

Space Physics Project , FYS3610

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Introduction

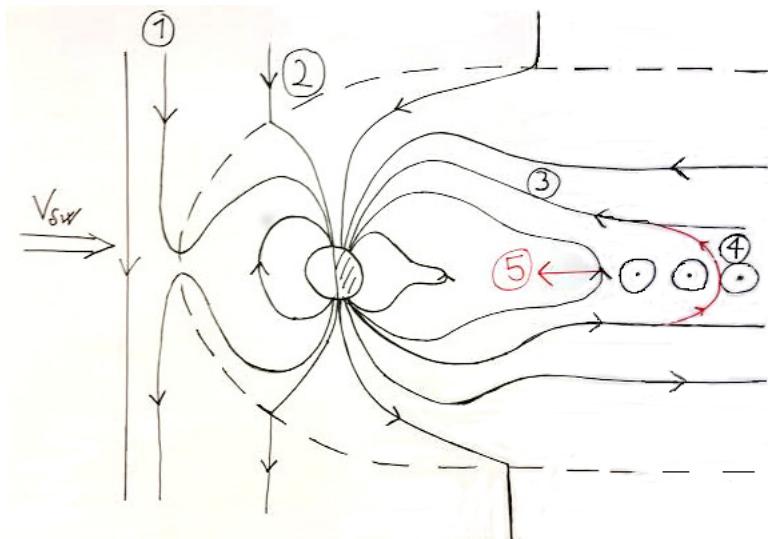
In this project we are going to analyse the ionospheric and auroral dynamics and its relation to the solar wind conditions. The two most relevant physical mechanisms are magnetic reconnection and frozen in magnetic fields. We will also investigate the Earth's currents system. Magnetic fields and current flows are governed by Magnetohydrodynamics, and coupled through Maxwell's equations. Finally, we will compare our observations to theory. We will investigate in a specific timeslot of 6 hours from the 6. January 2011 from 18:00 o'clock to the 07. January 2011 00:00 o'clock. We will not go into a detailed mathematical description of the different phenomena, but we will provide a clear overview of a few aspects of the ionospheric and auroral dynamics.

1 Theory part

In order to understand the movements and changes in the system of the earth and the sun, we want to introduce some basic concepts of space physics. A really important steady state picture is the Dungey cycle. It tries to explain what is going on with respect to the earth's magnetic field, the interaction of the earth's magnetic field with the Interplanetary Magnetic Field (hereafter IMF) and the current flow. One of the preconditions for the Dungey cycle is that the IMF is orientated southward. The Dungey cycle is as already mentioned a steady state picture. This means that the magnetic flux added to the open magnetic field lines through day side reconnection is equal to the magnetic flux subtracting from the open magnetic field lines through night side reconnection. However, this model does not fully resolve ionospheric and auroral dynamics at small timescales. (by small timescales we mean compared to the full cycle). We observe different phases such as adding magnetic flux (growing of the polar cap), some kind of explosive expansion (subtraction of magnetic flux from the polar cap) and a recovery phase. This phenomena is called a "Auroral substorm". In addition to that, we can extract information about what is going on with the earth-sun system by looking at different data sets containing information about the magnetic field on earth, the solar wind bulk speed, the auroral intensity, the current density and further data from different physical measurements. Based on these measurements we will interpret the received data and relate it to the theoretical predictions and the models we introduced in the theoretical part.

1.1 The Dungey cycle - a steady state picture

Figure 1: Dungey cycle; numbers belong to the description in section 1.1; the dashed line is the magneto-pause; usual lines describe magnetic field lines



If the magnetic field is northward orientated ($B_z > 0$ in GSM-coordinates) then there is no magnetic reconnection on the dayside. The IMF can not mix with the terrestrial magnetic field because everything is frozen in. The solar wind will be deflected slightly and flow around the Earth, and stretch far into space anti-sunward. As long as the magnetic field of Earth and the IMF have magnetic field lines which have parallel orientation. Then, everything stays frozen in and the field lines must simply follow the IMF flow past the Earth. However, if the IMF is southward orientated ($B_z < 0$), we have anti-parallel orientated magnetic field lines at the boundary of the terrestrial magnetic field and the IMF (at the magnetopause). Therefore magnetic reconnection occurs at the magnetopause, which causes a cyclic reconnection pattern at day and night-side. If we assume a steady state picture, we call the cycle, which only happens for southward IMF, the dungey cycle. Steady state picture means in this case that the rate of adding open magnetic

flux to the polar cap on day side is equal to the subtracting of open magnetic flux on the night side. A figure describing the dungey cycle with its different stages can be found in figure 1. The different numbers in this figure are explained below. The different parts of the cycle can be described as follows:

- 1 Opposite polarity field lines reconnect at magnetopause (if IMF southward orientated)
- 2 Magnetic field lines are still connected to the solar wind. (again frozen in) Therefore, the magnetic field lines are draged along with the solar wind in the direction of the magnetotail. (magnetic tension force)
- 3 Magnetic flux is added to the tail and compresses the plasma sheet. The magnetic pressure is opposed by the thermal pressure of the plasma. (magnetic pressure increases, plasma sheet is compressed into a thin sheet.)
- 4 Magnetic reconnection is triggered in the tail. The magnetic reconnection occurs far down the tail.
- 5 Reconnected magnetic field lines return to the day side, while some plasma is ejected away from the Earth. We can imagine these field lines as magnetic rubber bands with high tension, snapping back towards the Earth.

1.2 Auroral substorm

As already mentioned in the introduction, the assumption of a steady state picture is not always a appropriate description of what happens in reality. To specify this, the rate of adding magnetic flux and the rate of subtracting magnetic flux from the polar cap are in general not equal. However, after studying data from different measurements, Akasofu could state out in 1964 different phases of a cycle so called:

- 1 Growth phase
- 2 Expansion phase
- 3 Recovery phase

Compare this to figure 2, where the formation of a sub storm is indicated. In this figure number 5 belongs to the near Earth neutral line.

1.2.1 Growth phase

The day side reconnection dominates and the aurora moves to lower latitudes. The auroral activity increases in the cusp and open magnetic flux is added to the polar cap. The magnetic flux, which is added due to the day side reconnection makes the tail radius increase.(because the field lines are draged tail-wards with the IMF)

1.2.2 Expansion phase

Sudden onset of night-side reconnection at the near Earth neutral line (NENL). The magnetic reconnection at NENL leads to more explosive unloading of the stored energy. Sudden brightening of the night side aurora. Poleward and westward movement of the aurora (on average). Open magnetic field lines are being closed.

1.2.3 Recovery phase

The night side reconnection continues at smaller rates and the auroral display quietens. Furthermore, the polar cap contracts and the magnetosphere returns to pre growth phase "equilibrium".

