

ULX spectra revisited: Are accreting, highly magnetized neutron stars the engines of ultraluminous X-ray sources?

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ABSTRACT

Aims. In light of recent discoveries of pulsating ultraluminous X-ray sources (ULXs) and recently introduced theoretical schemes, that propose neutron stars (NSs) as the central engines of ULXs, we revisit the spectra of eighteen well known ULXs, in search of indications that favor this newly emerging hypothesis.

Methods. We examine the spectra from high quality *XMM-Newton* and *NuSTAR* observations. We use a combination of elementary black body and multicolor disk black body (MCD) models, to diagnose the predictions of classic and novel theoretical models of accretion onto NSs. We re-interpret the well established spectral characteristics of ULXs in terms of accretion onto lowly or highly magnetized NSs, and explore the resulting parameter space for consistency.

Results. We confirm the, previously noted, presence of the low energy (≤ 6 keV) spectral rollover and argue that it could be interpreted as due to thermal emission. The spectra are well described by a double thermal model consisting of a “hot” (≥ 1 keV) and a “cool” (≤ 0.7 keV) multicolor black body (MCB). Under the assumption that the “cool” MCD emission originates in a disk truncated at the neutron star magnetosphere, we find that all ULXs in our sample are consistent with accretion onto a highly magnetized ($B \geq 10^{12}$ G) neutron star. We note a strong correlation between the strength of the magnetic field, the temperature of the “hot” thermal component and the total unabsorbed luminosity. Examination of the *NuSTAR* data supported this interpretation and also confirmed the presence of a weak, high energy (≥ 15 keV) tail, most likely the result of modification of the MCB emission by inverse Compton scattering. We also note that the apparent high energy tail, may simply be the result of mis-modeling of MCB emission with an atypical temperature (T) vs radius (r) gradient, using a standard MCD model with a fixed gradient of $T \sim r^{-0.75}$.

Conclusions. We have offered a new and robust physical interpretation for the dual-thermal spectra of ULXs. We find that the best-fit derived parameters of our model, are in excellent agreement with recent theoretical predictions that favor super-critically accreting NSs as the engines of a large fraction of ULXs. Nevertheless, the considerable degeneracy between models and the lack of unequivocal evidence cannot rule out other equally plausible interpretations. Deeper broadband observations and time resolved spectroscopy are warranted to further explore this newly emerging framework.

Key words. Keywords should be given

1. Introduction

Ultraluminous X-ray sources (ULXs) are off-nuclear, extragalactic X-ray sources with isotropic luminosities that exceed the Eddington limit for a stellar mass black hole ($M_{\text{BH}} \lesssim 20 M_{\odot}$; see [Feng & Soria 2011](#) and references therein, and also [Bachetti 2016](#) and [Kaaret et al. 2017](#) for up-to-date and comprehensive reviews). It was initially suggested that ULXs were rare instances of intermediate mass black holes accreting at sub-Eddington rates ([Colbert & Mushotzky 1999](#); [Makishima et al. 2000](#); [Kaaret et al. 2001](#); [Miller et al. 2003](#)). Essentially a scaled up version of standard galactic BH X-ray binaries (BH-XRBs). However, it was quickly realized that a considerable fraction (if not all) of the ULX population can be powered by stellar mass BHs accreting at super-Eddington rates (e.g. [Gao et al. 2003](#); [Gilfanov et al. 2004](#); [Roberts et al. 2004](#); [Poutanen et al. 2007](#); [King 2009](#)). Furthermore, the astounding discovery of a pulsating ULX (PULX: [Bachetti et al. 2014](#)) demonstrated that ULXs can be powered by neutron stars (NSs). After the initial dis-

covery by [Bachetti et al.](#), two more PULXs have been detected ([Israel et al. 2017](#); [Furst et al. 2016](#)).

The discovery of NS-ULXs provided further support to the scenario of super-Eddington accretion onto a stellar mass compact object as the power source of ULXs, but also posed the crucial question of the potential prevalence of NSs as the engines of ULXs. Indeed, it has been known for some time that most of the ULXs do not transition through the phenomenological states of standard BH-XRBs, i.e. spectral transitions between hard and soft states and the appearance and quenching of relativistic jets (for a review of spectral and temporal properties of NS- and BH-XRBs see e.g. [Barret 2001](#); [Remillard & McClintock 2006](#); [Done et al. 2007](#); [Gilfanov 2010](#)). While some ULXs exhibit significant luminosity and spectral variations (e.g. [Swartz et al. 2004](#); [Sutton et al. 2013](#); [Feng et al. 2016](#)), they do not transition between the markedly different and characteristic “hard” and “soft” spectral states of nominal, sub-Eddington BH-XRBs (e.g. [Berghia et al. 2008](#); [Grisé et al. 2010](#); [Sutton et al. 2013](#)). Interestingly two hyperluminous X-ray sources (HLXs) – which have luminosities exceeding $\sim 10^{41}$ erg s⁻¹ and are

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strong IMBH candidates – seem to follow the transition pattern of stellar mass BH-XRBs (ESO 243-49 HLX-1: [Godet et al. 2009](#); [Servillat et al. 2011](#); [Webb et al. 2012](#), or M82 X-1: [Feng & Kaaret 2010](#)). In this work we only consider ULXs which are limited to luminosities $\lesssim 10^{41}$ erg/s.

In addition to the lack of state transitions, the spectra of numerous ULXs are significantly different from the typical spectra of BH-XRBs. More specifically, the majority of ULX spectra feature a notable spectral curvature above ~ 6 keV (e.g. [Stobbart et al. 2006](#); [Gladstone et al. 2009](#); [Sutton et al. 2013](#)). It has been proposed that perhaps the unusual spectra of ULXs correspond to a novel state of super-critical accretion, dubbed the *ultraluminous state* ([Gladstone et al. 2009](#)). In subsequent works it was posited that ULX spectra can be empirically classified into three classes based on their spectral shape: singly peaked *broadened disc* (BD) class and two-component hard ultraluminous (HUL) and soft ultraluminous (SUL) states ([Sutton et al. 2013](#)). Recognizing the physical mechanisms underlying the observed spectral characteristics of ULXs is a crucial step towards understanding accretion at super-Eddington states but also determining the nature of their accretor.

In the recent months, leading to this publication, an increasing number of compelling theoretical considerations – that point to NSs as plausible engines behind ULX emission – have been put forward by numerous authors (e.g. [King & Lasota 2016](#); [King et al. 2017](#); [Mushtukov et al. 2017](#)). Motivated by these findings and the apparent spectral and temporal similarities between ULXs and NS-XRBs we decided to revisit the spectra of known ULXs, in search of indications that may favor this newly emerging trend. More specifically, we ask whether the curvature of ULX spectra is due to hot thermal emission, rather than a hard power law with a cutoff at an improbably low energy and whether this can be physically interpreted in terms of emission from a super-critically accreting NS. To investigate this hypothesis, we have selected eighteen well known ULXs that have been studied by multiple authors in the past and were also included in the samples used in the seminal works of [Stobbart et al. \(2006\)](#) and [Sutton et al. \(2013\)](#), in which the spectral shape of ULXs was standardized and classified observationally. Below (Section 2) we briefly present the different interpretations of the spectral curvature in ULXs, the latest theoretical arguments for the nature of their central engine and the motivation behind our choice to revisit the spectra of known ULXs. In Section 3 we present the details of our data analysis and in Section 4 we discuss our findings and their implications with regard to the nature of the accretor in our sample and in ULXs in general.

2. Origin of the curvature in the spectra of ULXs

2.1. Optically thick, “warm” corona

Nominal BH-XRBs exhibit a power-law shaped tail towards high energies, during all spectral states. More specifically, during episodes of low-luminosity advection dominated accretion (also known as a *hard state*) the spectrum is dominated by a hard power law (spectral index of up to ~ 1.5) with an exponential cutoff at ~ 100 - 200 keV, while in the high-luminosity *soft state*, the power-law component becomes, less prominent and softer (spectral index exceeding ~ 2), but without an observable spectral cutoff (e.g., [Barret 2001](#); [Gallo et al. 2003](#); [Narayan 2005](#); [Remillard & McClintock 2006](#); [Done et al. 2007](#); [Gilfanov 2010](#)). As discussed above, the spectra of many ULXs (including all sources studied in this work) feature a spectral curvature and an abrupt drop in spectral counts at considerably lower en-

ergies than standard BH-XRBs. The spectral roll-over in ULXs was detected in *XMM-Newton* data of numerous sources (e.g. [Stobbart et al. 2006](#); [Sutton et al. 2013](#)) and was subsequently confirmed with the *NuSTAR* telescope (e.g. [Bachetti et al. 2013](#); [Walton et al. 2014, 2015](#); [Fürst et al. 2017](#)). It is often observed as low as ~ 3 keV.

In the past, many authors have modeled the unusual low-energy curvature of ULX spectra (including the ones revisited here) using a power-law spectrum with a low-energy cutoff and a relatively large e-folding energy (e.g. [Stobbart et al. 2006](#); [Roberts 2007](#); [Gladstone et al. 2009](#); [Sutton et al. 2013](#)). The uncommonly low-energy cutoff, is often linked to the presence of a corona of hot, thermal electrons with an unusually high Thomson scattering optical depth. Namely, multiple authors have considered that the shape of the high energy part of the spectrum is the result of thermal Comptonization of soft ($h\nu < 0.5$ keV) photons, by a corona of thermal electrons with an optical depth that often exceeds $\tau_T \approx 20$ (e.g. [Stobbart et al. 2006](#); [Roberts 2007](#); [Gladstone et al. 2009](#); [Pintore & Zampieri 2012](#); [Sutton et al. 2013](#); [Pintore et al. 2014](#)). While this configuration successfully reproduces the observed spectral shapes, its feasibility under realistic circumstances in the vicinity of critically accreting X-ray binaries may be problematic. More specifically, the high scattering depth, combined with the increased photon density will pause a significant burden to the coronal thermalization. This can be illustrated in the following simplified example, where we consider cooling due to Inverse Compton (IC) scatterings and Coulomb collisions as the sole thermalization mechanism. The cooling rate of thermal electrons due to multiple IC scatterings by photons with $h\nu \ll kT_e$ depends strongly on the scattering optical depth of the corona, i.e. the cooling timescale for large optical depth is (e.g. [Rybicki & Lightman 1979](#))

$$t_{\text{cool}} \approx \frac{m_e c^2 (R/c)}{4h\nu \tau_T^2}, \quad (1)$$

where m_e is the electron mass and R/c the characteristic source size, which can be inferred from variability considerations. It is obvious from eq. 1 alone that for the values of τ_T reported in these works, Compton cooling will be very rapid, i.e. $t_{\text{cool}} \ll R/c$. Below we illustrate that the cooling timescale will be too small to allow for electron thermalization.

