

The supermassive black hole coincident with the luminous transient ASASSN-15lh*

T. Krühler¹, M. Fraser^{2,3}, G. Leloudas^{4,5}, S. Schulze⁴, N. C. Stone⁶, S. van Velzen⁷, R. Amorin^{8,9}, J. Hjorth⁵, P. G. Jonker^{10,11}, D. A. Kann¹², S. Kim^{13,14}, H. Kuncarayakti^{15,16}, A. Mehner¹⁷, A. Nicuesa Guelbenzu¹⁸

¹ Max-Planck-Institut für extraterrestrische Physik, Gießenbachstraße, 85748 Garching, Germany

² School of Physics, O'Brien Centre for Science North, University College Dublin, Belfield, Dublin 4, Ireland.

³ Institute of Astronomy, University of Cambridge, Madingley Road, Cambridge CB3 0HA, UK

⁴ Department of Particle Physics and Astrophysics, Weizmann Institute of Science, Rehovot 7610001, Israel

⁵ Dark Cosmology Centre, Niels Bohr Institute, University of Copenhagen, Juliane Maries Vej 30, 2100 København Ø, Denmark

⁶ Columbia Astrophysics Laboratory, Columbia University, New York, NY, 10027, USA

⁷ Department of Physics & Astronomy, The Johns Hopkins University, Baltimore, MD 21218, USA

⁸ Cavendish Laboratory, University of Cambridge, 19 JJ Thomson Avenue, Cambridge CB3 0HE, United Kingdom

⁹ Kavli Institute for Cosmology, University of Cambridge, Madingley Road, CB3 0HA, United Kingdom

¹⁰ SRON, Netherlands Institute for Space Research, Sorbonnelaan 2, 3584 CA, Utrecht, The Netherlands

¹¹ Department of Astrophysics/IMAPP, Radboud University Nijmegen, P.O. Box 9010, 6500 GL Nijmegen, The Netherlands

¹² Instituto de Astrofísica de Andalucía (IAA-CSIC), Glorieta de la Astronomía s/n, E-18008, Granada, Spain

¹³ Instituto de Astrofísica, Facultad de Física, Pontificia Universidad Católica de Chile, Vicuña Mackenna 4860, 7820436 Macul, Santiago, Chile

¹⁴ Max-Planck-Institut für Astronomie, Königstuhl 17, 69117 Heidelberg, Germany

¹⁵ Finnish Centre for Astronomy with ESO (FINCA), University of Turku, Väisäläntie 20, 21500 Piikkiö, Finland

¹⁶ Tuorla Observatory, Department of Physics and Astronomy, University of Turku, Väisäläntie 20, 21500 Piikkiö, Finland

¹⁷ European Southern Observatory, Alonso de Córdova 3107, Vitacura, Casilla 19001, Santiago 19, Chile

¹⁸ Thüringer Landessternwarte Tautenburg, Sternwarte 5, 07778 Tautenburg, Germany

October 2, 2017

ABSTRACT

The progenitors of astronomical transients are linked to a specific stellar population and galactic environment, and observing their host galaxies hence constrains the physical nature of the transient itself. Here, we use imaging from the Hubble Space Telescope, and spatially-resolved, medium resolution spectroscopy from the Very Large Telescope obtained with X-Shooter and MUSE to study the host of the very luminous transient ASASSN-15lh. With these observations, we demonstrate that the dominant stellar population at the transient site is old (around 1 to 2 Gyr), without signs of recent star-formation. The ionized gas-phase emission comes from three time-invariable, narrow emission line components of collisionally-excited metal and Balmer lines. The ratios of emission lines in the Baldwin-Phillips-Terlevich diagnostic diagram indicate that the ionization source is a weak Active Galactic Nucleus with a black hole mass of $M_{\bullet} = 5^{+8}_{-3} \cdot 10^8 M_{\odot}$, derived through the M_{\bullet} - σ relation. The different components show spatial and velocity offsets on scales of 1 kpc and 500 km s⁻¹, respectively, which we explain by gas kinematics in the narrow line region. The location of the central emission-line component is coincident with that of the transient within a total positional uncertainty of 170 pc. Using the positional coincidence as well as other similarities with hosts of Tidal Disruption Events, we strengthen the argument that the transient emission observed as ASASSN-15lh is likely related to the disruption of a star around a supermassive black hole, which is most probably spinning with a Kerr parameter $a_{\bullet} > 0.4$.

Key words. stars: individual: ASASSN-15lh, galaxies: supermassive black holes

1. Introduction

One of the most remarkable astronomical transients reported in recent years was ASASSN-15lh at a redshift of $z = 0.232$, first discovered by the All-Sky Automated Survey for Supernovae (ASAS-SN; Shappee et al. 2014) on 2015-06-24 (Nicholls et al. 2015). ASASSN-15lh is characterized by an exceptional brightness at peak ($M_u \sim -23.5$ mag), relatively high black-body temperatures over a long period (between 12 000 K and 25 000 K in the first 300 days) and a relatively fast variability timescale of a few tens of days. Its radiated energy ($E_{\text{rad}} \sim 2 \times 10^{52}$ erg

within 300 days), rapid temporal evolution, a strong rebrightening at 120 days (Brown et al. 2016), and largely featureless optical spectra (Dong et al. 2016; Leloudas et al. 2016) make ASASSN-15lh hard to classify within the established scheme of transient phenomena, and it has prompted a number of theoretical models (e.g., Metzger et al. 2015; Bersten et al. 2016; Kozyreva et al. 2016; Coughlin & Armitage 2017; Sukhbold & Woosley 2016) trying to explain the remarkable observational features of ASASSN-15lh.

ASASSN-15lh was initially suggested (Dong et al. 2016; Godoy-Rivera et al. 2017) to be a hydrogen-poor super-luminous supernova (SLSN), an explosive phenomenon from the collapse of a massive star (Quimby et al. 2011). Other authors, however,

* Based on observations at ESO, Program IDs: 097.D-1054, 297.B-5035, 099.D-0115

have subsequently disagreed with this classification (Leloudas et al. 2016; Margutti et al. 2017). They argued that the observations are more consistent with those of thermal Tidal Disruption Events (TDEs, e.g., Komossa & Greiner 1999; Komossa & Bade 1999; Gezari et al. 2008; van Velzen et al. 2011; Gezari et al. 2012; Arcavi et al. 2014), the luminous emission from the accretion of stellar material after the disruption of a star by a supermassive black hole (SMBH) at the center of a galaxy (e.g., Rees 1988; Evans & Kochanek 1989). This possibility was initially considered unfeasible, primarily because the black-hole mass M_{\bullet} of the central black hole in the massive galaxy that hosted ASASSN-15lh seemed too high (Prieto et al. 2015). Above the Hills mass $M_{\text{Hills}} \sim 10^8 M_{\odot}$, the tidal radius of a solar-mass star lies within the minimum pericenter for parabolic orbits around a non-rotating black hole, and no electro-magnetic radiation should escape these kind of encounters (Hills 1975). However, the limits on the maximum black-hole mass that can produce luminous emission are relaxed for spinning black-holes (Kesden 2012), such that a TDE remained a physically viable model to produce ASASSN-15lh (Leloudas et al. 2016), possibly also explaining the relatively low volumetric rate of similar events (van Velzen 2017).

There are four fundamental shortcomings with SLSN-related models to explain the observations of ASASSN-15lh. The first problem is the absence of broad metal absorption lines that are typically observed in SLSN in their optical spectra. One of the most salient features in early SLSN spectra, for example, arises from an O II-doublet at 4100 Å and 4400 Å (Quimby et al. 2011). An absorption line at the expected wavelength of the bluer O II transition is detected in early ASASSN-15lh spectra, however, this line must be accompanied by the redder transition also (Mazzali et al. 2016), which would have clearly been detected in the available data if it were present (Leloudas et al. 2016; Margutti et al. 2017). Secondly, the UV absorption line spectrum of ASASSN-15lh is dominated by narrow high-ionization lines like N V or O VI, which is unusual for phenomena related to the collapse of massive stars. Thirdly, the temporal evolution of the transient would be unprecedented for SLSN, and its temperature and bolometric luminosity are inconsistent with previous observations of SLSN and the theoretical expectation of an expanding photosphere for these events. Finally, the galactic environment of ASASSN-15lh is in stark contrast to previous hydrogen-poor SLSN hosts, which have exclusively been strongly star-forming, and generally low-stellar mass and metallicity (e.g., Leloudas et al. 2015; Perley et al. 2016; Chen et al. 2016). However, the extreme observational characteristics of ASASSN-15lh are clearly unprecedented not only for SLSN, but also among bona-fide, optically-selected TDEs.

The physical properties of galaxies observed to host TDEs are quite different from those of SLSNe. They exhibit a curious preference (Arcavi et al. 2014; French et al. 2016; Graur et al. 2017; Law-Smith et al. 2017) for E+A galaxies (Zabludoff et al. 1996). These galaxies are dominated by the stellar light from A-type stars, and show characteristically strong Balmer absorption. They are thought to be the result of a star-burst several hundred million years in the past, possibly triggered by a galaxy merger (Zabludoff et al. 1996). In fact, spatially-resolved spectroscopy of the nearby E+A galaxy hosting the TDE ASASSN-14li (Prieto et al. 2016) reveals distinctive tidal tails in ionized gas emission, likely the result of a recent galaxy merger and ionization from an active galactic nucleus (AGN). The reason for the over-representation of E+A galaxies within the TDE host sample is an active subject of discussion, with physical processes such as stellar dynamics around binary black holes or increased stellar

densities in the galaxy center being discussed in the recent literature (e.g., Stone & van Velzen 2016; French et al. 2017).

Because of the strong differences in the physical properties of the host galaxies of SLSNe and TDEs, and the constraints that the environment can imply for the physical nature of the transient, we have obtained sensitive long-slit, medium-resolution and spatially resolved integral field unit (IFU) spectroscopy of the field of ASASSN-15lh. These observations are used to pinpoint the exact position of ASASSN-15lh within its host, reveal emission lines from ionized hydrogen and metal lines, constrain the mass of the central black-hole and probe the large-scale environment.

Throughout the paper, we adopt a flat Λ CDM cosmology with Planck parameters ($H_0 = 67.3 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_m = 0.315$, $\Omega_\Lambda = 0.685$, Planck Collaboration 2014), a Chabrier (2003) initial mass function (IMF), and report errors at the 1σ confidence level.