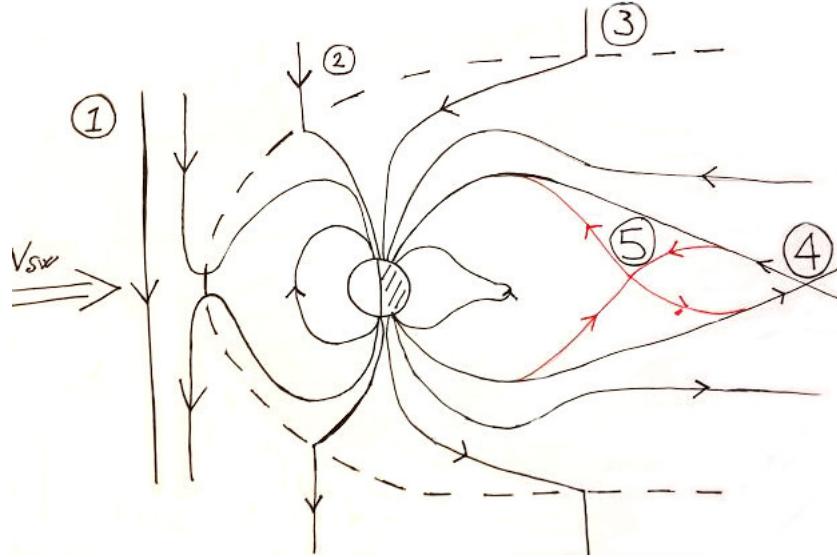
1.3 Currents

When we are dealing with the interaction of the earth's magnetic field with the IMF, it seems to be reasonable to investigate how the currents close to earth behave. The reason for this is due to the connection of magnetic fields and flowing currents. We know from Maxwell's equations that flowing electric currents cause a magnetic field. In particular, this relation is given by Ampere's law from Maxwell's equations:

$$\nabla \times \vec{B} = \mu_0 \cdot \vec{J} + \mu_0 \epsilon_0 \frac{\partial}{\partial t} \vec{E} \quad (1)$$

We can neglect the time dependence of the electric field in comparison with the \vec{J} -term. Therefore, we can derive the current flow by looking at the magnetic field. (in particular by looking at the curl of the magnetic field)(see measurement AMPERE) In general, we can express the current flow in a given system with given magnetic field and electric field as follows:

Figure 2: Auroral substorm



$$\vec{j} = en(\vec{v}_i - \vec{v}_e) = \sigma_p \vec{E}'_{\perp} + \sigma_H \frac{\vec{E}'_{\perp} \times \vec{B}}{B} + \sigma_{\parallel} \vec{E}_{\parallel} \quad (2)$$

Here, σ_p is the Pedersen conductivity, σ_H is the Hall conductivity and σ_{\parallel} is a conductivity, which belongs to the electric field. We notice that the different currents caused by different conductivities flow in orthogonal directions with respect to each other. The orientation of \vec{E} means parallel or orthogonal to the magnetic field. A discrete expression of the conductivities will not be needed in this report.

$$\vec{E} = \vec{E}_{\perp} + \vec{E}_{\parallel} \quad (3)$$

$$\vec{E}'_{\perp} = \vec{E}_{\perp} + \vec{v}_n \times \vec{B} \quad (4)$$

$$\vec{E}'_{\perp} \approx \vec{E}_{\perp} \quad (5)$$

1.3.1 Earth's current system

The magnetic field of the terrestrial magnetosphere is produced by superposition of magnetic fields from a variety of sources.[6][p.400] Near the earth's surface, the biggest contribution comes from the currents within the earth's liquid core. However, electric currents like the solar dynamo, the equatorial electrojet, the convection electrojets and the sub storm electrojets play also an important role. [6][p.400]. We want to give a short overview about the current system of the earth. We want to limit us to currents, which can actually be measured by the instruments used in this report. The currents, which flow near the polar cap can be described as follows: On the polar cap, we find Hall currents, Pedersen currents and Field-aligned currents.

The characterisation of these currents is shown in 1.3. When we want to describe the ionospheric currents, it is important to relate the gyro-frequency of electrons (ω_e) and ions (ω_i) to their collision frequency (γ_e, γ_i). The altitude profile of the collision-frequency of the two particles are not the same. This is due to different mass of the particles and therefore different particle-densities of the particles at given altitude. When electrons and ions gyrate much faster than they collide, both $\vec{E} \times \vec{B}$ -drift and there is no total current. This is the case for altitudes above 120 km. For altitudes between 80 km and 120 km, ions collide more often than they gyrate. Therefore, ions move in the direction of \vec{E} . The electrons still gyrate faster than they collide. Therefore, electrons still $\vec{E} \times \vec{B}$ -drift. In total, we can observe a current. Below 80 km altitude, electrons and ions collide more often than they gyrate. Both particles move along \vec{E} . There is no total current below 80 km altitude. The layer from 80 km to 120 km is called the dynamo layer. Magnetic reconnection on day and night side enrolls the so called twin cell convection. The twin cell convection belongs to a $\vec{E} \times \vec{B}$ -movement. Since this drift is in general not charge dependent, there is no total current, when electrons and ions both $\vec{E} \times \vec{B}$ -drift. As described above, there exists a layer, where electrons $\vec{E} \times \vec{B}$ -drift and ions move in the direction of \vec{E} . In this layer, we can observe the so called electron hall-currents. The movement of the electron hall currents is twin cell shaped. (compare chapter 1.4) This movement leads to a electric field, which can be described by $\vec{E} = -\vec{v} \times \vec{B}$. When only the ions move along \vec{E} , this leads to another current called the Pedersen current. Looking at the Pedersen-current-distribution, we can see that the Pedersen currents are not source free. However, it is necessary to close current systems. At the divergences of the Pedersen-currents, we promote currents. These currents are field aligned currents

and close the Pedersen-currents. In the following, we shortly want to locate the different currents. The electron Hall currents flow at lower altitudes than the Pedersen currents. The flow of the Hall currents is also associated with the so called twin cell convection. The ionospheric Pedersen currents are connected to the Field aligned currents (FAC). The FAC are on one side closed the magnetopause current and on the other side closed by the ring current. These field aligned currents were discovered by a Norwegian scientist, Kristian Birkeland in 1908. All current systems must to be closed, which is also the case. The current flow on the polar cap can be seen in figure 3.

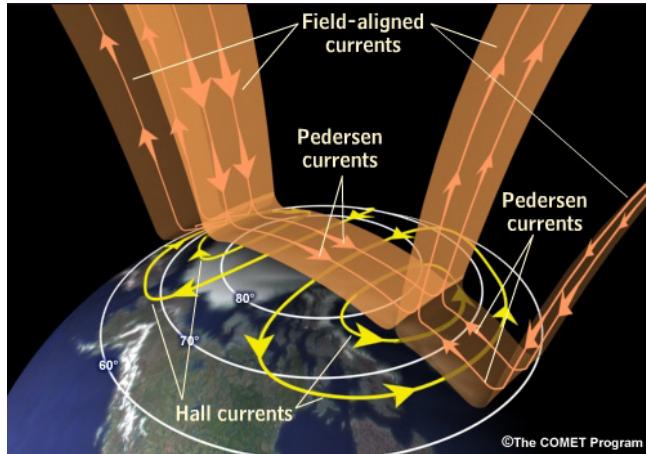


Figure 3: electric currents on the polar cap[1]

1.4 Twin cell convection

If we look at the Dungey cycle ($B_z < 0$), we can figure out the movement of the open/closed field lines at the polar caps, which gives us the movement of the particles at the polar cap. Let us first explain, what leads to the so called "Twin-cell"-convection. If we have southward IMF, the orientation of the IMF and the earth's magnetic field are opposite to each other and as we discussed in chapter 1.1, magnetic reconnection on the day side of the earth can occur. The magnetic reconnection at the day side adds open magnetic field lines to the polar cap. In other words, day side reconnection adds magnetic flux to the polar cap. The deformation of the usually round open/closed-field boundary leads to a movement, which tries to "smooth out" the little bump in the boundary. Magnetic field lines naturally want to orientate in a straight line.

