Statistics The IFT School on Cosmology Tools

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► Statistics ≡ Inference

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- ▶ How to extract information about scientific models from data.
- ► Most cosmologists work in a *Bayesian* framework of inference, although *Frequentist* methods are also sometimes used.

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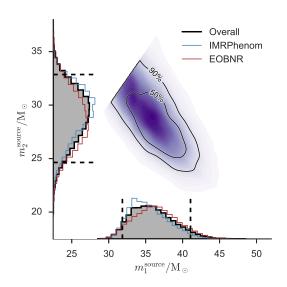
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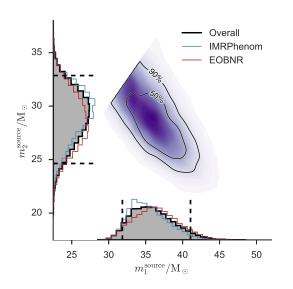
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- ► These are called *credible intervals*, state that we are e.g. 90% confident of the value lying in this range.
- ▶ More importantly, these are *summary statistics*.

LIGO binary merger

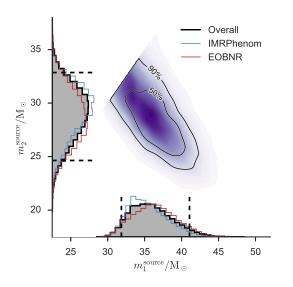


LIGO binary merger



Summary statistics summarise a full probability distribution.

LIGO binary merger



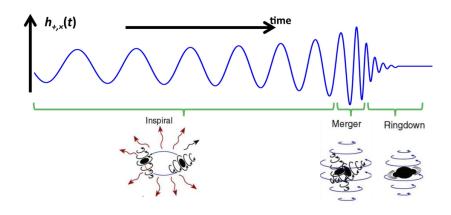
- Summary statistics summarise a full probability distribution.
- One goal of inference is to produce these probability distributions.

Extended example of inference: LIGO

► We will introduce the key concepts by discussing an extended example of the inference process.

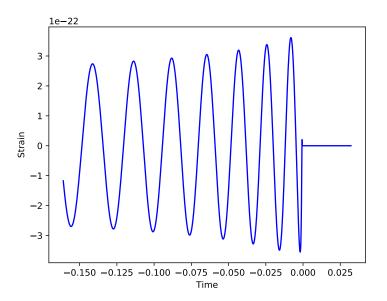
Theory

Extended example of inference: LIGO



The model M

Extended example of inference: LIGO



Extended example of inference: LIGO

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Theoretical signal depends on:

 $ightharpoonup m_1, m_2$: mass of binary

Extended example of inference: LIGO

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Extended example of inference: LIGO

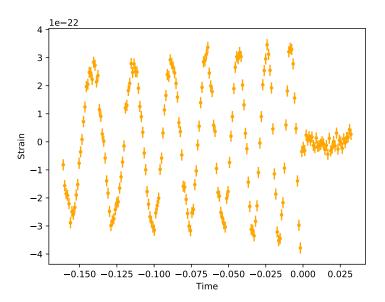
- $ightharpoonup m_1, m_2$: mass of binary
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- $ightharpoonup \Phi_c, t_c$: phase and time of coalescence

Extended example of inference: LIGO

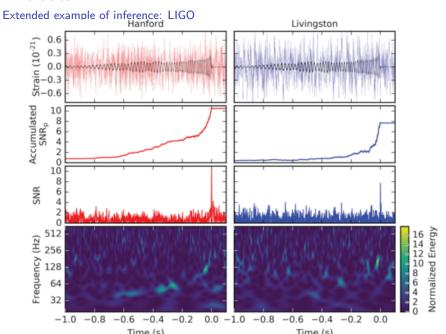
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- $ightharpoonup i, heta_{
 m sky}$: inclination and angle on sky (orbital parameters)

The data D

Extended example of inference: LIGO



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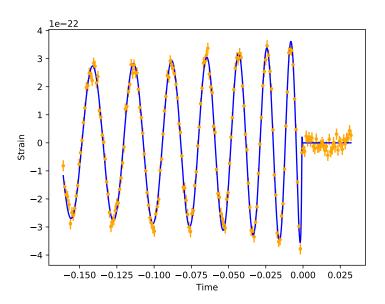
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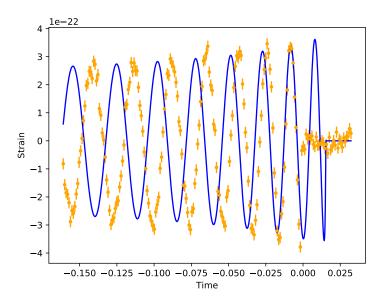
- \blacktriangleright $(t_i, h_i \pm \sigma_i)$: strain observed
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- ▶ We normally work with log-likelihoods, which turn $\prod \rightarrow \sum$.

