



# Cosmologie

L. Vacher -

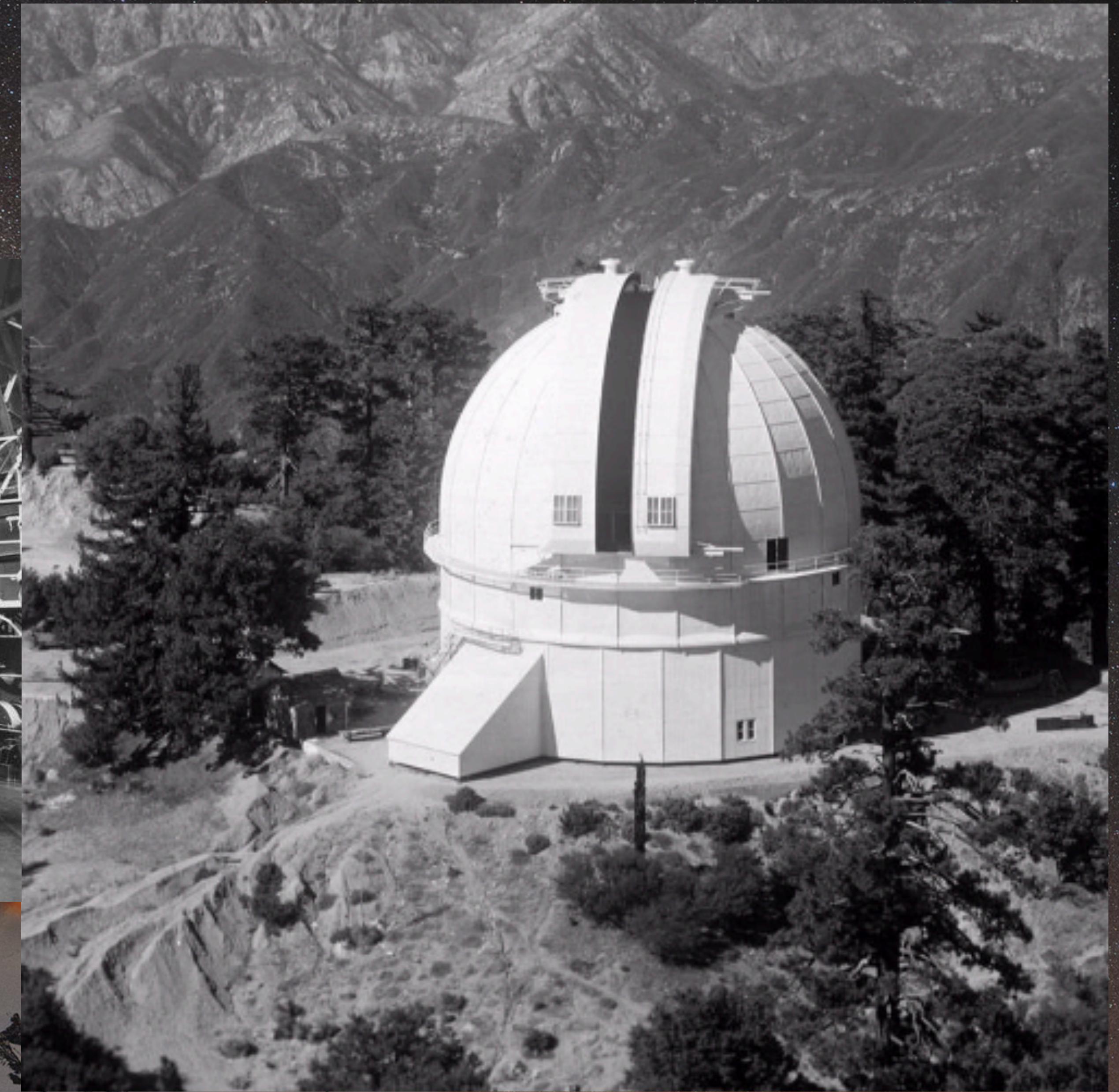
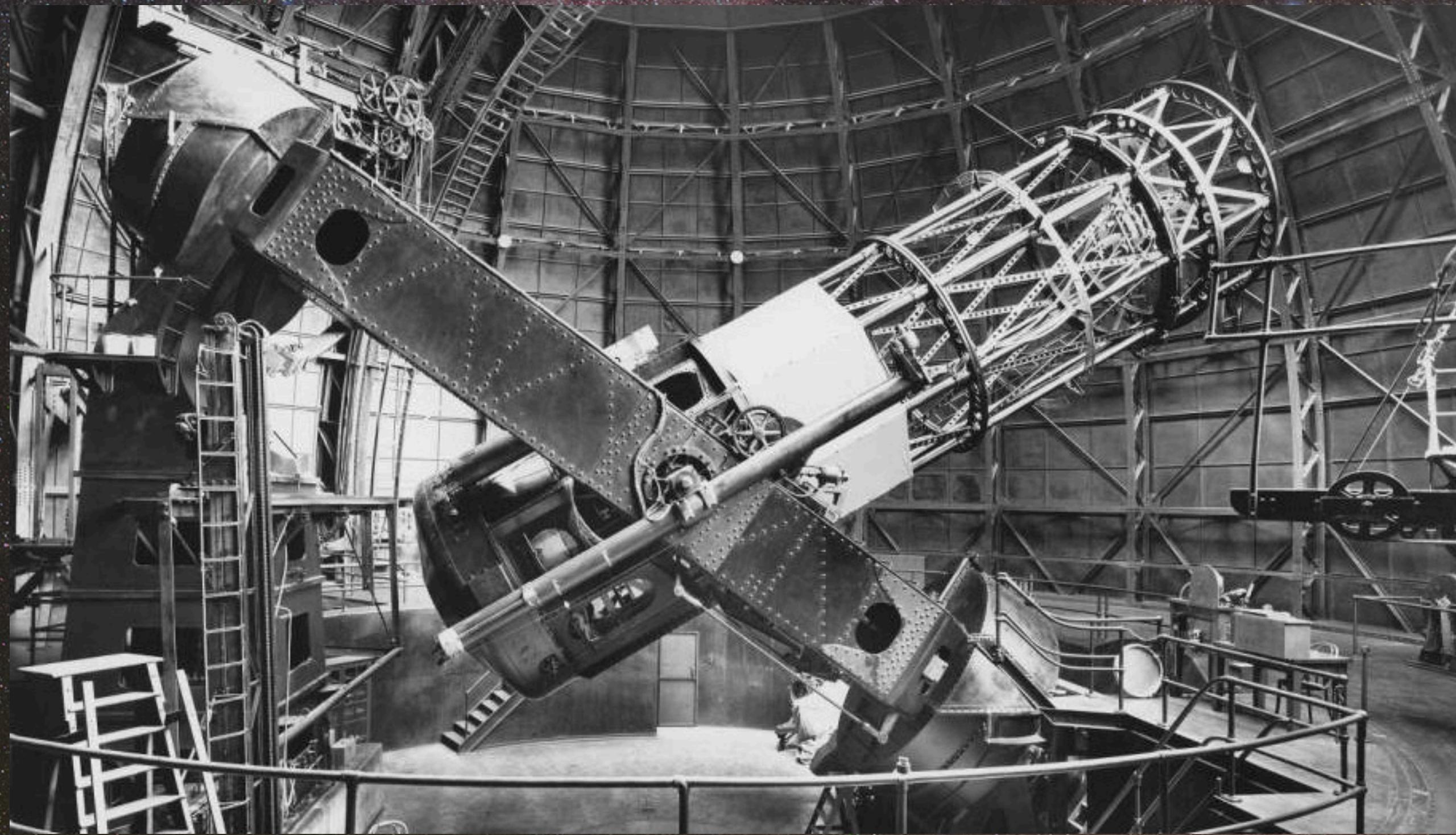