If we now look at the night side events, we know that magnetic reconnection occurs in the tail.(dungey cycle) Reasons for this happening are for example that through day side reconnection, magnetic flux is added to the tail, the magnetic pressure increases and leads in the end to magnetic reconnection in the tail. The magnetic reconnection to the tail closes open magnetic field lines. This means that the polar cap shrinks a bit on the night side. In other words, magnetic flux is subtracted at the night side of the earth from the polar cap. The subtraction leads to a little dint in the polar cap. The magnetic field lines tend to "smooth out" the little dint, which leads to a corresponding movement.

In total, we expect now a movement, which is well known as the twin cell convection. The electric potential on the poles will give us the movement of the charged particles. The currents move along potential lines. The equipotential lines are therefore also the trajectories of the charged particles. The currents are called "Hall"-currents. The hall-currents can be found in an altitude of approximately 110 km. (electro-jets)

2 Observations

All in all, we use five different sources and data to draw an image of what is going on with the interaction of the IMF with the earth's magnetic field. We will just shortly list the used methods in this reports. A detailed description of each of the methods can be found in the separated chapters. First of all, we use a satellite called ACE (Advanced composition explorer) in order to get knowledge about the magnetic field of the IMF at earth's position, which tells us whether day side magnetic reconnection occurs or not . We also use The Ground-based magnetometer to get data from the magnetic field on earth. Depressions in the magnetic field can be used to detect auroral sub storms. Moreover, we want to look at data from SuperDARN in order to know the movement of charged particles in the ionosphere. The data from AMPERE gives us a map of the Region 1 and Region 2 currents. Finally, the All-sky-Camera at Svalbard helps us to monitor and the intensity of the substorms.

2.1 ACE

In this section, we want to discuss the data from the ACE. (Advanced Composition Explorer) The ACE satellite orbits the Lagrangian point L_1 .

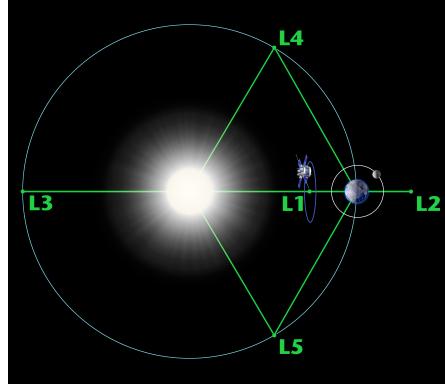


Figure 4: position of ACE [2]

This satellite sits somewhere on the line between the Sun and the Earth, where the gravitational forces are in equilibrium. (see figure 4) The solar wind particles move radially outwards from the sun. Therefore, the data from ACE can be used to describe the IMF at earth's position. Based on the solar wind bulk velocity and the distance from the satellite to the Earth, we make an approximation for when the solar wind reaches the Earth.

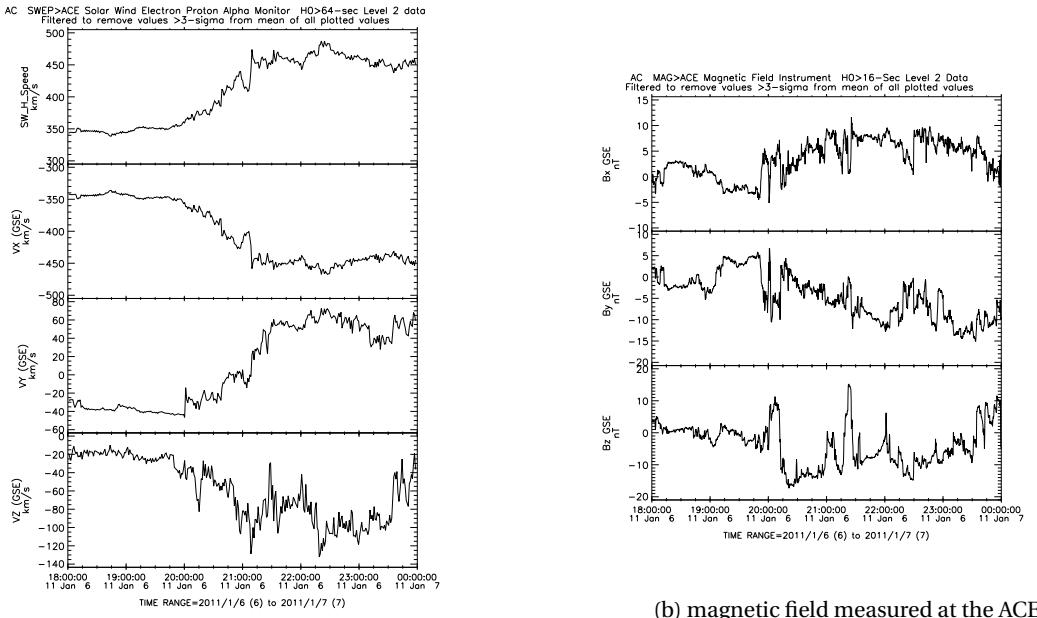


Figure 5: ACE data

We can therefore predict what the IMF looks like at Earth at a given time on earth. This is important because we know from the chapter 1.1, that the B_z tells us something about magnetic reconnection on the day side. If $B_z > 0$ (or northward IMF orientation), there is no magnetic reconnection on the day side. There is no open magnetic flux added to the polar cap. However, if the B_z component is negative (or southward orientation of IMF) magnetic reconnection on the day side can occur and the Dungey cycle kicks in. It is therefore very important to get data from the B_z component of the IMF. The other components of the IMF can be used at a later point to say something about the specific shape of the twin-cell convection. In figures 5a and 6, we find data from the solarwind speed and the ACE position. Both together will be used to determine the time delay of the IMF. The coordinate system GSE (Geocentric Solar Ecliptic). If we assume an average velocity of the solar wind (see 5a V_x) of $400 \frac{m}{s}$ and an average distance of $1.4802 \cdot 10^6 \text{ km}$ from ACE to Earth, we get an average time delay of 1 hour. (With values ranging from 55 minutes to 70 minutes.) From 18 o'clock to approximately 20:15 o'clock, the IMF at ACE's position is slightly positive. We can see that