The Likelihood: well matched

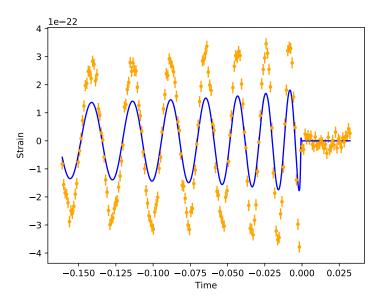
Extended example of inference: LIGO



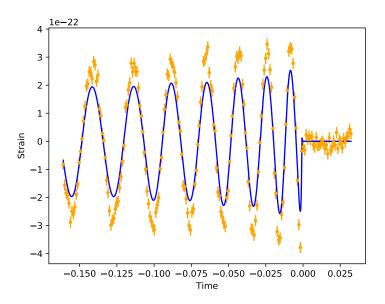
The Likelihood: coalescence off



The Likelihood: too large luminosity distance



The Likelihood: incorrect inclination



Extended example of inference: LIGO

► Likelihood

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▶ Likelihood ≡ Probability of data, given model parameters:

$$L(\Theta) \equiv P(D|\Theta, M)$$

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Likelihood × Prior

$$Posterior = \frac{Likelihood \times Prior}{Evidence}$$

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- ▶ Most Bayesian approaches are sensitive to this, and rightly so.

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- Still extremely important.

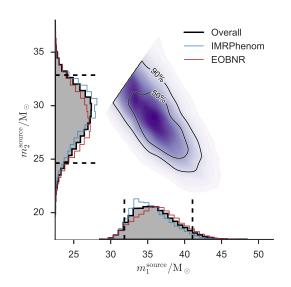
Extended example of inference: LIGO

Cannot plot the full posterior distribution:

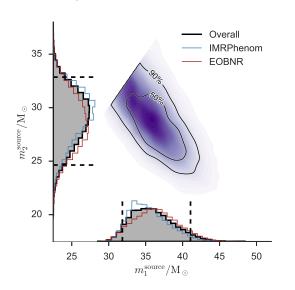
$$\mathcal{P}(\Theta) \equiv P(m_1, m_2, \theta, \phi, r, \Phi_c, t_c, i, \theta_{\rm sky} | D, M)$$

► Can plot 1D and 2D *marginalised* distributions e.g:

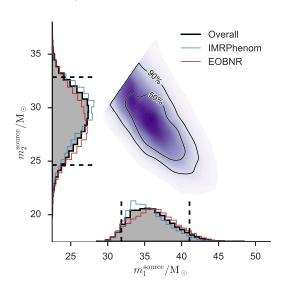
$$\begin{split} P(m_1, m_2 | D, M) &= \\ \int P(m_1, m_2, \theta, \phi, r, \Phi_c, t_c, i, \theta_{\rm sky} | D, M) \, d\theta \, d\phi \, dr \, d\Phi_c \, dt_c \, di \, d\theta_{\rm sky} \end{split}$$



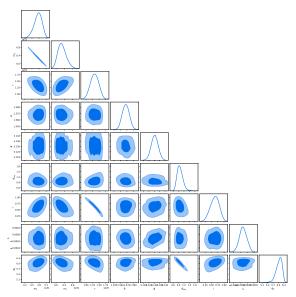
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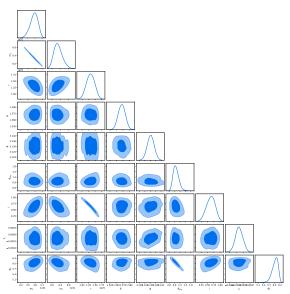
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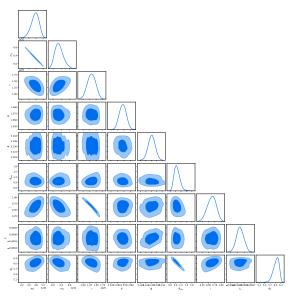
- May do this for each pair of parameters
- Generates a triangle plot



Extended example of inference: LIGO



Does give insight



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- Not the full picture

Extended example of inference: LIGO

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- Scientifically speaking, this is only half the story.
- ► In general, we will have several competing models that describe the data, and we want to know which is the "best".