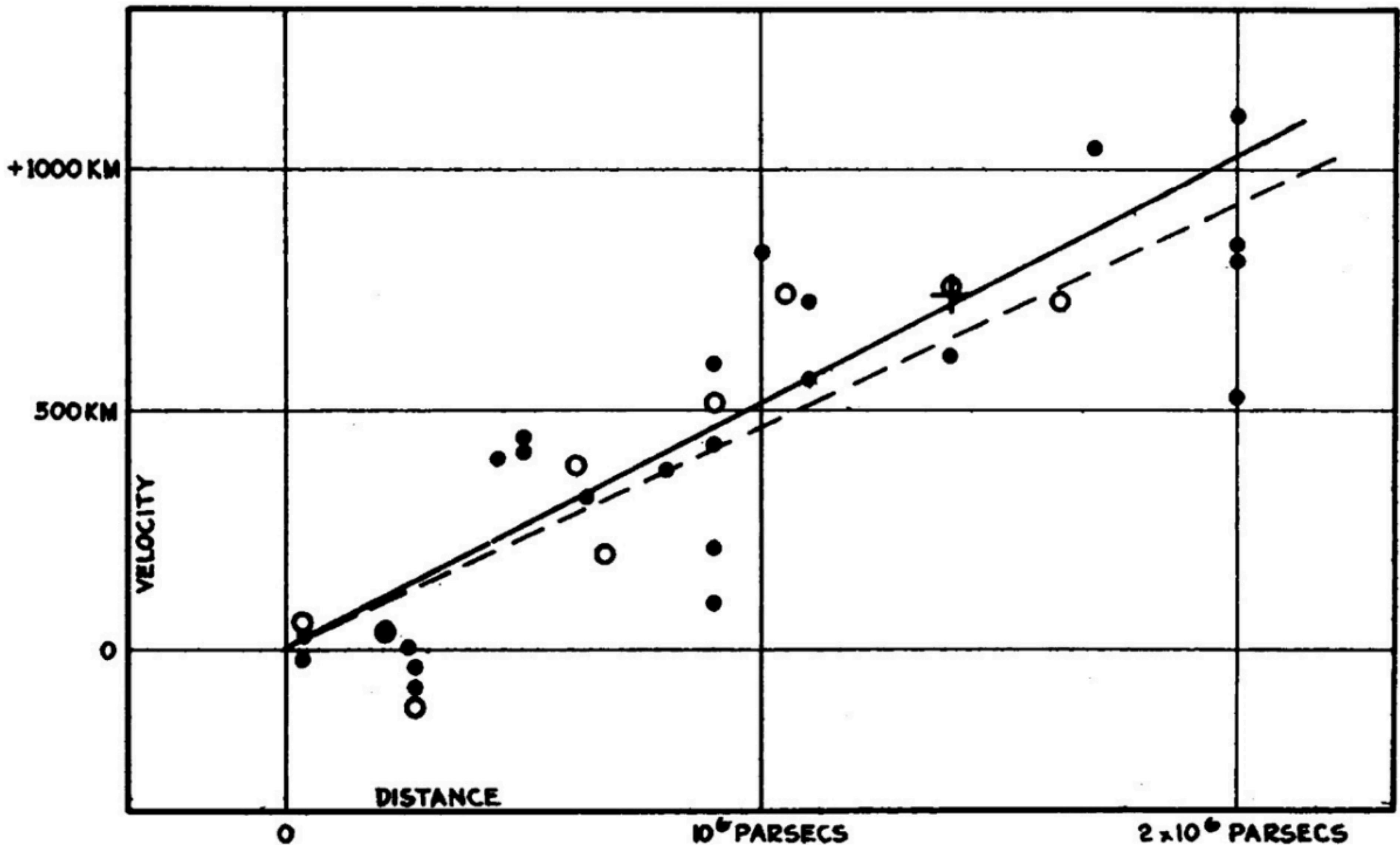
Edwin Hubble



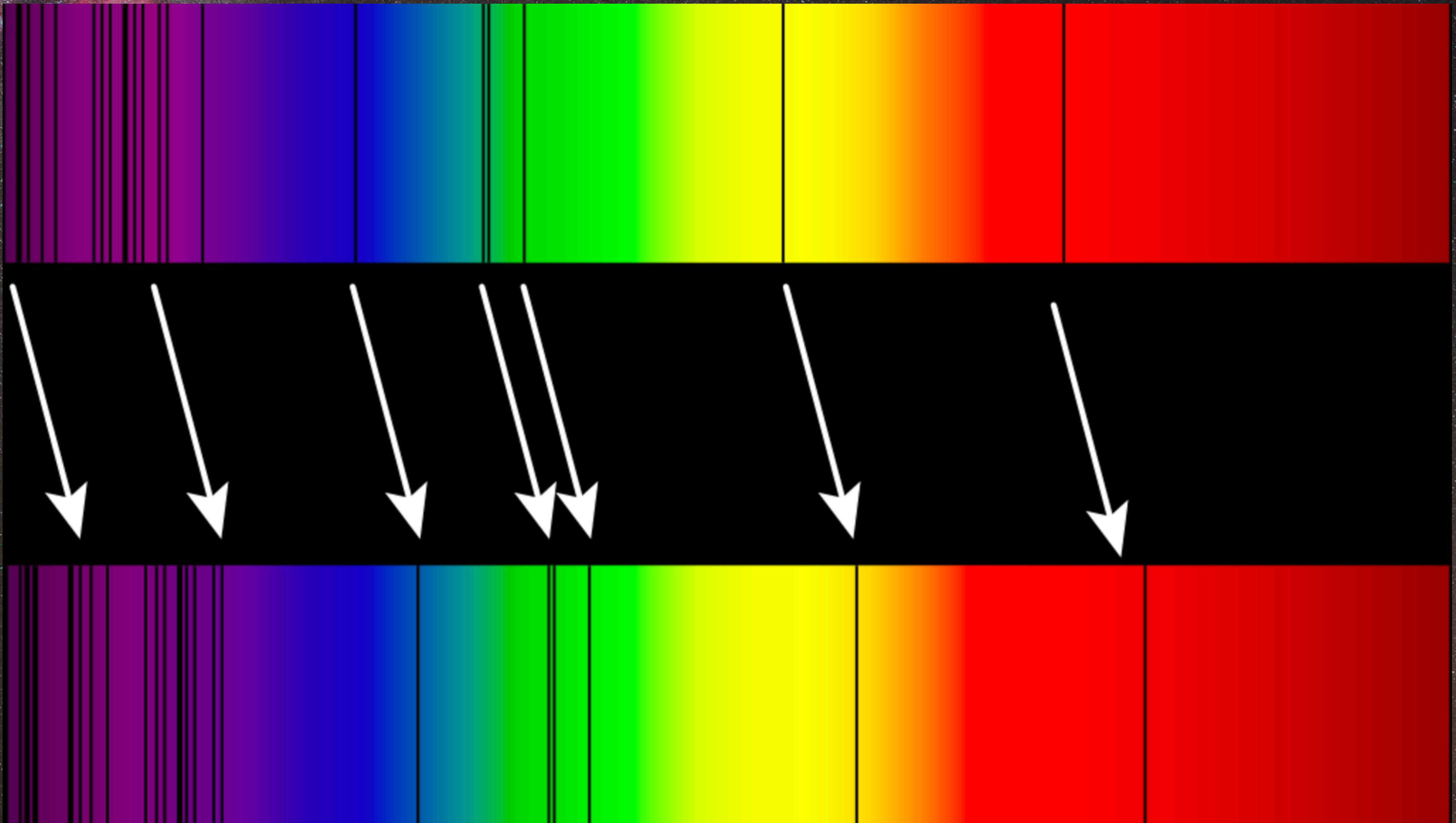
Milton Humason





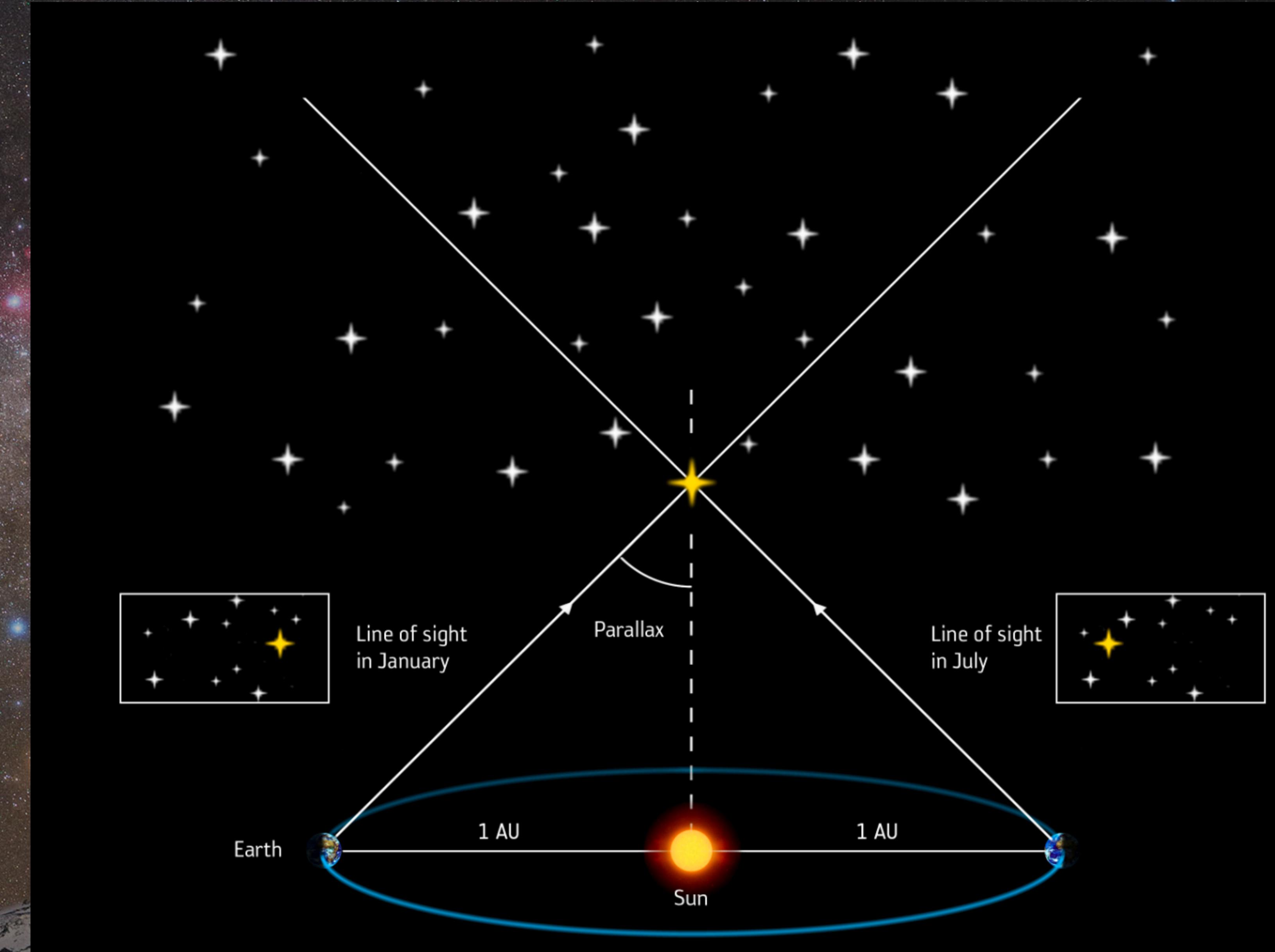


# Vitesse?



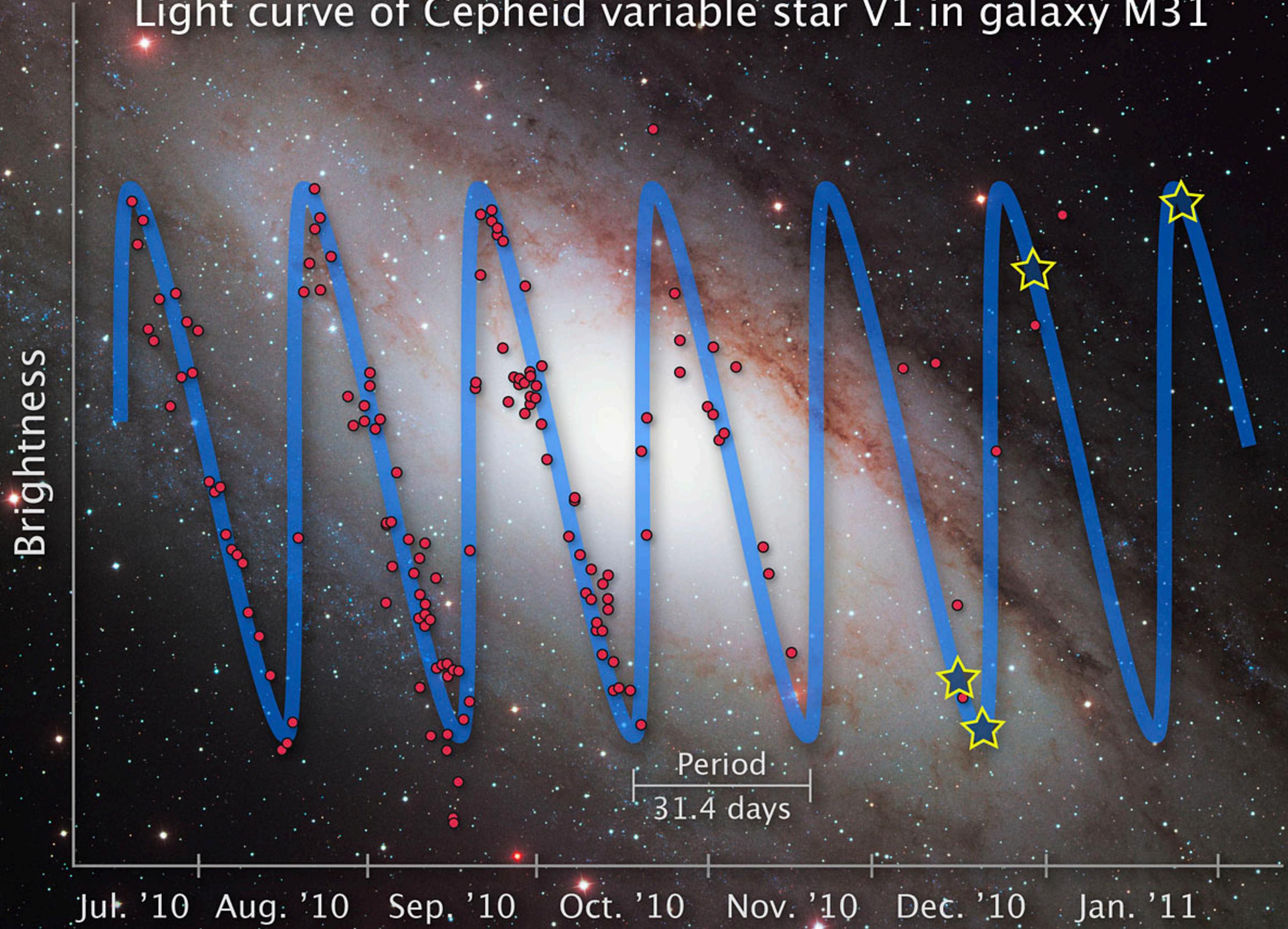
# Distance?

Parallaxes ?

