{1/2} + \frac{\kappa\Gamma}{(\Gamma-1)} (\Sigma_{\text{disk}}^0/q)^{\Gamma-1}, \\ &= -\frac{GM(1-x)}{\sqrt{c^2+r^2}} + \frac{GM}{a} \left[-\left(\frac{x}{q} \right) \left(\frac{1-\xi}{2} \right)^{1/2} + \frac{\varepsilon(x/q)^{\Gamma-1}}{3(\Gamma-1)} \left(\frac{1-\xi}{2} \right)^{3(\Gamma-1)/2} \right], \\ &= \frac{GMx}{aq} \left[-\frac{a}{c} \frac{q(1-x)}{x} \frac{1}{\sqrt{1+(r/c)^2}} - \left(\frac{1-\xi}{2} \right)^{1/2} + \frac{\varepsilon(x/q)^{\Gamma-2}}{3(\Gamma-1)} \left(\frac{1-\xi}{2} \right)^{3(\Gamma-1)/2} \right], \\ (v_\theta^0)^2 &= r \frac{\partial \Psi^0}{\partial r} = r \frac{\partial \Psi^0}{\partial \xi} \frac{\partial \xi}{\partial r} = 4 \left(\frac{r}{a} \right)^2 \left(\frac{1-\xi}{2} \right)^2 \frac{\partial \Psi^0}{\partial \xi} = (1+\xi)(1-\xi) \frac{\partial \Psi^0}{\partial \xi}, \\ &= \frac{GM}{a} \left[\frac{a(1-x)}{c} \left(\frac{r}{c} \right)^2 \left(\frac{1}{1+(r/c)^2} \right)^{3/2} + \frac{x}{q} \left(\frac{1+\xi}{2} \right) \left(\frac{1-\xi}{2} \right)^{1/2} \left(1 - \varepsilon \left(\frac{x}{q} \right)^{\Gamma-2} \left(\frac{1-\xi}{2} \right)^{\frac{3\Gamma}{2}-2} \right) \right], \\ &= \frac{GMx}{aq} \left[\frac{a}{c} \frac{q(1-x)}{x} \left(\frac{r}{c} \right)^2 \left(\frac{1}{1+(r/c)^2} \right)^{3/2} + \left(\frac{1+\xi}{2} \right) \left(\frac{1-\xi}{2} \right)^{1/2} \left(1 - \varepsilon \left(\frac{x}{q} \right)^{\Gamma-2} \left(\frac{1-\xi}{2} \right)^{\frac{3\Gamma}{2}-2} \right) \right], \end{aligned}$$

with

$$\frac{1-\xi}{2} = \frac{a^2}{r^2+a^2} = \frac{1}{1+(r/a)^2}; \quad \frac{1+\xi}{2} = \frac{r^2}{r^2+a^2} = \frac{(r/a)^2}{1+(r/a)^2}.$$

Note that for $x = 1$ (no bulb) and $q = 1$ (purely self-gravitating system), we recover the expression of Toomre

$$v_\theta^0 = \left(\frac{GM}{a} \right)^{1/2} \left(\frac{1+\xi}{2} \right)^{1/2} \left(\frac{1-\xi}{2} \right)^{1/4} \sqrt{1 - \varepsilon \left(\frac{1-\xi}{2} \right)^{\frac{3\Gamma}{2}-2}}.$$

The perturbative equation at 1st order are

$$\begin{aligned} \frac{\partial v_r^p}{\partial t} + \frac{v_\theta^0}{r} \frac{\partial v_r^p}{\partial \theta} - 2 \frac{v_\theta^0 v_\theta^p}{r} &= -\frac{\partial \Psi^p}{\partial r}, \\ \frac{\partial v_\theta^p}{\partial t} + v_r^p \frac{\partial v_\theta^0}{\partial r} + \frac{v_\theta^0}{r} \frac{\partial v_\theta^p}{\partial \theta} + \frac{v_r^p v_\theta^0}{r} &= -\frac{1}{r} \frac{\partial \Psi^p}{\partial \theta}, \\ \Delta \Phi_{\text{sg}}^p &= 4\pi G \frac{1}{q} \Sigma_{\text{disk}}^p \delta(z), \\ \psi^p &= \frac{\kappa \Gamma}{q^{\Gamma-1}} (\Sigma_{\text{disk}}^0)^{\Gamma-2} \Sigma_{\text{disk}}^p. \end{aligned}$$