The thermalization will occur primarily through energy exchange between high energy electrons and the thermal background of electrons and protons in the coronal plasma. For a high scattering optical depth (i.e. $\tau_T > 5$), one can plausibly assume that the main mechanism for the energy exchange will be Coulomb interactions between electrons and electrons, and electrons and protons, ignoring – for example – collective interactions between particles (e.g. [Begelman & Chiueh 1988](#)). If the relativistic electrons exchange energy more rapidly than they cool due to multiple scatterings, then the plasma will thermalize efficiently. The timescale of the Coulomb energy exchange will be

$$t_{\text{exch}} \sim \frac{E}{|dE/dt|}, \quad (2)$$

where E is the electron energy and dE/dt is the Coulomb cooling rate ([Gould 1975](#); [Frankel et al. 1979](#); [Coppi & Blandford 1990](#)). For relativistic electrons the Coulomb rate can be rewritten as (see also [Coppi 1999](#)),

$$\frac{dE}{dt} \sim -\tau_T \ln \Lambda \left(\frac{R}{c} \right)^{-1} \quad (3)$$

where $\ln \Lambda$ is the usual Coulomb logarithm. From the above approximations, it becomes evident that while the Coulomb cooling becomes more rapid, as optical depth increases ($t_{\text{exch}} \sim \tau_T^{-1}$), the τ_T^{-2} dependence of the Compton cooling timescale, results in a corona in which energetic photons cool down before they can thermalize. As τ_T increases, a thermal corona becomes more and more difficult to sustain. The problem of coronal thermalization is well known, and it can also become significant in the low τ_T regime (e.g. Coppi 1999, and references therein), particularly for hot ($kT_e \gtrsim m_e c^2$) coronas. As a result the presence of hybrid thermal/non-thermal electron distributions (e.g. Coppi 1999; Belmont et al. 2008; Malzac & Belmont 2009) are usually assumed. However, the high scattering optical depths that were claimed in earlier ULX literature (e.g. Stobbart et al. 2006; Roberts 2007; Gladstone et al. 2009; Sutton et al. 2013), present major sustainability issues even when considering hybrid, “warm” coronas with lower electron temperatures. It is only under very tight restrictions that coronas with $\tau_T > 5$ can be sustained (e.g. Róžańska et al. 2015). We must stress here, that the issues concerning the physical plausibility of the optically-thick corona model have been noted by the community since quite early (e.g. Middleton et al. 2011; Kajava et al. 2012), with the majority of later works, only using the power law with the exponential cutoff as an empirical description of the hard emission of ULXs, rather than a physical interpretation.

2.2. Optically thick winds

The most widely accepted interpretation of the spectral curvature of ULXs invokes the presence of strong, optically thick winds. Namely, under the assumption that ULXs are accreting BHs in the super-Eddington regime, then the spectral shape of the emission may be strongly influenced by the presence of massive, optically thick outflows. King & Pounds (2003) and Poutanen et al. (2007) argued that the curvature of the spectra (of at least some) ULXs can be interpreted in terms of reprocessing of the primary emission in the optically thick wind. In this scenario the soft thermal emission is associated with the wind itself, and the hard emission is also thermal and it originates in the hot, innermost parts of an accretion disk. Kajava & Poutanen (2009) followed up the predictions of Poutanen et al. (2007) and by studying the spectra of eleven known ULXs they claimed that the temperature of the soft thermal component decreased with its luminosity (i.e. $L_{\text{soft}} \sim T_{\text{in}}^{-3.5}$), in agreement with the prediction for emission from an optically thick wind. However, subsequent studies by (e.g. Miller et al. 2013) found that the luminosity of the soft component correlates positively with temperature, approximately following the $L_{\text{soft}} \sim T_{\text{in}}^4$ relations expected for standard accretion disks. However, in a recent study of numerous longterm observations of HO IX X-1, Luangtip et al. (2016) show that – with increasing luminosity – the source spectra evolve from a two component spectrum to a (seemingly) single component, thermal-like spectrum. The authors argue that the apparent heating of the soft disk component may be model depended. An artifact caused by neglecting to properly account for the evolving spectra.

In addition to the prediction of the ULX spectral shape, the optically thick wind model also offers an interpretation for the short timescale variability noted in many sources (e.g. Sutton et al. 2013). More specifically, Middleton et al. (2015a) combined the arguments of King & Pounds (2003) and Poutanen et al. (2007) with predictions for wind inhomogeneity (e.g. Takeuchi et al. 2013) and mass accretion rate fluctuations (e.g. Ingram & van der Klis 2013) to account for spec-

tral and timing variability of nine ULXs. The authors made a well-founded case for super-Eddington accretion onto stellar-mass BHs as the driving force behind ULXs. In this scheme the fractional variability noted by previous authors, is explained in terms of a “clumpy” wind partially obscuring the hard component, which appears to fluctuate. In the same context the different empirical states indicated by Sutton et al. (2013), are the result of different orientations between the observer and the disk/wind structure (see Fig. 1 of Middleton et al. 2015a). Building on these considerations several authors (e.g. Walton et al. 2014, 2015; Luangtip et al. 2016) have modeled ULX spectra extracted from *XMM-Newton* and *NuSTAR* data using a dual thermal model, in which the soft thermal emission is attributed to the optically thick wind and the hotter component to emission from the inner parts of a hot accretion disk. The presence of the strong outflows are also supported by the uniform optical counterparts of numerous ULXs, which indicate a hot wind origin (e.g. Pakull & Mirioni 2002; Fabrika et al. 2015), but also the presence of wind or jet blown, radio “bubbles” around some ULXs (e.g. Pakull & Mirioni 2003; Soria et al. 2006; Cseh et al. 2012). The strongest indication of an outflow lies in the discovery of soft X-ray spectral residuals near ~ 1 keV (e.g. Middleton et al. 2014, 2015b; Pinto et al. 2016, 2017), which have been interpreted as a direct signature of their presence. While broad emission- and absorption-like residuals, near the ~ 1 keV mark, have been observed in the spectra of numerous NS- and BH-XRBs, during different states and at different mass accretion rates (e.g. Boirin & Parmar 2003; Díaz Trigo et al. 2006; Ng et al. 2010; Kolehmainen et al. 2014), the absorption lines detected in the spectra of NGC 1313 X-1 and NGC 5408 X-1 (Pinto et al. 2016) and more recently in NGC 55 ULX (Pinto et al. 2017) seem to be strongly blue-shifted, indicating outflows with velocities reaching $0.2 c$.

2.3. The case of accreting Neutron Stars

The recent discoveries of the three PULXs, has established the fact that ULXs can be powered by accretion onto NSs. This realization is perhaps not surprising, considering that a mechanism that allows accretion at super-Eddington rates, onto highly magnetized ($B \gtrsim 10^{12}$ G) NSs, has been put forward since the late seventies (Gnedin & Sunyaev 1973; Basko & Sunyaev 1975, 1976). However, these works were not aimed at discussing super-Eddington accretion in the context of ULXs. The authors were attempting to resolve the complications that arise from the fact that, when material is accreted onto a very small area on the surface of the NS, the Eddington limit is considerably lower than the $\sim 1.8 \times 10^{38}$ erg/s, which corresponds to isotropic accretion onto a NS. Therefore, persistent X-ray pulsars with luminosities exceeding \sim a few 10^{37} erg/s, where in fact breaking the (local) Eddington limit. More specifically, in high-B accreting NSs, the accretion disk is interrupted by the magnetic field near the magnetospheric radius, at which point the accreted material is guided by the magnetic field lines onto a small area centered around the NS magnetic poles (e.g. Pringle & Rees 1972; Romanova et al. 2012). The resulting formation, is known as an accretion column. Due to the high anisotropy of the photoelectron scattering cross-section, in the presence of a strong magnetic field (Canuto et al. 1971; Lodenquai et al. 1974), the emission from the accretion column is concentrated in a narrow (“pencil-”) beam, which is directed parallel to the magnetic field lines (and hence the magnetic field axis, Basko & Sunyaev 1975). However, at high mass accretion rates a radiation dominated shock is formed at \sim few km above the surface of the NS.

As accretion rate exceeds a critical value (corresponding to a critical luminosity of $\sim 10^{37}$ erg/sec, e.g. [Basko & Sunyaev 1976](#); [Mushtukov et al. 2015b](#)), the accretion funnel is suffused with high-density plasma which is gradually sinking in the gravitational field of the NS ([Basko & Sunyaev 1976](#); [Wang & Frank 1981](#)). As a result the accretion funnel, in the direction parallel to the magnetic field axis, becomes optically thick and the emerging X-ray photons mostly escape from its – optically thin – sides, in a “fan-beam” pattern (see e.g. Fig.1 [Schönherr et al. 2007](#)). In recent refinements of this mechanism, it has been demonstrated that depending on the magnetic field strength and the pulsar spin it can facilitate luminosities of the order of 10^{40} erg/sec ([Mushtukov et al. 2015a](#)).

Observations of multiple X-ray Pulsars have yielded an empirical description of the primary emission of the accretion column as a very hard power law (spectral index $\lesssim 1.8$) with a low energy ($\lesssim 10$ keV) cutoff (e.g. [Caballero & Wilms 2012](#)). While a general, self consistent description of the spectral shape of the accretion column emission has not been achieved yet, several authors have attempted to reproduce it (e.g. [Nagel 1981](#); [Meszaros & Nagel 1985](#); [Burnard et al. 1991](#); [Hickox et al. 2004](#); [Becker & Wolff 2007](#)). More specifically, [Becker & Wolff \(2007\)](#) have reproduced the spectrum, assuming bulk and thermal Comptonization of seed Bremsstrahlung, black body and cyclotron photons.

While, in the high field regime, super-Eddington accretion can be efficiently sustained, lowly magnetized NSs can only reach moderately super-Eddington luminosities and only in the soft state of the so-called Z-sources (e.g. [Muno et al. 2002](#); [Done et al. 2007](#); [Lin et al. 2009](#)). When the accretion rate reaches and exceeds the Eddington limit, strong outflows are expected to inhibit higher accretion rates. Nevertheless, in a recent publication [King & Lasota \(2016\)](#) argue that super-Eddington accretion onto lowly magnetized NSs, can be maintained – along with powerful outflows – in a similar fashion to super-Eddington accretion onto BHs ([King et al. 2001](#)). Therefore, a considerable fraction of (non-pulsating) ULXs may be the result of beamed emission from NSs with relatively weak magnetic fields ($B < 10^{11}$ G), accreting at high mass transfer rates.

Intriguingly, one of the first ([Mitsuda et al. 1984](#)) and most frequently observed spectral characteristics associated with the soft state of Z-sources, is the presence of two thermal components (e.g. [White et al. 1988](#); [Mitsuda et al. 1989](#); [Barret 2001](#); [Lin et al. 2007](#); [Revnivtsev et al. 2013](#)). The “cool” thermal component, most likely, originates in a thin [Shakura & Sunyaev](#) disk and the additional “hot” thermal component corresponds to emission produced in hot optically thick plasma on the the surface of the NS, known as a *boundary layer* ([Sunyaev & Shakura 1986](#); [Sibgatullin & Sunyaev 2000](#)). Therefore, the presence of a dual thermal spectrum in XRBs strongly suggests emission from a solid surface, indicating a lowly magnetized NS. Nevertheless, in a new publication by [Mushtukov et al. \(2017\)](#) it is argued that in accreting high-B NSs, the – normally optically thin (e.g. [Basko 1980](#); [Nagase et al. 1992](#); [Ebisawa et al. 1996](#)) – material trapped in the Alfvén surface becomes optically thick as the luminosity exceeds $\sim 5 \cdot 10^{39}$ erg/s. The emission of the resulting structure will have a quasi-thermal spectrum at temperatures exceeding 1 keV. Combined with the soft thermal emission from a truncated accretion disk, the spectra of highly magnetized NSs may also feature the same dual thermal shape as high state Z-sources ([Mushtukov et al. 2017](#), see more details in Sections 3 and 4). Based on these considerations, it becomes apparent that the reanalysis of ULX spectra is warranted. To this end, we will explore the relevance of models, usually implemented in the

modeling of emission from NS-XRBs, in the context of ULXs. More importantly, we will investigate whether our best-fit models yield parameter values that are physically meaningful and in accordance with the predictions for the emission from ULXs.

3. Spectral extraction and data analysis

We have analyzed archival *XMM-Newton* observations of eighteen ULXs presented in Table 1. We have selected sources that have been studied extensively in the past and are confirmed ULXs. Furthermore, specific datasets were selected to have a sufficient number of counts to allow robust discrimination between the different models used to describe their spectral continuum. Except for these conditions, sources were chosen randomly. Therefore, the ULX sample analyzed in this work is not complete. Nevertheless, the purpose of this work is not to revisit the entire ULX catalog, but to demonstrate that a significant fraction of ULXs follow a specific pattern (presented below). For this purpose, our source sample is sufficiently extensive. For six of these sources we also analyzed archival *NuSTAR* data.