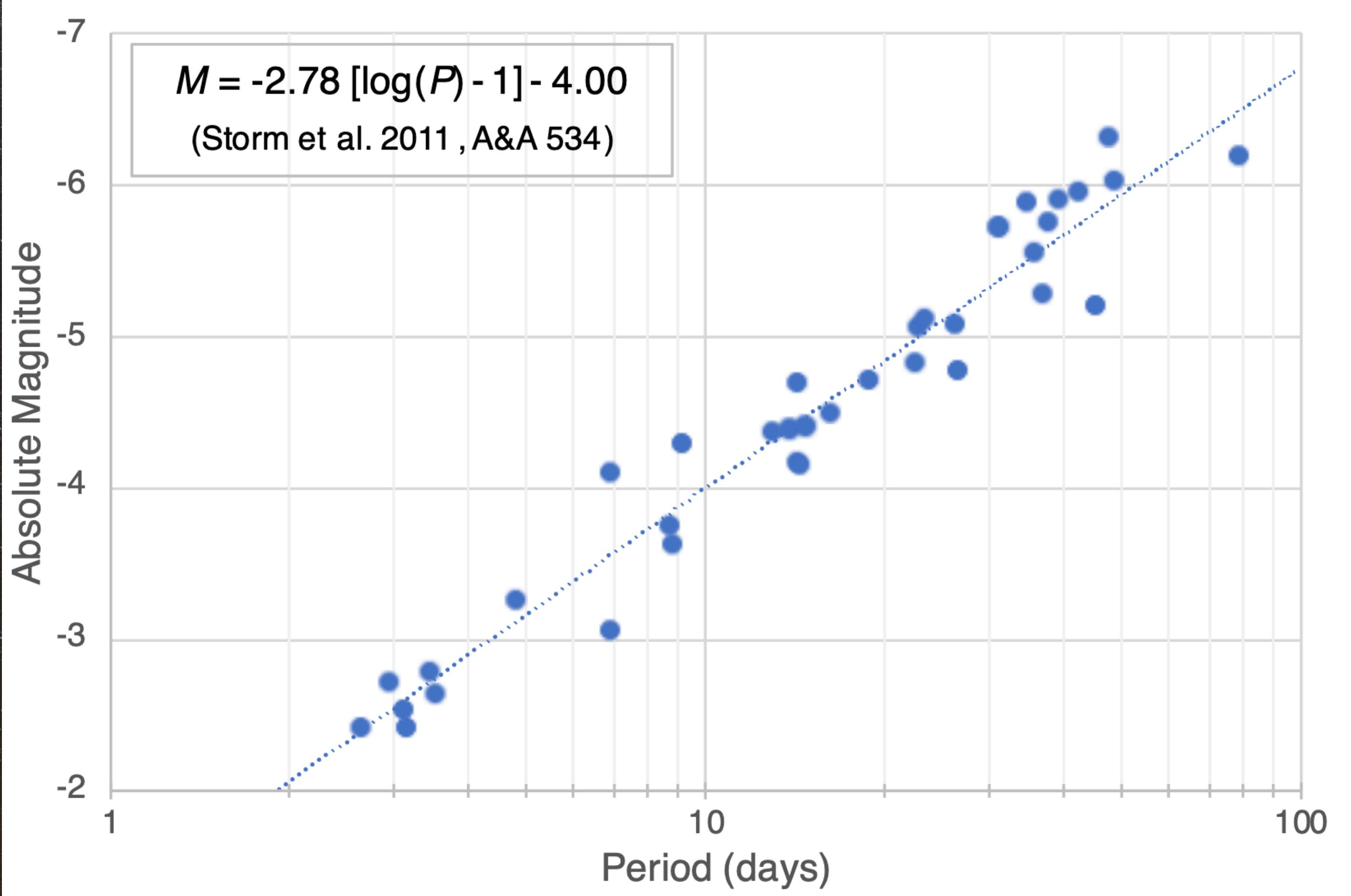


# Distance? Céphéides

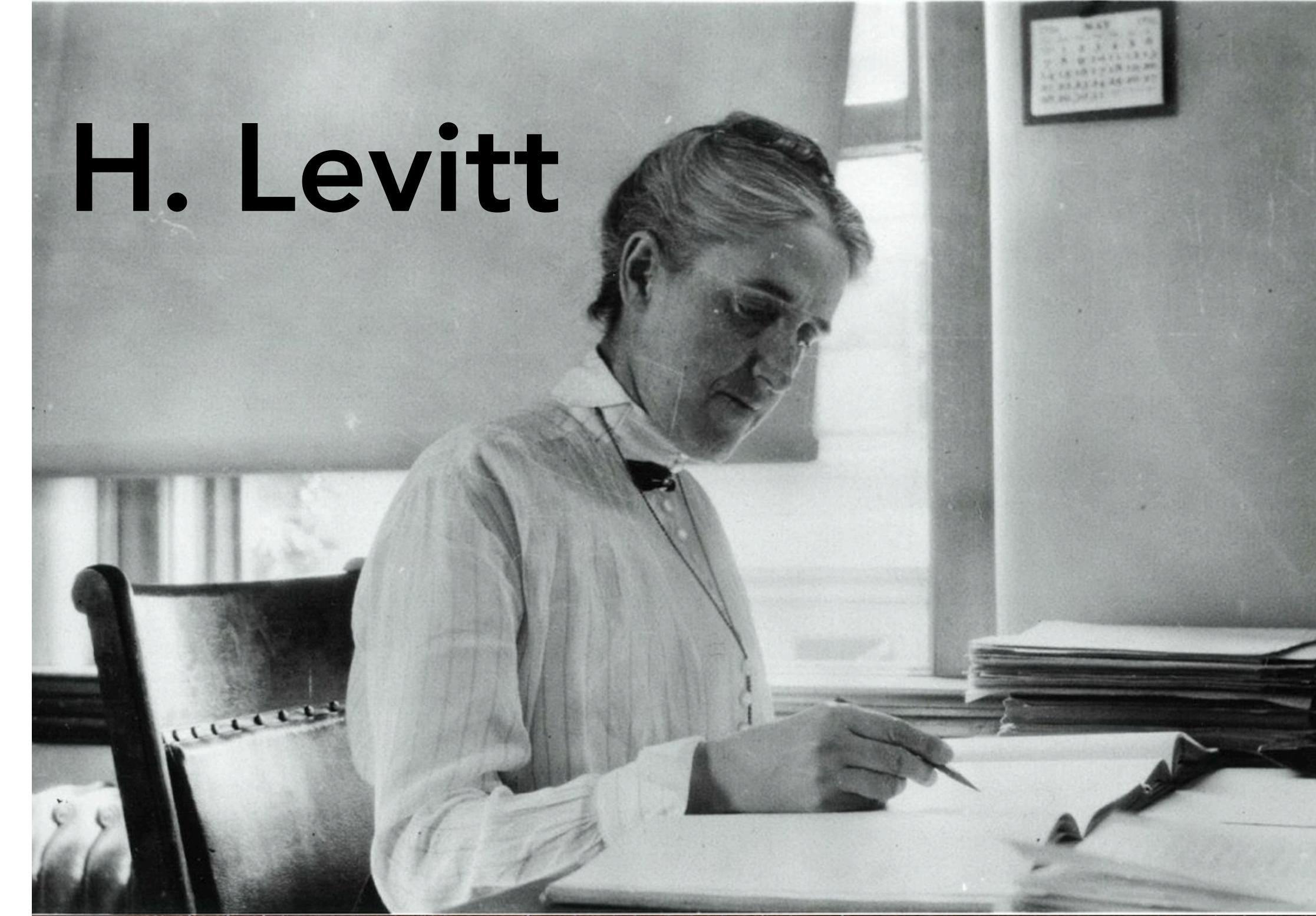
Light curve of Cepheid variable star V1 in galaxy M31



# Distance? Céphéides



H. Levitt

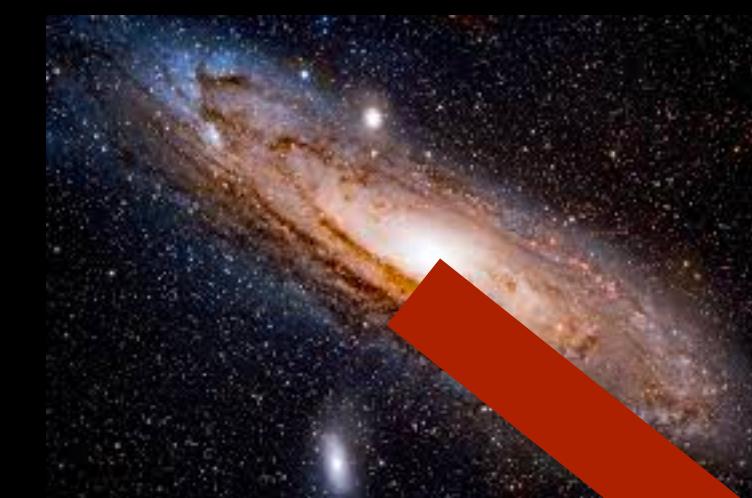
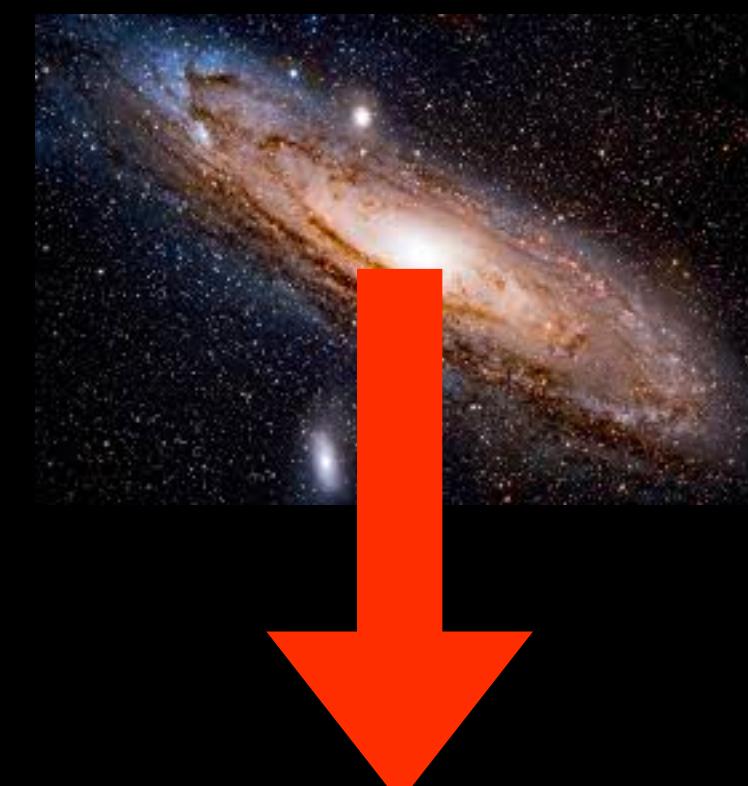
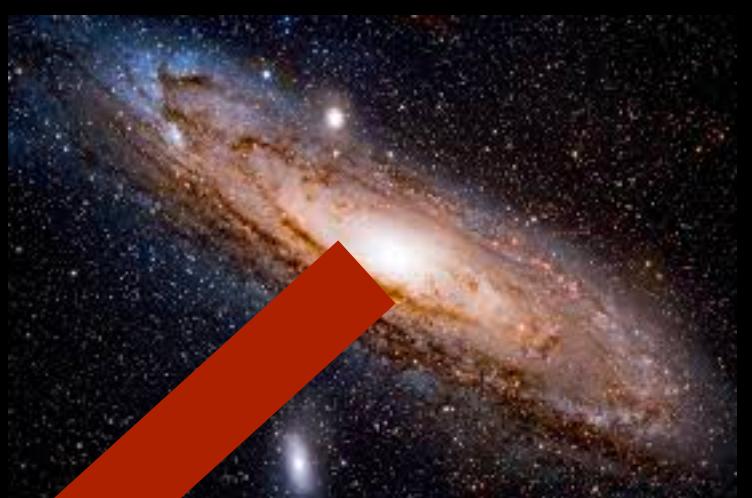
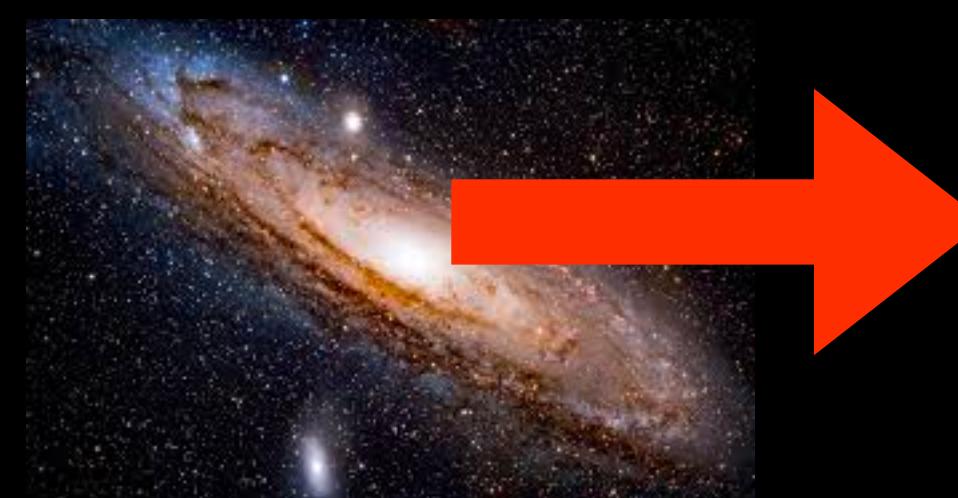
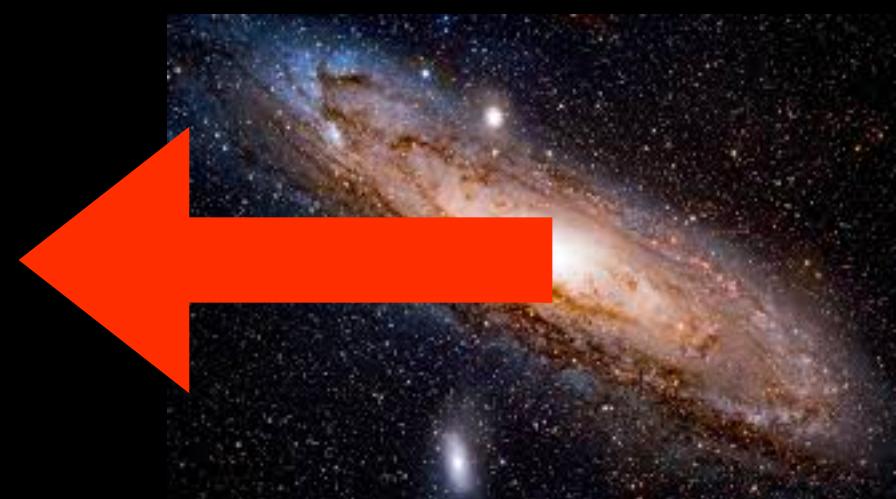
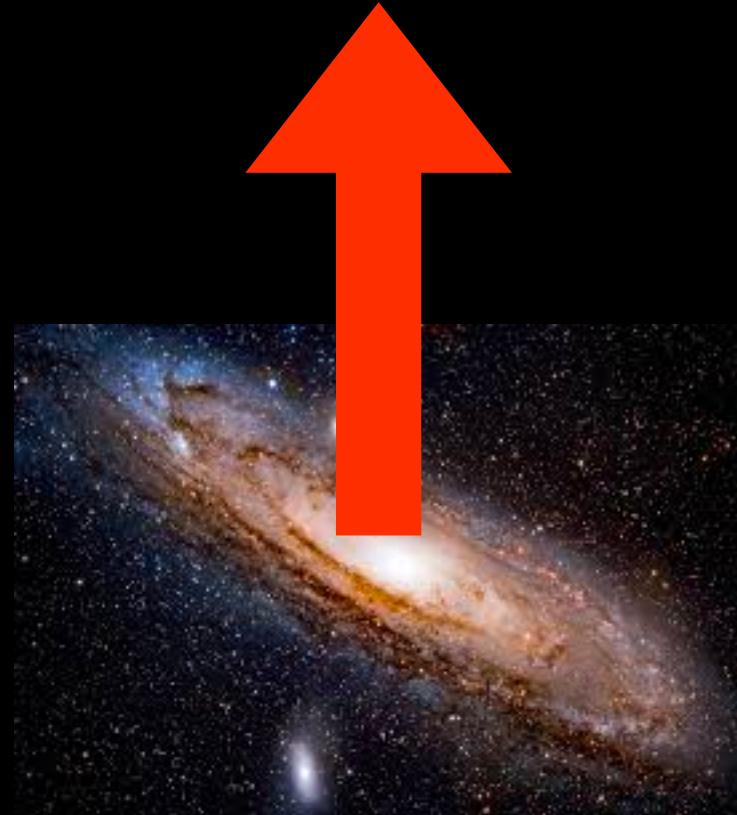
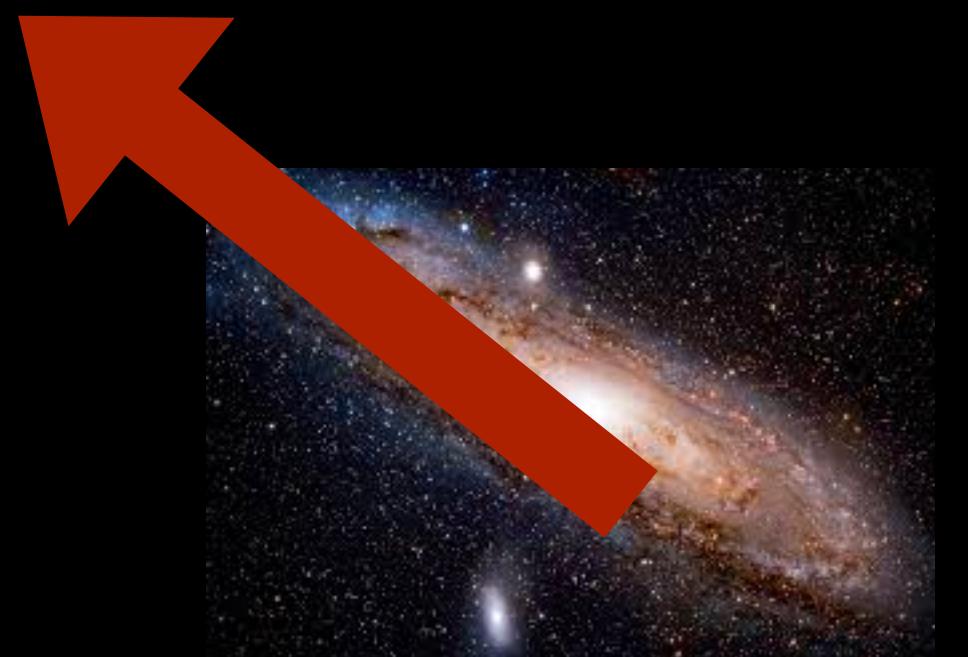