2. Observations and data analysis

2.1. Hubble Space Telescope imaging

The field of ASASSN-15lh was observed with Hubble Space Telescope (HST) Wide-Field Camera 3 (WFC3) under program 14346 (PI: C. Kochanek). A total of six exposures of 416 s integration time each were obtained on 2016-08-11 (433 days after peak brightness on 2015-06-05; Dong et al. 2016) in the F606W filter through Director’s Discretionary Time and made public in the HST archive. We downloaded the processed and CTE-corrected data and drizzled them onto a single frame with a pixel scale of $0''.025 \text{ px}^{-1}$ or 93 pc px^{-1} at the redshift of ASASSN-15lh. The transient is clearly detected at high significance as a bright point-source ($\text{FWHM} = 0''.07$) above the continuum emission of the galaxy (Fig. 2). Tying the WFC3 astrometry to ten sources from the Gaia DR1 catalog (Gaia Collaboration et al. 2016a,b), we measure a position of R.A. (J2000) = 22:02:15.4263, Decl. (J2000) = $-61:39:34.910$ in the astrometric reference frame defined by Gaia. The positional uncertainty is dominated by the root-mean-square difference to the astrometric tie-objects, which is 8 mas in each coordinate (30 pc co-moving). A cut-out from this image with a size of $1'$ by $1'$ and centered around the transient position is shown in the central panel of Figure 1.

The host is elongated along the North-East, South-West direction (Figure 2), and together with the lack of recent star-formation (Section 3.6) and broad-band photometric colors, it displays the typical characteristics of early-type elliptical (E4 in this case) or lenticular (S0) galaxies (e.g., Blanton & Moustakas 2009). Within hierarchical structure formation, these galaxies are often thought to be the result of mergers in particular in group environments where galaxy encounters are frequent (e.g., Bournaud et al. 2005; Bekki & Couch 2011). However, other physical processes like ram pressure stripping or quenching of spirals due to starvation could play a role as well, so that the formation of particular S0s remains the subject of active research.

2.2. X-Shooter long-slit spectroscopy

We initiated ground-based observations of ASASSN-15lh with X-Shooter (Vernet et al. 2011), a cross-dispersed long-slit spectrograph mounted at ESO’s Very Large Telescope (VLT) Unit Telescope (UT) 2. X-Shooter operates in three arms, simultaneously covering the wavelength range of 3000 Å to 25 000 Å with a resolution between 30 km s^{-1} and 60 km s^{-1} depending on the

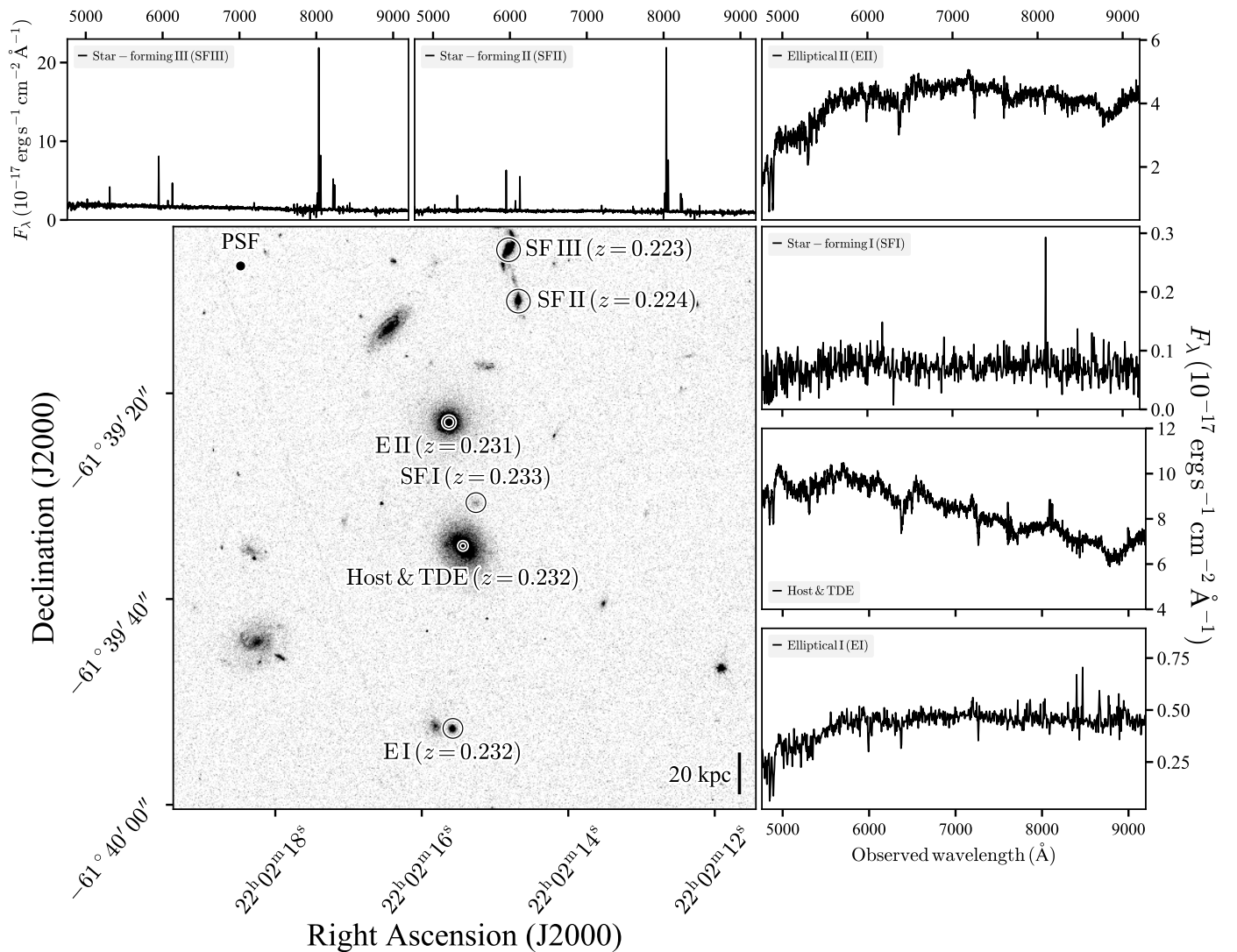


Fig. 1: HST WFC3 image in the F606W filter in the *main panel*. The image spans approximately 60'' by 60'', which is 220 kpc by 220 kpc at the redshift $z = 0.232$ of ASASSN-15lh and corresponds to the field-of-view of our MUSE integral field spectroscopy (Section 2.3). The six panels above and to the right of the WFC3 image show extracted MUSE spectra of the host (plus transient) and the five additional galaxies at a similar redshift. These galaxies are marked SFI, SFII, and SFIII for the three star-forming galaxies as well as EI and EII for two passive ellipticals in the main image. Other galaxies are either fore- or background sources. The size of the MUSE point-spread function, as well as a physical scale at $z = 0.232$ are indicated in the top left, and bottom right corner of the image respectively.

slit width and arm. In total, we obtained approximately 5700 s of integration split over two nights (2016-07-02 and 2016-08-02, which is 393 days and 424 days after peak brightness) through program 297.B-5035 (PI: M. Fraser). We used X-Shooter slit widths of 1'0 (3000 \AA to 5500 \AA), 0'9 (5500 \AA to 10 000 \AA) and 0'9 (10 000 \AA to 25 000 \AA) respectively, which were centered on the transient and oriented along the parallactic angle. Given that the transient aligns with the brightest part of the galaxy, these spectra are hence a superposition of transient and galaxy light.

The X-Shooter spectroscopy was reduced in a similar manner as described in detail in Krühler et al. (2015), making use of the ESO pipeline in its version 2.7.1 (Goldoni et al. 2006; Modigliani et al. 2010) and custom-written methods and tools. The data were flux-calibrated against the nightly spectrophotometric standard, which was LTT7987 on 2016-07-02 and EG274 on 2016-08-02, and extracted using variance-weighting. The signal-to-noise of the final spectrum is between 20 and 30

per spectral bin of size 0.4 \AA in the observed wavelength range between 3800 \AA and 9700 \AA , and somewhat lower above and below.

2.3. MUSE integral-field spectroscopy

We also used the Multi-Unit Spectroscopic Explorer (MUSE, Bacon et al. 2010) at VLT UT4 to obtain spatially-resolved spectroscopy of the field around ASASSN-15lh. MUSE is a state-of-the-art integral field unit (IFU) with an unprecedented combination of sensitivity, spatial sampling (spaxel size of 0'2 x 0'2), wavelength coverage (4750 \AA to 9350 \AA) and resolving power ($R = 1500$ to $R = 3000$ increasing from blue to red wavelengths). We obtained IFU spectroscopy for 3600 s of total integration on 2016-08-26, or 448 days after peak, under program 097.D-1054 (PI: S. Kim). The VLT/MUSE PSF of this epoch defines the effective spatial resolution and has a FWHM of 0'8 at 8000 \AA .

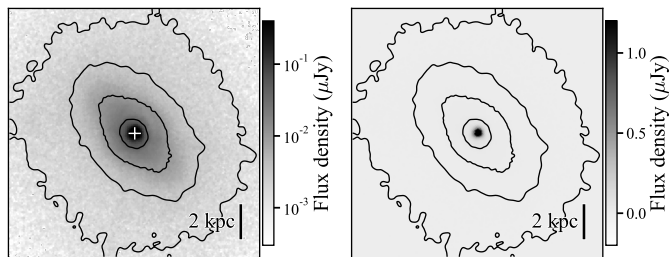


Fig. 2: Zoom in on the HST WFC3/F606W image to highlight the ASASSN-15lh host galaxy morphology (left panel, logarithmic scaling) and the transient (right panel, linear scaling). Both panels show the same image, sky area and contours, but have different color-bars. They span approximately $4''.2$ by $4''.2$, which is 16 kpc by 16 kpc at the redshift of ASASSN-15lh. North is up and East is to the left, and a physical scale at $z = 0.232$ is indicated in the bottom right corner of the image respectively. The position of the transient is marked with a white cross in the left panel.

A second epoch of MUSE spectroscopy was obtained on 2017-06-28 (754 days after peak) under ESO program 099.D-0115 (PI: T. Krühler). A total of 2800 s of exposure time on source lead to a similar depth as in the earlier epoch, but with a somewhat worse PSF with a FWHM= $1''.0$ at 8000 Å. Given the slightly better spatial resolution of the data set from 2016, most of the quantities, results and plots are derived using the earlier MUSE spectroscopy. The observations from the second epoch yield fully consistent results.

