

Space Physics Project , FYS3610

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Introduction

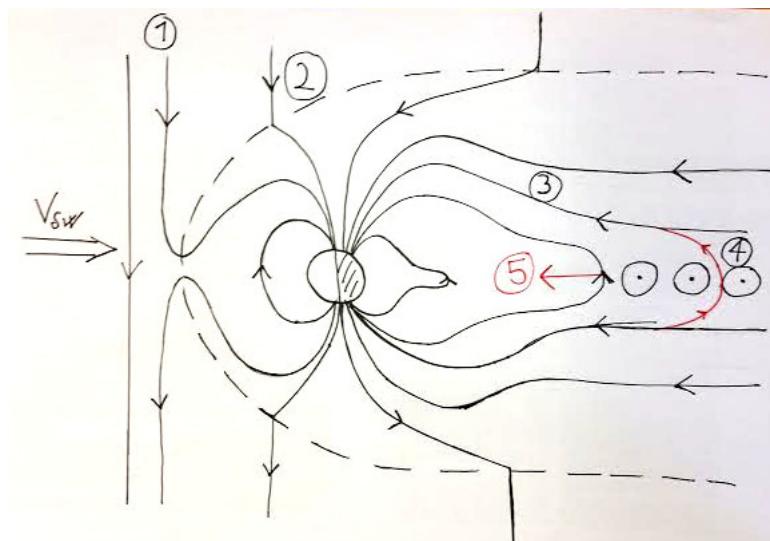
In this project we are going to analyse the ionospheric and auroral dynamics and its relation to the solar wind conditions. Furthermore, we will compare our observations to theory. We will investigate in a specific timeslot of 6 hours from the 6. January 2011 from 18:00 o'clock to the 07. January 2011 00:00 o'clock.

1 Theory part

In order to understand the movements and changes in the system of the earth and the sun, we want to introduce some basic concepts of space physics. A really important steady state picture of what is going on with respect to the earth's magnetic field, the interaction of the earth's magnetic field with the IMF and the consequences of the earth's environment (such as current flows or aurora) is the Dungey cycle. (the Dungey cycle is only valid when the IMF is southward orientated) The Dungey cycle is as already mentioned a steady state picture. This means that the magnetic flux added to the open magnetic field lines through day side reconnection is equal to the magnetic flux subtracting from the open magnetic field lines through night side reconnection. However, this cycle is in reality not a steady state picture. We observe different phases such as adding magnetic flux (growing of the polar cap), some kind of explosive expansion (subtraction of magnetic flux from the polar cap) and a recovery phase. This phenomena is called a "Auroral substorm". In addition to that, we can extract information about what is going on with the system earth-sun by looking at different data sets containing information about the magnetic field on earth, the solar wind bulk speed, the auroral intensity, the current density and further data from different physical measurements. Based on this measurements we will interpret the received data and relate it to the theoretical predictions and the models we introduced in the theoretical part.

1.1 The Dungey cycle - a steady state picture

Figure 1: Dungey cycle



If the magnetic field is northward orientated ($B_z > 0$) nothing happens at the boundary of the terrestrial magnetic field and the magnetic field of the sun/ the solar wind. IMF can not mix with the terrestrial magnetic field because everything is frozen in. Since the terrestrial magnetic field and the IMF are not anti-parallel orientated, magnetic reconnection does not occur and hence, everything is and stays frozen in. However, if the IMF is southward orientated ($B_z < 0$), we do have anti-parallel orientated magnetic field at the boundary of terrestrial magnetic field and the IMF (at the magnetopause). Therefore, magnetic reconnection occurs at the magnetopause, which causes a cyclic reconnection pattern at day/night-side. If we assume a steady state picture, we call the cycle, which happens for southward IMF, the dungey cycle. The different parts of the cycle can be described as follows:

- 1 opposite polarity field lines reconnect at magnetopause (if IMF southward orientated)
- 2 magnetic field lines are still connected to the solar wind. (frozen in again) Therefore, the magnetic field line is draged along with the solar wind in the direction of the tale. (magnetic tension force)
- 3 magnetic flux is added to the tail and compresses the plasma sheet (magnetic pressure increases, plasma sheet gets thinner and thinner)

4 magnetic reconnection occurs in the tail.

5 reconnected magnetic field lines return to the day side (imagination of "rubber bands" help to understand this)

1.2 Auroral substorm

As already mentioned in the introduction, the assumption of a steady state picture is not always an appropriate description of what happens in reality. To specify this, the rate of adding magnetic flux and the rate of subtracting magnetic flux from the polar cap are in general not equal. However, after studying data from different measurements, Akasofu could state out in 1964 different phases of a cycle so called:

- 1 growth phase
- 2 expansion phase
- 3 recovery phase

1.2.1 growth phase

The day side reconnection dominates and the aurora moves to lower latitudes. The auroral activity increases in the cusp and open magnetic flux is added to the polar cap. The magnetic flux, which is added due to the day side reconnection makes the tail radius increase (because the field lines are dragged tail-wards with the IMF).