3.1. *XMM-Newton* Spectral Extraction

For the *XMM-Newton* data, we only considered the EPIC-pn detector which has the largest effective area of the three EPIC detectors, in the full 0.3–10 keV bandpass, and had registered more than ~ 1000 photons for each of the observations considered, thus providing sufficient statistics to robustly discriminate between different spectral models while ensuring simplicity and self-consistency in our analysis. Therefore the following description of data analysis refers only to this instrument. The data were handled, using the *XMM-Newton* data analysis software SAS version 15.0.0. and the calibration files released¹ on January 22nd, 2016. In accordance with the standard procedure, we filtered all observations for high background flaring activity, by extracting high-energy light curves ($10 < E < 12$ keV) with a 100 s bin size. By placing appropriate threshold count rates for the high-energy photons, we filtered out time intervals that were affected by high particle background. During all observations pn was operated in Imaging Mode. In the majority of our sources, spectra were extracted from a circular region with a radius $> 35''$, centered at the core of the point spread function (psf) of each source. We thus ensured the maximum encircled energy fraction² within the extraction region. This was not possible in the case of NGC 4861 ULX1, M81 X-6 and NGC 253 ULX2 where we used spectral extraction apertures of $18''$, $13.8''$ and $12.5''$ in order to avoid contamination by adjacent sources or due to the proximity of our source to the edge of the chip³. The extraction and filtering process followed the standard instructions provided by the *XMM-Newton* Science Operations Centre (SOC⁴). More specifically, spectral extraction was done with SAS task `evselect`, using the standard filtering flags (`#XMMEA_EP && PATTERN<=4` for pn) and SAS tasks `rmfgen` and `arfgen` were respectively used to create the redistribution matrix and ancillary file. All spectra were regrouped to have at least 25 counts per bin and analysis

¹ XMM-Newton CCF Release Note: XMM-CCF-REL-332

² See *XMM-Newton* Users Handbook §3.2.1.1

http://xmm-tools.cosmos.esa.int/external/xmm_user_support/documentation/uhb/onaxisxraypsf.html

³ When part of the psf lies in a chip gap, effective exposure and encircled energy fraction may be affected

⁴ <http://www.cosmos.esa.int/web/xmm-newton/sas-threads>

was performed using the `xspec` spectral fitting package, version 12.9.0 (Arnaud 1996).

3.2. *NuSTAR* spectral extraction

The *NuSTAR* data were processed using the version 1.6.0 of the *NuSTAR* data analysis system, (*NuSTAR* DAS). We downloaded all public *NuSTAR* datasets using the “`heasarc_pipeline`” scripts (Multimission Archive TeamOAC, in prep.). These have already been processed to obtain L1 products. We then ran `nuproducts` using a 50'' extraction region around the main source and a 50-80'' extraction region for background, in the same detector as the source when possible, as far as we could to avoid contributions from the point-spread function (PSF) wings. We applied standard PSF, alignment, and vignetting corrections. Spectra were rebinned in order to have at least 30 counts bin⁻¹ to ensure the applicability of the χ^2 statistics. All sources in our sample dominate the background up to 20-30 keV. The models we use are relatively simple, and the physical interpretation does not change considerably for a change of best-fit parameters of 10-20%, and so we do not need an extremely precise modeling of the background.

3.3. Spectral analysis

3.3.1. *XMM-Newton*

The spectral continuum was modeled twice, firstly using a combination of a multicolor disk black body (MCD) and a black body component (`diskbb+bbbody`), with the black body (`bbbody`) acting as the hot thermal component and secondly using two MCD components (`diskbb+diskbb`). The first model was used because it is the most widely used model describing the spectra of NS-XRBs in the high-accretion, “soft” state. Our choice for the second model was based on the recent theoretical predictions by Mushtukov et al. (2017), where it is argued that critically accreting NSs with a high magnetic field (i.e. $B \gtrsim 10^{12}$ G), can become engulfed in an optically thick toroidal envelope which is the result of accreting matter moving along the magnetic field lines. Emission from the optically thick envelope is predicted to have a multicolor black body spectrum, with a temperature exceeding ~ 1 keV (more details in section 4). We model this hot thermal emission using the `diskbb` model, because it is the simplest and most reliable multi-temperature black body model in `xspec`, however we stress that we do not expect this emission to originate from a disk. Therefore, the inner disk radius corresponding the hot `diskbb` component, has no physical meaning and is not tabulated. In both models the cool disk component is modeled as a geometrically thin, optically thick Shakura & Sunyaev (1973) disk, which is expected to extend inwards until it reaches the surface of the NS, unless it is disrupted by strong outflows or a strong magnetic field (more details in section 4). We did not model intrinsic and/or host galaxy absorption separately from the Galactic absorption, but used one component for the total interstellar absorption, which was modeled using `tbnew_gas`, the latest improved version of the `tbabs` X-ray absorption model (Wilms et al. 2011).

For the dual MCD model, we assume that the disk becomes truncated at \sim the magnetospheric radius at which point the material follows the magnetic field lines to form the hypothetical envelope. Under this assumption, we also estimate the strength of the magnetic field (B), assuming that the inner radius of the “cool” `diskbb` coincides with the magnetospheric radius (R_{mag}) and using the expression for R_{mag} given by Lai (2014; see

also eq. 1 from Mushtukov et al. 2017). The complete `xspec` model used in the spectral fits is `tbnew_gas(cflux*diskbb + cflux*(diskbb or bbody))`. Where `cflux` is a convolution model that is used to calculate the flux of the two thermal components. Some of the sources exhibited strong residuals in the 0.5-1.2 keV region, commensurate with X-ray emission lines from hot, optically thin plasma. The emission features were modeled using the `mekal` model, which models the emission spectrum of a hot diffuse gas. The best fit parameters of the continuum were not sensitive to the modeling of these features (e.g. using a gaussian instead of `mekal`), however they are strongly required by the fit ($\delta\chi^2 > 15$ for two dof, in all sources). More specifically plasma temperature for Ho IX X-1 was ~ 0.92 keV, for M81 X-6 ~ 1.09 keV, for M83 ULX ~ 0.95 keV, NGC 4736 ULX ~ 0.40 keV and for NGC 5204 X-1 ~ 1.08 keV. While soft X-ray atomic features may be crucial to our understanding of the nature of ULXs (namely, the presence of strong winds and the chemical composition of the accreted material, e.g.: Middleton et al. 2015a,b; Pinto et al. 2016), they are not the focus of this work and are only briefly discussed in Section 4, but not studied further.

Given the high ($\gtrsim 1$ keV) temperatures of the hot thermal component in both double-thermal models, it is expected that electron scattering will have a significant effect on the resulting spectrum, as it becomes comparable to free-free absorption. Therefore, the actual emission will be radiated as a “diluted” black-body, which when modeled using a prototypical thermal model like `diskbb` or `bbbody` will result in temperature and radius estimations which will deviate from their “true” values. This issue is commonly addressed by considering a correction factor (f_{col}) that approximately accounts for the spectral modification (London et al. 1986; Lapidus et al. 1986; Shimura & Takahara 1995). This factor is often referred to as a *color correction factor* and detailed calculations, combined with multiple observations have demonstrated that it depends weakly on the size⁵ of the emitting region and the mass accretion rate (e.g. Shimura & Takahara 1995). Therefore in the 1st approximation it can be considered independent of these parameters and its value is estimated between ~ 1.5 and ~ 2.1 (e.g. Zimmerman et al. 2005, and references therein). The color correction factor affects both the temperature and normalization of the thermal models (i.e. `diskbb` and `bbbody`), with the corrected values given by

$$T_{\text{cor}} = \frac{T}{f_{\text{col}}} \quad (4)$$

and

$$R_{\text{in,cor}} = R_{\text{in}} f_{\text{col}}^2 \quad (5)$$

where T is the temperature of the MCD component and R_{in} is the inner radius. Although the spectral hardening effects are expected to be noticeable, particularly in the hot thermal component, we have decided not to include the color correction in any of our calculations and to tabulate and plot the values of all quantities of interest as provided by our best fits. The reader is, however, advised to note that the value of our results may vary by a value of $\sim f_{\text{col}}$.

The value of the inferred inner disk radius is also dependent on the viewing angle (i) of the accretion disk (i.e. $R_{\text{in}} \sim 1/\sqrt{\cos i}$). This dependence may become important in the estimation of R_{in} if the accretion disk is viewed at a large inclination angle (i.e. edge-on view). Nevertheless, since we have no indications

⁵ or inner radius in the case of an accretion disk

Table 1: Observation Log

| Source | Distance ^a | ObsID | Date | Duration ^b | Rate ^c | Obs. Mode | Filter | Position |
|-------------------|-----------------------|-------------|------------|-----------------------|-------------------------------------|-----------|--------|----------|
| | Mpc | | | s | 10 ⁻¹ ct s ⁻¹ | | | |
| <i>XMM-Newton</i> | | | | | | | | |
| Ho II X-1 | 3.27±0.60 | 0200470101 | 2004-04-15 | 44130 | 6.97±0.04 | FF | Medium | On-Axis |
| Ho IX X-1 | 3.77±0.80 | 0200980101 | 2004-09-26 | 75900 | 14.5±0.05 | LW | Thin | On-Axis |
| IC 342 X-1 | 2.73±0.70 | 0206890201 | 2004-08-17 | 16970 | 3.96±0.05 | EFF | Medium | On-Axis |
| M33 X-8 | 0.91±0.50 | 0102640101 | 2000-08-04 | 7144 | 55.1±0.28 | FF | Medium | On-Axis |
| M81 X-6 | 3.61±0.50 | 0111800101 | 2001-04-22 | 79660 | 4.38±0.02 | SW | Medium | On-Axis |
| M83 ULX | 4.66±0.70 | 0110910201 | 2003-01-27 | 19760 | 1.20±0.03 | EFF | Thin | ~6.5' |
| NGC 55 ULX | 1.60±0.20 | 0028740201 | 2001-11-14 | 34442 | 1.25±0.07 | FF | Thin | On-Axis |
| NGC 253 ULX2 | 3.56±0.80 | 0152020101 | 2003-06-19 | 65010 | 2.57±0.02 | FF | Thin | On-Axis |
| NGC 253 XMM2 | 3.56±0.80 | 0152020101 | 2003-06-19 | 64850 | 2.41±0.02 | FF | Thin | ~4.1' |
| NGC 1313 X-1 | 4.25±0.80 | 0106860101 | 2000-10-17 | 19880 | 6.89±0.01 | FF | Medium | ~5.4' |
| NGC 1313 X-2 | 4.25±0.80 | 0405090101 | 2006-10-15 | 80470 | 6.19±0.03 | FF | Medium | ~4.0' |
| NGC 4190 ULX1 | 2.83±0.10 | 0654650301 | 2010-11-25 | 11070 | 12.6±0.11 | FF | Medium | On-Axis |
| NGC 4559 X-1 | 7.31±0.20 | 0152170501 | 2003-05-27 | 33950 | 2.72±0.03 | FF | Medium | On-Axis |
| NGC 4736 ULX1 | 4.59±0.80 | 0404980101 | 2006-11-27 | 35040 | 1.89±0.02 | FF | Thin | On-Axis |
| NGC 4861 ULX | 7.00±1.00 | 0141150101 | 2003-06-14 | 13180 | 0.73±0.02 | FF | Medium | On-Axis |
| NGC 5204 X-1 | 4.76±0.90 | 0405690101 | 2006-11-15 | 7820 | 9.67±0.11 | FF | Medium | On-Axis |
| NGC 5907 ULX | 17.1±0.90 | 0729561301 | 2014-07-09 | 37480 | 3.30±0.03 | FF | Thin | On-Axis |
| NGC 7793 P13 | 3.58±0.70 | 0748390901 | 2014-12-10 | 41970 | 4.85±0.04 | FF | Thin | ~4.0' |
| <i>NuSTAR</i> | | | | | | | | |
| Ho II X-1 | 3.27±0.60 | 30001031005 | 2013-09-17 | 111104 | 0.38±0.01 | – | None | ~3.1' |
| Ho IX X-1 | 3.77±0.80 | 30002033003 | 2012-10-26 | 88030 | 1.29±0.13 | – | None | ~1.6' |
| IC 342 X-1 | 2.73±0.70 | 90201039002 | 2016-10-16 | 49173 | 1.26±0.02 | – | None | On-Axis |
| NGC 1313 X-1 | 4.25±0.80 | 30002035002 | 2012-12-16 | 100864 | 6.19±0.03 | – | None | ~2.8' |
| NGC 1313 X-2 | 4.25±0.80 | 30002035002 | 2012-12-16 | 100864 | 6.19±0.03 | – | None | On-Axis |
| NGC 5907 ULX | 17.1±0.90 | 80001042002 | 2014-07-09 | 57113 | 0.26±0.01 | – | None | On-Axis |

Notes. ^(a) All distance estimations are from [Tully et al. \(2013\)](#) ^(b) Of filtered pn observation. ^(c) Corresponding to the pn spectrum of the source.

for a high viewing angle (e.g. dips⁶ in their light curves or spectral absorption features resulting from an edge-on view of the accretion disk atmosphere) in any of our sources, we have selected a value of $i = 60$ deg for all sources in our list.