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What does data tell us about our model M_i ?

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Model averaging:

Multiple models with posterior on the same parameter: $P(y|M_i, D)$

$$P(y|D) = \sum_{i} P(y|M_i, D)P(M_i|D)$$

$$\mathcal{L}(\Theta) = P(D|\Theta, M)$$

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$$\Theta_{\Lambda\mathrm{CDM}} = (\Omega_b h^2, \Omega_c h^2, 100\theta_{MC}, \tau, \ln(10^{10} A_s), n_s)$$

Another example.

$$D = \{C_{\ell}^{(\text{Planck})}\}$$

$$M = \Lambda \text{CDM}$$

$$\Theta = \Theta_{\Lambda \text{CDM}}$$

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 $\mathcal{L}(\Theta) = P(D|\Theta, M)$

Another example.

$$\begin{split} D = & \{C_{\ell}^{(\text{Planck})}\} \\ M = & \Lambda \text{CDM} \\ \Theta = & \Theta_{\Lambda\text{CDM}} + \Theta_{\text{Planck}} \\ \Theta_{\Lambda\text{CDM}} = & (\Omega_b h^2, \Omega_c h^2, 100\theta_{MC}, \tau, \ln(10^{10} A_s), n_s) \end{split}$$

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$$\mathcal{L}(\Theta) = P(D|\Theta, M)$$

$$D = \{C_{\ell}^{(\text{Planck})}\}$$

$$M = \Lambda \text{CDM} + \text{extensions}$$

$$\Theta = \Theta_{\Lambda \text{CDM}} + \Theta_{\text{Planck}}$$

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$$\begin{split} \mathcal{L}(\Theta) &= P(D|\Theta, M) \\ D &= \{C_{\ell}^{(\text{Planck})}\} \\ M &= & \text{ACDM} + \text{extensions} \\ \Theta &= & \Theta_{\Lambda\text{CDM}} + \Theta_{\text{Planck}} + \Theta_{\text{extensions}} \\ \Theta_{\Lambda\text{CDM}} &= & (\Omega_b h^2, \Omega_c h^2, 100\theta_{MC}, \tau, \ln(10^{10}A_s), n_s) \\ \Theta_{\text{Planck}} &= & (y_{\text{cal}}, A_{217}^{ClB}, \xi^{tSZ-ClB}, A_{143}^{tSZ}, A_{100}^{PS}, A_{143 \times 217}^{PS}, A_{217}^{PS}, A_{100}^{tSZ}, A_{100}^{dust TT}, A_{143 \times 217}^{dust TT}, A_{217}^{dust TT}, C_{100}, C_{217}) \\ \Theta_{\text{extensions}} &= & (n_{\text{run}}) \end{split}$$

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$$\begin{split} \mathcal{L}(\Theta) &= P(D|\Theta, M) \\ D &= \{C_{\ell}^{(\text{Planck})}\} \\ M &= \Lambda \text{CDM} + \text{extensions} \\ \Theta &= \Theta_{\Lambda \text{CDM}} + \Theta_{\text{Planck}} + \Theta_{\text{extensions}} \\ \Theta_{\Lambda \text{CDM}} &= (\Omega_b h^2, \Omega_c h^2, 100\theta_{MC}, \tau, \ln(10^{10} A_s), n_s) \\ \Theta_{\text{Planck}} &= (y_{\text{cal}}, A_{217}^{CIB}, \xi^{tSZ-CIB}, A_{143}^{tSZ}, A_{100}^{PS}, A_{143}^{PS}, A_{143 \times 217}^{PS}, A_{217}^{PS}, \\ A^{kSZ}, A_{100}^{\text{dust}\,TT}, A_{143}^{\text{dust}\,TT}, A_{143 \times 217}^{\text{dust}\,TT}, A_{217}^{\text{dust}\,TT}, c_{100}, c_{217}) \\ \Theta_{\text{extensions}} &= (n_{\text{run}}, n_{\text{run,run}}, w) \end{split}$$