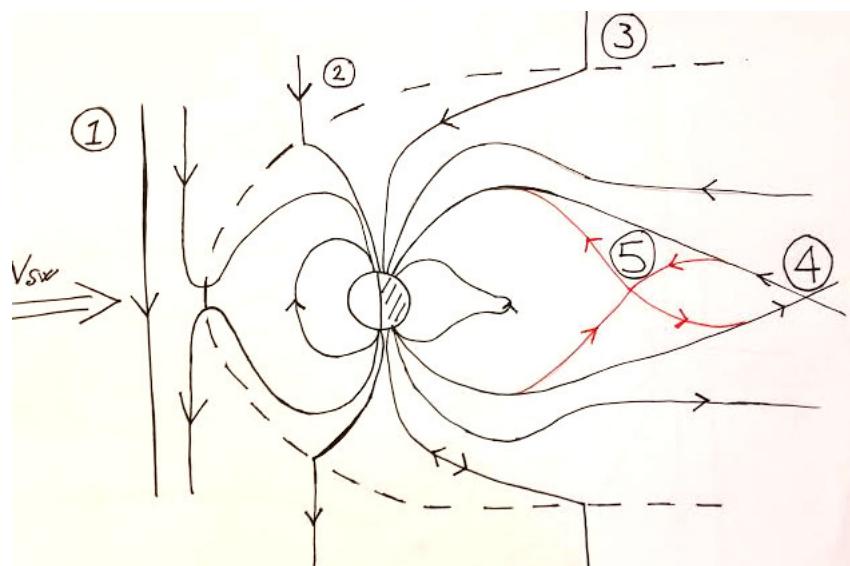
1.2.2 Expansion phase

Sudden onset of night-side reconnection at the NENL. Sudden brightening of the night side aurora. Poleward and westward movement of the aurora (on average). It closes open magnetic flux.

1.2.3 recovery phase

The night side reconnection continues at smaller rates and the auroral display quietens. Furthermore, the polar cap contracts and the magnetosphere returns to pro growth phase "equilibrium".

Figure 2: Auroral substorm



1.3 Currents

When we are dealing with the interaction of the earth's magnetic field with the IMF, it seems to be reasonable to investigate into how the currents close to earth behave. The reason for this is due to the connection of magnetic fields and flowing currents. We know from Maxwell's equations that flowing electric currents cause a magnetic field. In particular, this relation is given by Ampere's law (which follows from Maxwell's equations):

$$\nabla \times \vec{B} = \mu_0 \cdot \vec{J} \quad (1)$$

Therefore, we can derive the current flow by looking at the magnetic field. (in particular by looking at the curl of the magnetic field)(see measurement AMPERE)

In general, we can express the current flow in a given system with given magnetic field and electric field as follows:

$$\vec{j} = en(\vec{v}_i - \vec{v}_e) = \sigma_p \vec{E}'_{\perp} + \sigma_H \frac{\vec{E}'_{\perp} \times \vec{B}}{B} + \sigma_{\parallel} \vec{E}_{\parallel} \quad (2)$$

Here, σ_p is the Pedersen conductivity, σ_H is the Hall conductivity and σ_{\parallel} is a conductivity, which belongs to a current parallel to the electric field. We notice that the different currents caused by different conductivities flow in orthogonal directions with respect to each other. The orientation of \vec{E} means parallel or orthogonal to the magnetic field. A discrete expression of the conductivities will not be needed in this report.

$$\vec{E} = \vec{E}_{\perp} + \vec{E}_{\parallel} \quad (3)$$

$$\vec{E}'_{\perp} = \vec{E}_{\perp} + \vec{v}_n \times \vec{B} \quad (4)$$

$$\vec{E}'_{\perp} \approx \vec{E}_{\perp} \quad (5)$$

1.3.1 Earth's current system

We want to give a short overview about the current system of the earth. We want to limit us on currents, which can actually be measured by the measurements used in this report. The currents, which flow near the polar cap can be described as follows: On the polar cap, we find Hall currents, Pedersen currents and Field-aligned currents. The Electron Hall currents flow at lower altitudes than the Pedersen currents. The flow of the Hall currents is also associated with the so called twin cell convection. The ionospheric Pedersen currents are connected to the Field aligned currents (FAC). The FAC are on the one hand side closed with the magnetopause current and on the other hand side with the ring current. This field aligned currents are discovered by a Norwegian explorer called Kristian Birkeland in 1908. All current systems need to be closed, which is the case. The current flow on the polar cap can be seen in figure 3.