All best fit values for the absorbed dual-MCD model together with the estimations for the magnetic field and their classification, as proposed by [Sutton et al. \(2013\)](#), are presented in Table 2. The values for the absorbed MCD/black-body model are presented in Table 4. In Table 4 we also provide the estimations of the “spherization” radius ([Shakura & Sunyaev 1973](#)) of each source. Lastly, we note that the χ^2 values from the `tbnew_gas*(diskbb+bbbody)` fits were similar, albeit moderately higher than those of the dual MCD model and with moderately lower temperature of the hot component (kT_{BB} between ~0.9 keV and ~2.2 keV).

3.3.2. *NuSTAR*

In the *NuSTAR* spectra we ignored all channels below 4 keV, and thus we did not require the addition of the cool thermal component. The primary spectral component used to model all spectra is again a multi color disk black body (*diskbb*), the “hot” thermal component from the *XMM-Newton* fits. Furthermore, we

also look for the presence of a potential hard, non-thermal tail, which is usually detected in most XRBs, even in the soft state. A simultaneous, broadband fit of the combined *NuSTAR* plus *XMM-Newton* spectra is not explored in this paper. Recent, rigorous works have extensively studied the *XMM-Newton* (or *Swift*) + *NuSTAR* data that we revisit here (Ho II X-1: [Walton et al. 2015](#); HoIX X-1: [Luangtip et al. 2016](#); IC 342 X-1: [Rana et al. 2015](#); NGC 1313 X-1,X-2: [Bachetti et al. 2013](#); NGC 5907 X-1: [Fürst et al. 2017](#)) and have noted the presence of a spectral shape that can be modeled either as a hot thermal component or sharp cutoff with an additional, weak power-law tail. In this paper we do not seek to reproduce these analyses, but to discuss a possible novel interpretation of the spectral shape. The *NuSTAR* data are used with the purpose of confirming (or dismissing) the presence of these components in a comprehensive and consistent study. To this end, the separate analysis is swift and effective.

The *NuSTAR* data were modeled using a single *diskbb* model and an additional power law. The hard (>10 keV), power-law emission is faint, with less than 5% of observed photons registered above 20 keV, on average. Furthermore, the background contamination becomes predominant above 25 keV. Therefore, the slope or even the exact shape (i.e. the presence of an exponential cutoff) of the hard spectral tail cannot be constrained accurately. Both the thermal component and the power-law tail are required, in order to achieve an acceptable fit. More specifically, fitting the *NuSTAR* data with only the *diskbb* model results in pronounced residuals above ~15 keV (e.g. Figure 1) and a value for the reduced χ^2 that exceeds ~1.3 in all sources. Similarly, fit-

⁶ With the exception of NGC 55 ULX, which does show dips in its light curve ([Stobbart et al. 2004](#)). However, due to its likely supercritical accretion rates, the dips are not as constraining, for its inclination, as in typical XRBs.

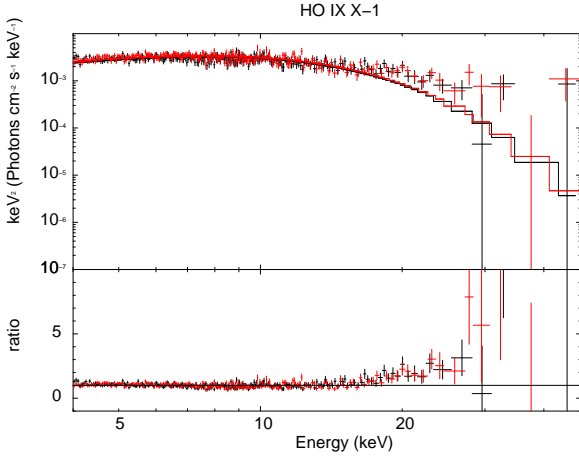


Fig. 1: Ho IX X-1: Unfolded spectrum, energy and data-vs-model ratio plot, for only the diskbb model. There are clear residuals above 20 keV.

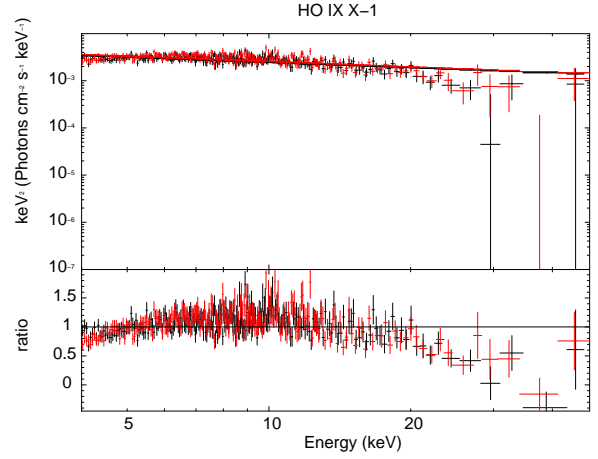


Fig. 2: Ho IX X-1: Unfolded spectrum, energy and data-vs-model ratio plot, for only the power law model. There is a clear curvature in the spectrum that cannot be described by the power law.

ting the *NuSTAR* spectra with just a power-law, results in residual structure characteristic of thermal emission (Figure 2) and reduced χ^2 values exceeding ~ 1.2 .

The temperature gradient of the accretion curtain will, most likely, differ from the $T \sim r^{-0.75}$ predicted by the standard thin disk MCD models like `diskbb` and this deviation will be enhanced further by electron scatterings. To diagnose the impact of this effect on the registered spectra, we also fitted them with the `xspec` model `diskpbb`, in which the disk temperature is proportional to r^{-p} and p is a free parameter. We find that the value of p is significantly smaller (on average $p \lesssim 0.53$) than that of a standard thin accretion disk. More interestingly, we find that the `diskpbb` fits did not require the addition of a power-law tail and yielded the same χ^2 values as the `diskbb` + `powerlaw` fits. Albeit with the notable exception of NGC 5907, which is the only pulsating ULX in our *NuSTAR* sample. In principle the `diskpbb` model could also be used to model the cool thermal emission, detected in the *XMM-Newton* data, since the inner disk parts may also become inflated due to the high accretion rates (see discussion in Section 4). Nevertheless, the addition of an extra free parameter in each thermal component will only add to the degeneracy between different models and will not provide any further insight to the physical parameters (i.e. temperature and size) of the emitting regions. Therefore, the `diskpbb` model is only used as a diagnostic for the geometry of the accretion curtain, and only for the *NuSTAR* data were its impact is much more significant. Its implications are discussed further in Section 4. While we have analyzed all available *NuSTAR* data for our sources, we only tabulate the results for those observations with the largest number of counts and for luminosities closer to the *XMM-Newton* observations. Nevertheless all *NuSTAR* observations produced – more or less – similar results to the ones presented here (see observation log in Table 1, for the *NuSTAR* observations analyzed in this work). Best fit values (including the value of p and the temperature of the `diskpbb` models) for the *NuSTAR* data are presented in Table 3.

4. Discussion

The use of a double thermal spectrum, with similar temperatures to those observed in the dual thermal spectra of soft state NS-XRBs, successfully describes the spectra of ULXs in our list and the unusual high energy roll-over of the ULX spectra can

be re-interpreted as the Wien tail of a hot (multicolor) black body component. The similarities between the spectral morphology of ULXs and those of NS-XRBs in the soft state are illustrated in Figure 3. We have plotted the *XMM-Newton* spectra of two known NS-XRBs (4U 1916-05 and 4U 1705-44: see appendix A.1) along with the spectra of two (non pulsating) ULXs from our sample. Dotted lines correspond to the dual thermal model (in this example it is an absorbed MCD plus black body model) – which in the case of the two NS-XRBs – is used to model emission from the boundary layer and the thin accretion disk. The same configuration is used to model the spectra of the two ULXs (in this example NGC 4559 X-1 and M81 X-6). M81 X-6 is in the BD state and NGC 4559 X-1 in the SUL state. We stress that the unfolded spectra presented in Figure 3, are modeled dependent. They are used here in order to illustrate the apparent similarities between the spectral shapes of ULXs and soft-state NS-XRBs and not to extract any quantitative information on the spectral parameters (see also, similar example plot in Sutton et al. 2013). The suitability of a double thermal model, for the spectra of ULXs, had been noted by Stobbart et al. (2006), but the model was dismissed, as it was difficult to explain the presence of a secondary thermal component in terms of an accreting black hole. To probe beyond this superficial similarity, we explore the parameter space of the different spectral fits with respect to theoretical expectations, and discuss our findings and their implications below.

More specifically, to investigate the case for (super-Eddington) accretion onto lowly magnetized NSs (e.g. King & Lasota 2016), we applied the `diskbb` plus `bbody` model that is often used to model NS-XRBs in the soft state. Indeed this model describes well the spectra of the sources in our list. However, the radius inferred from the hot black body fit, has a size that is approximately an order of magnitude larger⁷ than the size of the spreading layer on the surface of the NS (see Table 4). This is not surprising, since at such high accretion rates – and for a low magnetic field NS (as considered in the King & Lasota 2016 model) – the accretion disk will extend well beyond the “spherization” radius (R_{sph} : Shakura & Sunyaev 1973). The flow will be strongly super-Eddington and the material will, most

⁷ The boundary layer is expected to have a size of a few km (Lin et al. 2009 and Table A.1 in this work.)

