

Lecture 3 — Energy and Momentum

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- In general, the trace of a matrix g is the dimensionality of manifold described by g :

$$\text{Tr}(g) = \dim(\text{manifold}_g)$$

- Results from special relativity:

$$E^2 = (mc^2)^2 + (pc)^2 \rightarrow m^2 + p^2$$

$$p = mv\gamma \Rightarrow E = m\gamma$$

- In proper covariant notation, the four velocity of a particle on $x^\mu(t)$ is:

$$U^\mu = \frac{dx^\mu}{d\tau} \rightarrow \eta_{\mu\nu} U^\mu U^\nu = -1 \text{ (due to definition of } \tau \text{)}$$

- This makes the four momentum:

$$p^\mu = mU^\mu$$

$$P^0 = E$$

$$P^i = \vec{p} \rightarrow \begin{pmatrix} E \\ p^1 \\ p^2 \\ p^3 \end{pmatrix}$$

$$p_\mu p^\mu = -m^2$$

- Thus, we may write:

$$E^2 = m^2 + p^2$$

- With a Λ :

$$p^{\mu'} = \begin{pmatrix} \gamma m \\ mv\gamma \\ 0 \\ 0 \end{pmatrix} \text{ corresponding to } (-v)!$$

- Force in Spacetime

$$f^\mu = m \frac{d^2}{d\tau^2} x^\mu(\tau) = \frac{d}{d\tau} p^\mu(\tau)$$

- For electromagnetics, we may write:

$$f^\mu = qU^\lambda F_\lambda^\mu \text{ (the power of symmetry)}$$

- Particle to Fluid

- From individual particles, we get fluid elements with: ρ, P , viscosity,...

- Energy-Momentum Tensor (Stress-Energy Tensor)

- We want to derive $T^{\mu\nu}$, or the flux of p^μ across x^ν surface (Recall $R_{\mu\nu} - \frac{1}{2}Rg_{\mu\nu} = 8\pi GT_{\mu\nu}$)
- Since time is now a dimension, there is a flux for particles moving solely in time
- Imagine two cases:

- * At Rest

- Particles at rest have a flux in time = x^0
- $T^{00} = \rho = mn$, where m is the mass and n is the number density
- $T^{\mu'\nu'} = \Lambda_{\alpha}^{\mu'} \Lambda_{\beta}^{\nu'} T^{\alpha\beta}$

- * In a boosted frame

- There is now a flux across x and x'
- Particles are no longer “at rest”

- Dust: Particles at rest with respect to each other

- At rest:

$$T^{\mu\nu} = \begin{pmatrix} \rho & 0 & 0 & 0 \\ 0 & 0 & 0 & 0 \\ 0 & 0 & 0 & 0 \\ 0 & 0 & 0 & 0 \end{pmatrix} = \rho U^\mu U^\nu$$

- Perfect fluid (with density ρ and pressure p), the fluid is isotropic (same in all directions)

$$T^{\mu\nu} = \begin{pmatrix} \rho & 0 & 0 & 0 \\ 0 & p & 0 & 0 \\ 0 & 0 & p & 0 \\ 0 & 0 & 0 & p \end{pmatrix} = (\rho + p)U^\mu U^\nu + p\eta^{\mu\nu}$$

- * Pressure is used because energy and mass are then interchangeable
- * This kinetic energy really makes a box more “massive”
 - Imagine protons (1% quark mass)

- Equation of State

$$w = \frac{p}{\rho}$$

- Will appear in cosmology
- $w = 0$ for matter (“dust”) \rightarrow non-relative
- $w = 1/3$ for photons (CMB)
- $w = -1$ (?) for dark energy $T_\Lambda^{\mu\nu} = -\rho_{vac}\eta^{\mu\nu}$ (absolute energy is important)

- Conservation of Energy-Momentum

$$\partial_\mu T^{\mu\nu} = 0$$

- For $\nu = 0$ conservation of energy
- For $\nu = k = 1, 2, 3$ conservation of momentum

- Classical Field Theory

$$\text{Action: } S = \int dt \underbrace{L(q, \dot{q})}_{\text{Lagrangian}}$$

- $L = K - V$ (kinetic minus potential energy)
- Equations of motion:

$$\frac{\partial L}{\partial q} - \frac{d}{dt} \left(\frac{\partial L}{\partial \dot{q}} \right) = 0$$

- Move to a field:

$$q \rightarrow \{\Phi^i(x^\mu)\}$$

- * Quantize Φ to particles

$$S = \int d^4 \times \mathcal{L}(\Phi^i, \partial_\mu \Phi^i)$$

- Integrate the Lagrangian density
- Our field will be the metric of spacetime
- \mathcal{L} and symmetry are powerful tools