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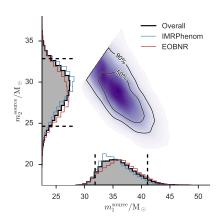
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- ▶ Parameter estimation: $L, \pi \to \mathcal{P}$: model parameters
- ▶ Model comparison: $L, \pi \to Z$: how good model is

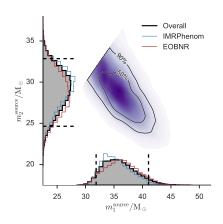
How to describe a high-dimensional posterior

How to describe a high-dimensional posterior



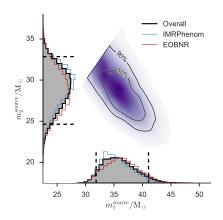
How to describe a high-dimensional posterior

In high dimensions, posterior P occupies a vanishingly small region of the prior π.



How to describe a high-dimensional posterior

- In high dimensions, posterior P occupies a vanishingly small region of the prior π.
- Sampling the posterior is an excellent compression scheme.



Marginalisation over the posterior

▶ Set of *N* samples $S = \{\Theta^{(i)} : i = 1, ... N : \Theta^{(i)} \sim \mathcal{P}\}$

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Parameter estimation

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- 3. Ensemble sampling (e.g. emcee).

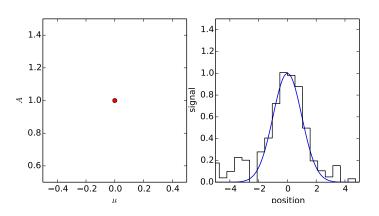
► Turn the *N*-dimensional problem into a one-dimensional one.

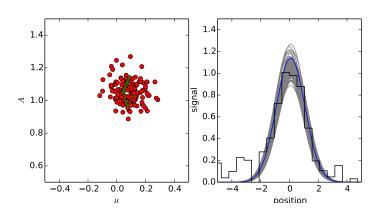
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 - 4. ... otherwise sometimes make step.





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 - ▶ Otherwise, reject, $w^{(i)} + = 1$ and repeat.

Struggles with...

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- stan is a fully fledged, rapidly developing programming language with HMC as a default sampler.

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- Can use information present in ensemble to guide proposals.
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- emcee is not the only (or even best) affine invariant approach.

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$$= \int P(D|\Theta, M)P(\Theta|M)d\Theta$$

$$= \langle \mathcal{L} \rangle_{\pi}$$

- MCMC fundamentally explores the posterior, and cannot average over the prior.
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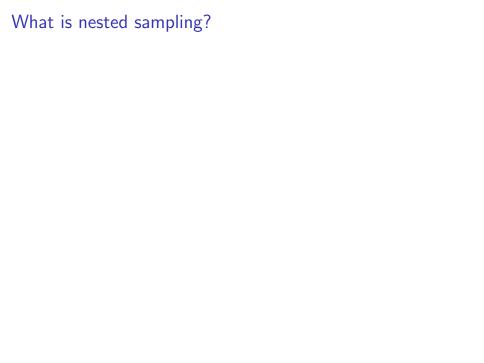
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- Uses ensemble sampling to compress prior to posterior.
- ▶ In doing so, it circumvents many issues (dimensionality, topology, geometry) that beset standard approaches.

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Requires one to be able to uniformly within a region, subject to a hard likelihood constraint.

Graphical aid

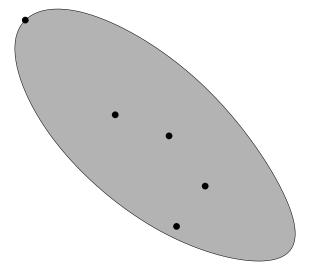
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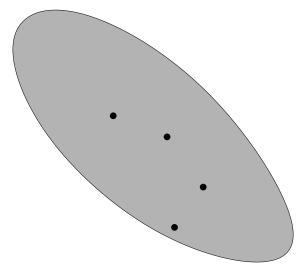
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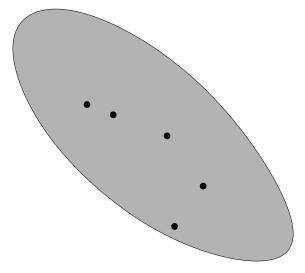
Nested Sampling Graphical aid