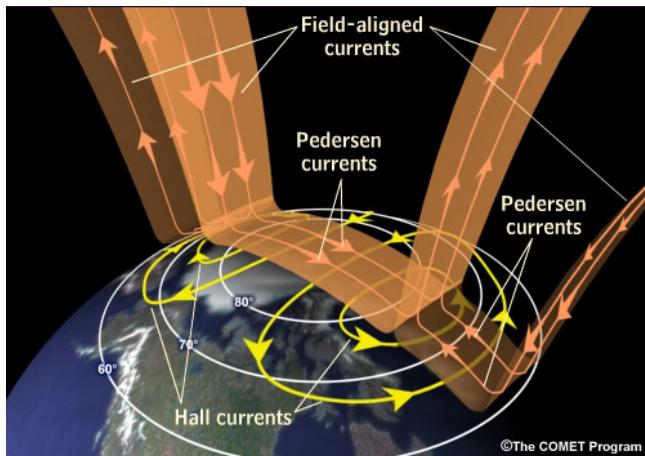


Figure 3: electric currents on the polar cap, source: http://www.windows2universe.org/earth/Magnetosphere/tour/tour_earth_magnetosphere_07.html&edu=high

1.4 Twin cell convection

If we look at the Dungey cycle ($B_z < 0$), we can figure out the movement of the open/closed field lines at the polar caps, which gives us the movement of the particles at the polar cap. Let us first explain, what leads to the so called "Twin-cell"-convection. If we have southward IMF, the orientation of the IMF and the earth's magnetic field are opposite to each other. Like discussed in chapter 1.1, magnetic reconnection on the day side of the earth can occur. The magnetic reconnection at the day side adds open magnetic field lines to the polar cap. In other words, day side reconnection adds magnetic flux to the polar cap. The deformation of the usually round open/closed-field boundary leads to a movement, which tries to "smooth out" the little bump in the boundary. Magnetic field lines naturally want themselves to orientate in a straight line.

Let us now look at the night side. If we look at the night side events, we know that magnetic reconnection occurs in the tail.(dungey cycle) Reasons for this happening are for example that through day side reconnection, magnetic flux is added to the tail, the magnetic pressure increases and leads in the end to magnetic reconnection in the tail. The magnetic reconnection to the tail closes open magnetic field lines. This means that the polar cap shrinks a bit

on the night side. In other words, magnetic flux is subtracted at the night side of the earth from the polar cap. The subtraction leads to a little dint in the polar cap. The magnetic field lines tend to "smooth out" the little dint, which leads to a corresponding movement.

In total, we expect now a movement, which is well known as the twin cell convection. The electric potential on the poles will give us the movement of the charged particles. The currents move along potential lines. The equipotential lines are therefore the trajectories of the charged particles. The currents are called "Hall"-currents. The hall-currents can be found in an altitude of approximately 110 km. (electro-jets)

2 Observations

All in all, we use five different sources and data to draw an image of what is going on with the interaction of the IMF with the earth's magnetic field. We will just shortly list the used methods in this report. A detailed description of each of the methods can be found in the separated chapters. First of all, we use a satellite called ACE (Advanced composition explorer) in order to get knowledge about the magnetic field of the IMF at earth's position, which tells us whether day side magnetic reconnection occurs or not. We also use Groundbasedmagnetometer to get data from the magnetic field on earth. Depressions in the magnetic field can finally be used to detect auroral sub storms. Moreover, we want to look at data from SuperDARN in order to know the movement of charged particles in the ionosphere. The data from AMPERE gives rise to Region 1 and Region 2 currents. Finally, the All-sky-Camera at Svalbard helps us to detect sub storms.

2.1 ACE

In this section, we want to discuss the data from the ACE. (Advanced Composition Explorer) The ACE satellite orbits the Lagrangian point L_1 .

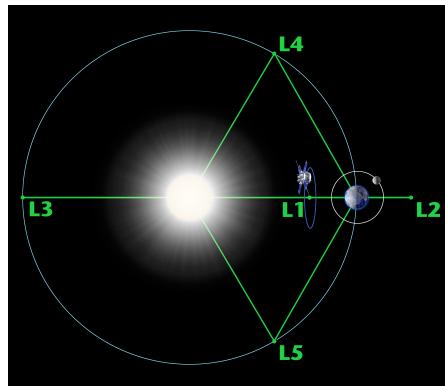
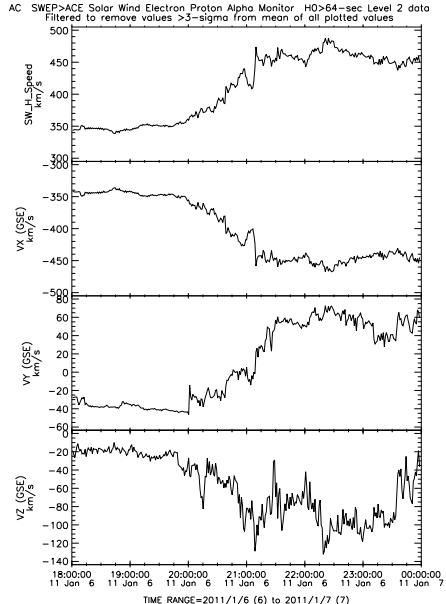