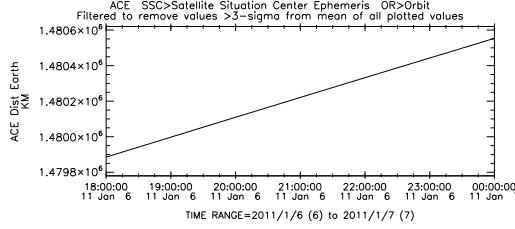
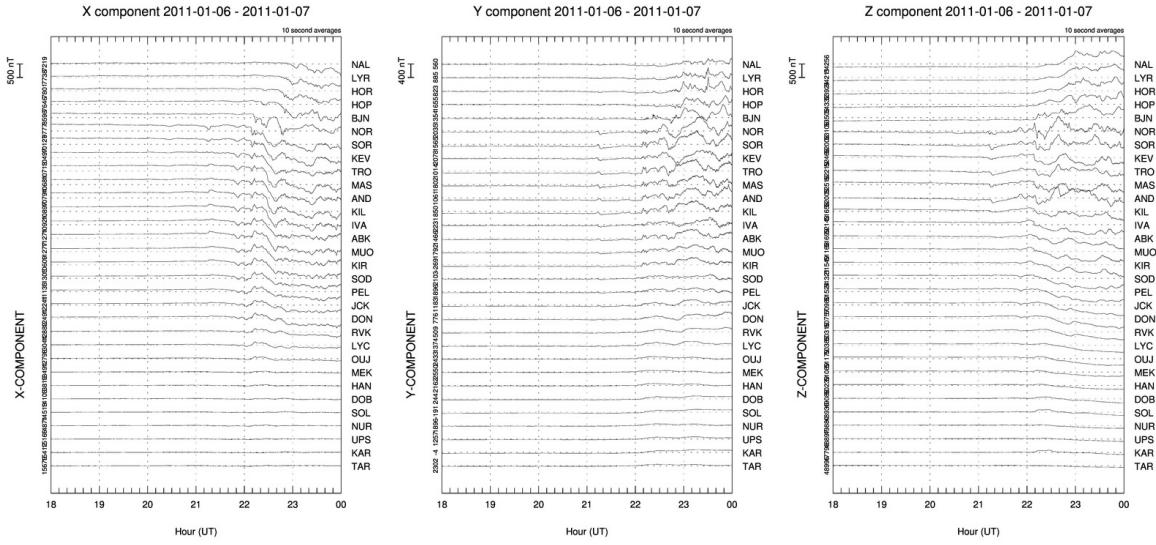


Figure 6: ACE distance from earth

the B_z -component of the IMF at the position of ACE turns negative at approximately 20:15 o'clock. (to $B_z \approx -12nT$) From there on, the IMF stays almost universally negative until 23:45 o'clock. At 21:20 o'clock, the B_z component turned for a short time to relatively high positive values. This sudden turn of the IMF might lead to what we call polar arcs, which are aurora that moves from the very high latitudes then southward towards the polar cap. Research suggest that these arcs might be somehow connected to auroral substorms, and that they are more prominent in times of northward IMF [7]. We do not know which type of polar arc this might be, as there are different types with specific characteristics, some of the different types are reviewed in [8].

2.2 Ground-based magnetometer

The Ground-based magnetometer we use is a collection of 32 station on the surface of the earth measuring the magnetic field. The collection is called IMAGE (International Monitor for Auroral Geomagnetic Effects). We have chosen a night side event, which means that we have chosen a time when the All-sky-camera in Svalbard is on the night side. If we have a look at our data, we should be aware of indications concerning auroral sub storms. During a sub storm, the westward electro-jets increase. The currents cause according to Ampere's rule a depression in the north-south component of the magnetic field on earth. After the depression, the magnitude of the north-south component slowly increases to its original value. Observing this depression is therefore a first hint of a sub storm. Let us now have a closer look at the data.



(a) Groundbasedmagnetometer x- (b) Groundbasedmagnetometer y- (c) Groundbasedmagnetometer z- direction direction direction

Figure 7: Data from Groundbasedmagnetometer, different components

In figure 7, you can see data from IMAGE. The figures 7a, 7b and 7c show different orthogonal components of the magnetic field. In figure 7a, you can see the south- north component of the magnetic field, which is the important component for us. The vertical axis of the plots show the time and the different horizontal lines are the magnetic field components measured at different stations on earth. On the right side, you can see abbreviations of the different stations. Note, that they are ordered concerning their latitude.

The first thing, which strikes the eye is that the magnetic field component is mostly constant from 18:00 o'clock until approximately 22:00 o'clock. At approximately 22:00 o'clock (depends on the latitude of the station), we can see

fluctuation of the magnetic field component and then a sudden depression of the magnetic field component. The first minimum of the measured magnetic field can be seen between 22:30 and 23:00 o'clock. This is a hint to a sub storm.(see chapter 3) We can also see a poleward moving of the depression especially for the 8 most northward stations (KEV to NAL). We can also see that the stations MEK to TAR don't measure a depression of the magnetic field. The stations KEV to OUJ measure the depression approximately at the same time.

2.3 SuperDARN

SuperDARN stands for Super Dual Auroral Radar Network and this network consists of more than 30 low-power HF radars that look into Earth's upper atmosphere. (from mid-latitudes to polar regions) The radars can among other things be used to measure the motion of charged particles in the ionosphere via the Doppler shift effect. (scattering of radiation) The motion of the plasma is interesting for us because we know what we expect from the plasma movement and would like to compare experimental data with theory. In the ionosphere at approximately 110 km altitude, we expect to see the twin cell convection if the orientation of the IMF is southward. The twin cell convection is therefore related to the dungey cycle. How can we visualize the motion of the hall currents? Unfortunately, we have not a complete data coverage all over the polar cap. Therefore, it seems to be useful to search for another way to visualize the motion of the plasma. The velocity of the hall currents is: $\vec{v} = \frac{\vec{E} \times \vec{B}}{B^2}$ and the formula for the electric potential is: $\vec{E} = -\vec{\nabla}\phi$. It now turns out that lines of equal potential are perpendicular to the electric field, which is also the case for the hall currents. It can be shown that indeed the electric potential can describe the motion of the hall currents, the currents move along equipotential lines. We use the data of the Radars (motion of plasma) to make an expansion in spherical harmonics for the electric potential. The plotting of the potential can then be used to describe the plasma flow.