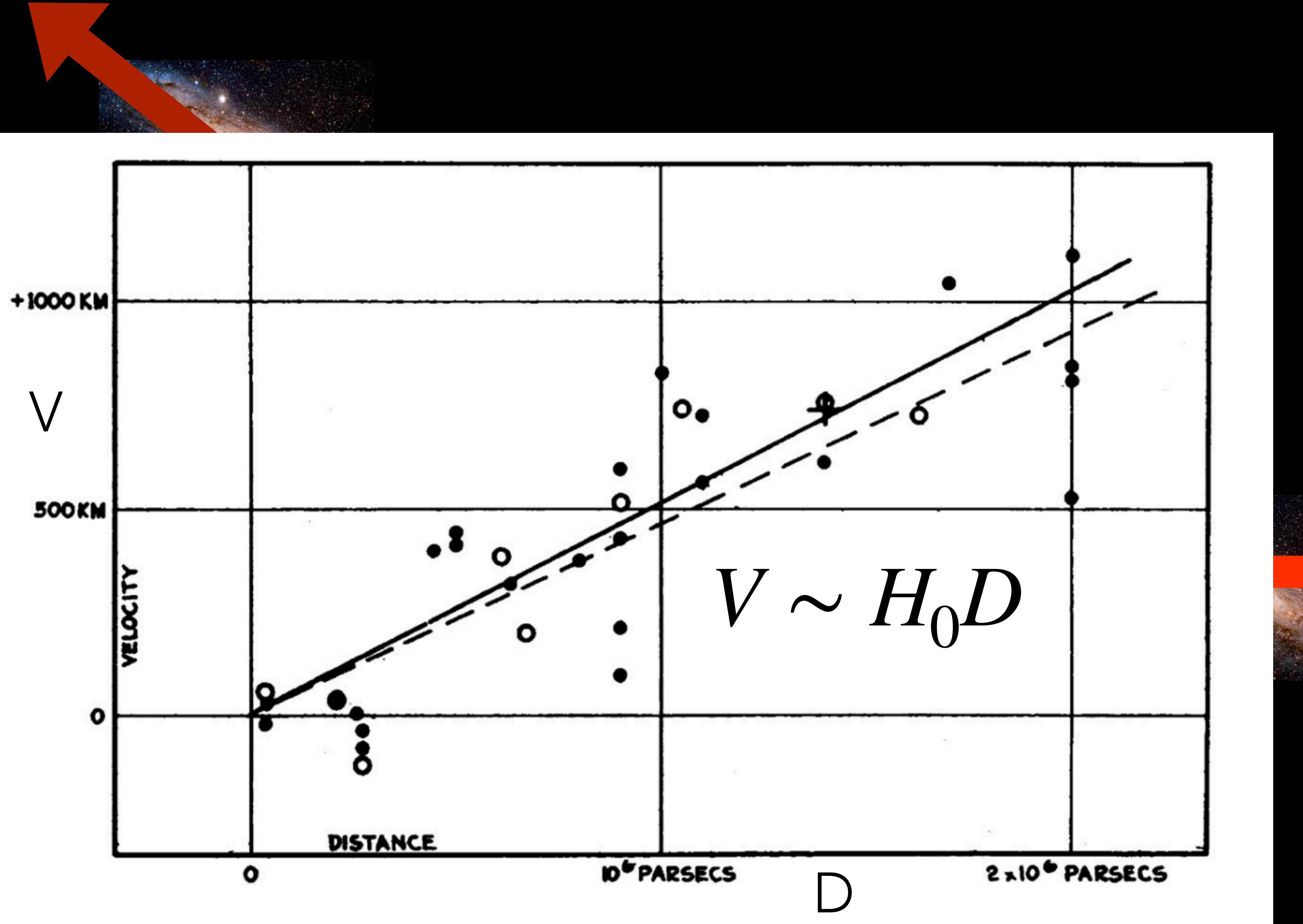
Première nouvelle: « nébuleuses » vs « galaxies »



Première nouvelle: « nébuleuses » vs « galaxies » (The Hubble extremely deep field)



