

Solutions MESA practicum 1

Getting started with MESA

3 Evolving a 1 M_{\odot} star

- 3.1**
- In the HR diagram, after the star has settled on the main sequence, the star expands slightly, moving up and right (almost constant T_{eff} , but higher L , following $L = 4\pi R^2 \sigma T^4$).
 - In the T - ρ plot, the centre of the star corresponds to the high- T , high- ρ end of the curve, while the surface is at low T and ρ . The dashed lines delineate regions where different equations of state dominate, i.e., ideal-gas pressure, radiation pressure and (non-relativistic or extremely relativistic) degeneracy pressure. We further see the T , ρ boundary lines for where certain burn conditions are met (i.e., H, He, C and O burning). The star is supported by the ideal gas law.
 - The protostar is fully convective (light blue), because the entire star is very cool, so radiative transport is inefficient. Because the whole star is convective, even a tiny superadiabaticity can carry a very large energy flux and the star is essentially ediabatic, and the luminosity of the fully convective star is practically independent of its structure, leading to the nearly vertical path in the HR diagram at low T_{eff} .
- (a) As the star keeps on contracting, the virial theorem dictates that the internal temperature will rise. As the internal temperature rises, the opacity (and thus ∇_{rad}) decreases, until radiative transport becomes more efficient: a radiative core develops (green region). This causes the star to move away from the Hayashi line, to higher effective temperatures.
- (b) The $\rho - T$ plot shows that the central part of the star has crossed the H burning limit. It is also highlighted with the nuclear energy production rate in yellow. In the HR diagram the model lies on the main sequence, at about $1 L_{\odot}$ and $\log T_{\text{eff}} = 3.76$, so the agreement is good.
- If the number of steps were proportional to the actual evolutionary timescale, we would expect far more steps on the main sequence, because the main sequence occurs on the H-burning nuclear timescale ($\sim \text{Gyr}$) which is orders of magnitude longer than Kelvin-Helmholtz contraction timescales. In MESA the opposite occurs, because the timestep control is set by how fast the stellar structure changes, not by the absolute physical time elapsed.
- 3.2**
- The **Summary_Burn** panel shows the temperature profile (yellow right axis) and burning properties, like the nuclear energy production rate and neutrino energy losses (2 shades of blue) as well as the density profile as functions of mass coordinate. The **Abundance** panel shows the composition profiles for different isotopes across the star. This clearly shows how core burning changes the abundances of different elements.
 - Abundance changes are caused by nuclear fusion in the core and later in shell-burning regions. On the main sequence, hydrogen is converted into helium via the pp chain.

As hydrogen is depleted, the helium mass fraction rises in the core. We can also see H shell burning occur. Although the pp chain dominates, there is some CNO cycle burning happening, as evidenced by the change in C, N and O abundances. Note though that the apparent dramatic changes at the bottom of the logarithmic scale correspond to very small absolute abundances.

- (a) The star reaches the end of the main sequence at about ~ 10 Gyr for a solar-type model. You can see this in the $\rho - T$ diagram from the inner core of the star no longer producing energy (not highlighted). You can see it in the Abundance plot from the core being fully made of He, and in the HR diagram when the star starts to climb towards the giant branch.
- (b) As the star leaves the main sequence, the core contracts and heats while the envelope expands and cools, shifting the star to the red giant branch. The $T-\rho$ profiles reflect this: central density rises, and the temperature increases in the contracting core. We also see that the core becomes inert: the nuclear energy and neutrino production drops in the core, but instead moves to a shell (spike in blue lines). Eventually the core becomes electron-degenerate before helium ignition: you can see this in the ρ T profile where the core has crossed the degeneracy EOS line.
- (c) During red giant evolution, the outer convective envelope deepens and dredges up material from regions previously processed by nuclear burning. This changes the surface abundances: We mostly see He^3 being moved to the surface, but the C/N ratio also changes on the surface.

4 Exploring PGSTAR options

4.1 Hand in your movie.