

The Constantly Changing Hubble Constant

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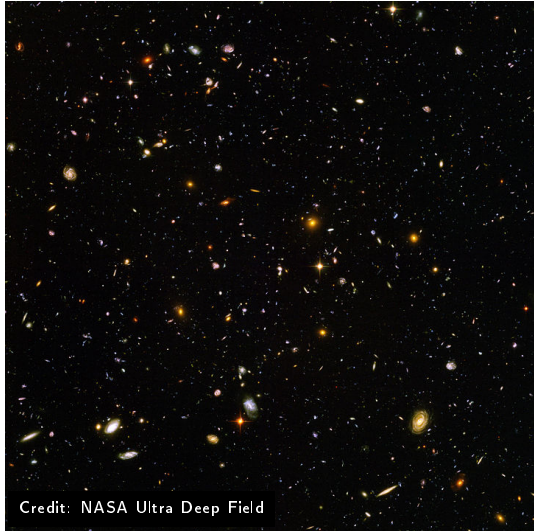
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5304 6366.

The Universe is expanding!

- ▶ But what does this actually mean?
- ▶ How do we know it is expanding?
- ▶ Why is it expanding?
- ▶ How fast is it expanding? How do we measure this?
- ▶ Is the expansion rate changing?

How do we know?

- ▶ Everywhere we look, distant galaxies are receding; more distant galaxies are receding faster.
- ▶ So either we are at the centre of a cosmic conspiracy, or all the space between all the galaxies is expanding.



Credit: NASA Ultra Deep Field

Is the solar system expanding? Are we expanding?

Go to www.menti.com (code 5304 6366) and choose:

1. Yes, a lot
2. Yes, but only a tiny amount
3. No

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Is the solar system expanding? Are we expanding?

- ▶ Other forces - molecular forces between the molecules in your body, and gravitational forces between the Sun and the planets - are far more than strong enough to overcome the effect of cosmic expansion.

- ▶ Gravity is even strong enough to keep the Andromeda Galaxy from receding from us.



Credit: David Dayag

- ▶ It's only the furthest objects - whose gravitational attraction to us is negligible - that recede.

What does *recession velocity* actually mean?

- ▶ We say 'distant galaxies are moving away from us'. This is informal language.
- ▶ They aren't really moving, they just appear to be - because the intervening space is expanding.
- ▶ Sometimes this distinction is important - for example, the recession velocity can exceed the speed of light.

Which 'Ed' first had the idea that the Universe is expanding?

Go to www.menti.com (code 5304 6366) and choose:

1. Edmond Halley
2. Edwin Hubble
3. Edgar Allan Poe

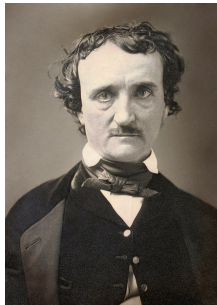
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History

- ▶ In 1848 Edgar Allan Poe published *Eureka*, which included a description of expanding space.



Credit: Public domain

- ▶ Expansion is not obvious without large telescopes and so isn't usually part of pre-modern cosmologies. Full understanding only came in the 20th century.

So how fast is the expansion?

- ▶ For every additional distance of one megaparsec, there's an additional recession velocity of about 70 kilometers per second.
- ▶ So the expansion speed is about 70 kilometers per second per megaparsec.
- ▶ One megaparsec is about three million light years. It's the typical distance between galaxies.
- ▶ 70 kilometers per second is about 150,000 miles per hour.

So how fast is the expansion?

Start with a distance:

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13.5 million years later it will be 1% longer:

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Start with a distance:

13.5 million years later it will be 1% longer:

Continental drift is about six times faster...

H_0

- ▶ The current expansion rate is called the *Hubble constant* or *Hubble parameter* and is denoted ' H_0 '.
- ▶ The ' H ' commemorates Edwin Hubble (1889-1953), who was one of the first to measure it.
- ▶ The ' 0 ' refers to today. The expansion rate was different in the distant past.



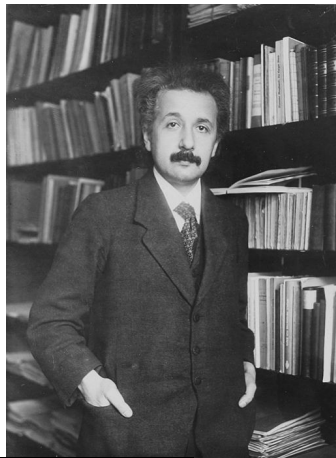
Credit: Johan Hagemeyer

Why does the Universe expand?