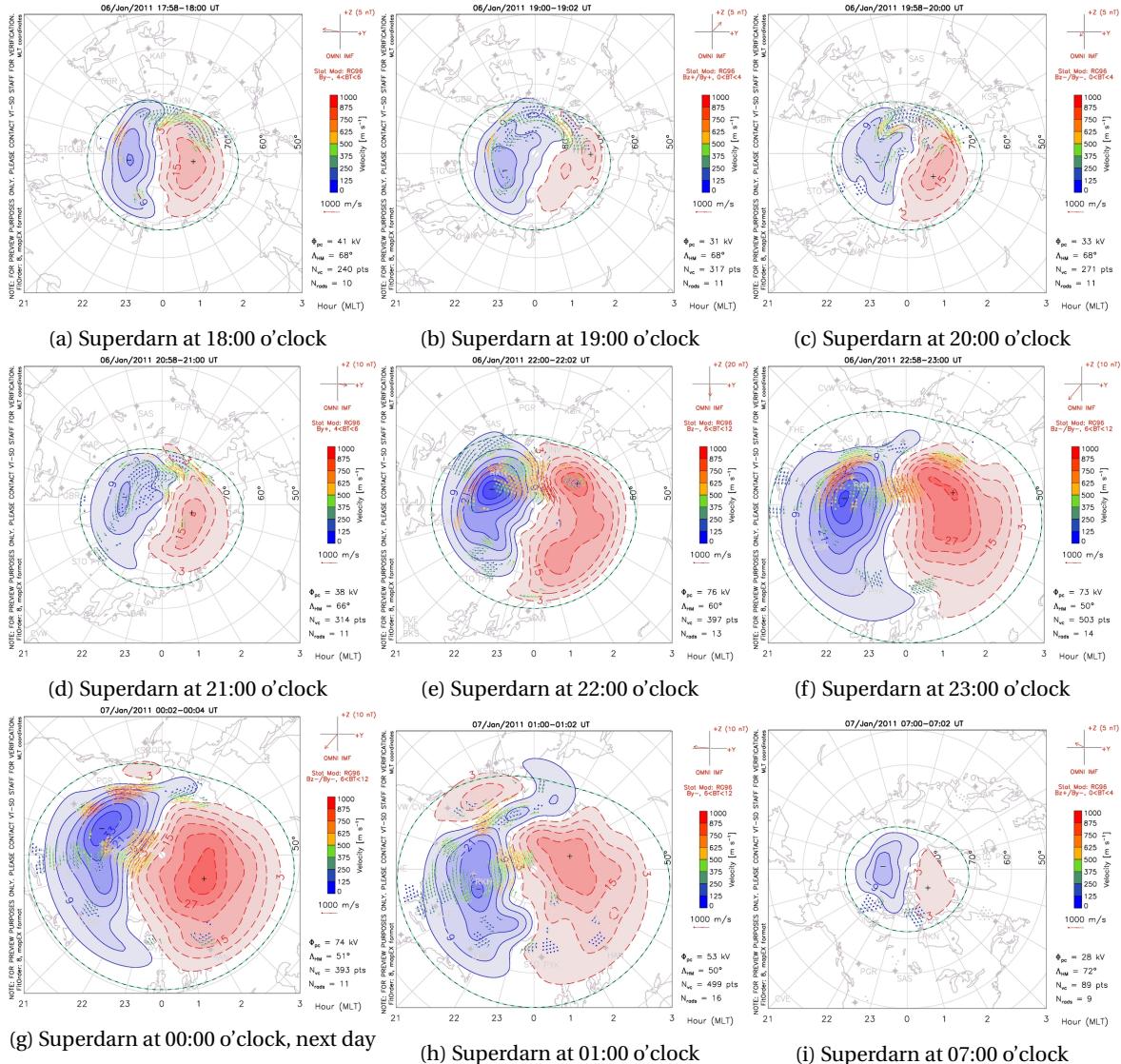


Figure 8: data from Superdarn for different times

The plot of the electric field on the polar cap is shown in figure 8 for different times. For all figures you can qualitatively see two regions. One region with a positive electric potential at 3 o'clock and one region with a negative electric potential at 9 o'clock. The orientation of these "cells" is qualitatively the same. (two cells are separated by a line, which goes from 12 o'clock to 6 o'clock) (look at figure 8) The strength of the electric potential varies a lot with time. We can compare the strength by comparing the contour lines of the potential. "The total strength of convection is characterised by the polar-cap potential drop" [6][p.302]. From 18 o'clock to 21 o'clock, we have approximately the

same strength of the electric potential. At this time, we also have approximately the same magnetic open flux through the polar cap. $\phi_{pc} = 31\text{ kV}$ to $\phi_{pc} = 41\text{ kV}$. The size of the polar cap stays the same in this time interval. We can also see that the area, where the electric potential is significantly different from 0 is fairly small and located at high latitudes. (above 70° north)(see figures 8a to 8d) At 22 o'clock, we can see a huge difference to the previous time step. The strength of the electric potential increases. The potential area is increased from about 66° to about 60° latitude. The open magnetic flux through the polar cap is now $\phi_{pc} = 76\text{ kV}$. (see figure 8e) At 23 o'clock on the 6th of January 2011, The area expands even more and the potential reaches its maximum from all the shown time frames. (expanded almost to 50° latitude) The strength of the electric field stays approximately the same as well as the open magnetic flux through the polar cap. (see figure 8f) At 00 o'clock on the 7th of January 2011, the properties described at 23 o'clock stay the same. However, we see on the top of figure 8g a small but separated blob of positive potential, which we couldn't observe at 23 o'clock. At 01 o'clock, we observe a weaker electric field than in the previous time step. We also observe some perturbations compared to the "two-cell-observation" from the previous steps.

2.4 AMPERE

AMPERE (Active Magnetosphere and Planetary Electrodynamics Response Experiment) is a earth observing system, which can be used to measure near-realtime magnetic field by 66 commercial satellites. The data from the magnetic field collected by AMPERE can be used to improve our understanding of space weather. In our project, we want to use AMPERE to determine the Birkeland currents. The Birkeland currents are field aligned currents. As described in the theory part, flowing electric currents cause a curl of the magnetic field. If there are such FAC, we should also see a curl in the magnetic field measured by the satellites. We don't want to have a look at the explicit magnetic field data, but directly at the simulated current flow based on the measured magnetic field of AMPERE. We expect to see two current regions. Region 1 currents flow in positive z-direction [GSM], while region 2 currents flow in the opposite, negative z-direction [GSM]. The configuration of the electric fields set up by the movement of magnetic flux tubes across the polar cap is such that we expect region 1 & 2 currents to lie inside each other like concentric circles. Furthermore we also expect that on the dayside the region one current is on the higher latitude while the region 2 currents are on the lower. Vice versa for the night side. We are interested to see how these currents change with respect to time.

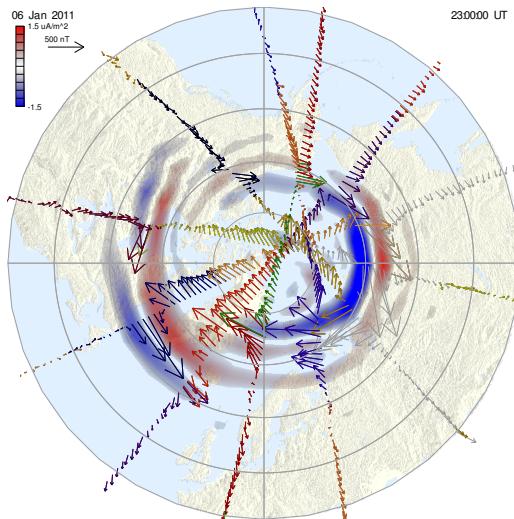


Figure 9: this figure shows the measured magnetic field vectors by the satellites at 23 o'clock of the 6th January and the field align currents in blue and red

In figure 9, we can see the measured magnetic field by the satellites visualized through little arrows. We convince ourselves that a curl in the magnetic field can be related to a area with blue or red colour, which belong to field aligned currents. Red colour means that the field aligned currents flow away from the surface of the earth and blue towards the earth.