Initial data processing was performed via the MUSE pipeline¹ (version 1.6.2, Weilbacher et al. 2014), which produces a fully reduced and sky-subtracted three-dimensional data cube, calibrated in wavelength, flux and the two astrometric dimensions. Starting with this pipeline-produced data cube, we use third-party software to correct for telluric absorption (molecfit; Smette et al. 2015) and sky-line residuals (zap; Soto et al. 2016), and our own software for the analysis.

We further correct the MUSE flux-scale through synthetic photometry of the star at R.A. (J2000) =22:02:11.92, Decl. (J2000) = -61:39:46.6 (Figure 1) and comparison to its r - and i -band magnitudes from Leloudas et al. (2016). To map the accurate, HST-derived position from Section 2.1 onto the MUSE data cube, we first reconstruct several images centered at various wavelengths from the MUSE integral field spectroscopy. The MUSE field-of-view contains a handful (5-7, depending on the wavelength range used to reconstruct the MUSE images) of comparison sources which we can then use to register the MUSE astrometry against the HST imaging using a linear transformation². The registration process using different re-constructed images yields consistent results within a typical scatter smaller than 40 mas (or 0.2 MUSE pixels). This places ASASSN-15lh in our MUSE data cube with a total accuracy better than 150 pc in each coordinate.

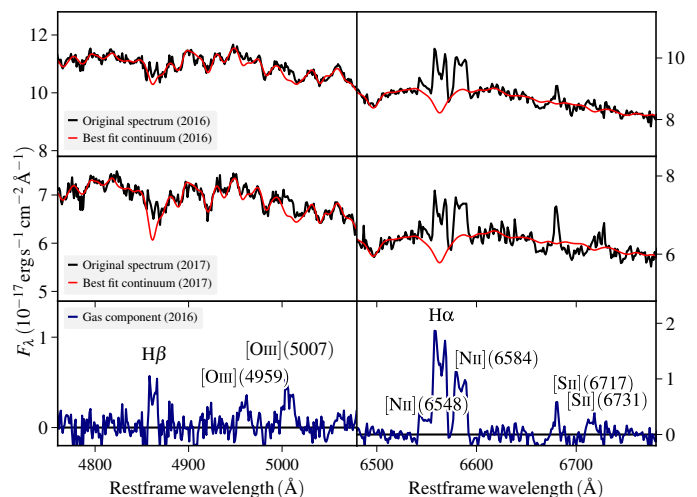


Fig. 3: Decomposition of continuum and gas emission. The upper four panels show the observed spectrum in black and the best fit continuum from pPXF in red in the wavelength range of H β and [O III] on the left, and H α and [N II] on the right side. The top row is derived from the MUSE spectra in 2016 (424 days after peak), the middle row shows the data from 2017 (754 days after peak). The lower two panels show the emission line contribution after subtraction of the continuum component for the 2016 data. A similar procedure applied to the 2017 data yields consistent results for the line fluxes and line ratios.

3. Results

3.1. Modeling the spectral continuum

Emission lines caused by the recombination of Hydrogen and decay of collisionally-excited states of metal ions provide information on the physical properties of the plasma, as well as its ionization source. In particular the ratios between Balmer and metal lines are fundamental to ascertain whether star-formation is present in the host. The Balmer lines, however, are a superposition of stellar absorption lines, emission lines from the ionized gas and a potential contribution of the transient itself.

We disentangle these different components by modeling the observed continuum with stellar templates, while emission from the transient is represented with a low-order polynomial. An illustrative example of this procedure using penalized pixel fitting (pPXF, Cappellari & Emsellem 2004; Cappellari 2017) is given in Fig. 3, where we show the central component of the host extracted from the MUSE cube to highlight already a couple of noteworthy features. Firstly, there are obvious detections of multiple emission lines, which correspond to the transitions of H β , [O III]($\lambda\lambda$ 4959, 5007), [N II]($\lambda\lambda$ 6548, 6584), H α and potentially [S II]($\lambda\lambda$ 6717, 6731), even though the significance of the latter is not particularly convincing (Figure 3) and depends somewhat on the details of the subtraction and telluric correction.

The superposition of the H α and [N II] complex is what was interpreted as a broad H α -emission in the lower-resolution spectra of Leloudas et al. (2016). The detected lines are constant between the two MUSE epochs while the transient significantly declines over the time period of one year, and we hence confirm the line emission as coming from the host (and not from the tran-

¹<http://www.eso.org/sci/software/pipelines/>

²The optical distortion of MUSE, visible as a trapezoidal shape of the field of view is already removed from the data cube through the astrometric correction applied within the MUSE pipeline.

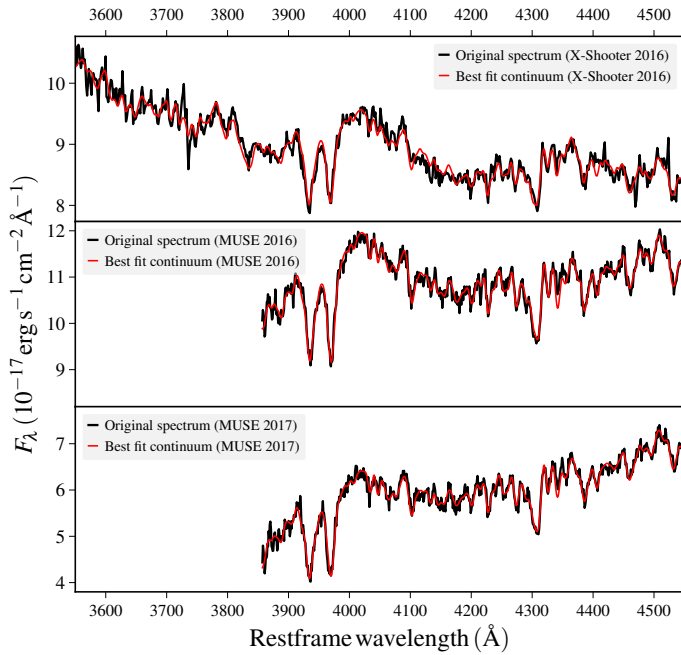


Fig. 4: Continuum model of transient plus stellar emission fit to the X-shooter and MUSE spectra for measuring the stellar kinematics in the wavelength range of the strong Ca II H+K doublet, H δ , H γ , and the G-band. The emission from the transient is clearly evident through the decrease in flux between 2016 and 2017, and a continuum that is substantially bluer than what would be expected from the Gyr-old stellar population (Sect. 3.3) of the host galaxy alone.

sient). We thus corroborate the identifications in Margutti et al. (2017) as well as their interpretation that the lines originate in the host galaxy. Secondly, each of these lines is obviously not well described with a single Gaussian line shape and shows strong velocity structure. We will return to the line shape and discuss it in detail in Sect. 3.4.

3.2. Stellar kinematics and mass of the central black hole

A similar procedure as in the previous section (Section 3.1, and Fig. 3) also returns the stellar kinematics by broadening and shifting template spectra until they match the observed data. We are in particular interested in the observed broadening of absorption lines to derive the velocity dispersion of the stellar component within the effective radius (σ_e), because it correlates well with the mass of the central black hole (Ferrarese & Merritt 2000; Gebhardt et al. 2000). Figure 4 shows medium resolution X-Shooter and MUSE spectra (instrumental resolution in this wavelength range $\sigma_{\text{inst}} \sim 25 \text{ km s}^{-1}$ for X-shooter and $\sigma_{\text{inst}} \sim 80 \text{ km s}^{-1}$ for MUSE) in the wavelength range of the strong Ca II H+K doublet, H δ , H γ , and the G-band, as well as the best fit continuum. The shown fits (Fig. 4) result in a luminosity-weighted line-of-sight velocity dispersion of $\sigma = 225 \pm 15 \text{ km s}^{-1}$, and fits to the data excluding the Balmer lines yield comparable values within errors. The effective radius R_e of the galaxy is $0''.4$, and the measured luminosity-weighted velocity dispersion within the X-shooter slit of $1''.0$ or $2.5 \times R_e$ can thus be considered as a fair proxy of σ_e (Gebhardt et al. 2000). The spatially resolved data from the MUSE data cube (Figs. 4 and 5) also shows good agreement between measurements within R_e and the X-shooter value.

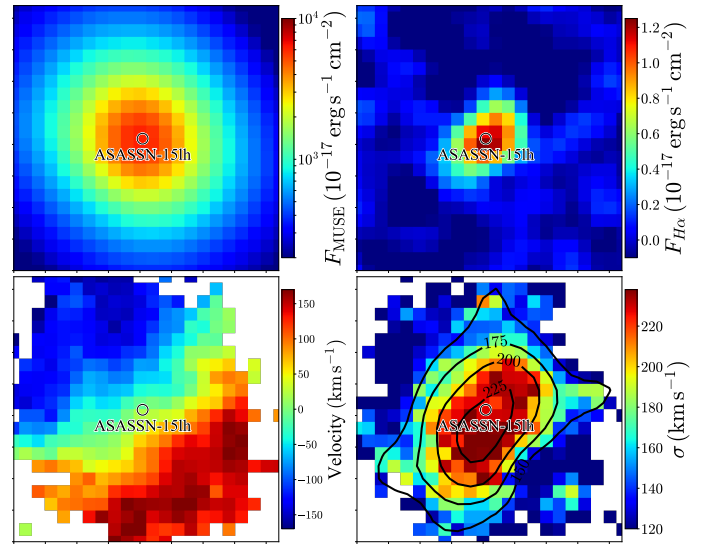


Fig. 5: In the top row, reconstructed images of the ASASSN-15lh host where the total intensity in the MUSE spectral range (4650 Å to 9300 Å) is shown in the left panel, and the central component of H α emission in the right panel. The bottom row shows the stellar velocity field (left panel) and velocity dispersion σ (right panel). Each of these maps is $6''.2$ by $6''.2$, which corresponds to 23 kpc by 23 kpc at the redshift of ASASSN-15lh. North is up and East is to the left.

This value of σ_e corresponds to a mass M_\bullet of the central black hole of $M_\bullet = 5.3^{+8.0}_{-3.0} \cdot 10^8 M_\odot$ (Eq. 3, 5 or 7 in Kormendy & Ho 2013). The quoted uncertainty is dominated by the intrinsic scatter in the M_\bullet - σ relation (~ 0.3 dex), and errors in measuring and inferring σ_e from our data do not contribute significantly. This mass estimate is similar to those obtained previously (Prieto et al. 2015; Leloudas et al. 2016), but has somewhat smaller uncertainties. The Eddington luminosity for this black hole mass is $L_{\text{Edd}} = 7^{+10}_{-4} \cdot 10^{46} \text{ erg s}^{-1}$, which means that ASASSN-15lh radiated at $\sim 5\%$ of the SMBH's Eddington ratio ($L_{\text{peak}} \sim 3 \cdot 10^{45} \text{ erg s}^{-1}$) at peak, which is quite similar to the bulk of previously studied TDEs (Hung et al. 2017; Wevers et al. 2017).