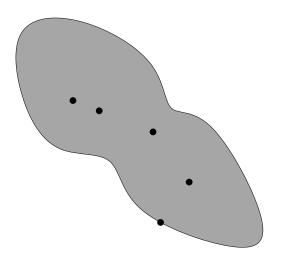
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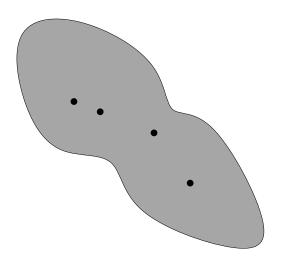
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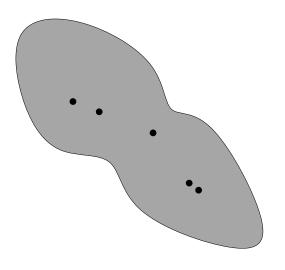


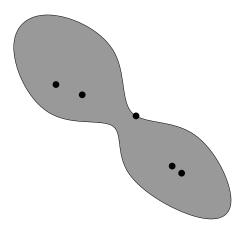


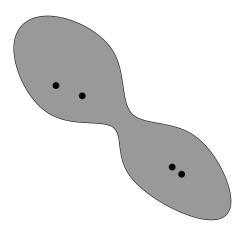


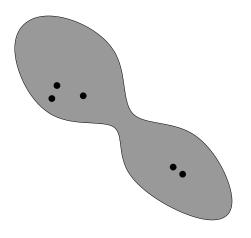


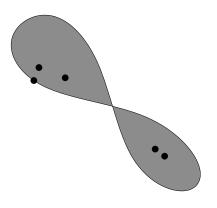


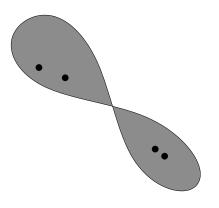


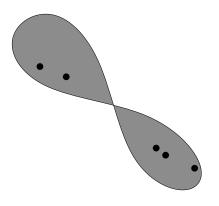


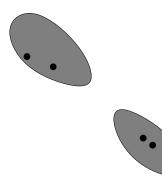


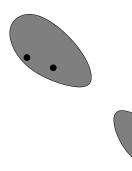


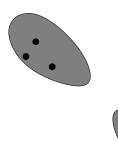


























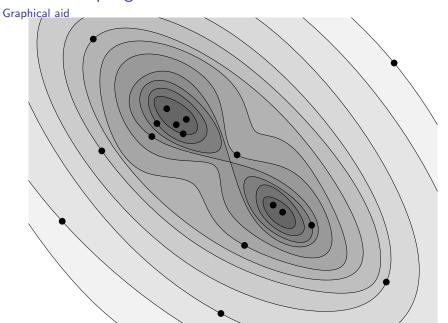






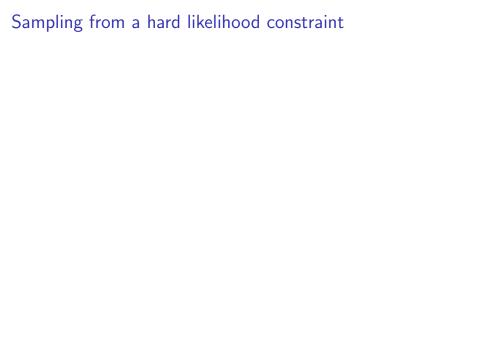






► The set of dead points are posterior samples with an appropriate weighting factor

- ➤ The set of dead points are posterior samples with an appropriate weighting factor
- They can also be used to calculate evidences, since it sequentially updates the priors.



Sampling from a hard likelihood constraint

"It is not the purpose of this introductory paper to develop the technology of navigation within such a volume. We merely note that exploring a hard-edged likelihood-constrained domain should prove to be neither more nor less demanding than exploring a likelihood-weighted space."

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► Most of the work in NS to date has been in attempting to implement a hard-edged sampler in the NS meta-algorithm.

Sampling within an iso-likelihood contour

Previous attempts

Rejection Sampling MultiNest; F. Feroz & M. Hobson (2009).