With this pre-knowledge, we can now have a closer look at the data. From 18 o'clock (see figure 10a) until 20 o'clock(see figure 10c), we don't see significant currents. At 21 o'clock (see figure 10d), we see a small increase of the currents indicated by a bit more intensive red and blue colour. However, it is hard to derive a structure of the current-flow. At 22 o'clock (see figure 10e), the current flow is a bit more intense than one hour before. We can now see four different areas. An inner circle consisting of a current coming down from space to earth on the right half on the circle and of currents flowing into space the left part of the circle. The outer circle consists of a current coming down from space to earth on the left half on the circle and of currents flowing into space on the right part of the circle. This basic structure can also be seen at 23 o'clock and at 00 o'clock.(see figure 10f and 10g) The currents reach their maximum at approximately 23 o'clock. From there on, the currents decrease. At 02 o'clock (see figure 10i), we can say that there

aren't any significant currents any more.

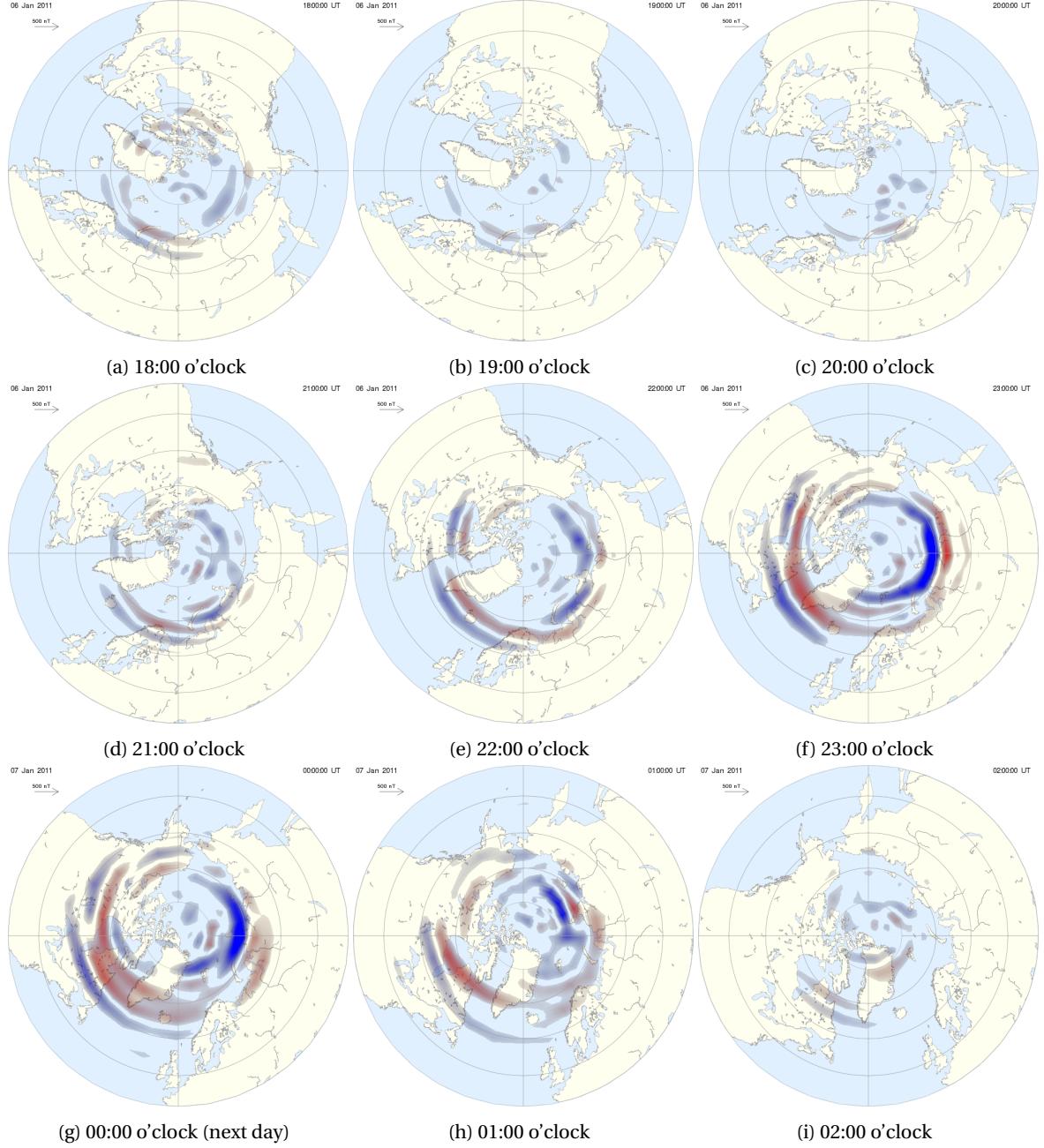


Figure 10: Data from AMPERE at different times

2.5 Svalbard All-Sky Imager

The Svalbard All-Sky Imager measures the intensity of two different wavelength all over the sky. Every pixel represents a small part of the sky. We look at two different wavelength ($\lambda_1 = 630\text{nm}$, $\lambda_2 = 557.7\text{nm}$). The reason for this is that we want to have a look at the most important constituents of the northern lights. λ_1 belongs to the relaxation from the 1D orbital to the 3p orbital of atomic oxygen and λ_2 belongs to the relaxation from the 1s orbital to the 1D orbital. We have different images for each of these wavelengths. (compare reference [6][p.469f]) in order to visualize the time-development of the intensity of this light, it is convenient to use a so called keogram. A keogram takes a slice from the center of each image at each time step and adds the different slices together in order to produce a timeline. Longitudinal movements are easily detected in keograms, however for the exact latitudinal movements one should check the full images. In the images, red areas represent relatively high intensity of the specific wavelength and black relatively low intensity. We should keep in mind that the intensities are rescaled for every picture in order to have a better visualization. The figure 11 shows a general overview of the whole day. (06.01.2011) Figure 11a shows the overview of the whole day for λ_2 and 11b shows the overview of the whole day for λ_1 . The interval from 00 o'clock to

18 o'clock is not interesting for us. We can see that from 18 o'clock to approximately 22 o'clock, the intensity on both images are really low; whereas from 22 o'clock to 00 o'clock (7th January) we can see that we have high intensity over all longitudes represented by the red colour.