Figure 5 summarizes galaxy properties through four maps of total intensity in the MUSE wavelength range (observed 4650 Å to 9300 Å), flux in the central component of the H α emission (Sections 3.1 and 3.4), line of sight velocity field V , and velocity dispersion³. The peak of the total intensity as well as the H α flux aligns well with the ASASSN-15lh transient, and the velocity field shows clear rotation in the direction of the photometric major axis (NE-SW direction). The stellar velocity dispersion rises towards the center and shows a somewhat elongated structure along the galaxy minor axis. Its center is slightly ($\sim 0''.2 \pm 0''.1$) offset⁴ from the peak of the total intensity which can be interpreted as a sign that the galaxy center has been disturbed by a past merger, but the significance of the offset is low. In general, the central velocity dispersion field of early-type galaxies is very

³The galaxy kinematics were derived after summing nine adjacent spaxels to obtain a sufficiently high S/N, and pixels in the σ and V maps are hence correlated.

⁴The total error is a combination from uncertainties in the measurement of σ , errors during centroiding and the correlation between individual spaxels due to seeing and spectral extraction

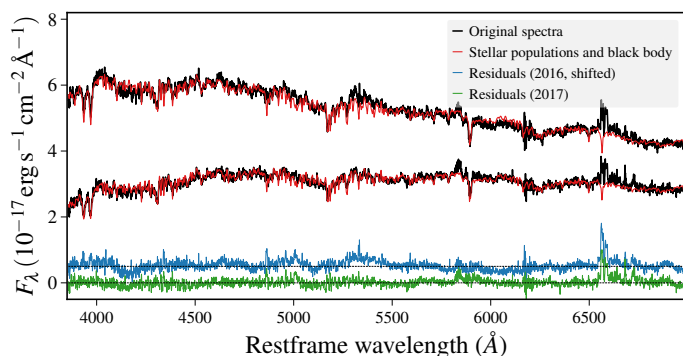


Fig. 6: A physical model of transient plus multiple stellar populations in comparison to the nuclear spectra (radius of $0''.5$) over the MUSE wavelength range. The 2016 data returns broad features in the residuals, for example at 4200 Å, 5300 Å, and 6000 Å, which are much less pronounced in the 2017 data, and represent hence likely deviations from the ASASSN-15lh spectrum from a pure black-body.

rich in features (Emsellem et al. 2004) and our limited spatial resolution (FWHM=3 kpc) prevents us here from making stronger claims with respect to the origin, nature and implications of the velocity dispersion field.

3.3. The central stellar population

We also performed a second set of fits to derive a more physical interpretation of the continuum by modeling the transient with a black-body component (as opposed to the low-order polynomial in the Section 3.1) and the galaxy with a superposition of templates from single stellar populations (Fig. 6). These fits with a black-body continuum turned out to be less accurate in disentangling emission lines from the continuum as they returned broad residuals in the subtracted spectrum. The residuals are likely deviations of the ASASSN-15lh spectrum from a pure black-body, as similar fits to spectra of the EI and EII galaxies (Fig. 1) show no residuals, and they are much weaker in the 2017 epoch (Fig. 6). Here, we used starlight (Cid Fernandes et al. 2005, 2009) and Bruzual & Charlot (2003) templates in a similar manner as we have described in detail elsewhere (Galbany et al. 2016; Krühler et al. 2017).

The detailed physical parameters of the central stellar population vary somewhat depending on the exact choice of the size of the spectral extraction, the best-fit transient black-body temperature, as well as the base list of galaxy templates. However, contributions from two stellar populations are always required for a satisfactory fit of the stellar component in the galaxy center: a dominant population with an age of around 1 Gyr to 2 Gyr (typically contributing around 60% to 80% to the composite galaxy spectrum), and a second population that is significantly older with an age between 10 Gyr and 13 Gyr and contributing between 10% and 30% of the stellar light⁵. The presence of a 1-2 Gyr old population is compatible to the oldest TDE hosts studied previously (French et al. 2016, 2017).

The strength of the H δ absorption line has often been used for a post-starburst, or E+A, classification in galaxies (Dressler

et al. 1999; Worthey & Ottaviani 1997), and in particular galaxies with TDEs (e.g., Arcavi et al. 2014; French et al. 2016; Law-Smith et al. 2017). For the ASASSN-15lh host, we estimate an $\text{EW}_{\text{H}\delta} \sim 2$ Å by measuring the optical depth of the H δ absorption in the spectra and comparing it to the continuum from galaxy models derived via pre-explosion photometry⁶ at the respective wavelength. Following Dressler & Gunn (1992), an $\text{EW}_{\text{H}\delta} \sim 2$ Å together with the absence of [O II]($\lambda 3727$) in our spectra (Figure 4) classifies the host as a passively-evolving galaxy.

3.4. Emission-line profiles

Once the continuum from stars and transient has been separated from the emission lines, we look more closely to the line shape and velocity structure of the H α and [N II] complex of Fig. 3 in order to study the ionization source, and its relation to the transient ASASSN-15lh itself.

In Fig. 7, we show a zoom in on the continuum-subtracted wavelength range of the respective transitions from our MUSE and X-Shooter spectroscopy. The line-shape of the individual transitions is complex, and each line transition is composed out of three components that are well-separated in wavelength space - a central, broader component (FWHM = 8 Å or 380 km s⁻¹), as well as a narrower (FWHM = 3 Å or 100 km s⁻¹) blue and red component offset by approximately 5 Å (250 km s⁻¹) in each direction. The redshifts of the three components are $z = 0.2310$, $z = 0.2318$ and $z = 0.2331$ with errors of around ± 0.0002 each.

The originally reported redshift was $z = 0.2326$ from narrow Mg II absorption lines (Dong et al. 2015), which we also confirm with our spectra. However, Mg II absorption does not necessarily yield the most accurate galaxy redshift, as the Mg II gas clouds are subject to random motion within the gravitational potential of the galaxy. Hence, we adopt 0.2318 ± 0.0002 measured in a heliocentric reference frame for the host's systemic redshift from galaxy emission and stellar absorption lines in the following. This change is insignificant for this or any of the previous articles about ASASSN-15lh.

The line shape of the H α line and each of the collisionally excited [N II]($\lambda\lambda 6548, 6584$) lines is identical within the measurement uncertainties. Also, the line shape is well-resolved, in particular through the medium-resolution X-Shooter data, and appears comparable between the two spectrographs⁷.

The emission lines from the H β and [O III] transitions are generally consistent with this picture (Fig. 3), in particular H β shows evidence for a similar line-profile. However, large statistical errors stemming from the bright background of galaxy and transient as well as systematic uncertainties from the continuum subtraction prevent us from performing a detailed kinematic analysis for any other lines except H α and [N II].

To derive line-fluxes and gas kinematics, we fit a superposition of three Gaussians for each of the three transitions simultaneously to the spectra shown in Fig. 7. The intrinsic line-width, broadened by the instrumental resolution, is tied between both instruments and the three transitions, and only the normalization is allowed to vary during the fit (while the MUSE spectra

⁶Even our latest spectra from 2017 still show a weak contamination from the transient so that a direct application of the Lick indices is not possible.

⁷The width of the instrumental resolution at the wavelength range of H α (observed 8100 Å) has a FWHM ~ 110 km s⁻¹ for MUSE and a FWHM ~ 35 km s⁻¹ for X-Shooter.

⁵Previously estimated stellar population ages of around 5 Gyr for the ASASSN-15lh host were based on pre-explosion broad-band photometry (Melchior et al. 2015; Leloudas et al. 2016; Dong et al. 2016), and as such only probe the integrated light of many stellar populations.

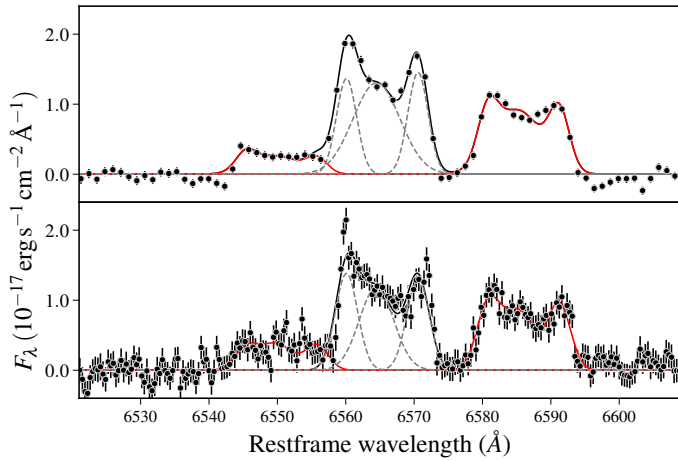


Fig. 7: Emission-line profiles and component decomposition for the $H\alpha$ and $[N II](\lambda 6548, 6584)$ complex. The *upper panel* shows a spectrum extracted for the central component (radius of 0.9 or 3.4 kpc around the galaxy center) of the host galaxy from the MUSE datacube, the *lower panel* shows the X-Shooter spectrum. The contribution from $[N II]$ is shown in red, the total sum of the emission line model in black. For $H\alpha$, we also plot the three individual Gaussian components with dashed lines.

should measure the full flux of the emission lines, the slit from X-shooter might lead to slit-losses).

3.5. Positional analysis of the emission-line components

The MUSE integral-field spectroscopy allows us to go beyond a standard kinematic analysis as done above, and perform a spatially-resolved analysis of the individual velocity components. Here, we created a continuum-subtracted data cube from the original MUSE spectroscopy by performing a fit similar to the one of Fig. 3, but now for each individual spaxel in the astrometrically-calibrated data (Sect. 2.3). Similar procedures have been used by us frequently in the past on MUSE data (Galbany et al. 2016; Prieto et al. 2016; Krühler et al. 2017), and allow us to combine and visualize the spatial information of the MUSE maps with the velocity information of the emission line kinematics.

Figure 8 shows the channel maps of the continuum-subtracted spectroscopy at the center of the ASASSN-15lh host, and in the wavelength range of $H\alpha$. Each of the three panels shows the reconstructed image in the given wavelength range, whereas the rightmost panel is a subtraction between the bluest and reddest component. The position of the transient as derived through the HST-to-MUSE astrometric alignment is indicated by a cross.