- ▶ Science is not so good with 'why?' questions...
- ▶ There's an *initial condition*: the Universe started expanding at the Big Bang.
- ▶ The later behaviour of the expansion (does it slow down? speed up?) then depends, essentially via gravity, on *what's in the Universe*.

Why does gravity play a role?

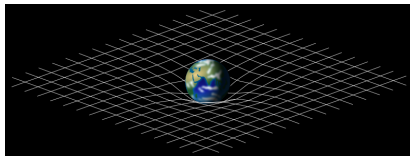
- ▶ *General relativity*, our modern theory of gravity, is due to Einstein (1916).



Credit: Paul Ehrenfest

Why does gravity play a role?

- ▶ Einstein's theory 'promotes' spacetime from being a passive background to be a physical entity in itself, capable of change.
- ▶ Mass/Energy
(remember - they are the same - $E = mc^2$)
bends spacetime,
essentially changing
distances and angles.



Credit: Mysid

Why does gravity play a role?

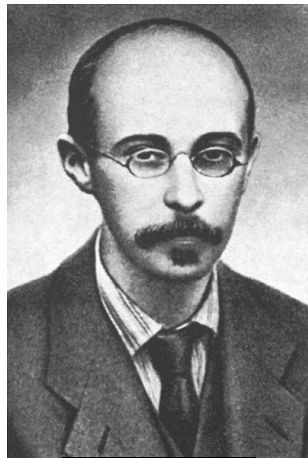
- ▶ This 'changing of distances and angles' works locally; the distorted spacetime governs how objects move, and this leads e.g. to the apple falling from the tree.

Why does gravity play a role?

- ▶ This 'changing of distances and angles' works locally; the distorted spacetime governs how objects move, and this leads e.g. to the apple falling from the tree.
- ▶ But it also works on the Universe as a whole - mass/energy can cause distances to change *everywhere* in the Universe - and in particular can lead to increasing distances everywhere. This is the expansion that we see.

Contents of Universe control expansion

- It was Alexander Friedmann (Алекса́ндр Алекса́ндрович Фри́дман) (1888-1925) who first realised this (1922); he found a *dynamical* solution to Einstein's equations.



Credit: Public domain

What if we go backwards in time?

- ▶ George Lemaître (1894-1966) realised that if the Universe was expanding then it must, at an earlier stage, have been very small; he thereby invented the idea of the 'Big Bang'.



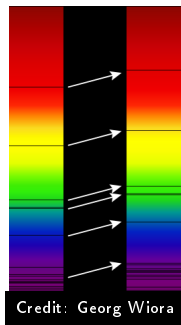
Credit: Public domain

How do we measure the expansion rate?

- ▶ In theory it's easy: find a distant galaxy, measure its recession velocity and its distance, and take the ratio.
- ▶ Example: a galaxy is receding at 1600 kilometers per second and is 20 megaparsecs away;
then H_0 is 80 kilometers per second per megaparsec.

Can use *redshift* to measure recession velocity

- ▶ Light from distant galaxies gets stretched by the expansion; this makes it turn redder.
- ▶ It's fairly easy to measure the amount of red-shifting, as spectral lines are a convenient reference point. The redshift then immediately gives the velocity.



V M Slipher

- ▶ The first redshifts for galaxies (known then as nebulae) were made in 1912 by Vesto Slipher (1875 - 1969) at the Lowell Observatory in Flagstaff, Arizona.



Recession velocity for galaxy M77

From Slipher's 1917 'Lowell Observatory Bulletin 80':

These recent spectrograms have been measured for radial velocity. The results for the four spectrograms — two of 1913 and two of 1917 — are as follows:

Plate 1913, November 6,	Velocity -1060 km.
Plate 1913, November 22, 23,	Velocity -1150
Plate 1917, November 6, 7, 8,	Velocity -1080
Plate 1917, November 12 to 16, two prisms	Velocity -1130 s
	Velocity -1145 c
	Velocity -1135 c
	<hr/>
Mean Velocity	-1120 km.