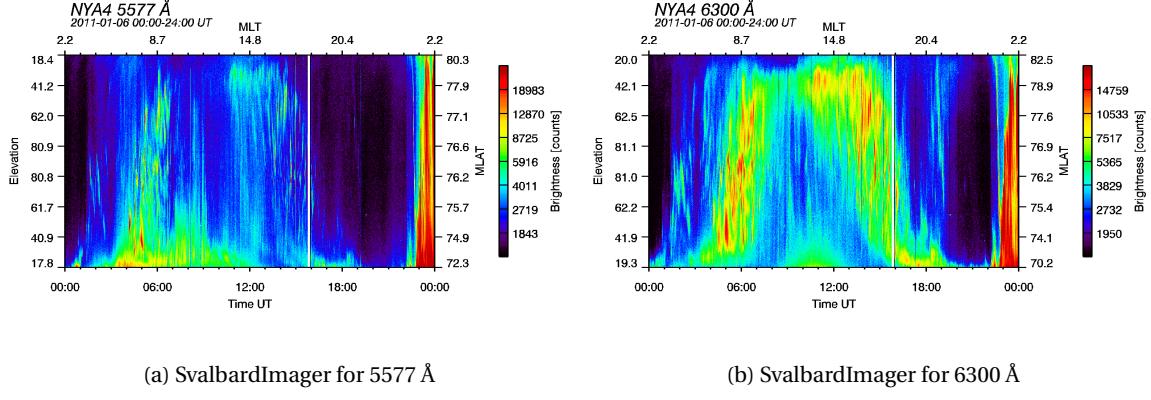


Figure 11: Overview over the Svalbard-Imagers

In the figure 11a, we can see keograms from the whole period, with one hour intervals. In figure 12a to 12d, we don't see anything special. The red area at below 20° is associated with the polar oval. At 22:30, we can see the first movement of an area of high intensity of light from north to south and at 22:45, we see a higher intensity of the light up to 76.5°latitude. (see figure 12e) At 23 o'clock, we can see a lot of areas with high light-intensity. The longest time interval of high light intensity all over the latitudes is at approximately 23:30 o'clock nd lasts about 5-10 minutes. At 23:45 o'clock, the intensity is low and at 23:55, higher intensities are again measured at high latitudes. In figure (12e) we observe an interesting filament which could be a polar arc. We can see it moving from the very high latitudes at 22:30, then 15 minutes later it seem to merge with the polar cap. At the same time as the polar arc merges with the polar cap, over all intensity is on the rise, and shortly after we observe a very dynamic and active auroral display.

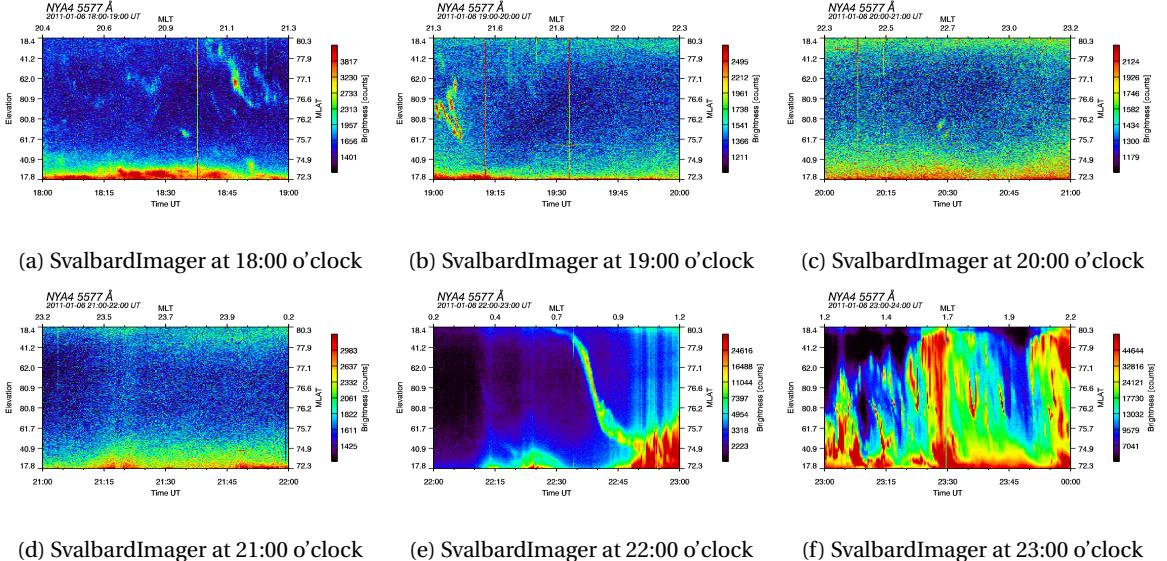


Figure 12: Ketograms from the SvalbardImager for $\lambda = 5577 \cdot 10^{-10} \text{ m}$, different times

The Svalbard imager for the wavelength of $\lambda = 630 \text{ nm}$ shows qualitatively the same as the imager of $\lambda = 577.7 \text{ nm}$. This can be seen in figure 13. The figures 13a to 13f show each intervals of 1 hour.

3 Discussion

In this section, we want to explain our observations. We will also try to connect our observation with what we expect from theory. It is also very interesting to see, whether the data from the different measurements can be interpreted in a sense-full, physical way.

The Dungey cycle, auroral substorms and the twin-cell convection are all dependent on the IMF. Especially the sign of the B_z -component is really important because it determines, whether magnetic reconnection occurs or not. (compare 1.1) Therefore, it makes sense to look first at the data from ACE. If we assume that the solar wind needs on average 1 hour to travel from ACE to earth, we can say that the B_z -component is positive from 19 o'clock to 21:15 o'clock at earth's position. At approximately 21:15, the anti-parallel magnetic field should reach Earth. While the B_z component is negative, we expect that magnetic reconnection on day side kicks in. When magnetic reconnection on day side occurs, open magnetic flux is added to the polar cap. If magnetic reconnection on the night side doesn't take place, the polar cap should grow. Southward orientation of the IMF is needed in order enter the growth phase of a substorm. On average, we observed a southward orientation of the magnetic field until 01:45 o'clock of the 7th of January 2011. However, we can see a sudden peak in the B_z component at 22:20 o'clock at Earth. The peak could be the coupled to the polar arc like filament we observe approximately one hour later with the All-sky imager.

Next, we want to look for signs of a substorm in process. A good method to detect substorms are to look at the data from the ground-based-magnetometers. A substorm normally produce westward electro-jets, which lead to a magnetic field according to Amperes law. In the end, if a substorm kicks in, this leads to a depression of the south-north component of the magnetic field measured on the ground. A eastward eletrojet would lead to a increase of the south-north component of the magnetic field measured on the ground. We now have a look at the data from 2.4. As described there, we can see a depression of the south-north component of the magnetic field at approximately 22 o'clock. The depression reaches a minimum at approximately 22:30 o'clock. (depending on the specific measuring station) This depression is a strong hint to a substorm happening in between 22 o'clock and 23 o'clock. At the stations more southward, you can see that there is first a increase of the south-nor-magnetic field and than a depression. The most likely interpretation is that there is a local current flowing eastward, which leads to an increase of the magnetic field. The reasons for that can be various. Depending on the station, eastward electrojets must occur. This depends also on the shape of the twin cell convection at the specific time. An important parameter to look at would be the B_y -component of the IMF at this time because this determines the specific shape of the twin cell convection.