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Suffers in high dimensions

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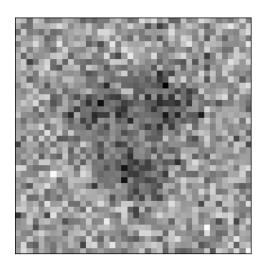
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- Current "state-of-the-art".
- PolyChord 2.0 imminent.

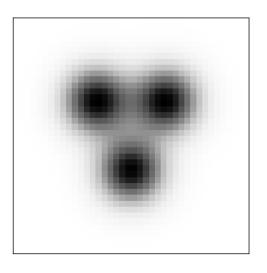
Object detection

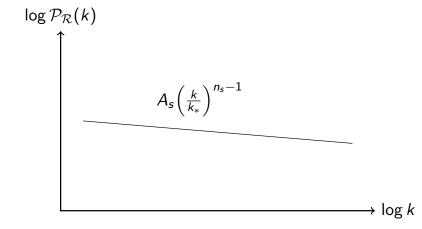
Toy problem

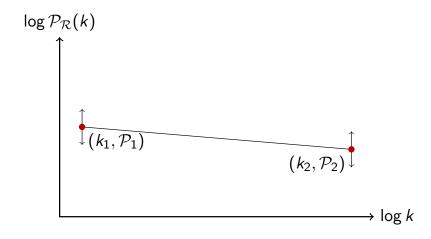


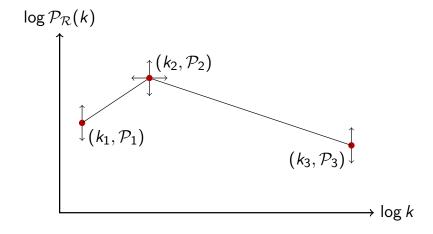
Object detection

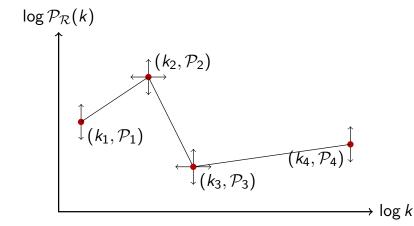
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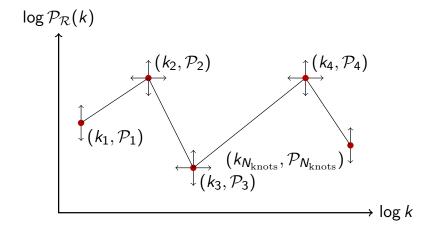