Modern value (from NASA/IPAC Extragalactic database):

BASIC DATA for MESSIER 077 ([Back to INDEX](#))

Helio. Radial Velocity	:	1137 +/-	3 km/s	
Redshift	:	0.003793 +/- 0.000010		2014MNRAS.440..696A

Can use Cepheids to measure distance

- ▶ Cosmic distances can be estimated using variable stars called *Cepheids*. We know how much light they produce (it's related to their period) - so they are *standard candles*. Their observed brightness then tells us their distance.
- ▶ The brightness/period relation was discovered by Henrietta Swan Leavitt (1868 - 1921) at Harvard.



Credit: Public Domain

How far away is that Cepheid?

A and B are two identical Cepheids; A is 100 parsecs distant; B is 4% of the brightness of A . How far is B ?

Go to www.menti.com (code 5304 6366) and choose:

1. 500 parsecs
2. 2500 parsecs
3. 4 parsecs

How far away is that Cepheid?

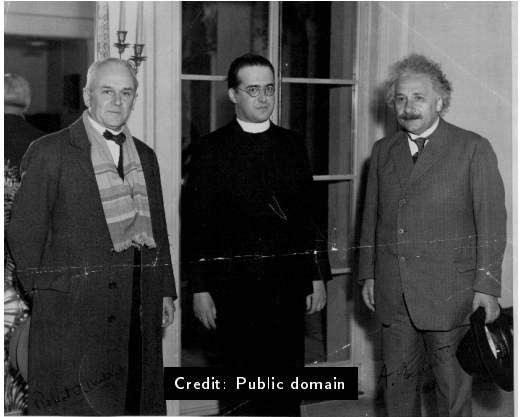
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Recession velocity vs distance relationship

It was Lemaître who first suggested a linear relationship between recession velocity and distance.



Extract from Lemaître's paper of 1927

l'instant où la lumière est reçue et à l'instant où elle est émise. v est la vitesse de l'observateur qui produirait le même effet. Lorsque la source est suffisamment proche nous pouvons écrire approximativement

$$\frac{v}{c} \approx \frac{R_2 - R_1}{R_1} \approx \frac{dR}{R} \approx \frac{R'}{R} dt \approx \frac{R'}{R} r$$

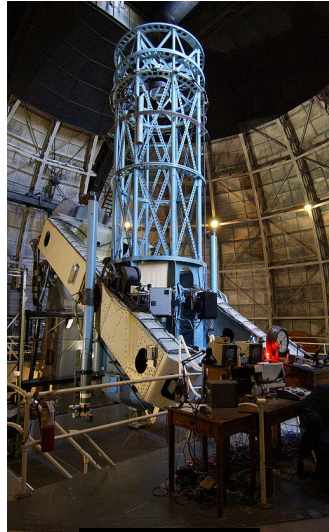
où r est la distance de la source. Nous avons donc

$$\frac{R'}{R} \approx \frac{v}{cr} \quad (23)$$

Les vitesses radiales de 43 nébuleuses extra-galactiques sont données par Strömberg (⁵).

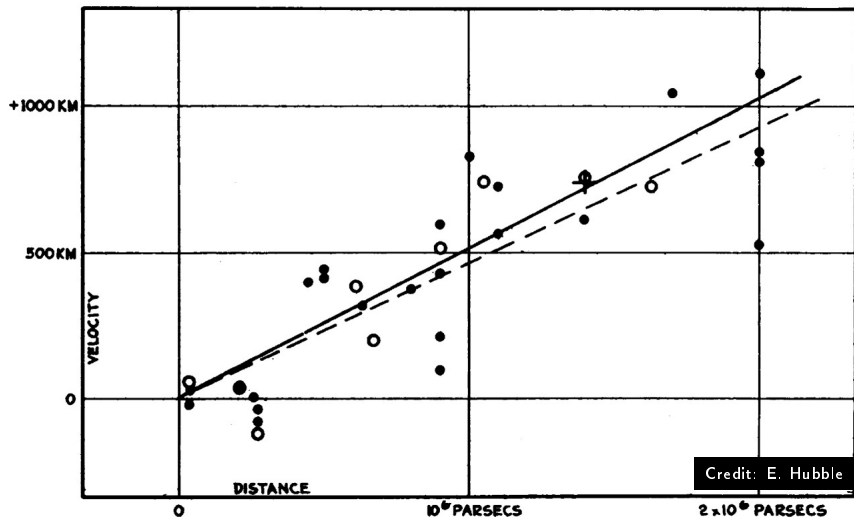
Recession velocity vs distance relationship

Hubble then provided supporting evidence - he used the 100 inch Hooker Telescope at Mt Wilson to spot Cepheids in a range of galaxies.