Furthermore, we want to have a look at the data from SuperDARN. (see chapter 2.3) We notice, that the shape of the electric potential fit the theoretical predictions. We can see two cells, which can be characterized through a region with negative electric potential (blue) and a region with positive electric potential (red). In the region with negative electric potential, the charged particles move clockwise and in the positive region the particles move anti-clockwise. Also the orientation of the two cells fit. The direction to the sun is in the figures 8 at 12 o'clock. (referring to the hours of a day which, letting the Earth rotate so the system is relative to the sun.) The twin cells should be separated by a line from approximately 6 o'clock to 12 o 'clock. (directions) This depends also on the components of the IMF. The orientation of the twin cells seems to be right.

Looking at the time development of the electric field, we can now connect the increase of the electric potential from 21 to 22 o'clock with other physical phenomena. From section 2.1, we know that the negative turn of the B_z component of the IMF is at about 21:45 o'clock on earth. A southward IMF results in magnetic reconnection on day side. Magnetic reconnection on the day side is directly coupled to the twin-cell convection pattern. The twin cell convection should increase, which is exactly what we can see in the data. From the Ground-based magnetometer, we got a strong hint of an auroral substorm kicking in between 22:00 o'clock and 23:00 o'clock. Sub storms decrease the total magnetic flux in the polar cap. We can see that from 22 o 'clock to 23 o'clock, that ϕ_{pc} stays approximately constant. All in all, we can say that the theory matches the experimental data concerning this measurement.

We want now to discuss the data from the Measurement AMPERE.(observations described in chapter 2.4) The curl of the magnetic field measured parallel to the surface of the earth determines the current flow perpendicular to the surface to the earth. The field aligned currents are coupled to the Dungey cycle. To be more precise, The hall currents are driven by the dungey cycle. The hall currents lead to a electric field at a specific altitude. Pedersen currents (mostly ions) follow this electric field. The pederson currents are not source free. At the spots, where it come to a accumulation of charge, we expect field aligned currents to close this current system. Depending on the sign of the accumulated charge, the field aligned currents flow into space or in the direction of the earth's surface. Derived from this idea, e expect region 1 and region 2 currents. Both can be seen in the data. Since the field aligned currents are as described in chapter 1.3.1 coupled to the twin cell convection, we expect also an increase of FAC when twin cell convection increases. The data support this argument since we can see an increase of FAC from 21 o'clock to 00 o'clock, which fits to the development of the twin cell convection. This means that we have qualitatively a good match between theory matches and experiment.

Finally, let us look at the Svalbard-imager. The first significant development in the imager can be seen at approximately 22:35 o'clock. A small brightened area moves equator-ward and westward. We could not relate the movement of the auroral arc to a specific physical phenomena, but it seems to trigger the onset of the first substorm.

From reference [6][p.419, Fig. 13.14.], we know that "sudden changes in the solar-wind dynamic pressure and,

more frequently, northward turnings of the IMF can trigger the sudden unloading of stored energy".

At approximately 22:20, the IMF- B_z -component on Earth turns suddenly northward and southward again. It is possible that this sudden change triggered caused the polar arc which then in turn triggered the the substorm. At 22:45 o'clock, we can see an increase in intensity of both wavelength for lower latitudes. This seems to be the first substorm detected this night. The time of the substorm seems to be reasonable. In the time before the substorm, we know from the data gathered from ACE that the B_z component on the IMF was negative from 22:15 o'clock and then on. In this time, energy should have been stored in the magnetotail of the earth by adding magnetic flux.

4 Conclusion

5 Additional plots

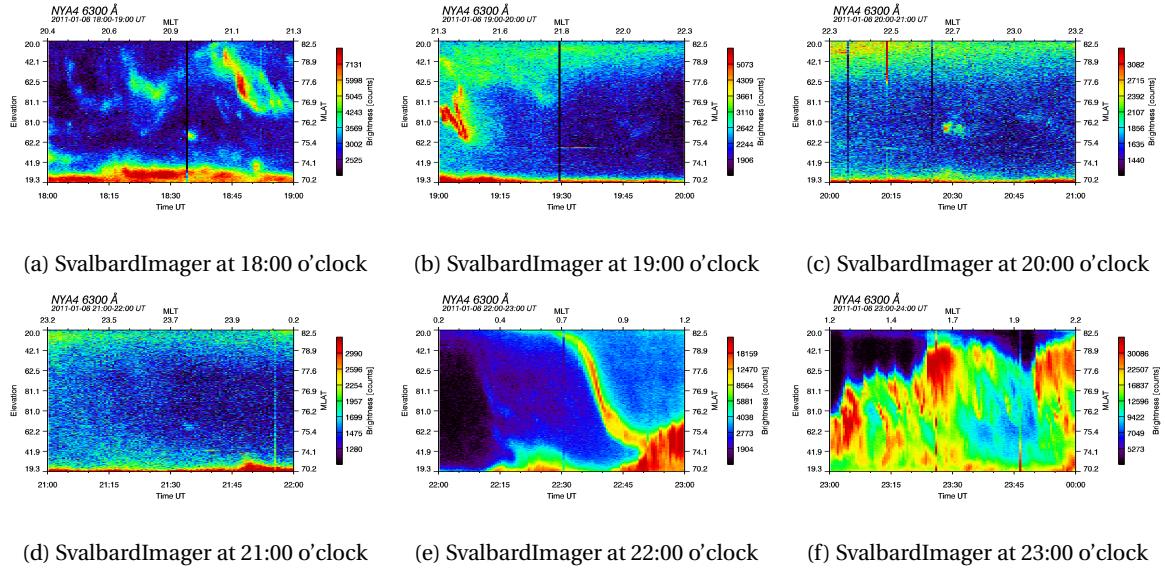


Figure 13: Ketograms from the SvalbardImager for $\lambda = 6300 \cdot 10^{-10} \text{ m}$, different times

References

- [1] http://www.windows2universe.org/earth/Magnetosphere/tour/tour_earth_magnetosphere_07.html&edu=high, 27.11.2015
- [2] http://www.nasa.gov/mission_pages/sunearth/news/ace-15th.html, 27.11.2015
- [3] Gerd W. Prölss: Physics of the Earth's Space Environment (An Introduction). Bonn dec 2003
- [4] Sandholt, P.E., C.J. Farrugia, J. Moen, Ø. Noraberg, B. Lybekk, T. Sten, and T.L. Hansen, A classification of dayside auroral forms and activities as a function of IMF orientation, *J. Geophys. Res.*, 103, 23,325-23,345, 1998.
- [5] Ruohoniemi, J. M., R. A. Greenwald, Dependencies of high-latitude plasma convection: Consideration of interplanetary magnetic field, seasonal and universal time factors in statistical patterns, *J. Geophys. Res.*, 110, A09204, doi:10.1029/2004JA010815, 2005.
- [6] Kievelson and Russel, Introduction to Space Physics, Chapters 9, 13, and 14
- [7] J.S. Murphree, J.B. Austin, D.J. Hearn, L.L. Cogger, R.D. Elphinstone, J. Woch Satellite observations of polar arcs, *J. Atmospheric and Terrestrial Physics* Volume 56, Issue 2, February 1994, Pages 265–284
- [8] L. Zhu, R.W. Schunk, J.J. Sojka, Polar cap arcs: a review, Revised 15 July 1996, Accepted 2 August 1996, Available online 5 June 1998, Page 8