Deuxième nouvelle: Les galaxies s'éloignent



Deuxième nouvelle: Les galaxies s'éloignent

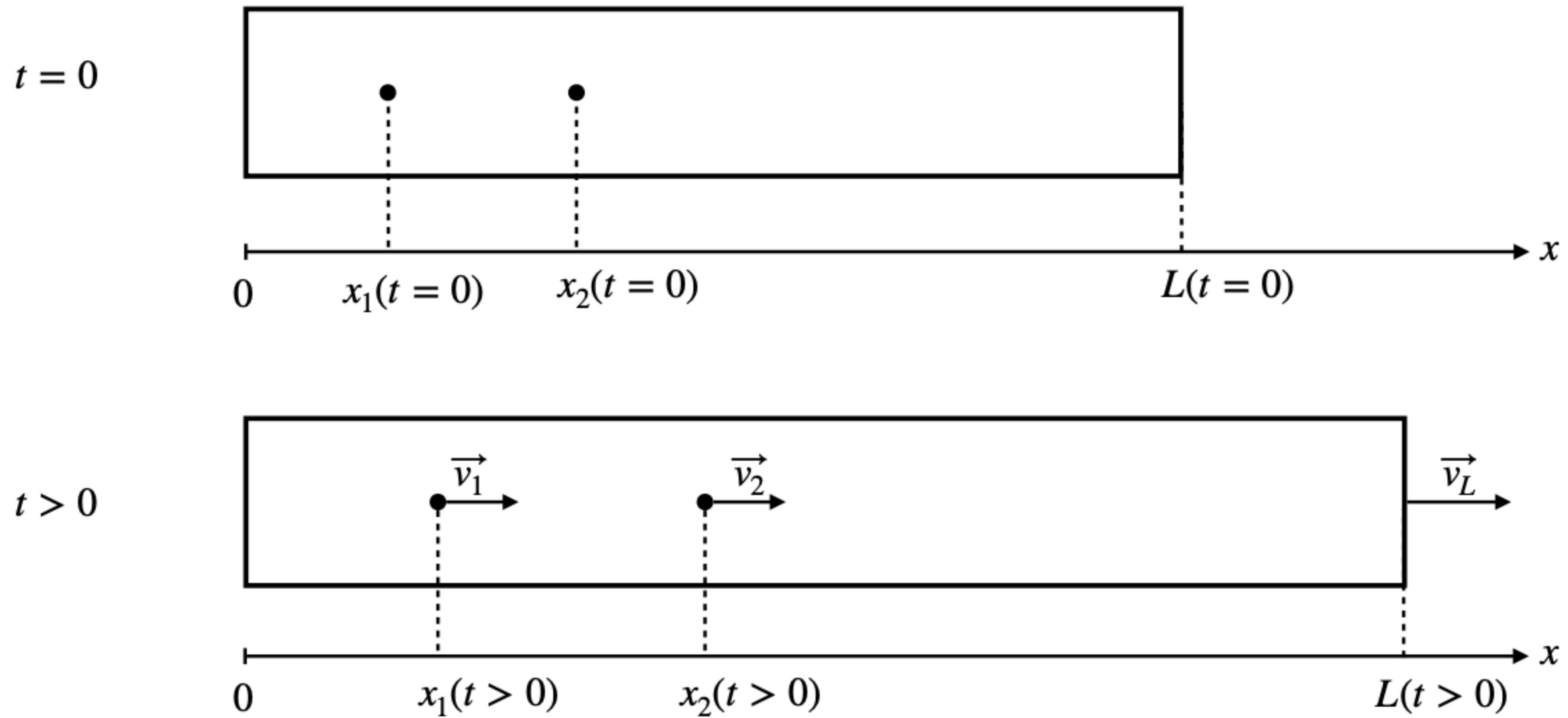
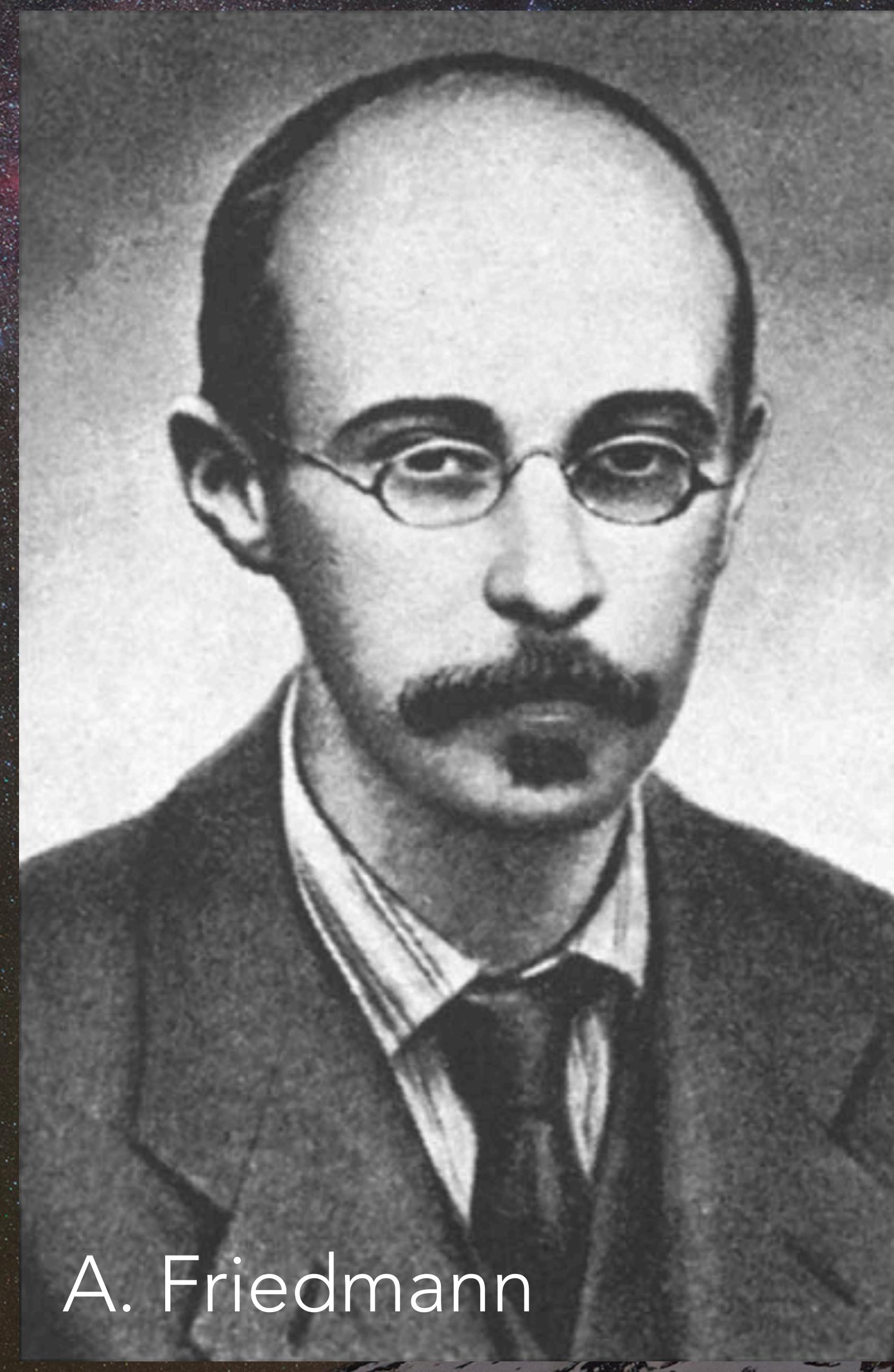
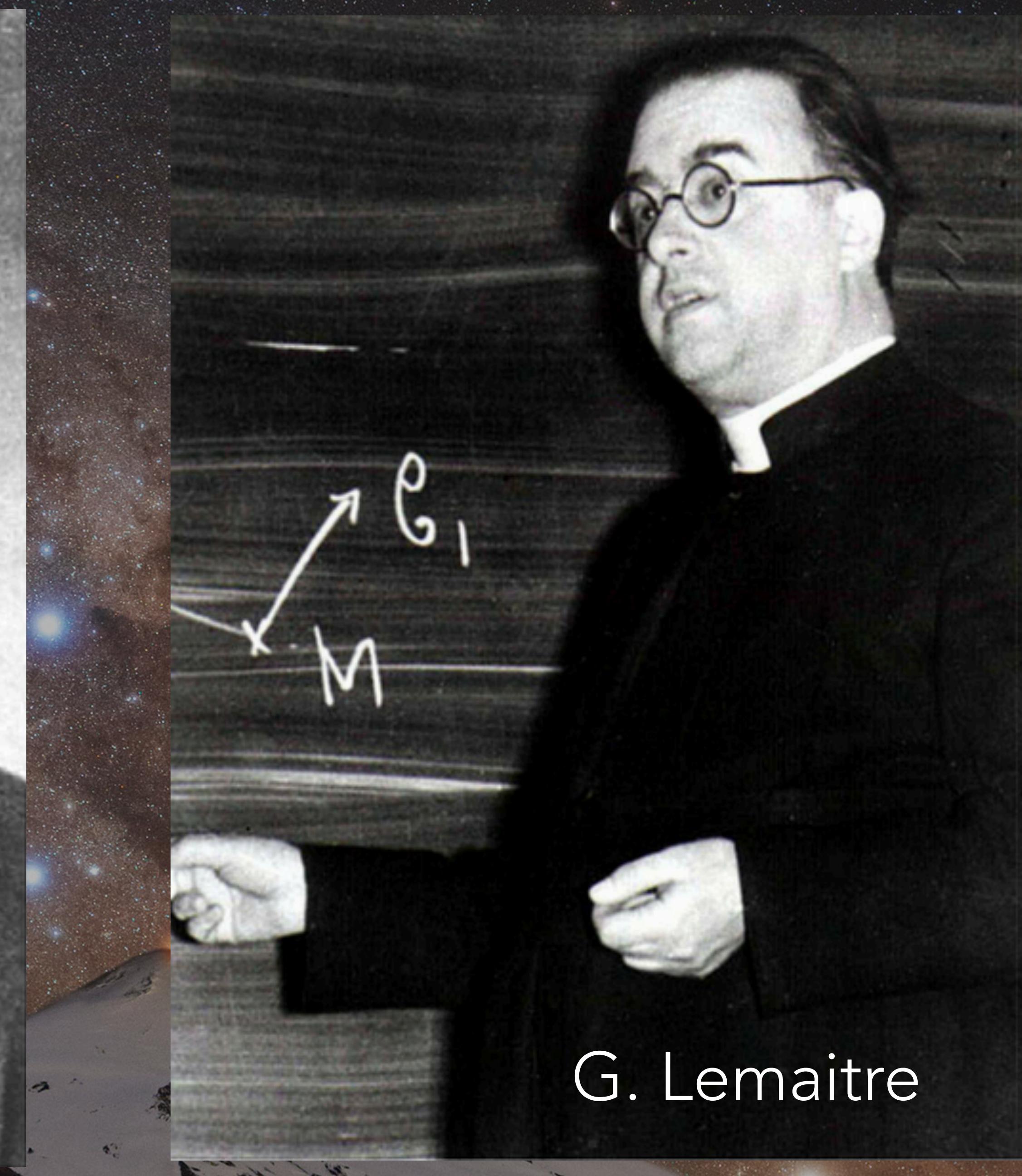


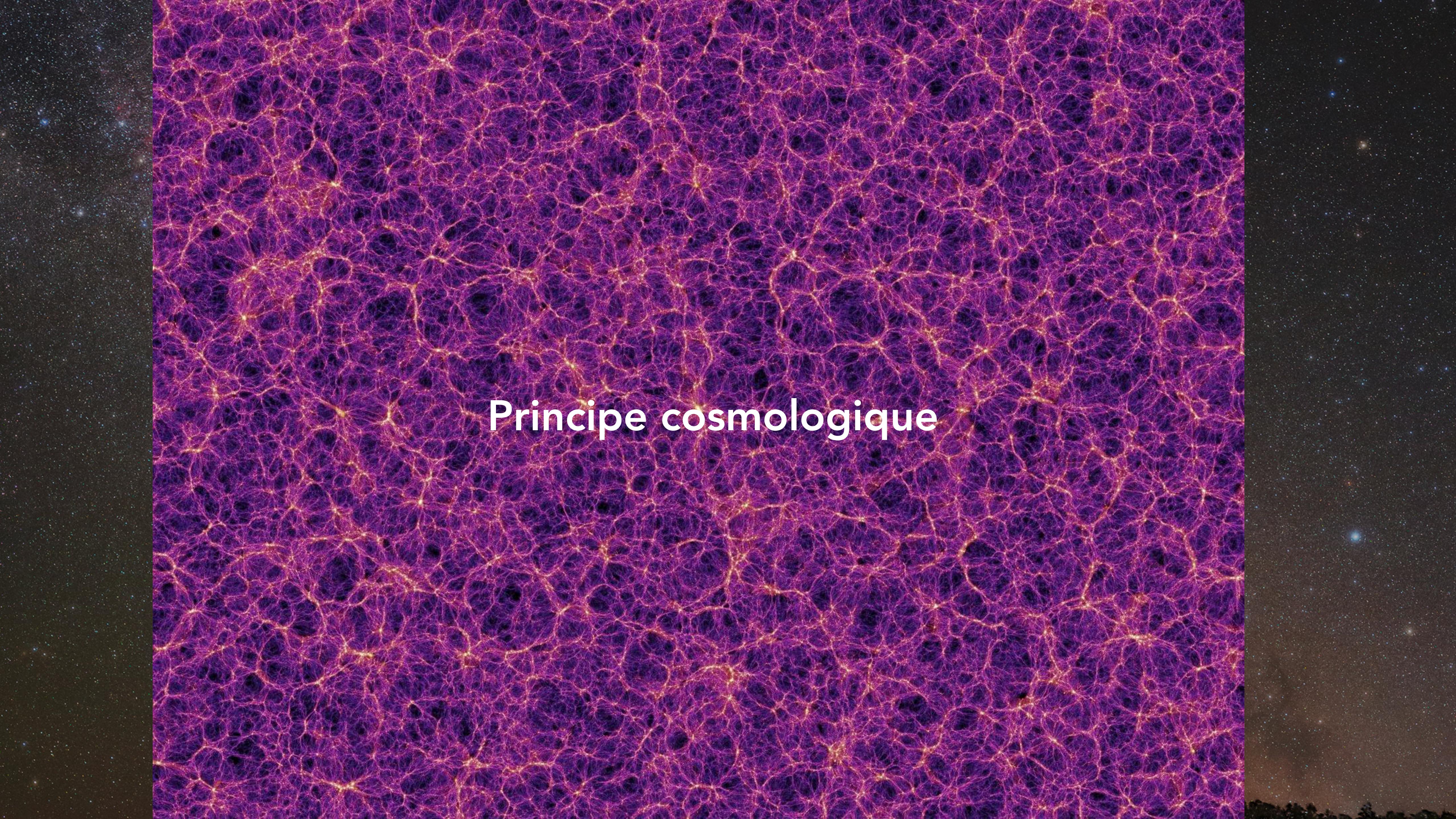
FIGURE 1 – Representations schématiques de l'élastique en expansion à  $t = 0$  et à  $t > 0$