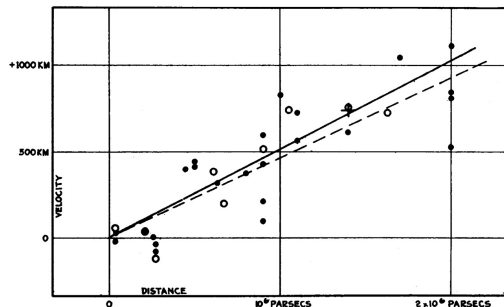


Credit: Ken Spencer

Hubble's plot - from 1929 PNAS paper



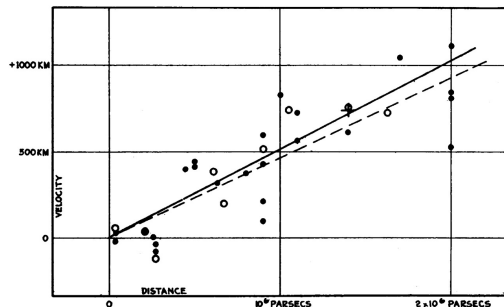
What value for H_0 does this plot give?



Go to www.menti.com (code 5304 6366) and choose:

1. 70 km per second per megaparsec
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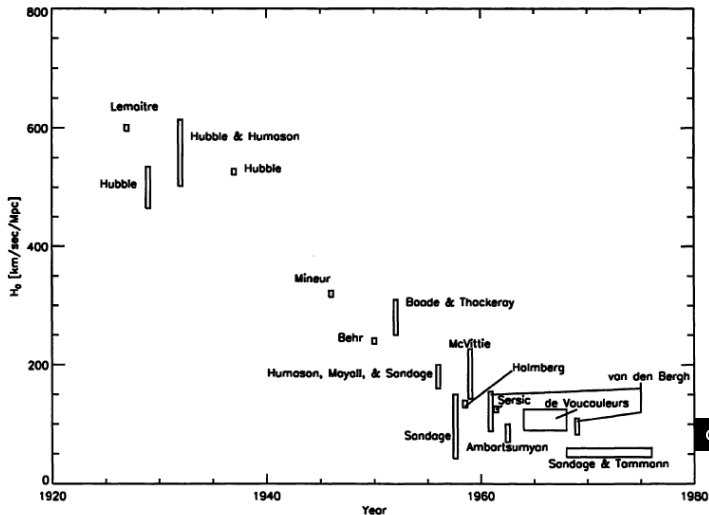
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Falling estimates of H_0 1920 - 1980



Credit: V. Trimble

Distances had been underestimated

Several revisions (distance up, H_0 down) followed:

1. Two types of Cepheids
2. Interstellar absorption of light

3. Scott effect (visible galaxies tend to be brighter than average)



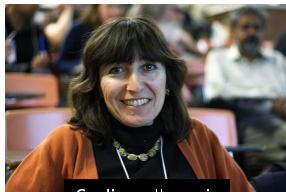
Credit: MacTutor

4. Regions of hot gas were being misidentified as bright stars

Hubble Space Telescope

- ▶ The Hubble Space Telescope was built (in part) to measure H_0 .

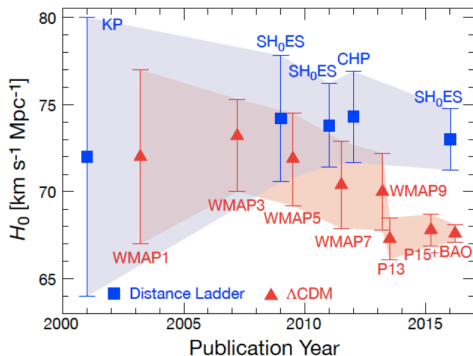
- ▶ Lead investigator Wendy Freedman.



Credit: pollosaurio

- ▶ Used supernovae as additional standard candles.
- ▶ $H_0 = 72 \pm 8$ kilometers per second per megaparsec.