Its evident that the three velocity components are not only separated in velocity space, but are also located at different positions. In addition, the blue and red components are offset from the transient location. The velocity separation between the blue and red component is 500 km s^{-1} , and the spatial offset is 1.8 ± 0.3 MUSE spaxel or 0.36 ± 0.06 , corresponding to a projected distance of 1.3 ± 0.2 kpc.

The central component is placed symmetrically between the red and blue emission peaks (both spatially, as well as in velocity space), and is consistent with the transient position within the combined astrometric uncertainty of 0.22 MUSE spaxels, or 45

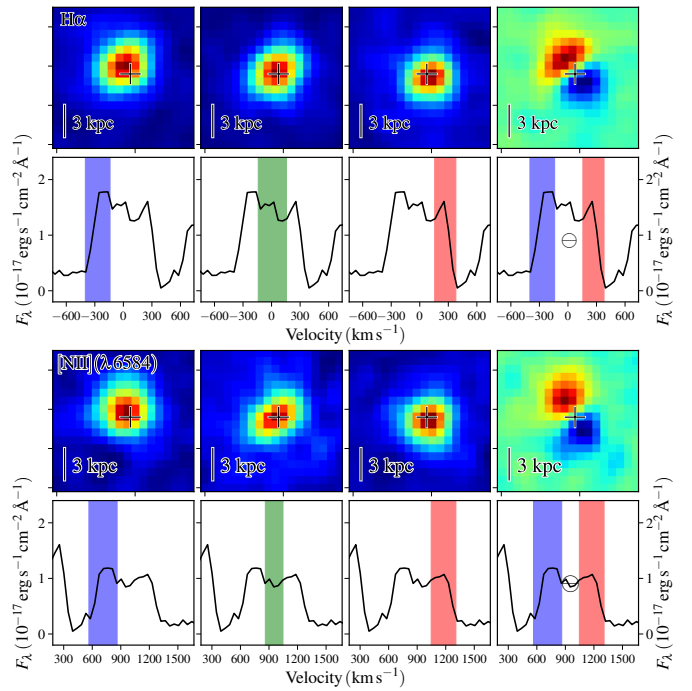


Fig. 8: Channel maps of the velocity components. The two *upper rows* show in the three leftmost panels the reconstructed images in a given velocity range (indicated by the coloured region in each spectrum) from the continuum-subtracted MUSE map for $H\alpha$. The right panel shows a subtraction between the blue and red components. The position of the transient is indicated by a cross. The two *lower rows* show the same for $[N II]$. All images have been smoothed by a Gaussian kernel with a size of the FWHM of 2 px by 2 px for illustration purposes. The physical scale of the images is indicated by the bar in the lower-left corner of each image. North is up, and East is to the left in all images.

mas, which corresponds to a physical scale of 170 pc at $z = 0.232$.

3.6. Ionization source

The strength and ratios of various collisionally-excited and recombination lines of metal and hydrogen ions trace the physical conditions in the gas-phase as well as the origin of the radiation that was ionizing the gas in the first place. A useful and widely-used diagnostic plot is the Baldwin-Phillips-Terlevich (BPT) diagram (Baldwin et al. 1981), which discriminates between ionizing flux coming from the hard radiation of AGN or shocks or $H II$ regions where the UV-flux is dominated by massive stars through the ratios of $[N II]/H\alpha$ and $[O III]/H\beta$.

The BPT diagnostic is frequently used for host galaxies of transient objects, and different classes of objects occupy very different phase-space. For example, the hosts of γ -ray bursts or super-luminous supernovae typically reside in the high $[O III]/H\beta$, low $[N II]/H\alpha$ regime (Krühler et al. 2015; Leloudas et al. 2015), characteristic of young starbursts at low metallicity. In contrast, the nearby ($d_L = 90 \text{ Mpc}$) TDE ASASSN-14li (Holoien et al. 2016) showed an extended structure of ionized gas with emission-line ratios that imply ionization from an AGN (Prieto et al. 2016).

The ionized regions of ASASSN-15lh are plotted in the BPT diagram in Fig. 9, where we use the SDSS DR7 spectroscopy (Abazajian et al. 2009) with line fluxes from the MPA/JHU cata-

log as a background sample. The total fluxes of the four emission lines ($H\alpha$, $H\beta$, $[N II](\lambda 6584)$, $[O III](\lambda 5007)$) and thus their ratios are rather well constrained, in particular when adding the constraint that $[O III](\lambda 5007)/[O III](\lambda 4959)=3$ in the fit. However, for the $[O III]$ and $H\beta$ lines, the individual components are not easy to separate and the flux ratio has hence large uncertainties.

It is evident that all individual components of the line emission, and their integrated flux, are located in the part of the BPT diagram that is occupied by low-luminosity AGN and shock ionization/excitation, similar to many other TDE hosts (French et al. 2017). And even though the measurement error, especially for the $[O III]/H\beta$ ratio is substantial, it is clear that all components occupy a region in the plot that is offset from the star-forming sequence of SDSS galaxies.

In particular the central component, which is positionally coincident with the transient, has a high value of $[N II]/H\alpha$, inconsistent with pure star-formation. Our measurements of line ratios hence require that at least a significant fraction of the ionization is coming from AGN or shocks (Cid Fernandes et al. 2011). Similarly, a classification based on the equivalent width (EW) of $H\alpha$, as well as the $[N II]/H\alpha$ ratio (Cid Fernandes et al. 2011) shows the central region of the ASASSN-15lh host to be consistent with a weak AGN ($[N II]/H\alpha > 0.4$, and $3 \text{ \AA} < EW_{H\alpha} < 6 \text{ \AA}$), sometimes referred to as low-ionization nuclear emitting regions or LINERs.

Given that the $H\alpha$ emission is more compact than the stellar emission, the exact value of $EW_{H\alpha}$ depends on the size of the spectral extraction. Using central spectra extracted in the region of the line emission (0.9 radius), $EW_{H\alpha} = 3.2 \pm 0.4 \text{ \AA}$, while it is $EW_{H\alpha} = 1.3 \pm 0.2 \text{ \AA}$ when considering the full extend of the galaxy.

The leading theoretical model to explain LINERs is photo-ionization from a central, low-luminosity AGN (e.g., Ho 2008, for a review). In contrast to other ionization sources (young stars, fast shocks, evolved stars), a narrow-line region of a central AGN would naturally explain the line-ratios, the kinematics, and, as shown in Section 4.1, the spatial offset between the three observed kinematic components seen in the host of ASASSN-15lh.

We conclude that the observed line emission is most likely the radiation from gas photo-ionized by AGN radiation, and neither from star-formation nor the transient itself. Deriving exact limits on the SFR from the $H\alpha$ emission is complex: while we can exclude from the BPT diagram that all of the ionized gas emission comes from star-formation, we can not strictly exclude this for a smaller fraction. Assuming that the star-formation powered $H\alpha$ line flux is less than half of the observed flux, the SFR would be $SFR \lesssim 0.1 M_{\odot} \text{ yr}^{-1}$ (Kennicutt 1998). This implies that the host is a quiescent galaxy, lying at least two orders of magnitude below the star-formation rates of similarly massive galaxies on the main sequence (e.g., Peng et al. 2010; Whitaker et al. 2012).

3.7. Constraints on the bolometric AGN luminosity

Summing over all spectral components, the total $[O III]$ or $H\alpha$ luminosity of the gas photo-ionized by the AGN is $L_{[O III]} \sim 10^{40} \text{ erg s}^{-1}$ or $L_{H\alpha} \sim 3 \times 10^{40} \text{ erg s}^{-1}$, respectively. These values correspond to an X-ray luminosity of around $L_X \sim 10^{41} \text{ erg s}^{-1}$ in the 2-10 keV energy range, or a bolometric luminosity of $L_{bol} \sim 10^{42} - 10^{43} \text{ erg s}^{-1}$ assuming average correction factors (Ho 2008; Lamasstra et al. 2009; Lusso et al. 2012). Even though these luminosity estimates are uncertain by at least an order of

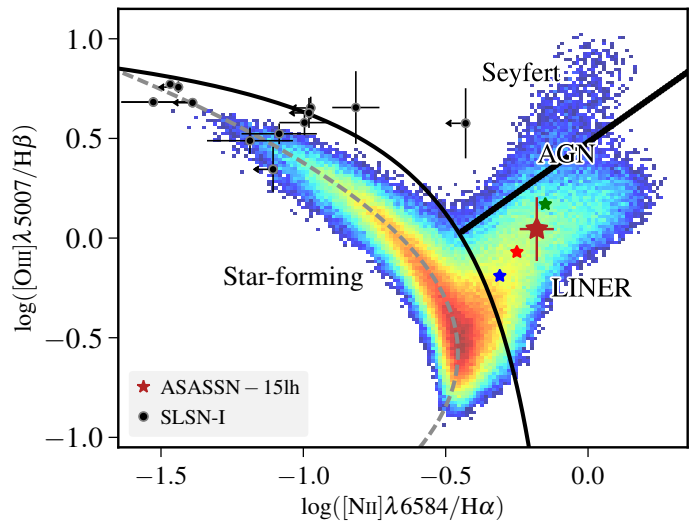


Fig. 9: The host of ASASSN-15lh in the BPT diagram. The large, red star represents the total integrated flux, while the three smaller stars show the three individual components in their respective color, i.e., the green is the central component. Error-bars of the individual components are not shown to enhance clarity in the figure. To indicated the size of the respective error bars we note that the location of the smaller stars is consistent with each other within 1σ . Black data are hydrogen-poor SLSN from Leloudas et al. (2015), with limits indicated by arrows. The black solid lines are differentiation lines between star-forming galaxies and AGN from Kewley et al. (2013) at $z \sim 0.23$, and between LINERs and Seyferts from Cid Fernandes et al. (2010). The grey dashed line is the ridge line, i.e., the line with the highest density of star-forming galaxies in SDSS (Brinchmann et al. 2008).

magnitude, the Eddington ratio λ_{edd} of the central black hole is far below unity ($\lambda_{edd} \sim 10^{-5} - 10^{-4}$) for the AGN emission.

The inferred X-ray luminosity of a weak AGN is comparable to the emission observed by Margutti et al. (2017), who estimate a 0.3-10 keV luminosity of $L_X \sim 10^{41} - 10^{42} \text{ erg s}^{-1}$. Given that this measurement is derived in a somewhat larger energy interval, and the X-ray source seems relatively soft, we can not rule out that the observed X-ray flux is caused by steady accretion of a low-luminosity, pre-existing AGN. Given that the pre-transient, AGN-dominated X-ray emission estimated here, as well as the recent X-ray detection (Margutti et al. 2017) are both uncertain by at least an order of magnitude, the available constraints on X-ray fluxes are also fully consistent with substantial, order of magnitude, X-ray variability.