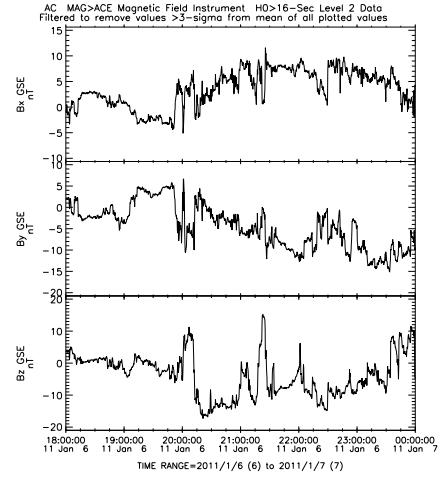
Figure 4: position of ACE, source: http://www.nasa.gov/mission_pages/sunearth/news/ace-15th.html

This Lagrangian point sits exactly between the earth and the sun.(see figure 4) The solar wind particles move radially outwards from the sun. Therefore, the data from ACE can be used to describe the IMF at earth's position when we take the time into account, which is needed from the solar wind to travel from the position of the ACE to earth. If we take a look at the position of ACE and the solar-wind speed at the position of ACE, we can calculate the time the IMF needs to travel from ACE to earth.

Finally, we know what the IMF looks like at earth's position at a given time on earth. This is important because we know from the chapter 1.1, that the B_z tells us something about magnetic reconnection on the day side. If $B_z > 0$ (or northward IMF orientation), there is no magnetic reconnection on the day side. There is no open magnetic flux added to the polar cap. However, if the B_z component is negative (or southward orientation of IMF) magnetic reconnection on the day side can occur and the dungey cycle kicks in. It is therefore very important to get data from the B_z component of the IMF. The other components of the IMF can be used at a later point to tell something about the specific shape of the twin-cell convection. In figures 5a and 6, we find data from the solarwindspeed and the ACE position. Both together will be used to determine the time delay of the IMF. The coordinate system GSE (Geocentric Solar Ecliptic). If we assume an average velocity of the solar wind (see 5a V_x) of $400 \frac{m}{s}$ and an average distance of $1.4802 \cdot 10^6$ km from ACE to earth, we get an average time delay of 1 hour and.(range between 55 minutes and 70 minutes) This is the time, the solar wind needs to travel from ACE to earth.



(a) Solarwind data



(b) magnetic field measured at the ACE

Figure 5: ACE data

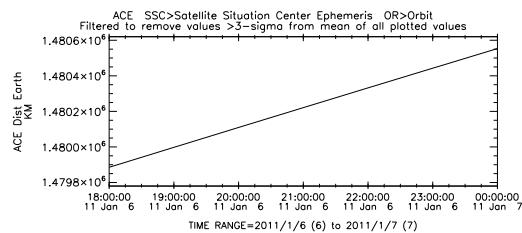


Figure 6: ACE distance from earth

2.2 Groundbasedmagnetometer

The Groundbasedmagnetometer we use is a collection of 32 station on the surface of the earth measuring the magnetic field. (by the use of magnetometers) The collection is called IMAGE (International Monitor for Auroral Geomagnetic Effects). We have chosen a night side event, which means that we have chosen a time when the All-sky-camera in Svalbard is on the night side. If we have a look at our data, we should be aware of indications concerning auroral sub storms. During a sub storm, the westward electro-jets increase. The currents cause according to Ampere's rule a depression in the north-south component of the magnetic field on earth. After the depression, the magnitude of the north-south component slowly increases to its original value. Observing this depression is therefore a first hint of a sub storm. Let us now have a closer look at the data.