More recent values - diverging opinion



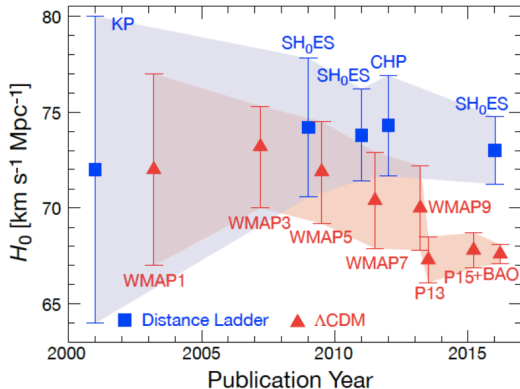
Credit: W. Freedman

- ▶ Blue results use standard candles: $H_0 = 74 \pm 1.4$ kilometers per second per megaparsec.
- ▶ Red results use data from the early Universe: 67.4 ± 0.6 kilometers per second per megaparsec.

H_0 from early Universe data

- ▶ For 400,000 years after the Big Bang, the Universe was opaque. Then it became transparent, and we can today observe (red-shifted) light from then (the *Cosmic Microwave Background*).
- ▶ We can calculate the physical size of ripples in the hot gas of the early universe (it is the physics of sound waves), and observe their angular size on the sky. Trigonometry then gives us the distance from which we are viewing the light.
- ▶ Redshift is easy - the spectrum has a prominent peak.
- ▶ Finally, correct for the evolution of H_0 from then to now.

Hubble tension

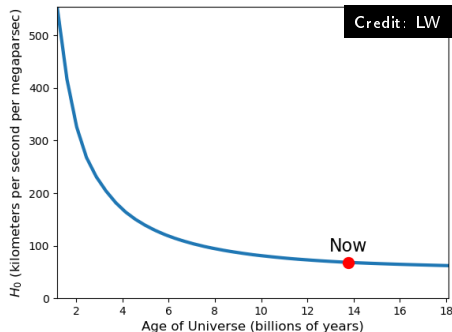


Credit: W. Freedman

Should we be worried by the discrepancy? It could indicate that our theory is wrong. But it might also just be a systematic error in one of the measurements.

Hubble constant - past and future

- ▶ H_0 isn't constant over cosmic timescales. The evolution is controlled by the contents of the Universe.



- ▶ Matter - mostly dark matter, but also hydrogen and helium - tends to push H_0 towards zero (no expansion).
- ▶ But the Universe also contains dark energy, and this opposes this trend. In the distant future the expansion will stabilise near 57 kilometers per second per megaparsec.

Bibliography

1. Nussbaumer, H. and Bieri, L., *Discovering the Expanding Universe*, Cambridge University Press, 2009
2. Trimble, V., H_0 : *The Incredible Shrinking Constant*, 1925-1975, Publications of the Astronomical Society of the Pacific, 108, 730, 1996
3. Way, M. J. and Hunter, D. eds., *Origins of the Expanding Universe: 1912-1932*, Astronomical Society of the Pacific Conference Series, 2013

Where to find this document

- ▶ You can find the presentation at
<https://tinyurl.com/bycke8v6>

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13. Lemaître paper extract: Lemaître, G., *Annales de la Société scientifique de Bruxelles*, A47, 49-59, 1927. Reprinted in Nussbaumer and Bieri, *Discovering the Expanding Universe*, Cambridge University Press, 2009
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15. Hubble plot: Source: Hubble, H., "A relation between distance and radial velocity among extra-galactic nebulae", *PNAS* March 15, 1929 15 (3) 168-173. From <https://www.pnas.org/content/15/3/168/tab-figures-data>
16. Trimble plot: Source: Trimble, V., "H0: The incredible shrinking constant", 1925-1975. 1996 *PASP* 108 1073. From <https://iopscience.iop.org/article/10.1086/133837/meta>
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19. Freedman plot: Source: Friedmann, W. L. "Cosmology at a Crossroads: Tension with the Hubble Constant." *Nature Astronomy* 1 (2017): 0169. From <https://arxiv.org/abs/1706.02739>
20. H0 History plot: Source: Original work. From <https://github.com/LorneWhiteway/HistoryOfH0Talk/blob/master/h0.png>