4. Discussion

4.1. The nature of the ionized-gas emission

As described in the previous paragraphs (Section 3.4 and 3.5), the emission-line profile observed in the ASASSN-15lh host galaxy nucleus is complex, and consists of multiple components with spatial and velocity offsets (Figures 7 and 8). The components are aligned along a single direction, and all are likely to be ionized by an AGN. The double-peaked nature of the $H\alpha$ and $[N II]$ line initially led us to suspect a binary AGN as origin of the emission. Giant early-type galaxies such as the ASASSN-15lh host are thought to be the result of galaxy mergers (e.g., Hopkins et al. 2006, and references therein). Because supermassive black holes (SMBHs) arguably reside in the centers of all mas-

sive galaxies (e.g., Kormendy & Ho 2013, for a review), a good fraction of spheroidal galaxies should also host binary SMBHs. If active, for example through nuclear accretion induced by the merger, the double SMBH will appear as a binary AGN. The most convincing cases of binary AGN on kpc scales, as would be the case here, have been imaged as double point source in hard X-ray (e.g., Komossa et al. 2003; Bianchi et al. 2008) or radio emission (Fu et al. 2011; Müller-Sánchez et al. 2015).

However, this scenario seems unable to fully explain our observations: while the line-shape from our MUSE data could conceivably be explained with two components only, the higher-resolution X-Shooter data clearly demonstrates the presence of an even more complex structure (Fig. 7). We would hence need to invoke three aligned and active SMBH, which we consider too contrived to explore any further. Only the medium-resolution X-Shooter data (FWHM=35 km s⁻¹) has allowed us to convincingly rule out this possibility. Based on the spectral resolution of MUSE (FWHM=150 km s⁻¹) alone, we would have probably considered a binary AGN as the cause of the line profile more seriously.

Instead, narrow-line gas kinematics in a rotating disk, or galactic winds driven by the AGN offer a much more natural explanation for the observed kpc-scale emission in Balmer and collisionally-excited metal lines (e.g., Shen et al. 2011). Indeed, complex narrow-line regions (NLRs) are explained with AGN outflows at least for some nearby galaxies (e.g., Fischer et al. 2011) but the differentiation between rotating NLR disks and genuine outflows is not always trivial (Shen et al. 2011; Müller-Sánchez et al. 2015) in general. In our case, we observe a relatively symmetrical line-profile, as well as a geometry that is aligned along the major axis of the galaxy (Figs. 2 and 8), both of which argues for an origin in a rotating disk of ionized gas. Clearly, the stellar population is much more extended than the emission lines (Figure 5), and neither the red nor blue component of the emission lines has a kinematic counterpart to absorption lines in the core. However, the stellar velocity field co-rotates with ionized gas even though on much larger scales (Figure 5). We are hence lead to suggest that the emission line shape and position of all components is caused by an AGN at the position of the central component with a rotating disk of ionized gas, but clearly not multiple black holes on kpc scales. **The observed velocity and radius of the narrow-line disk then constrains the enclosed mass M_{enc} . Assuming that the disk is viewed nearly edge-on, using a circular rotation velocity of $v_{\text{H}\alpha} \sim 250 \text{ km s}^{-1}$, radius $r_{\text{H}\alpha} \sim 650 \text{ pc}$ (Section 3.5) and following Gezari et al. (2003), we derive $M_{\text{enc}} \sim 9 \times 10^9 M_{\odot}$, a factor 20 larger than the mass of the SMBH, and a factor 10 lower than the total stellar mass of the galaxy Leloudas et al. (2016).**

The [O III] luminosity of Seyferts correlates with the size of the narrow-line region (e.g., Schmitt et al. 2003). This correlation would predict a radius of 200 pc, significantly smaller than the observed radius of 650 pc. Or conversely, the observed NLR radius corresponds to a [O III] luminosity of $L_{[\text{O III}]} \sim 10^{41.5} \text{ erg s}^{-1}$. This mismatch possibly indicates that the central ionizing source was more active in the past, leaving behind only an extended narrow-line region.

4.2. The origin of the transient

Having pinpointed the location of the transient to the position of a weak AGN (and therefore a supermassive black hole), it is reasonable to relate both of these phenomena. The ASASSN-15lh environment thus constrains on the nature of ASASSN-15lh it-

self, with three different scenarios typically discussed in the pertinent literature in similar cases (e.g., van Velzen et al. 2011; Drake et al. 2011; Holoien et al. 2014; Vinkó et al. 2015): a very luminous core collapse supernova in the nuclear region of the host, a tidal disruption event or intrinsic variability from the AGN.

The first of those, a luminous supernova at the nucleus of a passive galaxy is somewhat contrived for ASASSN-15lh in several ways: firstly, there are the inconsistencies of the spectral properties and temporal evolution of the transient with other SLSNe or luminous SNe of type IIn as mentioned in the introduction or discussed in Leloudas et al. (2016). In addition, the association of a massive-star related phenomenon with an early-type host, a \sim Gyr old stellar population, and no obvious signs of recent star-formation does not seem very viable. Finally, the positional coincidence with a NLR of a weak AGN and its location in the BPT diagram (Fig. 9) strongly suggest that ASASSN-15lh is the result of a physical phenomenon closely related to the SMBH and not star-formation.

AGN, and to some extent also non-active SMBHs such as the one in the center of the Milky-Way, are known to be variable throughout the electromagnetic spectrum on various time-scales (e.g., Ulrich et al. 1997; Baganoff et al. 2001). Typically, these forms of AGN variability are considered to be stochastic resulting from changes in the accretion rate, and are therefore clearly not able to explain the dramatic ASASSN-15lh variability and spectral evolution. However, much more extreme variations have been observed in changing-look quasars on shorter time-scales (e.g., Shappee et al. 2014; LaMassa et al. 2015; Gezari et al. 2017), even though also here the discrimination between AGN activity and TDEs is not always trivial (Merloni et al. 2015). Two lines of evidence indicate that such an AGN-related event is not the origin of ASASSN-15lh itself. Firstly, no strong X-ray variability contemporaneous with large UV variability for ASASSN-15lh is observed, despite regular and simultaneous monitoring (Brown et al. 2016; Margutti et al. 2017). And secondly, the nebular line emission is constant in the various epochs of our spectroscopic monitoring. There is no evidence for appearing or disappearing of broad emission lines nor an increase in the continuum from a Seyfert 1 in our spectra, arguing against the interpretation of ASASSN-15lh as a changing-look quasar.

In the light of our detailed observation of the transient's environment, it thus appears that the explanation of ASASSN-15lh as a TDE (Leloudas et al. 2016; Margutti et al. 2017) remains the most plausible one. TDEs within galaxies with regions of AGN-related ionization and excitation have been discovered before, with ASASSN-14li as the best-observed event (van Velzen et al. 2016; Holoien et al. 2016; Prieto et al. 2016). A more recent candidate is reported in Blanchard et al. (2017). Of course, a TDE is not the only explanation that would physically relate the transient with a SMBH in the center. Moriya et al. (2017), for example, propose that some luminous nuclear transients within AGN are due to an interaction between accretion disk winds and clouds in the broad-line region. Because of the non-detection of AGN broad lines within the ASASSN-15lh host, a direct application of this scenario to our situation seems not straight forward, however, it serves to illustrate that other physical mechanisms apart from TDEs are still clearly interesting to explore for the remarkable transient emission observed as ASASSN-15lh.

4.3. The spin of the supermassive black hole

As noted by other authors already (e.g. Prieto et al. 2015; Dong et al. 2016; Leloudas et al. 2016), and shown in Figure 10, the

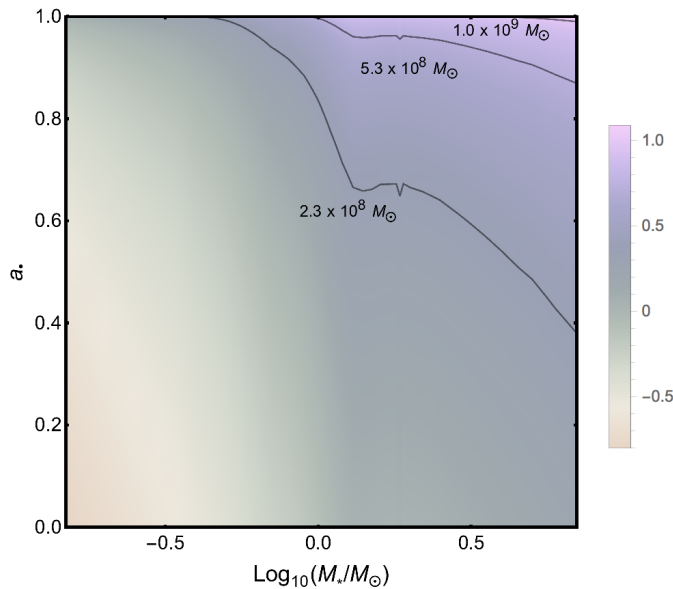


Fig. 10: Dependence between minimum spin parameter a_* (y-axis), stellar mass of the disrupted star (x-axis) and SMBH mass in $10^8 M_\odot$ in color. Shown in black are also three tracks of specific SMBH masses, showing that there is no viable solution for TDEs for zero black-hole spin, or SMBH masses greater than $10^9 M_\odot$ within the constraints on the other parameters from this work. This plot assumes optimal orbital orientation between disruption axis and the spin vector, uses the ZAMS mass-radius relation for stars with $Z = 0.008$ from Girardi et al. (2000), the formalism of Kesden (2012) and follows the calculations detailed in the supporting material of Leloudas et al. (2016).

SMBH mass of $M_\bullet = 5.3^{+8.0}_{-3.0} \cdot 10^8 M_\odot$ remains too high for a non-spinning black hole to disrupt stars of $< 2.5 M_\odot$ (the most massive stars with lifetimes shorter than the approximate age of the starburst from Section 3.3). These parameters corresponds to a limit⁸ on the black-hole spin $a_* \gtrsim 0.4$, when assuming $M_\odot < 2.5 M_\odot$ and $M_\bullet > 2 \cdot 10^8 M_\odot$ (Figure 10).

The argument can also be reversed for the highest SMBH masses in the confidence interval of the $M_\bullet = 5.3^{+8.0}_{-3.0} \cdot 10^8 M_\odot$ measurement. Even for a maximally-spinning black-hole, SMBH masses of $M_\bullet > 10^9 M_\odot$ would not lead to luminous emission from the disruption of a $2.5 M_\odot$ star. Or, in other words, ASASSN-15lh can not be a TDE if the SMBH has $M_\bullet > 10^9 M_\odot$ (Fig. 10).