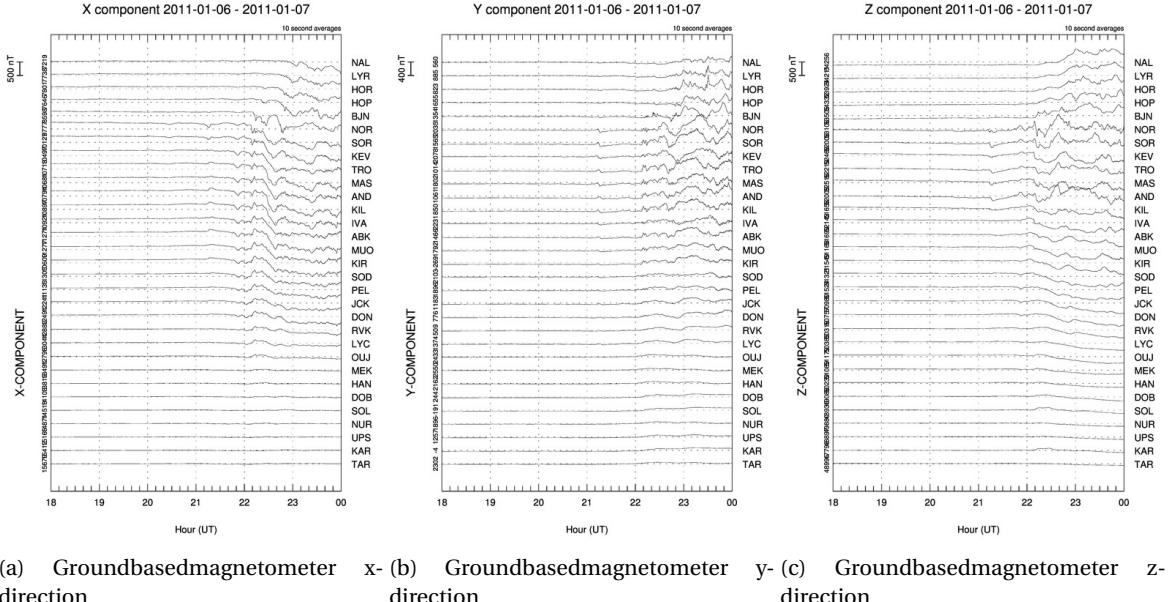


Figure 7: Data from Groundbasedmagnetometer, different components

In figure 7, you can see data from IMAGE. The figures 7a, 7b and 7c show different orthogonal components of the magnetic field. In figure 7a, you can see the south- north component of the magnetic field, which is the important component for us. The vertical axis of the plots show the time and the different horizontal lines are the magnetic field components measured at different stations on earth. On the right side, you can see abbreviations of the different stations. Note, that they are ordered concerning their latitude. The first thing, which strikes the eye is that the magnetic field component is mostly constant from 18:00 o'clock until approximately 22:00 o'clock. At approximately 22:00 o'clock (depends on the latitude of the station), we can see fluctuation of the magnetic field component and then a sudden depression of the magnetic field component. This is a strong hint to a sub storm. We can also see a poleward moving of the depression especially for the 8 most northward stations (KEV to NAL). We can also see that the stations MEK to TAR don't measure a depression of the magnetic field. The stations KEV to OIJ measure the depression approximately at the same time.

2.3 SuperDARN

SuperDARN stands for Super Dual Auroral Radar Network and this network consists of more than 30 low-power HF radars that look into Earth's upper atmosphere. (from mid-latitudes to polar regions) The radars can among others be used to measure the motion of charged particles in the ionosphere via the Doppler shift effect. (scattering of radiation) The motion of the plasma is interesting for us because we know what we expect from the orientation of the IMF is southward. The twin cell convection is therefore related to the dungey cycle. How can we visualize the motion of the hall currents? Unfortunately, we have not a complete data coverage all over the polar cap. Therefore, it seems to be useful to search for another way to visualize the motion of the plasma. The velocity of the hall currents is: $\vec{v} = \frac{\vec{E} \times \vec{B}}{B^2}$ and the formula for the electric potential is: $\vec{E} = -\nabla\phi$. It now turns out that lines of equal potential are perpendicular to the electric field, which is also the case for the hall currents. It can be shown that indeed the electric potential can describe the motion of the hall currents. The currents move then on equal potential lines. We use the data of the Radars (motion of plasma) to make a expansion in spherical harmonics for the electric potential. The plotting of the potential can then be used to describe the plasma flow.