Table 2: Best fit parameters of the dual MCD model for the *XMM-Newton* observations. All errors are in the 1σ confidence range.

| Source | nH | kT _{disk} | R _{disk} ^a | kT _{hot} | L/L _{edd} ^b | L _{hot} /L _{disk} | B | Clas. ^c | red. χ^2/dof |
|---------------|--|--|---------------------------------------|--|--|-------------------------------------|--|--------------------|-------------------|
| | [$\times 10^{21}$ cm ²] | keV | [km] | keV | | | 10^{12} G | | |
| Ho II X-1 | 0.40 ^d | 0.42 ^{+0.02} _{-0.01} | 1421 ⁺¹²⁸ ₋₁₁₂ | 1.65 ^{+0.06} _{-0.05} | 69.8 ^{+21.3} _{-17.7} | 0.74 \pm 0.03 | 22.7 ^{+3.27} _{-2.98} | SUL | 1.22/137 |
| Ho IX X-1 | 0.64 ^{+0.15} _{-0.13} | 0.48 ^{+0.05} _{-0.04} | 557 ⁺¹⁵⁴ ₋₁₀₈ | 3.15 ^{+0.10} _{-0.09} | 78.5 ^{+27.3} _{-22.2} | 5.62 \pm 0.32 | 4.64 ^{+2.08} _{-1.42} | HUL | 1.03/156 |
| IC 342 X-1 | 7.89 ^{+0.60} _{-0.53} | 0.48 ^{+0.06} _{-0.05} | 431 ⁺¹³² _{-97.8} | 2.84 ^{+0.10} _{-0.09} | 21.8 ^{+9.70} _{-7.45} | 2.38 \pm 0.17 | 1.56 ^{+0.98} _{-0.55} | HUL | 1.02/102 |
| M33 X-8 | 0.95 \pm 0.01 | 0.50 ^{+0.08} _{-0.07} | 237 ^{+70.4} _{-47.8} | 1.35 ^{+0.06} _{-0.06} | 11.1 ^{+8.93} _{-7.73} | 4.22 \pm 1.11 | 0.42 ^{+0.36} _{-0.13} | BD | 0.99/121 |
| M81 X-6 | 1.63 ^{+0.13} _{-0.15} | 0.93 ^{+0.23} _{-0.30} | 145 ^{+18.6} _{-19.2} | 1.72 ^{+0.16} _{-0.27} | 24.8 ^{+6.55} _{-5.74} | 1.27 \pm 0.81 | 0.27 ^{+0.14} _{-0.06} | BD | 1.13/124 |
| M83 ULX | 0.11 ^{+0.93} _{-0.10} | 0.37 ^{+0.11} _{-0.12} | 411 ^{+62.3} _{-76.1} | 1.23 ^{+0.16} _{-0.13} | 6.25 ^{+1.61} _{-1.39} | 3.25 \pm 2.18 | 0.80 ^{+0.36} _{-0.37} | – | 0.92/44 |
| NGC 55 ULX | 2.40 ^{+0.39} _{-0.60} | 0.28 \pm 0.02 | 1163 ⁺¹⁰⁸ ₋₁₆₁ | 0.82 \pm 0.02 | 10.1 ^{+3.22} _{-3.22} | 1.12 \pm 0.10 | 6.06 ^{+3.04} _{-2.81} | SUL | 1.08/89 |
| NGC 253 ULX2 | 3.59 ^{+0.59} _{-0.60} | 0.24 ^{+0.05} _{-0.03} | 1263 ⁺⁶⁶¹ ₋₅₉₄ | 1.64 \pm 0.04 | 18.5 ^{+7.56} _{-6.08} | 8.99 \pm 3.72 | 9.89 ^{+8.34} _{-6.64} | BD | 0.99/120 |
| NGC 253 XMM2 | 1.17 ^{+0.19} _{-0.17} | 0.53 ^{+0.08} _{-0.07} | 239 ^{+63.4} _{-45.8} | 1.54 ^{+0.14} _{-0.10} | 10.2 ^{+4.40} _{-3.51} | 2.60 \pm 0.64 | 0.41 ^{+0.21} _{-0.13} | BD | 1.13/108 |
| NGC 1313 X-1 | 2.42 ^{+0.24} _{-0.23} | 0.32 \pm 0.02 | 1612 ⁺²⁸² ₋₂₃₆ | 2.33 ^{+0.11} _{-0.10} | 52.3 ^{+16.1} _{-13.3} | 2.41 \pm 0.17 | 24.3 ^{+7.91} _{-5.91} | SUL | 0.88/113 |
| NGC 1313 X-2 | 1.83 ^{+0.11} _{-0.10} | 0.65 ^{+0.07} _{-0.08} | 313 ^{+52.3} _{-41.4} | 2.25 ^{+0.16} _{-0.12} | 50.2 ^{+17.5} _{-13.3} | 3.23 \pm 0.81 | 1.44 ^{+0.45} _{-0.32} | BD | 0.95/144 |
| NGC 4190 ULX1 | 0.99 ^{+0.58} _{-0.47} | 0.50 ^{+0.22} _{-0.12} | 352 ⁺²⁹⁹ ₋₁₆₂ | 1.95 ^{+0.20} _{-0.10} | 36.2 ^{+2.58} _{-2.16} | 6.88 \pm 3.02 | 1.52 ^{+2.78} _{-0.96} | BD | 1.00/112 |
| NGC 4559 X-1 | 0.80 ^{+0.33} _{-0.31} | 0.26 ^{+0.03} _{-0.02} | 2609 ⁺⁸⁰² ₋₆₀₄ | 1.41 ^{+0.08} _{-0.07} | 42.5 ^{+1.59} _{-1.63} | 1.99 \pm 0.73 | 49.1 ^{+29.3} _{-18.2} | SUL | 1.05/84 |
| NGC 4736 ULX1 | 3.54 ^{+1.03} _{-1.62} | 0.40 ^{+0.08} _{-0.09} | 513 ⁺³⁷³ ₋₁₇₈ | 1.18 ^{+0.13} _{-0.08} | 19.5 ^{+5.30} _{-4.46} | 2.34 \pm 0.51 | 1.96 ^{+1.23} _{-0.73} | BD | 1.09/68 |
| NGC 4861 ULX | 1.24 ^{+1.07} _{-0.90} | 0.25 ^{+0.09} _{-0.05} | 1639 ⁺⁵¹² ₋₃₆₀ | 1.13 ^{+0.44} _{-0.20} | 12.6 ^{+2.51} _{-2.13} | 1.52 \pm 0.47 | 11.4 ^{+6.96} _{-4.00} | – | 1.09/22 |
| NGC 5204 X-1 | 0.18 ^{+0.41} _{-0.18} | 0.36 ^{+0.04} _{-0.05} | 1428 ⁺⁵⁶⁷ ₋₃₇₂ | 1.41 ^{+0.08} _{-0.12} | 48.8 ^{+13.1} _{-12.5} | 1.22 \pm 0.16 | 19.0 ^{+11.6} _{-6.85} | SUL | 1.09/64 |
| NGC 5907 ULX | 0.75 ^{+0.10} _{-0.09} | 0.40 ^{+0.08} _{-0.06} | 2497 ⁺²⁵⁷ ₋₃₂₄ | 2.53 ^{+0.11} _{-0.13} | 483 ^{+39.7} _{-40.9} | 5.51 \pm 0.66 | 184 ^{+11.3} _{-12.9} | – | 1.09/83 |
| NGC 7793 P13 | 1.25 ^{+0.27} _{-0.26} | 0.33 ^{+0.04} _{-0.03} | 671 ⁺²⁰³ ₋₁₆₀ | 3.27 ^{+0.13} _{-0.12} | 34.6 ^{+10.6} _{-8.78} | 11.0 \pm 0.35 | 4.17 ^{+2.46} _{-1.58} | – | 1.08/141 |

Notes. ^(a) Radius inferred from cool diskbb component by solving $K=(R_{\text{disk}}/D_{10\text{kpc}})^2 \cos i$, for R_{disk} (the inner radius of the disk in km). K is the normalization of the cool diskbb model, $D_{10\text{kpc}}$ is distance in units of 10 kpc and i is the inclination.

^(b) “Bolometric” luminosity (L) in the 0.1–100 keV range, extrapolated from the best-fit model. L_{edd} is the Eddington luminosity for an isotropically accreting NS with a mass of $1.4 M_{\odot}$.

^(c) From [Sutton et al. \(2013\)](#). BD: broadened disc state, HUL: hard ultraluminous state, SUL: soft ultraluminous (SUL) state.

^(d) Parameter frozen at total galactic HI column density ([Dickey & Lockman 1990](#)).

Table 3: Best fit parameters for the *NuSTAR* observations. All errors are in the 1σ confidence range.

| Source | kT _{hot} | K _{hot} ^a | Γ | K _{po} ^b | L/L _{edd} ^c | red. χ^2/dof | p | kT _{hot} | red. χ^2/dof |
|--------------|--|--|--|--|--|-------------------|---|--|-------------------|
| | keV | [10^{-3}] | | [10^{-3}] | | diskbb+po | | keV | diskbb |
| Ho II X-1 | 2.11 ^{+0.17} _{-0.11} | 5.09 ^{+2.21} _{-1.91} | 2.41 ^{+0.21} _{-0.26} | 0.98 ^{+0.07} _{-0.05} | 52.1 ^{+8.85} _{-7.36} | 0.98/227 | \downarrow 0.507 | 3.31 ^{+0.10} _{-0.09} | 1.01/228 |
| Ho IX X-1 | 3.19 ^{+0.14} _{-0.13} | 2.74 ^{+0.36} _{-0.46} | 2.38 ^{+0.12} _{-0.11} | 2.12 ^{+0.07} _{-0.05} | 86.9 ^{+40.8} _{-33.0} | 1.01/440 | 0.542 \pm 0.009 | 4.33 ^{+0.21} _{-0.19} | 1.02/441 |
| IC 342 X-1 | 2.37 ^{+0.14} _{-0.11} | 9.81 ^{+3.19} _{-2.51} | 2.58 ^{+0.16} _{-0.18} | 3.58 ^{+1.60} _{-1.33} | 39.1 ^{+18.9} _{-13.3} | 1.00/296 | \downarrow 0.508 | 3.41 ^{+0.07} _{-0.10} | 1.04/297 |
| NGC 1313 X-1 | 2.64 ^{+0.44} _{-0.37} | 2.41 ^{+1.61} _{-0.71} | 2.80 ^{+0.29} _{-0.23} | 1.58 ^{+0.29} _{-0.23} | 36.0 ^{+14.8} _{-12.3} | 1.02/122 | \downarrow 0.505 | 3.52 ^{+0.14} _{-0.13} | 1.08/123 |
| NGC 1313 X-2 | 1.94 ^{+0.19} _{-0.25} | 1.46 ^{+0.63} _{-0.43} | 4.26 ^{+0.49} _{-0.22} | 1.18 ^{+0.34} _{-0.31} | 14.7 ^{+6.08} _{-5.00} | 0.98/61 | \downarrow 0.522 | 1.77 ^{+0.08} _{-0.07} | 0.98/62 |
| NGC 5907 ULX | 2.51 ^{+0.14} _{-0.12} | 2.41 ^{+0.56} _{-0.66} | 1.68 ^{+0.37} _{-0.41} | 0.05 ^{+0.04} _{-0.06} | 457 ^{+49.6} _{-46.5} | 1.02/107 | 0.579 ^{+0.06} _{-0.04} | 3.39 ^{+0.08} _{-0.07} | 1.25/108 |

Notes. ^(a) Where K_{hot} is the normalization parameter for the diskbb component. Namely $K_{\text{hot}}=(R_{\text{hot}}/D_{10\text{kpc}})^2 \cos i$, where R_{hot} is the inner radius of the disk in km, $D_{10\text{kpc}}$ is the distance in units of 10 kpc and i is the inclination.