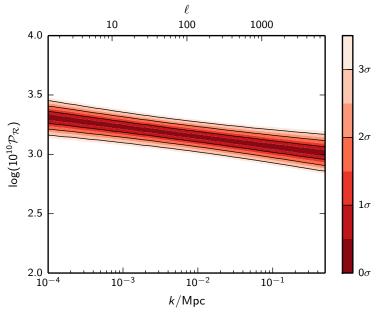


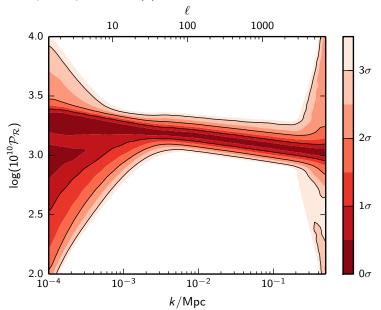


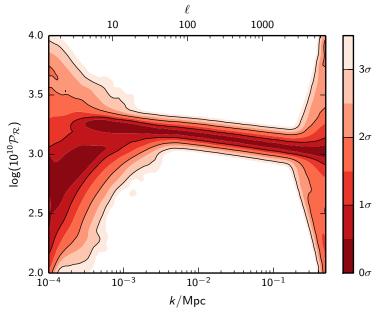


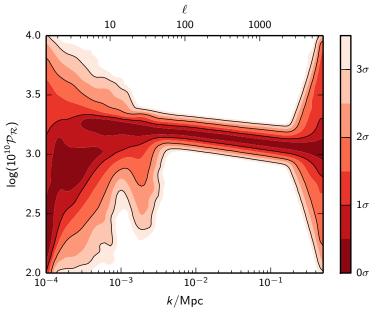


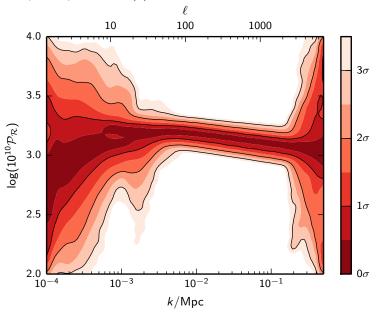


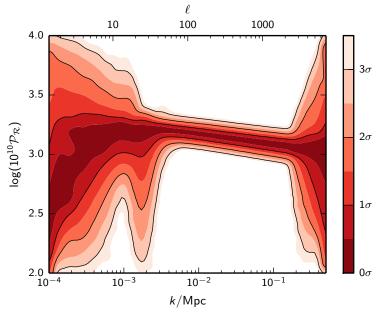


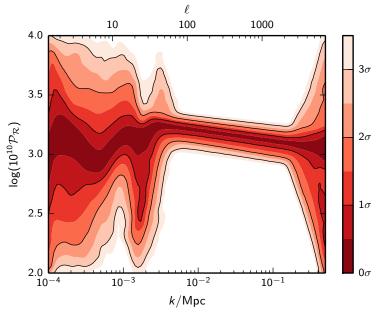


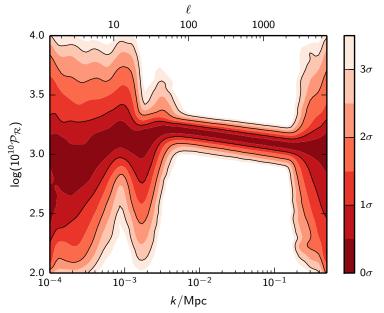


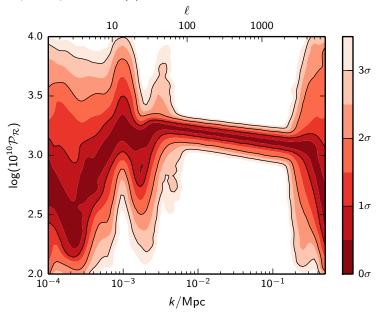




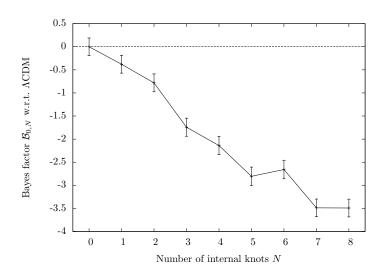




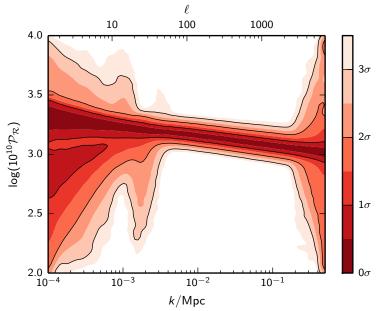




Bayes Factors



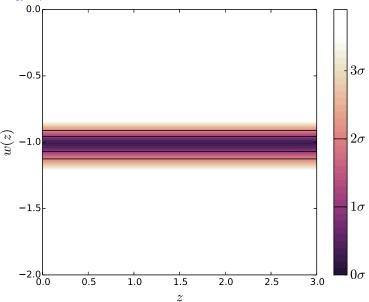
Marginalised plot



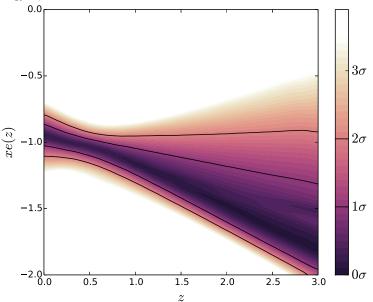
Same thing, but for Dark energy equation of state w(z) (quintessence).