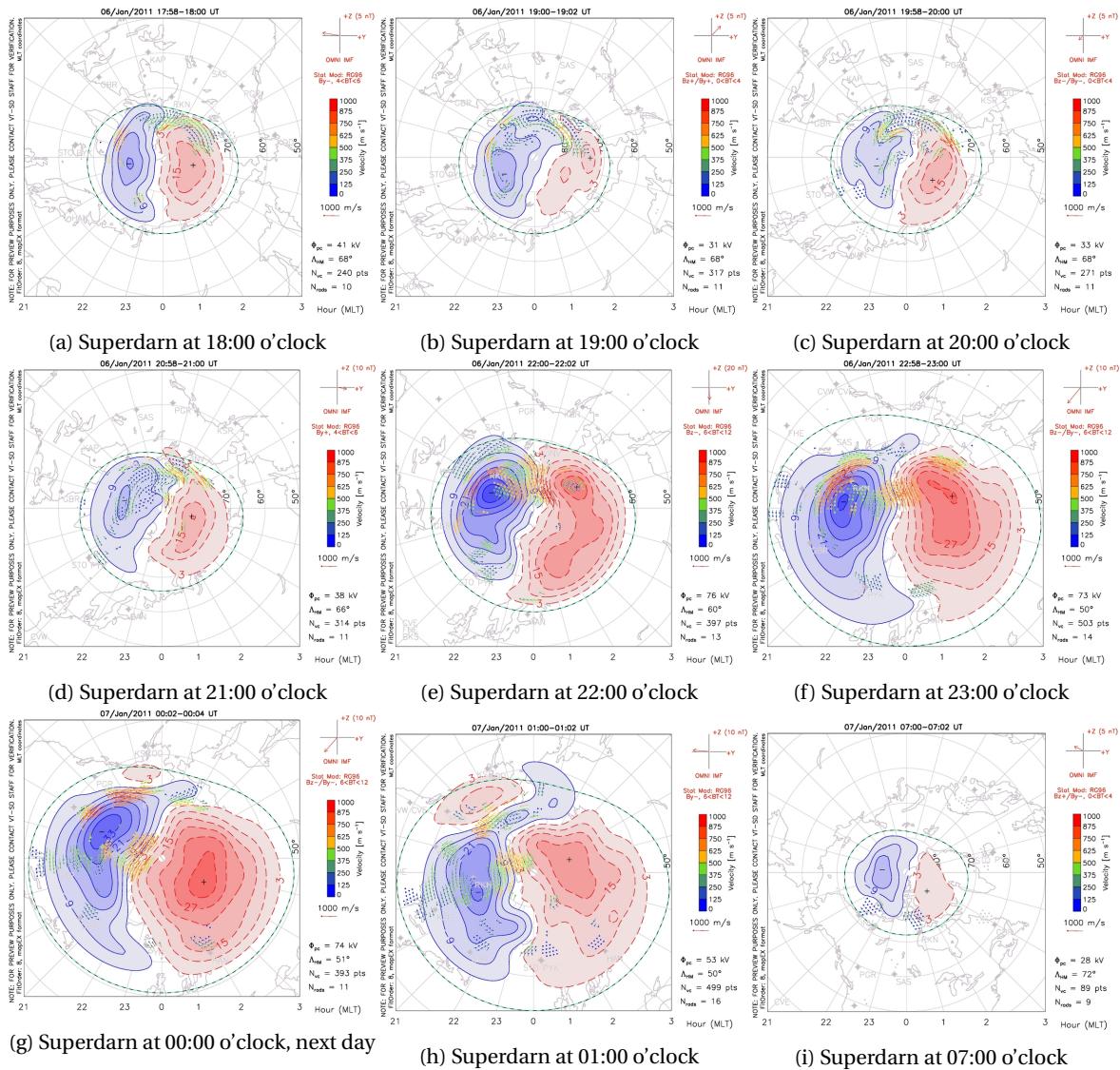


Figure 8: data from Superdarn for different times

2.4 AMPERE

AMPERE (Active Magnetosphere and Planetary Electrodynamics Response Experiment) is a earth observing system, which can be used to measure near-realtime magnetic field by 66 commercial satellites. The data from the magnetic field collected by AMPERE can be used to improve our understanding of space weather. In our project, we want to use AMPERE to determine the Birkeland currents. The Birkeland currents are field aligned currents. As described in the

theory part, flowing electric currents cause a curl of the magnetic field. If there are such FAC, we should also see a curl in the magnetic field measured by the satellites. We don't want to have a look at the explicit magnetic field data, but directly at the simulated current flow based on the measured magnetic field of AMPERE. This means that we expect to see the Region 1 currents at high latitudes and the Region 2 currents at lower latitude. In addition to that, we are interested to see how the currents change with respect to time.