^(b) Power-law component normalization constant: photons/keV/cm²/s at 1 keV

^(c) Luminosity in the 3.–78 keV range, extrapolated from the best-fit model. L_{edd} is the Eddington luminosity for an isotropically accreting NS with a mass of $1.4 M_{\odot}$

likely, be ejected away from the surface of NS (e.g. [King et al. 2017](#)). In this case the hot thermal component may be the result of emission of the inner disk layers, exposed by the strong outflows. In this case the hot thermal component would correspond to the stripped, inner accretion disk and the soft thermal component, as proposed in the optically thick wind scenario. Indeed, the best fit values for the size of soft thermal component are in agreement with the R_{sph} for most of the sources in our list, allud-

ing to the exciting possibility of accreting NSs powering a large fraction of ULXs. However, it is surprising that the dual thermal spectrum would be almost indistinguishable between the BH-ULX and the NS-ULXs, given that in this framework the maximum temperature of the accretion disk should exceed ~ 4 keV in the case of the NSs (perhaps even higher given the very high accretion rates of particular sources). On the other hand, there is yet no strictly defined mechanism to account for the hot black

Table 4: Best fit parameters of the MCD plus black body model for the *XMM-Newton* observations. All errors are in the 1σ confidence range.

| Source | nH | kT _{in} | R _{in} | kT _{bb} | R _{bb} | R _{sph} ^a | red. χ^2/dof |
|---------------|--|--|---------------------------------------|--|--|--|-------------------|
| | [$\times 10^{21}$ cm ²] | K | [km] | K | km | km | |
| Ho II X-1 | 0.40 ^b | 0.57 ^{+0.02} _{-0.01} | 724 ^{+76.5} _{-62.1} | 1.26 ^{+0.17} _{-0.18} | 93.5 ^{+28.3} _{-22.7} | 974 ⁺²⁹⁷ ₋₂₄₇ | 1.25/137 |
| Ho IX X-1 | 0.33 ^{+0.10} _{-0.08} | 0.91 ^{+0.04} _{-0.03} | 198 ^{+38.3} _{-42.2} | 1.95 ^{+0.27} _{-0.25} | 59.2 ^{+18.8} _{-17.7} | 1120 ⁺³⁸⁴ ₋₃₁₀ | 1.06/156 |
| IC 342 X-1 | 7.42 ^{+0.44} _{-0.41} | 0.61 \pm 0.05 | 302 ^{+61.1} _{-52.7} | 1.66 ^{+0.14} _{-0.16} | 72.8 ^{+23.0} _{-19.8} | 304 ⁺¹³⁵ ₋₁₂₃ | 1.05/101 |
| M33 X-8 | 0.82 ^{+0.03} _{-0.02} | 0.70 ^{+0.06} _{-0.07} | 202 ^{+31.8} _{-41.0} | 1.11 ^{+0.10} _{-0.18} | 62.1 ^{+30.8} _{-44.5} | 155 ⁺¹⁹³ ₋₁₀₈ | 1.01/121 |
| M81 X-6 | 1.61 ^{+0.16} _{-0.12} | 1.03 ^{+0.33} _{-0.42} | 155 ^{+20.8} _{-22.2} | 1.51 ^{+0.46} _{-0.35} | 38.2 ^{+8.12} _{-9.78} | 346 ^{+91.4} _{-89.1} | 1.13/124 |
| M83 ULX | 0.39 ^b | 0.38 ^{+0.16} _{-0.08} | 482 ^{+86.8} _{-91.8} | 0.83 ^{+0.09} _{-0.11} | 88.3 ^{+20.0} _{-16.3} | 87.2 ^{+22.5} _{-19.4} | 0.95/44 |
| NGC 55 ULX | 2.16 ^{+0.58} _{-0.51} | 0.33 ^{+0.03} _{-0.02} | 986 ⁺¹⁰⁸ ₋₂₀₁ | 0.65 ^{+0.16} _{-0.15} | 128 ^{+18.3} _{-17.7} | 153 ^{+32.1} _{-29.8} | 1.11/89 |
| NGC 253 ULX2 | 2.35 ^{+0.58} _{-0.51} | 0.95 ^{+0.14} _{-0.12} | 121 ⁺³⁰² ₋₂₅₇ | 1.39 ^{+0.16} _{-0.15} | 51.8 ^{+14.3} _{-15.7} | 258 ⁺¹⁰⁵ _{-84.8} | 1.05/120 |
| NGC 253 XMM2 | 1.09 ^{+0.19} _{-0.21} | 0.65 ^{+0.09} _{-0.10} | 218 ^{+64.3} _{-47.6} | 1.17 ^{+0.11} _{-0.09} | 54.2 ^{+22.0} _{-16.6} | 142 ^{+61.4} _{-49.0} | 1.14/108 |
| NGC 1313 X-1 | 1.83 ^{+0.18} _{-0.17} | 0.44 ^{+0.21} _{-0.20} | 778 ⁺¹⁵⁶ ₋₁₃₃ | 1.39 ^{+0.34} _{-0.43} | 52.8 ^{+22.0} _{-17.7} | 730 ⁺²²⁵ ₋₁₈₆ | 1.09/113 |
| NGC 1313 X-2 | 1.70 \pm 0.11 | 0.86 ^{+0.11} _{-0.08} | 261 ^{+43.2} _{-35.6} | 1.66 ^{+0.05} _{-0.05} | 54.6 ^{+15.8} _{-15.8} | 700 ⁺²⁰⁴ ₋₂₀₄ | 1.01/144 |
| NGC 4190 ULX1 | 0.56 ^{+0.32} _{-0.31} | 0.87 ^{+0.42} _{-0.38} | 256 ⁺²⁰⁸ ₋₁₂₀ | 1.51 ^{+0.17} _{-0.19} | 85.0 ^{+10.2} _{-9.06} | 505 ^{+31.8} _{-30.1} | 1.06/112 |
| NGC 4559 X-1 | 0.15 ^b | 0.43 ^{+0.04} _{-0.02} | 1133 ⁺³²⁵ ₋₄₁₂ | 1.04 ^{+0.06} _{-0.10} | 169 ^{+21.5} _{-18.1} | 626 ^{+77.9} _{-72.7} | 1.07/84 |
| NGC 4736 ULX1 | 0.86 ^{+0.33} _{-0.31} | 0.61 ^{+0.32} _{-0.18} | 269 ⁺²⁰⁹ ₋₁₂₂ | 0.99 \pm 0.06 | 76.6 ^{+11.3} _{-10.7} | 272 ^{+73.9} _{-62.2} | 1.12/68 |
| NGC 4861 ULX | 1.15 ^{+0.92} _{-0.81} | 0.33 ^{+0.11} _{-0.08} | 2398 ⁺⁷²⁸ ₋₅₈₅ | 0.93 ^{+0.31} _{-0.18} | 205 ^{+44.7} _{-29.7} | 176 ^{+35.2} _{-21.1} | 1.10/22 |
| NGC 5204 X-1 | 0.14 ^b | 0.43 ^{+0.08} _{-0.12} | 1262 ⁺³⁸⁸ ₋₂₉₅ | 0.99 ^{+0.13} _{-0.15} | 152 ^{+31.0} _{-39.1} | 681 ⁺²¹¹ ₋₁₇₄ | 1.15/65 |
| NGC 5907 ULX | 6.04 ^{+3.01} _{-2.52} | 0.74 \pm 0.11 | 608 ⁺¹³³ ₋₁₀₁ | 1.38 ^{+0.12} _{-0.09} | 224 ^{+49.3} _{-57.7} | 6740 ⁺⁵⁵⁴ ₋₅₇₁ | 1.09/83 |
| NGC 7793 P13 | 0.66 \pm 0.14 | 0.63 \pm 0.04 | 251 ⁺¹⁰⁸ ₋₁₁₅ | 1.70 \pm 0.05 | 59.7 ^{+22.0} _{-19.1} | 483 ⁺¹⁴⁸ ₋₁₂₂ | 1.10/141 |

Notes. ^(a) Spherization radius (e.g. King & Lasota 2016): $R_{\text{sph}} = 27/4\dot{m}R_g$, where $R_g = GM/c^2$ and $\dot{m} = \dot{M}/\dot{M}_{\text{Edd}}$.

^(b) Parameter frozen at total galactic HI column density (Dickey & Lockman 1990).

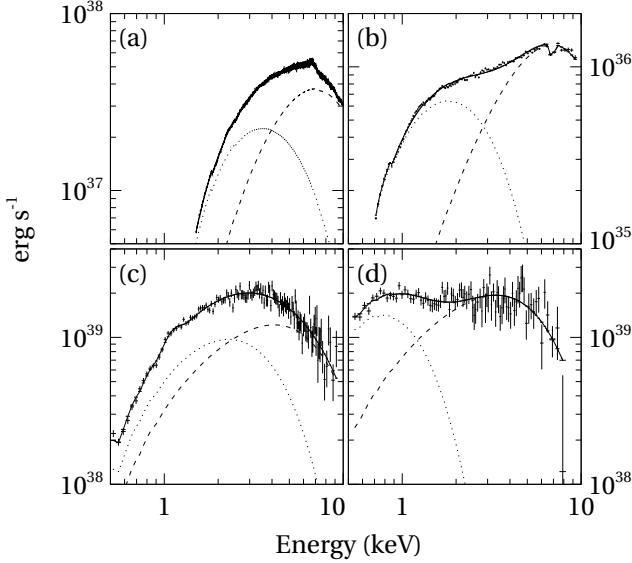


Fig. 3: Example, unfolded spectra of two ULXs from our sample and two well known NS-XRBs in the soft state. (a) 4U 1705-44, during a soft state. (b) Double thermal spectrum from NS-LMXB 4U 1916-05. (c) Apparent, dual thermal emission from ULX M81 X-6, at similar temperatures (see table 2) as 4U 1705-44. (d) Similarly shaped spectrum from NGC 4559 X-1.

body emission for the super-Eddington regime in accreting NS, and a more precise treatment may be able to resolve this apparent discrepancy. The fact still remains, that the homogeneous fit parameters hint at the possibility of most (if not all) sources in our sample belonging to one uniform population. It is cer-

tainly plausible that this is a population of NS-XRBs instead of BH-XRBs. This implication, becomes even more intriguing when we consider the fact that two of our sources (the pulsators NGC 5907 ULX and NGC 7793 P13) are almost certainly powered by highly magnetized NSs, but – unlike sub-Eddington NS-XRBs – the spectra of pulsating and non-pulsating ULXs are remarkably similar.

The possibility of highly magnetized NSs powering more ULXs in our list becomes even more relevant when we consider that the most reliable and thoroughly established mechanism for sustained super-Eddington accretion episodes, is the funneling of material onto the magnetic poles of high-B NSs. Indeed, in a recent publication Pintore et al. (2017), indicate that since the hard emission from many ULXs can be described by a combination of a hard power law and an exponential cutoff – as expected for the emission of the accretion funnel – this could be considered as an indication in favor of highly magnetized NSs powering a significant fraction of ULXs. However, this claim is problematic, since – in the presence of a high magnetic field – the photons are expected to be concentrated in a narrow beam, most likely, following the fan-beam emission diagram. Therefore, as the NS rotates, the emission should be registered in the form of characteristic pulsations.

More importantly, the shape of their pulse profile has a complex shape comprised of two or more characteristic sharp peaks (e.g. Nagel 1981; White et al. 1983; Mészáros 1992; Paul et al. 1996, 1997; Rea et al. 2004; Vasilopoulos et al. 2013, 2014; Koliopanos & Gilfanov 2016). This picture is further complicated by the fact that – most likely – a fraction of the fan-beam emission is scattered by fast electrons at the edge of accretion column and subsequently beamed towards the surface of the NS (Kaminker et al. 1976; Lyubarskii & Syunyaev 1988; Poutanen et al. 2013), off of which it is reflected, resulting in

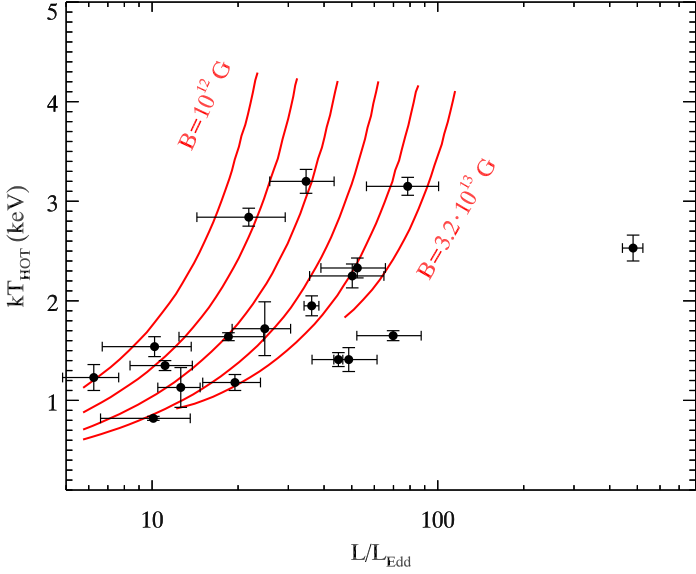


Fig. 4: Unabsorbed luminosity (in the 0.5-10 keV range) versus the temperature (in keV) of the hot multicolor disk component for the *XMM-Newton* data (Table 2). The (red) solid curves correspond to internal temperature (T_{in}) of the accretion curtain vs total luminosity, as predicted by Mushtukov et al. (2017) (see their Fig. 3). Different curves correspond to different magnetic field strength. From left to right it is 10^{12} , 2×10^{12} , 4×10^{12} , 8×10^{12} , 1.6×10^{13} and 3.2×10^{13} G.