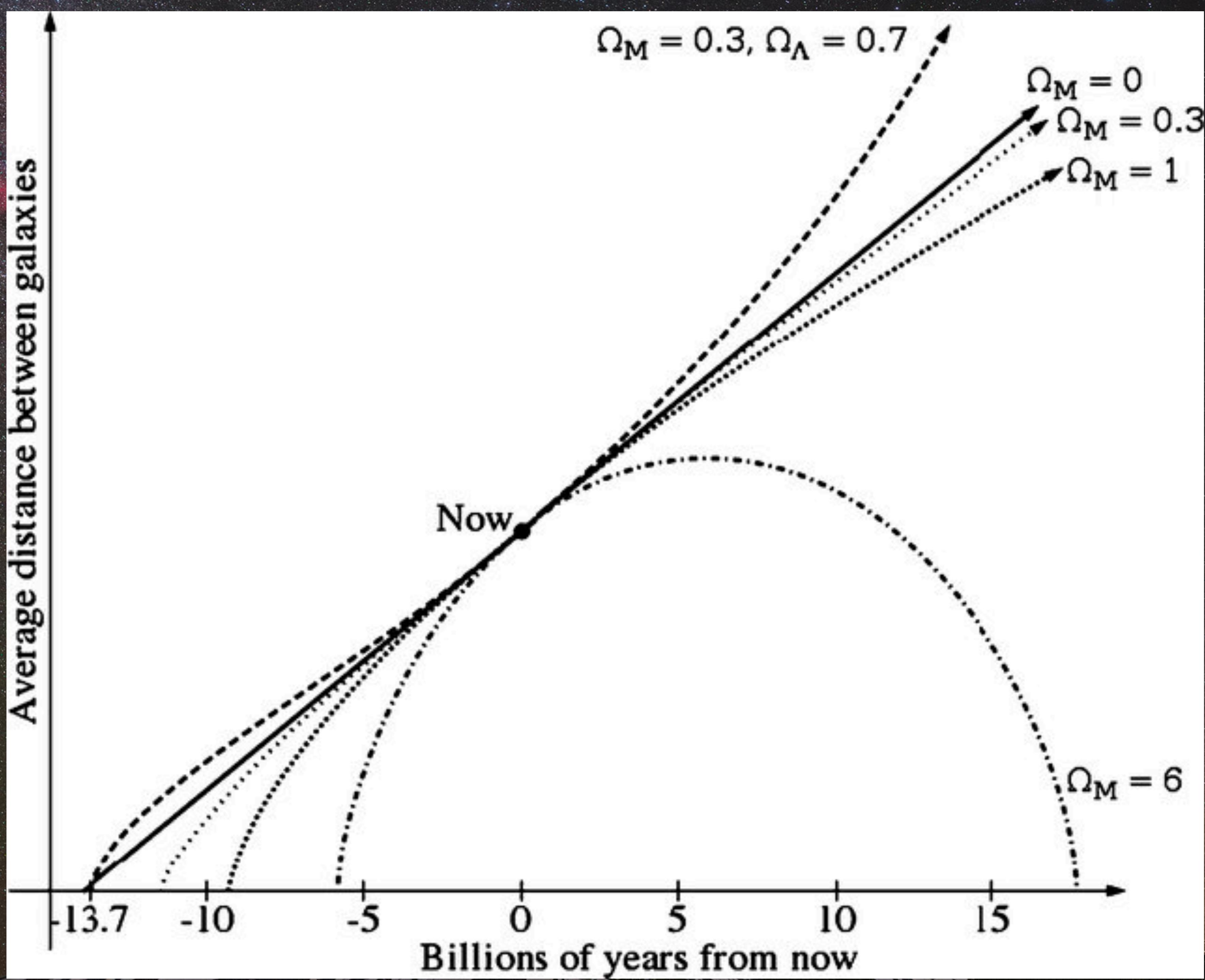
A. Friedmann



G. Lemaître



Principe cosmologique



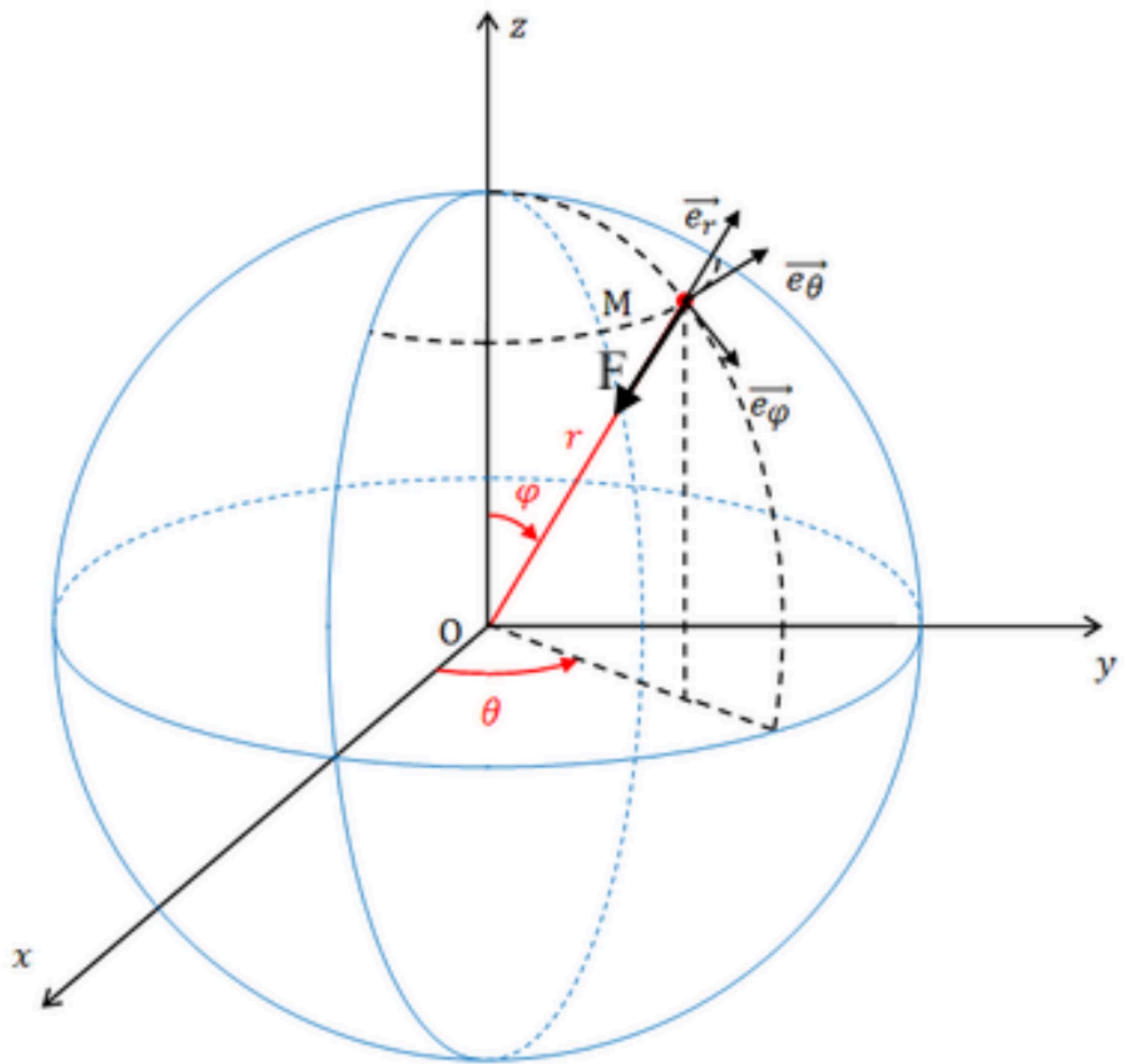
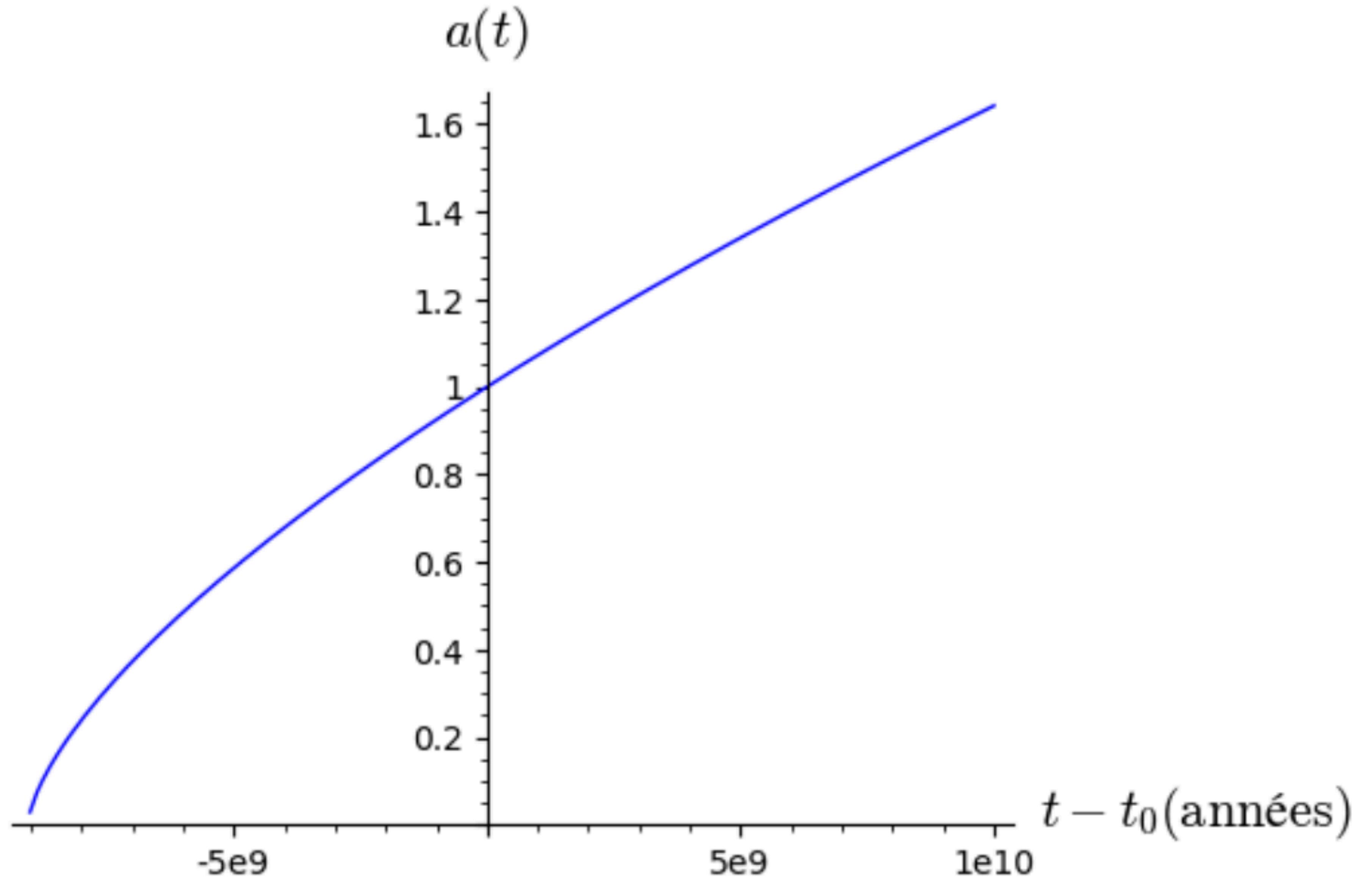
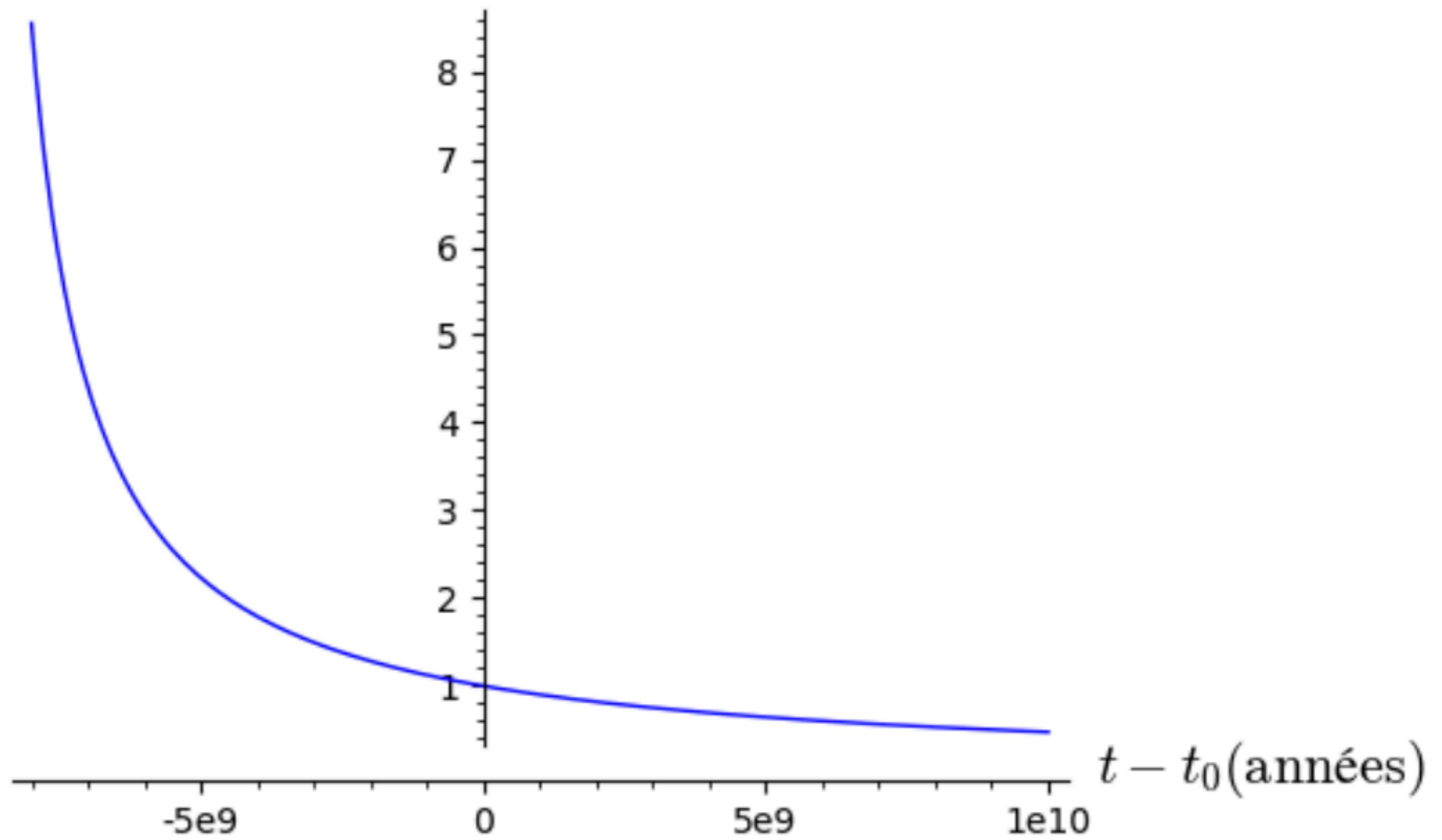