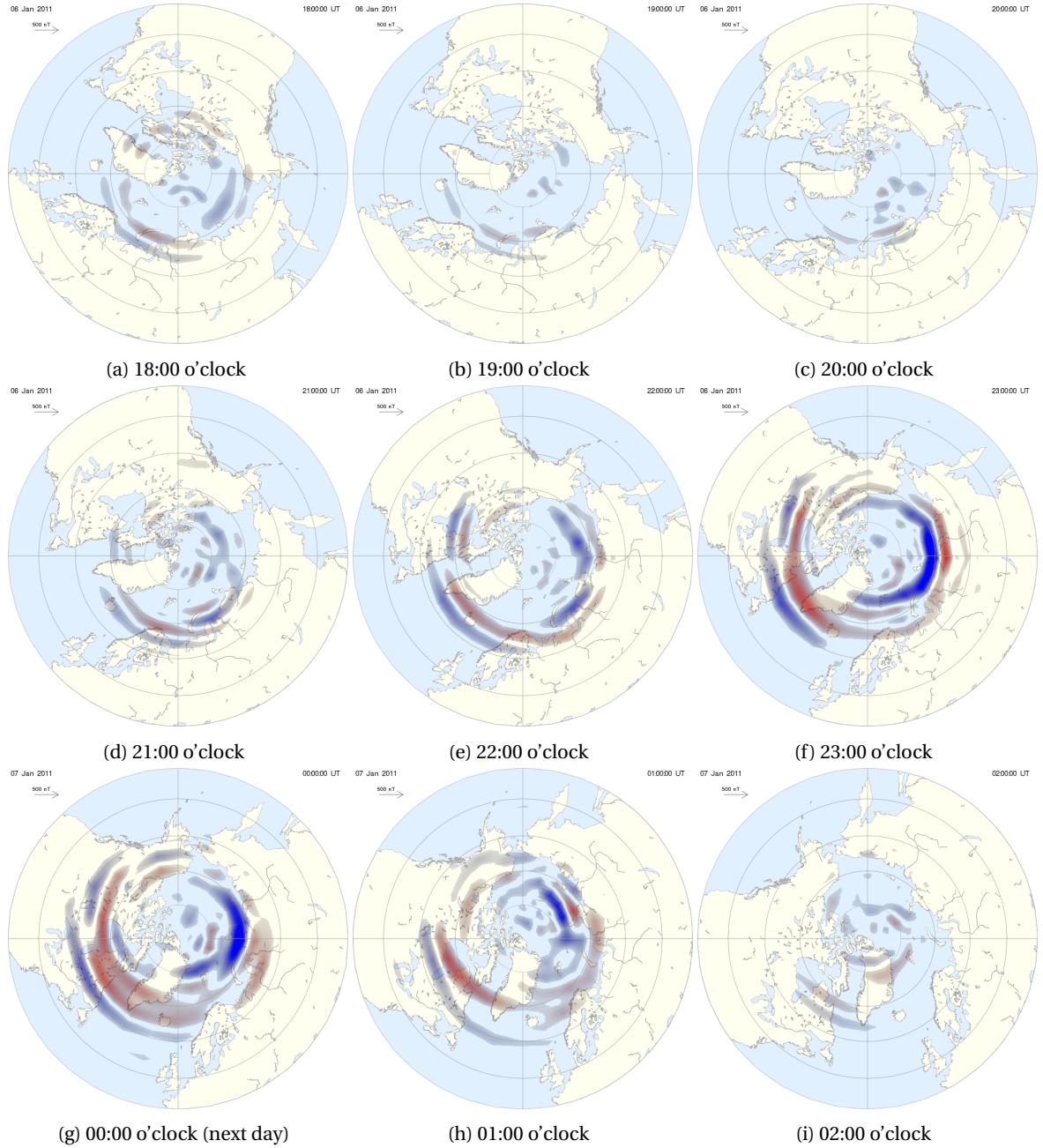
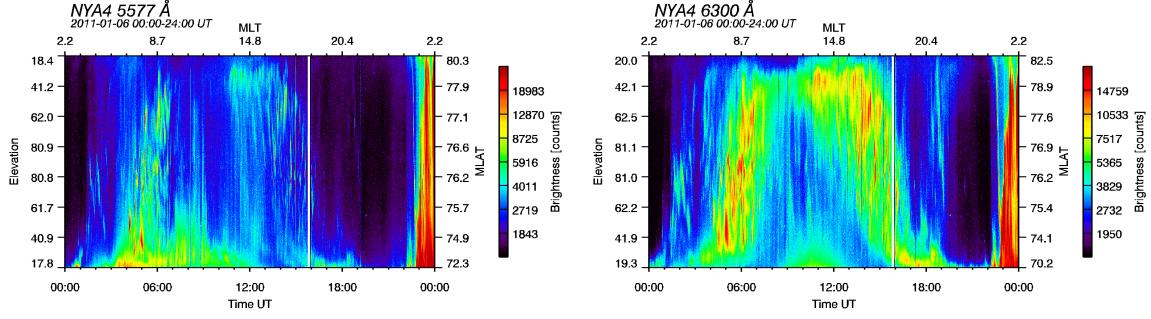