An alternative way for creating luminous emission from a stellar disruption in such a massive galaxy could be a tight binary SMBH where the two black holes have very different mass, and the lower-mass secondary is sufficiently light to produce a bright TDE (Coughlin & Armitage 2017). This scenario, however, would leave the ≈ 10 times higher luminosity of ASASSN-15lh compared to other TDEs unexplained. In addition, SMBH binaries are rare and probably contribute only $\sim 3\%$ of the cosmic TDE rate (Wegg & Nate Bode 2011), and of those binary-induced disruptions, the fraction of TDEs caused by low-mass secondaries is only $\sim 1\%$ (Wegg 2013).

⁸This limit is based on the observational quantities, and does not take into account that disruptions of lower mass stars are more likely - which would increase the required Kerr parameter.

4.4. The large-scale environment

The co-moving area of 220 kpc by 220 kpc shown in Fig. 1 contains a number of galaxies with a similar redshift as ASASSN-15lh. Within 200 km s^{-1} , there are two early-type galaxies (EI and EII in Fig. 1) at 40 kpc and 70 kpc projected distance, respectively, as well as a small satellite (20 kpc projected distance) to the north of the host that shows signs of star-formation through the detection of the $H\alpha$ emission line (SFI). At somewhat larger distance (100 kpc) and velocity offset (1800 km s^{-1}), two star-forming galaxies (SFII and SFIII) are connected through tidal tails and are therefore strongly interacting.

Remarkably, these six galaxies align along a specific direction (North-South). No other galaxies in the field of view with measured redshift are similarly close to ASASSN-15lh in velocity space, but these four galaxies (or six, when including the two more distant merging galaxies) constitute a significant galaxy overdensity. The ASASSN-15lh host is the most massive ($M_\star = 10.95^{+0.15}_{-0.11} M_\odot$) member observed, and possibly the node of a larger gravitationally bound system. Such an overdense environment is frequently observed around E+A galaxies as well (Goto 2005), and interactions will be common. In fact, a major merger between two gas-rich spirals (SFII and SFIII) is even directly observed within this association.

The massive, early-type host of ASASSN-15lh is thus plausibly the result of a previous interaction, where the age of the youngest stellar population observed (1–2 Gyr, Sect. 3.3) indicates the typical time-scale of the merger. This scenario is broadly consistent with the LINER-like signatures in the central component (e.g., Cid Fernandes et al. 2004), indicating the presence of a faint AGN, as well as the early-type morphology and other galaxy properties such as the lack of current star-formation.

4.5. Comparison to other TDE hosts

It is useful to compare the ASASSN-15lh host to other, well-observed and more nearby TDE hosts with similar data. The galaxy of ASASSN-14li, for example, has comparable line ratios in the narrow-line region, suggestive of a similar ionization and excitation process of the emission lines. In fact, the emission-line properties of the sample of TDE hosts studied by French et al. (2017) are in general in good agreement with the line properties observed here.

In contrast to ASASSN-14li, strong tidal tails are not observed here, neither in stellar light or in narrow emission lines. This might indicate that the ASASSN-15lh host is in a later stage after the galaxy-galaxy interaction, with more time to relax into an undisturbed morphology, or simply that our spectroscopic observations are not deep enough to probe features of a past merger. In fact, most TDE hosts including ASASSN-14li have a relatively symmetric distribution of the stellar light (Law-Smith et al. 2017), indicating that a previous galaxy interaction is not directly obvious in most cases.

The TDE rate seems to be significantly enhanced in E+A galaxies (Arcavi et al. 2014; French et al. 2016), which are thought to be observed few 100 Myrs after a star-burst, likely induced through a galaxy merger (Goto 2005). In our case, the stellar population seems somewhat older (of order Gyr) but generally consistent with the distribution of ages and other properties in TDE host samples (French et al. 2017; Law-Smith et al. 2017). It also displays LINER-like signatures that are also often present in E+A galaxies (Yang et al. 2006). This is again quite similar to the TDE hosts studied in French et al. (2017), which in many cases are also LINERS.

These considerations thus indicate a common evolutionary path for the ASASSN-15lh host and other galaxies with TDEs. A galaxy-galaxy interaction leads to a star-burst and possibly subsequent AGN activity, where the star-formation has ceased to the present day, and the SMBH only accretes at a very low rate ($\lambda_{\text{edd}} \sim 10^{-5} - 10^{-4}$) in our case). The nuclear starburst leads to a high stellar density, which in turn increases the TDE rate due to the short relaxation time for two-body interactions (Metzger & Stone 2016; Stone & van Velzen 2016; Graur et al. 2017). The evolutionary stage is set by the age of the youngest stellar component, and indicates a time-scale of around a Gyr since the star-burst, somewhat older than the average timescales of other TDE hosts.

5. Summary and Conclusions

We have presented here HST imaging as well as spatially-resolved, medium resolution spectroscopy from the ESO VLT instruments (X-Shooter and MUSE) of the environment of the luminous transient ASASSN-15lh. Based on these data, we conclude that:

- (i) The spectrum of the galaxy nucleus consists of three components: a stellar component with at least two different stellar populations (1-2 Gyr and 10 Gyr), the transient emission that is reasonably well described with a black-body of $\sim 13\,000$ K (at 430 days after peak), and constant, narrow (~ 500 km s $^{-1}$) line emission from ionized gas.
 - (ii) The line emission is related to the transitions of H α , H β , [N II], and [O III], and splits up into three components that are separated spatially by 1.3 kpc and in velocity by 500 km s $^{-1}$.
 - (iii) From their position in the BPT diagram, we show that the line ratios are consistent with LINER-like excitation, and we demonstrate that ionization by a weak AGN, and not star-formation, is the likely origin of the observed emission lines.
 - (iv) The central emission-line component is positionally coincident with the ASASSN-15lh transient, and we suggest that this is also the position of the central supermassive black hole in the host galaxy. The mass of the black hole derived from the stellar velocity dispersion σ is $M_{\bullet} = 5.3^{+8.0}_{-3.0} \cdot 10^8 M_{\odot}$. We argue that the origin of the complex emission line shape is best explained by kinematical effects (a rotating circumnuclear gas disk, or less-likely, outflows) in the AGN narrow-line region.
 - (v) The spatial association of the transient with a supermassive black hole, together with no detectable star-formation, leads us to favor a physical mechanism for the transient that involves the supermassive black hole, and from those specifically the tidal disruption of a star by a spinning SMBH.
 - (vi) The observed physical properties of the ASASSN-15lh environment are in striking contrast to those of explosive phenomena related to the death of very massive stars like hydrogen-poor SLSNe (e.g., Lunnan et al. 2014; Schulze et al. 2016) or long GRBs (e.g., Krühler et al. 2015; Lyman et al. 2017; Schady 2017). We consider them much more akin to those of TDEs, even though the post-starburst timescale of 1 to 2 Gyr observed here is slightly longer than the average of other TDE hosts (e.g., French et al. 2017; Law-Smith et al. 2017), and the host's stellar and SMBH mass significantly higher.
- This interpretation is consistent with the galaxy morphology and spectrum as a passively-evolving galaxy, the galaxy's location in an overdense larger-scale environment, and the faint AGN emission in its center. All this suggests that the ASASSN-15lh host went through a very active phase of star-formation and AGN activity roughly one or two Gyr in the past, possibly triggered by a galaxy-galaxy interaction.

Acknowledgements. We are very grateful to the referee for a timely and very constructive report as well as I. Arcavi, L. Christensen, J. Greiner, P. Schady, and L. Izzo for helpful comments on the manuscript, which increased the clarity and quality of the manuscript. T.K. acknowledges support through the Sofja Kovalevskaja Award to P. Schady from the Alexander von Humboldt Foundation of Germany. M.F. acknowledges the support of a Royal Society - Science Foundation Ireland University Research Fellowship. N.C.S. received financial support from NASA through Einstein Postdoctoral Fellowship Award Number PF5-160145, and thanks the Aspen Center for Physics for its hospitality during the completion of this work. R.A. acknowledges support from the ERC Advanced Grant 695671 "QUENCH". We acknowledge the use of NumPy and SciPy (Walt et al. 2011) for computing and matplotlib (Hunter 2007) for creating the plots in this manuscript. We thank ESO's Director's Discretionary Time Committee for allocating telescope time for this project, and the observing staff on Paranal for support in obtaining the MUSE and X-Shooter data.