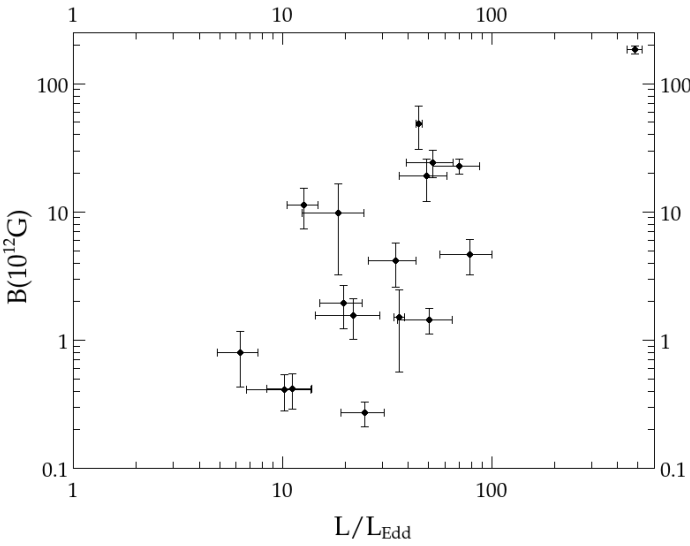


Fig. 5: Magnetic field strength vs. Luminosity for the dual MCD model. All values are taken from Table 2 (columns 6 and 8).

a secondary “polar” beam, which further complicates the pulse shape (e.g. Trümper et al. 2013; Koliopanos & Vasilopoulos 2017). All but three known ULXs lack any evidence of pulsations and the three pulsating sources (NGC 5907, M82 X-2 and NGC 7793 P13) have very smooth and simple sinusoidal pulse profiles. Therefore, very serious doubts are cast on the interpretation of the ULX spectra as due to direct emission from the accretion column.

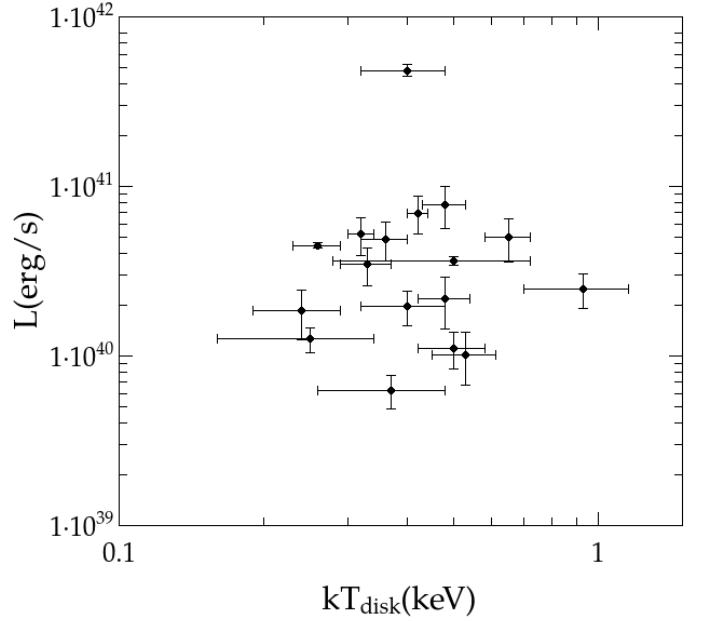


Fig. 6: Luminosity vs. disk temperature for the dual MCD model. All values are taken from Table 2 (columns 3 and 6).

This contradiction, appears to be resolved in a new publication by Mushtukov et al. (2017). In this work, the authors demonstrate that highly magnetized NSs, accreting at high mass accretion rates, can become engulfed in a closed and optically thick envelope (see their Figure 1). As the primary, beamed emission of the accretion funnel is reprocessed by the optically thick material, the original pulsation information is lost. However, if the latitudinal gradient is sufficiently pronounced – and depending on the viewing angle and inclination of the accretion curtain – the emission may be registered as smooth sinusoidal pulses (this could be the case of the three PULXs). More interestingly, the reprocessed emission is expected to have a multicolor black body (MCB) spectrum with a high temperature ($\gtrsim 1.0$ keV). The hot MCB component will be accompanied by a cooler ($\lesssim 0.5$ keV), thermal component, which originates in a truncated accretion disk. More specifically, the accretion disk is expected to extend uninterrupted, until it reaches at the $\sim R_{mag}$, where the material follows the magnetic field lines to form the optically thick curtain. In this description the characteristic double thermal spectra of NS-XRBs can coexist with high-B super-Eddington accretion, thus setting NSs as excellent candidates for powering ULXs.

In light of these findings, we remodeled the *XMM-Newton* spectra of the eighteen ULXs, using two MCD components. Indeed, the dual MCD model yields marginally better fits than the MCD/black-body fit, in all sources. The best fit values for the temperature and inner radius of the cool MCD component, indicate the presence of a strong magnetic field in all the ULXs in our sample. Namely, their values are consistent with a truncated accretion disk. If we assume that the disk is truncated close to the magnetospheric radius (i.e. to the first approximation $R_{in}=R_{mag}$), we estimate that the intensity of the magnetic field, exceeds 10^{12} G in most sources in our list. More importantly, the temperatures of the hot MCD component and the fit-derived luminosities (see Table 2) occupy the same parameter space, as predicted by Mushtukov et al. (2017: Their figure 3 and also Figure 4 in this work). Namely, the best-fit values for kT_{hot} appear to follow the theoretical curves predicted

in Mushtukov et al. (2017) and as a general trend, sources with stronger magnetic fields are more luminous (see Figure 5) and have a hotter accretion curtain. The observed correlation between the magnetic field strength and the source luminosity is in agreement with the predictions of Mushtukov et al. (2015a), where the accretion luminosity of magnetized neutron stars, in the super-critical regime, is discussed.

Following this scheme, we also place NGC 5907 ULX in the magnetar regime ($B \sim 1.8 \cdot 10^{14}$ G) which is in agreement with the findings of Israel et al. (2017). As Israel et al. also point out, such a high value of the magnetic field is puzzling⁸, since the source should be repeatedly entering the propeller regime (Illarionov & Sunyaev 1975; Stella et al. 1986). However, we must underline the fact that the magnetic field values presented in this work are estimated based on the crude assumption that the R_{mag} is equal to the truncation radius of a thin Shakura-Sunyaev disk accreting onto a bipolar magnetic field. Such as, the derived values should be treated as indications of a strongly magnetized accretor, but not considered at face value. A more realistic treatment of specific sources could yield B-field values of up to a factor of 5-6 times lower. For instance, if we re-estimate the magnetic values using the latest considerations of Chashkina et al. (2017) – where it is shown that in the radiation pressure dominated regime, the size of the magnetosphere is independent of the mass accretion rate (see their equation 61 and also equations 39 and 41) – we end up with a value of $B \sim 3.5 \cdot 10^{13}$ G for NGC 5907 ULX and a factor of $\sim 30 - 470\%$ lower magnetic field values for the other sources in our list. Nevertheless, the main outcome of our analysis remains. The best-fit parameters are consistent with our underlying assumption of a high magnetic field, which reinforces the plausibility of this scenario.

An additional implication of the high magnetic fields is the requirement that these sources are very young. Depending on the initial value of the magnetic field (which in this scenario should be at magnetar levels), the initial spin period and the mass accretion rate, these sources are most likely younger than $\sim 5 \times 10^6$ yr, if we assume that they currently have a magnetic field of the order of 10^{12} G (e.g. Ghosh & Lamb 1979; Zhang & Kojima 2006; Pan et al. 2016). Indications in favor of a relatively young age (of the order of ~ 10 Myr) can be maintained for sources HoII X-1, HoIX X-1, IC342 X-1, M81 X-6, NGC 1313 X-2, NGC 253 XMM2, NGC 253 ULX2, NGC 4559, NGC 4736 and NGC 5204, which are associated with young stellar environments and star forming regions (Soria et al. 2005; Liu et al. 2007; Grisé et al. 2008; Berghea et al. 2008, 2010; Grisé et al. 2011; Berghea et al. 2013). Furthermore, sources HoII X-1, HoIX X-1, IC342 X-1, M81 X-6, NGC 1313 X-2 and NGC 5204 have optical counterparts indicating that they are very young objects (Zampieri et al. 2004; Kaaret et al. 2004; Ramsey et al. 2006), while the nebula of IC 342 X-1 and HoIX X-1 indicate activity of less than ~ 1 Myr (Pakull & Mirioni 2002; Kaaret et al. 2004; Abolmasov et al. 2007; Feng & Kaaret 2008; Cseh et al. 2012). However, estimation of the stellar companion’s age, based on the optical counterpart, can be hindered by the fact that its emission may originate in the (irradiated) outer accretion disk, rather than the photosphere of the donor star (e.g. Grisé et al. 2012; Tao et al. 2012a,b). Furthermore, the magnetic field values inferred from the spectral fitting of some of our sources, would require even younger ages than those derived from observations – i.e. for $B \gtrsim 10^{13}$ G and assuming standard magnetic field decay (e.g. Frank et al. 2002; Zhang & Kojima

2006). Therefore, investigation for further indications of the presence of magnetic field in ULXs and more accurate estimation of the magnetic field strength is required to explore this intriguing scenario.

If the “cool” MCD component, indeed originates in a truncated accretion disk, we would expect a positive correlation between the disk temperature (T_{disk}) and the bolometric luminosity (L) as argued by Miller et al. (2013). In Figure 6 we have plotted R_{disk} vs T_{disk} , however, since the accretion disk is expected to become truncated at different values of R_{disk} (which in our scheme correspond to different B-field strength and mass accretion rates of different sources), there is a large scatter in the derived values and an accurate estimation of the $L \sim T$ relation cannot be attempted. Regarding our choice to model the cool thermal emission using a thin disk model, we must note that while in all sources analyzed in this work the R_{mag} is larger than R_{sph} (for any plausible value of f_{col}) and therefore the accretion disk could be assumed to remain thin, the R_{mag} values are only nominally larger than R_{sph} and – more importantly – the fit inferred luminosities of the disk component are super-Eddington, suggesting that the disk will, most likely, be geometrically thick. In this case we would expect that the advection will perturb the thin MCD spectrum which we have used to model the cool thermal component. Nevertheless, this will not introduce any significant qualitative difference in our results (see e.g. Straub et al. 2013 and also discussion in Mushtukov et al. 2017) and therefore – as with the hot thermal component – the diskbb model is sufficient for the purposes of this work. Another potential issue of the geometrically thick disk is the expected emergence of outflows due to the radiation pressure, which may put the stability of this mechanism into question. However, in the presence of strong magnetic fields the accreting, optically thick material will remain bound, even for luminosities of the order of $\sim 10^{40}$ erg/s (Mushtukov et al. 2017; Chashkina et al. 2017).

We also note, that the numerically estimated curves in Figure 3 of Mushtukov et al., refer to the internal temperature (T_{in} : is the temperature of the inner boundary of the emission curtain, which faces the NS) of the geometrically thick accretion envelope. In the optically thick regime, T_{in} is related to T_{out} (which corresponds to the observed T_{hot}) as,

$$T_{\text{in}} \approx T_{\text{out}} \tau^{\frac{1}{4}} \quad (6)$$

Therefore, the values of T_{hot} presented in Figure 4, should be multiplied by a factor of $\approx 1.8 - 2.1$ (corresponding to an optically thick corona, i.e. $\tau \approx 10 - 20$), in order to represent the internal temperature (T_{in}). However, as stated in section 3, in our temperature estimations we have ignored the spectral hardening which would have produced color corrected temperatures given by $T_{\text{cor}} = T_{\text{hot}}/f_{\text{col}}$, with f_{col} ranging between 1.5 and 2.1. This notable consistency between the color correction factor and the relation between internal and external temperature in an optically thick accretion envelope, further reinforces its plausibility and with it, our confidence in the observational verification of this new scheme proposed by Mushtukov et al. (2017).