We define $X^p(r, \theta, t) = \sum_{m \in \mathbb{Z}} X_m^p(r, t) e^{im\theta}$ and look for a temporal dependency in $e^{-i\omega t}$. Aoki & Iye say that there is this following correspondance between surface density and gravitational potential through the Poisson equation:

$$(\Sigma_{\text{disk}})_m^p(r, \theta) = \frac{xM}{2\pi a^2} \left(\frac{1-\xi}{2} \right)^{3/2} \sum_{n=|m|}^{\infty} a_n^m \widehat{P_n^{|m|}}(\xi) e^{-i\omega t}, \quad (12)$$

$$(\Phi_{\text{sg}})_m^p(r, \theta) = -\frac{GMx}{aq} \left(\frac{1-\xi}{2} \right)^{1/2} \sum_{n=|m|}^{\infty} \frac{a_n^m}{2n+1} \widehat{P_n^{|m|}}(\xi) e^{-i\omega t}, \quad (13)$$

$$(\psi)_m^p(r, \theta) = \kappa \Gamma \left(\frac{M}{2\pi a^2} \right)^{\Gamma-1} \left(\frac{x}{q} \right)^{\Gamma-1} \left(\frac{1-\xi}{2} \right)^{3/2(\Gamma-1)} \sum_{n=|m|}^{\infty} a_n^m \widehat{P_n^{|m|}}(\xi) e^{-i\omega t}, \quad (14)$$

$$(\Psi)_m^p(r, \theta) = \frac{GMx}{aq} \left(\frac{1-\xi}{2} \right)^{1/2} \sum_{n=|m|}^{\infty} \left[\frac{\varepsilon}{3} \left(\frac{x}{q} \right)^{\Gamma-2} \left(\frac{1-\xi}{2} \right)^{\frac{3\Gamma}{2}-2} - \frac{1}{2n+1} \right] a_n^m \widehat{P_n^{|m|}}(\xi) e^{-i\omega t}. \quad (15)$$

Let us decompose the velocity components based on their equilibrium expression:

$$(v_r)_m^p = i \frac{m}{|m|} \left(\frac{GMx}{aq} \right)^{1/2} \left(\frac{1+\xi}{2} \right)^{-1/2} \left(\frac{1-\xi}{2} \right)^{1/4} \sum_{n=|m|}^{\infty} b_n^m \widehat{P_n^{|m|}}(\xi) e^{-i\omega t}, \quad (16)$$

$$(v_\theta)_m^p = \left(\frac{GMx}{aq} \right)^{1/2} \left(\frac{1+\xi}{2} \right)^{-1/2} \left(\frac{1-\xi}{2} \right)^{1/4} \sum_{n=|m|}^{\infty} c_n^m \widehat{P_n^{|m|}}(\xi) e^{-i\omega t}. \quad (17)$$

Letting $X_m^p = X^1 e^{-i\omega t}$, this yields the set of equations

$$i(-\omega + m\Omega)(\Sigma_{\text{disk}})^1 + \frac{1}{r} \frac{d(r\Sigma_0(v_r)^1)}{dr} + \frac{im\Sigma^0(v_\theta)^1}{r} = 0, \quad (18)$$

$$\frac{d(\Psi)^1}{dr} + i(-\omega + m\Omega)(v_r)^1 - 2\Omega(v_\theta)^1 = 0, \quad (19)$$

$$im \frac{(\Psi)^1}{r} + \frac{\alpha^2}{2\Omega}(v_r)^1 + i(-\omega + m\Omega)(v_\theta)^1 = 0, \quad (20)$$

where $\Omega = v_\theta^0/r$ is the angular velocity and $\alpha^2 = 4\Omega^2[1 + r/(2\Omega) \cdot (d\Omega/dr)]$ is the epicyclic frequency.