Figure 9: Data from AMPERE at different times

2.5 Svalbardimager



(a) SvalbardImager for 5577 Å

(b) SvalbardImager for 6300 Å

Figure 10: Overview over the Svalbard-Imagers

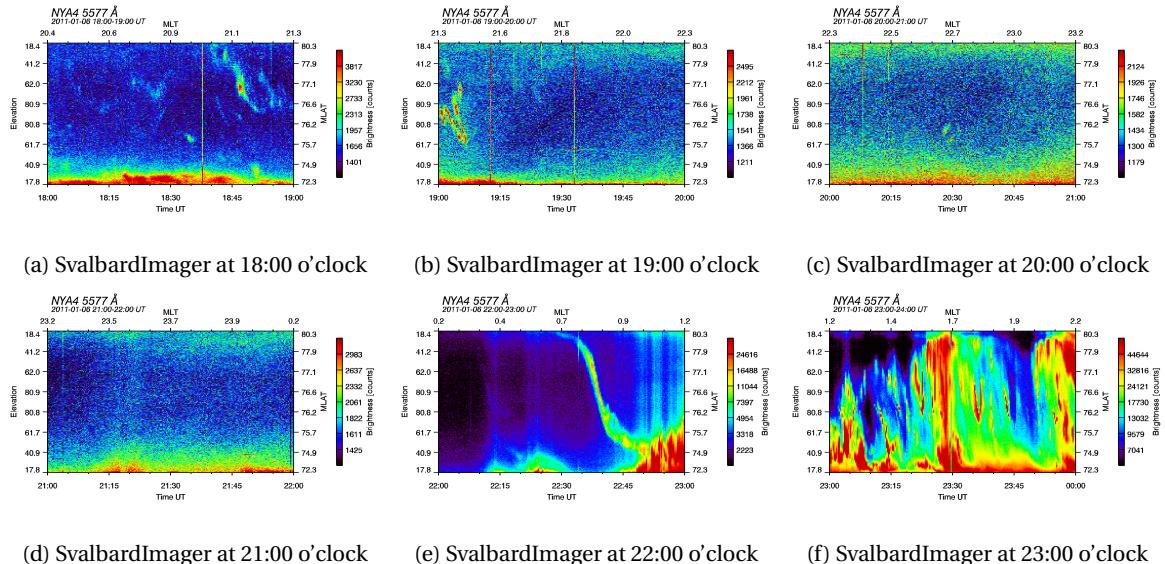


Figure 11: Ketograms from the SvalbardImager for $\lambda = 5577 \cdot 10^{-10} \text{ m}$, different times

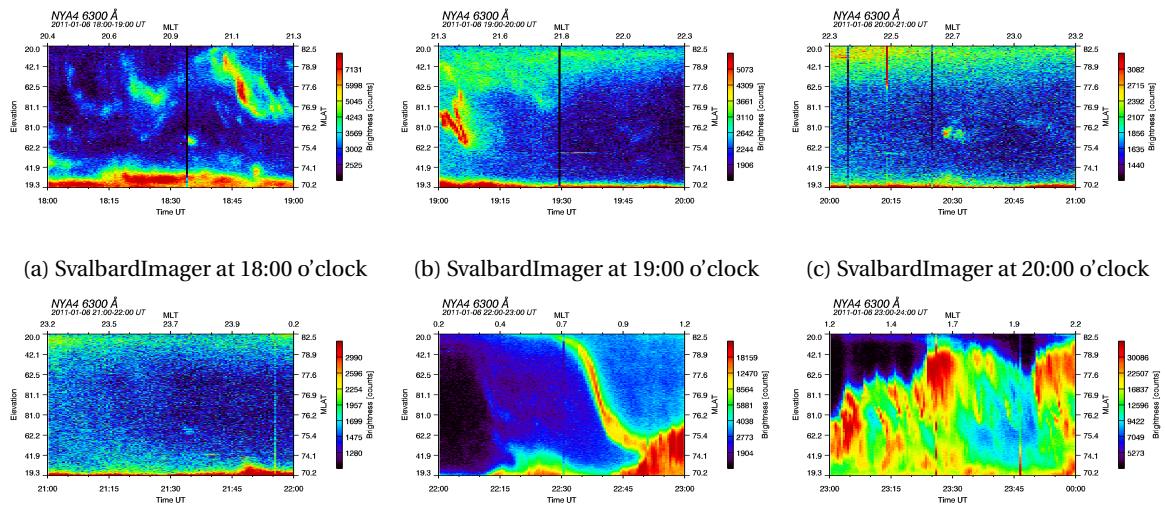


Figure 12: Ketograms from the SvalbardImager for $\lambda = 6300 \cdot 10^{-10} \text{ m}$, different times

3 discussion

4 conclusion