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- ▶ Data used is Planck 2015, BOSS DR 11, JLA supernovae and BOSS Ly α data

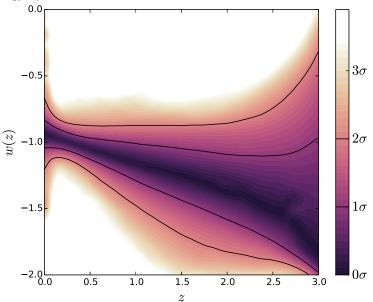
Flat, variable w



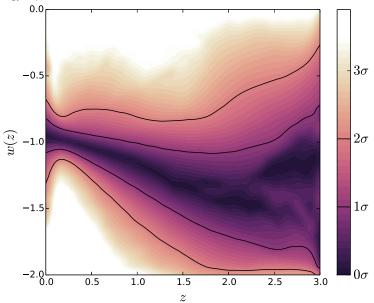
Tilted



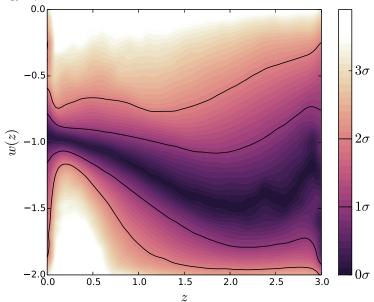
1 internal node



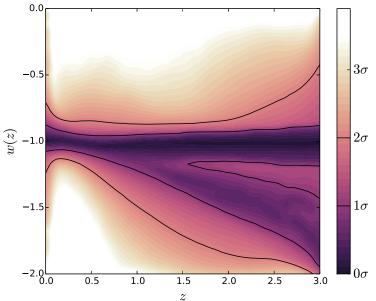
2 internal nodes



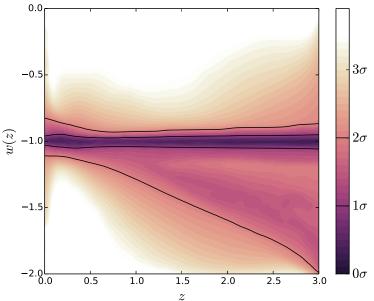
3 internal nodes



Marginalised plot - just extension models

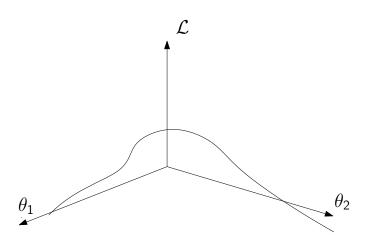


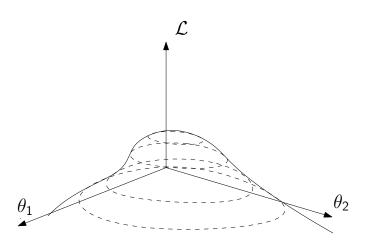
Marginalised plot - including LCDM

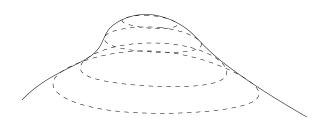


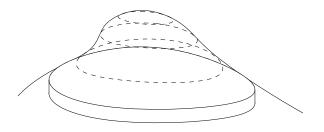
Nested Sampling

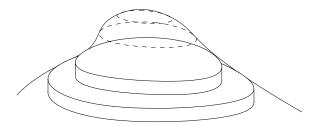
Calculating evidences

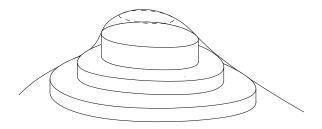


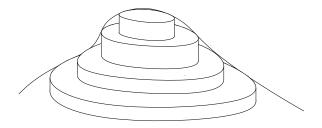


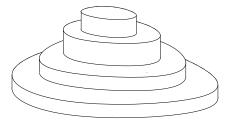


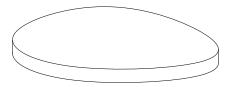




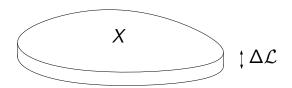


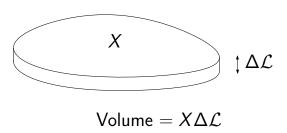


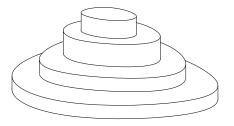


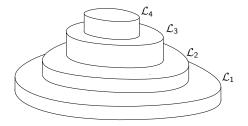


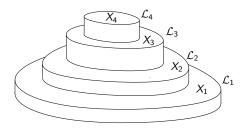


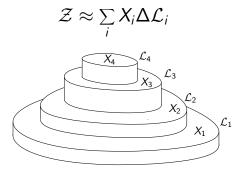












Exponential volume contraction

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$$X_{i+1} \approx \frac{n}{n+1} X_i, \qquad X_0 = 1 \tag{2}$$