These intriguing findings are also supported by the *NuSTAR* observations. Indeed, analysis of the *NuSTAR* data confirms the presence of the < 10 keV roll over, observed in the *XMM-Newton* data (see e.g. Figure 2). More importantly, when the spectral curvature is modeled as hot MCB emission the temperatures of the diskbb component, in the *NuSTAR* data, are in agreement with the *XMM-Newton* observations, particularly in those sources that were observed at similar luminosity. The presence of a thermal-like component in the *NuSTAR* spectra of ULXs has

⁸ In Israel et al. (2017) a multi-pole component is proposed as a resolution.

also been noted for sources Circinus ULX5 (Walton et al. 2013) and NGC 5204 X-1 (Mukherjee et al. 2015), further supporting the case for emission from hot, optically thick material.

Indications for the presence of optically thick material, at the boundary of the magnetosphere, can also be found in sources that lie below the Eddington limit. Several X-ray pulsars (in the sub-Eddington regime) exhibit a characteristic spectral “soft excess”, which is well described by a black body distribution at a temperature of $\sim 0.1\text{--}0.2\text{ keV}$ (e.g. Hickox et al. 2004, and references therein). This emission has been attributed to reprocessing of hard X-rays by optically thick material, in the vicinity of the magnetosphere. The size of the reprocessing region is considerably larger than the inner edge of a standard accretion disk and it is argued that it may partially cover the primary hard emission from the accretion column (e.g. Zurita Heras et al. 2006; La Palombara & Mereghetti 2006; Reig et al. 2009; Sidoli et al. 2015). It is plausible that the predictions of Mushtukov et al. (2017) are – in essence – an expansion of these arguments to the super-Eddington regime, where the optically thick material engulfs the entire magnetosphere obfuscating most (or all) of the primary hard emission. The resulting accretion envelope has a temperature that is an order of magnitude higher than that of the soft excess in the sub-Eddington sources.

The *NuSTAR* data also reveal the presence of a weak power law above $\sim 15\text{ keV}$. The hard emission, which is often present in the spectra of NS-XRBs in the soft state (e.g. Barret 2001; Done et al. 2007; Lin et al. 2009, and references therein), has also been noted by Walton et al. (2014), Walton et al. (2015) and Fürst et al. (2017) in *NuSTAR* data of Ho IX X-1, Ho II X-1 and NGC 5907 ULX, respectively. The thermal emission of the accretion curtain may be modified by IC scatterings from a photoionised atmosphere, analogous to an accretion disk corona (e.g. Sunyaev & Titarchuk 1980; Haardt & Maraschi 1993). The presence of this highly ionized plasma is also supported by the detection of emission-like features in some of the observations in our list (see Section 3). We note that the presence of similar, broad emission-like features centered at $\sim 1\text{ keV}$ have also been detected in the spectra of “nominal” X-ray pulsars at lower accretion rates (e.g. Ramsay et al. 2002; Sidoli et al. 2015; La Palombara et al. 2016), which also exhibit the soft excess.

We must also remark the possibility that the apparent power-law tail may in fact be an artifact, resulting from modeling the MCB emission of the quasi-spherical accretion curtain, with a MCD model. In Section 3, it was noted that when we model the hot thermal component with a MCB model where T_{hot} is proportional to r^{-p} and p is left to vary freely, the *NuSTAR* spectra can be successfully fitted without the requirement for the hard tail. More specifically, in all cases the value of p is less than ~ 0.58 , which – in the context of accretion disks – points to an “inflated”, advective, slim disk (e.g. Kubota & Makishima 2004, and references therein). In this case, it is fairly plausible that the radically different geometry of the accretion curtain, will cause a significant deviation from the temperature gradient of $T \sim r^{-0.75}$ assumed by the standard MCD model used in our fits, resulting in an underestimation of the hard emission, which appears as an excess above $\sim 15\text{ keV}$.

The hypothesis of an “obscured”, highly magnetized NS, as the central engine in ULXs may also resolve the contradiction regarding the only known (to this date) detection of a relativistic jet in a ULX, in Ho II X-1 (Cseh et al. 2015). While collimated jets are often detected in BH-XRBs and AGN they are only present during the low-accretion, non-thermal *hard state*. In the case of Ho II X-1, though, the collimated jet is detected in a high accretion state, during which the spectrum is dominated by

thermal emission, which is in stark contrast to the BH-XRB/jet paradigm. This contradiction is resolved when we consider a high-B NS as the accretor, which – for sufficiently high values of the magnetic field – can power collimated, relativistic jets at high accretion rates (Parfrey et al. 2017, 2016). However, in this framework the presence of the jet is also contingent upon the NS spin period. Only a limited set of parameters would yield a powerful jet in the high accretion rate regime, which may explain the lack of a jet in most known ULXs. It is also certainly plausible that the non-detection of radio jets, may be the result of lack of sensitivity, since the expected flux would most likely lie in the few μJy range or less. Given the above discussion and the bulk of theoretical expectations, strong outflows should be expected for any type of accretor (i.e. BH, highly or lowly magnetized NS). The more pertinent question, then, would be whether the soft thermal component originates in a hot optically thick wind component, close to the accretor or the truncated accretion disk. Therefore, given the recent considerations regarding the different candidates for the ULX accretors, it is important that the evolution of the $L_{\text{soft}} \sim T_{\text{in}}$ relation for specific sources is revisited.

The “universal”, power-law shaped luminosity function of ULXs and HMXBs (e.g. Gilfanov et al. 2004; Swartz et al. 2004; Mineo et al. 2012) may also be interpreted as favoring NS-powered ULXs. More specifically, the smooth shape of the HMXB luminosity function up to $\log L \sim 40.5$, strongly implies that ULXs are composed of ordinary HXMBs with stellar mass accretors. Since most HMXBs are powered by NSs (e.g. Liu et al. 2006; Belczynski & Ziolkowski 2009) and also most ULXs are found in star-forming regions (e.g. Feng & Soria 2011, and references therein) that favor the evolution of NS-HMXBs, it is reasonable to postulate that most ULXs are indeed NS-XRBs. It is also of great interest to investigate if there are fundamentally different characteristics between ULXs and sources that lie above the $\sim 10^{40}\text{ erg/s}$ break in the luminosity function (Mineo et al. 2012). Indeed the two brightest HLXs – M82 X-1 and ESO 243-49 HLX-1 – do not feature the spectral cutoff of ULXs (e.g. Dewangan et al. 2006; Farrell et al. 2009) and also appear to transition between the empirical BH-XRB accretion states (e.g. Godet et al. 2009; Feng & Kaaret 2010, although Brightman et al. 2016 recently indicated that M82 X-1, during episodes of high accretion, can be modeled as a stellar mass BH, accreting at super-Eddington rates). The differing aspects between sources above and below the luminosity break, indicate a different type of accretor between HLXs and ULXs. Within the scheme discussed in this work, this could mean that while most ULXs have NS accretors, HLXs harbor either supercritically accreting stellar-BHs or sub-Eddington accreting IMBHs. Nevertheless, given the very small sample of HLX sources, such hypotheses remain strictly in the realm of speculation.

5. Conclusion

We have presented an alternative interpretation of the X-ray spectra of eighteen well-known ULXs, which provides physically meaningful spectral parameters. More specifically, from the analysis of the *XMM-Newton* and *NuSTAR* spectra, we note that the curvature above $\sim 5\text{ keV}$ – found in the spectra of most ULXs – is consistent with the Wien tail of thermal emission in the $>1\text{ keV}$ range. Furthermore the high quality *XMM-Newton* spectra confirm the presence of a secondary cooler, thermal component. These findings are in agreement with the analysis presented in previous works (e.g. Walton et al. 2014, 2015; Luangtip et al. 2016). However, in contrast to the currently ac-

cepted paradigm, we propose that the dual thermal spectrum may be the result of accretion onto a highly magnetized NS, as predicted in recent theoretical models (Mushtukov et al. 2017), in which the hot thermal component originates in an optically thick envelope that engulfs the entire NS, at the boundary of the magnetosphere and the soft thermal component originates in an accretion disk, that becomes truncated at approximately the magnetospheric radius. We claim that this finding offers an additional and compelling argument in favor of neutron stars as more suitable candidates for powering ULXs, as has been recently suggested (King & Lasota 2016; King et al. 2017). In light of this interpretation, the ultraluminous state classification put forward by Sutton et al. (2013) can be re-interpreted in terms of different temperatures and relative flux contribution of the two thermal components, which result in the different spectral morphologies.

Nevertheless, we stress that there is considerable degeneracy between different models, that can fit the spectra equally well and so far there are no observational features – such as cyclotron lines, or transitions to the propeller regime (e.g. Pringle & Rees 1972; Lamb et al. 1973; Illarionov & Sunyaev 1975) that will conclusively favor this hypothesis over other comprehensive and equally plausible interpretations (i.e. optically thick outflows from critically accreting black holes). Furthermore, the presence of strong outflows, is also expected in the case of accreting high field NSs, which may account for the soft thermal emission, in the NS scenario as well. Given the encouraging results of this work, further examination of this scenario is warranted. To this end the fractional variability of ULXs (which is addressed by the BH super-Eddington wind model) should be reviewed in the context of the highly magnetized NS model, and the possibility of aperiodic flux variation due to the rotating accretion curtain must be explored further. Moreover, deeper broadband observations that will also allow precise phase resolved spectroscopy of the pulsating sources, as well as long term monitoring – in sources for which this is feasible – are necessary, in order to further probe this newly emerging paradigm.

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Appendix A: Spectral extraction and analysis of the NS-XRBs

XMM-Newton spectra for NS-XRBs 4U 1705-44 and 4U 1916-05 were extracted and analyzed using the same procedures as described in Section 3. Both sources were fitted with an absorbed MCD plus black body model (`xpsec model tbnew_gas(diskbb+bbbody)`), where the MCD model was used for the “cool” thermal emission of the truncated accretion disk and the black body for the “hot” thermal emission, expected to originate in the boundary layer formed on the surface of the NS. Best fit values are presented in Table A.1. The inner disk radius inferred from the normalization parameter of the `diskbb` component was estimated using the expression given in Table 2. We assumed an inclination of $i = 60^\circ$ for 4U 1705-44 and $i = 80^\circ$ for 4U 1916-05 (this is an edge-on viewed source Boirin et al. 2004). The size of the black body emitting region (column 6) was estimated using the expression given in Table 4.

Distances of 8 kpc and 9 kpc were assumed for 4U 1705-44 and 4U 1916-05 respectively.

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Table A.1: Best fit values for *XMM-Newton* observations 0085290301 and 0551270201 of sources 4U 1916-05 and 4U 1705-44, respectively

| Source | nH | kT _{in} | R _{in} | T _{BB} | R _{BB} |
|------------|-------------------------------|------------------|-------------------------|-----------------|-------------------------|
| | $\times 10^{21} \text{ cm}^2$ | keV | km | keV | km |
| 4U 1705-44 | 22.5 \pm 0.13 | 1.14 \pm 0.03 | 11.8 $^{+0.37}_{-0.38}$ | 1.76 \pm 0.01 | 5.36 $^{+0.22}_{-0.19}$ |
| 4U 1916-05 | 2.15 \pm 0.03 | 0.66 \pm 0.01 | 13.2 $^{+0.31}_{-0.12}$ | 1.75 \pm 0.01 | 1.23 $^{+0.02}_{-0.01}$ |

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