FIGURE 3 –



$$H(t)/H_0$$



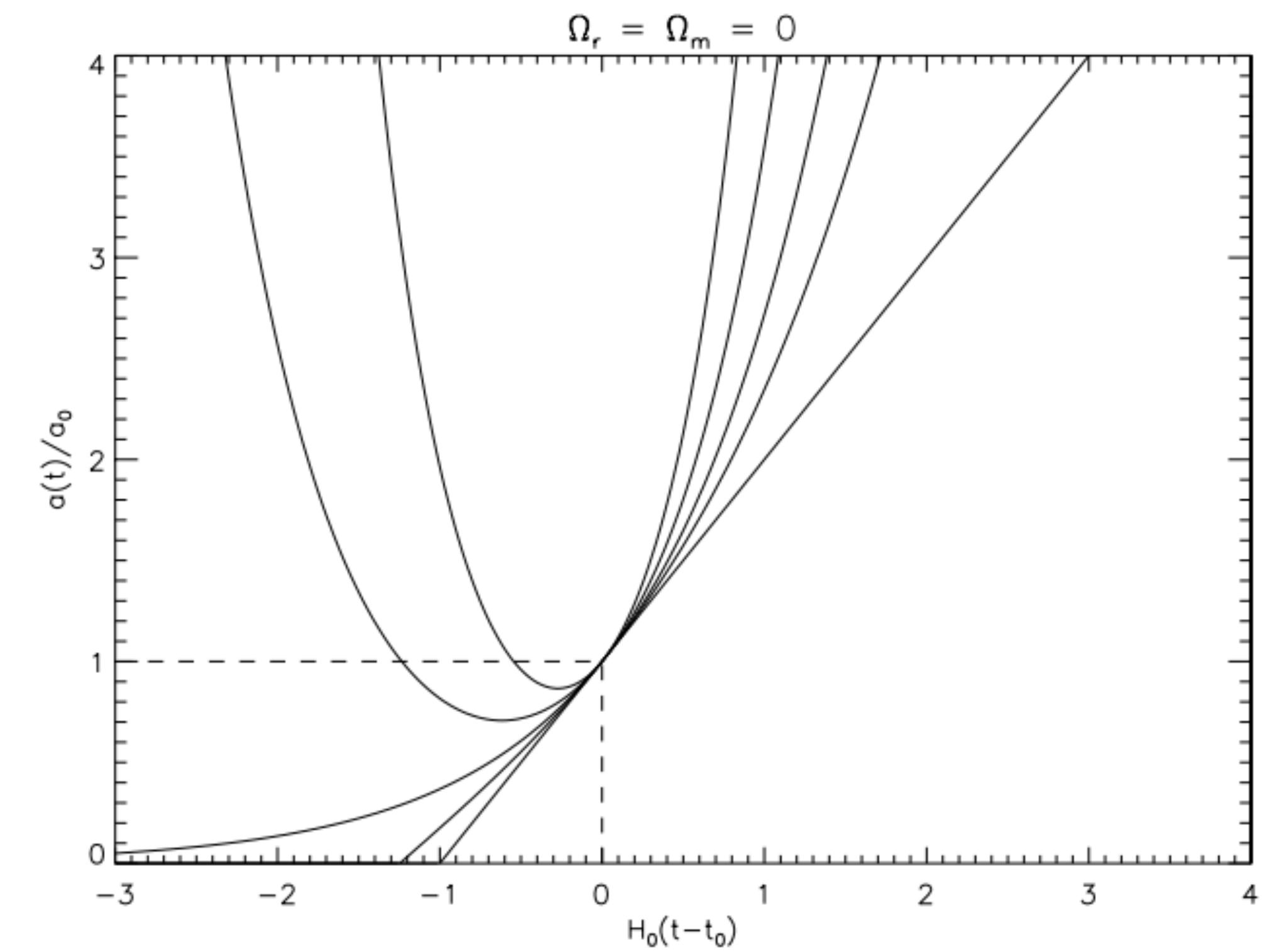
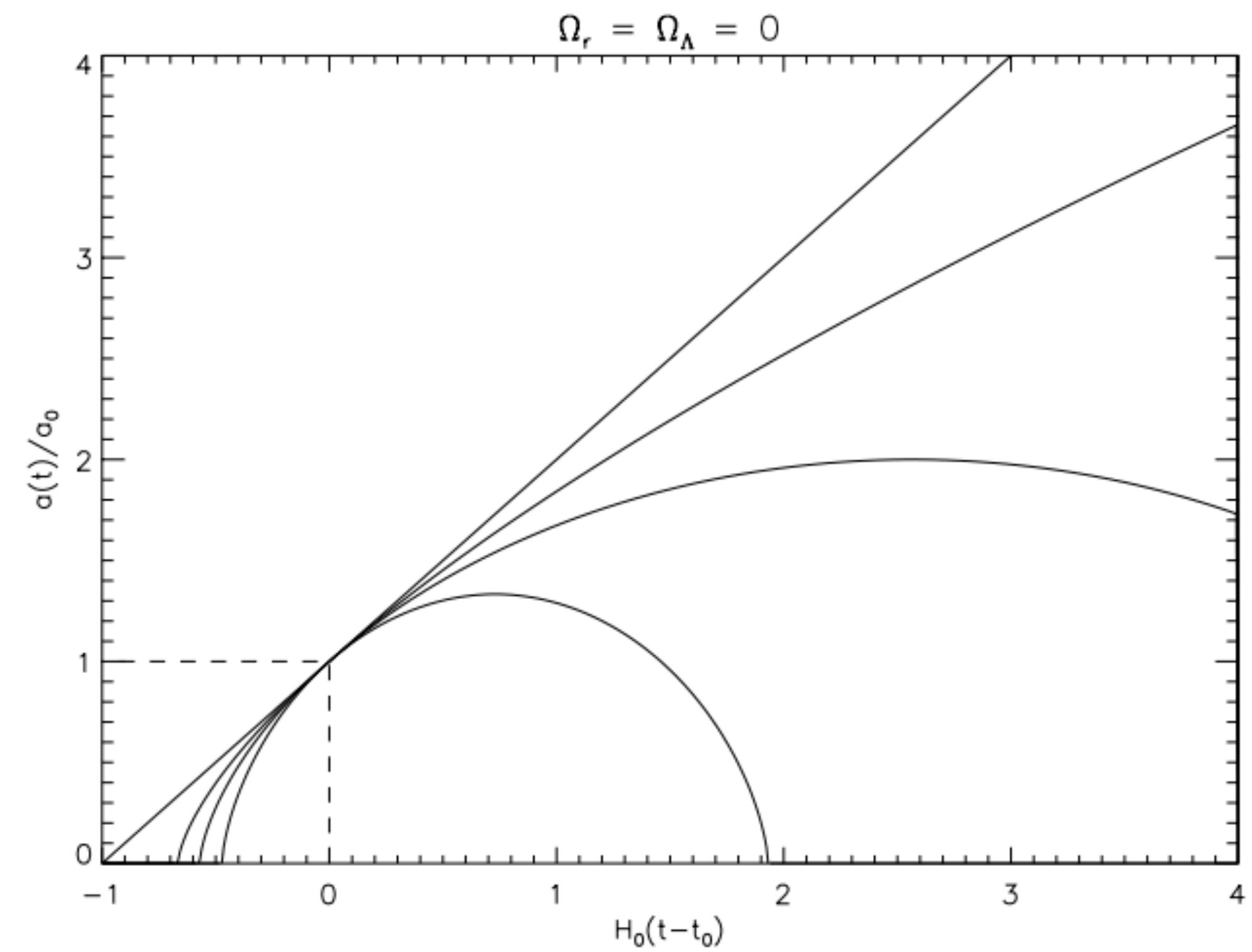
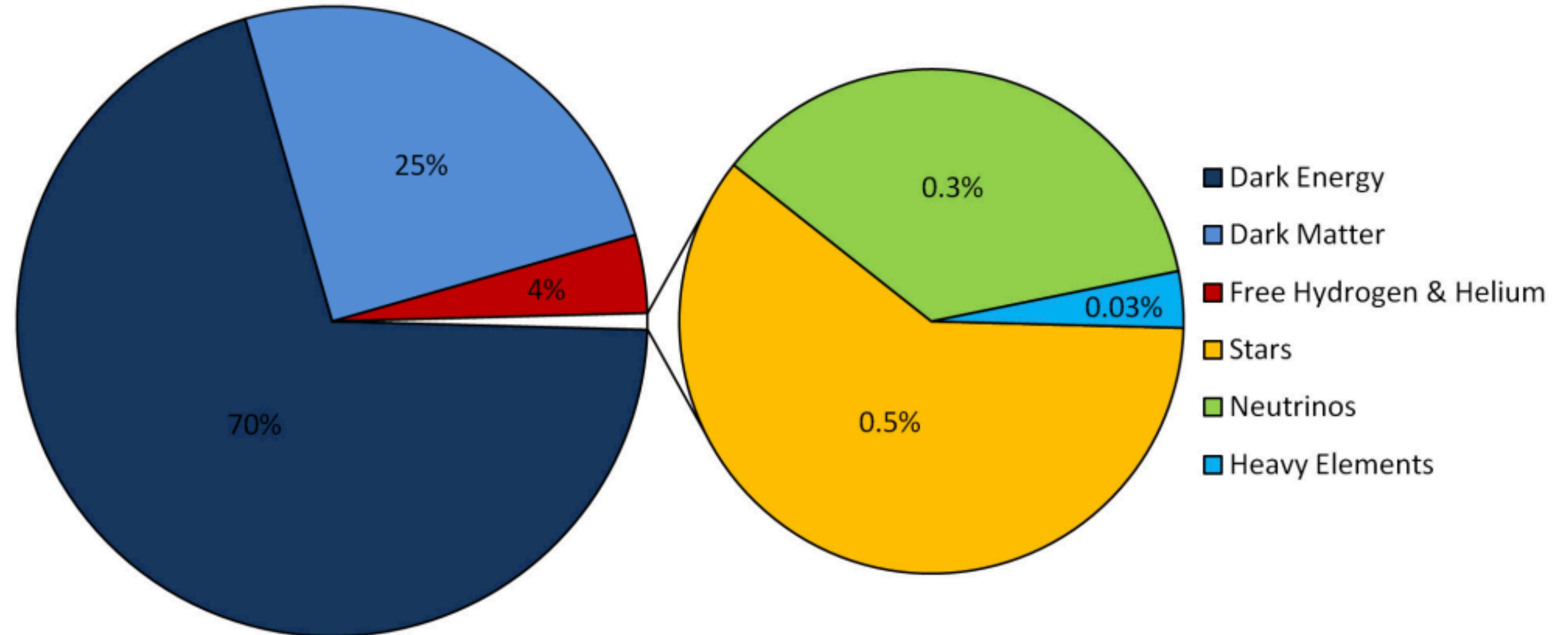
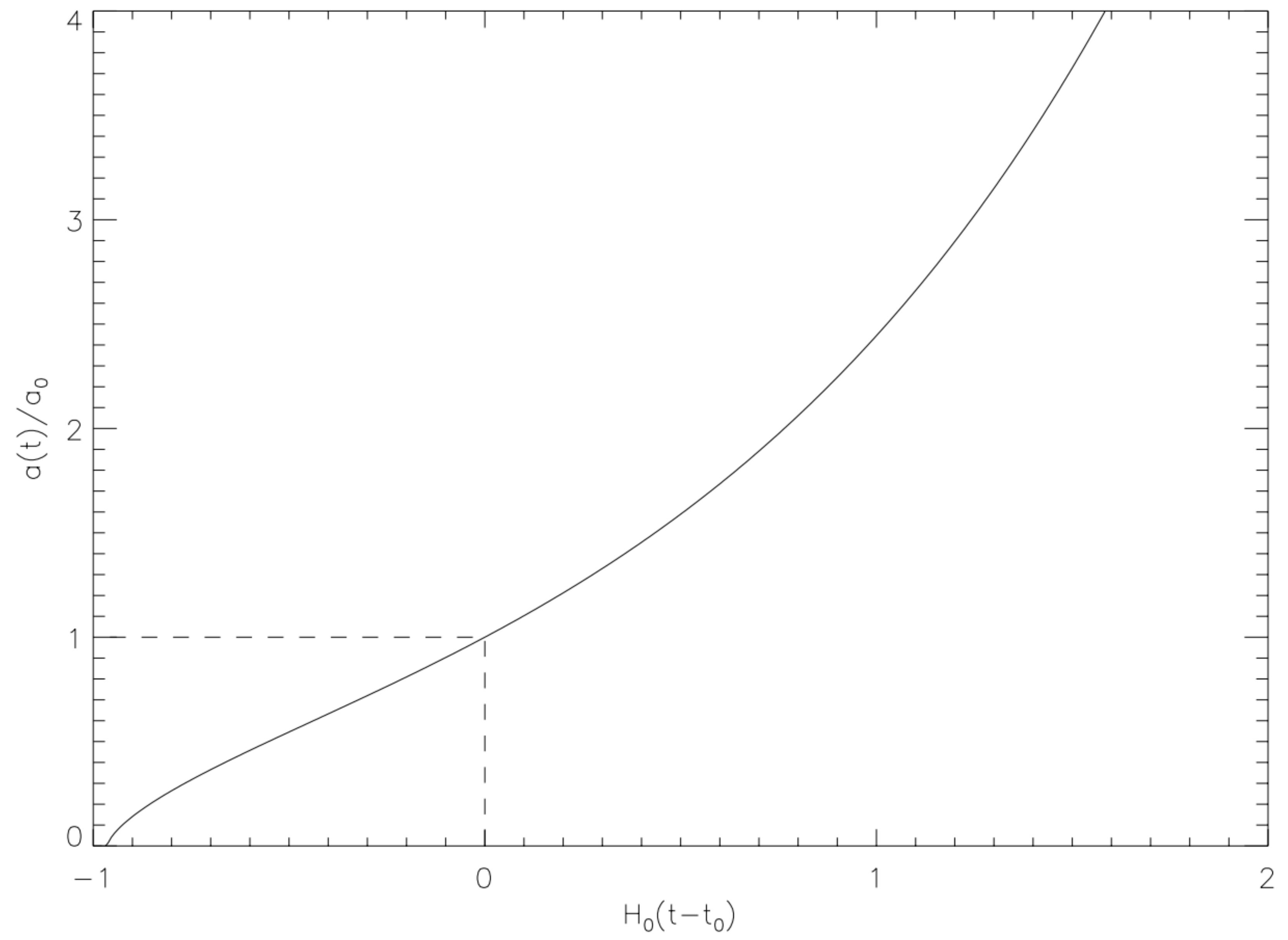
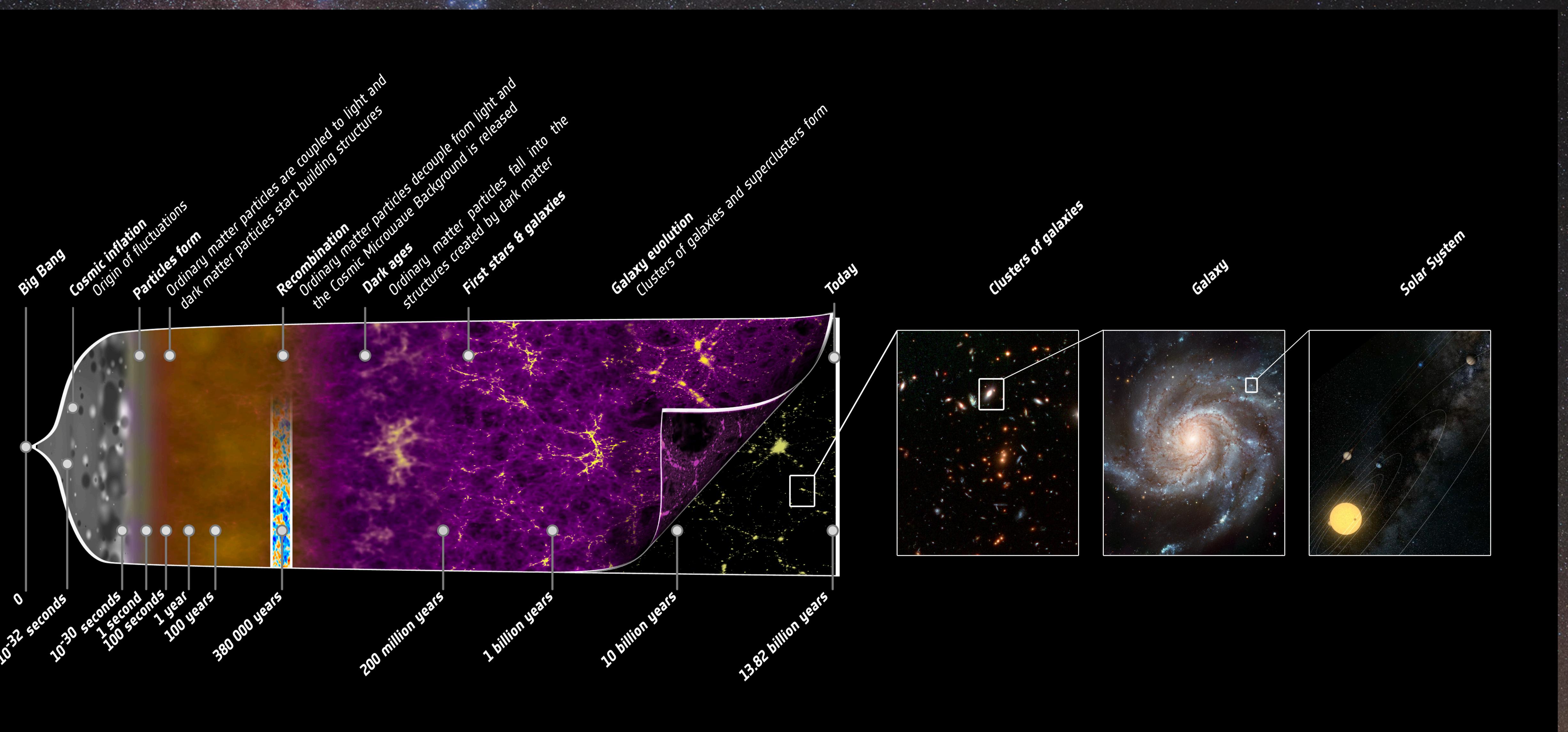
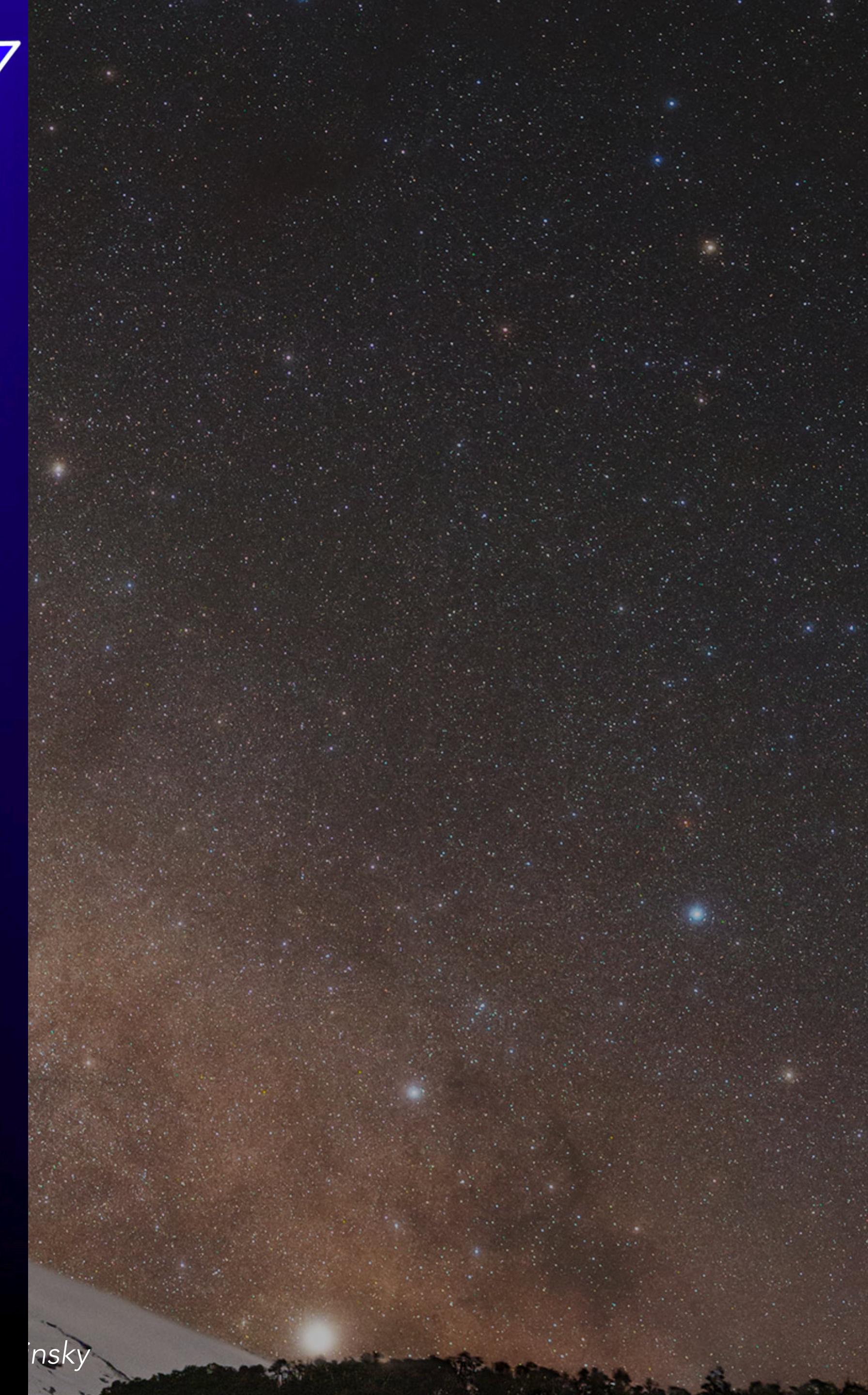
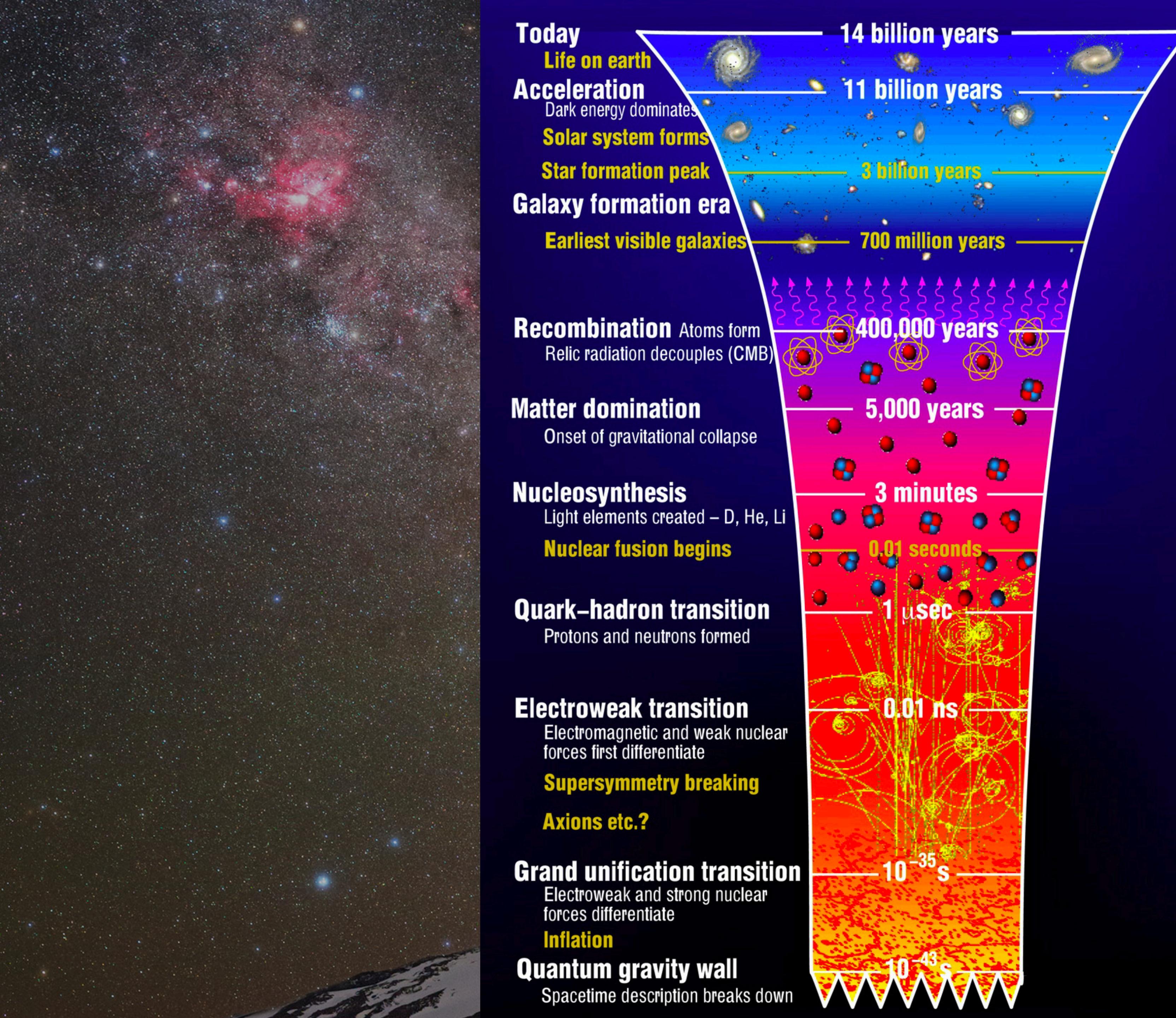


FIGURE 2 – Evolution de l'univers, représentée par son facteur d'échelle  $a$  en fonction du temps par rapport à l'époque actuelle. *A gauche* : Univers de matière ( $\Omega_r = \Omega_\Lambda = 0$ ), pour  $\Omega_m = 0, 1, 2, 4$ . *A droite* : Univers de De Sitter ( $\Omega_r = \Omega_m = 0$ ) pour  $\Omega_\Lambda = 0, 0.5, 1, 2, 4$ .









# flat $\Lambda$ CDM