References

- Abazajian, K. N., Adelman-McCarthy, J. K., Agüeros, M. A., et al. 2009, *ApJS*, 182, 543
- Arcavi, I., Gal-Yam, A., Sullivan, M., et al. 2014, *ApJ*, 793, 38
- Bacon, R., Accardo, M., Adjali, L., et al. 2010, in *Proc. SPIE*, Vol. 7735, Ground-based and Airborne Instrumentation for Astronomy III, 773508
- Baganoff, F. K., Bautz, M. W., Brandt, W. N., et al. 2001, *Nature*, 413, 45
- Baldwin, J. A., Phillips, M. M., & Terlevich, R. 1981, *PASP*, 93, 5
- Bekki, K. & Couch, W. J. 2011, *MNRAS*, 415, 1783
- Bersten, M. C., Benvenuto, O. G., Orellana, M., & Nomoto, K. 2016, *ApJ*, 817, L8
- Bianchi, S., Chiaberge, M., Piconcelli, E., Guainazzi, M., & Matt, G. 2008, *MNRAS*, 386, 105
- Blanchard, P. K., Nicholl, M., Berger, E., et al. 2017, *ArXiv e-prints* [arXiv:1703.07816]
- Blanton, M. R. & Moustakas, J. 2009, *ARA&A*, 47, 159
- Bournaud, F., Jog, C. J., & Combes, F. 2005, *A&A*, 437, 69
- Brinchmann, J., Pettini, M., & Charlot, S. 2008, *MNRAS*, 385, 769
- Brown, P. J., Yang, Y., Cooke, J., et al. 2016, *ApJ*, 828, 3
- Bruzual, G. & Charlot, S. 2003, *MNRAS*, 344, 1000
- Cappellari, M. 2017, *MNRAS*, 466, 798
- Cappellari, M. & Emsellem, E. 2004, *PASP*, 116, 138
- Chabrier, G. 2003, *PASP*, 115, 763
- Chen, T.-W., Smartt, S. J., Yates, R. M., et al. 2016, *MNRAS*, in press [arXiv:1605.04925]
- Cid Fernandes, R., González Delgado, R. M., Schmitt, H., et al. 2004, *ApJ*, 605, 105
- Cid Fernandes, R., Mateus, A., Sodré, L., Stasińska, G., & Gomes, J. M. 2005, *MNRAS*, 358, 363
- Cid Fernandes, R., Schoenell, W., Gomes, J. M., et al. 2009, in *Rev. Mex. Astron. Astrofis. Conf. Ser.*, Vol. 35, 127–132
- Cid Fernandes, R., Stasińska, G., Mateus, A., & Vale Asari, N. 2011, *MNRAS*, 413, 1687
- Cid Fernandes, R., Stasińska, G., Schlickmann, M. S., et al. 2010, *MNRAS*, 403, 1036
- Coughlin, E. R. & Armitage, P. J. 2017, *ArXiv e-prints* [arXiv:1705.04689]
- Dong, S., Shappee, B. J., Prieto, J. L., et al. 2015, *The Astronomer's Telegram*, 7774
- Dong, S., Shappee, B. J., Prieto, J. L., et al. 2016, *Science*, 351, 257
- Drake, A. J., Djorgovski, S. G., Mahabal, A., et al. 2011, *ApJ*, 735, 106
- Dressler, A. & Gunn, J. E. 1992, *ApJS*, 78, 1
- Dressler, A., Smail, I., Poggianti, B. M., et al. 1999, *ApJS*, 122, 51
- Emsellem, E., Cappellari, M., Peletier, R. F., et al. 2004, *MNRAS*, 352, 721
- Evans, C. R. & Kochanek, C. S. 1989, *ApJ*, 346, L13
- Ferrarese, L. & Merritt, D. 2000, *ApJ*, 539, L9
- Fischer, T. C., Crenshaw, D. M., Kraemer, S. B., et al. 2011, *ApJ*, 727, 71
- French, K. D., Arcavi, I., & Zabludoff, A. 2016, *ApJ*, 818, L21
- French, K. D., Arcavi, I., & Zabludoff, A. 2017, *ApJ*, 835, 176
- Fu, H., Zhang, Z.-Y., Assef, R. J., et al. 2011, *ApJ*, 740, L44
- Gaia Collaboration, Brown, A. G. A., Vallenari, A., et al. 2016a, *A&A*, 595, A2
- Gaia Collaboration, Prusti, T., de Bruijne, J. H. J., et al. 2016b, *A&A*, 595, A1
- Galbany, L., Anderson, J. P., Rosales-Ortega, F. F., et al. 2016, *MNRAS*, 455, 4087
- Gebhardt, K., Bender, R., Bower, G., et al. 2000, *ApJ*, 539, L13
- Gezari, S., Basa, S., Martin, D. C., et al. 2008, *ApJ*, 676, 944
- Gezari, S., Chornock, R., Rest, A., et al. 2012, *Nature*, 485, 217
- Gezari, S., Halpern, J. P., Komossa, S., Grupe, D., & Leighly, K. M. 2003, *ApJ*, 592, 42
- Gezari, S., Hung, T., Cenko, S. B., et al. 2017, *ApJ*, 835, 144
- Girardi, L., Bressan, A., Bertelli, G., & Chiosi, C. 2000, *A&AS*, 141, 371

- Godoy-Rivera, D., Stanek, K. Z., Kochanek, C. S., et al. 2017, MNRAS, 466, 1428
- Goldoni, P., Royer, F., François, P., et al. 2006, in SPIE Conf. Ser., Vol. 6269
- Goto, T. 2005, MNRAS, 357, 937
- Graur, O., French, K. D., Zahid, H. J., et al. 2017, ArXiv e-prints [arXiv:1707.02986]
- Hills, J. G. 1975, Nature, 254, 295
- Ho, L. C. 2008, ARA&A, 46, 475
- Holoien, T. W.-S., Kochanek, C. S., Prieto, J. L., et al. 2016, MNRAS, 455, 2918
- Holoien, T. W.-S., Prieto, J. L., Bersier, D., et al. 2014, MNRAS, 445, 3263
- Hopkins, P. F., Hernquist, L., Cox, T. J., et al. 2006, ApJS, 163, 1
- Hung, T., Gezari, S., Blagorodnova, N., et al. 2017, ApJ, 842, 29
- Hunter, J. D. 2007, Computing In Science & Engineering, 9, 90
- Kennicutt, Jr., R. C. 1998, ARA&A, 36, 189
- Kesden, M. 2012, Phys. Rev. D, 85, 024037
- Kewley, L. J., Maier, C., Yabe, K., et al. 2013, ApJ, 774, L10
- Komossa, S. & Bade, N. 1999, A&A, 343, 775
- Komossa, S., Burwitz, V., Hasinger, G., et al. 2003, ApJ, 582, L15
- Komossa, S. & Greiner, J. 1999, A&A, 349, L45
- Kormendy, J. & Ho, L. C. 2013, ARA&A, 51, 511
- Kozyreva, A., Hirschi, R., Blinnikov, S., & den Hartogh, J. 2016, MNRAS, 459, L21
- Krühler, T., Kuncarayakti, H., Schady, P., et al. 2017, A&A, 602, A85
- Krühler, T., Malesani, D., Fynbo, J. P. U., et al. 2015, A&A, 581, A125
- LaMassa, S. M., Cales, S., Moran, E. C., et al. 2015, ApJ, 800, 144
- Lamastra, A., Bianchi, S., Matt, G., et al. 2009, A&A, 504, 73
- Law-Smith, J., Ramirez-Ruiz, E., Ellison, S. L., & Foley, R. J. 2017, ArXiv e-prints [arXiv:1707.01559]
- Leloudas, G., Fraser, M., Stone, N. C., et al. 2016, Nature Astronomy, 1, 0002
- Leloudas, G., Schulze, S., Krühler, T., et al. 2015, MNRAS, 449, 917
- Lunnan, R., Chornock, R., Berger, E., et al. 2014, ApJ, 787, 138
- Lusso, E., Comastri, A., Simmons, B. D., et al. 2012, MNRAS, 425, 623
- Lyman, J. D., Levan, A. J., Tanvir, N. R., et al. 2017, MNRAS, 467, 1795
- Margutti, R., Metzger, B. D., Chornock, R., et al. 2017, ApJ, 836, 25
- Mazzali, P. A., Sullivan, M., Pian, E., Greiner, J., & Kann, D. A. 2016, MNRAS, 458, 3455
- Melchior, P., Drlica-Wagner, A., Bechtol, K., et al. 2015, The Astronomer's Telegram, 7843
- Merloni, A., Dwelly, T., Salvato, M., et al. 2015, MNRAS, 452, 69
- Metzger, B. D., Margalit, B., Kasen, D., & Quataert, E. 2015, MNRAS, 454, 3311
- Metzger, B. D. & Stone, N. C. 2016, MNRAS, 461, 948
- Modigliani, A., Goldoni, P., Royer, F., et al. 2010, in SPIE Conf. Series, Vol. 7737, SPIE Conf. Series
- Moriya, T. J., Tanaka, M., Morokuma, T., & Ohsuga, K. 2017, ApJL, in press [arXiv:1706.06855]
- Müller-Sánchez, F., Comerford, J. M., Nevin, R., et al. 2015, ApJ, 813, 103
- Nicholls, B., Holoien, T. W.-S., Stanek, K. Z., et al. 2015, The Astronomer's Telegram, 7642
- Peng, Y.-j., Lilly, S. J., Kovač, K., et al. 2010, ApJ, 721, 193
- Perley, D. A., Quimby, R. M., Yan, L., et al. 2016, ApJ, 830, 13
- Planck Collaboration. 2014, A&A, 571, A16
- Prieto, J. L., Krühler, T., Anderson, J. P., et al. 2016, ApJ, 830, L32
- Prieto, J. L., Shappee, B. J., Dong, S., et al. 2015, The Astronomer's Telegram, 7776
- Quimby, R. M., Kulkarni, S. R., Kasliwal, M. M., et al. 2011, Nature, 474, 487
- Rees, M. J. 1988, Nature, 333, 523
- Schady, P. 2017, Royal Society Open Science, 4, 170304
- Schmitt, H. R., Donley, J. L., Antonucci, R. R. J., et al. 2003, ApJ, 597, 768
- Schulze, S., Krühler, T., Leloudas, G., et al. 2016, MNRAS, submitted [arXiv:1612.05978]
- Shappee, B. J., Prieto, J. L., Grupe, D., et al. 2014, ApJ, 788, 48
- Shen, Y., Liu, X., Greene, J. E., & Strauss, M. A. 2011, ApJ, 735, 48
- Smette, A., Sana, H., Noll, S., et al. 2015, A&A, 576, A77
- Soto, K. T., Lilly, S. J., Bacon, R., Richard, J., & Conseil, S. 2016, MNRAS, 458, 3210
- Stone, N. C. & van Velzen, S. 2016, ApJ, 825, L14
- Sukhbold, T. & Woosley, S. E. 2016, ApJ, 820, L38
- Ulrich, M.-H., Maraschi, L., & Urry, C. M. 1997, ARA&A, 35, 445
- van Velzen, S. 2017, ArXiv e-prints [arXiv:1707.03458]
- van Velzen, S., Anderson, G. E., Stone, N. C., et al. 2016, Science, 351, 62
- van Velzen, S., Farrar, G. R., Gezari, S., et al. 2011, ApJ, 741, 73
- Vernet, J., Dekker, H., D'Odorico, S., et al. 2011, A&A, 536, A105
- Vinkó, J., Yuan, F., Quimby, R. M., et al. 2015, ApJ, 798, 12
- Walt, S. v. d., Colbert, S. C., & Varoquaux, G. 2011, Computing in Science and Engg., 13, 22
- Wegg, C. 2013, PhD thesis, California Institute of Technology
- Wegg, C. & Nate Bode, J. 2011, ApJ, 738, L8
- Weilbacher, P. M., Streicher, O., Urrutia, T., et al. 2014, in ASP Conf. Ser., Vol. 485, Astronomical Data Analysis Software and Systems XXIII, ed. N. Manset & P. Forshay, 451
- Wevers, T., van Velzen, S., Jonker, P. G., et al. 2017, ArXiv e-prints [arXiv:1706.08965]
- Whitaker, K. E., van Dokkum, P. G., Brammer, G., & Franx, M. 2012, ApJ, 754, L29
- Worthey, G. & Ottaviani, D. L. 1997, ApJS, 111, 377
- Yang, Y., Tremonti, C. A., Zabludoff, A. I., & Zaritsky, D. 2006, ApJ, 646, L33
- Zabludoff, A. I., Zaritsky, D., Lin, H., et al. 1996, ApJ, 466, 104